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JWST-IPA: Chemical Inventory and Spatial Mapping of Ices in the Protostar HOPS 370 —Evidence for an Opacity Hole and Thermal Processing of Ices

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Abstract

The composition of protoplanetary disks, and hence the initial conditions of planet formation, may be strongly influenced by the infall and thermal processing of material during the protostellar phase. The composition of dust and ice in protostellar envelopes, shaped by energetic processes driven by the protostar, serves as the fundamental building material for planets and complex organic molecules. As part of the JWST General Observers program, “Investigating Protostellar Accretion,” we observed an intermediate-mass protostar HOPS 370 (OMC2-FIR3) using NIRSpec integral field unit and Mid-Infrared Instrument medium-resolution spectroscopy. This study presents the gas and ice phase chemical inventory revealed with the JWST in the spectral range of $\sim 2.9\text{--}28\ \mu\text{m}$ and explores the spatial variation of volatile ice species in the protostellar envelope. We find evidence for the thermal processing of ice species throughout the inner envelope. We present the first high-spatial resolution ($\sim 80\ \text{au}$) maps of key volatile ice species H_2O , CO_2 , $^{13}\text{CO}_2$, CO , and OCN^- , which reveal a highly structured and inhomogeneous density distribution of the protostellar envelope, with a deficiency of ice column density that coincides with the jet/outflow shocked knots. Further, we observe high relative crystallinity of H_2O ice around the shocked knot seen in the H_2 and OH wind/outflow, which can be explained by a lack of outer colder material in the envelope along the line of sight due to the irregular structure of the envelope. These observations show clear evidence of thermal processing of the ices in the inner envelope, close to the outflow cavity walls, heated by the luminous protostar.



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Unified Astronomy Thesaurus concepts: Protostars (1302); Astrochemistry (75); Star formation (1569); Stellar jets (1607); James Webb Space Telescope (2291)

1. Introduction

Planetary systems are the natural by-products of the star formation process (S. W. Stahler & F. Palla 2004; L. Hartmann 2009). Star formation begins with the collapse of a slowly rotating molecular cloud, followed by the formation of a central protostar and a circumstellar disk due to angular momentum conservation (R. K. Ulrich 1976; P. Cassen & A. Moosman 1981; S. Terebey et al. 1984; F. H. Shu et al. 1987; L. Hartmann 2009). The circumstellar disk serves as the birthplace of planets and is widely known as a protoplanetary disk (P. J. Armitage 2020; also see recent review on planet formation by J. Drażkowska et al. 2023). The origin of the chemical composition of disks and planets, however, has been a matter of lengthy debate.

The debate over the origin centers around two main hypotheses: 1) chemical inheritance and evolution and 2) reset. Chemical inheritance refers to the idea that chemical composition is inherited from the molecular cloud from which the protoplanetary disk is formed and should therefore be similar to the composition of the molecular cloud (see, e.g., L. I. Cleves et al. 2014; K. Furuya et al. 2017; K. Altwegg et al. 2019; S. S. Jensen et al. 2021). However, the inherited chemical composition may also be influenced by the properties of the formation environment. For instance, external factors such as ultraviolet (UV) radiation and cosmic-ray flux could alter the composition of the envelope. (e.g., J. K. Jørgensen et al. 2006; J. E. Lindberg & J. K. Jørgensen 2012; J. E. Lindberg et al. 2014, 2017; B. A. L. Gaches & S. S. R. Offner 2018; W. R. M. Rocha & S. Pilling 2018; S. S. Jensen et al. 2021). In contrast, the reset hypothesis posits that the interstellar chemical composition is destroyed/reset and that the chemical inventory depends only on the chemical reactions taking place in the protoplanetary disks during the later stages of planet formation. The idea of a reset is supported by meteoritic evidence (see for details M. J. Mumma & S. B. Charnley 2011; K. M. Pontoppidan et al. 2014). Moreover, several studies in the literature suggest contributions of both inheritance and reset (D. Brownlee et al. 2006; R. Visser et al. 2009; F. J. Ciesla 2010; M. Rubin et al. 2018; M. N. Drozdovskaya et al. 2019), but the relative importance of inheritance versus reset is poorly constrained (K. I. Öberg & E. A. Bergin 2021). The problem, however, might not be dichotomous as intermediate stages of star formation, particularly the protostellar phase, experience energetic events that can trigger chemical reactions and change the chemical makeup of the natal envelope.

The current epoch of star and planet formation within our Galaxy serves as a natural laboratory to understand the origin and formation mechanisms of our solar system and other planetary systems more generally. Most importantly, the earliest phases of star formation can provide insights into the chemical evolution of the seed material for the stars and planets. The protostellar phase, often referred to as the primary accretion phase (e.g., P. Andre et al. 1993, 2000; N. J. I. Evans et al. 2009; S. S. R. Offner & C. F. McKee 2011; W. J. Fischer et al. 2017; M. Narang et al. 2023; S. Federman et al. 2023; J. J. Tobin & P. D. Sheehan 2024), plays a vital role in accumulating/transporting gas and icy dust from the molecular

cloud onto the star and the protostellar/protoplanetary disk for planet formation (the birthplace of planets; see also C. F. Manara et al. 2018; Ł. Tychoniec et al. 2020). It is during this phase that most of the stellar mass is accreted (W. J. Fischer et al. 2017), and the initial conditions for planet formation are set. Hints of the early onset of planet formation, such as radial substructures and grain growth detected by the Atacama Large Millimeter/submillimeter Array (ALMA) in protostellar disks, have also been observed during the protostellar phase (see, e.g., ALMA Partnership et al. 2015; D. Harsono et al. 2018; P. D. Sheehan & J. A. Eisner 2018; D. M. Segura-Cox et al. 2020; P. D. Sheehan et al. 2020; N. Ohashi et al. 2023; J. Drażkowska et al. 2023, and references therein). Simulations have also shown that protostellar disks are more plausible sites for planetesimal formation than protoplanetary (Class II) disks because of their large gas and dust reservoirs (Y. Tsukamoto et al. 2017; Y. A. Tanaka & Y. Tsukamoto 2019). Further, A. J. Cridland et al. (2022) demonstrated with their simulations that during protostellar envelope collapse, grain growth and hydrodynamic drag can enhance the dust-to-gas ratio in the disk, which can start the planetesimal formation due to streaming instabilities (see also S. Koga et al. 2022; J. Drażkowska et al. 2023; A. Miotello et al. 2023).

While the protostellar phase plays a pivotal role in accretion, its impact on altering the chemical composition of dust and ice remains underexplored. Energetic processes occurring during this phase, such as the enhanced protostellar luminosity during the accretion outbursts, UV emission from accretion and outflow shocks, shocks driven by jets/outflows, and X-rays from coronal activity, have the potential to significantly modify the physical structure and chemical composition of ice species within the natal envelope (see, e.g., F. A. van Broekhuizen et al. 2005; H. G. Arce et al. 2008; H. J. Kim et al. 2012; R. Visser et al. 2012, 2015; M. N. Drozdovskaya et al. 2014, 2015; A. O. Benz et al. 2016; A. Karska et al. 2018; A. Mirocha et al. 2021; J. Kim et al. 2022). Further, frequent accretion bursts can trigger and accelerate the formation reactions of complex organic molecules (COMs) in the envelope (V. Taquet et al. 2016). COMs have also been detected in the jet/outflow shocks from protostars (e.g., R. Bachiller & M. Pérez Gutiérrez 1997; H. G. Arce et al. 2008; L. E. Kristensen et al. 2010; E. Mendoza et al. 2014; C. Codella et al. 2015; B. Lefloch et al. 2017).

All the processes discussed above heat the protostellar envelope and thermally process the ices, leading to the sublimation and crystallization of ice species and, more importantly, triggering chemical reactions. Thermal processing of ices has been studied extensively both observationally and in the laboratory (e.g., S. A. Sandford & L. J. Allamandola 1988; A. G. G. M. Tielens et al. 1991; P. A. Gerakines et al. 1995; P. Ehrenfreund et al. 1997, 1999; J. E. Chiar et al. 1998; A. C. A. Boogert et al. 2000, 2008; K. M. Pontoppidan et al. 2008; K. I. Öberg et al. 2009; also see review article by A. C. A. Boogert et al. (2015) and references therein). Further, thermal processing can alter and form new chemical species in the envelope by opening new pathways for chemical reactions (see, e.g., R. T. Garrod & E. Herbst 2006; C. J. Bennett et al. 2007; R. T. Garrod et al. 2008; N. Sakai & S. Yamamoto 2013;

K. Altwegg et al. 2020). Therefore, it is important to understand the impact of these energetic processes in the protostellar phase in shaping the system’s chemistry. For a detailed discussion on chemistry during the protostellar phase, the reader is referred to J. K. Jørgensen et al. (2020), K. I. Öberg & E. A. Bergin (2021), C. Ceccarelli et al. (2023), and references therein.

The launch of the James Webb Space Telescope (JWST) has opened up the infrared spectral window rich in the spectral features of key volatile carriers, such as water (H_2O), carbon dioxide (CO_2), and carbon monoxide (CO), with a superior sensitivity and angular resolution compared to the Infrared Space Observatory and Spitzer. In particular, the Integral Field Units of Near InfraRed Spectrograph (NIRSpec; T. Böker et al. 2022; P. Jakobsen et al. 2022) and Mid-Infrared Instrument (MIRI; G. H. Rieke et al. 2015; G. S. Wright et al. 2015) offer a unique opportunity not only to investigate the chemical composition of a protostellar envelope but also to observe the spatial variations in its distribution. This technological advancement has paved the way for a spatial investigation of protostellar processes that could shape/change the chemistry of volatiles within the envelope.

Investigating Protostellar Accretion (IPA) across the mass spectrum is a JWST Cycle 1 medium General Observers program (PID: 1802; P.I.: S. Thomas Megeath; T. Megeath et al. 2021; S. A. Federman et al. 2024; M. Narang et al. 2024; D. A. Neufeld et al. 2024; A. E. Rubinstein et al. 2024; D. M. Watson et al. 2025, in preparation) to observe five protostars in a broad luminosity range ($L_{\text{bol}} \sim 0.2\text{--}10^4 L_{\odot}$). IPA uses NIRSpec/IFU and MIRI/MRS in 2×2 mosaicing mode to obtain spectral cubes with a field of view of $6'' \times 6''$. The program aims to study protostellar accretion, jets/outflows (S. A. Federman et al. 2024; M. Narang et al. 2024), and to characterize the chemical properties of the protostellar environment (N. G. C. Brunken et al. 2024; P. Nazari et al. 2024; K. Slavicinska et al. 2024). Other programs with JWST have also focused on the icy chemistry (e.g., Y.-L. Yang et al. 2022; M. K. McClure et al. 2023; W. R. M. Rocha et al. 2024) and mass accretion/ejection processes of the protostars (e.g., H. Beuther et al. 2023; D. Harsono et al. 2023; T. P. Ray et al. 2023; K. D. Assani et al. 2024; A. Caratti o Garatti et al. 2024; L. Francis et al. 2024; J. D. Green et al. 2024; B. Nisini et al. 2024; Ł. Tychoniec et al. 2024).

In this study, we investigate the spatial distribution of ice species and evidence for thermal processing within the envelope of the intermediate-mass protostar HOPS 370 (aka OMC2-FIR3), which is the second most luminous source in the IPA sample. It is a Class 0/I intermediate-luminosity protostar with a bolometric luminosity (L_{bol}) of $315.7 L_{\odot}$ (total luminosity $\sim 511 L_{\odot}$) and a bolometric temperature (T_{bol}) of 71.5 K (E. Furlan et al. 2016; J. J. Tobin et al. 2020a). The central protostar is deeply embedded in the envelope. It is located at a distance of 390 pc in the northern part of the Integral Shaped Filament in the OMC-2 region with close to an edge-on geometry (inclination angle $\sim 71^\circ$; S. Federman et al. 2023; J. J. Tobin et al. 2020a, 2020b). Multiwavelength observations have revealed the presence of a strong bipolar jet/outflow (see, e.g., S. Takahashi et al. 2008; B. González-García et al. 2016; M. Osorio et al. 2017; S. A. Federman et al. 2024). Our JWST observations have revealed the presence of a shocked knot in the northern part of the protostellar jet (see S. A. Federman et al. 2024; D. A. Neufeld et al. 2024).

In this paper, we present the chemical inventory of gas and ice species and map the spatial distribution of volatile ice species present in the envelope of HOPS 370. We also search for evidence of the thermal processing of ice species in the protostellar envelope. In Section 2, we discuss the observation and data reduction. In Section 3, we present the observational results for HOPS 370. In Section 4, we delve into an in-depth discussion of the findings and their implications and summarize our findings in Section 5.

2. Observations and Data Reduction

The NIRSpec IFU observations of HOPS 370 were carried out on 2022 October 16, using the F290LP/G395M filter/grating ($R \sim 700\text{--}1300$). The spectral coverage for these observations is $\sim 2.87\text{--}5.01 \mu\text{m}$. Employing a 2×2 mosaicing mode with a 10% overlap and a four-point-dither pattern, we achieved a $6'' \times 6''$ field of view (FOV). Given HOPS 370’s brightness, we employed the NRSIRS2RAPID readout pattern, incorporating 15 groups per integration and one integration. The effective exposure time for these observations amounted to 3501.3 s.

The raw data obtained from the Mikulski Archive for Space Telescopes (MAST) was reduced using JWST pipeline version 1.9.5.dev7+gbf7d3c9b and CRDS context files “jwst_1069.pmap.” To achieve science-ready products we applied custom bad pixel masks and astrometry corrections. For further details regarding the data reduction and astrometry refinement processes, the reader is referred to S. A. Federman et al. (2024).

HOPS 370 was observed on 2023 March 15, with MIRI/MRS, in all four channels covering a spectral range from ~ 4.9 to $28 \mu\text{m}$ at spectral resolving power ranging from $R \sim 4000\text{--}1500$ (O. C. Jones et al. 2023). Mosaicing settings for these observations were similar to those of NIRSpec with the four-point-dither pattern. In MIRI observations, we have incorporated 50 groups per integration and two integrations with the FASTR1 readout pattern. The effective exposure time for these observations was 4484.5 s. Dedicated observation of background with similar exposure parameters was also taken for background subtraction. The data reduction process for MIRI/MRS is discussed in D. A. Neufeld et al. (2024).

Our FOVs (NIRSpec and MIRI) of HOPS 370 also include another source, [MGM2012] 2301 (S. T. Megeath et al. 2012), toward the south of HOPS 370.

3. Results

The spectral images obtained using NIRSpec and MIRI reveal the gas and dust structure of the inner 1200 au surrounding the protostar and its chemical richness (S. A. Federman et al. 2024; D. A. Neufeld et al. 2024; A. E. Rubinstein et al. 2024). We present the chemical inventory and spatial ice distribution in this section.

3.1. Morphology

Figure 1 shows the continuum and line maps of HOPS 370 that reveal the morphology of the inner protostellar environment. The star symbol in all the panels of Figure 1 marks the ALMA continuum disk position ($05^{\text{h}}35^{\text{m}}27^{\text{s}}.63\text{--}05^{\circ}09'34''.42$ (ICRS); J. J. Tobin et al. 2020b). The highly inclined orientation (approximately 72° inclination; J. J. Tobin et al. 2020a; S. A. Federman et al. 2024) reveals the presence of hourglass-shaped scattered-light cavities with the northern

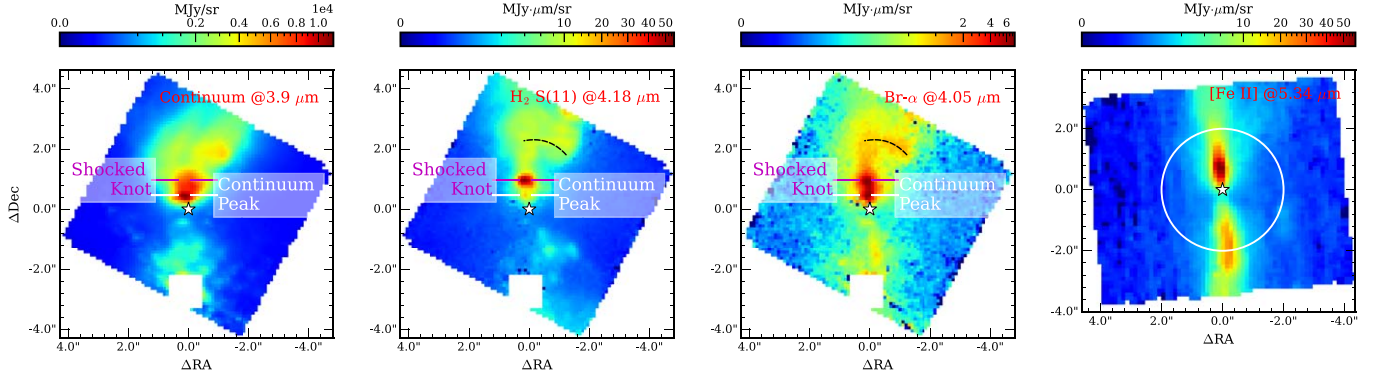


Figure 1. Morphology of the HOPS 370 protostellar environment seen in the continuum map at $3.9 \mu\text{m}$, H_2 S(11) map at $4.18 \mu\text{m}$, Br- α map at $4.05 \mu\text{m}$, and [Fe II] map at $5.34 \mu\text{m}$. All the line maps are continuum subtracted. The white star marks the ALMA $870 \mu\text{m}$ continuum position in all the panels. The white circular aperture in the right panel is used for 1D spectrum extraction. The locations of the continuum peak at $3.9 \mu\text{m}$ and the shocked knot are shown using the horizontal white and magenta color lines, respectively. The dashed black line in the middle panels marks the fishhook-like structure in Br- α and H_2 . The bright foreground pre-main-sequence star, [MGM2012] 2301, is masked to enhance the contrast.

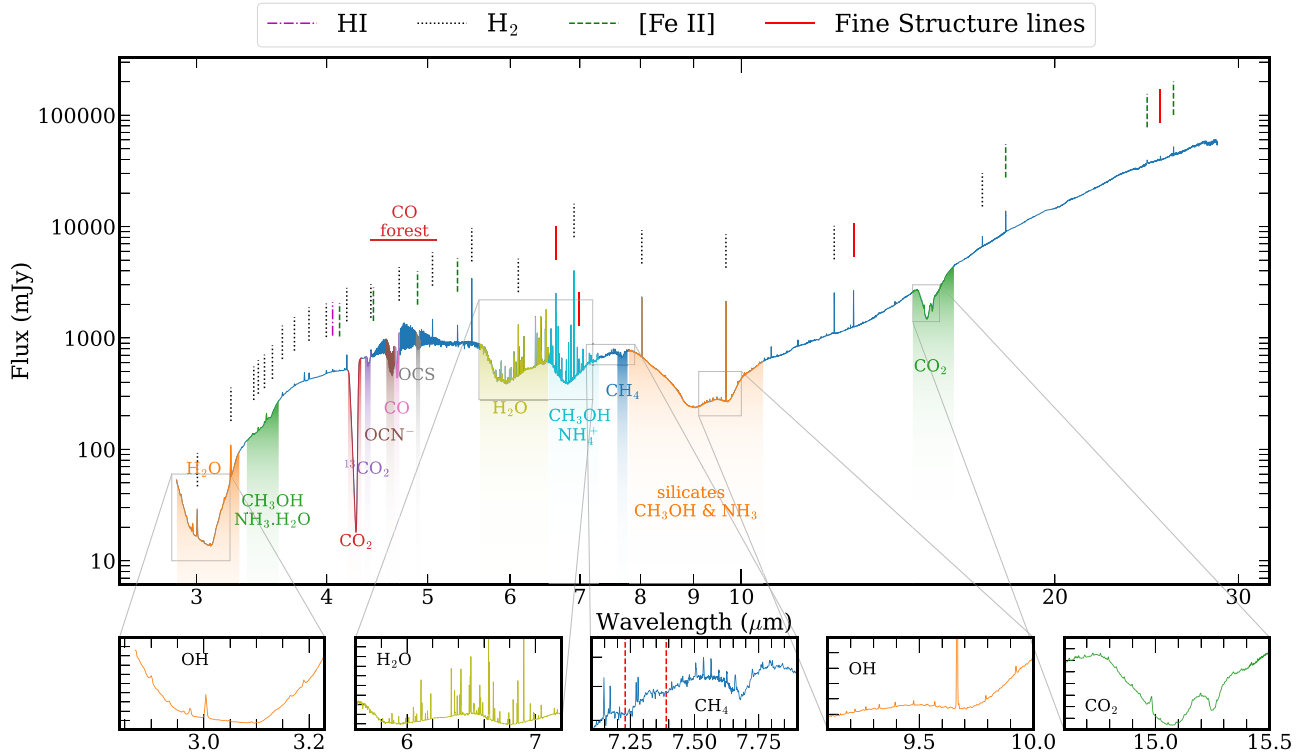


Figure 2. Extracted 1D spectrum using a $2''$ aperture centered at the ALMA continuum position of HOPS 370. Prominent absorption features are gradient-shaded in colors while the detected emission lines are marked using vertical line markers. Molecular lines of H_2O , CH_4 , OH, and CO_2 are displayed in the insets.

cavity brighter relative to the southern cavity (Figure 1, left panel). Asymmetric continuum emission has also been seen in other protostars (e.g., in TMC 1A, D. Harsono et al. 2023; K. D. Assani et al. 2024). The northern cavity shows an east-west asymmetry with brighter emission to the west. Bright structured emission in H_2 is surrounded by a weaker hourglass-shaped H_2 emission (middle-left panel). In the northern cavity, a shocked knot is detected in both OH and H_2 (see also S. A. Federman et al. 2024; D. A. Neufeld et al. 2024). The shocked knot is $\sim 0.5''$ north of the scattered continuum peak (see Figure 1). The Br-Alpha line peaks at the shock knot and the continuum source (middle-right panel). There is a fishhook-like structure in Br- α and H_2 emission in the northern

cavity ($\sim 2''$ away from the ALMA continuum position). A highly collimated jet is traced in [Fe II] at $5.34 \mu\text{m}$, this also peaks at the position of the shock knot (right panel; see also S. A. Federman et al. 2024; D. A. Neufeld et al. 2024). This collimated jet is also seen in Br-alpha and, in the northern cavity, H_2 (S. A. Federman et al. 2024).

3.2. The Chemical Inventory

In Figure 2, we present the combined NIRSpect and MIRI spectrum extracted from a $2''$ radius aperture centered on the ALMA continuum peak (see Figure 1). This spectrum covers the wavelength range from 2.87 to $28 \mu\text{m}$ and reveals a rich chemical inventory in both the gas and ice phases.

3.2.1. Gas Species

HOPS 370 shows a diverse set of gas-phase molecular emission lines. We detect many rovibrational lines from the P and R branches of CO that exhibit both extended and jet-like structures (S. A. Federman et al. 2024; A. E. Rubinstein et al. 2024). Pure rotational lines of H_2 are marked in dotted black vertical lines in Figure 2. We have detected pure rotational H_2 lines from S(1) to S(19). The H_2 lines appear to have originated both in the jet and its warm gas in the cavity, with emission peaking at the shock knot (Figure 1; see also S. A. Federman et al. 2024; D. A. Neufeld et al. 2024; H. Tyagi et al. 2025, in preparation). Other sources observed with JWST also show similar morphology in H_2 and appear to originate in the winds (e.g., N. Arulanantham et al. 2024; I. Pascucci et al. 2024; J. A. Sturm et al. 2024). We have also detected H_2O , OH, CH_4 , and CO_2 line forests, zoomed in the inset in Figure 2 (details about these lines will be discussed in P. Manoj et al. (2025, in preparation), D. A. Neufeld et al. (2025, in preparation), and A. E. Rubinstein et al. (2025, in preparation); see also L. Francis et al. 2024; D. A. Neufeld et al. 2024). The HCN rovibrational lines from the P, Q, and R branches between the wavelength range of 13.7–14.3 μm are also detected. Further, we detect HI, [Fe II], and other fine-structure emission lines (see Figure 2). Detailed line lists of detected molecular and atomic/ionic species are given in Tables A1 and A2, respectively. In Table A1, we do not present lines of CO, nor the lines at NIRSpec wavelengths as they have been already presented by A. E. Rubinstein et al. (2024) and S. A. Federman et al. (2024).

3.2.2. Ice Species

Many ice absorption features, highlighted in Figure 2, are detected at high signal-to-noise ratios (SNRs). Dominant ice absorption features in the HOPS 370 spectrum include the OH stretching mode of H_2O at 3.0 μm , CO_2 stretching mode at 4.27 μm , $^{13}\text{CO}_2$ stretching mode at 4.38 μm (N. G. C. Brunken et al. 2024), OCN^- at 4.62 μm , CO stretching mode at 4.67 μm , C–O stretch mode of OCS at 4.90 μm , H_2O bending mode at 6.0 μm , $\text{CH}_3\text{OH}/\text{NH}_4^+$ at 6.85 μm , CH_4 at 7.67 μm , and double-peaked CO_2 bending mode at 15.2 μm . We detected the 9.7 μm silicate absorption feature blended with the NH_3 umbrella and CH_3OH absorption features. We also detected two weak absorption features at 7.23 and 7.38 μm probably due to C–H deformation bands of $\text{C}_2\text{H}_5\text{OH}/\text{HCOOH}$, and $\text{CH}_3\text{CHO}/\text{HCOO}^-$ respectively (marked by red dashed lines in the third inset of the bottom row in Figure 2; also see A. C. A. Boogert et al. 2015; W. R. M. Rocha et al. 2024 for detailed discussion on these species). Absorption bands of $\text{NH}_3:\text{H}_2\text{O}$, CH_3OH at ~ 3.5 μm , and HDO ice at ~ 4.1 μm have also been detected (K. Slavicinska et al. 2024). Additionally, we have a tentative detection of CH_3CN and $\text{C}_2\text{H}_5\text{CN}$ (P. Nazari et al. 2024).

The presence of a narrow absorption feature at 3.1 μm due to crystalline H_2O and a double-peaked 15.2 μm CO_2 absorption feature suggests hints of thermal processing (A. C. A. Boogert et al. 2015). Furthermore, a double-peaked $^{13}\text{CO}_2$ absorption feature, attributed to ice segregation due to thermal processing, was observed by N. G. C. Brunken et al. (2024). However, the origin of thermal processing remains unclear based on the spectrum at a single position. A comprehensive list of all the detected ice species is given in Table B1.

3.3. Spatial Variation of Ice Spectra

The 1D NIRSpec+MIRI spectrum, centered on the protostellar position, in Figure 2 shows signs of thermal processing, prompting further investigation into its localization and physical origin. IFU observations offer a comprehensive data set for this investigation, particularly in the NIRSpec observation, where extended scattered light provides an extended background continuum to be absorbed by the ice species present in the protostellar envelope (see Figure 1, first panel from left).

To investigate the spatial variations in absorption features of ice species, we extract spectra at various positions in the protostellar envelope using circular apertures with a radius of 0.6. Our extraction apertures are displayed in Figure 3 overlaid on the 3.9 μm continuum image (referenced by numbers and colors in the top-right panel). Extracted spectra, plotted in the subpanels, share the same color as its extraction aperture in the right top panel. The overall spectra look qualitatively similar in all the apertures with some minor changes in absorption features, indicating that the dominant absorption ice species are common throughout the envelope.

The heating of CO_2 ice mixtures in the envelope can sublimate the most volatile ice species (distillation) and rearrange molecular bonds (segregation), which leads to the formation of pure CO_2 ice (double-peaked absorption in the 4.38 and 15.2 μm feature indicates segregation/distillation of CO_2 ; see also, P. Ehrenfreund et al. 1997; A. C. A. Boogert et al. 2000; K. M. Pontoppidan et al. 2008). In our observations, $^{13}\text{CO}_2$ (displayed in the insets in Figure 3) also shows double-peaked absorption in all the apertures, suggesting CO_2 ice distillation/segregation due to thermal processing throughout the envelope. Apertures 0, 1, 2, and 3 show the absorption peak at 3.1 μm due to crystalline H_2O . Further, we detect a relatively weaker 4.67 μm CO absorption with respect to the 4.27 μm CO_2 absorption feature in all the apertures. All the signatures of ice absorption features suggest that the protostellar envelope has undergone thermal processing.

3.4. Optical Depth Maps

To investigate the spatial distribution of the ice species, we generated peak optical depth maps from the NIRSpec observations (Figures 4 and 5). We first subtracted all the emission lines from the data cube and generated a line-free spectral cube (for details, see A. E. Rubinstein et al. 2024), which is then used to generate optical depth maps.

Next, we fitted a local continuum using low-order (second order for $^{13}\text{CO}_2$, CO, and OCN^- ; fourth order for CO_2) polynomials to the ice absorption features (except H_2O) at each spaxel. Local continua are used because scattering can produce large wing emissions, e.g., the blue wing of 4.27 μm CO_2 (E. Dartois et al. 2022, 2024), which makes fitting a global continuum challenging. Determining the continuum for H_2O at 3.0 μm poses additional challenges due to the incomplete spectral coverage shortward of 2.87 μm of our observations. To address this, a straight-line fit is applied between anchor points at 4–4.18 μm and ~ 2.87 μm . While fitting a straight-line continuum using an anchor point at 2.87 μm for the H_2O may underestimate the true optical depth, it facilitates comparison of relative strengths across the IFU. We calculated the optical

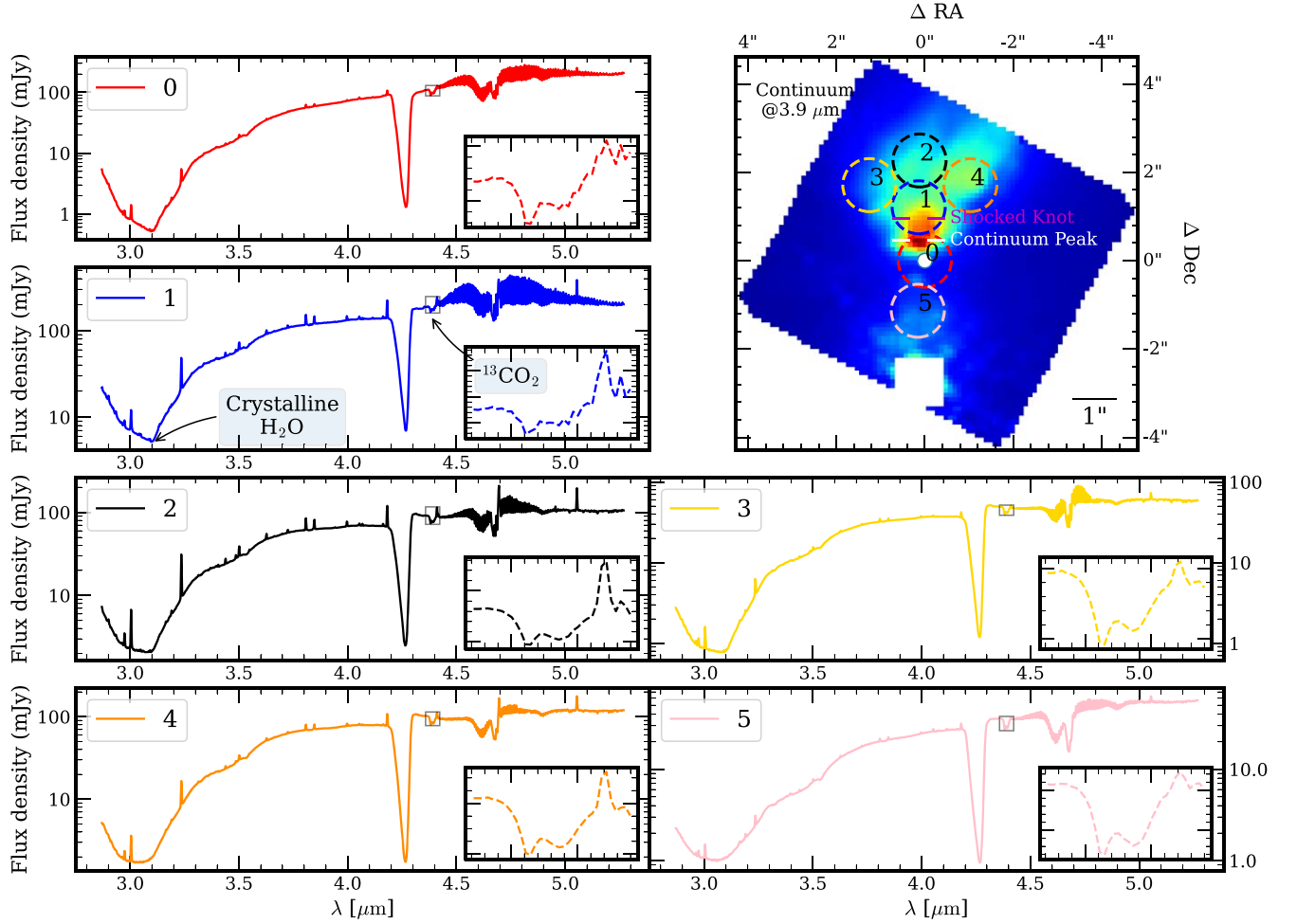


Figure 3. Spatial variation of the NIRSpect spectrum extracted from the circular apertures of $0''.6$ radius that are shown overlaid on the $3.9 \mu\text{m}$ continuum image. The numbers and colors corresponding to apertures are used to point to the extracted spectrum. Other color schemes are similar to Figure 1. Insets in the spectra panels display the $^{13}\text{CO}_2$ ice feature. The [MGM2012] 2301 is masked in the continuum image to enhance the contrast.

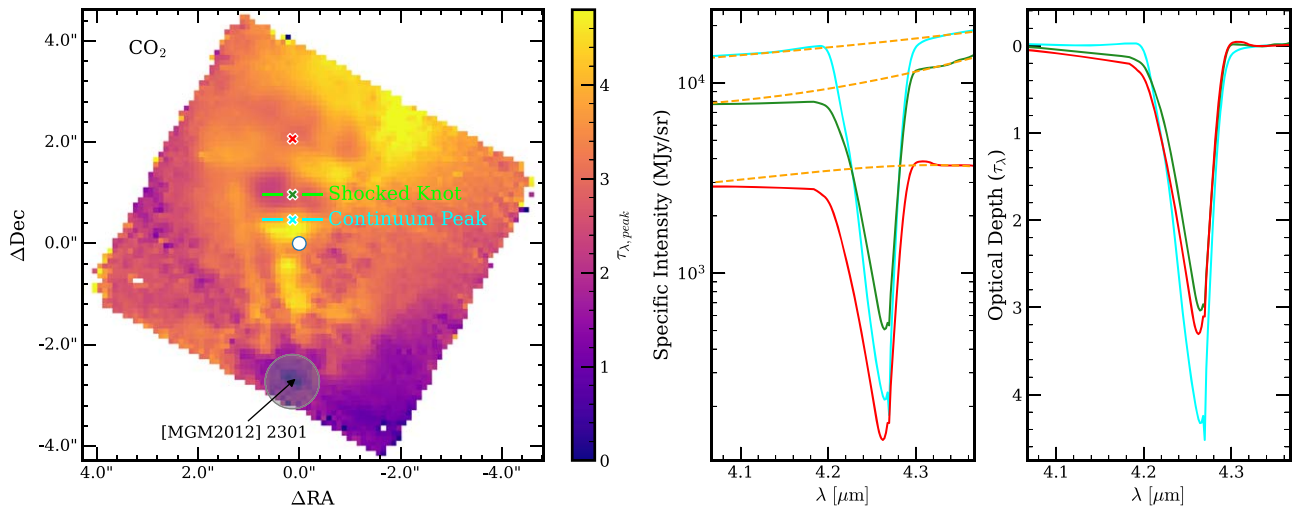


Figure 4. Peak optical depth map of CO_2 and illustration of optical depth calculation for each spaxel. Left: peak optical depth map of CO_2 at $4.27 \mu\text{m}$, with the white circle indicating the ALMA $870 \mu\text{m}$ continuum peak. The locations of the $3.9 \mu\text{m}$ continuum peak and the shocked knot are shown using the horizontal aqua and lime color lines, respectively. Cyan, green, and red crosses mark spaxel positions corresponding to the fit results displayed in the middle and right panels. The gray-shaded region in the south marks the location of the bright foreground source [MGM2012] 2301, causing PSF artifacts in the map. Middle: specific intensity of marked spaxel (“x”) as a function of the wavelength. The colors of the plotted lines correspond to the same color “x” spaxel. The dashed orange lines represent the fitted continuum to the respective spaxel. Right: calculated optical depths for the marked spaxels.

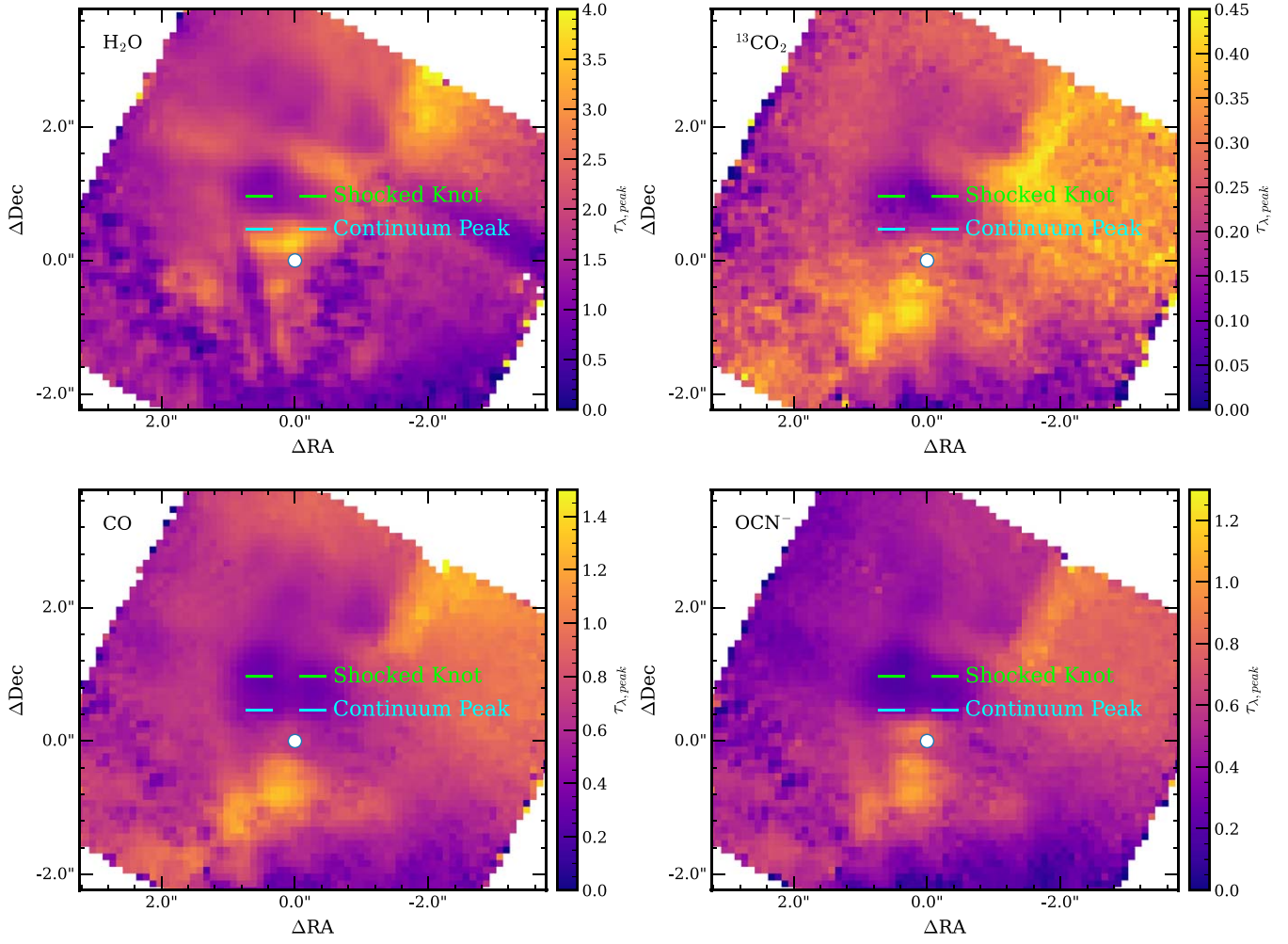


Figure 5. Peak optical depth maps (zoomed in) of H_2O at $3.0\ \mu\text{m}$, $^{13}\text{CO}_2$ at $4.38\ \mu\text{m}$, CO at $4.67\ \mu\text{m}$, and OCN^- at $4.61\ \mu\text{m}$. The white circle marks the ALMA continuum position. The southern region of all the maps is affected by the PSF artifacts of [MGM2012] 2301, marked in Figure 4.

depth at each spaxel using the following equation,

$$\tau_\lambda = -\ln \left[\frac{I_{\lambda,\text{obs}}}{I_{\lambda,\text{cont}}} \right], \quad (1)$$

where $I_{\lambda,\text{obs}}$ is the observed specific intensity value and $I_{\lambda,\text{cont}}$ is the continuum estimated from the fit.

Figure 4 illustrates the results of optical depth calculations for CO_2 at $4.27\ \mu\text{m}$. The left panel shows the peak optical depth map. The middle and right panels showcase spectral profiles with fitted continuum and the corresponding optical depth, respectively, for spaxels marked with crosses on the optical depth map. The peak optical depth maps of other species (H_2O at $3.0\ \mu\text{m}$, $^{13}\text{CO}_2$ at $4.38\ \mu\text{m}$, CO at $4.67\ \mu\text{m}$, and OCN^- at $4.61\ \mu\text{m}$) are shown in Figure 5.

Peak optical depth maps of all the species show more or less similar morphology. Interestingly, a low optical depth region around the shocked knot is observed (roughly $0.8''$, ($\sim 312\ \text{au}$) north of the ALMA continuum position). This elliptical-shaped opacity deficit region/hole is visible in maps of all the ice species. There is another extended low optical depth region $\sim 2''$ ($\sim 780\ \text{au}$) northwest of the ALMA continuum position, separated by a band of extinction. A relatively high optical depth region is located at the scattered-light continuum peak, indicating a change in scattered-light emission structure at the

peak ice absorption wavelength. Additionally, we find another high optical depth region, northwest of the ALMA continuum position off the edge of the outflow cavity (near $\Delta\text{R.A.}, \Delta\text{decl.} = -2.0'', 2.0''$).

In the southern part of the FOV, spatial information is obscured by the point-spread function (PSF) artifacts due to [MGM2012] 2301, a bright pre-main-sequence star saturated at NIRSpec wavelengths (see also S. A. Federman et al. 2024). However, the $^{13}\text{CO}_2$ and CO maps (top-right and bottom-left panels of Figure 5, respectively), which are relatively less affected by artifacts, suggest an increase in the optical depth, roughly following an hourglass geometry in the southern envelope similar to scattered light.

In summary, the optical depth maps of the ice species examined show similar spatial structures in the northern cavity (see Figures 4 and 5). The presence of the low optical depth region/hole in all the ice species suggests a common origin for these structures.

3.5. Low Extinction around Shocked Knot

Motivated by the low optical depth regions highlighted in Figures 4 and 5, we further investigated the scattered-light emission across the FOV at the peak CO_2 ice absorption wavelength ($4.27\ \mu\text{m}$), as illustrated in Figure 6. We use mean channel maps to compare the observed intensity structures at

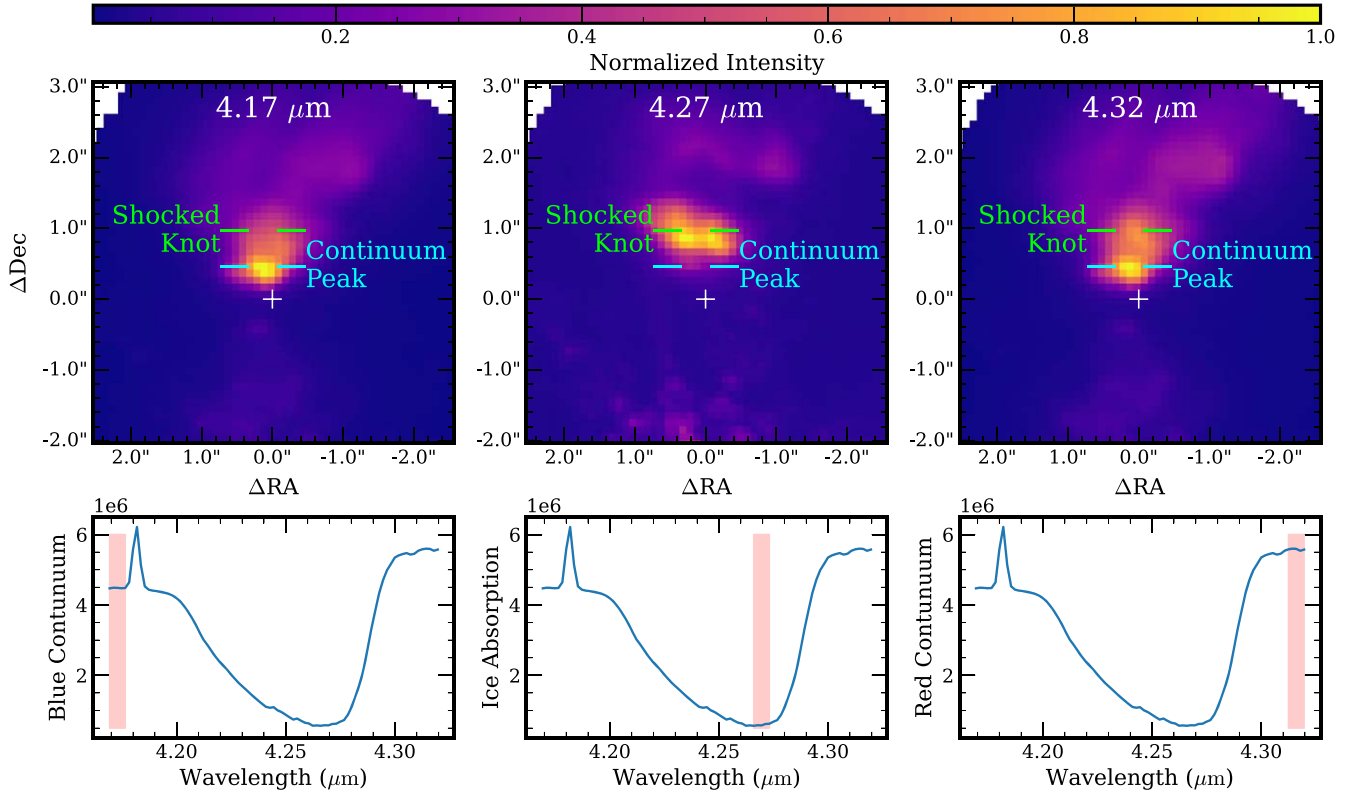


Figure 6. Top: normalized mean channel maps from NIRSpect observations showing the differences in observed intensities between the continuum and CO₂ ice absorption wavelengths. The mean wavelength corresponding to each map is indicated at the top of each panel. Bottom: The red-shaded regions in the spectra below each map highlight the spectral range used to create the mean channel maps. The white “+” symbol denotes the position of the ALMA continuum. Note: The source [MGM2012] 2301 is masked in this figure.

the ice absorption-free continuum and peak ice absorption wavelengths.

In Figure 6, the left and right column panels display the mean channel maps at ice absorption-free wavelengths on the blue and red sides of the 4.27 μm CO₂, respectively. The middle panel in the top row displays the observed mean channel map at the peak ice absorption wavelengths. Moreover, red-shaded regions overlaid on the extracted spectrum below each map highlight the spectral regions corresponding to the displayed mean channel maps.

The scattered-light continuum in the left and right panels in Figure 6 show similar hourglass geometry with the northern part being brighter. The emission peaks in both panels are marked by white horizontal bars. However, the emission structure in the middle panel differs significantly from the left and right panels, showing a shifted emission peak and disappearance of the hourglass geometry.

A similar change in the observed intensity structures at the peak ice absorption wavelengths compared to the scattered continuum is seen for all the ice species (also evident from the optical depth maps in Figure 5), suggesting that all the scattered light in FOV except at the shocked knot position is getting significantly absorbed at peak ice absorption wavelengths, and only the light at the shocked knot position escapes the protostellar envelope at those wavelengths. The continuum escaping region at peak absorption wavelengths almost perfectly matches the low-opacity region discussed in Section 3.4 which further confirms the presence of an opacity hole in the envelope.

3.6. Mapping the Excess Crystallization of Water Ice

To investigate the region of excess heating as indicated and discussed in Section 3.3, we map the ratio between crystalline and amorphous H₂O ice within the FOV. This was done by following the procedure discussed in Section 3.4. Initially, optical depth maps were generated at 3.0 μm (the peak absorption wavelength for amorphous H₂O) and 3.1 μm (the peak absorption wavelength for crystalline H₂O). The ratio of the optical depth map at 3.1 μm to that at 3.0 μm then, is a measure of excess absorption due to crystalline H₂O. Figure 7 shows the ratio map, clearly indicating that the ratio is higher toward the shocked knot, with the highest water crystallization coincident with the shocked knot. The southern part of the ratio map is noisy due to the PSF artifacts of [MGM2012] 2301.

3.7. Spatial Decomposition Maps of ¹³CO₂, OCN[−], and CO

To further investigate the envelope environment, we show the decomposed maps of ¹³CO₂, OCN[−], and CO in Figure 8. The decomposed maps are generated by fitting the observed absorption features using laboratory spectra of contributing species in the respective ice absorption bands. These maps cover the inner 4'' $\times$ 4'' region.

To generate the decomposed maps, the 4'' $\times$ 4'' region was divided into a 2D grid with a cell size of 3 pixel $\times$ 3 pixels ($\sim 0.3 \times 0.3$). This binning approach, employed instead of utilizing individual pixels, enhances the SNR for subsequent decomposition analysis. Spectra were extracted from each 0.3×0.3 square cell. Subsequently, we fit a low-order (third

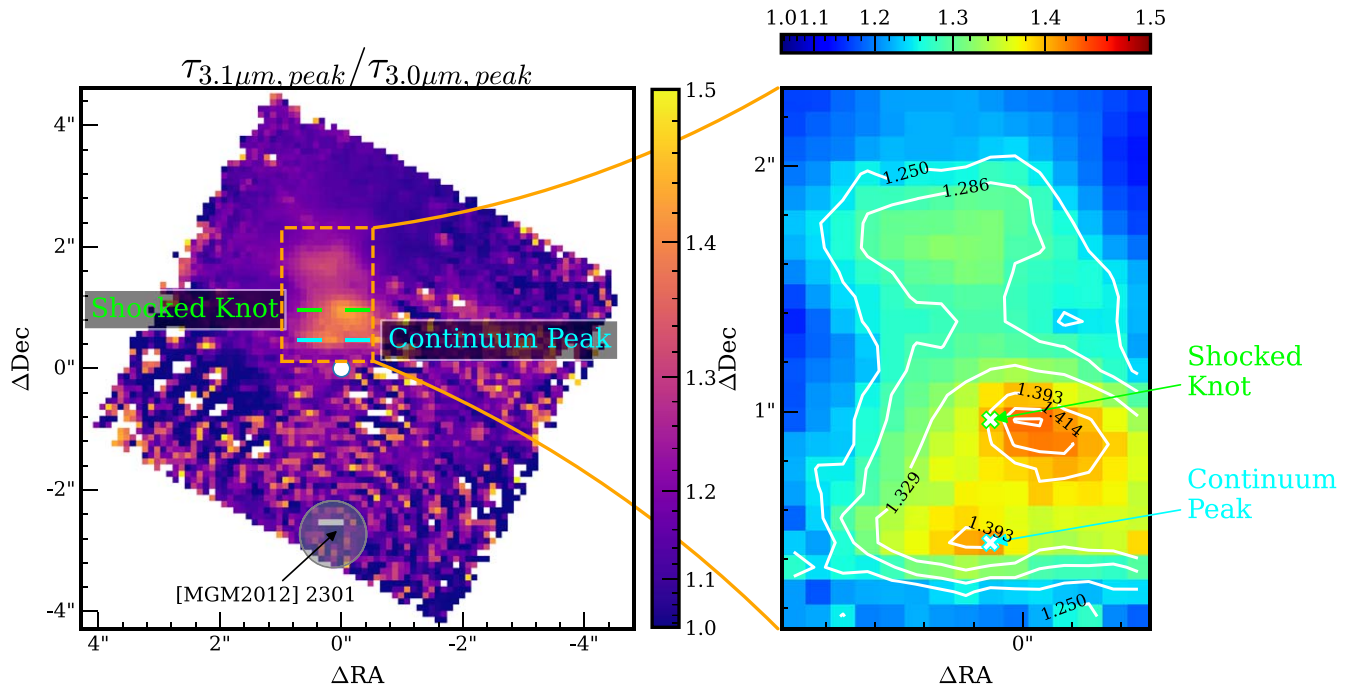


Figure 7. Left: ratio of peak optical depth map of H₂O at 3.1 μ m (crystalline H₂O) to peak optical depth map of H₂O at 3.0 μ m (amorphous H₂O). The white circle marks the ALMA continuum position. Right: zoomed-in ratio map, marked by an orange dashed rectangle in the left panel, in color scale. White contours are the ratio contours. The lime and aqua crosses mark the location of the shocked knot and 3.9 μ m continuum peak.

order for ¹³CO₂; second order for CO and OCN[−]) polynomial local continuum in the extracted spectra to obtain the optical depths for each cell. These derived optical depth values were then used in the decomposition of the absorption features.

For the decomposition of ¹³CO₂, we followed the methodology outlined by N. G. C. Brunken et al. (2024), where they demonstrated the necessity of including warm and hot ice species in the fitting process for HOPS 370. In particular, we utilized laboratory spectra of pure CO₂ at 80 K, CO₂:H₂O (1:10) at 160 K, and CO₂:CH₃OH (3:1) at 105 K, obtained from the Leiden Ice Database for Astrochemistry³¹ (W. R. M. Rocha et al. 2022), to model the observed ¹³CO₂ absorption feature at 4.38 μ m (Figures 8(a) and (b)).

Due to the feature overlap, the decomposition of OCN[−] and CO ice was done simultaneously (Figures 8(c) and (d)). We used a mixture of Gaussian and Lorentzian functions to fit the absorption feature (see for details K. M. Pontoppidan et al. 2003; F. A. van Broekhuizen et al. 2005; A. C. A. Boogert et al. 2022, and references therein). Components for both the OCN[−] and CO are described in Table 1. See P. Nazari et al. (2024) for details about OCN[−] ice.

Figures 8(a) and (c) show the peak optical depth maps of various fitted components of ¹³CO₂, OCN[−], and CO. Both figures demonstrate an excess of polar species relative to apolar species across the entire FOV. The dominance of the polar CO component relative to the apolar CO component across FOV and the requirement of hot/warm ice components to fit the ¹³CO₂ absorption profile suggests thermal processing throughout the envelope (see Section 4.1). The maps also show a minimum in the optical depth coincident with the bright knot in the H₂ 0–0 located $\sim 1''$ north of the ALMA continuum position.

4. Discussion

The energetic processes in the protostellar phase, such as protostellar irradiation, episodic accretion-drive outbursts, and shocks driven by jets/outflows, can potentially change the chemical composition or the physical structure of the envelope material. Direct observation of these phenomena in action has been previously limited due to a lack of sensitivity and angular resolution. However, the unique combination of JWST NIRSpec+MIRI’s sensitivity, angular resolution, and broad spectral coverage (2.8–28 μ m) has finally made these detailed studies feasible.

Our latest observations of HOPS 370 using JWST NIRSpec and MIRI reveal the chemical inventory and previously unseen details of the ice distribution within the envelope. In this section, we discuss the results obtained in the Section 3.

4.1. Warm Envelope

As demonstrated in Section 3, multiple ice absorption signatures (e.g., weak CO ice, double-peaked 4.38 μ m CO₂, double-peaked 15.2 μ m CO₂, and the presence of crystalline water) provide evidence for significant envelope heating in HOPS 370 (K. M. Pontoppidan et al. 2008; A. C. A. Boogert et al. 2015; J. Kim et al. 2022; N. G. C. Brunken et al. 2024). Some of these signatures, e.g., double-peaked ¹³CO₂ ice features are quite extended (see Figure 3). Additionally, CO ice absorption is much weaker relative to CO₂ everywhere in the envelope (see Figures 4 and 5). All of these together suggest the thermal processing of ice species throughout the envelope in our FOV.

Additionally, the decomposed maps of CO, and OCN[−] (see Figure 8(c)) show that the polar mixtures of these ice species have higher optical depths across the envelope relative to the apolar mixtures. The sublimation temperatures of polar mixtures are relatively higher than those of apolar mixtures (e.g., the sublimation temperatures of polar CO and apolar CO

³¹ <https://icadb.strw.leidenuniv.nl/>

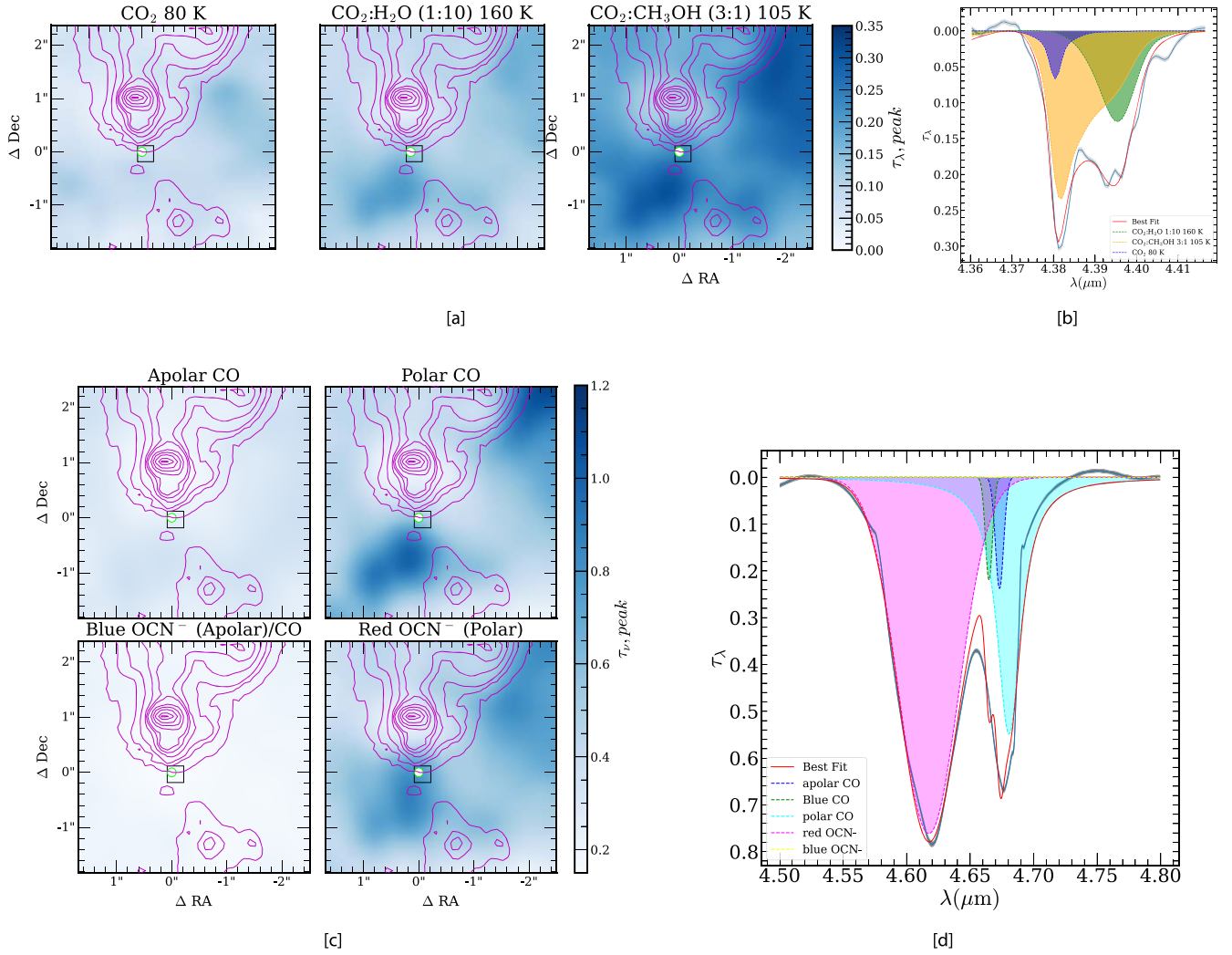


Figure 8. (a) Decomposed maps of $^{13}\text{CO}_2$. The figure shows three components from left to right: CO_2 at 80 K, $\text{CO}_2:\text{H}_2\text{O}$ (1:10) at 160 K, and $\text{CO}_2:\text{CH}_3\text{OH}$ (3:1) at 105 K. All panels use color maps to represent the peak optical depth. A lime circle marks the ALMA continuum position, and the magenta color contour overlaid displays the H_2 0–0 S(11) emission in the FOV. The contour levels are 0.015, 0.03, 0.05, 0.1, 0.15, 0.2, 0.35, 0.5, 0.67, 0.85, $0.97\times$ (Maximum of H_2 0–0 S(11) emission). The densest contours ($1''$ north to the ALMA continuum) mark the shocked knot. (b) Fitting example of the decomposition of $^{13}\text{CO}_2$ ice for the black cell marked in the peak optical depth maps in (a). (c) Decomposed maps of CO and OCN^- ice. The top row displays the apolar (left) and polar (right) components of CO, while the bottom row shows the apolar OCN^-/CO (left) and polar OCN^- (right). All panels use color maps to represent the peak optical depth. The overlaid lime circle and magenta color contours are similar to those on the panel (a). (d) Fitting example of decomposition of CO and OCN^- ice features for the black cell marked in the peak optical depth maps in (c).

are ~ 90 K and ~ 20 K, respectively, see also A. C. A. Boogert et al. 2015). The combination of weak CO ice absorption, dominated by the polar mixture, along with the presence of strong rovibrational CO lines in emission (see Figure 3), strongly suggest the possibility of CO ice sublimation in the inner envelope up to ~ 800 au (the extent of the decomposed maps) from the central protostar.

To explore the potential cause of the thermal processing of ice species over this spatial extent, we calculate the radial profile of the dust temperature in the envelope. For simplicity, we assume the envelope to be spherical and compute the equilibrium temperature of a dust grain at a distance R from the central protostar. Assuming dust grains to be perfect blackbodies, the temperature of the dust grain (T_{dust}) at a distance R from the protostar with luminosity L_* is given by:

$$T_{\text{dust}}(R) = \left[\frac{L_*}{16\pi\sigma R^2} \right]^{1/4}. \quad (2)$$

Figure 9 shows the radial profile of dust temperature. A total luminosity of $511L_{\odot}$ was used for this calculation

(E. Furlan et al. 2016; J. J. Tobin et al. 2020b). The radial profile of dust temperature shows that dust may have temperatures above 20 K and 90 K for radial distances less than ~ 4000 and ~ 200 au, respectively. These temperatures are high enough to explain the extended ice heating observed in HOPS 370's envelope. Hence, the envelope heating to this spatial extent in HOPS 370 can likely be attributed to its high total luminosity, $511L_{\odot}$, due to the large accretion rate of $\dot{M}_{\text{acc}} \sim 2.25 \times 10^{-5} M_{\odot} \text{ yr}^{-1}$ (J. J. Tobin et al. 2020b). However, the exact mechanism responsible for the high accretion rate in HOPS 370, whether it is in a steady state or in a burst phase, is still a matter of debate (J. J. Tobin et al. 2020b). Nevertheless, given the simplistic assumptions made in the calculations, a detailed radiative transfer modeling study is required to understand the details of the thermal structure of the envelope in the future.

4.2. Is There a Local Ice Deficit in the Inner Envelope?

Instead of a smooth decrement in the envelope dust density as a function of the distance from the central protostar, the peak

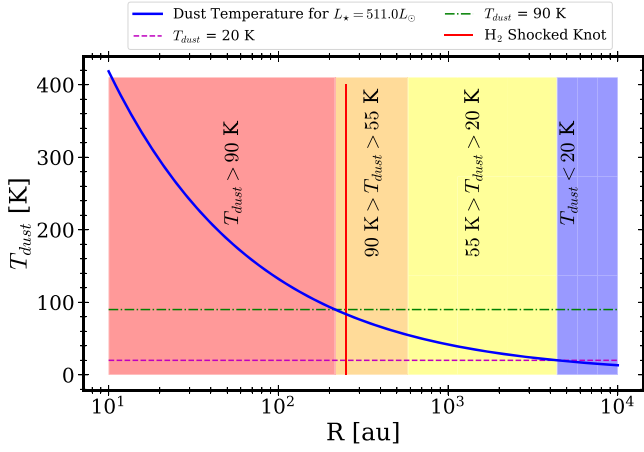


Figure 9. Radial profile of dust temperature of the HOPS 370 envelope. The color-shaded regions mark various dust temperature zones. The red vertical line represents the inclination corrected distance of the H₂ shocked knot from the central protostar.

Table 1
CO and OCN[−] Component Profiles

Component	Profile	Center (μm)	FWHM (μm)
Apolar OCN [−] /CO	Gaussian	4.5969	0.0317
Polar OCN [−]	Gaussian	4.6174	0.0554
Blue CO	Gaussian	4.6648	0.0065
Apolar CO	Gaussian	4.6731	0.0076
Polar CO	Lorentzian	4.6806	0.0232

optical depth maps of volatiles such as H₂O, CO₂, ¹³CO₂, CO, and OCN[−] reveal a highly structured spatial distribution of ice species column density in the envelope (see Figures 4 and 5). Further, channel maps in Figure 6 show a remarkable difference in the observed intensity distribution at peak ice absorption wavelengths compared to ice absorption-free/continuum wavelengths. The normalized channel maps show that, at peak ice absorption wavelengths, not all background scattered light is absorbed by the ices in the envelope. The scattered light leaks through an elliptical patch, spatially overlapping with the H₂ and OH shocked knots. These signatures suggest relatively less absorption of light by ice species at the shocked knot position. There is another low-opacity bar-like region northwest of the shocked knot position (~2'' north of the ALMA continuum position).

To confirm the lack of absorption offered by ice species at these low-opacity regions, we generated extinction (A_V) and column density maps using the CO₂ ice absorption band. To produce a column density map of CO₂, we use the following equation at each spaxel,

$$N_{\text{CO}_2} = \frac{1}{A} \int \tau_\nu d\nu, \quad (3)$$

where $A = 1.1 \times 10^{-16} \text{ cm mol}^{-1}$ is the density-corrected band strength (P. A. Gerakines et al. 1995; M. Bouilloud et al. 2015). Next, we generate the extinction map using the KP5 extinction law (K. M. Pontoppidan et al. 2024).

We used the relation $A_\lambda = 2.5 \log(e) N \kappa_{\text{ext}}(\lambda)$, where $\kappa_{\text{ext}}(\lambda)$ and N represent the mass absorption coefficient and mass column density, respectively (K. M. Pontoppidan et al. 2024).

Then,

$$\frac{A_V}{A_\lambda} = \frac{\kappa_{\text{ext}}(V)}{\kappa_{\text{ext}}(\lambda)}. \quad (4)$$

At each spaxel, we used the optical depth (τ_λ) at 4.27 μm of CO₂ ice absorption feature to calculate A_λ . Using Equation (4), we were able to calculate A_V at each spaxel. We present the extinction and column density maps obtained using the 4.27 μm CO₂ feature in the left and middle panels of Figure 10, respectively. The shocked knot marked by magenta-colored lines coincides with the lower extinction and lower column density regions.

We independently measure the extinction values toward the shocked knot using H₂ 0–0 S(*J*) rotational lines (see Figure C1; for details, see M. Narang et al. 2024; D. A. Neufeld et al. 2024; H. Tyagi et al. 2025, in preparation), which provides a measure of extinction independent of ice optical depth. Four circular apertures of radius 0.4'' labeled as 0, 1, 2, and 3 are depicted in black, magenta, red, and orange, respectively in the middle panel of Figure 10—were used to calculate A_V . We find that aperture 1, which spatially coincides with the low optical depth region and the shocked knot, exhibits a lower A_V value by ~2 mag compared to aperture 0 and 2 (see dot markers in the right panel of Figure 10). Similarly, aperture 3 has a lower extinction than aperture 0, and 2 has a higher extinction than aperture 1. The H₂ rotational diagrams at apertures 0, 1, 2, and 3 are presented in Appendix C.

In the same panel, we have marked the mean A_V values of these apertures calculated using the CO₂ ice with a cross marker for comparison. Both measurements exhibit a similar trend; however, the A_V values obtained from CO₂ ice are consistently higher than those obtained from the H₂ analysis. The difference may come from differences between the actual and adopted extinction law. Alternatively, a weak component of foreground H₂ may also lower the extinction in this band.

Despite the difference in the absolute extinction values, this test independently verifies that indeed there is a lack of column density of material at the shocked knot position and the Brα fishhook-like structure. We call this low optical depth region around the shocked knot an *extinction minimum* in the envelope of HOPS 370. In Appendix D, we further explore the effects of scattered light on the optical depth of the CO₂ absorption feature.

Below, we explore the possible explanations for the observed substructures in the envelope ice density.

4.2.1. Can Intense UV Radiation from the Strong Shocked Knot Produce Localized Excess Heating?

The bright knot in the outflow, apparent in the H₂ and Brα maps, spatially coincides with the ice column density deficiency structure of the envelope. We also observe increased relative crystallization of H₂O at the shocked knot position (see Figure 7) suggesting excess heating at the shocked knot position (marked by magenta lines in all the figures). D. A. Neufeld et al. (2024) have reported a suprathermal OH emission that peaks at the same knot position and attributed the suprathermal OH emission to the photodissociation of water by UV photons. Furthermore, the second ice deficit column density structure (roughly 2'' north of the ALMA continuum position) also spatially coincides with the fishhook-like structure of Brα emission (see third panel of Figure 1), again demonstrating the presence of UV photons at the ice deficit position. These findings, taken together, suggest that UV photons originating from the shocked emission in jets/outflows

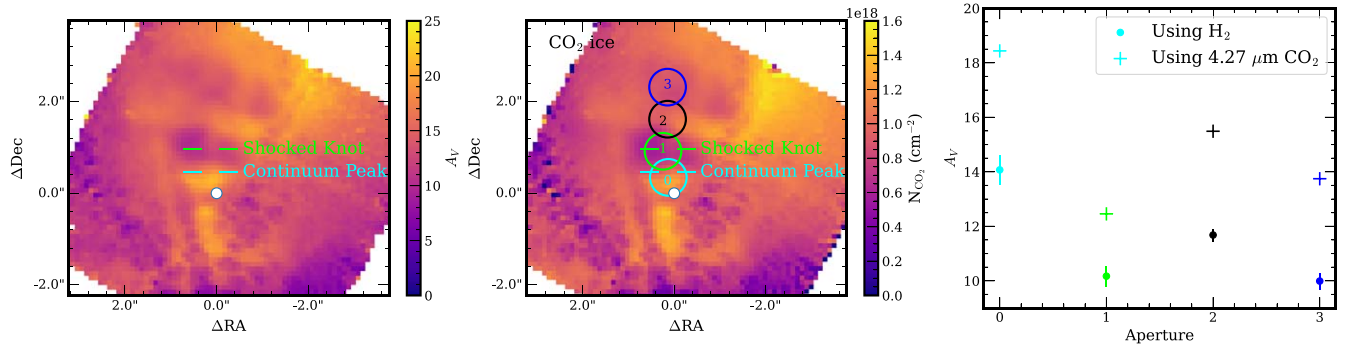


Figure 10. Extinction and column density maps using the $4.27 \mu\text{m CO}_2$ feature. The left panel shows the visual extinction (A_V) map using KP5 and the peak optical depth of the CO_2 feature. The middle panel shows the column density map of CO_2 , derived using the band strength of CO_2 at $4.27 \mu\text{m}$. The shocked knot peak of H_2 and the continuum peak are marked by lime and aqua color lines, respectively. The white-filled circle marks the ALMA continuum position. Aqua, lime, black, and blue circles, marked by numbers 0, 1, 2, and 3 are the apertures used to calculate the extinction independently using H_2 rotational lines. The right panel shows the measured extinction (A_V) corresponding to the apertures marked in the middle panel using H_2 rotational lines and $4.27 \mu\text{m CO}_2$ feature in cross and dot markers, respectively.

could be responsible for excess heating of the inner envelope wall. This heating likely leads to the ice deficit through the sublimation of icy grain mantles and enhances the relative crystallization of H_2O ice.

However in such a scenario, if we assume the shocked knot is located on the collimated jet axis, one would expect more uniform heating of the envelope instead of the heating concentrated in a localized region. Further, the low optical depth regions show low A_V values calculated using H_2 lines suggesting low total (ice + dust) column densities in these regions (see Figure 10). For this hypothesis to hold true, UV photons originating from the shocked knot would need also to destroy/sublimate dust. However, given the high sublimation temperature of silicate dust grains ($\sim 1500 \text{ K}$) and the fact that only a small fraction of UV photons escape the shocked knot region (D. A. Neufeld et al. 2024), this scenario is less likely to explain our observations.

4.2.2. Variable Extinction Due to Foreground Material or Initial Conditions

The localized ice deficit and variable extinction observed in HOPS 370 (see Figures 4 and 5) could be attributed to the initial conditions and collapse mechanism of its envelope, or possibly to variable foreground extinction.

A simple case of variable foreground extinction of the molecular cloud could in principle also explain our observations. However, the protostellar envelope is the densest part along our line of sight and column density is lower by a factor of ~ 2 at the shocked knot (see Figure 10). Given these facts, this scenario is unlikely to explain our observations.

In the case of initial conditions, we assume that the initial core density has density irregularities and asymmetries. Then, the structures in the envelope column density might arise due to a rapid turbulent collapse of the envelope which did not give enough damping time to smooth out the initial density irregularities (see also J. J. Tobin et al. 2010). In this scenario, the density deficit could be concentrated in specific regions, such as the inner cavity wall and the envelope's outer edge, or it could potentially exist throughout the envelope along our line of sight. Additionally, this case suggests an irregular supply of envelope material to the protostellar disk, which may affect episodic accretion due to disk instability. In this case, these regions would represent the initial “seeds” that might originate the gravitational instabilities at the base of accretion bursts.

Assuming that the density deficit occurs in the colder part of the envelope, the relatively high crystallization of H_2O ice at the

shocked knot can be explained. In this scenario, our observations would probe the inner heated regions of the envelope, while the line of sight would be depleted of the colder material in the outer parts of the envelope. This absence of colder material would manifest as localized excess heating within the envelope. Figure 11 displays a simplified cartoon depicting such a scenario.

5. Conclusion

Our JWST NIRSpec IFU and MIRI MRS observations of the intermediate-mass protostar HOPS 370 (OMC2-FIR3) demonstrate that during the protostellar phase the envelope can be significantly heated, thereby providing conditions for further processing of building material for the protostellar/protoplanetary disk. This study presents a rich inventory of gas-phase and ice-phase species in the spectral range from 2.8 to $28 \mu\text{m}$. Further, we present the first high spatial resolution ($\sim 80 \text{ au}$) maps of ice species in the protostellar envelope using NIRSpec data. Below we summarize our key findings from these observations:

1. The envelope of HOPS 370 exhibits clear evidence of thermally processed ices, seen in various ice absorption features: crystalline H_2O , double-peaked $^{13}\text{CO}_2$, and the double-peaked $15.2 \mu\text{m CO}_2$.
2. The decomposed maps of CO and OCN^- show that the polar mixture of ice species dominates the entire inner envelope, indicating additional evidence for thermal processing. Weak CO ice absorption compared to CO_2 together with strong CO rovibrational gas-phase lines further suggest CO ice sublimation in the envelope.
3. The optical depth and column density maps of volatile ice species reveal a nonuniform and structured density distribution of the protostellar envelope with a minimum in ice column density and line-of-sight extinction coincident with a prominent shock-heated knot in an outflow jet.
4. Dust extinction measured with H_2 rotational lines shows a region of lower extinction, and hence a lower dust column density region, at the shocked knot position.
5. The enhanced relative crystallization of H_2O observed around the shocked knot position suggests local excess heating or deficit of outer colder envelope material along the line of sight.
6. We propose that the irregular structure in the envelope has resulted in a lower column density/extinction line of sight, and the observed emission along this sightline is

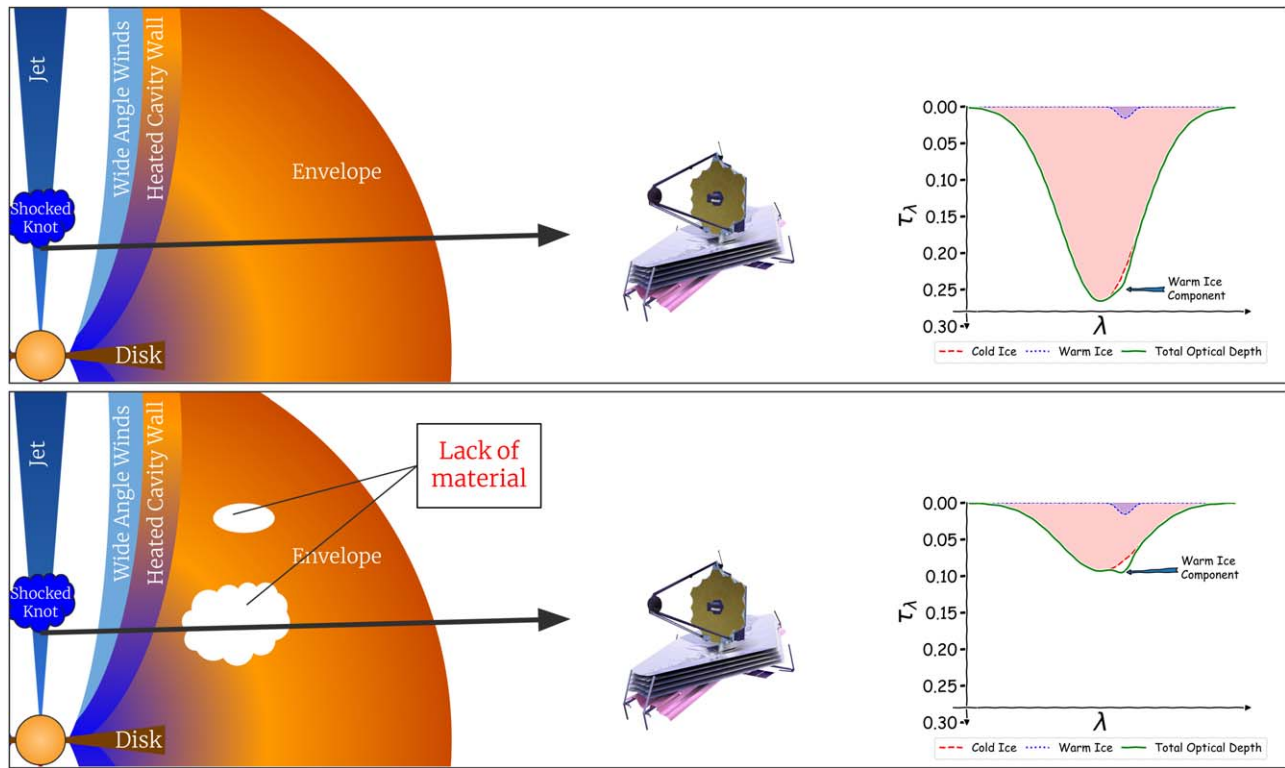


Figure 11. Simple cartoon diagram depicting the difference between the theoretical (top panel) vs. observed/proposed (bottom panel) density structure of the protostellar envelope. Cartoon models of the optical depth of an ice feature consisting of warm and cold ice are also shown for both cases. In the ice cartoon models, the green line shows the observed absorption profile by the JWST while the blue dotted line and green dashed lines represent the warm and cold ice components. The image is not to scale.

dominated by the inner thermally processed ices near the outflow cavity wall. This shows clear evidence of the thermal processing of the ice closer to the inner cavity wall. The spatial coincidence of the shocked knot seen in H_2 and $Br\alpha$ emission along this line of sight appears to be a chance alignment.

Although the high luminosity of HOPS 370 dominates the envelope heating, the role and contribution of the outflow shocks in the thermal processing of the protostellar envelope have not been fully explored yet. Future observations of low-luminosity protostars (where protostellar luminosity is inefficient in heating the envelope) with JWST may reveal the role of outflow shocks in the thermal processing of ices in the envelope.

6. Data Availability

All of the data presented in this article were obtained from MAST at the Space Telescope Science Institute. The specific observations analyzed can be accessed via doi:[10.17909/3kky-t040](https://doi.org/10.17909/3kky-t040).

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Facility: JWST (NIRSpec, MIRI).

Software: Astropy (Astropy Collaboration et al. 2013, 2018, 2022); CARTA (A. Comrie et al. 2021); Matplotlib (J. D. Hunter 2007); NumPy (C. R. Harris et al. 2020); SciPy (P. Virtanen et al. 2020).

Appendix A Line List

Table A1 presents the line list of detected molecular species. The atomic/ionic line list is presented in Table A2.

Table A1
Detected Molecular Lines

Species	Wavelength (μm)	Transition, ν	Transition, J
OH	2.9343	X3/2 1-X3/2 0	PP 4.5ee
OH	2.9346	X3/2 1-X3/2 0	PP 4.5ff
OH	2.9603	X1/2 1-X1/2 0	PP 4.5ff
OH	2.9604	X1/2 1-X1/2 0	PP 4.5ee
OH	2.9700	X3/2 1-X3/2 0	PP 5.5ee
OH	2.9704	X3/2 1-X3/2 0	PP 5.5ff
OH	3.0415	X1/2 1-X1/2 0	PP 6.5ff
OH	3.0418	X1/2 1-X1/2 0	PP 6.5ee
OH	3.0481	X3/2 1-X3/2 0	PP 7.5ee
OH	3.0487	X3/2 1-X3/2 0	PP 7.5ff
OH	3.0850	X1/2 1-X1/2 0	PP 7.5ff
OH	3.0855	X1/2 1-X1/2 0	PP 7.5ee
OH	3.0906	X3/2 1-X3/2 0	PP 8.5ee
OH	3.0913	X3/2 1-X3/2 0	PP 8.5ff
OH	3.1354	X3/2 1-X3/2 0	PP 9.5ee
OH	3.1363	X3/2 1-X3/2 0	PP 9.5ff
H ₂	5.0530	0-0	S(8)
H ₂	5.5116	0-0	S(7)
H ₂	5.8112	1-1	S(7)
H ₂ O	5.8608	010-000	2 2 0-1 1 1
H ₂ O	5.8830	010-000	2 2 1-1 1 0
H ₂ O	5.8967	010-000	5 0 5-4 1 4
H ₂ O	5.9356	010-000	4 1 4-3 0 3
H ₂ O	5.9701	010-000	4 0 4-3 1 3
H ₂ O	5.9904	010-000	3 1 3-2 0 2
H ₂ O	6.0142	010-000	2 2 1-2 1 2
H ₂ O	6.0490	010-000	2 1 2-1 0 1
H ₂ O	6.0521	010-000	3 0 3-2 1 2
H ₂ O	6.0670	010-000	2 2 0-2 1 1
H ₂ O	6.0757	010-000	3 2 1-3 1 2
H ₂	6.1088	0-0	S(6)
H ₂ O	6.1140	010-000	3 1 2-3 0 3
H ₂ O	6.1166	010-000	1 1 1-0 0 0
H ₂ O	6.1434	010-000	2 0 2-1 1 1
H ₂ O	6.1596	010-000	2 1 1-2 0 2
H ₂ O	6.1857	010-000	1 1 0-1 0 1
H ₂ O	6.3447	010-000	1 0 1-1 1 0
H ₂ O	6.3705	010-000	2 0 2-2 1 1
H ₂ O	6.3740	010-000	2 2 1-3 1 2
H ₂ O	6.3905	010-000	1 1 1-2 0 2
H ₂ O	6.4094	010-000	3 1 2-3 2 1
H ₂ O	6.4165	010-000	3 0 3-3 1 2
H ₂ O	6.4205	010-000	0 0 0-1 1 1
H ₂ O	6.4338	010-000	5 1 4-5 2 3
H ₂	6.4388	1-1	S(6)
H ₂ O	6.4848	010-000	2 1 2-2 2 1
H ₂ O	6.4925	010-000	2 1 2-3 0 3
H ₂ O	6.4978	010-000	1 0 1-2 1 2
H ₂ O	6.5195	010-000	3 2 1-3 3 0
H ₂ O	6.5226	010-000	3 1 3-3 2 2
H ₂ O	6.5423	010-000	3 2 2-3 3 1
H ₂ O	6.5554	010-000	4 2 3-4 3 2
H ₂ O	6.5676	010-000	2 0 2-3 1 3
H ₂ O	6.5738	010-000	4 1 4-4 2 3
H ₂ O	6.5904	010-000	3 1 3-4 0 4
H ₂ O	6.6287	010-000	4 2 3-5 1 4
H ₂ O	6.6421	010-000	1 1 0-2 2 1
H ₂ O	6.6723	010-000	1 1 1-2 2 0
H ₂ O	6.6836	010-000	4 1 4-5 0 5
H ₂ O	6.7078	010-000	4 0 4-5 1 5
H ₂ O	6.7123	010-000	6 1 6-6 2 5
H ₂ O	6.7235	010-000	2 1 1-3 2 2

Table A1
(Continued)

Species	Wavelength (μm)	Transition, ν	Transition, J
H ₂ O	6.7745	010-000	5 1 5-6 0 6
H ₂ O	6.7865	010-000	5 0 5-6 1 6
H ₂ O	6.7934	010-000	3 1 2-4 2 3
H ₂ O	6.8265	010-000	2 1 2-3 2 1
H ₂ O	6.8531	010-000	4 1 3-5 2 4
H ₂ O	6.8577	010-000	2 2 0-3 3 1
H ₂ O	6.8642	010-000	2 2 1-3 3 0
H ₂	6.9098	0-0	S(5)
H ₂ O	6.9600	010-000	3 2 1-4 3 2
H ₂ O	6.9934	010-000	3 2 2-4 3 1
H ₂ O	7.0242	010-000	3 1 3-4 2 2
H ₂ O	7.0451	010-000	3 3 1-4 4 0
H ₂ O	7.1182	010-000	5 2 3-6 3 4
H ₂ O	7.1471	010-000	4 2 3-5 3 2
H ₂ O	7.1647	010-000	4 3 1-5 4 2
H ₂ O	7.1714	010-000	4 3 2-5 4 1
H ₂ O	7.2079	010-000	4 4 1-5 5 0
H ₂ O	7.2726	010-000	4 1 4-5 2 3
H ₂	7.2802	1-1	S(5)
CH ₄	7.4748	00011F2-00001A1	6F2 2-5F1 1
CH ₄	7.4761	00011F2-00001A1	6F1 1-5F2 1
CH ₄	7.4791	00011F2-00001A1	6E 1-6F2 2
CH ₄	7.4797	00011F2-00001A1	6F2 1-5F1 2
CH ₄	7.5034	00011F2-00001A1	5A2 1-4A1 1
CH ₄	7.5044	00011F2-00001A1	5F2 1-4F1 1
CH ₄	7.5051	00011F2-00001A1	5E 1-4E 1
CH ₄	7.5070	00011F2-00001A1	5F1 1-4F2 1
CH ₄	7.5335	00011F2-00001A1	4F2 1-3F1 1
CH ₄	7.5343	00011F2-00001A1	4F1 1-3F2 1
CH ₄	7.5354	00011F2-00001A1	4A1 1-3A2 1
CH ₄	7.5634	00011F2-00001A1	3E 1-2E 1
CH ₄	7.5638	00011F2-00001A1	3F1 1-2F2 1
CH ₄	7.5940	00011F2-00001A1	2F2 1-1F1 1
CH ₄	7.6253	00011F2-00001A1	1A2 1-0A1 1
CH ₄	7.6554	00011F2-00001A1	6A2 1-6A1 1
CH ₄	7.6561	00011F2-00001A1	3A1 1-3A2 1
CH ₄	7.6570	00011F2-00001A1	4F1 2-4F2 1
CH ₄	7.6575	00011F2-00001A1	5F2 3-5F1 2
CH ₄	7.6581	00011F2-00001A1	1F2 1-1F1 1
CH ₄	7.6587	00011F2-00001A1	2F1 1-2F2 1
CH ₄	7.6590	00011F2-00001A1	6F2 3-6F1 1
CH ₄	7.6595	00011F2-00001A1	5E 2-5E 1
CH ₄	7.6603	00011F2-00001A1	2E 1-2E 1
CH ₄	7.6605	00011F2-00001A1	3F1 2-3F2 1
CH ₄	7.6628	00011F2-00001A1	6F1 3-6F2 2
CH ₄	7.6637	00011F2-00001A1	3F2 1-3F1 1
CH ₄	7.6652	00011F2-00001A1	4E 1-4E 1
CH ₄	7.6673	00011F2-00001A1	4F2 2-4F1 1
CH ₄	7.6704	00011F2-00001A1	4A2 1-4A1 1
CH ₄	7.6724	00011F2-00001A1	5F1 2-5F2 1
CH ₄	7.6759	00011F2-00001A1	5F2 2-5F1 1
CH ₄	7.6802	00011F2-00001A1	6A1 1-6A2 1
CH ₄	7.6831	00011F2-00001A1	6F1 2-6F2 1
CH ₄	7.6842	00011F2-00001A1	6E 2-6E 1
CH ₄	7.6907	00011F2-00001A1	0F2 1-1F1 1
CH ₄	7.7239	00011F2-00001A1	1E 1-2E 1
CH ₄	7.7257	00011F2-00001A1	1F1 1-2F2 1
CH ₄	7.7582	00011F2-00001A1	2F2 2-3F1 1
CH ₄	7.7612	00011F2-00001A1	2F1 2-3F2 1
CH ₄	7.7651	00011F2-00001A1	2A1 1-3A2 1
CH ₄	7.7914	00011F2-00001A1	3A2 1-4A1 1
CH ₄	7.7943	00011F2-00001A1	3F2 2-4F1 1
CH ₄	7.7965	00011F2-00001A1	3E 2-4E 1
CH ₄	7.8027	00011F2-00001A1	3F1 3-4F2 1
CH ₄	7.8280	00011F2-00001A1	4F2 4-5F1 1
CH ₄	7.8318	00011F2-00001A1	4F1 3-5F2 1
CH ₄	7.8408	00011F2-00001A1	4E 2-5E 1
CH ₄	7.8429	00011F2-00001A1	4F2 3-5F1 2
CH ₄	7.8642	00011F2-00001A1	5E 3-6E 1
CH ₄	7.8653	00011F2-00001A1	5F1 4-6F2 1
CH ₄	7.8692	00011F2-00001A1	5A1 1-6A2 1

Table A1
(Continued)

Species	Wavelength (μm)	Transition, ν	Transition, J
CH ₄	7.8804	00011F2-00001A1	5F1 3-6F2 2
CH ₄	7.8841	00011F2-00001A1	5F2 4-6F1 1
CH ₄	7.8875	00011F2-00001A1	5A2 2-6A1 1
H ₂	8.0257	0-0	S(4)
H ₂	8.4534	1-1	S(4)
OH	9.1337	X1/2 0-X1/2 0	RR 45.5ee
OH	9.1384	X1/2 0-X1/2 0	RR 44.5ee
OH	9.1546	X1/2 0-X1/2 0	RR 43.5ff
OH	9.1813	X1/2 0-X1/2 0	RR 42.5ff
OH	9.2186	X1/2 0-X1/2 0	RR 41.5ff
OH	9.2664	X1/2 0-X1/2 0	RR 40.5ff
OH	9.3251	X1/2 0-X1/2 0	RR 39.5ff
OH	9.3947	X1/2 0-X1/2 0	RR 38.5ff
OH	9.4757	X1/2 0-X1/2 0	RR 37.5ff
OH	9.5686	X1/2 0-X1/2 0	RR 36.5ff
H ₂	9.6656	0-0	S(3)
OH	9.6726	X3/2 0-X3/2 0	RR 36.5ee
OH	9.6738	X1/2 0-X1/2 0	RR 35.5ff
OH	9.7921	X1/2 0-X1/2 0	RR 34.5ff
OH	9.9241	X1/2 0-X1/2 0	RR 33.5ff
OH	10.0707	X1/2 0-X1/2 0	RR 32.5ff
OH	10.0707	X1/2 0-X1/2 0	RR 32.5ff
H ₂	10.1778	1-1	S(3)
OH	10.2329	X1/2 0-X1/2 0	RR 31.5ff
OH	10.4120	X1/2 0-X1/2 0	RR 30.5ff
OH	10.6093	X1/2 0-X1/2 0	RR 29.5ff
OH	10.8263	X1/2 0-X1/2 0	RR 28.5ff
OH	11.0475	X3/2 0-X3/2 0	RR 28.5ff
OH	11.0522	X1/2 0-X1/2 0	RR 27.5ee
OH	11.0647	X1/2 0-X1/2 0	RR 27.5ff
OH	11.3267	X1/2 0-X1/2 0	RR 26.5ff
OH	11.6103	X3/2 0-X3/2 0	RR 26.5ee
OH	11.6147	X1/2 0-X1/2 0	RR 25.5ff
OH	11.6147	X1/2 0-X1/2 0	RR 25.5ff
OH	11.9314	X1/2 0-X1/2 0	RR 24.5ff
OH	12.2556	X3/2 0-X3/2 0	RR 24.5ff
H ₂	12.2796	0-0	S(2)
OH	12.6574	X3/2 0-X3/2 0	RR 23.5ee
OH	12.6646	X1/2 0-X1/2 0	RR 22.5ff
OH	13.0807	X3/2 0-X3/2 0	RR 22.5ee
OH	13.0893	X1/2 0-X1/2 0	RR 21.5ff
OH	13.5492	X3/2 0-X3/2 0	RR 21.5ee
OH	13.5596	X1/2 0-X1/2 0	RR 20.5ff
HCN	13.7598	0110-0000	R 4e
HCN	13.8160	0110-0000	R 3e
HCN	13.8726	0110-0000	R 2e
HCN	13.9297	0110-0000	R 1e
HCN	13.9873	0110-0000	R 0e
HCN	14.0295	0110-0000	Q 10e
HCN	14.0324	0110-0000	Q 9e
HCN	14.0350	0110-0000	Q 8e
HCN	14.0373	0110-0000	Q 7e
HCN	14.0393	0110-0000	Q 6e
HCN	14.0410	0110-0000	Q 5e
HCN	14.0425	0110-0000	Q 4e
HCN	14.0436	0110-0000	Q 3e
HCN	14.0445	0110-0000	Q 2e
HCN	14.0451	0110-0000	Q 1e
OH	14.0692	X3/2 0-X3/2 0	RR 20.5ee
OH	14.0818	X1/2 0-X1/2 0	RR 19.5ff
HCN	14.1630	0110-0000	P 2e
HCN	14.2225	0110-0000	P 3e
HCN	14.2826	0110-0000	P 4e
CO ₂	14.5528	01101-00001	R 24e
CO ₂	14.5867	01101-00001	R 22e
CO ₂	14.6208	01101-00001	R 20e
OH	14.6482	X3/2 0-X3/2 0	RR 19.5ee
CO ₂	14.6549	01101-00001	R 18e

Table A1
(Continued)

Species	Wavelength (μm)	Transition, ν	Transition, J
OH	14.6636	X1/2 0-X1/2 0	RR 18.5ff
CO ₂	14.6892	01101-00001	R 16e
CO ₂	14.7235	01101-00001	R 14e
CO ₂	14.7580	01101-00001	R 12e
CO ₂	14.7925	01101-00001	R 10e
CO ₂	14.8271	01101-00001	R 8e
CO ₂	14.8618	01101-00001	R 6e
CO ₂	14.8966	01101-00001	R 4e
CO ₂	14.9315	01101-00001	R 2e
CO ₂	14.9564	01101-00001	Q 34e
CO ₂	14.9595	01101-00001	Q 32e
CO ₂	14.9624	01101-00001	Q 30e
CO ₂	14.9652	01101-00001	Q 28e
CO ₂	14.9664	01101-00001	R 0e
CO ₂	14.9677	01101-00001	Q 26e
CO ₂	14.9701	01101-00001	Q 24e
CO ₂	14.9722	01101-00001	Q 22e
CO ₂	14.9742	01101-00001	Q 20e
CO ₂	14.9760	01101-00001	Q 18e
CO ₂	14.9777	01101-00001	Q 16e
CO ₂	14.9791	01101-00001	Q 14e
CO ₂	14.9803	01101-00001	Q 12e
CO ₂	14.9814	01101-00001	Q 10e
CO ₂	14.9823	01101-00001	Q 8e
CO ₂	14.9830	01101-00001	Q 6e
CO ₂	14.9835	01101-00001	Q 4e
CO ₂	14.9838	01101-00001	Q 2e
CO ₂	15.0191	01101-00001	P 2e
CO ₂	15.0543	01101-00001	P 4e
CO ₂	15.0896	01101-00001	P 6e
CO ₂	15.1249	01101-00001	P 8e
CO ₂	15.1604	01101-00001	P 10e
CO ₂	15.1959	01101-00001	P 12e
CO ₂	15.2316	01101-00001	P 14e
CO ₂	15.2673	01101-00001	P 16e
OH	15.2674	X3/2 0-X3/2 0	RR 18.5ff
CO ₂	15.3031	01101-00001	P 18e
CO ₂	15.3390	01101-00001	P 20e
CO ₂	15.3749	01101-00001	P 22e
OH	16.0214	X3/2 0-X3/2 0	RR 17.5ee
OH	16.0451	X1/2 0-X1/2 0	RR 16.5ff
OH	16.8404	X3/2 0-X3/2 0	RR 16.5ee
OH	16.8703	X1/2 0-X1/2 0	RR 15.5ff
H ₂	17.0362	0-0	S(1)
OH	17.7693	X3/2 0-X3/2 0	RR 15.5ee
OH	17.8076	X1/2 0-X1/2 0	RR 14.5ff
OH	18.7921	X3/2 0-X3/2 0	RR 14.5ff
OH	18.8299	X3/2 0-X3/2 0	RR 14.5ee
OH	18.8487	X1/2 0-X1/2 0	RR 13.5ee
OH	18.8795	X1/2 0-X1/2 0	RR 13.5ff
OH	20.0085	X3/2 0-X3/2 0	RR 13.5ff
OH	20.0497	X3/2 0-X3/2 0	RR 13.5ee
OH	20.0819	X1/2 0-X1/2 0	RR 12.5ee
OH	20.1151	X1/2 0-X1/2 0	RR 12.5ff
OH	21.4199	X3/2 0-X3/2 0	RR 12.5ff
OH	21.4649	X3/2 0-X3/2 0	RR 12.5ee
OH	21.5169	X1/2 0-X1/2 0	RR 11.5ee
OH	21.5525	X1/2 0-X1/2 0	RR 11.5ff
OH	23.0739	X3/2 0-X3/2 0	RR 11.5ff
OH	23.1234	X3/2 0-X3/2 0	RR 11.5ee
OH	23.2048	X1/2 0-X1/2 0	RR 10.5ee
OH	23.2433	X1/2 0-X1/2 0	RR 10.5ff
OH	25.0351	X3/2 0-X3/2 0	RR 10.5ff
OH	25.0899	X3/2 0-X3/2 0	RR 10.5ee
OH	25.2164	X1/2 0-X1/2 0	RR 9.5ee
OH	25.2579	X1/2 0-X1/2 0	RR 9.5ff

Note. Molecular line information is taken from HITRAN (I. E. Gordon et al. 2022).

Table A2
Detected Atomic and Ionic Lines

Species	Wavelength (μm)	Transition
[Fe II]	5.341	$^4F_{9/2}-^6D_{9/2}$
[Ni II]	6.635	$^2D_{3/2}-^2D_{5/2}$
[Fe II]	6.7235	$^4F_{9/2}-^6D_{7/2}$
[Ar II]	6.985	$^2P_{1/2}-^2P_{3/2}$
[Co II]	10.523	$^3F_3-^3F_4$
[Ni II]	10.682	$^4F_{7/2}-^4F_{9/2}$
[Cl I]	11.333	$^2P_{1/2}-^2P_{3/2}$
[Ni II]	12.729	$^4F_{5/2}-^4F_{7/2}$
[Ne II]	12.814	$^2P_{1/2}-^2P_{3/2}$
[Fe II]	17.936	$^4F_{7/2}-^4F_{9/2}$
[Fe II]	24.522	$^4F_{5/2}-^4F_{7/2}$
[S I]	25.245	$^3P_1-^3P_2$
[Fe II]	25.985	$^6F_{7/2}-^6F_{9/2}$

Note. Line information is taken from NIST (A. Kramida & Y. Ralchenko 1999).

Appendix B

Detected Ice Species

We present a list of detected ice species in Table B1.

Table B1
Detected Ice Species

Absorption Band (μm)	Absorbing Ice	Identification
3.0	H ₂ O	OH stretching
3.47	NH ₃ ·H ₂ O	N–OH hydrate
3.53	CH ₃ OH	CH stretching
4.1 (a)	HDO	OD stretching
4.27	CO ₂	CO stretching
4.38 (b)	¹³ CO ₂	CO stretching
4.45 (c)	CH ₃ CN/C ₂ H ₅ CN (tentative)	CN stretching
4.62	OCN [−]	CN stretching
4.67	CO	CO stretching
4.9	OCS	CO stretching
6	H ₂ O	H ₂ O Bending
6.85	CH ₃ OH/NH ₄ ⁺	C–H deformation/NH ₄ ⁺ bend
7.23	C ₂ H ₅ OH/HCOOH	CH deformation
7.38	CH ₃ CHO/HCOO [−]	CH deformation
7.67	CH ₄	CH ₄ deformation
9.7	Silicates blended with NH ₃ , CH ₃ OH	SiO stretching NH stretching, CO stretching
15.2	CO ₂	CO ₂ bending

References. (a) K. Slavicinska et al. (2024), (b) N. G. C. Brunken et al. (2024), (c) P. Nazari et al. (2024).

Appendix C

H₂ Rotational Diagrams

Figure C1 displays the H₂ rotational diagrams used to calculate the extinction towards the envelope.

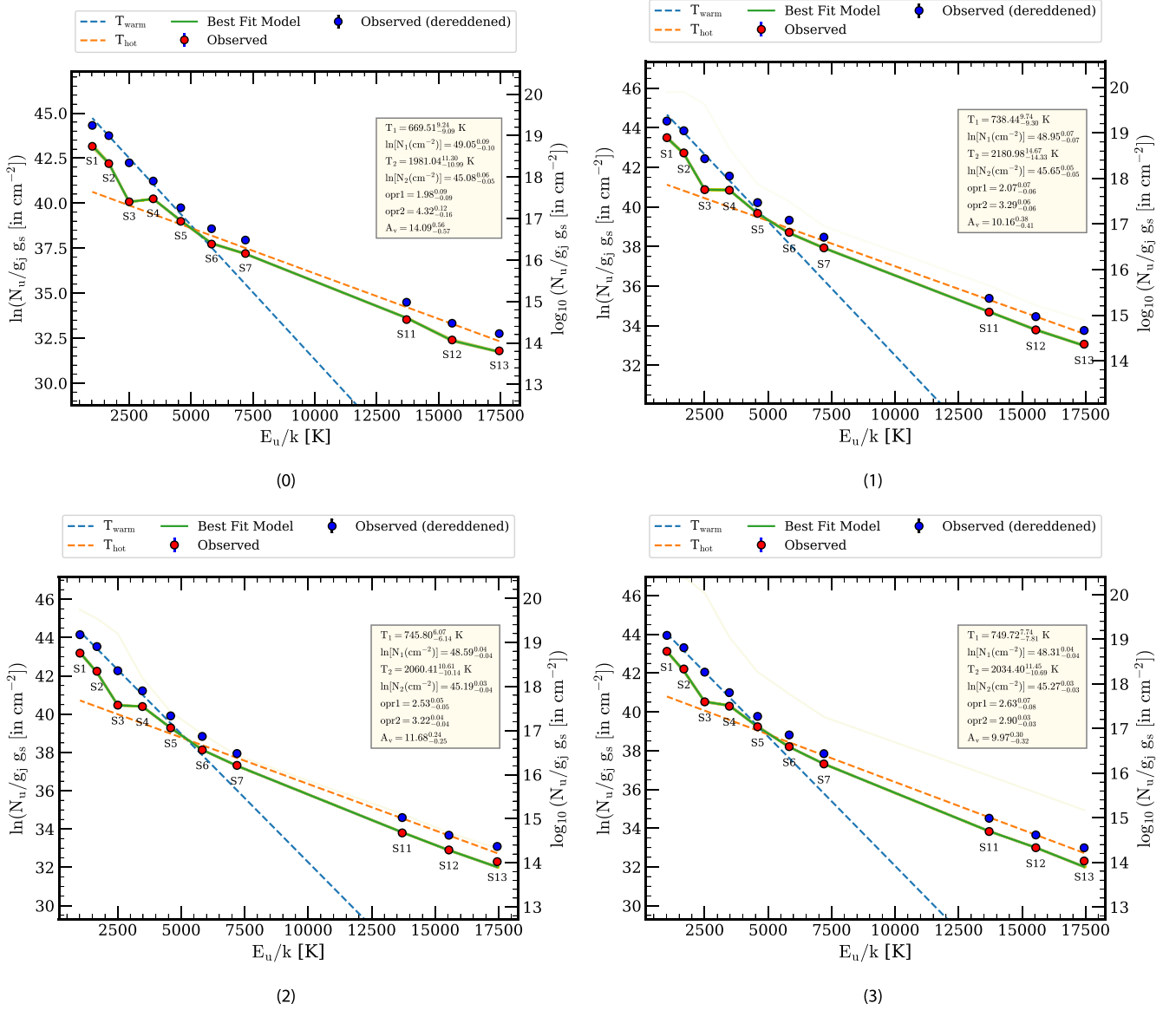


Figure C1. H₂ rotational diagram for apertures 0, 1, 2, and 3 displayed in Figure 10. Red points are the observed data points while the blue points are de-reddened points using the KP5 extinction law. Blue and orange dashed lines correspond to cold and hot temperature gas considered in the fitting model. The green line is the best fit for the observed rotational diagram.

Appendix D

CO₂ and ¹³CO₂ Ratio Maps

To further investigate the nature of the hole in the ¹³CO₂ absorption, we examine the CO₂/¹³CO₂ ratio in HOPS 370's envelope. We generated a ¹³CO₂ column density map using a band strength (density-corrected), $A = 1.15 \times 10^{-16} \text{ cm mol}^{-1}$ (P. A. Gerakines et al. 1995; M. Bouilloud et al. 2015, similar to the CO₂ column density map discussed in Section 4.2). The column density ratio map of CO₂/¹³CO₂ is shown in the top panel of Figure D1. We calculated the mean molecular ratios of CO₂/¹³CO₂ in the aperture marked by magenta and black colors at the shocked knot and the continuum peak in Figure D1, yielding values of 63.5 ± 9.8 and 62.0 ± 8.9 , respectively. Moving away from the shocked knot in both the north and south directions, we observed a decreasing trend in the CO₂/¹³CO₂ ratios (see bottom panel of Figure D1).

The CO₂/¹³CO₂ ratio can vary depending on the band strengths used in the column density calculation. Figure D1 (bottom panel) compares ratios calculated with band strengths from P. A. Gerakines et al. (1995; density-corrected values, shown in multiple colors) and M. Bouilloud et al. (2015; in gray; $A = 6.8 \times 10^{-17} \text{ cm mol}^{-1}$ for CO₂ and $A = 7.6 \times 10^{-17} \text{ cm mol}^{-1}$ for ¹³CO₂). The mean ratio difference between the two sets is a factor of 1.17. For this discussion, we adopt the density-corrected band strengths of P. A. Gerakines et al. (1995) as M. Bouilloud et al. (2015) suspected that their band strength values are lower due to possible saturation of 4.27 CO₂ absorption feature. For a detailed discussion on the band strength uncertainties, the reader is referred to N. G. C. Brunken et al. (2025, in preparation).

The obtained molecular ratios of CO₂/¹³CO₂ in HOPS 370's envelope, closer to the continuum peak and shocked knot, are similar to the average local interstellar medium (ISM) ratio of 69 (A. C. A. Boogert et al. 2000), as well as the local dark cloud ratio of 69–87 measured in the Chameleon I star-forming region (M. K. McClure et al. 2023). However, farther from the continuum peak and shocked knot the ratio decreases by a factor of ~ 1.3 – 1.7 . It is worth noting that A. C. A. Boogert et al. (2000) reported variations in the CO₂/¹³CO₂ ratio in some sources. Along the lines of sight toward GC 3, W 33A, and HH 100, they measured low CO₂/¹³CO₂ ratios of 52, 53, and 52, respectively.

Alternatively, the variation in the ratio may result from radiative transfer effects. J. A. Sturm et al. (2023) suggested that the saturation of the main isotopologue of the CO₂ ice absorption band by scattered light could lead to a lower CO₂/¹³CO₂ ratio. They observed a molecular ratio of 14 in the edge-on protoplanetary disk HH 48 NE. As mentioned in Section 3, our optical depth calculations also utilize scattered light as a background continuum. Therefore, it is plausible that low CO₂/¹³CO₂ ratios farther away from the central protostar in our observations may be attributed to the partial saturation of the CO₂ main isotopologue. This can lower the observed optical depth of the CO₂ feature, resulting in an underestimate of the column density in the protostellar envelope.

Despite this potential issue with the CO₂ optical depth. The presence of the low-opacity hole is clearly supported by other evidence. The hole is seen in multiple ice including ¹³CO₂ and H₂O. In addition, the independent extinction measurement using H₂ rotational lines agrees with the low extinction at the shocked knot position.

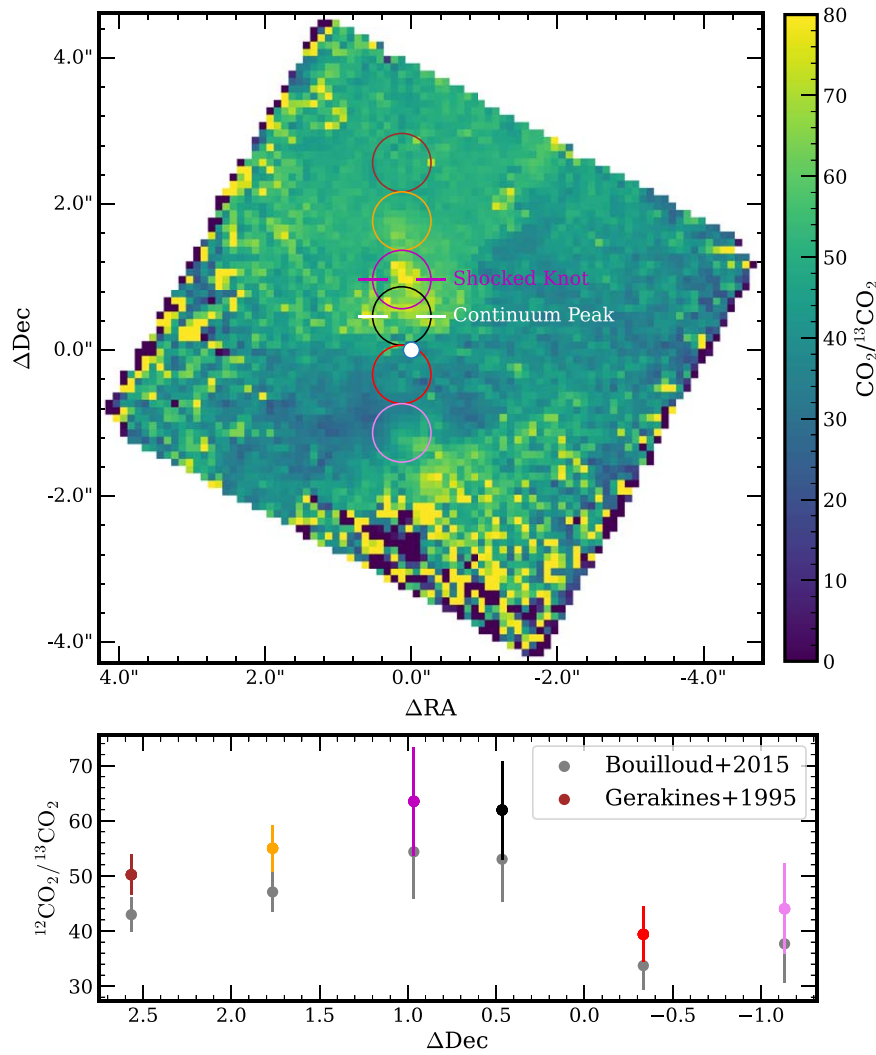


Figure D1. Top panel: CO_2 ($4.27 \mu\text{m}$) to $^{13}\text{CO}_2$ ($4.38 \mu\text{m}$) ratio map. Shocked knot and continuum peaks are marked by magenta and white color lines, respectively. Colored apertures marked on the map are used to calculate the mean $\text{CO}_2/^{13}\text{CO}_2$ ratio plotted in the bottom panel. Bottom panel: Mean $\text{CO}_2/^{13}\text{CO}_2$ ratio of marked apertures in the top panels as a function of the angular distance of the aperture from the ALMA continuum position. Multicolored and gray markers represent ratio values calculated using the density-corrected band strength values from P. A. Gerakines et al. (1995) and M. Bouilloud et al. (2015), respectively.

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