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The *JWST* FRESCO survey: legacy NIRC*am*/grism spectroscopy and imaging in the two GOODS fields

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ABSTRACT

We present the *JWST* cycle 1 53.8 h medium program FRESCO, short for ‘First Reionization Epoch Spectroscopically Complete Observations’. FRESCO covers 62 arcmin² in each of the two GOODS/CANDELS fields for a total area of 124 arcmin² exploiting *JWST*’s powerful new grism spectroscopic capabilities at near-infrared wavelengths. By obtaining ~ 2 h deep NIRC*am*/grism observations with the *F444W* filter, FRESCO yields unprecedented spectra at $R \sim 1600$ covering 3.8–5.0 μm for most galaxies in the NIRC*am* field of view. This setup enables emission line measurements over most of cosmic history, from strong PAH lines at $z \sim 0.2$ –0.5, to Pa α and Pa β at $z \sim 1$ –3, He I and [S III] at $z \sim 2.5$ –4.5, H α and [N II] at $z \sim 5$ –6.5, up to [O III] and H β for $z \sim 7$ –9 galaxies. FRESCO’s grism observations provide total line fluxes for accurately estimating galaxy stellar masses and calibrating slit-loss corrections of NIRS*pec*/MSA spectra in the same field. Additionally, FRESCO results in a mosaic of *F182M*, *F210M*, and *F444W* imaging in the same fields to a depth of ~ 28.2 mag (5σ in $0''.32$ diameter apertures). Here, we describe the overall survey design and the key science goals that can be addressed with FRESCO. We also highlight several, early science results, including: spectroscopic redshifts of Lyman break galaxies that were identified almost 20 yr ago, the discovery of broad-line active galactic nuclei at $z > 4$, and resolved Pa α maps of galaxies at $z \sim 1.4$. These results demonstrate the enormous power for serendipitous discovery of NIRC*am*/grism observations.

Key words: surveys – galaxies: evolution – galaxies: formation – galaxies: high-redshift – dark ages, reionization, first stars.

1 INTRODUCTION

Revealing the dramatic build-up of galaxies from $z > 6$ to the peak of star formation at $z \sim 2$ –3 is one of astronomy’s great achievements with the *Hubble Space Telescope* (*HST*) and *Spitzer Space Telescopes*. Data from these observatories revealed that the first 1 Gyr of cosmic history ($z > 6$) was a time of rapid change: soon after the birth of the first stars from metal-free primordial gas, the galaxy population grew rapidly both in star formation and stellar mass (e.g. Finkelstein 2016; Stark 2016; Oesch et al. 2018; Bouwens et al. 2023). After $z \sim 8$, the star formation and stellar mass density grew ~ 10 – $30\times$ up to the peak of cosmic star formation at $z \sim 2$ –4 (see Fig. 7; e.g. Madau & Dickinson 2014; Bouwens et al. 2015; Song et al. 2016; Davidzon et al. 2017; Furtak et al. 2021; Stefanon et al. 2021).

Even though this general picture of early galaxy build-up is well established, the foundation it stands on remains uncertain. In

particular, *HST*-based analyses only probe the rest-frame ultraviolet (UV) of $z > 3$ galaxies, which can lead to a bias against dusty or old galaxies (e.g. Casey, Narayanan & Cooray 2014; Casey et al. 2018; Wang et al. 2019; Barrufet et al. 2023; Rodighiero et al. 2023; Xiao et al. 2023). Additionally, broad-band rest-frame optical imaging can be contaminated by extremely strong emission lines that appear to be common at $z \geq 4$, complicating accurate measurements of the underlying continuum and inferences of stellar masses and ages (e.g. Schaerer & de Barros 2009; Labbé et al. 2013; Stark et al. 2013; Bisigello et al. 2019; Stefanon et al. 2022; Endsley et al. 2023).

The most significant shortcoming before the advent of *JWST*, however, remained the lack of spectroscopically confirmed redshifts. In particular, studies exploring the earliest galaxies in the epoch of reionization (EoR) with spectroscopy were sparse before *JWST* due to disappearing Ly α lines in the neutral IGM (e.g. Stark et al. 2010; Treu et al. 2013; Pentericci et al. 2014), and due to low Ly α fractions in general. Consequently, at $z > 7$, only a handful of luminous galaxies had spectroscopic confirmations, based on Ly α or weak rest-frame UV emission (e.g. Oesch et al. 2015; Stark et al. 2017), or from ALMA ISM lines (e.g. Inoue et al. 2016; Hashimoto et al.

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2018; Bakx et al. 2020; Bouwens et al. 2022; Schouws et al. 2022). This caused extra uncertainty given that the inferences on galaxy abundances at high redshift are statistical in nature. The agreement between candidate catalogues from different teams often remained limited, even when based on the same data: <60 per cent overall, dropping to ~ 30 per cent at the faint end (e.g. compare catalogues from Bouwens et al. 2015; Finkelstein et al. 2015). Similarly, blind spectroscopic surveys, e.g. with VIMOS or MUSE, up to $z \sim 6$ consistently showed that a non-negligible fraction (up to 60 per cent) of genuine high-redshift galaxies can be missed in traditional Lyman Break colour–colour selections due to scatter in their observed colours (e.g. Le Fèvre et al. 2015; Inami et al. 2017).

With the advent of multiplexed infrared spectroscopy of the *JWST*, all these issues that plagued the exploration of early galaxies can finally be overcome. Already after just a few months of observations with *JWST*, extremely early galaxy candidates are now identified out to $z \sim 12$ –16 (e.g. Finkelstein et al. 2022; Harikane et al. 2022; Naidu et al. 2022; Adams et al. 2023; Atek et al. 2023; Bouwens et al. 2023; Donnan et al. 2023) and spectroscopically confirmed to $z \sim 13$ (e.g. Curtis-Lake et al. 2022; Arrabal Haro et al. 2023; Bunker et al. 2023). Additionally, rest-frame optical photometry and galaxy selections are enabled beyond $z > 3$ (e.g. Barrufet et al. 2023; Nelson et al. 2023; Pérez-González et al. 2023; Rodighiero et al. 2023).

In this paper, we present the *JWST* FRESCO survey, short for ‘First Reionization Epoch Spectroscopically Complete Observations’ (GO-1895; PI Oesch). FRESCO fully exploits the unprecedented grism spectroscopic ability of *JWST* at ~ 4 –5 μm to obtain a *complete* sample of early star-forming emission line galaxies down to $\sim 2 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ across cosmic history in two $\sim 62 \text{ arcmin}^2$ *F444W* NIRCam/grism mosaics in the extragalactic legacy fields CANDELS/Deep in GOODS-South and North (Fig. 1 and Table 1).

FRESCO is designed to obtain a spectrum with $R \sim 1600$ for every reionization-era galaxy in the field down to ~ 0.2 –0.5 L_{*UV} , probing the strong rest-frame optical emission lines [O III] + H β at $z = 6.7$ –9.0 and H α + [N II] at $z = 4.9$ –6.6, in addition to other emission lines across the full cosmic history. This includes dust-insensitive Paschen lines at $z \sim 1$ –3, as well as PAH 3.3 μm lines at $z < 0.5$. Therefore, the FRESCO observations enable accurate measurements of early star formation and stellar mass build up; provide detailed insights into early, low-metallicity star formation; reveal small-scale 3D clustering and measure the contribution of mergers to early galaxy assembly. These key science cases will be discussed in this paper, after first introducing the survey design.

This paper is organized as follows: in Section 2, we introduce the observational design of the FRESCO grism and imaging survey. In Section 3, we describe the main scientific goals of FRESCO, before we end with a summary in Section 4. Throughout this paper a standard Lambda cold dark matter cosmology is adopted with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$.

2 OBSERVATIONAL DESIGN

The primary goals of FRESCO are to resolve two of the key limitations of our current understanding of the galaxy stellar mass build-up during the epoch of reionization at $z > 6$: (1) uncertain redshifts and outlier fractions, and (2) unknown emission line contribution to rest-frame optical broad-band photometry. In the following, we outline the survey design to overcome these limitations.

2.1 Slitless grism spectra

The core of FRESCO is NIRCam slitless grism observations, which result in a spectrum for every source in the field of view (see e.g. Sun

et al. 2023). FRESCO follows in the footsteps of numerous successful grism surveys conducted with HST such as 3D-HST (Brammer et al. 2012; Momcheva et al. 2016), GLASS (Treu et al. 2015), FIGS (Pirzkal et al. 2018), CLEAR (Estrada-Carpenter et al. 2019), WISP (Atek et al. 2010), and GRAPES (Pirzkal et al. 2004), but with greater sensitivity, at longer wavelength, and with spectral resolution a factor $10\times$ higher.

Unlike traditional slit-based spectroscopy that can only target pre-selected sources, slitless spectra effectively provide a complete sample of emission line selected galaxies. With *JWST*, the NIRISS instrument provides slitless spectra up to $\sim 2 \mu\text{m}$. The NIRCam camera (Rieke, Kelly & Horner 2005; Rieke et al. 2023) includes a grism mode at longer wavelengths, providing spectra at relatively high resolution, $R \sim 1600$, up to 5 μm . In cycle 1, this mode is only used by a few programs, including the GTO/EIGER survey (Kashino et al. 2023; Matthee et al. 2023b), as well as the ERS/CEERS program (Finkelstein et al. 2022), who are all using the *F356W* filter. As we will show below, such grism observations at 3–5 μm are extremely powerful probes of galaxy build-up across cosmic history.

2.1.1 Choice of filter and grism

Grism observations can use the same set of medium and wide filters as NIRCam imaging in its long-wavelength (LW) channel. The choice of filter determines the wavelength range of the spectra. The key redshift range that FRESCO targets is $z > 6$, i.e. the reionization epoch, where spectroscopic redshifts can be determined from strong rest-optical emission lines such as H α + [N II] and [O III] + H β . FRESCO therefore obtains 4–5 μm grism spectra with the *F444W* filter. This setup allows us to probe H α + [N II] lines for galaxies at $z = 4.9$ –6.6 and [O III] + H β lines from $z = 6.7$ –9.0. If bright enough, the [O II] line could in principle even be detected in sources at $z = 9.3$ –12.5. Hence, the *F444W* grism spectra result in near complete coverage of the epoch of reionization out to some of the most distant galaxies known in the Universe.

NIRCam is equipped with two different grisms in the LW channel that provide the same spectral resolution, but have a dispersion direction that is rotated by 90 degree relative to each other (GrismR and GrismC). FRESCO is using only one of them: GrismR. The advantages of taking data with both grisms are a better handle on contamination along the dispersion direction and unambiguous identification of sources in the direct image from which the spectra originate. Unfortunately, using GrismC results in a disproportionate increase of overheads, given that out-of-field images need to be taken at 35 arcsec offsets, which results in separate visits. As an example, if both GrismC and GrismR were used, FRESCO’s survey time would have increased by 15 h (~ 30 per cent) without any gain in sensitivity. Even with only one grism observation, most shortcomings can be mitigated, as demonstrated by the extremely successful *HST* grism survey 3D-*HST*, from which the community has built extensive experience to deal with overlapping spectra (Brammer et al. 2012). Due to these excessive overhead costs, FRESCO is obtained with GrismR observations only, similarly to the GTO/EIGER survey (GTO-1243; PI Lilly).

2.1.2 NIRCam/grism mosaic and wavelength coverage

The wavelength coverage of the NIRCam/grism depends on the location of the source on the detector. At a wavelengths of 3.95 μm , the *F444W* grism and the direct image coincide. However, at other wavelengths, the positions are shifted by 9.85 \AA per pixel, varying only slightly across the field. This means that only a certain fraction of sources from the direct image will obtain complete spectral coverage.

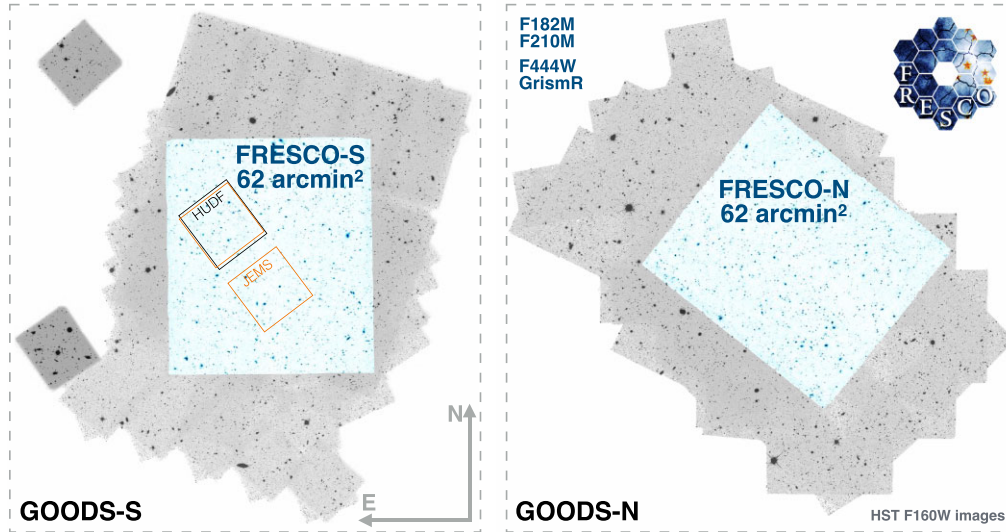


Figure 1. FRESKO was designed to cover the majority of the two CANDELS/Deep regions in the GOODS-S and -N fields, which are among the most valuable extragalactic legacy fields in the full sky. Among others, the FRESKO fields overlap with the medium-band survey JEMS (orange) that covers the original HUDF/XDF (black). The vast amount of ancillary data available in these fields made these fields the obvious choice for spectroscopic follow-up. At the highest redshifts, more than 40 per cent of all known $z \sim 7$ –8 candidate galaxies lie in these fields, with GOODS-N featuring an overdensity of $z \sim 7$ –8 candidates. Additional overdensities are known to exist at lower redshifts, which will be mapped out with these data. The FRESKO mosaics are obtained in a 4×2 mosaic with partial overlap for optimal wavelength coverage. The total area amounts to $7.3 \times 8.5 \text{ arcmin}^2$ both with grism spectroscopy and medium band imaging, which further enhances the legacy value of these fields.

Table 1. Summary of the FRESKO fields and areas.

Field	RA	Dec.	Area (arcmin^2)
FRESKO-S	03:32:31	−27:48:04	60.9 (grism) 62.4 (imaging)
FRESKO-N	12:36:46	+ 62:14:45	60.9 (grism) 62.4 (imaging)

Additionally, the grism spectra cover sources from outside the field of view. The GrismR spectra of detector A and B are dispersed in the opposite direction. By obtaining a second pointing offset by $\sim 1.8 \text{ arcmin}$, FRESKO achieves coverage over 3.9 – $4.4 \mu\text{m}$ over almost the full field of view and partial coverage from 3.9 – $5 \mu\text{m}$ over 4.9 arcmin in x -direction. In the y -direction, FRESKO obtained 4 pointings with minimal overlap (thus spanning 8.5 arcmin). Therefore, each FRESKO field covers $6.7 \times 8.5 \text{ arcmin}^2$ (57 arcmin^2) at 3.9 – $4.4 \mu\text{m}$ and $4.9 \times 8.5 \text{ arcmin}^2$ (42 arcmin^2) at 3.9 – $5 \mu\text{m}$ (corresponding to 73 per cent of the full mosaic shown in Fig. 1). More details on the spectral coverage and the spectral extraction will be provided in Brammer et al. (in preparation).

2.1.3 Exposure times and sensitivity

FRESKO is designed to reach a 5σ emission-line sensitivity for NIR-Cam/grism observations of $\sim 2 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ (for compact sources), which is the expected [O III]5008 flux of sources at 1 mag below L_* at $z = 7$ –9 (based on the LF derived from De Barros et al. 2019; see also Matthee et al. 2023b). This is achieved with eight grism exposures of 880s taken with the MEDIUM2 readout mode with 9 groups for each pointing. This results in a total grism exposure time of 7043 s at each position (see also Table 2). Four large scale dithers are used, but no sub-pixel dithering. The grism exposures are taken in two sets of these four-point dithers, in each set using a different short-wavelength filter for imaging (see Section 2.2). An

Table 2. Summary of exposure times per FRESKO pointing. In total, 2 mosaics of 4×2 pointings were obtained. Line sensitivity estimated for compact sources, integrated over full extension of the line. Imaging depths measured as the 5σ rms in $0''.32$ diameter apertures, without corrections to total fluxes.

Filter	Exposures	Time (s)	5σ Depth
F444W grism	8	7043	$2 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$
F444W imaging	3	934	28.2 mag
F210M imaging	8	3522	28.2 mag
F182M imaging	11	4456	28.4 mag

example of the resulting FRESKO imaging and spectroscopy in the GOODS-S field is shown in Fig. 2.

2.1.4 Grism data reduction

Details of the NIRCam/grism data reduction will be provided in a data release paper (Brammer et al., in preparation). Briefly, we use the publicly available `grizli` code,¹ with slightly modified sensitivity curves and spectral traces based on the v4 grism configuration files provided by Nor Pirzkal.² `grizli` is used to reduce all images and spectra and to align them to *Gaia*-matched reference frames. Given that we are mostly interested in emission line detections, continuum-subtracted spectra are created using a running median filter along each row. Following Kashino et al. (2023), the filter uses a 12 pixel central gap to minimize self-subtraction. For each source in the imaging catalogue, optimally extracted spectra are then produced based on the individual input exposures. The spectral catalogues will be released together with individual science papers and with a future data release paper.

¹<https://github.com/gbrammer/grizli>

²<https://github.com/npirzkal/GRISMCONF>

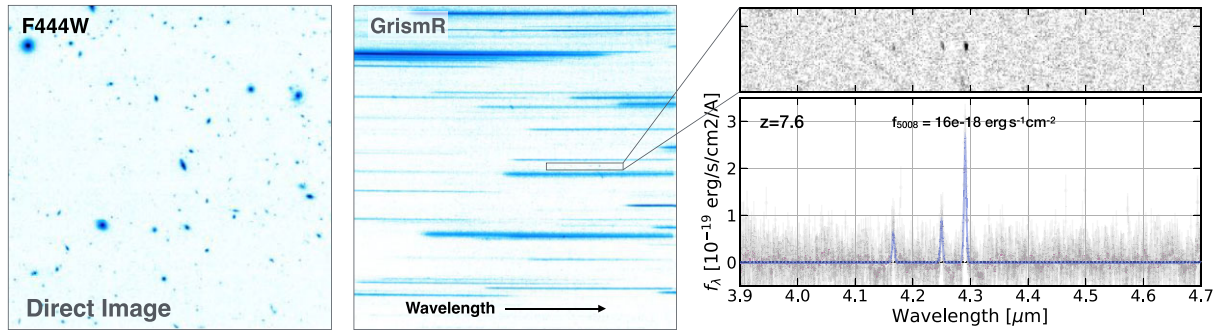


Figure 2. Example of FRESCO’s spectroscopic strategy. The left-hand panel shows a portion of the *F444W* direct image in the FRESCO-S field. The middle panel shows the associated GrismR data, from which FRESCO can identify EoR galaxies as emission line sources. Down to UV continuum magnitudes of 26.8, $z = 6.7\text{--}9.0$ sources are expected to show three emission lines from the [O III] doublet as well as $H\beta$, resulting in unambiguous redshift identifications. For example, the right-hand panels show the 2D and the extracted 1D spectra of a source in the FRESCO-S field with an unambiguous grism redshift at $z = 7.6$.

2.2 FRESCO’s Imaging Legacy

Simultaneous with the grism spectra, FRESCO obtains NIRC*am* medium band images in two short-wavelength filters (*F182M* and *F210M*) to extend the space-based legacy data in these fields to 1.8 and 2.1 μm down to 28.3 mag (at 5σ). These medium-band filters seamlessly extend the NIR wavelength coverage of the *HST* data in the CANDELS/Deep area, complementing the GTO/JADES NIRC*am* imaging. These data improve the UV spectral slope measurement of $z > 7$ galaxies and, at lower redshift, allow the community to push rest-frame optical imaging analyses beyond the peak of cosmic star formation, from $z \sim 2.5$ to beyond $z \sim 4$.

After the grism exposures, three direct and out-of-field images are taken in the *F444W* filter. These images are needed to associate the spectra with individual galaxies. The exposure times of these direct images are set to produce a $>4\sigma$ detection of every single galaxy for which a significant emission-line detection is expected. This ensures that spectra can be associated with the imaged galaxies in the field. FRESCO thus obtained 934s exposures in *F444W*, split over three offset positions, leading to a 5σ -depth of 28.2 mag (see Table 2).

The medium-band images are taken with the same readout mode as the grism exposures (MEDIUM2, 9 groups), resulting in an exposure time of 3522s each. The *F182M* filter is used during the three *F444W* imaging exposures, which are taken with the SHALLOW4 readout pattern with 6 groups. It thus receives an additional integration of 934s. The total exposure times per pointing for each filter are listed in Table 2.

2.3 Target fields

FRESCO is split into two separate fields with an equal setup covering two of the most well-studied fields in the sky: the central regions of the two GOODS fields (North and South; Giavalisco, Steidel & Macchetto 1996). These were also covered by the Deep tier of the CANDELS survey (Grogin et al. 2011; Koekemoer et al. 2011). In GOODS-South, the FRESCO field further covers the HUDF/XDF field (Beckwith et al. 2006; Ellis et al. 2013; Illingworth et al. 2013). Additional imaging over these fields was taken over the years by a very large number of programs. A complete listing of *HST* programs can be found on the Hubble Legacy Field (HLF) release page³ (see also Whitaker et al. 2019; Illingworth et al. 2016).

³<https://archive.stsci.edu/prepds/hlf/>

These fields are thus among the most valuable extragalactic legacy fields, especially for distant galaxy science. More than 40 per cent of all *HST*-selected $z \sim 7\text{--}8$ candidate galaxies from blank-field surveys lie in these two areas alone (e.g. Bouwens et al. 2015; Finkelstein et al. 2015). One of the first sources spectroscopically confirmed with $\text{Ly}\alpha$ at $z > 7$ lies in GOODS-N ($z = 7.5$; Finkelstein et al. 2013) within a spectroscopically confirmed overdensity of $z \sim 7.5$ galaxies (Jung et al. 2020). GOODS-N also contains the most distant confirmed galaxy with *HST*: GN-z11 (Oesch et al. 2016; Bunker et al. 2023), a dusty, red QSO at $z \sim 7$ (Fujimoto et al. 2022), as well as an enigmatic dusty source, HDF850.1, embedded in another overdensity at $z \sim 5.2$ (Walter et al. 2012). By targeting both GOODS fields, FRESCO thus enables a first estimate of the diversity of cosmic structures across cosmic history, well into the EoR.

Both of the FRESCO fields are (partially) observed with the proprietary NIRC*am*/GTO data from the JADES survey (Eisenstein et al. 2023; see also Helton et al. 2023), and with two of the MIRI GTO programs (Rieke et al., in preparation; Ostlin et al., in preparation; see also Rieke et al. 2015; Rinaldi et al. 2023). Additionally, the FRESCO data overlap with the public programs *JWST* Extragalactic Medium-band Survey (JEMS; Williams et al. 2023) and NGDEEP (Bagley et al. 2023).

2.4 Data acquisition

The FRESCO survey design results in a total science time of 35.5 h, for a total of 53.8 h including overheads. The FRESCO-S data in the GOODS-South field were acquired between 2022 November 13 and 18, at an orientation of ~ 0 deg (V3PA in the range 353.9–0.6). The FRESCO-N data in the GOODS-North field were obtained between 2023 February 11 and 13, at an orientation of 230.5 (V3PA in the range 230.37–230.59). The central position of FRESCO-N was offset slightly from the centre of the CANDELS/Deep region in order to capture a few high-profile objects and an expected overdensity of $z \sim 7.5$ galaxies, as discussed in the previous section.

3 SCIENTIFIC GOALS

FRESCO exploits the unique capability of the NIRC*am*/grism to obtain deep spectra at $\sim 4\text{--}5$ μm for the entire survey field. This enables a wealth of both targeted science, as well as serendipitous discovery. Below, we discuss some of the most important scientific questions the community can address with these data.

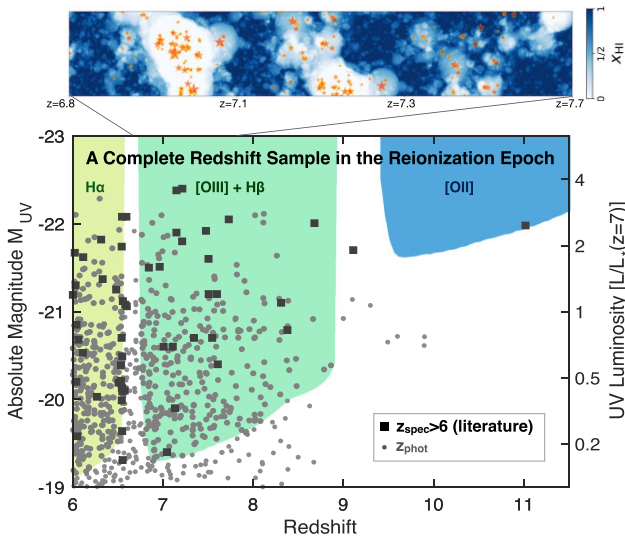


Figure 3. The average spectroscopic sensitivity of the FRESCO survey to observe strong rest-frame optical emission lines during the epoch of reionization. Before the advent of *JWST*, less than 1 per cent of known sources at $z > 6$ had spectroscopically confirmed redshifts. FRESCO is obtaining complete samples with spec- z down to $M_{UV} = -19$ ($z \sim 6$) to -20 ($z \sim 8$) thanks to strong $H\alpha$ and $[O\text{III}] + H\beta$ lines, respectively. This allows us to map out the large-scale structure during cosmic reionization. Simulations predict that ionized regions first appear around overdensities as shown by a slice through the DRAGONS simulation (top panel). The 3D large-scale structure during the epoch of reionization will for the first time be revealed thanks to FRESCO redshifts.

3.1 Spectroscopic census of early Galaxy build-up

3.1.1 Spectroscopic redshifts in the heart of reionization

The first 1 Gyr of cosmic history that constitute the cosmic reionization epoch remain a key unknown in our understanding of the Universe’s evolution (see recent reviews Dayal & Ferrara 2018; Robertson 2022). Despite enormous efforts, the galaxies in the EoR have eluded almost all attempts at spectroscopic characterization before *JWST*, apart from a handful of especially bright (or lensed) galaxies. Only two dozen ‘normal’ galaxies had confirmed spectroscopic redshifts at $z > 7$, out of more than 2000 candidates that were known from prime extragalactic HST legacy fields alone (Fig. 2; e.g. Bouwens et al. 2015; Finkelstein et al. 2015; Atek et al. 2018). The main reason for this difficulty was that the primary line for redshift confirmations, $\text{Ly}\alpha$, is severely attenuated at $z > 6$ due to absorption in the largely neutral intergalactic medium (e.g. Dijkstra 2014; Mason et al. 2018). With *JWST*, this critical shortcoming can finally be overcome through rest-frame optical spectroscopy. FRESCO probes the $[O\text{III}] + H\beta$ emission lines of galaxies at $z = 6.7\text{--}9.0$ as well as $H\alpha$ for sources at $z = 4.9\text{--}6.7$. The spectroscopically confirmed redshifts allow us to probe the small-scale clustering and 3D correlation function of galaxies during the reionization epoch (e.g. Endsley et al. 2020). From simulations, the progress of reionization is expected to be correlated with the underlying density field (see Fig. 3; Leonova et al. 2022; Qin et al. 2022).

Fig. 4 demonstrates the power of FRESCO’s NIRCcam/grism spectra to measure the redshifts of EoR galaxies. The two highlighted sources are Lyman break galaxies that were identified almost 20 years ago by Bouwens et al. (2004) from the HST/NICMOS images over the Hubble Ultra Deep Field (Thompson et al. 2005). With FRESCO,

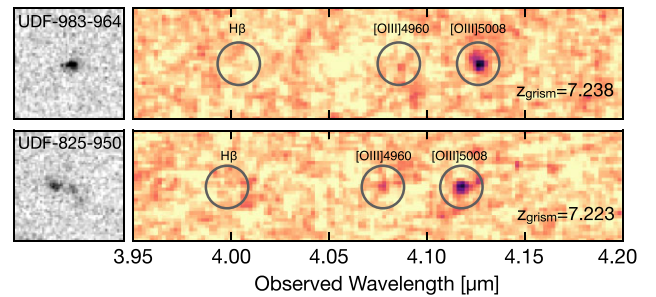


Figure 4. An example of how the NIRCcam/grism data from the FRESCO survey can finally obtain spectroscopic redshift confirmation for sources that have been identified as high-redshift galaxy candidates almost 20 yr prior. The two galaxies shown here have already been detected with the NICMOS camera in Bouwens et al. (2004). The left stamps correspond to the $F444W$ direct image, while the right-hand panels show the continuum-subtracted grism data. The $[O\text{III}]$ doublet is significantly detected in both cases resulting in an unambiguous redshift confirmation. The $H\beta$ line is found to be weaker for both sources. This example demonstrates the power of NIRCcam/grism observations to obtain an emission line selected census of galaxies in the epoch of reionization.

their high-redshift nature is confirmed through the detection of the $[O\text{III}]\lambda\lambda 4960, 5008$ doublet at $z_{\text{grism}} = 7.238$ and $z_{\text{grism}} = 7.223$, respectively. Importantly, these objects were not selected prior to the observations: all objects in the field are observed spectroscopically, without the need to pre-select targets. This is the main advantage of grism observations and enables a very wide array of science.

3.1.2 Accurate stellar masses in the heart of reionization

From *HST* + *Spitzer* observations, it has been known for several years that the emission line strengths of $[O\text{III}]$ or $H\alpha$ are rapidly increasing towards higher redshift (e.g. Schaerer & de Barros 2010; Fumagalli et al. 2012; Labbé et al. 2013), resulting in observed-frame equivalent widths that are $>5000 \text{ \AA}$ on average (e.g. Smit et al. 2014; Roberts-Borsani et al. 2016; De Barros et al. 2019; Endsley et al. 2021). This evolution has dramatic consequences for stellar mass estimates that can remain uncertain by factors up to 5–10 \times (Fig. 5) – a problem that still applies to NIRCcam imaging. Bisigello et al. (2019) estimate *median* correction factors up to 0.87 dex, if strong emission lines are not accounted for. FRESCO provides emission line measurements for all sources in the field, enabling precise corrections of the broad-band fluxes for emission line contamination on a source-by-source basis. A prominent example is shown in Fig. 5. The source in question still has one of the most distant $\text{Ly}\alpha$ detections at $z_{\text{Ly}\alpha} = 7.5$ (Finkelstein et al. 2013; Jung et al. 2020). Thanks to this, it was clear that the excess in the IRAC 4.5- μm band was due to strong rest-frame optical emission lines. Nevertheless, the strength of these lines was uncertain. The FRESCO data shows that the source is actually composed of two components with slightly offset velocities. For both components, the contribution of these emission lines to the broad-band photometry can now be assessed separately. With these corrections, FRESCO thus provides accurate stellar mass measurements over two key deep fields. Additionally, it will enable the derivation of statistical correction factors for other imaging surveys.

While these strong rest-optical emission lines can complicate stellar mass estimates, they enable efficient spectroscopic confirmations. Therefore, FRESCO’s spectra continue to propel the spectroscopic frontier into the heart of cosmic reionization by obtaining redshifts

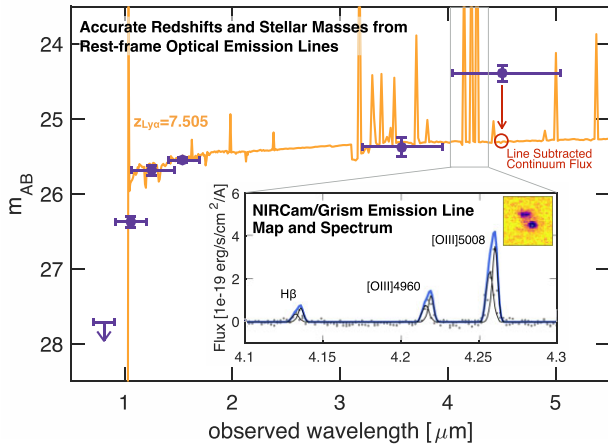


Figure 5. At $z > 5$, the rest-frame optical [O III] + H β lines are found to have observed-frame EWs of $>5000 \text{ \AA}$ – on average (Labbé et al. 2010). This enables extremely efficient spectroscopic confirmation (inset). However, such strong lines contaminate broad-band photometry, as clearly seen in the IRAC CH2 photometry of this source. This can result in stellar mass uncertainties of up to 5–10 \times . FRESCO’s spectra enable us to correct *JWST*’s broad-band photometry to obtain true rest-frame optical continuum and hence stellar mass measurements.

down to 0.2–0.5 L_* at $z \sim 7$ –9 (see Fig. 2). This critically enables measurements of the UV luminosity and stellar mass functions based on pure spectroscopic samples at $z \gtrsim 5$.

3.1.3 Metal-free star formation and the build-up of metals at $z > 3$

A major question for extragalactic surveys with *JWST* remains whether it is possible to detect primordial (PopIII/zero-metallicity) star formation. While candidate PopIII galaxies at $z \sim 7$ have been claimed in the literature (e.g. Sobral et al. 2015), none have been confirmed (see e.g. Bowler et al. 2017; Matthee et al. 2017). FRESCO has a unique discovery potential for extremely low metallicity candidates at $z > 6$. The medium-band imaging at 1.8–2 μm will result in improved UV-continuum slope measurements at $z \sim 7$ –9 enabling the identification of especially blue galaxies for which the grism [O III] and H β line ratios provide an initial gas-phase metallicity estimate ([O III]/H β decreases at $<0.2 Z_\odot$; e.g. Maiolino et al. 2008; Inoue 2011; Curti et al. 2020). FRESCO thus has the capability to provide promising, extremely metal-poor candidates for future NIRSspec follow-up. At $z \sim 4.9$ –6.6, FRESCO continues to trace the build-up of metals through [N II]/H α line ratios in order to constrain the mass–metallicity relations and models of early chemical enrichment. As such, FRESCO data allows us to trace the transition from extremely metal poor star formation at $z > 6$ to more enriched conditions that are found at later times.

3.1.4 The prevalence of AGNs among early galaxies

Another central question of extragalactic astronomy is the abundance of active galactic nuclei (AGNs) in the early Universe. These sources could provide a non-negligible contribution to reionization (e.g. Giallongo et al. 2015; Madau & Haardt 2015), and are expected to explain the abundances of supermassive black holes at later times (e.g. Volonteri 2010). Interestingly, several of the most luminous $z \sim 7$ –9 galaxies show indications for AGN activity, based on their ground-based rest-UV spectra, suggesting that the AGN fraction

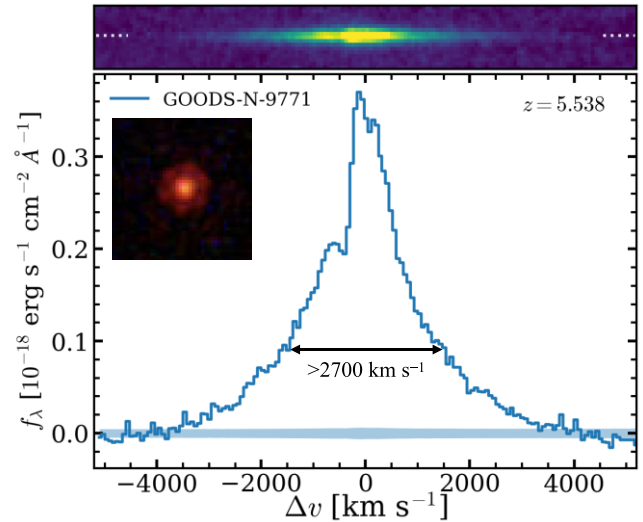


Figure 6. An example of a serendipitously discovered broad-line, low-mass AGN from FRESCO spectra. The extremely broad H α + [N II] lines identify these galaxies as hosting massive black holes (for more information, see Matthee et al. 2023a). NIRCcam/grism observations are uniquely capable of discovering such sources and confirming their number densities in the future. Given the high spectral resolution of $R \sim 1600$, such observations also have the power to reveal kinematic information for less extreme galaxies.

could be significant (Laporte et al. 2017; Sobral et al. 2018). Owing to the $R \sim 1600$ spectra, FRESCO can test this scenario directly by enabling the identification of potential AGN through broad emission lines across the full redshift range $z \sim 5$ –9 based on H α , H β , or [O III] as well as through indirect methods such as the mass excitation diagram (after a recalibration to $z \sim 7$ –9 based on *JWST* spectra; Juneau et al. 2011; Trump et al. 2013).

First *JWST* spectra have already revealed a significant population of low-mass AGN (e.g. Harikane et al. 2023; Kocevski et al. 2023; Übler et al. 2023), and more candidates are identified through NIRCcam imaging (e.g. Labbe et al. 2023). NIRCcam/grism observations have the potential to provide a complete sample of such sources. A prominent example from the FRESCO data is shown in Fig. 6, with an extremely broad H α emission line at $z \sim 5$. This source is part of a larger sample of such galaxies identified in NIRCcam/grism data (Matthee et al. 2023a).

3.2 Star formation across cosmic history

3.2.1 Unbiased star-formation rate indicators

Star formation rate (SFR) is a fundamental observable property of galaxies that is required to trace the growth and formation of galaxies throughout cosmic time. Gold standard SFR indicators are hydrogen recombination lines, as they emerge from H II regions around the most massive and recently born stars and, hence, trace the almost-instantaneous SFR in galaxies (Kennicutt & Evans 2012). Aside from Ly α , which suffers from uncertainties associated with resonant scattering, Balmer optical lines, such as H α and H β are the bright H lines and good tracers of star formation activity. However, these lines are highly affected by dust attenuation, and the uncertainties associated with their attenuation correction factors (e.g. uncertainties in the nebular reddening measurements and attenuation curve assumptions) hamper their potential as accurate SFR diagnostics (e.g. Fanelli, O’Connell & Thuan 1988; Reddy et al. 2015; Shivaeei et al.

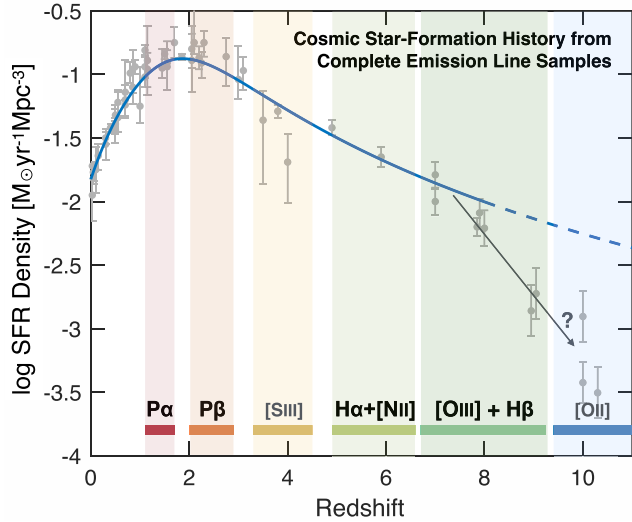


Figure 7. FRESCO is designed to return a complete sample of star-forming galaxies based on emission lines across the majority of cosmic history. FRESCO 4–5 μm spectra (a) directly constrain the debated accelerated evolution of the star formation rate density at $z > 8$ based on [O III] + H β samples at $z = 6.9\text{--}9.0$; (b) follow galaxy build-up through the reionization epoch with H α + [N II] measurements at $z = 4.9\text{--}6.6$; and (c) probe the peak of cosmic SF ($z \sim 1\text{--}3$) through dust-insensitive Pa α and Pa β maps. In the local Universe, FRESCO even covers PAH lines at 3.3 μm for a small number of galaxies (see also Fig. 8).

2018, among many more). Pa α , on the other hand, is an instantaneous and dust-insensitive SFR indicator, owing to its longer wavelength, making it an important SFR diagnostic, particularly in dusty star-forming galaxies at the peak epoch of cosmic star formation history. FRESCO takes advantage of the unprecedented near-IR capabilities of *JWST* to observe Paschen lines (Pa α and Pa β) and trace optically thick star formation in large samples of galaxies at $z \sim 1\text{--}3$, which is typically missed in optical surveys (see also Finkelstein et al. 2011; Cleri et al. 2022; Reddy et al. 2023). Another tracer of obscured star formation activity that is accessible to FRESCO is the 3.3 μm feature at $z \lesssim 0.5$, which is the emission from PAH dust grains (see e.g. Genzel & Cesarsky 2000; Kim et al. 2012).

The FRESCO spectra at $\sim 4\text{--}5 \mu\text{m}$ thus have the capability to trace star formation across the entire cosmic history (see Fig. 7). To illustrate this, Fig. 8 shows a collage of extracted emission line spectra in the FRESCO-S field. This includes 333 sources that previously had spectroscopic redshifts measured from the literature up to $z \sim 6$, in addition to 27 sources for which FRESCO detected bright [O III] + H β lines at $z \sim 7$. The spectra are sorted by redshift. Given the availability of previous spectroscopic redshifts the Pa α lines at $z \sim 1\text{--}2$ dominate the figure. Several overdensities are clearly visible. The figure shows the power of NIRCcam/grism observations to probe star formation and clustering across cosmic history.

3.2.2 Spatial distribution of star formation at cosmic noon

A unique feature of grism observations is that they result in spatially resolved maps of emission lines and hence star formation. For instance, H α maps from *HST* grism observations of $z \sim 1\text{--}2$ galaxies have been used to reveal star-forming clumps and centrally suppressed star formation in massive galaxies, consistent with inside-out growth (e.g. Wuyts et al. 2013; Nelson et al. 2016). This appears to hold, even when accounting for dust gradients inferred

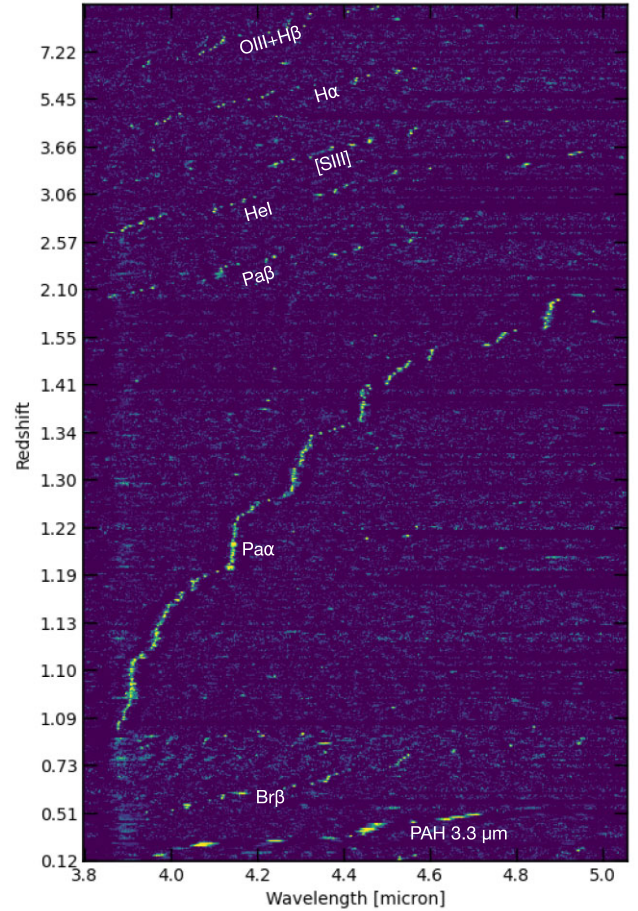


Figure 8. Line coverage at different redshifts in the NIRCcam/grism observations of FRESCO. Shown is a stack of 333 2D spectra in the GOODS-South field for sources with pre-existing spectroscopic redshifts from the literature up to $z \sim 6.5$. Beyond that, we show examples of new H β and [O III] lines measured for galaxies at $z > 6.8$. Given the large number of spectroscopic redshifts in this field at $z \sim 1\text{--}2$, Pa α lines currently dominate this figure. Galaxy overdensities are clearly visible. Full spectral extractions of the FRESCO NIRCcam/grism data will be described in Brammer et al. (in preparation).

from resolved broad-band colours or stacked Balmer decrements (e.g. Tacchella et al. 2018). However, ALMA observations of the 870 μm dust-continuum emission show a different picture: the dust continuum in many massive galaxies is centrally concentrated (Tadaki et al. 2020). These observations support a different scenario, in which violent dissipative events fuel nuclear star formation. Given the uncertainties in H α dust corrections and complications in the interpretation of the dust-continuum emission (e.g. due to different heating mechanisms, dust temperature gradients, insensitivity to diffuse emission), an instantaneous and dust-insensitive tracer of star formation is critically needed. FRESCO provides exactly this, with spatially resolved measurements of the Paschen series lines at $z \sim 1\text{--}3$. These measurements allow us to finally answer the question: do massive galaxies at cosmic noon form their last stars in their centres or in their outskirts before they quench?

A first example of a resolved Pa α map in the FRESCO-S field is shown in Fig. 9. The *grizli*-generated Pa α map indeed reveals a very compact star-forming core for this galaxy that is completely absent in the shorter wavelength imaging. This is embedded in

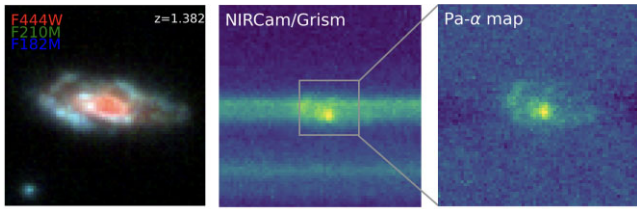


Figure 9. The spatial distribution of SFR from emission line maps across cosmic history as probed by FRESCO. NIRCcam/grism data uniquely provide spatially resolved line maps at 4–5 μm . Shown is the Pa α map from a dusty $z = 1.38$ galaxy in the FRESCO-S field. The left-hand panel shows an RGB composite of the FRESCO F182M, F210M, and F444W images. The central panel shows the NIRCcam/grism data, from which we extract an emission line map at the Paschen- α wavelength using the `grizli` tool. The FRESCO data set thus reveals the sites of star formation in galaxies.

more extended star formation throughout the galaxy. NIRCcam/grism observations such as from FRESCO enable systematic studies of the spatial distribution of star formation in cosmic noon galaxies.

3.2.3 Illuminating the complex ISM in assembling galaxies

At $z > 5$, galaxies are rapidly assembling, resulting in highly complex interstellar media, as demonstrated by ALMA observations of EoR galaxies, where spatial offsets between the rest-UV continuum, and different ISM lines such as [C II] or [O III] 88 μm are common (e.g. Maiolino et al. 2015; Carniani et al. 2017). FRESCO provides detailed H α (H β) maps of a large sample of galaxies at $z = 4.9$ –6.6 ($z = 6.7$ –9.0) to constrain how and where within galaxies stars are forming during the epoch of reionization (for first examples see the EIGER program Matthee et al. 2023b). Comparisons between H α and rest-optical continuum sizes provide a first test of inside-out growth during an epoch when galaxies are still assembling – rather than fully settled in – (cold) discs, or may reveal homologous growth with a similar pace of stellar build-up across all radii.

3.2.4 Illuminating the dark side of star formation

Over the past few years it has become more and more clear that *HST*-based surveys have been missing a potentially important population of massive, star-forming galaxies at $z > 3$ that are sufficiently obscured such that they remained undetected at rest-UV wavelengths (e.g. Caputi et al. 2015; Franco et al. 2018; Alcalde Pampliega et al. 2019; Wang et al. 2019; Fudamoto et al. 2021; Xiao et al. 2023). These galaxies have been identified either through Spitzer/IRAC or through ALMA observations. However, given the limited photometry available, their photometric redshifts and thus also their nature remained highly uncertain. With *JWST* we can now finally probe the rest-frame optical wavelengths up to $z \sim 9$, which brings this optically faint galaxy population into view (e.g. Barrufet et al. 2023; Pérez-González et al. 2023; Rodighiero et al. 2023). FRESCO’s imaging and spectroscopy is especially powerful to characterize this enigmatic population of galaxies. The grism spectra enable – for the first time – to obtain spectroscopic redshifts for such galaxies thanks to emission line searches. An example of this is shown in Fig. 10 of a galaxy that has previously been detected only at sub-millimeter wavelengths (Cowie et al. 2018; Yamaguchi et al. 2019; Gómez-Guijarro et al. 2022; Xiao et al., in preparation).

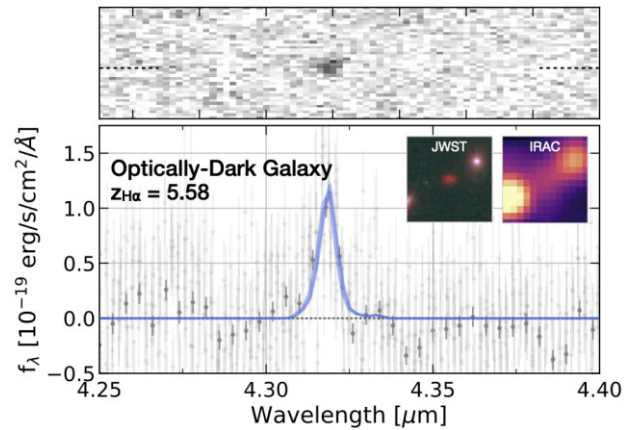


Figure 10. FRESCO’s imaging and spectroscopy has the power to reveal the nature of the enigmatic population of optically faint galaxies. Shown is an example galaxy from Xiao et al. (in preparation) in the FRESCO-S field. The source was known from ALMA, but its counterpart could not be identified in the low-spatial resolution IRAC imaging. FRESCO’s F444W images clearly reveal this galaxy, and with the NIRCcam/grism spectra its redshift can be determined through the H α emission line.

3.2.5 The emergence of quiescent galaxies

At $z < 2$, the massive galaxy population is dominated by dead, quiescent galaxies, with no significant ongoing star formation (e.g. Muzzin et al. 2013; Davidzon et al. 2017). However, it is still unknown when such quiescent galaxies first appeared in the Universe. Deep ground-based *K*-band surveys demonstrated that quiescent galaxies existed since $z \sim 4$, however, their exact number densities remain debated (Straatman et al. 2015; Davidzon et al. 2017; Forrest et al. 2020; Valentino et al. 2020), mainly because *HST*-based surveys are limited to rest-frame optical observations at $z < 2.5$. First searches with *JWST* are already finding and confirming quiescent, low star formation rate galaxies up to $z \sim 5$ –7 (see e.g. Carnall et al. 2023; Looser et al. 2023; Valentino et al. 2023). A large population of such sources in the very early Universe challenges simulations (e.g. Wellons et al. 2015; Merlin et al. 2019; Hartley et al. 2023), such that the timing of when massive galaxies grow, shut off their star formation and turn quiescent is a very powerful constraint on galaxy evolution models.

The short-wavelength imaging from FRESCO is designed to complement the JADES/GTO broad-band data in these fields and seamlessly extends *HST*’s legacy data to enable accurate selections of quiescent galaxies up to $z \sim 4$ from Balmer break measurements (see Fig. 11). In the short-wavelength channel, the imaging filters are the same as obtained with the JEMS survey over the HUDF/XDF.

3.3 Legacy science

The CANDELS/Deep fields in the centres of GOODS-S and -N are two of the most studied areas in the sky with the most comprehensive ancillary data. FRESCO builds on these data and further increases the legacy value of these fields beyond that of the GTO teams by enabling a more complete census of emission line galaxies through blind spectroscopy. Among others, FRESCO provides emission samples for future spectroscopic follow-up with NIRSpec. Thus, FRESCO’s zero-proprietary data enable a wealth of legacy science to be performed by the community. First science papers by the community using FRESCO imaging and spectroscopy

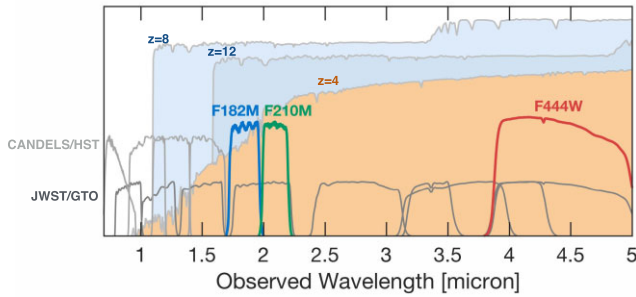


Figure 11. FRESCO obtains images in two medium-band filters at $\sim 2 \mu\text{m}$ (and a direct image in $F444W$). These data seamlessly extend *HST*'s legacy imaging over these fields and complement the GTO data resulting in vast improvements in UV-continuum slope measurements for $z > 6$ galaxies and enabling the selection of quiescent galaxies up to $z \sim 4$.

have already been published as pre-prints (Laporte et al. 2022; Helton et al. 2023).

4 SUMMARY

This paper provided an overview of the medium program FRESCO – a NIRCcam grism spectroscopy and imaging survey of two fields in the GOODS-North and -South fields, respectively. The power of NIRCcam/grism observations are clearly demonstrated with this survey: the grism provides $R \sim 1600$ spectroscopy of all sources in the field of view in its LW channel, while imaging is obtained in a short-wavelength filter at the same time. Adopting this strategy with the $F444W$ filter, FRESCO obtained deep spectra covering 3.8–5.0 μm , reaching average line sensitivities of $2 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ (5σ). At the same time, deep $F182M$ and $F210M$ images are obtained reaching ~ 28.2 mag rms (5σ in $0''.32$ diameter apertures), together with $F444W$ direct images reaching similar depth. The FRESCO observations thus significantly enhance the rich ancillary data set in these key legacy fields, enabling a vast amount of science by the community. The spectral coverage allows one to probe emission line galaxies across almost the full cosmic history. We have highlighted a few science cases, that showcase the enormous power for serendipitous discovery of such NIRCcam/grism observations. Future, wider-area grism programs would thus be able to obtain large, complete samples of rare classes of galaxies that will be very difficult to follow-up, e.g. with NIRSpec spectroscopy through targeted programs. We hope that the community will use the FRESCO data set for a large number of scientific discoveries as well as inspiration for many further *JWST* programs, including NIRCcam/grism observations.

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DATA AVAILABILITY

The *HST* and *JWST* image mosaics of the FRESCO fields are already released at MAST as a High Level Science Product via <https://doi.org/10.17909/gdyc-7g80>. The spectra are still being calibrated and will be made available together with an upcoming data paper (Brammer et al., in preparation). For updates, please check our webpage: <https://jwst-fresco.astro.unige.ch/> or the MAST page <https://archiv.e.stsci.edu/hlsp/fresco/>.

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APPENDIX A: INFORMATION OF EXAMPLE GALAXIES

In Table A1, we provide more details for the sources that are shown in the main text.

Table A1. Source information for the example galaxies shown in Figs 4, 5, 6, 9, 10.

Fig	RA	Dec.	z_{grism}
4(a)	03:32:38.81	-27:47:07.4	7.237 ± 0.001
4(b)	03:32:39.53	-27:47:17.7	7.223 ± 0.001
5(a)	12:36:37.92	62:18:08.7	7.507 ± 0.001
5(b)	12:36:37.87	62:18:08.5	7.498 ± 0.001
6	12:37:07.44	62:14:50.3	5.538 ± 0.001
9	03:32:34.04	-27:50:29.1	1.3820 ± 0.0002
10	03:32:28.91	-27:44:31.5	5.579 ± 0.001

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