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z-GAL: A NOEMA spectroscopic redshift survey of bright *Herschel* galaxies

I. Overview

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ABSTRACT

Using the IRAM Northern Extended Millimetre Array (NOEMA), we conducted a Large Programme (*z*-GAL) to measure redshifts for 126 bright galaxies detected in the *Herschel* Astrophysical Large Area Survey (H-ATLAS), the HerMES Large Mode Survey (HeLMS), and the *Herschel* Stripe 82 (HerS) Survey. We report reliable spectroscopic redshifts for a total of 124 of the *Herschel*-selected galaxies. The redshifts are estimated from scans of the 3 and 2-mm bands (and, for one source, the 1-mm band), covering up to 31 GHz in each band, and are based on the detection of at least two emission lines. Together with the Pilot Programme, where 11 sources had their spectroscopic redshifts measured, our survey has derived precise redshifts for 135 bright *Herschel*-selected galaxies, making it the largest sample of high-*z* galaxies with robust redshifts to date. Most emission lines detected are from ¹²CO (mainly from $J = 2-1$ to 5-4), with some sources seen in [CI] and H₂O emission lines. The spectroscopic redshifts are in the range $0.8 < z < 6.55$ with a median value of $z = 2.56 \pm 0.10$, centred on the peak epoch of galaxy formation. The linewidths of the sources are large, with a mean value for the full width at half maximum ΔV of $590 \pm 25 \text{ km s}^{-1}$ and with 35% of the sources having widths of $700 \text{ km s}^{-1} < \Delta V < 1800 \text{ km s}^{-1}$. Most of the sources are unresolved or barely resolved on scales of ~ 2 to $3''$ (or linear sizes of ~ 15 – 25 kpc, unlensed). Some fields reveal double or multiple sources in line emission and the underlying dust continuum and, in some cases, sources at different redshifts. Taking these sources into account, there are, in total, 165 individual sources with robust spectroscopic redshifts, including lensed galaxies, binary systems, and over-densities. This paper presents an overview of the *z*-GAL survey and provides the observed properties of the emission lines, the derived spectroscopic redshifts, and a catalogue of the entire sample. The catalogue includes, for each source, the combined continuum and emission lines' maps together with the spectra for each of the detected emission lines. The data presented here will serve as a foundation for the other *z*-GAL papers in this series reporting on the dust emission, the molecular and atomic gas properties, and a detailed analysis of the nature of the sources. Comparisons are made with other spectroscopic surveys of high-*z* galaxies and future prospects, including dedicated follow-up observations based on these redshift measurements, are outlined.

Key words. galaxies: high-redshift – galaxies: ISM – submillimeter: galaxies – gravitational lensing: strong – radio lines: ISM

1. Introduction

Over the last two decades, surveys in the far-infrared and submillimetre wavebands have opened up a new window into our understanding of the formation and evolution of galaxies, revealing a population of massive, dust-enshrouded galaxies forming stars at enormous rates in the early Universe (see, e.g., reviews by Blain et al. 2002; Carilli & Walter 2013; Casey et al. 2014; Hodge & da Cunha 2020). In particular, the wide-field extragalactic imaging surveys carried out with the *Herschel* Space Observatory have increased the number of known dust-obscured star-forming galaxies (DSFGs) from hundreds to several hundreds of thousands (Eales et al. 2010; Oliver et al. 2012; Lutz et al. 2011; Elbaz et al. 2011; Ward et al. 2022). Including the all-sky *Planck*-HFI (High Frequency Instrument; *Planck Collaboration Int. XXVII* 2015) and the South Pole Telescope (SPT) surveys (Vieira et al. 2010), today, we have large samples of luminous DSFGs that are among the brightest in the Universe, whose rest-frame 8–1000 μm luminosities (L_{IR}) exceed a few $10^{12} L_{\odot}$, and include numerous examples of strongly lensed systems (Negrello et al. 2010, 2017; Wardlow et al. 2013; Bussmann et al. 2013, 2015; *Planck Collaboration Int. XXVII*

2015; Spilker et al. 2016; Nayyeri et al. 2016; Bakx et al. 2018) and rare cases of extreme galaxies such as hyper-luminous infrared galaxies (HyLIRGs) with $L_{\text{IR}} > 10^{13} L_{\odot}$ reaching star formation rates of $1000 M_{\odot} \text{ yr}^{-1}$ or more (see, e.g., Ivison et al. 2013, 2019; Fu et al. 2013; Oteo et al. 2016; Riechers et al. 2013, 2017; Wang et al. 2021).

The exact nature of these luminous DSFGs is still debated (Narayanan et al. 2015), although many of them are mergers (Tacconi et al. 2008; Engel et al. 2010). Compared to local ultraluminous infrared galaxies, DSFGs are more luminous and several orders of magnitude more numerous by volume. DSFGs are most commonly found around $z \sim 2$ – 4 (e.g., Danielson et al. 2017), where the cosmic densities of star formation and black hole accretion peak (Madau & Dickinson 2014) and the gas accretion onto galaxies reaches its maximum (Walter et al. 2020). DSFGs therefore play a critical role in the history of cosmic star formation as the locus of the physical processes driving the most extreme phases of galaxy formation and evolution.

Detailed follow-up studies of these galaxies require precise estimates of their distances to investigate their nature and physical properties. Photometric redshifts, which are based on spectral

energy distribution templates designed for infrared-bright high- z galaxies, provide useful indications of their redshift range; however, in most cases, this method remains too imprecise for efficient dedicated follow-up observations (see, e.g., Casey et al. 2012; Chakrabarti et al. 2013; Ivison et al. 2016; Bendo et al. 2023). The dusty nature of these luminous galaxies makes optical and near-infrared spectroscopy using even the largest ground-based facilities difficult, if not impossible (Chapman et al. 2005), and, in the case of gravitational lensing, the optical spectra detect the foreground lensing galaxies rather than the background high- z lensed object.

Measuring molecular or atomic emission lines via sub-millimetre spectroscopy, which are unambiguously related to the (sub-)millimetre source and are not affected by dust extinction, has proved to be the most reliable method to derive secure redshifts for high- z dust-obscured objects (e.g., Weiß et al. 2009; Swinbank et al. 2010). The increase in the instantaneous frequency range of receivers operating at millimetre and sub-millimetre facilities has made this method more efficient over time and recent dedicated redshift surveys have provided crucial information that enable targeted studies to explore the nature of DSFGs in the early Universe. From the first sub-millimetre galaxies (SMGs) detected by the Submillimetre Common-User Bolometer Array (SCUBA) in the continuum (Smail et al. 1997; Hughes et al. 1998), whose redshifts were only measured a decade later (e.g., Chapman et al. 2005; Weiß et al. 2009; Walter et al. 2012), to subsequent observations of tens of bright sources from the *Herschel* wide surveys that enabled the measurement of their redshifts (e.g., Harris et al. 2012; Lupu et al. 2012; Riechers 2011), sub-millimetre spectroscopy was established as the prime method to precisely derive unambiguous distances to high- z DSFGs.

Due to their increased sensitivity, millimetre and sub-millimetre interferometers, in particular the IRAM Northern Extended Millimetre Array (NOEMA) and the Atacama Large Millimetre/submillimetre Array (ALMA), play a central role in measuring redshifts of high- z DSFGs selected from wide field surveys. Using ALMA, a spectroscopic redshift survey of 81 high- z strongly gravitationally lensed galaxies selected in the SPT 2500 deg² cosmological survey was presented by Reuter et al. (2020) – see, also, Weiß et al. (2013) and Strandet et al. (2016). Initially selected at 1.4 mm, the SPT-selected DSFGs have redshifts spanning the range $1.9 < z < 6.9$ with a median of $z = 3.9 \pm 0.2$. This survey contains roughly half of the spectroscopically confirmed $z > 5$ DSFGs, making it the most complete and largest high- z sample of DSFGs prior to the sample we present herein. A survey of molecular gas in SMGs selected from 870 μm continuum surveys of the COSMOS, UDS, and ECDFS fields was performed using ALMA and NOEMA by Birkin et al. (2021) resulting in spectroscopic redshifts for 45 SMGs (from the 61 originally selected) with a median redshift of $z = 3.0 \pm 0.2$. Spectroscopic redshift measurements of 17 bright SMGs selected from the AS2COSMOS survey are reported in Chen et al. (2022) with redshifts in the range $z = 2\text{--}5$ with a median of $z = 3.3 \pm 0.3$. For a sample of 24 strongly lensed high- z galaxies, identified in the *Planck* survey, measurements of their spectroscopic redshifts, followed by detailed studies of the physical gas properties through multi spectral line ¹²CO and [CI] observations, were performed using a suite of single dish telescopes (e.g., IRAM 30 m, GBT, APEX, and LMT) by Cañameras et al. (2015) and Harrington et al. (2016, 2021).

A number of very bright and/or very red *Herschel*-selected high- z sources have been observed using various facilities,

including the GBT, the IRAM 30 m telescope, CARMA, NOEMA, and ALMA, yielding secure redshifts for about ~ 70 sources (Cox et al. 2011; Harris et al. 2012; Wardlow et al. 2013; Ivison et al. 2016; Riechers et al. 2013, 2017, 2021; Negrello et al. 2017; Fudamoto et al. 2017; Bakx et al. 2018, 2020a; Nayyeri et al. 2016, and references therein). Recently, Urquhart et al. (2022) presented a spectroscopic redshift survey of 85 high- z *Herschel*-selected galaxies from the southern fields of the H-ATLAS survey (*Herschel* Astrophysical Terahertz Large Area Survey – e.g., Eales et al. 2010; Valiante et al. 2016), listed in Bakx et al. (2018, 2020a), using ALMA and the Atacama Compact Array (ACA, also known as the Morita Array), and they derived robust redshifts for 71 DSFGs that were identified in the fields of 62 of the *Herschel*-selected sources.

The correlator PolyFiX on NOEMA, with its capability to process a total, instantaneous single polarisation bandwidth of 31 GHz, alleviates one of the main problems related to the measurement of redshifts of high- z DSFGs, namely the large overheads that are currently required in spectral-scan mode. This was demonstrated by Neri et al. (2020) who conducted a Pilot Programme to measure redshifts for 13 bright galaxies detected in the northern fields of the H-ATLAS survey with $S_{500\mu\text{m}} \geq 80$ mJy and reported reliable spectroscopic redshifts for 11 of the selected sources. In one field, two galaxies with different redshifts were detected and, in two cases, the sources were found to be binary galaxies. The spectroscopic redshifts were derived from scans of the 3 and 2 mm bands, and based on the detection of at least two emission lines. The spectroscopic redshifts are in the range $2.08 < z < 4.05$ with a median value of $z = 2.9 \pm 0.6$. The linewidths of the sources were large, with a median value for the full width at half maximum (FWHM) of 800 km s^{-1} , potentially indicating cases of HyLIRGs and extreme mergers. In addition to the detailed information obtained on their structure and kinematics (mainly through emission from mid- J ¹²CO transitions), the large instantaneous bandwidth of NOEMA provides, for each source, an exquisite sampling of the underlying dust continuum emission in the 3 and 2 mm (i.e. 90 and 150 GHz) bands. Together with the studies mentioned above, the number of *Herschel*-selected high- z galaxies with precise redshifts amounts to ~ 150 sources in total.

Building upon this successful NOEMA Pilot Programme and taking advantage of the new capabilities of NOEMA (ten to 12 antennas and PolyFiX; see, e.g., Neri & Sánchez-Portal 2022), here we present the NOEMA z -GAL project, a comprehensive redshift survey of 126 of the brightest ($S_{500\mu\text{m}} \geq 80$ mJy) DSFGs with photometric redshifts $z_{\text{phot}} \geq 2$ selected from the *Herschel* surveys in northern and equatorial fields and for which no reliable redshift measurements were previously available. Secure redshifts, based on at least two emission lines, were derived for all but two of the selected sources. This makes the NOEMA z -GAL project, together with the results from the Pilot Programme, the largest of these spectroscopic redshift surveys to date.

The results of this comprehensive redshift survey are presented in a series of three papers that report on various aspects of the available data: (i) the current paper, the first in the series, describes the sample of selected *Herschel* sources, outlines the observational strategy, reviews the data reduction, and presents the basic results of the spectroscopic data, including the measurement of the spectroscopic redshifts, which was the main goal of this study; (ii) Paper II (Ismail et al. 2023) discusses the continuum properties of the sources in the sample and presents a detailed study of the derived dust properties and their evolution with redshift, based on the available spectral energy

distributions including the *Herschel*, SCUBA-2, and NOEMA data; and (iii) Paper III (Berta et al. 2023) reports on the physical properties of the sources derived from the molecular and atomic gas (^{12}CO , [CI], and H_2O) and dust, including the excitation conditions, the depletion timescales and first estimates of the stellar masses. In a subsequent study (Paper IV), Bakx et al. (in prep.) will report on the nature of the sources and identify, using ancillary optical, near-infrared and radio data, gravitationally lensed sources, buried Active Galactic Nuclei (AGN) starbursts, proto-clusters, and hyper-luminous galaxies.

The structure of this paper, the first of the z -GAL series, is as follows. In Sect. 2, we describe the sample selection, the observations, and the data reduction. In Sect. 3, we present the main results including the redshift measurements, the spectral properties of the sources, and their morphology. A catalogue of the entire sample is presented in Appendix A including, for each source, combined continuum and emission line maps together with spectra of all the emission lines that were detected. In Sect. 4, we compare the spectroscopic and photometric redshifts, discuss the widths of the CO emission lines, and present the general properties of the z -GAL sources. Finally, in Sect. 5, we summarise the main conclusions and outline future prospects.

Throughout this paper, we adopt a spatially flat ΛCDM cosmology with $H_0 = 67.4 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_M = 0.315$ (Planck Collaboration VI 2020).

2. Sample selection and observations

2.1. Sample selection

The sources of the NOEMA z -GAL Large Programme were chosen from three *Herschel* surveys.

Firstly, we selected galaxies from the HerMES Large Mode Survey (HeLMS) and the *Herschel* Stripe 82 (HerS) Survey based on the source selection reported by Nayyeri et al. (2016). Together, these partially overlapping surveys cover 372 deg^2 on the sky. From the list of 77 sources with $S_{500\mu\text{m}} \geq 100 \text{ mJy}$ listed in Nayyeri et al. (2016), we retained 60 sources for which no spectroscopic redshift data was available (43 from HeLMS and 17 from HerS). One of these sources (HeLMS-33) turned out to be a blazar (J0030–021) and was subsequently dropped. Only one source for which no spectroscopic redshift was available in the original list was overlooked in the selection and consequently could not be included in the observations, namely HeLMS-53. In total, we observed 59 *Herschel*-selected sources from HeLMS and HerS.

Secondly, we selected sources from the *Herschel* Bright Sources (HerBS) sample, which contains the 209 galaxies detected in the H-ATLAS survey with $500\mu\text{m}$ flux densities $S_{500\mu\text{m}} > 80 \text{ mJy}$ and photometric redshifts $z_{\text{phot}} > 2$ (Bakx et al. 2018). Most of these galaxies (189 in total) have been observed at $850\mu\text{m}$ using SCUBA-2, and 152 were detected. The SCUBA-2 flux densities originally reported by Bakx et al. (2018) were revised by Bakx et al. (2020c) together with the photometric redshifts estimated for each of the sources.

From the HerBS sample, we selected all the sources from the 121 high-redshift galaxies in the North Galactic Pole (NGP) and equatorial (GAMA) fields of H-ATLAS (covering 170.1 and 161.6 deg^2 , respectively), which had no spectroscopic redshift measurements. In addition to the 11 sources of the Pilot Programme (Neri et al. 2020), only 22 sources in the HerBS sample had available spectroscopic redshifts based on previous multi-line detections and 4 sources are listed with spectroscopic redshifts based on tentative, single line detections (see refer-

ences in Bakx et al. 2018, and Table A.5). In a subsequent study, Bakx et al. (2020a) presented spectroscopic measurements of four sources selected from the HerBS catalogue; two of these sources had robust redshifts derived from multi-line detections and were not included in the z -GAL Large Programme (HerBS-52 and HerBS-64), while two other sources (HerBS-61 and HerBS-177) were kept in the final sample.

In the selection process, we eliminated 14 sources: two blazars – HerBS-16 (J1410+020) and HerBS-112 (J1331+30) – and 12 weak sources ($S_{500\mu\text{m}} \leq 95 \text{ mJy}$) with low-quality SPIRE data and/or no SCUBA-2 detection, based upon the experience gained in the Pilot Programme (Neri et al. 2020). The majority of the selected sources are reliably detected in the SCUBA-2 maps with $>3\sigma$ $850\mu\text{m}$ flux densities and positionally coincident with the *Herschel* source, and have good-quality SPIRE data at $500\mu\text{m}$ with a signal-to-noise ratio $S/N \geq 10$. In addition, one source, that was part of the z -GAL Pilot Programme (HerBS-204), and for which no spectroscopic redshift could be derived by Neri et al. (2020), was added to the sample of the z -GAL Large Programme; and one source, HerBS-173, for which no emission line was detected neither at 2 or 3-mm, was dropped in the final selection. In total, we observed 67 *Herschel*-selected galaxies from H-ATLAS.

The number of HerBS, HeLMS and HerS-selected galaxies observed in the z -GAL Large Programme, that cover the northern and equatorial regions, is therefore 126. Together with the 11 HerBS-selected sources for which spectroscopic redshifts are available from the z -GAL Pilot Programme (Neri et al. 2020), the complete z -GAL survey amounts to a total of 137 *Herschel*-selected galaxies. Figure 1 shows the distribution of the 137 z -GAL-selected sources on *Herschel*/SPIRE colour maps of the H-ATLAS, HeLMS and HerS fields.

The *Herschel* (and SCUBA-2) flux densities, as well as the photometric redshifts z_{phot} (Sect. 3.3), are listed in Table A.1 for the HeLMS and HerS sources and in Table A.2 for the HerBS sources that were selected for the z -GAL Large Programme¹. The 11 sources of the z -GAL Pilot Programme are listed separately in Table A.3. Finally, the sources for which spectroscopic redshifts were available are listed in Table A.5 and those that were not retained in the final selection or not observed are listed in Table A.4.

Figure 2 displays the *Herschel*/SPIRE colour-colour plot ($S_{500\mu\text{m}}/S_{250\mu\text{m}}$ vs. $S_{500\mu\text{m}}/S_{350\mu\text{m}}$) for the z -GAL and Pilot Programme sources together with sources listed in Bakx et al. (2018) and Nayyeri et al. (2016) with spectroscopic redshifts and the sources that were excluded from the z -GAL selection. The colour-colour plot displays the selection function resulting from the flux limits in the samples together with the redshift track of a typical DSFG, the Cosmic Eyelash (SMM J2135–010 – Ivison et al. 2010; Swinbank et al. 2010), and indicates the expected photometric redshifts covered by the selected sources.

2.2. Observations

We used NOEMA to observe the 126 *Herschel*-selected bright DSFGs (Tables A.1 and A.2) in the NGP and equatorial fields and derive their redshifts by scanning the 3 and 2 mm bands to search for at least two emission lines. The observations of the z -GAL Large Programme were carried out under project M18AB (PIs: P. Cox, T. Bakx and H. Dannerbauer). Starting on December 10, 2018, NOEMA continuously and regularly

¹ For the names of the sources, we here adopt HeLMS and HerS to be consistent with the names in the HerBS catalogue.

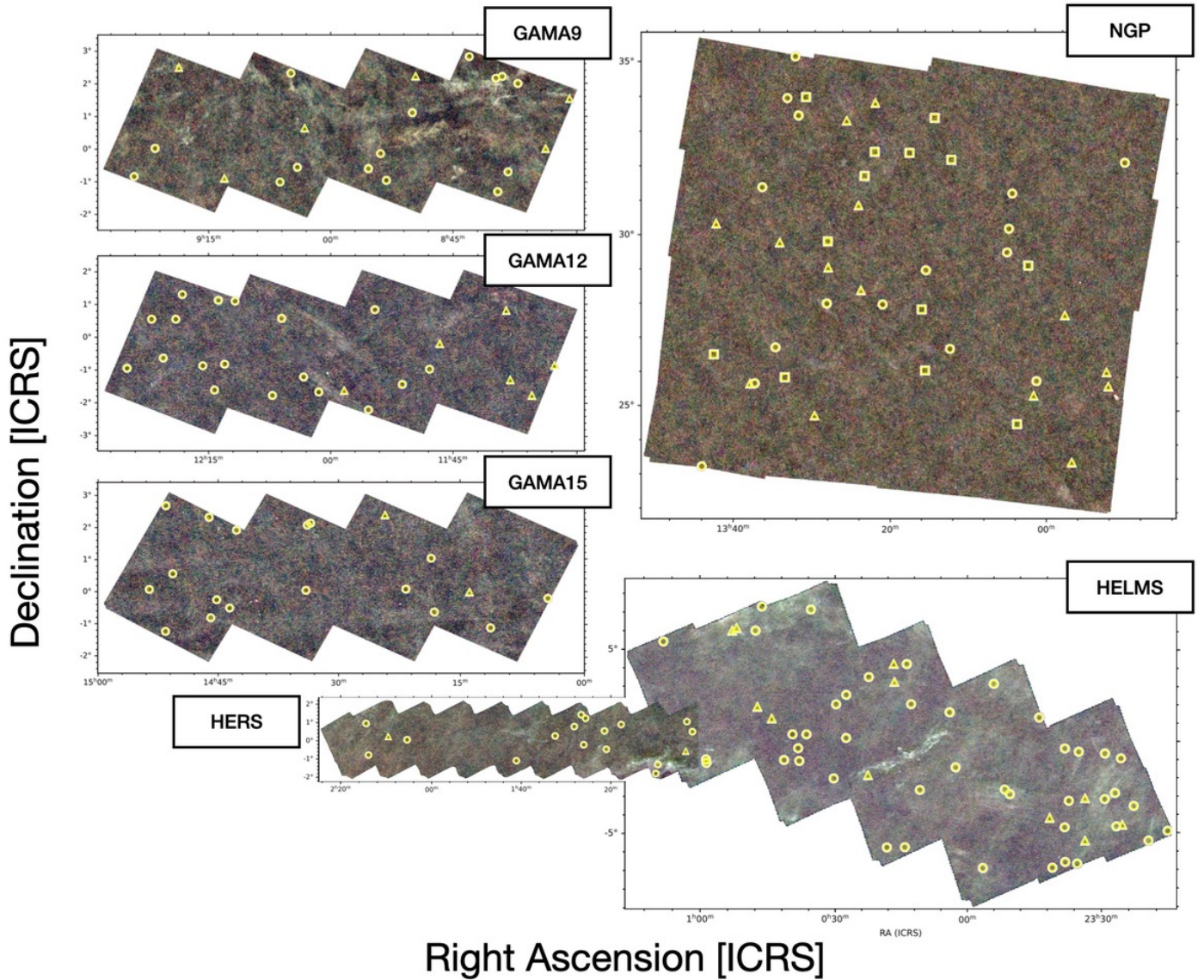


Fig. 1. *Herschel*/SPIRE colour maps of the H-ATLAS (GAMA9, GAMA12, GAMA15, and NGP), HeLMS and HerS fields with the positions of the 137 sources that were observed with NOEMA in the z -GAL project and the Pilot Programme (Neri et al. 2020), highlighted with yellow circles and squares, respectively. The yellow triangles indicate the sources for which spectroscopic redshifts were already available (see Table A.5). The maps display the distribution of the z -GAL sources across the sky and show the proximity of some of the sources that enabled for them to be grouped together during the NOEMA observations, resulting in a more efficient use of telescope time (Sect. 2.2).

observed the 126 *Herschel*-selected sources for a two-year period, using nine to ten antennas in two configurations (C and D) of the array. The z -GAL Large Programme was completed in July 26, 2020 with 190 h of observing time used. A subsequent Director’s Discretionary Time (DDT) project (D20AB) started soon thereafter with the aim of measuring an additional emission line for 13 sources that had an emission line falling outside the frequency settings originally defined for the Large Programme and secure for these sources unambiguous spectroscopic redshifts. On January 15 2021, all of the 126 *Herschel*-selected galaxies were reported in at least two emission lines and, henceforth, robust redshifts measurements were potentially acquired for the entire originally proposed sample. Following these observations and after a careful analysis of the data, a dozen sources were found to still have ambiguous redshift solutions (Sect. 3.1). Additional observations were then performed under the same DDT project to measure a third emission line in these sources. This removed remaining ambiguities, giving robust spectro-

scopic redshifts for 124 sources of the 126 selected sources, thereby successfully achieving the original goal of this project.

Observing conditions were on average excellent during the entire time frame of this project with an atmospheric phase stability of typically 10–40 deg RMS and 2–5 mm of precipitable water vapour. The correlator was operated in the low-resolution mode to provide spectral channels with a nominal resolution of 2 MHz. The NOEMA antennas are equipped with 2SB receivers that cover a spectral window of 7.744 GHz in each sideband and polarisation. Since the two sidebands are separated by 7.744 GHz, two frequency settings are necessary to span a contiguous spectral range of 31 GHz. At 3 mm, we adjusted the spectral survey to cover a frequency range from 75.874 to 106.868 GHz. At 2 mm, we then selected two frequency windows that covered as well as possible the candidate redshifts allowed by the emission lines detected at 3 mm, across the frequency range from 127.377 to 158.380 GHz (Fig. 3). Additional observations to secure the redshifts for a dozen

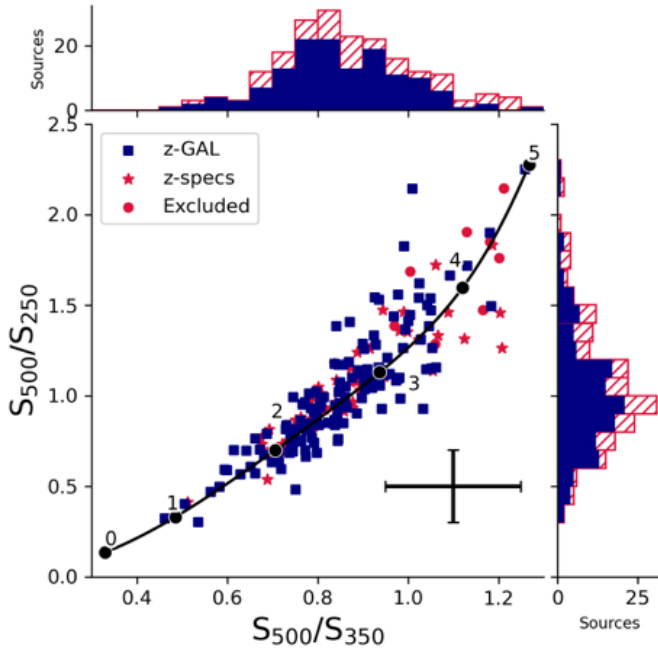


Fig. 2. *Herschel*/SPIRE colour-colour plot ($S_{500\mu\text{m}}/S_{250\mu\text{m}}$ vs. $S_{500\mu\text{m}}/S_{350\mu\text{m}}$) for the *z*-GAL and Pilot Programme sources (shown as blue squares). Typical uncertainties on the colours are shown in the bottom right corner of the figure. The sources listed in Bakx et al. (2018) and Nayyeri et al. (2016) with spectroscopic redshifts (labelled ‘*z*-spec’ and listed in Table A.5) are included for comparison and indicated with a red star. The sources that were excluded from the *z*-GAL selection (labelled “Excluded”) are indicated with a red dot (see Table A.4). The side histograms display the distribution of the values as a function of the number of sources. The overlaid curve shows the redshift track expected for a galaxy with a spectrum energy distribution comparable to the Cosmic Eyelash galaxy; the photometric redshifts are indicated along the curve (see Ivison et al. 2016).

sources were done at 3 mm and covered the frequency ranges from 70.689 to 78.433 GHz and 86.177 to 93.921 GHz. For one source (HeLMS-17), we performed additional observations at 1 mm across the frequency ranges from 208.884 to 216.628 and 224.372 to 232.116 GHz to obtain a robust redshift.

The wide spectral coverage of the NOEMA correlator ensures that a scan of both the 3 and 2 mm spectral windows can detect at least two ^{12}CO emission lines for every $z \gtrsim 1.2$ source, between $^{12}\text{CO}(2-1)$ and $^{12}\text{CO}(9-8)$, in each band, with the exception of a few redshift gaps. The gaps most relevant to the present observations are at 3 mm for $1.733 < z < 1.997$, and at 2 mm for $1.668 < z < 1.835$ (Fig. 3). The redshift range $1.733 < z < 1.835$ was not covered by any of the 3 and 2 mm settings. The spectral coverage of these observations also includes the [CI] ($^3\text{P}_1-^3\text{P}_0$) and [CI] ($^3\text{P}_2-^3\text{P}_1$) fine-structure lines, two transitions of water, namely ortho- $\text{H}_2\text{O}(4_{23}-3_{30})$ and para- $\text{H}_2\text{O}(2_{11}-2_{02})$, and $\text{HCO}^+(3-2)$ and $\text{HCN}(3-2)$, which were all detected in sources selected for this survey (Tables A.6–A.8). In total, 358 emission lines were detected in the combined *z*-GAL survey and Pilot Programme. These include among the brightest lines in the sub-millimetre range and the most likely species (van der Werf et al. 2010) to be picked up in redshift searches (see also, e.g., Bakx & Dannerbauer 2022). Exploring both the 3 and 2 mm spectral bands is therefore a prerequisite for detecting, in most cases, at least two ^{12}CO emission lines in the $2 < z < 6$ *Herschel* bright galaxies selected for this survey and deriving reliable spectroscopic redshifts.

All observations were carried out in track-sharing mode by cyclically switching between groups of galaxies within a track, as was made possible due to the proximity of many sources to each other (Fig. 1). Track-sharing allowed us to observe efficiently by not repeating pointing and some calibration observations for sources belonging to a group of five to seven galaxies in a small field (within 10° to 15°). The sky positions of the 126 sources were distributed in five large fields, namely: H-ATLAS NGP, GAMA9, GAMA12 and GAMA15, and HerMES HeLMS/HerS (H Field) shown in Fig. 1. In summary, the sources were ordered in 21 groups of five to seven sources. On average, we spent ~ 1.5 h of total observing time per source, including overheads and calibration. Two different configurations (C and D) of the array were used, yielding angular resolutions between $1''.2$ and $3''.5$ at 2 mm, and $1''.7$ and $\sim 6''$ at 3 mm. Observations were started by first observing sources in one track in the lower 3 mm frequency setting. Galaxies that did not show a robust line detection were then observed again in the upper 3 mm frequency setting. For every line detection, the most probable redshift was estimated taking into account the photometric redshift. The galaxies were subsequently observed in one of the two 2 mm frequency settings, and when no line was detected, observed again in the second setting. To resolve the cases of ambiguous redshift determinations, additional frequency windows were observed according to inclusion or exclusion criteria based on photometric redshifts and relative intensities of the detected lines (Sect. 3.1).

2.3. Data reduction

Phase and amplitude calibrators were chosen in close proximity to the track-shared sources in order to reduce uncertainties in astrometric and relative amplitude calibration. MWC349 and LkH α 101 were used as absolute flux calibrators. The positions of the source correspond to the barycenter of the emission region in the pseudo-continuum maps (Appendix A.1). The precision of the position relies on the S/N, and we typically achieved an absolute astrometric accuracy of $< 0''.2$ (Tables A.6 and A.7). The absolute flux calibration was estimated to be accurate to within 10%. The data were calibrated, averaged in polarisation, imaged, and analysed in the GILDAS software package².

The *uv*-tables were produced from calibrated visibilities in the standard way and cleaned using natural weighting and support masks defined on each of the detected sources. Continuum maps were produced for each sideband and setup, excluding the spectral ranges that include emission lines. Continuum extraction was performed on cleaned continuum maps, after correcting for primary beam attenuation, using polygonal apertures of variable size and shape, and a pseudo-continuum map as guidance. This was achieved independently for each source according to the procedure described in Appendix A.1.

In case multiple sources were detected in the field of view, the fluxes of each individual object were extracted. Source positions were computed as the barycenter of the flux distribution in the sky of the given object, within the extraction polygon. The spectra of each source were extracted from cleaned cubes within polygonal apertures defined by hand on the emission line channels to enhance the S/N of the spectral detection. The emission lines of the majority of the 165 sources revealed in the *Herschel* fields were detected with good signal-to-noise ratios. Only in a few cases, the sources (10 in total) displayed at least one emission line with $S/N \lesssim 3$.

² <http://www.iram.fr/IRAMFR/GILDAS/>

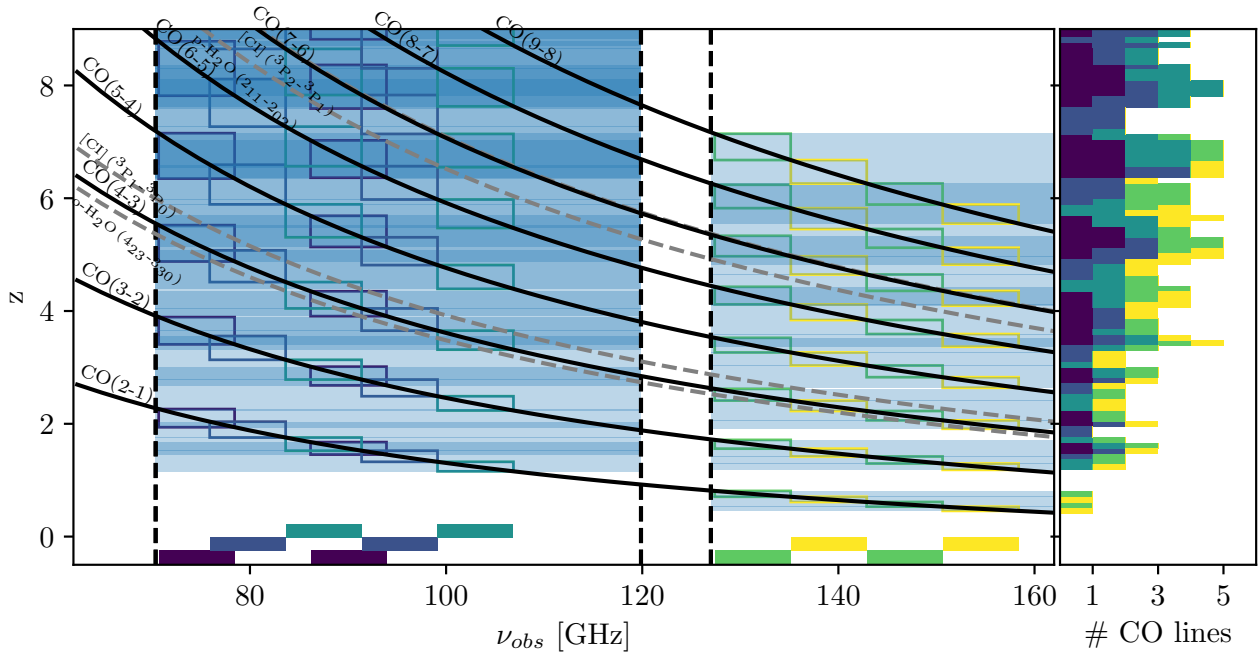


Fig. 3. Spectral coverage of all the molecular and atomic lines searched for in the z -GAL survey as a function of redshift in the 3 and 2 mm atmospheric windows covering the frequency ranges 70.689–106.868 GHz and 127.377–158.380 GHz. The boxes at the bottom show the LSB and USB frequency settings used in the z -GAL Large Programme for the 3-mm band (blue and dark green, respectively) and the 2-mm band (light green and yellow, respectively); those that were selected for the additional observations in the 3-mm band are highlighted in indigo (the set-up at 1 mm used for HeLMS-17 is not displayed – see text for further details). The lines detected in the z -GAL survey include the following: the ^{12}CO emission lines (solid black) from $J = 2-1$ to $8-7$; the $[\text{CI}] (^3\text{P}_1-^3\text{P}_0)$ and $[\text{CI}] (^3\text{P}_2-^3\text{P}_1)$ fine-structure lines (492 and 809 GHz rest frequencies, respectively, shown as grey dotted lines); and two water transitions, namely, the ortho- H_2O ($4_{23}-3_{30}$) and para- H_2O ($2_{11}-2_{02}$) (rest frequencies at 448.001 and 752.033 GHz, respectively, also shown as grey dotted lines). To enhance readability, the right panel exclusively displays the ^{12}CO emission lines that can be detected for a given redshift in each of the frequency settings (identified with their corresponding colour). The 2 mm frequency windows were selected to optimally cover the range of spectroscopic redshifts based on the 3 mm observations. The dark blue zones identify the redshift ranges where at least two emission lines are detected at 3 or 2 mm, while the light blue zones indicate the redshift ranges where only one line is present. The wide frequency coverage enabled by NOEMA enables the detection of at least one emission line in each receiver band, except for a few small redshift gaps (further details are provided in Sect. 2.2).

The simultaneous fitting of all the detected lines of a particular source, utilising single or double Gaussian profiles, offers the advantage of improved S/N compared to fitting the lines individually (Appendix A.1). This approach benefits from the combined information on the line properties and thereby enhances the accuracy of the fitting process. Additionally, by assuming, for all the detected lines, common velocity widths and velocity separations for double Gaussian profiles, the simultaneous fitting method optimises the precision of the measurements. For the majority of the sources, the line widths have been estimated with $S/N \gtrsim 5$, with the exception of three sources, namely: HerBS-72, HerBS-124 E and HerBS-194 S (for which $3.5 < S/N < 4.5$). The errors on the derived redshift values, which are estimated from the differences in centroid positions of the fitted profiles, are typically of the order of a few 10^{-5} – 10^{-4} .

The spectroscopic redshifts and the extracted properties of the spectral lines are listed in Tables A.6 and A.7, together with the measurements of the sources of the z -GAL Pilot Programme from Neri et al. (2020) in Table A.8. The combined continuum and emission line images of all the z -GAL sources are displayed, together with the detected molecular and atomic emission lines, in the series of figures in the catalogue presented in Appendix A. Detailed comments on individual sources are given in Appendix A.2.

In addition to the emission lines, the continuum flux densities of the sources were extracted from up to ten available polarisation-averaged 7.744 GHz wide sidebands over the fre-

quency range from 76.3 to 206.5 GHz. All the z -GAL sources were detected in the continuum in at least four sidebands. The NOEMA continuum measurements are presented and discussed in detail in Paper II, including tables listing the continuum flux densities for the entire z -GAL sample³.

3. Results

The observations from the z -GAL project resulted in a large data set containing a wealth of information about the *Herschel*-selected sources. In addition to unambiguous spectroscopic redshifts measurements, the NOEMA data provide key information on the morphology and the multiplicity of the sources, and high-quality measurements of the properties of the emission lines (widths and line fluxes) as well as on the underlying dust continuum. In order to display the richness of the data and present in the clearest possible way an overview of the z -GAL survey, we built a catalogue (Appendix A) that shows, for each of the z -GAL sources, the combined continuum and emission lines' maps together with all the molecular and atomic emission lines

³ All z -GAL data, including line emission and continuum data products, are available from the z -GAL collaboration and can be viewed on the IRAM Large Programme Archive web page: <https://iram-institute.org/science-portal/proposals/lp/completed/a-noema-spectroscopic-redshift-survey-of-bright-herschel-galaxies/>

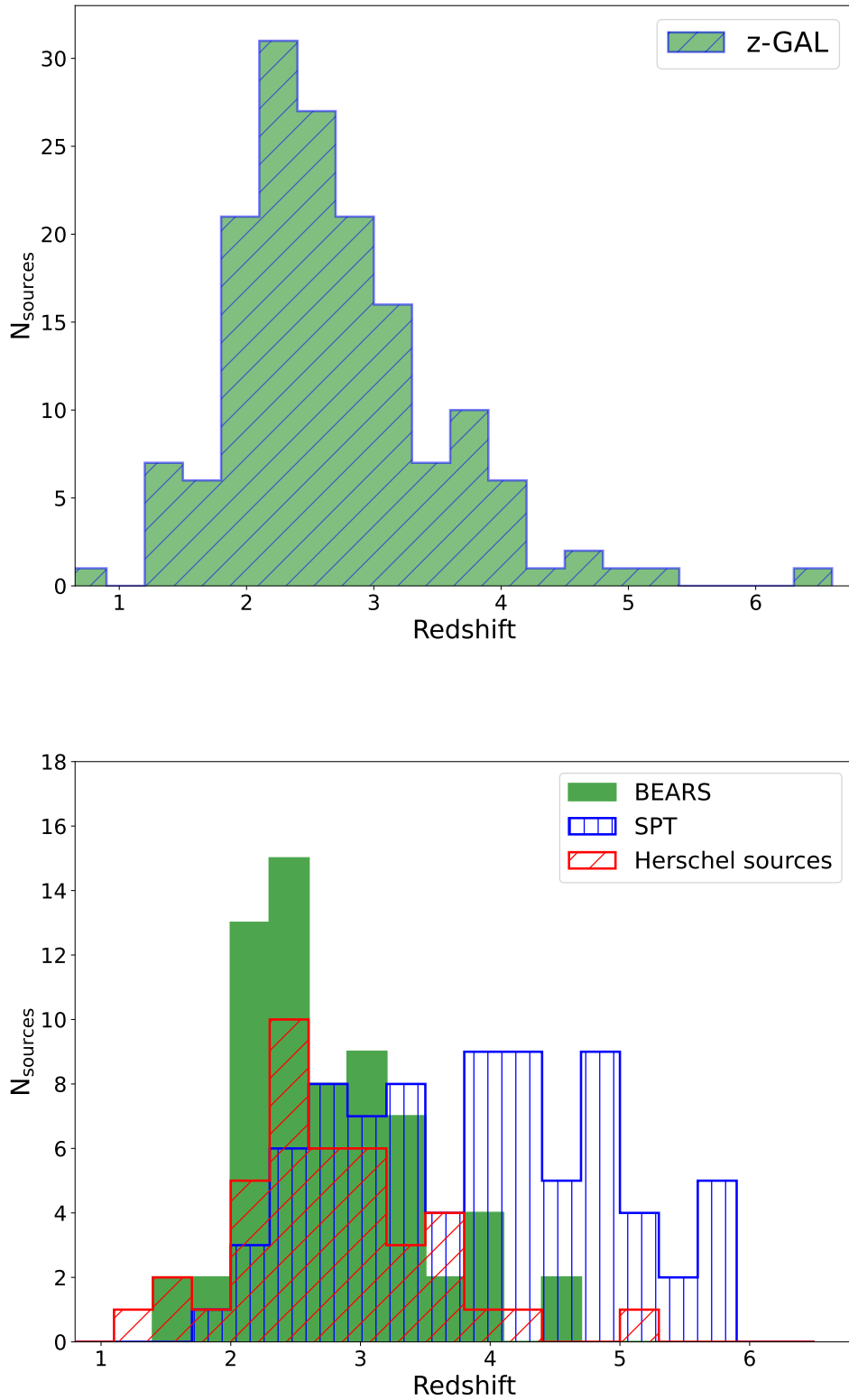


Fig. 4. Spectroscopic redshift distribution of the z -GAL survey compared to other high- z galaxy samples. Upper panel: Spectroscopic redshift distribution for the 165 individual sources identified in the fields of the 137 bright *Herschel*-selected galaxies of the z -GAL sample and the Pilot Programme (Neri et al. 2020). All the sources with robust spectroscopic redshifts were detected in at least two emission lines, mostly from CO (Tables A.6–A.8). The spectroscopic redshifts are in the range $0.8 < z < 6.5$ with a median value of $z = 2.56 \pm 0.10$. Bottom panel: Spectroscopic redshift distribution of the SPT sample (Reuter et al. 2020), the BEARS sample (Urquhart et al. 2022), and the *Herschel* sources for which z_{spec} values are listed in Bakx et al. (2018) and Nayyeri et al. (2016) with updates provided in Table A.5.

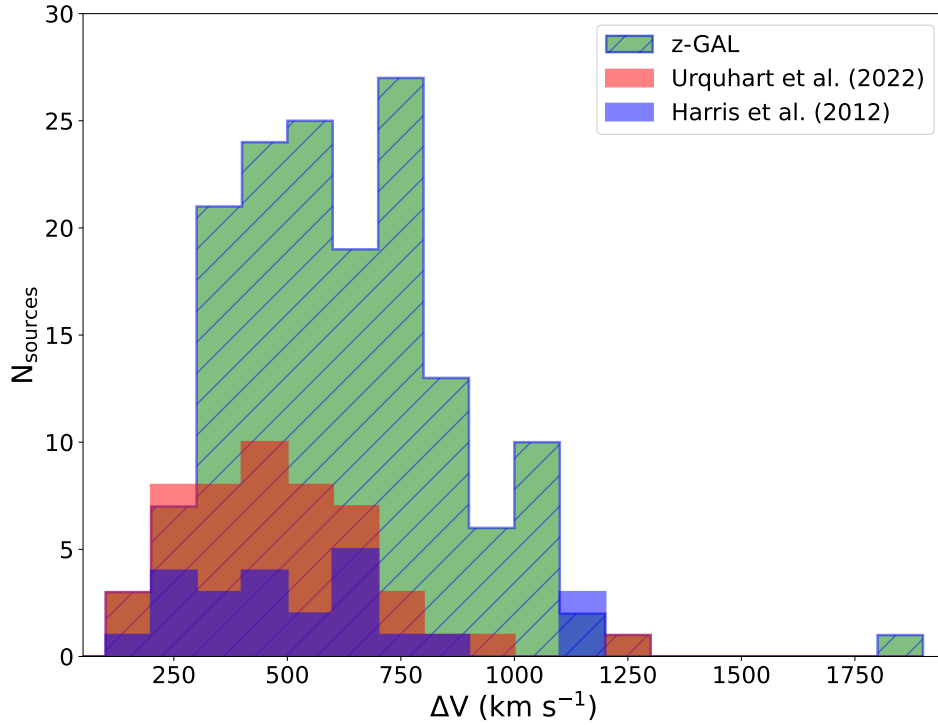


Fig. 5. Distribution of the FWHM for all the individual sources identified in the fields of the 137 bright *Herschel*-selected galaxies of the *z*-GAL sample and the Pilot Programme (Neri et al. 2020; Stanley et al. 2023) for which spectroscopic redshift were measured – see Tables A.6–A.8. The linewidths of the sources are on average large with a median value for the FWHM of $590 \pm 25 \text{ km s}^{-1}$ (Sect. 3.2). For comparison, the distributions of the FWHM for the BEARS sample (Urquhart et al. 2022) and the *Herschel*-selected galaxies reported in Harris et al. (2012, and references therein) are shown in the red and blue histograms, respectively. The median values of the FWHM are $465 \pm 37 \text{ km s}^{-1}$ for the BEARS sample and $485 \pm 73 \text{ km s}^{-1}$ for the sample of Harris et al. (2012) – further details are provided in Sect. 3.2.

that were detected in the HeLMS, HerS and HerBS sources (Figs. A.1–A.3, respectively). The continuum maps and emission lines detected in the sources from the Pilot Programme have been presented in Neri et al. (2020). The sources’ positions, the spectroscopic redshifts and emission lines properties are tabulated in Table A.6 for the HeLMS and HerS sources and in Tables A.7 and A.8 for the HerBS sources.

In this section, we summarise the main results on the spectroscopic redshifts (Sect. 3.1) and the line widths of the sources (Sect. 3.2) and present a comparison between the spectroscopic and photometric redshifts in Sect. 3.3. The morphology and general properties of the sources will be discussed in Sect. 4.

3.1. Spectroscopic redshifts

The main goal of the *z*-GAL project has been achieved by measuring, in the originally selected sample of 137 *Herschel*-selected sources, robust spectroscopic redshifts for 135 sources based on the detection of at least two emission lines in the 3 and 2 mm spectral bands. Taking into account the individual sources that were identified in some of the fields as double or multiple, as well as sources with different redshifts in the same field (e.g., HerBS-43), there are, in total, 165 individual sources with robust redshifts, including lensed galaxies, binary systems (such as HerBS-70 and HerBS-95) and clusters (e.g., HerBS-150) in the *z*-GAL sample (Sect. 4.1). The sources span the redshift range $0.8 < z < 6.5$, with a median of $z_{\text{median}} = 2.56 \pm 0.10$ and are centred around the peak of cosmic evolution (Fig. 4), which is consistent with the original selection of sources with $z_{\text{phot}} \gtrsim 2$. The majority of the *z*-GAL sources are at $z < 3.5$ (i.e. 86%) and only 21 sources are at $z > 3.5$.

As shown in the bottom panel of Fig. 4, the spectroscopic redshift distribution of the *z*-GAL sources is comparable to both the distributions of the *Herschel*-selected sample of southern H-ATLAS sources, the BEARS survey (Urquhart et al. 2022), and the *Herschel*-selected sources for which spectroscopic redshifts were measured in previous studies (as listed in Nayyeri et al. 2016; Bakx et al. 2018). This similitude reflects the fact that the wavelengths and the source selection criteria of these surveys are comparable to those of the *z*-GAL survey.

This is in contrast with the spectroscopic redshift distribution of the SPT-selected galaxies reported in Reuter et al. (2020), which is clearly different than all the samples based on *Herschel*-selected galaxies (Fig. 4), with the SPT sources peaking at significantly higher redshifts. The SPT-selected galaxies show a flat distribution between $z = 2.5$ and $z = 5.0$, with a major fraction (76%) of the sample at $z > 3$ and a median redshift of $z_{\text{median}} = 3.9 \pm 0.2$. The SPT galaxies were selected from a survey performed at a longer wavelength than the *Herschel* surveys and the difference in redshifts between the SPT and *Herschel*-selected galaxies is consistent with expectations for the selected wavelengths of these surveys (see, e.g., Reuter et al. 2020; Béthermin et al. 2015, and references therein). The systematic study of galaxies selected from the *Herschel* and SPT surveys thus offers an opportunity to gather critical complementary information on galaxy populations at different epochs of cosmic evolution, with *Herschel*-selected sources probing the peak of star formation activity around $2 < z < 3$, while the SPT-selected galaxies provide crucial information about the nature of star formation at earlier epochs.

For all but two (HerBS-82 and HerS-19) of the selected sources, we detected at least two emission lines, mostly from ^{12}CO

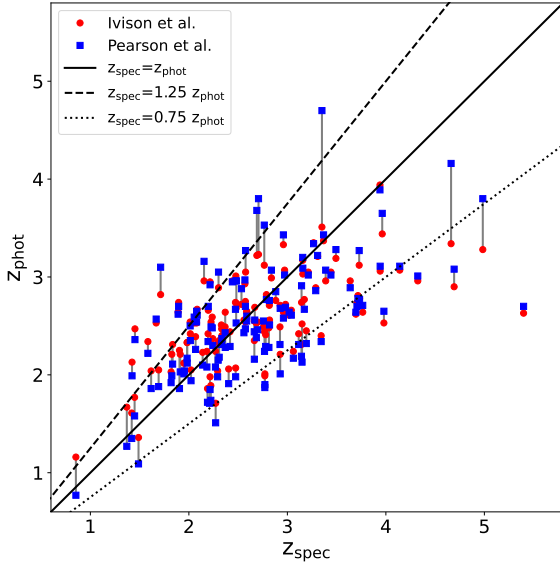


Fig. 6. Comparison of the spectroscopic redshifts of the sources in the z -GAL sample that are single and multiple (where each source has a similar redshift) with the photometric redshifts derived using the methods of Ivison et al. (2016) and Pearson et al. (2013) – blue and red symbols, respectively (as listed in Tables A.6–A.8 and A.1–A.3, respectively). For each source, the two photometric redshift values are linked by a grey vertical line positioned at the spectroscopic redshift value. The solid line shows where the photometric and spectroscopic redshifts are equal and the dashed lines indicate the $\pm 25\%$ deviations. The significant spread here displayed is consistent with previous findings (e.g., Ivison et al. 2016; Reuter et al. 2020; Urquhart et al. 2022) and demonstrates that photometric redshifts, based on *Herschel* and, when available, SCUBA-2 850 μm continuum flux densities, are only indicative.

ranging from the $J = 2-1$ to the $8-7$ transition and, in some case, from H_2O emission lines, [CI] and, in one case, HCO^+ and HCN (Tables A.6 and A.7). The majority of the sources, that is 85%, were detected in two transitions of ^{12}CO emission, only 8% were detected in three ^{12}CO emission lines and 21% in ^{12}CO and lines of other species (water or atomic carbon). In total, 40 sources were detected in $^{12}\text{CO}(2-1)$ and 109 sources in $^{12}\text{CO}(3-2)$. The CO emission lines are all relatively strong, resulting in high signal-to-noise ratios ($S/N > 5$) that provide the necessary quality to derive precise and reliable redshifts. In addition, significant information is available about the properties of the molecular gas, including the morphology and dynamics for sources that are resolved, and the physical conditions of the molecular gas (Paper III).

Based on the successful outcome of z -GAL, the following remarks can be made regarding the derivation of robust redshifts: (i) reliable spectroscopic redshift need most of the time the detection of at least two lines and one line is never enough; (ii) photometric redshifts are only indicative and sometimes ambiguous. Their values depend on the SED wavelength coverage and adopted template for the SED used in the analysis and are found to be, on average, within 20–30% of the z_{spec} values (Sect. 3.3); (iii) most emission lines detected are from ^{12}CO (mostly $J = 2-1$ to $5-4$), with some sources seen in H_2O and [CI] emission lines; for one source at $z = 0.85$, the two detected lines are $\text{HCN}(3-2)$ and $\text{HCO}^+(3-2)$; and (iv) some targets required a third line to lift any ambiguity, as different redshift solutions were compatible with the available photometry. This is mainly due to the degeneracy of the CO ladder, resulting in frequencies of the emission lines that are equally spaced. As an example, two emission lines were detected at 97.75 and 145.60 GHz in the source HerBS-108 (Table A.7). These lines could be identified either with $^{12}\text{CO}(2-1)$

and $^{12}\text{CO}(3-2)$ or with $^{12}\text{CO}(4-3)$ and $^{12}\text{CO}(6-5)$ indicating spectroscopic redshifts of $z_{\text{spec}} = 1.25$ or $z_{\text{spec}} = 3.71$, respectively. The photometric redshift for HerBS-108 ($z_{\text{phot}} = 2.4$; Table A.2) lies in between these two values. The measurement of a third line, detected at 73.3 GHz and identified with $^{12}\text{CO}(3-2)$, provided the required information to derive a robust spectroscopic redshift of $z_{\text{spec}} = 3.7166$ (Table A.7). It should be noted that, in many of those cases, the relative intensities of the CO emission lines can provide additional useful constraints, as some solutions seem to be incompatible with the expected excitation conditions in high- z dusty star-forming galaxies. Although, differential lensing could result in unusual CO line ratios (see, e.g., the case of HeLMS-19 described in Appendix A.2), for all the z -GAL sources where a third line was measured to disentangle between two spectroscopic redshift solutions, the CO line fluxes were always found to be compatible with an excitation akin to dusty starburst galaxies.

3.2. Line widths

The z -GAL sources display an exceptionally wide distribution of line widths, as shown in the histogram in Fig. 5, with a first ensemble of sources ($\sim 65\%$ of the sample) with a FWHM of $\Delta V < 700 \text{ km s}^{-1}$ and a second one (representing 35% of the sample) displaying broad lines, extending from 700 km s^{-1} to 1800 km s^{-1} for the most extreme case (HerBS-147 – see Appendix A.2). The mean value for the FWHM of the z -GAL sample is $590 \pm 25 \text{ km s}^{-1}$. Many of the broad line sources show profiles that are asymmetrical or double-peaked, with separation between the peaks of up to a few 100 km s^{-1} , indicative of merger systems and/or rotating disks (as illustrated in the catalogue in Appendix A).

Compared to other samples of high- z dusty star forming galaxies, the z -GAL sample stands out by the large fraction of sources with very large line widths. This is illustrated in Fig. 5 where the z -GAL sample is compared to the BEARS sample Urquhart et al. (2022) and the sources reported in Harris et al. (2012), that have mean values of the FWHM of $465 \pm 37 \text{ km s}^{-1}$ and $485 \pm 73 \text{ km s}^{-1}$, respectively, both distinctly smaller than the mean of the z -GAL sample. In the case of the *Herschel*-selected H-ATLAS sources in the BEARS sample, only ten sources display line widths in excess of 700 km s^{-1} (Fig. 5). The unlensed DSFGs from Bothwell et al. (2013) have a mean value of the FWHM of $510 \pm 80 \text{ km s}^{-1}$, with individual values ranging from 200 to 1400 km s^{-1} , very similar to the mean value of $540 \pm 40 \text{ km s}^{-1}$ derived for a sample of 45 high- z DSFGs presented by Birkin et al. (2021). For the SPT sources, the weighted-average FWHM is $370 \pm 130 \text{ km s}^{-1}$, a value that is comparable to the average of lensed DSFGs and Main Sequence $z < 2$ galaxies (Aravena et al. 2016, and references therein).

3.3. Comparison to photometric redshifts

As has been shown in previous studies, the spectroscopic redshifts of bright high- z dusty galaxies can significantly differ from the photometric redshifts based on the available measurements of the dust continuum flux densities (e.g., Neri et al. 2020; Reuter et al. 2020; Urquhart et al. 2022). Deriving redshifts using sub-millimetre SEDs of galaxies with known redshifts and dust temperatures as templates remains uncertain (e.g., Jin et al. 2019; Bendo et al. 2023). This is particularly true when using SPIRE data alone because the 250, 350, and 500 μm bands are close to the peak of the observed SED for $2 < z < 4$ galaxies. In order to homogenise the photometric redshifts of all three HerBS, HerMES and HeRS samples and show differences when using different analysis, we here adopted the methods

Table 1. Multiplicity of the z -GAL sources.

Number of sources detected within the NOEMA field Emission lines in fields with multiple sources	Nb. of fields	Source names
1	102	
2		
Emission lines for both sources (same z_{spec})	6	HeLMS-1 & 19; HerS-3; HerBS-61, 92 & 116
Emission lines for both sources (different z_{spec})	12	HeLMS-26, 40; HerS-18; HerBS-43 & 70 HerBS-76, 95, 124, 162, 194, 199 & 204
Emission lines for one source with second source only detected in the continuum	5	HeLMS-32, 35 & 43; HerBS-74 & 89
3		
Emission lines for all sources (different z_{spec})	5	HeLMS-17; HerBS-38, 109, 150 & 205
Emission lines for two sources (different z_{spec}) with the third source only detected in the continuum	4	HerS-19; HerBS-53, 91 & 187
Emission lines for one source with two other sources only detected in the continuum	3	HerS-7; HerBS-105 & 180

Notes. All the 137 NOEMA fields are centred on the coordinates of the *Herschel*-selected sources. The names of the *Herschel* sources are only given for the fields where two or three sources were detected. Detailed comments on these sources are given in Appendix A.2.

recommended by Ivison et al. (2016) and Pearson et al. (2013). The former method is based on functions fitted to the SED of the Cosmic Eyelash (J2135–0102) and two composite SEDs of multiple sub-millimetre galaxies; this method has been shown to perform better than other methods to estimate photometric redshifts (see discussion and references in Blain 1999; Bendo et al. 2023). The latter method is based on the sums of two blackbody SEDs (based only on *Herschel* flux densities) and derived from H-ATLAS sources with known redshifts. The z_{phot} values derived for the sources of z -GAL sample using both methods are listed in Tables A.1–A.3.

The z_{phot} values are on average consistent within 20–30% of the z_{spec} values, with deviations that become larger for redshifts $z > 4.5$; only for a few sources are the two values comparable within 10% (Fig. 6). This is in agreement with previous studies that have shown that the average uncertainty in a photometric redshift estimate is $\Delta(z)/(1+z) = 0.12$ for sources for which prior information indicates that they are at $z > 1$ (Pearson et al. 2013). The poor accuracy and reliability of redshifts derived from sub-millimetre continuum photometry is due to the degeneracy between temperature, the spectral index β , and redshifts (e.g., Bendo et al. 2023). In addition, the presence of multiple sources in the field, particularly when the spectroscopic redshifts of the sources are different, introduces another level of inaccuracy in the photometric redshift derived from the *Herschel* and SCUBA-2 flux densities (e.g., HerBS-109 or HerBS-53). The derived values of the redshifts based on continuum measurements alone remain therefore indicative and not precise enough to follow up efficiently with targeted observations of molecular or atomic gas.

4. Discussion

In this section, we describe general properties of the sample of high- z bright *Herschel*-selected galaxies observed in z -GAL. We first comment on the morphology of the sources and highlight some of their properties as a guide to the catalogue presented in Appendix A and then focus on global characteristics of the sample, based on relationships between the infrared luminosity, the CO luminosity and the line widths of the sources. Finally, we place the main findings of this overview in the context of the sample selection and will discuss the implications for probing the activity of DSFGs at the peak of cosmic evolution.

4.1. Morphology of the z -GAL sources

The main goal of this section is to highlight the diversity of the galaxies observed in the z -GAL project, comment on the multiplicity of the sources and derive first conclusions on the sources' properties, that will be further analysed in Bakx et al. (in prep.) based on multi-wavelength ancillary data and a more in-depth exploitation of the NOEMA data. Detailed comments on individual sources are given in Appendix A.2 and the combined continuum and emission line maps are shown in Figs. A.1–A.3 for the HeLMS, HerS and HerBS sources, respectively.

All the sub-millimetre bright galaxies in the z -GAL sample remained unresolved in the $17''$ – $35''$ beam of the *Herschel*-SPIRE and $13''$ beam of the SCUBA-2 observations. The angular resolution of NOEMA, between $1''2$ to $3''5$ at 2 mm, enabled us to measure with high precision the location of each of the *Herschel*-selected sources and, in some cases, to resolve them or to reveal multiple counterparts in the NOEMA field. A summary of the sources' morphology is provided in Table 1, which highlights the number of double and triple systems found in the z -GAL survey

In the majority of the cases (102 in total), the *Herschel* field contains only one source in the NOEMA data. In a few cases, unrelated sources are detected in the field of view of these high- z galaxies (Appendix A.2). For instance, about $15''$ to the east of HerBS-74, a galaxy at $z = 2.559$, a strong continuum emission source is detected that corresponds to a foreground galaxy seen in the HST near-infrared image (Borsato et al. 2023). In the case of HerBS-89a, a starburst at $z = 2.95$, a weak source is detected in the continuum about $6''$ to the east (Neri et al. 2020; Berta et al. 2021); no emission lines are seen in the NOEMA data and there is no corresponding source in the SDSS catalogue.

Some of these unresolved high- z galaxies are likely lensed (Sect. 4.2) and higher angular resolution follow-up observations are needed to confirm their precise morphology. Examples of such gravitationally lensed sources, that have been recently confirmed, include HerBS-89a and HerBS-154 from the Pilot Programme (Neri et al. 2020) where $0''3$ resolution follow-up observations at 1-mm with NOEMA revealed in these sources $1''0$ and $2''0$ diameter Einstein rings (Berta et al. 2021, and Butler et al., in prep.).

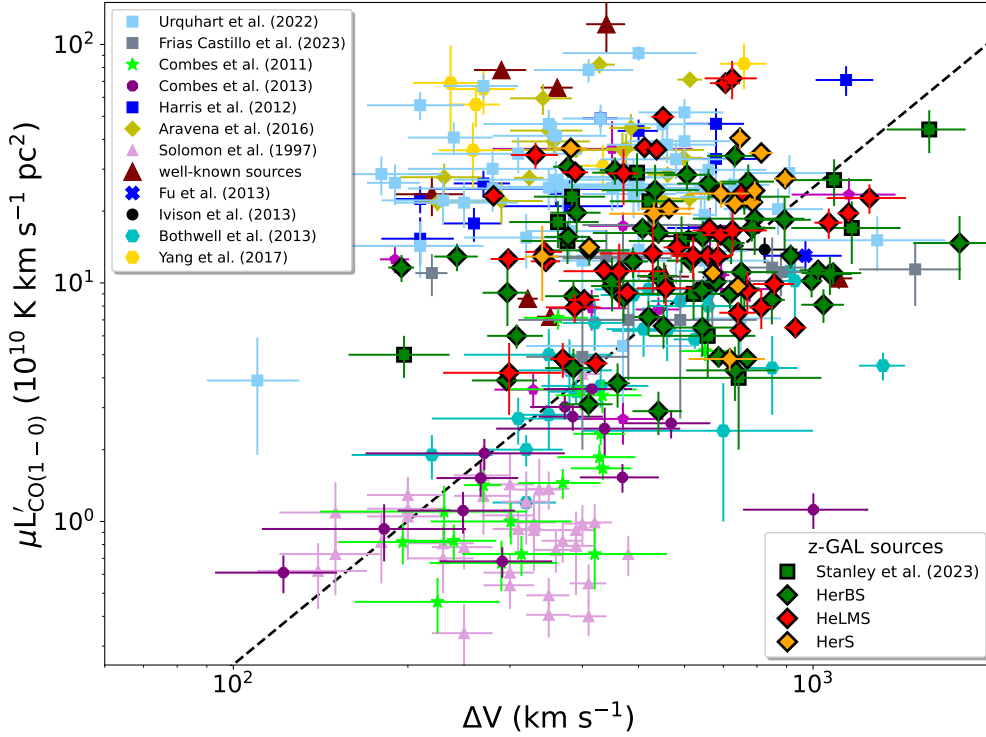


Fig. 7. ^{12}CO luminosity ($L'_{\text{CO}(1-0)}$) plotted against the linewidth (ΔV) of the ^{12}CO emission line for the z -GAL sources. The HerS, HeLMS and HerBS samples are identified with orange, red and green diamond shaped symbols, respectively, where the values for $L'_{\text{CO}(1-0)}$ are from Paper III. The sources of the Pilot Programme from Neri et al. (2020) are shown as green squares with the $^{12}\text{CO}(1-0)$ luminosity values from Stanley et al. (2023). The z -GAL sample is compared to high- z lensed and unlensed galaxy samples (Harris et al. 2012; Bothwell et al. 2013; Yang et al. 2017; Urquhart et al. 2022; Frias Castillo et al. 2023), individual galaxies including the two binary hyper-luminous systems H-ATLAS J084933 at $z = 2.41$ (Ivison et al. 2013) and HXMM01 at $z = 2.308$ (Fu et al. 2013), well-known high- z sources, both lensed (IRAS F10214, Eyelash, Cloverleaf, APM 08279 and the Cosmic Eyebrow) and unlensed (BR1202N and S, and BR1335), from Carilli & Walter (2013, and references therein) and Dannerbauer et al. (2019), and local ultra-luminous infrared galaxies (ULIRGs; Combes et al. 2011, 2013; Solomon et al. 1997) with corresponding symbols as indicated in the upper left panel. The sources of Harris et al. (2012) are the ones reported in Table 1 of that paper, except for H-ATLAS J084933 where we used the follow-up measurements of Ivison et al. (2013). No correction for amplification was applied to the ^{12}CO luminosities. With the exception of the sources of this paper and the sources selected from Yang et al. (2017), Bothwell et al. (2013) and Urquhart et al. (2022), all the other sources used in this plot have been measured in $^{12}\text{CO}(1-0)$ or in $^{12}\text{CO}(2-1)$. Corrections for excitation were applied for sources for which only ^{12}CO transitions higher than $J = 2-1$ are available, using the ratios listed in Carilli & Walter (2013). The dashed line shows the best-fitting power-law fit derived from the data for the unlensed DSFGs, $L'_{\text{CO}(1-0)} = 10^{5.4} \times \Delta V^2$ (e.g., Bothwell et al. 2013; Zavala et al. 2015).

A few z -GAL sources were resolved in the NOEMA observations. Remarkable cases include HeLMS-19 that displays two close components at the same redshift ($z_{\text{spec}} = 4.688$) tracing the arcs of a large Einstein ring seen in the HST near-infrared image (Borsato et al. 2023), and HerBS-204, a pair of interacting galaxies at $z_{\text{spec}} = 3.293$ linked by extended emission, which is reminiscent of the Antennae galaxy (NGC 4038/NGC 4039).

Other fields contain multiple sources separated by $5''$ to $10''$ with redshifts that differ slightly and separations that are too large to be consistent with gravitational amplification. These sources point to genuine binary or triple systems. Examples are the binary systems HerBS-70 and HerBS-95 at $z_{\text{spec}} = 2.31$ and $z_{\text{spec}} = 2.97$ (Neri et al. 2020, the former is part of a proto-cluster; Bakx et al., in prep.), and the triple systems HerBS-150 and HerBS-205 at $z_{\text{spec}} = 2.960$ and 3.67 , respectively.

In total, we identified 184 individual sources in the 137 fields centred on the *Herschel*-selected sources (Table 1), 169 of which are detected in emission lines. Of these sources, precise redshifts were measured for 165 of them; the sources for which only tentative redshifts were derived are HerBS-82, the west and south-east sources of HerS-19, and the source to the west of HerS-18, likely tracing a foreground galaxy (further details are provided in Appendix A.2).

4.2. General properties of the z -GAL sample

We present in this sub-section the global properties of the z -GAL sample and place the findings of this survey in the context of other studies of high- z DSFGs and local ultra-luminous infrared galaxies (ULIRGs). The two relationships that will be used are $L'_{\text{CO}(1-0)}$ vs. ΔV and L_{IR} vs. $L'_{\text{CO}(1-0)}$ (Figs. 7 and 8). Detailed analysis of these relationships are presented in Paper III for the latter to derive the depletion timescales and in Bakx et al. (in prep.) for the former to constrain the nature of the sources.

We compute the ^{12}CO luminosities of the sources (in $\text{K km s}^{-1} \text{pc}^2$) using the standard relation given by Solomon & Vanden Bout (2005),

$$L'_{\text{CO}} = 3.25 \times 10^7 S_{\text{CO}} \Delta V \nu_{\text{CO}}^{-2} D_L^2 (1+z)^{-1}, \quad (1)$$

where $S_{\text{CO}} \Delta V$ is the velocity-integrated ^{12}CO line flux in Jy km s^{-1} , ν_{CO} the rest frequency of the ^{12}CO emission line in GHz, and D_L the luminosity distance in Mpc. The $L'_{\text{CO}(1-0)}$ luminosities were derived using the lowest available $J \rightarrow (J-1)$ ^{12}CO transition for each z -GAL source after applying a correction for the excitation adopting the median brightness temperature ratios values listed in Carilli & Walter (2013), and applying similar corrections, where needed, for sources taken from the literature (see Berta et al. 2023 for the tabulated values). For the sources

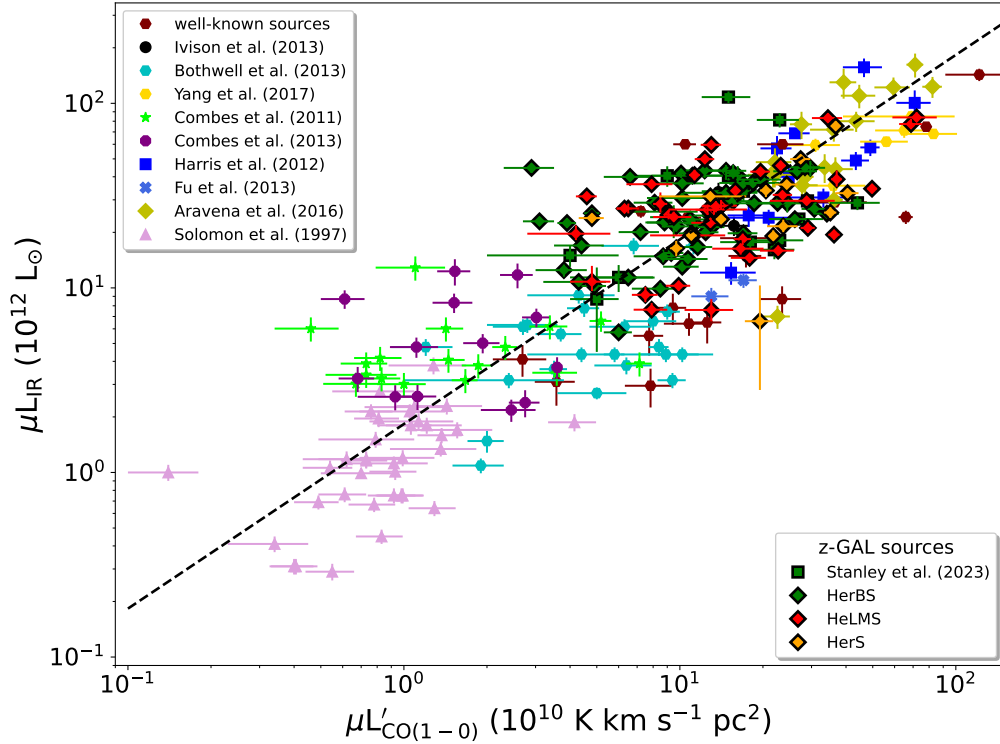


Fig. 8. Infrared luminosity (L_{IR} ; 8–1000 μm) as a function of $L'_{\text{CO}(1-0)}$ for the z -GAL sample, including the sources from the Pilot Programme (Neri et al. 2020; Stanley et al. 2023). The sources are identified using the same symbols as in Fig. 7. The infrared luminosities are derived from Paper II (Sect. 4.2). For comparison, we also plot high- z lensed and unlensed galaxies, as well as local ultra-luminous infrared galaxies (ULIRGs) from the literature (also shown in Fig. 7) with references and corresponding symbols given in the upper left panel. No correction for amplification was applied to the ^{12}CO and infrared luminosities. The dashed line shows the linear relationship between L_{IR} and $L'_{\text{CO}(1-0)}$ for a gas depletion timescale of $\tau_{\text{dep}} = 0.2$ Gyr (Paper III).

of the Pilot Programme, we used the $L'_{\text{CO}(1-0)}$ values and the line widths from Stanley et al. (2023). All the $L'_{\text{CO}(1-0)}$ values were recalculated for the cosmology adopted in this study.

The total infrared luminosities L_{IR} (8–1000 μm) of the z -GAL sources are based on the 50–1000 μm luminosities derived and listed in Paper II after applying a correction of 0.7 following Berta et al. (2013) – readers can refer to Paper III for more details.

4.2.1. The $L'_{\text{CO}(1-0)}$ – ΔV relationship

Figure 7 shows the relation between the apparent ^{12}CO luminosities, $L'_{\text{CO}(1-0)}$, and the FWHM (ΔV) of the ^{12}CO emission lines for all the z -GAL sources detected in $^{12}\text{CO}(1-0)$ up to (5–4) emission lines (corresponding to a total of 122 sources). The z -GAL sample is compared to high- z lensed and unlensed DSFGs, as well as local ULIRGs from the literature (see figure caption for details and references). With the exception of the sources selected from Urquhart et al. (2022), Yang et al. (2017) and Bothwell et al. (2013), all the other sources selected from the literature have been observed in $^{12}\text{CO}(1-0)$ or in $^{12}\text{CO}(2-1)$. In the case of the BEARS sample of Urquhart et al. (2022), we retained, to be consistent with the approach adopted for z -GAL, only those sources for which two emission lines were detected and for which unambiguous spectroscopic redshifts could be derived (50 sources in total).

This relationship, also known as the Tully-Fischer relationship, has first been based on observations of the 21 cm line of atomic hydrogen in nearby galaxies (Tully & Fisher 1977) and was proposed to be used for CO observations by Dickey & Kazes (1992) – see also Schöniger & Sofue (1994)

and references therein. The ^{12}CO line emission traces the kinematics of the potential well in which a galaxy’s molecular gas lies, and can therefore provide a measure of the dynamical mass of the galaxy, modulo any inclination or dispersion effects. The CO Tully-Fischer relationship therefore provides a rough estimate of the correlation between the gas mass and the galaxy dynamics. As was first pointed out by Harris et al. (2012), strong lensing boosts the apparent CO luminosity without affecting the line width, placing the gravitationally amplified sources above the roughly quadratic relationship between $L'_{\text{CO}(1-0)}$ and ΔV that is observed for the unlensed sources. Since then, this relationship has therefore been used in various studies of the molecular gas traced by CO in high- z DSFGs (e.g., Bothwell et al. 2013; Carilli & Walter 2013; Aravena et al. 2016; Yang et al. 2017; Neri et al. 2020; Birkin et al. 2021; Urquhart et al. 2022; Stanley et al. 2023). However, the usefulness of this effect for measuring the lensing magnification has been debated in the literature, and the consensus is that deriving exact values for the magnification factors, based on the Tully-Fischer relationship and using unresolved CO data, remains unreliable (e.g., Aravena et al. 2016; Stanley et al. 2023). Notwithstanding this caution, the separation between strongly lensed and unlensed (or weakly lensed) sources is clearly present in the $L'_{\text{CO}(1-0)}$ versus ΔV relationship including high- z galaxies and, in principle, could be used to distinguish sources that are strongly lensed from intrinsically hyper-luminous sources.

The most obvious feature in the relationship shown in Fig. 7 is indeed the difference in position between the unlensed or weakly lensed galaxies that lie along the power-law fit derived from the data for the unlensed DSFGs (Bothwell et al. 2013;

Zavala et al. 2015) and the sources that are strongly lensed (see references in the figure caption). Interestingly, the majority of the z -GAL sources are located in the upper part of the relationship indicating that several of them are strongly lensed, similarly to what is observed for the sources from the BEARS sample. Many of the sources that are farthest away from the relationship, that is with the largest apparent $L'_{\text{CO}(1-0)}$ luminosities, have line widths $\lesssim 700 \text{ km s}^{-1}$, and include sources that are located amongst the most magnified high- z galaxies known to date. In contrast, the sources with the larger line widths (that are dominated by the z -GAL sample) tend to be distributed along or close to the quadratic relationship for the unlensed sources. These sources could correspond to hyper-luminous galaxies ($L_{\text{IR}} \gtrsim 10^{13} L_{\odot}$) akin to the binary sources H-ATLAS J084933 (Ivison et al. 2013) and HXMM01 (Fu et al. 2013). The number of these potential HyLIRGs is remarkable amongst the sample of z -GAL bright *Herschel*-selected galaxies and more detailed observations are required to explore in greater detail their nature and derive their intrinsic properties.

4.2.2. The $L_{\text{IR}}-L'_{\text{CO}(1-0)}$ relationship

Figure 8 shows the relationship between $L'_{\text{CO}(1-0)}$ and L_{IR} for the z -GAL sample, compared to other samples of high- z DSGFs and local ULIRGs from the literature (references are provided in the figure caption). The infrared luminosity is a measure of the on-going star formation activity, while the CO emission is a tracer of the molecular gas available to form stars. The $L_{\text{IR}}-L'_{\text{CO}(1-0)}$ relation is therefore an indication of the fraction of the molecular gas (M_{mol}) that is converted to new stars, measuring the star formation efficiency, and equivalent to the integrated Kennicutt-Schmidt relation (Kennicutt 1998). There is a tight linear correlation between L_{IR} and $L'_{\text{CO}(1-0)}$ with the z -GAL sources populating the high end of the infrared luminosities ($6 \times 10^{12} < L_{\text{IR}}/L_{\odot} < 8 \times 10^{13}$) and the region corresponding to the most massive galaxies ($4 < L'_{\text{CO}(1-0)}/10^{10} \text{ K km s}^{-1} \text{ pc}^2 < 80$), with $M_{\text{mol}} > 10^{11} M_{\odot}$, intermediate between the high- z strongly lensed galaxies (e.g., the SPT sample; Aravena et al. 2016) and the unlensed high- z and local ULIRGs (Bothwell et al. 2013; Combes et al. 2011; Solomon et al. 1997). It should be noted that this relationship will hold when correcting the lensed sources for the magnification, as the amplification will affect similarly the ^{12}CO and the dust continuum emission (provided that they are co-spatial) and shift the data points along the relationship. The gas depletion timescales for the z -GAL sample are estimated and discussed in Paper III and are shown to be in the range between 0.1 and 1.0 Gyr, populating the part of the main sequence of star formation and the locus of starburst outliers.

4.3. Selection effects

Based on the previous sections, we subsequently place the results in the context of the sources' selection criteria and highlight some of the main findings with respect to model predictions (e.g., Negrello et al. 2010).

The z -GAL sources were selected based on a flux density threshold at $500 \mu\text{m}$, using $S_{500 \mu\text{m}} > 100 \text{ mJy}$ for the HeLMS and HerS sources (Nayyeri et al. 2016) and $S_{500 \mu\text{m}} > 80 \text{ mJy}$ for the HerBS sources (Bakx et al. 2018). This selection resulted in a sample containing amongst the brightest high- z galaxies, including potentially gravitationally lensed galaxies. Indeed, Negrello et al. (2010) demonstrated that strongly lensed DSGFs dominate over unlensed galaxies for $S_{500 \mu\text{m}} > 100 \text{ mJy}$, a flux density region where the count of unlensed DSGFs decreases steeply. Using a galaxy evolution model for the HerBS fields

specifically, Bakx et al. (2020b) showed that $\sim 70\%$ of the HerBS sources are predicted to be lensed.

From the z -GAL sample selection criteria, aimed at identifying infrared bright high- z galaxies that are potentially lensed (see, e.g., Negrello et al. 2017), the majority of the selected sources are expected to be single and unresolved in the NOEMA observations. The number of fields for which single sources were detected amounts to 102 (Table 1), representing 75% of the z -GAL sample, which is in line with this expectation. The results from the z -GAL survey further support the predictions from Negrello et al. (2010) and Bakx et al. (2020b) by the large number of sources of the sample that are located well-above the power-law fit derived from the data for the unlensed DSGFs in the $L'_{\text{CO}(1-0)}$ vs. ΔV relationship. It is interesting to note that this effect is more pronounced for the sources selected from the HeLMS and HerS samples than those from the HerBS sample, with about half of the HerBS sources distributed along or below the linear relation (Fig. 7). In addition, there are more HerBS sources with line widths in excess of 700 km s^{-1} (33 in total) than for the HeLMS and HerS sources (22 in total), with nearly $4\times$ as much HerBS sources with $\Delta V > 900 \text{ km s}^{-1}$ than HeLMS and HerS sources (15 and 4, respectively). As noted in Aravena et al. (2016), unlensed DSGFs tend to display larger linewidths and these differences could indicate that the sources selected from the HerBS sample contain a significant number of potential hyper-luminous galaxies.

Another notable difference between the HeLMS and HerS samples and the HerBS sample is the number of NOEMA fields centred on the *Herschel* positions where multiple sources are revealed. Not taking into account resolved lensed sources (e.g., HeLMS-19 and HerBS-91) and only retaining the binary and triple systems (with spectroscopic redshifts), there are 16 HerBS sources that turned out to be multiple systems, compared to only 4 HeLMS and 2 HerS sources (Table 1 and Appendix A.2).

Follow-up observations at higher angular resolution than the current NOEMA data are essential to fully explore the nature of the z -GAL sources, further probe their properties and assess their true nature. The z -GAL spectroscopic survey has allowed for significant progress in our understanding and classification of these bright *Herschel*-selected galaxies, centred on the peak epoch of cosmic galaxy formation, revealing lensed galaxies, multiple systems, interacting galaxies and potential HyLIRGs. The sources' characteristics extracted from the currently available data will hopefully serve as a foundation to study with better spatial resolution and sensitivity some of these exceptional DSGFs as well as critical aspects of galaxy evolution.

5. Conclusions

In this paper we introduce and present an overview of a Large Programme, z -GAL, using NOEMA to measure reliable redshifts for a sample of 126 bright *Herschel*-selected galaxies. From this original sample, spectroscopic redshifts were measured for 124 sources. Together with the z -GAL Pilot Programme (Neri et al. 2020), where 11 sources had their spectroscopic redshifts measured, this amounts to 135 dusty, bright *Herschel*-selected galaxies with precise redshifts, making it the largest sample of high- z galaxies with unambiguous redshifts to date and more than doubling the number of *Herschel*-selected galaxies with robust spectroscopic redshifts.

The results and conclusions are summarised as follows:

- Precise spectroscopic redshifts (z_{spec}) were established for 135 of the *Herschel*-selected galaxies based on the detection of at least two emission lines, mostly from ^{12}CO ranging

from the $J = 2-1$ to the $8-7$ transition. In addition, for 28 sources, we also report the detection of the emission of the atomic carbon fine-structure line [CII] ($^3P_1-^3P_0$), and for eight sources, the detection of water emission lines (mostly in the para- H_2O ($2_{11}-2_{02}$) transition and for one source in the ortho- H_2O ($4_{23}-3_{30}$) transition). For one source, the redshift was derived based on the detection of the HCN($3-2$) and HCO^+ ($3-2$) emission lines. The derived spectroscopic redshifts are in the range $0.8 < z < 6.5$ with a median value of $z_{\text{spec}} = 2.56 \pm 0.10$ and a tail in the distribution to $z > 3.5$.

- The majority of the z -GAL targets are individual sources (102 in total), which are unresolved or barely resolved on scales of 15–25 kpc. In total, 35 fields display multiplicity, including double and triple sources with lensed systems (e.g., HeLMS-19), mergers (e.g., HerBS-204), protoclusters (e.g., HerBS-150 and HerS-19), and fields with sources at different redshifts (e.g., HerBS-38). Taking these sources into account, the total number of individual sources in the z -GAL sample amounts to 184, of which 165 have robust spectroscopic redshifts.
- The photometric redshifts (z_{phot}), which depend on the wavelength coverage and the adopted template for the SED, are only indicative and are, on average for the sources studied in this paper, within 20–30% of the z_{spec} values here reported.
- The linewidths of the sources are large, with a mean value for the CO FWHM of $590 \pm 25 \text{ km s}^{-1}$ and with $\sim 35\%$ of the sources having widths of $700 \text{ km s}^{-1} < \Delta V < 1800 \text{ km s}^{-1}$. Compared to other high- z dusty star-forming galaxies, the z -GAL sample stands out by the large fraction of sources with large line widths. About half of the sources display double-peaked profiles indicative of merger systems and/or rotating disks.
- The wealth of information on the molecular and atomic emission lines in the z -GAL offers the possibility to analyse their physical properties in detail and study their evolution across cosmic time. This analysis is presented in [Paper III](#).
- In addition to the emission lines, the measurements of the underlying dust continuum at different frequencies across the 3 and 2-mm bands provides a solid foundation to explore the dust properties of the z -GAL sample and compare it to other available samples. This study is presented in [Paper II](#).
- Based on the location of the z -GAL sources in the $L'_{\text{CO}(1-0)}$ versus ΔV relationship, we conclude that a majority of the selected high- z galaxies are gravitationally amplified and that the sample includes HyLIRGs. Precise measurements of the amplification factors and the derivation of the properties of these sources will require higher resolution follow-up observations in the sub-millimetre and at near-infrared wavelengths to study the characteristics of the lensed high- z galaxies and the foreground amplifying galaxies.

The observations presented in this study have enabled one of the largest comprehensive measurement of redshifts of high- z dusty star-forming galaxies, providing a useful database for exploring in detail the properties of these sources and, using follow-up observations, the properties of the lensing systems in the case of gravitational amplification. This extended sample provides, together with other already available redshift measurements, a sizeable and homogeneous ensemble of ~ 300 bright *Herschel* selected galaxies with reliable redshifts, which will allow us to increase the number of known lensed galaxies at the peak of cosmic evolution, to provide a new sample of HyLIRGs, and to identify additional rare objects.

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References

- Aravena, M., Spilker, J. S., Bethermin, M., et al. 2016, *MNRAS*, 457, 4406
- Bakx, T. J. L. C., & Dannerbauer, H. 2022, *MNRAS*, 515, 678
- Bakx, T. J. L. C., Eales, S. A., Negrello, M., et al. 2018, *MNRAS*, 473, 1751
- Bakx, T. J. L. C., Dannerbauer, H., Frayer, D., et al. 2020a, *MNRAS*, 496, 2372
- Bakx, T. J. L. C., Eales, S., & Amvrosiadis, A. 2020b, *MNRAS*, 493, 4276
- Bakx, T. J. L. C., Eales, S. A., Negrello, M., et al. 2020c, *MNRAS*, 494, 10
- Bendo, G. J., Urquhart, S. A., Serjeant, S., et al. 2023, *MNRAS*, 522, 2995
- Berta, S., Lutz, D., Nordon, R., et al. 2013, *A&A*, 555, L8
- Berta, S., Young, A. J., Cox, P., et al. 2021, *A&A*, 646, A122
- Berta, S., Stanley, F., Ismail, D., et al. 2023, *A&A*, 678, A28 (Paper III)
- B ethermin, M., De Breuck, C., Sargent, M., & Daddi, E. 2015, *A&A*, 576, L9
- Birkin, J. E., Weiss, A., Wardlow, J. L., et al. 2021, *MNRAS*, 501, 3926
- Blain, A. W. 1999, *MNRAS*, 309, 955
- Blain, A. W., Smail, I., Ivison, R. J., Kneib, J.-P., & Frayer, D. T. 2002, *Phys. Rep.*, 369, 111
- Borsato, E., Marchetti, L., Negrello, M., et al. 2023, *MNRAS*, submitted
- Bothwell, M. S., Smail, I., Chapman, S. C., et al. 2013, *MNRAS*, 429, 3047
- Bussmann, R. S., P erez-Fournon, I., Amber, S., et al. 2013, *ApJ*, 779, 25
- Bussmann, R. S., Riechers, D., Fialkov, A., et al. 2015, *ApJ*, 812, 43
- Ca ameras, R., Nesvadba, N. P. H., Guery, D., et al. 2015, *A&A*, 581, A105
- Carilli, C. L., & Walter, F. 2013, *ARA&A*, 51, 105
- Casey, C. M., Berta, S., B ethermin, M., et al. 2012, *ApJ*, 761, 140
- Casey, C. M., Narayanan, D., & Cooray, A. 2014, *Phys. Rep.*, 541, 45
- Chakrabarti, S., Magnelli, B., McKee, C. F., et al. 2013, *ApJ*, 773, 113
- Chapman, S. C., Blain, A. W., Smail, I., & Ivison, R. J. 2005, *ApJ*, 622, 772
- Chen, C.-C., Liao, C.-L., Smail, I., et al. 2022, *ApJ*, 929, 159
- Combes, F., Garc a-Burillo, S., Braine, J., et al. 2011, *A&A*, 528, A124
- Combes, F., Garc a-Burillo, S., Braine, J., et al. 2013, *A&A*, 550, A41
- Cox, P., Krips, M., Neri, R., et al. 2011, *ApJ*, 740, 63
- Danielson, A. L. R., Swinbank, A. M., Smail, I., et al. 2017, *ApJ*, 840, 78
- Dannerbauer, H., Harrington, K., D az-S anchez, A., et al. 2019, *AJ*, 158, 34
- Dickey, J. M., & Kazes, I. 1992, *ApJ*, 393, 530
- Eales, S., Dunne, L., Clements, D., et al. 2010, *PASP*, 122, 499
- Elbaz, D., Dickinson, M., Hwang, H. S., et al. 2011, *A&A*, 533, A119
- Engel, H., Tacconi, L. J., Davies, R. I., et al. 2010, *ApJ*, 724, 233
- Frias Castillo, M., Hodge, J., Rybak, M., et al. 2023, *ApJ*, 945, 128
- Fu, H., Cooray, A., Feruglio, C., et al. 2013, *Nature*, 498, 338
- Fudamoto, Y., Ivison, R. J., Oteo, I., et al. 2017, *MNRAS*, 472, 2028

- Harrington, K. C., Yun, M. S., Cybulski, R., et al. 2016, *MNRAS*, 458, 4383
- Harrington, K. C., Weiss, A., Yun, M. S., et al. 2021, *ApJ*, 908, 95
- Harris, A. I., Baker, A. J., Frayer, D. T., et al. 2012, *ApJ*, 752, 152
- Hodge, J. A., & da Cunha, E. 2020, *Roy. Soc. Open Sci.*, 7, 200556
- Hughes, D. H., Serjeant, S., Dunlop, J., et al. 1998, *Nature*, 394, 241
- Ismail, D., Beelen, A., Buat, V., et al. 2023, *A&A*, 678, A27 (Paper II)
- Iverson, R. J., Swinbank, A. M., Swinyard, B., et al. 2010, *A&A*, 518, L35
- Iverson, R. J., Swinbank, A. M., Smail, I., et al. 2013, *ApJ*, 772, 137
- Iverson, R. J., Lewis, A. J. R., Weiss, A., et al. 2016, *ApJ*, 832, 78
- Iverson, R. J., Page, M. J., Cirasuolo, M., et al. 2019, *MNRAS*, 489, 427
- Jin, S., Daddi, E., Magdis, G. E., et al. 2019, *ApJ*, 887, 144
- Kennicutt, R. C., Jr. 1998, *ApJ*, 498, 541
- Lupu, R. E., Scott, K. S., Aguirre, J. E., et al. 2012, *ApJ*, 757, 135
- Lutz, D., Poglitsch, A., Altieri, B., et al. 2011, *A&A*, 532, A90
- Madau, P., & Dickinson, M. 2014, *ARA&A*, 52, 415
- Maresca, J., Dye, S., Amvrosiadis, A., et al. 2022, *MNRAS*, 512, 2426
- Narayanan, D., Turk, M., Feldmann, R., et al. 2015, *Nature*, 525, 496
- Nayyeri, H., Keele, M., Cooray, A., et al. 2016, *ApJ*, 823, 17
- Negrello, M., Hopwood, R., De Zotti, G., et al. 2010, *Science*, 330, 800
- Negrello, M., Amber, S., Amvrosiadis, A., et al. 2017, *MNRAS*, 465, 3558
- Neri, R., & Sánchez-Portal, M. 2022, *Eur. Phys. J. Web Conf.*, 265, 00002
- Neri, R., Cox, P., Omont, A., et al. 2020, *A&A*, 635, A7
- Oliver, S. J., Bock, J., Altieri, B., et al. 2012, *MNRAS*, 424, 1614
- Oteo, I., Iverson, R. J., Dunne, L., et al. 2016, *ApJ*, 827, 34
- Pearson, E. A., Eales, S., Dunne, L., et al. 2013, *MNRAS*, 435, 2753
- Planck Collaboration VI. 2020, *A&A*, 641, A6
- Planck Collaboration Int. XXVII. 2015, *A&A*, 582, A30
- Reuter, C., Vieira, J. D., Spilker, J. S., et al. 2020, *ApJ*, 902, 78
- Riechers, D. A. 2011, *ApJ*, 730, 108
- Riechers, D. A., Bradford, C. M., Clements, D. L., et al. 2013, *Nature*, 496, 329
- Riechers, D. A., Leung, T. K. D., Iverson, R. J., et al. 2017, *ApJ*, 850, 1
- Riechers, D. A., Cooray, A., Pérez-Fournon, I., & Neri, R. 2021, *ApJ*, 913, 141
- Schöniger, F., & Sofue, Y. 1994, *A&A*, 283, 21
- Smail, I., Iverson, R. J., & Blain, A. W. 1997, *ApJ*, 490, L5
- Solomon, P. M., & Vanden Bout, P. A. 2005, *ARA&A*, 43, 677
- Solomon, P. M., Downes, D., Radford, S. J. E., & Barrett, J. W. 1997, *ApJ*, 478, 144
- Spilker, J. S., Marrone, D. P., Aravena, M., et al. 2016, *ApJ*, 826, 112
- Stanley, F., Jones, B. M., Riechers, D. A., et al. 2023, *ApJ*, 945, 24
- Strandet, M. L., Weiß, A., Vieira, J. D., et al. 2016, *ApJ*, 822, 80
- Swinbank, A. M., Smail, I., Longmore, S., et al. 2010, *Nature*, 464, 733
- Tacconi, L. J., Genzel, R., Smail, I., et al. 2008, *ApJ*, 680, 246
- Tully, R. B., & Fisher, J. R. 1977, *A&A*, 54, 661
- Urquhart, S. A., Bendo, G. J., Serjeant, S., et al. 2022, *MNRAS*, 511, 3017
- Valiante, E., Smith, M. W. L., Eales, S., et al. 2016, *MNRAS*, 462, 3146
- van der Werf, P. P., Isaak, K. G., Meijerink, R., et al. 2010, *A&A*, 518, L42
- Vieira, J. D., Crawford, T. M., Switzer, E. R., et al. 2010, *ApJ*, 719, 763
- Walter, F., Decarli, R., Carilli, C., et al. 2012, *Nature*, 486, 233
- Walter, F., Carilli, C., Neeleman, M., et al. 2020, *ApJ*, 902, 111
- Wang, L., Gao, F., Best, P. N., et al. 2021, *A&A*, 648, A8
- Ward, B. A., Eales, S. A., Pons, E., et al. 2022, *MNRAS*, 510, 2261
- Wardlow, J. L., Cooray, A., De Bernardis, F., et al. 2013, *ApJ*, 762, 59
- Weiß, A., Kovács, A., Coppin, K., et al. 2009, *ApJ*, 707, 1201
- Weiß, A., De Breuck, C., Marrone, D. P., et al. 2013, *ApJ*, 767, 88
- Yang, C., Omont, A., Beelen, A., et al. 2017, *A&A*, 608, A144
- Zavala, J. A., Yun, M. S., Aretxaga, I., et al. 2015, *MNRAS*, 452, 1140
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Appendix A: Catalogue of the z -GAL sources

The following catalogue presents, for all the z -GAL sources, the combined continuum and emission line maps together with the molecular and atomic lines that were detected in each galaxy identified in the NOEMA field around the *Herschel* position. A series of tables lists the derived spectroscopic redshifts for each of the sources and the properties (widths and line fluxes) of all the emission lines that were detected.

The *Herschel* bright high- z galaxies that were selected for the z -GAL project are listed in Tables A.1, A.2 and A.3, that include the *Herschel* coordinates, the SPIRE and, when available, the SCUBA-2 flux densities, together with the estimated photometric redshifts. The sources without z_{spec} , that were not selected in (or dropped from) z -GAL, are presented in Table A.4 and the sources with previously available z_{spec} measurements are listed in Table A.5.

The main results of the z -GAL survey are given in Tables A.6, A.7 and A.8 that include, for all the sources identified in the NOEMA fields, the spectroscopic redshifts and the emission lines properties for the HeLMS and HerS sources, and the HerBS sources from the Large and Pilot Programmes, respectively.

A.1. Presentation of the catalogue

For all the sources, the continuum and line emission observed at 3 and 2 mm (including for one source, HeLMS-17, the emission lines detected at 1-mm) were combined to increase the signal-to-noise of the maps and reveal the morphology and the multiplicity of the galaxies detected in each of the fields around the position of the *Herschel*-selected source. Thus, as a first approximation, it was assumed that the line and the continuum emission originate from the same region.

The maps of the z -GAL sources were created using a multi-step procedure. First, pure continuum uv -tables were generated for each sideband and each source using their line-free frequency ranges. Every uv -table was then Fourier-transformed adapting natural weighting to produce a dirty map and cleaned to measure the continuum flux of each source in all sidebands. The spectral index of each source was then determined over the frequency range covered by the z -GAL observations, by fitting a single power-law function ($S_\nu \sim \nu^\alpha$) to the continuum fluxes with weighting factors taking into account the sensitivity achieved in each sideband. In a second step, pseudo-continuum uv -tables were generated by taking both the continuum and line emission into account. To preserve source sizes and maximise sensitivity, we proceeded as follow: (i) the uv -coordinates of the visibilities of each sideband were scaled to the highest frequency sideband by a factor ν_S/ν_S^H , where ν_S^H and ν_S were the respective central frequencies of the highest frequency sideband and of the one to scale; (ii) the real and imaginary parts of the visibilities of the individual uv -tables were scaled by a factor $(\nu_S/\nu_S^H)^{-\alpha}$; and (iii) the respective visibility weights were corrected by a factor $(\nu_S/\nu_S^H)^{2\alpha}$. By concatenating the scaled and re-weighted uv -tables of each sideband, a single uv -table was then created for each source and Fourier-transformed to generate a dirty map using natural weighting. Finally, the mapped emission regions were cleaned down to the $1\text{-}\sigma$ noise level. Given the compactness of the pseudo-continuum emission regions, no corrections were made for the primary beam attenuation in the various sidebands.

The basic strategy that was adopted to measure the redshifts of the z -GAL sources was to perform least squares Gaussian fits

to the identified spectral lines. The algorithm used derives the spectroscopic redshift of a source and its associated error by fitting simultaneously up to four spectral lines. This has the significant advantage that the redshift of a source is not measured independently on different spectral lines and that spectral profiles that are noisier can also be taken into account. All spectral profiles of the z -GAL sources were fitted with double-peaked Gaussians when the single Gaussian profile appeared unsuitable. To reduce the number of degrees of freedom, the algorithm assumes for each source the same velocity width, Δw , for the set of detected spectral lines. For a source with double-peaked profiles, the algorithm also assumes that the velocity spacing, Δs , between the fitted Gaussian profiles is the same for all detected lines. The line widths of the double-peaked profiles were estimated as $\Delta v = \Delta w + \Delta s$. This method enabled to derive with great precision the spectroscopic redshifts for each source, yielding typical errors of a few $\sim 10^{-4}$. The determination of the redshifts, line intensities, line widths, velocity separations, and their respective errors, was based solely on the rest frequencies of the lines and the signal-to-noise with which they were detected. The results of this analysis are presented in Tables A.6 and A.7.

The catalogue is organised following the order adopted in the tables of this paper with the HeLMS sources shown in Fig. A.1, the HerS sources in Fig. A.2 and the HerBS sources in Fig. A.3. The continuum maps and the emission lines detected in the sources from the Pilot Programme can be found in Neri et al. (2020), while their spectroscopic redshifts and line properties are listed in Table A.8. The source name and the derived spectroscopic redshift are indicated along the top of each series of panels displaying the emission lines of that source. The emission lines, detected in the 2 mm (top) and 3 mm (bottom) bands, are identified in the upper (left or right) corner of each spectrum; for the triple sources, the emission lines at 2 and 3 mm of the third source (e.g. HerBS-38) are shown in the bottom left and right panels, respectively. In the case of HeLMS-17, the ^{12}CO emission lines detected in the 1 mm and 2 mm bands are displayed in the top and bottom (left and right) panels, respectively, while the $\text{H}_2\text{O}(2_{11-2_{02}})$ emission line, which is detected at 1 mm in the Western source, is shown in the lower centre panel. The spectra are shown with the continuum subtracted and each emission line is centred at the zero velocity corresponding to its rest frequency. Fits to the emission line profiles are displayed as solid red lines. The contours of the combined continuum and emission line maps are plotted starting at 3σ . The synthesised beam is shown in the lower left corner of each map.

A.2. Comments on individual sources

In the following, we comment on individual sources highlighting their specificity and underlining some of the issues found in the fields where they are located. The sources here discussed are listed in their order of appearance in the catalogue.

- **HeLMS-1** - This is an extended source with a second peak $\sim 8''$ to the south that is resolved in the 2 mm dust continuum and the $^{12}\text{CO}(4-3)$ line emission. A detailed study of this source, which is amplified by a foreground lensing galaxy cluster, will be presented in Cox et al. (in prep.).
- **HeLMS-17** - To derive a robust spectroscopic redshift of this multiple source, additional observations were made at 1 mm, which enabled the detection of the $^{12}\text{CO}(6-5)$ and $\text{H}_2\text{O}(2_{11-2_{02}})$ emission lines. Together with the $^{12}\text{CO}(4-3)$

emission line detected at 2 mm, spectroscopic redshifts of $z_{\text{spec}}=2.2983$ for the single source to the East and of $z_{\text{spec}}=2.2972$ for the double source to the west were derived. It should be noted that the two components of the western source have the same redshift and that both are detected in the $^{12}\text{CO}(6-5)$ and the water emission lines. Observations performed at 3 mm covered the redshifted $^{12}\text{CO}(3-2)$ emission line (at ~ 104.8 GHz), although this line remained barely detected and deeper observations will be needed to extract the line properties of this transition in HeLMS-17.

- **HeLMS-19** - This source is resolved in two components to the east and west, separated by $\sim 5''$, that have been detected in the $^{12}\text{CO}(7-6)$, (5-4) and (4-3) emission lines, yielding spectroscopic redshifts that are nearly identical, namely $z_{\text{spec}}=4.6871$ and 4.6882 , respectively. The [CI](1-0) and [CI](2-1) emission lines shown in the catalogue have been integrated over the east and west components of this source to enhance the S/N of these relatively weak lines. The two components trace an Einstein ring that has been detected with the HST (Borsato et al. 2023). It is interesting to note that the line fluxes of the CO emission lines in HeLMS-19 have unusual ratios with the $^{12}\text{CO}(5-4)$ being weaker than both the $^{12}\text{CO}(3-2)$ and (7-6) emission lines (Table A.6). This could be a result of differential amplification and higher angular resolution observations are needed to further explore the spatial distributions of the CO emission lines and reveal if and how they differ.
- **HeLMS-26** - Two sources are detected to the east and the west in the field around the *Herschel* position. They have slightly different redshifts ($z_{\text{spec}}=2.6899$ and 2.6875 , respectively) and are separated by $\sim 12''$ corresponding to a linear distance of ~ 100 kpc.
- **HeLMS-32** - A second source, tentatively detected in the continuum (3σ), is seen $\sim 7''$ to the east of the $z_{\text{spec}}=1.7153$ source.
- **HeLMS-35** - A weak source, only detected in the continuum, is present $\sim 6''$ to the east of the main $z_{\text{spec}}=1.6684$ source.
- **HeLMS-40** - Two unresolved sources are revealed in the NOEMA field, HeLMS-40 W1 & W2, with slightly different redshifts of $z_{\text{spec}}=3.1445$ and 3.1395 , respectively, and offset to the west with respect to the *Herschel* position. The two sources are separated by $\sim 4''$, corresponding to a linear distance of ~ 30 kpc.
- **HeLMS-43** - East to the main source at $z_{\text{spec}}=2.2912$, there is a weaker source that is only detected in the continuum.
- **HeLMS-57** - This source displays an arc-like structure extending towards the south-west with a size of $\approx 6''$.
- **HerS-3** - HerS-3 is extended over $\sim 10''$ in the north-east and south-west direction with two distinct peaks detected in the $^{12}\text{CO}(3-2)$ and (5-4) transitions, yielding a similar redshift of $z_{\text{spec}}=3.0607$ for both components.
- **HerS-7** - This source, close in position to the *Herschel* coordinates, is detected in the continuum and the $^{12}\text{CO}(2-1)$ and $^{12}\text{CO}(4-3)$ emission lines, yielding a spectroscopic redshift of $z_{\text{spec}}=1.9838$. A second weaker component, $\sim 3.5''$ to the east, is revealed in the higher angular resolution 2 mm data; the detection of an emission line at the same frequency (154.5 GHz) as the main component indicates that both are at the same redshift. Two other sources in the field are only detected in the continuum emission, one $\sim 11''$ to the north and a weaker source $\sim 12''$ to the West.
- **HerS-8** - The galaxy in the field, at $z_{\text{spec}} = 2.2431$ based on the detection of the $^{12}\text{CO}(4-3)$ and (3-2) emission lines, is offset by $\sim 8''$ to the north with respect to the *Herschel* position.
- **HerS-9** - This source is the only one detected in the HCN(3-2) and $\text{HCO}^+(3-2)$ emission lines yielding an unambiguous spectroscopic redshift of $z_{\text{spec}} = 0.8530 \pm 0.0001$. HerS-9 coincides in position with a bright galaxy seen in the HST (Borsato et al. 2023). The SDSS DR12 spectroscopic redshift of this galaxy agrees precisely with the NOEMA derived value, namely $z_{\text{phot}} = 0.853$ (Nayyeri et al. 2016).
- **HerS-18** - The field around HerS-18 reveals two distinct sources to the east and the west. The source to the East is detected in the $^{12}\text{CO}(2-1)$ and (3-2) emission lines, yielding a robust redshift of $z_{\text{spec}}=1.6926$. The source to the west is only seen in one emission line that is detected at a frequency of 147.4 GHz and is therefore unrelated to the east component for which the $^{12}\text{CO}(2-1)$ and (3-2) emission lines are detected at frequencies of 182.4 and 85.6 GHz, respectively. In the field of HerS-18 there is a foreground galaxy seen in the SDSS at $z_{\text{phot}} = 0.495 \pm 0.075$ (Nayyeri et al. 2016). The identification of this emission line is therefore tentatively associated with the $^{12}\text{CO}(2-1)$ transition that would indicate a spectroscopic redshift of $z_{\text{spec}}=0.56$, in agreement with the value from the SDSS. However, additional observations are needed to confirm this result.
- **HerS-19** - The NOEMA observations reveal a triple, possibly quadruple system with an unresolved source to the west, an extended structure with a strong component in the south-east, a weak source in the north-east and a tentatively detected source in the centre of the field. In the frequency range covered by the 3 mm observations (75.894 to 83.638 GHz and 91.382 to 99.126 GHz), well defined emission lines are detected in the south-east and west components at 82.9 and 83.6 GHz; the latter line is found at the border of the bandwidth and, as a result, its profile is only partially covered. At 2 mm, HerS-19 was observed in the frequency ranges 127.377 to 135.135 GHz and 142.865 to 150.623 GHz. The continuum emission of the west and south-east components are dominating the spectra at 2 mm (with total flux densities of 2.1 and 1.5 mJy at 146 and 131 GHz, respectively - see Paper II). However, despite a detailed analysis of the available data, no emission lines were found at 2 mm that could unambiguously be identified with ^{12}CO transitions (or with other species) compatible in frequency with the emission lines detected at 82.9 and 83.6 GHz, thereby preventing us to derive a solid redshift solution for HerS-19. The lack of any clear emission line(s) in the frequency ranges that were covered from 91.382 to 150.623 GHz could be compatible with a redshift of $z \sim 3.16$, if the emission line detected around 83 GHz is identified with the CO(3-2) transition. However, additional observations are needed to definitely confirm

this redshift solution, which is close to the photometric redshift of HerS-19, ($z_{\text{phot}}=3.03$ –Table A.1), by covering the 2-mm frequency ranges 135.134-142.892 and 150.622-158.380 GHz, that would enable to detect the $^{12}\text{CO}(5-4)$ emission line (Fig. 3).

- **HerBS-38** - In the field of HerBS-38, there are three well-separated sources. The south-eastern and western sources have redshifts of $z_{\text{spec}}=2.4775$ and 2.4158 , respectively. These redshifts are based on the detection of the CO(3-2) emission line at 99.4 and 101.3 GHz and the [CI](1-0) and $^{12}\text{CO}(4-3)$ emission lines at 141.5 and 136.5 GHz in the south-east and west sources (with only half of the $^{12}\text{CO}(4-3)$ profile detected in the west source). The south-eastern and western sources are separated by $\sim 20''$, corresponding to a linear distance of ~ 160 kpc, with the difference in redshift indicating a comoving radial distance of ~ 70 Mpc. The third source to the north-east is detected in two emission lines at frequencies of 99.4 and 106 GHz, with no other line seen in the frequency ranges that were covered in the observations. Both lines have relatively low S/N. The derived spectroscopic redshift of $z_{\text{spec}}=6.5678$, based on the identification of the above lines with the $\text{H}_2\text{O}(2_{11}-2_{02})$ and CO(7-6) emission lines, appears to be the only solution and therefore the most likely. The detection of a third CO emission line in this source would, however, definitely confirm this result.
- **HerBS-53** - Two sources are revealed to the east and west in the NOEMA field with $^{12}\text{CO}(2-1)$ and (3-2) emission lines yielding spectroscopic redshifts of $z_{\text{spec}}=1.4219$ and 1.4236 , respectively. The distance of the sources is $\sim 5''$, corresponding to a projected linear distance of ~ 40 kpc. The closeness of their redshifts indicates that these sources are most likely related. To the north, another source is detected in the continuum only, with no emission lines found in the frequency ranges covered by the observations (i.e. 75.8-106.8 GHz and 127.3-158.3 GHz).
- **HerBS-61** - This source is resolved at 2 mm into two components separated by $\sim 2''-3''$ that are both at the same redshift ($z_{\text{spec}}=3.7293$). Only the combined spectra are shown to improve the S/N of the $^{12}\text{CO}(6-5)$, (4-3) and (3-2) emission lines and reveal the details of the line profiles. This spectroscopic redshift is in excellent agreement with the value of $z_{\text{spec}} = 3.7275$ derived by Bakx et al. (2020a).
- **HerBS-74** - The NOEMA field is dominated by the strong continuum emission of another source, $\sim 10''$ to the East, that coincides with a bright nearby galaxy seen in the HST (Borsato et al. 2023). No emission lines were detected with NOEMA in this foreground galaxy. The source at high redshift is centred near the *Herschel* coordinates. It is weak in the continuum and detected in the emission lines of $^{12}\text{CO}(4-3)$ and (3-2), yielding a secure spectroscopic redshift of $z_{\text{spec}} = 2.5610$.
- **HerBS-76** - This source is extended in the east-west direction. To the east, a main component displays strong, narrow ($\Delta(V) = 200 \pm 15 \text{ km s}^{-1}$) $^{12}\text{CO}(4-3)$ and (3-2) emission lines. The western component shows weaker $^{12}\text{CO}(4-3)$ and (3-2) emission lines with significantly larger line widths ($\Delta(V) = 742 \pm 149 \text{ km s}^{-1}$). The derived spectro-

scopic redshifts are slightly different with $z_{\text{spec}}=2.3302$ and 2.3319 for the eastern and western components, respectively.

- **HerBS-82** - This source is unresolved and shows a strong, double-peaked emission line centred at a frequency of 150.7 GHz with a line width of $181 \pm 15 \text{ km s}^{-1}$ and a line flux of $4.07 \pm 0.38 \text{ Jy km s}^{-1}$. A second broader ($1584 \pm 338 \text{ km s}^{-1}$) emission line (with a line flux of $1.43 \pm 0.38 \text{ Jy km s}^{-1}$ and a 20x times weaker peak flux density) is detected at 97.7 GHz. By identifying the 150.7 GHz emission line with the $^{12}\text{CO}(6-5)$ transition, the lower frequency emission could correspond to the $\text{H}_2\text{O}(4_{23}-3_{30})$ transition resulting in a spectroscopic redshift of $z_{\text{spec}}=3.5871$. However, the notable difference in the profiles of the two lines makes this identification doubtful; in addition, this solution is far off from the estimated photometric redshift of $z_{\text{phot}} \sim 2.49$ (Table A.7).

Another possibility is that the 150.7 GHz emission line traces the $^{12}\text{CO}(5-4)$ transition, yielding a spectroscopic redshift of $z_{\text{spec}} = 2.8277$. In this case, the [CI] emission line should have been detected at 128.5 GHz; however, nothing is seen at that frequency in the available NOEMA data. Therefore the identification with the $^{12}\text{CO}(4-3)$ is favoured, which would correspond to a spectroscopic redshift of $z_{\text{spec}} = 2.0583$ and would be compatible with the fact that no other line, with a similar profile as the 150.7 GHz emission line, is detected in the frequency ranges that were covered for this high- z source (Fig. 3). This result should be confirmed by measuring another emission line, for example of a lower CO transition.

Finally, the emission line detected at 97.7 GHz is likely unrelated to the *Herschel*-detected source. HerBS-82 is indeed centred on a bright, foreground lensing galaxy seen in the HST (Borsato et al. 2023) and the emission line could therefore trace the molecular gas of this nearby galaxy. Identifying this emission line with the $^{12}\text{CO}(1-0)$ transition would result in a spectroscopic redshift of $z_{\text{spec}} = 0.1788$ for the foreground lensing galaxy.

- **HerBS-91** - The combined continuum and emission lines map of this source reveals three components. The eastern and centre components are detected in the $^{12}\text{CO}(4-3)$ and (3-2) emission lines and in [CI(1-0)] (the catalogue shows the combined spectrum) yielding a similar redshift for both, namely $z_{\text{spec}}=2.4047$. The western source is only detected in the continuum.
- **HerBS-92** - The field of HerBS-92 reveals two main components that are separated by $\sim 15''$ and have the same redshift ($z_{\text{spec}} = 3.2644$). The weak $^{12}\text{CO}(5-4)$ and (3-2) emission lines have been integrated over both components in order to increase the S/N of the spectra. South of the *Herschel* position, a weaker source is detected in the continuum only.
- **HerBS-105** - The field of HerBS-105 reveals three components, two of which are only detected in the continuum. The bright component to the south-west is detected in the continuum and in the $^{12}\text{CO}(5-4)$ and (4-3) emission lines yielding a redshift of $z_{\text{spec}} = 2.6684$. The two other sources are detected to the north-east in the continuum: the strongest of the two is located $\sim 6''$ from the south-western component, while the weaker (4σ) component is $\sim 9''$ further to the east.

- **HerBS-109** - Three distinct components are detected in the field of this *Herschel* source. The north-western and southern components have comparable spectroscopic redshifts of $z_{\text{spec}}=1.5850$ and 1.5843 , respectively, based on the detection of the $^{12}\text{CO}(3-2)$ and $\text{CO}(2-1)$ emission lines. They are separated by $12''$ corresponding to a projected distance of ~ 100 kpc. The component to the north-east is unrelated to this binary system with a derived spectroscopic redshift of $z_{\text{spec}}=2.8385$ based on the detection of the $^{12}\text{CO}(5-4)$ and $(3-2)$ emission lines.
- **HerBS-116** - Two components to the east and west, separated by $\sim 4''$, are clearly identified at 2 mm but remain unresolved in the lower angular resolution 3 mm observations. The spectra of both components have therefore been combined. Two emission lines are detected in each component, corresponding to the $^{12}\text{CO}(5-4)$ and $(3-2)$ transitions, yielding a spectroscopic redshift of $z_{\text{spec}}=3.1573$ for this system.
- **HerBS-124** - Within the field of this source, there are two sources separated by $\sim 4''$ that are both detected in the continuum and the emission lines of $^{12}\text{CO}(4-3)$ and $\text{CO}(3-2)$, resulting in redshifts of $z_{\text{spec}} = 2.2781$ and 2.2772 for the eastern and western components, respectively. The eastern source is weak in both the continuum and line emission.
- **HerBS-127** - This source is detected in the continuum and in the emission lines of $^{12}\text{CO}(5-4)$ and $(3-2)$ and displays weak extended emission to the south. The spectroscopic redshift is $z_{\text{spec}}=3.1958$.
- **HerBS-136** - The field of HerBS-136 reveals one main component detected in the continuum and in the $^{12}\text{CO}(5-4)$ and $(3-2)$ emission lines, yielding a spectroscopic redshift of $z_{\text{spec}}=3.2884$. About $4''$ to the east, two weaker components are detected in the continuum, forming a $\sim 5''$ extended structure.
- **HerBS-147** - This galaxy, at $z_{\text{spec}}=3.1150$, displays the largest line width of the entire z -GAL sample with $\Delta V = 1789 \pm 300 \text{ km s}^{-1}$ for the $^{12}\text{CO}(5-4)$ and $(3-2)$ emission lines.
- **HerBS-150** - Three distinct sources are revealed in the east, centre and west in the combined continuum and emission line map of this field separated by $\sim 10''$ and $5''$. The eastern source is detected in the $^{12}\text{CO}(6-5)$ and $(4-3)$ emission lines, yielding a robust spectroscopic redshift of $z_{\text{spec}}=3.6732$. The centre and western components are each detected in only one line seen around the same frequency as the $\text{CO}(4-3)$ emission line in the eastern component ($98.6 < \nu_{\text{obs}} < 98.8$ GHz) so that the identification of this line with the $^{12}\text{CO}(4-3)$ is most likely. The derived spectroscopic redshifts are close to the eastern component with values of $z_{\text{spec}}=3.6787$ and 3.6682 , respectively. It should be noted that the emission line of the western source is only partially covered as it was detected at the edge of the bandwidth. Further observations will be needed to fully investigate the properties of this cluster of galaxies.
- **HerBS-162** - The combined continuum and emission line map shows a strong component (to the south-west) centred near the *Herschel* coordinates and a second weaker component to the north-east. Both components are detected in the $^{12}\text{CO}(4-3)$ and $(3-2)$ emission lines, yielding spectroscopic redshifts that are slightly different of $z_{\text{spec}}=2.4739$ and 2.4742 , respectively. The sources show distinct profiles with double peaked CO lines for the south-west component and single profiles for the north-east component, suggesting that these sources are distinct.
- **HerBS-177** - The spectroscopic redshift of $z_{\text{spec}} = 3.9625$ is in excellent agreement with the value of $z_{\text{spec}} = 3.9633$ that was derived for this source by [Bakx et al. \(2020a\)](#).
- **HerBS-180** - This field is dominated by a strong double peaked source to the north-east that is only detected in the continuum emission. No emission lines were found in this source in the frequency range that was covered (75.89-83.6 GHz, 91.38-99.12 GHz and 127.37-158.38 GHz). A weaker source to the south-west is detected in the $^{12}\text{CO}(3-2)$ and $(2-1)$ emission lines, resulting in a spectroscopic redshift of $z_{\text{spec}}=1.4527$.
- **HerBS-187** - Three sources are detected in this field. The eastern and western sources, which are separated by $\sim 4''$, are both detected in the $^{12}\text{CO}(2-1)$ and $(4-3)$ emission lines, yielding spectroscopic redshifts that are close to each other, namely $z_{\text{spec}}=1.8285$ and 1.8274 , respectively. The weaker southern source is only detected in the continuum and additional observations are needed to measure its redshift.
- **HerBS-194** - Two sources are detected in the field of HerBS-194 to the north and the south. They are separated by $\sim 15''$ and equidistant from the *Herschel*-source coordinates. Both sources are seen in the $^{12}\text{CO}(4-3)$ and $(3-2)$ emission lines and have spectroscopic redshifts that are close to each other, namely $z_{\text{spec}}=2.3335$ and 2.3316 , respectively. The projected distance between the sources is ~ 120 kpc.
- **HerBS-199** - Two sources are revealed in this field to the East (close to the *Herschel* position) and to the West with a separation of ~ 12 arcsec. They are both detected in the emission lines of $^{12}\text{CO}(4-3)$ and $^{12}\text{CO}(2-1)$ resulting in spectroscopic redshifts that are comparable, namely $z_{\text{spec}}=1.9248$ and 1.9197 , respectively. The projected distance between the sources is ~ 100 kpc.
- **HerBS-204** - This source is the only case in the z -GAL sample that shows two sources that are separated by $\sim 6''$ and are linked by a bridge of matter, with all components detected in the dust continuum and the emission lines of $^{12}\text{CO}(5-4)$ and $^{12}\text{CO}(4-3)$. The derived spectroscopic redshift is $z_{\text{spec}}=3.493$ and the projected separation between the two main components, and hence the extent of this interacting system is ~ 45 kpc.
- **HerBS-205** - The field of HerBS-205 reveals a triple system dominated by a source to the east that is resolved into two components (north-east and south-east) and another, weaker component $\sim 8''$ to the west. In all three sources, the $^{12}\text{CO}(3-2)$ and $^{12}(5-4)$ emission lines are clearly detected, yielding spectroscopic redshifts that are nearly identical with $z_{\text{spec}}=2.9600$ for the Eastern double source and $z_{\text{spec}}=2.9630$ for the western source, respectively. The projected distance between the Eastern and Western components is ~ 60 kpc.

Table A.1. z -GAL Sample: HeLMS and HerS sources.

Source name		Herschel Coordinates (J2000.0)		z_{phot}		Flux density Herschel		
Herschel Name	Other			(a)	(b)	$S_{250\mu\text{m}}$ (mJy)	$S_{350\mu\text{m}}$	$S_{500\mu\text{m}}$
HeLMS Sources								
HERMES J233441.0-065220	HeLMS-1	23:34:41.0	-06:52:20	2.25	1.86	431±6	381±7	272±7
HERMES J000215.9-012819	HeLMS-3	00:02:15.9	-01:28:29	1.61	1.35	643±7	510±6	258±7
HERMES J003929.6+002426	HeLMS-11	00:39:29.6	+00:24:26	2.73	2.96	140±7	157±7	154±8
HERMES J235601.5-071142	HeLMS-12	23:56:01.5	-07:11:42	2.49	2.43	178±7	184±6	154±7
HERMES J003619.8+002420	HeLMS-14	00:36:19.8	+00:24:20	2.04	1.86	251±6	247±6	148±7
HERMES J231857.2-053035	HeLMS-16	23:18:57.2	-05:30:35	2.71	2.76	143±7	183±7	146±8
HERMES J232558.3-044525	HeLMS-17	23:25:58.3	-04:45:25	2.33	2.15	190±6	189±6	142±8
HERMES J232210.3-033559	HeLMS-19	23:22:10.3	-03:35:59	2.90	3.08	114±6	160±7	134±8
HERMES J233728.8-045106	HeLMS-20	23:37:28.8	-04:51:06	2.42	2.34	162±6	178±7	132±8
HERMES J001800.1-060235	HeLMS-21	00:18:00.1	-06:02:35	1.99	1.87	206±6	186±7	130±7
HERMES J005841.2-011149	HeLMS-23	00:58:41.2	-01:11:49	1.36	1.09	391±7	273±6	126±8
HERMES J003814.1-002252	HeLMS-24	00:38:14.1	-00:22:52	3.28	3.80	82±6	120±6	126±7
HERMES J004124.0-010307	HeLMS-25	00:41:24.0	-01:03:07	2.23	2.10	178±6	186±7	125±8
HERMES J004747.1+061444	HeLMS-26	00:47:47.1	+06:14:44	3.22	3.68	85±7	119±6	125±8
HERMES J003758.0-010622	HeLMS-27	00:37:58.0	-01:06:22	2.64	2.71	125±7	144±6	124±8
HERMES J003009.1-020625	HeLMS-28	00:30:09.1	-02:06:25	2.72	2.88	114±6	135±6	122±7
HERMES J001027.1-024624	HeLMS-30	00:10:27.1	-02:46:24	2.03	1.92	185±6	170±6	121±7
HERMES J001353.5-060200	HeLMS-31	00:13:53.5	-06:02:00	2.12	2.01	178±7	176±6	120±7
HERMES J000336.9+014013	HeLMS-32	00:03:36.9	+01:40:13	2.82	3.10	103±6	112±6	119±7
HERMES J002719.5+001204	HeLMS-34	00:27:19.5	+00:12:04	1.71	1.51	248±6	20±7	116±8
HERMES J232500.1-005643	HeLMS-35	23:25:00.1	-00:56:43	2.53	2.57	122±6	132±7	114±8
HERMES J234314.0+012152	HeLMS-36	23:43:14.0	+01:21:52	2.53	2.65	115±6	115±6	113±8
HERMES J010801.8+053201	HeLMS-37	01:08:01.8	+05:32:01	2.49	2.55	122±6	120±6	113±7
HERMES J002208.1+034044	HeLMS-38	00:22:08.1	+03:40:44	1.86	1.72	190±6	157±6	113±7
HERMES J002936.3+020710	HeLMS-39	00:29:36.3	+02:07:10	3.12	3.53	81±6	107±6	112±7
HERMES J235332.0+031718	HeLMS-40	23:53:32.0	+03:17:18	2.74	2.91	102±6	123±7	111±7
HERMES J233633.5-032119	HeLMS-41	23:36:33.5	-03:21:19	2.40	2.37	130±6	131±6	110±7
HERMES J234014.6-070738	HeLMS-42	23:40:14.6	-07:07:38	2.12	2.04	158±6	154±6	110±8
HERMES J233420.4-003458	HeLMS-43	23:34:20.4	-00:34:58	2.04	1.98	156±7	141±5	109±8
HERMES J231447.5-045658	HeLMS-44	23:14:47.5	-04:56:58	1.67	1.27	220±8	141±7	106±8
HERMES J001226.9+020810	HeLMS-45	00:12:26.9	+02:08:10	2.63	2.70	107±6	142±6	106±7
HERMES J004622.3+073509	HeLMS-46	00:46:22.3	+07:35:09	3.05	3.27	82±9	113±9	105±10
HERMES J234951.6-030019	HeLMS-47	23:49:51.6	-03:00:19	1.89	1.74	186±7	167±6	105±8
HERMES J232833.6-031416	HeLMS-48	23:28:33.6	-03:14:16	3.51	4.70	49±6	104±6	105±8
HERMES J233721.9-064740	HeLMS-49	23:37:21.9	-06:47:40	1.98	1.85	173±6	161±7	105±8
HERMES J235101.7-024426	HeLMS-50	23:51:01.7	-02:44:26	2.53	2.59	112±6	124±6	105±7
HERMES J232617.5-025319	HeLMS-51	23:26:17.5	-02:53:19	2.96	3.16	86±6	109±6	104±7
HERMES J233727.1-002343	HeLMS-52	23:37:27.1	-00:23:43	1.86	1.71	182±6	157±6	104±8
HERMES J002718.1+023943	HeLMS-54	00:27:18.1	+02:39:43	3.23	3.80	69±6	87±6	103±7
HERMES J232831.8-004035	HeLMS-55	23:28:31.8	-00:40:35	2.76	2.85	95±7	120±6	102±7
HERMES J001325.7+042509	HeLMS-56	00:13:25.7	+04:25:09	2.96	3.07	89±6	98±6	102±7
HERMES J003519.7+072806	HeLMS-57	00:35:19.7	+07:28:06	2.33	2.17	134±6	135±7	101±8
HerS Sources								
HERS J012041.6-002705	HerS-2	01:20:41.6	-00:27:05	2.42	2.35	240±6	260±6	198±7
HERS J012751.1+004940	HerS-3	01:27:54.1	+00:49:40	2.24	2.17	253±6	250±6	191±7
HERS J012620.5+012950	HerS-5	01:26:20.5	+01:29:50	1.77	1.58	268±8	228±7	133±9
HERS J010133.8+003157	HerS-7	01:01:33.8	+00:31:57	2.33	2.12	165±7	154±6	122±7
HERS J010911.7-014830	HerS-8	01:09:38.9	-01:48:30	2.37	2.30	146±8	152±7	118±10
HERS J010911.7-011733	HerS-9	01:09:11.7	-01:17:33	1.16	0.77	393±8	220±8	118±9
HERS J011722.3+005624	HerS-10	01:17:22.3	+00:56:24	2.93	2.97	105±6	125±6	117±7
HERS J005847.3-010017	HerS-11	00:58:47.3	-01:00:17	3.34	4.16	63±8	116±7	115±9
HERS J012546.3-001143	HerS-12	01:25:46.3	-00:11:43	2.24	2.08	152±8	135±7	113±9
HERS J012521.0+011724	HerS-13	01:25:21.0	+01:17:24	2.07	1.98	165±8	153±7	114±10
HERS J014057.3-010547	HerS-14	01:40:57.3	-01:05:47	2.40	2.34	136±8	143±8	112±9
HERS J012106.9+003457	HerS-15	01:21:06.9	+00:34:57	2.89	3.05	94±6	130±7	110±7
HERS J021434.4+005923	HerS-16	02:14:34.4	+00:59:23	2.66	2.70	110±9	134±8	109±10
HERS J021402.6-004612	HerS-17	02:14:02.6	-00:46:12	2.61	2.64	110±8	130±8	105±9
HERS J013212.2+001754	HerS-18	01:32:12.2	+00:17:54	2.05	1.88	176±7	175±6	104±8
HERS J020529.1+000501	HerS-19	02:05:29.1	+00:05:01	2.83	3.03	89±6	112±6	102±8
HERS J010246.1+010543	HerS-20	01:02:46.1	+01:05:43	2.63	2.65	107±8	133±8	102±11

Notes. The source names and *Herschel* flux densities are from [Nayyeri et al. \(2016\)](#), adopting the names HeLMS-* and HerS-* for consistency with the names of the HerBS sources (Table A.2). The sources are listed in order of decreasing $S_{500\mu\text{m}}$ flux density. The photometric redshifts (z_{phot}) have been re-evaluated using the methods described by [Iverson et al. \(2016\)](#) and [Pearson et al. \(2013\)](#) listed as (a) and (b), respectively.

Table A.2. z -GAL Sample: HerBS sources.

Source name		Herschel Coordinates (J2000.0)		z_{phot}		Flux density			SCUBA-2
Herschel Name	Other			(a)	(b)	$S_{250\mu\text{m}}$	Herschel $S_{350\mu\text{m}}$	$S_{500\mu\text{m}}$	$S_{850\mu\text{m}}$
						(mJy)			
H-ATLAS 144608.6+021927	HerBS-38	14:46:08.6	+02:19:27	3.01	2.96	73.4±7.1	111.7±8.1	122.1±8.7	13.6±8.2
H-ATLAS 144556.1-004853	HerBS-46	14:45:56.1	-00:48:53	2.31	2.11	126.7±7.3	132.6±8.4	111.8±8.7	20.3±5.2
H-ATLAS 121301.5-004922	HerBS-48	12:13:01.6	-00:49:22	2.31	2.19	136.6±6.6	142.6±7.4	110.9±7.7	25.3±5.0
H-ATLAS 120319.1-011253	HerBS-50	12:03:19.1	-01:12:54	2.71	2.68	114.3±5.3	142.8±8.2	110.2±8.6	47.2±5.4
H-ATLAS 120709.2-014702	HerBS-51	12:07:09.2	-01:47:03	2.24	2.08	143.2±7.4	149.2±8.1	110.3±8.7	23.8±4.9
H-ATLAS 115112.2-012637	HerBS-53	11:51:12.3	-01:26:37	2.13	1.99	141.2±7.4	137.7±8.2	108.4±8.8	11.6±5.5
H-ATLAS 120127.6-014043	HerBS-61	12:01:27.7	-01:40:44	3.12	3.27	67.4±6.5	112.1±7.4	103.9±7.7	38.1±4.8
H-ATLAS 121542.7-005220	HerBS-62	12:15:42.6	-00:52:20	2.54	2.47	119.7±7.4	135.5±8.2	103.4±8.6	32.4±5.1
H-ATLAS 134422.6+231952	HerBS-65	13:44:22.6	+23:19:52	2.53	2.38	110.6±6.4	98.3±7.2	101.6±7.7	24.4±5.8
H-ATLAS 144512.1-001510	HerBS-72	14:45:12.1	-00:15:11	2.96	2.89	78.8±6.5	100.7±7.4	99.6±7.7	34.2±4.7
H-ATLAS 120600.7+003459	HerBS-74	12:06:00.7	+00:34:59	2.56	2.43	88.7±7.4	104.1±8.1	98.8±8.7	13.2±4.8
H-ATLAS 133534.1+341835	HerBS-76	13:35:34.1	+34:18:35	2.51	2.31	108.5±5.9	124.3±6.0	98.5±7.0	27.4±4.4
H-ATLAS 143352.4+020417	HerBS-78	14:33:52.4	+02:04:17	2.78	2.77	87.7±7.3	102.4±8.1	98.2±8.8	33.9±5.7
H-ATLAS 121144.8+010638	HerBS-82	12:11:44.8	+01:06:38	2.49	2.41	114.5±6.7	123.2±7.6	96.8±8.0	29.4±6.0
H-ATLAS 121812.8+011841	HerBS-83	12:18:12.8	+01:18:42	3.94	3.89	49.5±7.2	79.7±8.1	94.1±8.8	45.2±5.5
H-ATLAS 114752.7-005831	HerBS-85	11:47:52.8	-00:58:31	2.56	2.51	92.1±6.6	104.2±7.4	96.0±7.7	8.0±6.1
H-ATLAS 092135.4+000131	HerBS-91	09:21:35.7	+00:01:32	2.06	1.91	139.2±7.3	128.8±8.1	95.1±8.6	21.4±6.5
H-ATLAS 133808.9+255153	HerBS-92	13:38:08.9	+25:51:53	3.35	3.34	42.2±5.7	75.3±6.0	94.9±7.2	17.9±6.2
H-ATLAS 083932.2-011758	HerBS-105	08:39:32.2	-01:17:58	2.69	2.56	73.8±7.4	88.5±8.1	93.2±8.7	19.0±5.4
H-ATLAS 083817.4-004134	HerBS-108	08:38:17.4	-00:41:35	2.81	2.78	84.5±7.4	106.1±8.2	92.0±8.8	34.7±6.0
H-ATLAS 132900.4+281914	HerBS-109	13:29:00.4	+28:19:14	2.34	2.22	121.7±5.4	140.1±5.9	92.8±7.6	29.3±5.6
H-ATLAS 141832.9+010212	HerBS-110	14:18:33.0	+01:02:13	2.99	3.07	66.0±6.6	106.5±7.5	92.8±7.8	30.1±4.6
H-ATLAS 133538.3+265742	HerBS-115	13:35:38.3	+26:57:42	2.64	2.28	116.2±5	133.5±6.0	91.8±6.9	9.6±6.4
H-ATLAS 121348.0+010812	HerBS-116	12:13:48.0	+01:08:13	3.17	3.20	65.1±7.4	96.6±8.2	93.6±8.5	37.6±5.4
H-ATLAS 122158.5+003326	HerBS-124	12:21:58.5	+00:33:26	2.01	1.90	135.7±8.2	116.1±8.2	89.9±8.6	21±6
H-ATLAS 130432.2+295338	HerBS-125	13:04:32.2	+29:53:38	2.75	2.70	75.7±5.8	103.4±5.7	89.8±7.1	24.0±4.1
H-ATLAS 145135.2-011418	HerBS-126	14:51:35.3	-01:14:18	2.65	2.55	81.9±7.2	95.9±8.2	89.8±8.8	22.3±5.7
H-ATLAS 132128.6+282020	HerBS-127	13:21:28.6	+28:20:20	2.45	2.32	110.0±5.5	122.7±6.1	89.5±6.9	30.0±6.1
H-ATLAS 130414.6+303538	HerBS-128	13:04:14.6	+30:35:37	2.39	2.26	106.4±5.7	111.2±5.9	89.2±7.1	24.9±5.1
H-ATLAS 130053.8+260303	HerBS-129	13:00:53.8	+26:03:03	3.21	3.22	59.4±5.9	85.4±5.9	89.0±7.0	35.2±4.4
H-ATLAS 133440.4+353141	HerBS-134	13:34:40.4	+35:31:41	2.77	2.67	69.9±5.9	97.3±6.2	87.9±7.3	20.4±4.9
H-ATLAS 085308.5-005728	HerBS-136	08:53:08.6	-00:57:28	2.89	2.86	68.3±7.5	97.5±8.2	87.7±8.6	28.1±5.9
H-ATLAS 145337.2+000407	HerBS-137	14:53:37.2	+00:04:08	2.66	2.61	86.0±7.2	103.6±8.0	87.7±8.6	29.4±6.6
H-ATLAS 142140.3+000447	HerBS-140	14:21:40.4	+00:04:48	2.41	2.29	96.8±7.2	98.5±8.2	87.4±8.7	20.9±5.4
H-ATLAS 141810.0-003747	HerBS-143	14:18:10.1	-00:37:47	2.59	2.48	77.7±6.5	97.3±7.4	87.1±7.9	16.4±3.8
H-ATLAS 143403.5+000234	HerBS-147	14:34:03.6	+00:02:34	2.42	2.31	103.3±7.4	103.3±8.1	86.6±8.5	28.2±4.9
H-ATLAS 133827.6+313956	HerBS-149	13:38:27.6	+31:39:55	2.35	2.16	101.5±5.5	103.3±6.0	86.0±7.0	18.4±5.4
H-ATLAS 122459.1-005647	HerBS-150	12:24:59.0	-00:56:46	3.37	3.43	53.6±5	81.3±8.3	92.0±8.9	35.8±4.4
H-ATLAS 144243.4+015504	HerBS-153	14:42:43.4	+01:55:04	2.52	2.13	123.2±7.2	133.4±8.1	85.7±8.8	31.1±5.6
H-ATLAS 084957.7+010713	HerBS-157	08:49:57.7	+01:07:13	2.74	2.70	81.2±7.3	98.9±8.2	85.2±8.7	30.9±5.8
H-ATLAS 144334.3-003034	HerBS-162	14:43:34.3	-00:30:34	2.74	2.67	76.1±6.5	92.5±7.3	84.6±7.7	26.7±5.1
H-ATLAS 121416.3-013704	HerBS-164	12:14:16.3	-01:37:04	2.54	2.50	88.0±6.4	99.3±7.4	84.3±7.7	19.4±5.6
H-ATLAS 090613.8-010042	HerBS-165	09:06:13.8	-01:00:42	2.54	2.49	73.4±7.4	80.2±8.0	84.3±8.7	15.1±6.2
H-ATLAS 130341.5+313754	HerBS-167	13:03:41.5	+31:37:53	2.95	2.92	52.1±5.6	82.2±6.0	84.3±7.2	12.4±5.5
H-ATLAS 083859.3+021325	HerBS-169	08:38:59.3	+02:13:26	2.54	2.46	95.2±7.5	105.2±8.2	84.0±8.7	31.2±5.4
H-ATLAS 083945.0+021021	HerBS-171	08:39:45.1	+02:10:22	2.68	2.53	71.3±7.3	97.4±8.1	83.4±8.6	19.6±4.5
H-ATLAS 145040.5+003333	HerBS-172	14:50:40.5	+00:33:34	2.49	2.31	76.1±7.4	85.1±8.1	83.3±8.9	12.8±5.3
H-ATLAS 121900.8+003326	HerBS-175	12:19:00.8	+00:33:28	3.03	3.09	56.7±7.4	81.5±8.0	81.9±8.8	20.6±6.0
H-ATLAS 131222.2+270219	HerBS-176	13:12:22.2	+27:02:19	2.72	2.69	76.7±5.5	90.1±5.8	82.9±6.9	27.2±5.1
H-ATLAS 115433.6+005042	HerBS-177	11:54:33.6	+00:50:42	3.44	3.65	53.9±7.4	85.8±8.1	83.9±8.6	51.4±4.9
H-ATLAS 115521.0-021329	HerBS-179	11:55:21.0	-02:13:30	3.06	3.11	62.9±7.3	79.9±8.2	82.2±8.5	32.6±5.3
H-ATLAS 131539.2+292219	HerBS-180	13:15:39.2	+29:22:19	2.47	2.36	88.2±5.4	102.6±5.8	82.6±7.1	19.0±4.4
H-ATLAS 090453.2+022017	HerBS-183	09:04:53.3	+02:20:18	2.64	2.62	87.0±7.2	98.2±8.0	82.3±8.8	33.2±6.2
H-ATLAS 092408.8-005017	HerBS-185	09:24:08.8	-00:50:18	2.96	3.01	71.8±7.4	87.7±8.2	82.2±8.5	37.0±6.0
H-ATLAS 083705.2+020033	HerBS-187	08:37:05.2	+02:00:33	2.21	1.99	108.0±7.2	97.0±8.1	82.0±8.6	18.5±4.7
H-ATLAS 084259.9+024959	HerBS-188	08:42:59.9	+02:49:59	2.44	2.24	84.2±7.4	101.5±8.1	81.8±8.6	15.1±4.7
H-ATLAS 090405.3-003332	HerBS-190	09:04:05.3	-00:33:32	2.64	2.60	82.7±7.3	90.8±8.2	81.6±8.7	28.7±5.6
H-ATLAS 124753.3+322448	HerBS-191	12:47:53.3	+32:24:48	3.05	3.02	57.7±5.8	81.5±5.8	81.5±7.5	26.7±5.2
H-ATLAS 085352.0-000804	HerBS-193	08:53:52.1	-00:08:05	2.62	2.64	96.0±7.3	95.0±8.1	81.4±8.9	35.8±5.5
H-ATLAS 085521.1-003603	HerBS-194	08:55:21.2	-00:36:04	2.47	2.32	95.6±7.5	98.8±8.1	81.3±8.5	26.0±6.1
H-ATLAS 122034.2-003805	HerBS-197	12:20:34.1	-00:38:05	2.44	2.29	81.9±7.5	93.8±8.2	84.8±8.7	11.1±5.1
H-ATLAS 133352.2+334913	HerBS-199	13:33:52.2	+33:49:13	2.21	2.03	112.4±5.4	108.8±5.9	80.6±7.0	20.3±5.8
H-ATLAS 141117.8-010655	HerBS-201	14:11:17.9	-01:06:56	3.07	3.11	52.2±7.2	78.6±8.2	80.5±8.7	24.8±4.7
H-ATLAS 143328.4+020811	HerBS-202	14:33:28.4	+02:08:11	2.05	1.94	117.5±7.3	100.7±8.3	80.4±8.5	23.7±5.1
H-ATLAS 132909.5+300957	HerBS-204	13:29:09.5	+30:09:57	3.19	3.28	57.9±5.5	95.3±6.1	80.1±7.1	40.0±6.6
H-ATLAS 145132.7+024101	HerBS-205	14:51:32.8	+02:41:01	2.64	2.58	84.5±7.2	104.4±8.3	80.2±8.9	31.1±7.2
H-ATLAS 140421.7-001217	HerBS-206	14:04:21.8	-00:12:17	2.48	2.28	79.3±7.4	102.6±8.4	80.2±8.8	12.9±4.3

Notes. The source names and *Herschel* flux densities are from Bakx et al. (2018). The SCUBA-2 flux densities of the HerBS sources are the revised values from Bakx et al. (2020c). The sources are listed in order of decreasing $S_{500\mu\text{m}}$ flux density. For consistency with the HerS and HerMES sources, the photometric redshifts (z_{phot}) for the HerBS sources have been re-evaluated using the methods described by Ivison et al. (2016) and Pearson et al. (2013) listed as (a) and (b), respectively.

Table A.3. HerBS sources from the Pilot Programme (Neri et al. 2020).

Source name		<i>Herschel</i> Coordinates (J2000.0)		z_{phot}		Flux density			SCUBA-2
<i>Herschel</i> Name	Other			(a)	(b)	$S_{250\mu\text{m}}$	$S_{350\mu\text{m}}$	$S_{500\mu\text{m}}$	$S_{850\mu\text{m}}$
						(<i>Herschel</i>)			
						(mJy)			
H-ATLAS 133413.8+260458	HerBS-34	13:34:13.8	+26:04:58	2.55	2.44	136.1±5.4	161.0±5.5	126.5±6.8	34.0±5.7
H-ATLAS 132419.0+320752	HerBS-43	13:24:19.0	+32:07:52	3.05	3.02	84.4±4.9	116.0±5.2	115.4±6.3	37.0±5.1
H-ATLAS 133255.8+342208	HerBS-44	13:32:55.8	+34:22:08	2.21	2.01	164.3±5.8	186.8±5.8	114.9±7.2	25.3±4.5
H-ATLAS 131540.6+262322	HerBS-54	13:15:40.6	+26:23:22	2.96	2.95	94.0±5.7	116.1±6.1	108.6±7.1	44.7±4.6
H-ATLAS 130333.1+244643	HerBS-58	13:03:33.2	+24:46:43	2.67	2.61	99.0±5.5	111.5±5.9	104.5±7.1	30.5±5.0
H-ATLAS 130140.2+292918	HerBS-70	13:01:40.2	+29:29:18	2.35	2.18	119.6±5.8	136.8±5.8	100.0±7.1	21.9±5.5
H-ATLAS 131434.1+335219	HerBS-79	13:14:34.1	+33:52:19	2.53	2.54	103.4±6	115.3±6.0	97.9±7.3	28.5±5.0
H-ATLAS 131611.5+281219	HerBS-89	13:16:11.5	+28:12:19	3.33	3.43	71.8±5.7	103.4±5.7	95.7±7.0	52.8±4.3
H-ATLAS 134342.5+263919	HerBS-95	13:43:42.5	+26:39:19	3.07	3.02	61.9±5.7	101.3±5.7	94.7±7.6	27.4±6.2
H-ATLAS 131211.5+323837	HerBS-113	13:12:11.5	+32:38:37	2.82	2.77	80.7±5.9	103.4±6.0	92.0±7.0	32.0±5.2
H-ATLAS 132258.2+325050	HerBS-154	13:22:58.2	+32:50:50	2.74	2.70	79.1±5.6	87.9±5.9	85.6±7.2	28.8±4.2

Notes. The source names and *Herschel* flux densities are from Bakx et al. (2018). The SCUBA-2 flux densities of the HerBS sources are the revised values from Bakx et al. (2020c). The sources are listed in order of decreasing $S_{500\mu\text{m}}$ flux density. For consistency with the HerS and HerMES sources, the photometric redshifts (z_{phot}) for the HerBS sources have been re-evaluated using the methods described by Ivison et al. (2016) and Pearson et al. (2013) listed as (a) and (b), respectively.

Table A.4. Northern galactic plane and equatorial sources, without spectroscopic redshift measurements, not selected in (or dropped from) z -GAL from the HerBS, HeLMS, and HerS catalogues.

Source name		<i>Herschel</i> Coordinates (J2000.0)		Flux density			SCUBA-2		
<i>Herschel</i> Name	Other			$S_{250\mu\text{m}}$	$S_{350\mu\text{m}}$	$S_{500\mu\text{m}}$	$S_{850\mu\text{m}}$		
						(<i>Herschel</i>)			
						(mJy)			
HerBS sources									
H-ATLAS 141004.7+020306	HerBS-16	14:10:04.7	+02:03:06	119.4±7.3	151.0±8.4	176.0±8.7	86.0±4.7	Blazar	
H-ATLAS 091809.5+001929	HerBS-99	09:18:09.5	+00:19:29	93.2±7.4	116.6±8.2	94.3±8.7	23.9±5.0		
H-ATLAS 133108.4+303034	HerBS-112	13:31:08.4	+30:30:34	71.8±5.8	87.0±5.8	92.2±7.0	220.3±6.2	Blazar	
H-ATLAS 113833.8-014655	HerBS-119	11:38:33.8	-01:46:55	68.5±7.2	85.6±8.1	91.2±8.6	<i>12.5±5.5</i>		
H-ATLAS 142706.4+002258	HerBS-130	14:27:06.4	+00:22:58	119.4±7.3	118.7±8.1	88.8±8.6	15.5±5.1		
H-ATLAS 134441.5+240345	HerBS-133	13:44:41.5	+24:03:45	85.4±5.5	98.5±6.1	88.1±7.3	<i>17.4±5.9</i>		
H-ATLAS 134855.6+240745	HerBS-139	13:48:55.6	+24:07:45	76.9±5.9	82.9±5.9	87.4±6.8	<i>15.2±6.0</i>		
H-ATLAS 091454.0-010358	HerBS-142	09:14:54.0	-01:03:58	69.0±7.3	72.2±8.1	87.2±8.5	22.4±5.7		
H-ATLAS 133057.5+311734	HerBS-152	13:30:57.5	+31:17:34	47.7±5.6	53.4±6.0	85.8±6.9	14.1±5.0		
H-ATLAS 132329.9+311528	HerBS-158	13:23:29.9	+31:15:28	64.7±5.4	75.7±6.2	85.1±7.2	24.4±4.1		
H-ATLAS 122407.4-003247	HerBS-161	12:24:40.7	-00:32:47	56.5±7.3	75.7±8.1	82.4±8.8	<i>9.8± 4.3</i>		
H-ATLAS 145754.2+000018	HerBS-195	14:57:54.2	+00:00:18	70.3±7.3	92.7±8.1	81.0±8.8	17.3±5.3		
H-ATLAS 134403.1+242628	HerBS-196	13:44:03.1	+24:26:28	86.9±5.7	92.3±6.3	81.0±7.1	<i>13.9±6.3</i>		
H-ATLAS 141827.4-001703	HerBS-203	14:18:27.4	-00:17:03	117.2±6.5	116.4±7.4	80.2±7.6	15.0±4.8		
HeLMS and HerS sources									
HERMES 003032.1-021153	HeLMS-33	00:30:32.1	-02:11:53	81±7	98±6	118±8		Blazar	
HERMES 004532.6-000123	HeLMS-53	00:45:32.6	-00:01:23	48±7	85±6	103±8			

Notes. The source names and *Herschel* flux densities are from Bakx et al. (2018) and Nayyeri et al. (2016). The SCUBA-2 flux densities of the HerBS sources are the revised values from Bakx et al. (2020c). The SCUBA-2 flux densities indicated in italics are classified as non-detections, as discussed in Sect. 3 of Bakx et al. (2018). The sources are listed in order of decreasing $S_{500\mu}$ flux density. For the sources that are identified as blazars the reader is referred to Sect. 2.

Table A.5. Northern galactic plane and equatorial sources with previously available spectroscopic redshift measurements from the HerBS, HeLMS, and HerS catalogues.

Source name		Herschel Coordinates (J2000.0)		Flux density			SCUBA-2	z_{spec}
Herschel Name	Other			$S_{250\mu\text{m}}$	$S_{350\mu\text{m}}$	$S_{500\mu\text{m}}$	$S_{850\mu\text{m}}$	
(mJy)								
HerBS Sources								
H-ATLAS 134429.5+303034	HerBS-1	13:44:29.5	+30:30:34	461.9±5.8	465.7±6.5	343.3±7.1	90.3±5.3	2.30
H-ATLAS 114637.9-001132	HerBS-2	11:46:37.9	-00:11:32	316.0±6.6	357.9±7.4	291.8±7.7	98.2±5.4	3.26
H-ATLAS 132630.1+334408	HerBS-3	13:26:63.0	+33:44:08	190.5±5.6	281.3±5.9	278.6±7.5	73.8±5.0	2.95
H-ATLAS 083051.0+013225	HerBS-4	08:30:51.0	+01:32:25	248.5±7.5	305.3±8.1	269.1±8.7	98.7±5.4	3.63
H-ATLAS 125632.5+233627	HerBS-5	12:56:32.5	+23:36:27	209.3±5.6	288.5±6.0	264.0±7.0	97.6±5.0	3.56
H-ATLAS 132427.0+284450	HerBS-6	13:24:27.0	+28:44:50	342.3±5.6	371.0±5.9	250.9±6.9	38.1±5.6	1.68
H-ATLAS 132859.2+292327	HerBS-7	13:28:59.2	+29:23:27	268.4±4.4	296.3±4.8	248.9±5.9	77.4±6.5	2.78
H-ATLAS 084933.4+021442	HerBS-8	08:49:33.4	+02:14:42	216.7±7.5	248.5±8.2	208.6±8.6	36.3±5.8	2.41
H-ATLAS 125135.3+261458	HerBS-9	12:51:35.3	+26:14:58	157.9±5.9	202.2±6.0	206.8±6.9	83.5±5.5	3.68
H-ATLAS 113526.2-014606	HerBS-10	11:35:26.2	-01:46:06	278.8±7.4	282.9±8.2	204.0±8.6	73.2±5.3	3.13
H-ATLAS 133008.6+245900	HerBS-12	13:30:08.6	+24:59:00	271.2±5.4	278.2±5.9	203.5±6.9	67.7±6.4	3.11
H-ATLAS 142413.9+022303	HerBS-13	14:24:13.9	+02:23:03	112.2±7.3	182.2±8.2	193.3±8.5	89.7±5.4	4.24
H-ATLAS 141351.9-000026	HerBS-15	14:13:51.9	-00:00:26	188.6±7.4	217.0±8.1	176.4±8.7	44.0±4.9	2.48
H-ATLAS 090311.6+003907	HerBS-19	09:03:11.6	+00:39:07	133.2±7.4	186.1±8.2	165.2±8.8	87.7±6.3	3.04
H-ATLAS 132504.4+311534	HerBS-20	13:25:04.4	+31:15:34	240.6±5.4	226.6±6.0	164.9±7.3	26.2±4.9	1.84
H-ATLAS 133846.5+255055	HerBS-29	13:38:46.5	+25:50:55	159.0±5.8	183.1±6.0	137.6±7.5	36.9±5.8	2.34
H-ATLAS 132301.7+341649	HerBS-30	13:23:01.7	+34:16:49	124.1±5.6	144.5±6.0	137.0±7.2	48.2±4.7	2.19
H-ATLAS 125652.5+275900	HerBS-31	12:56:52.5	+27:59:00	133.9±5.8	164.1±6.0	131.8±7.4	41.5±4.9	2.79
H-ATLAS 091840.8+023048	HerBS-32	09:18:40.8	+02:30:48	125.7±7.2	150.7±8.2	128.4±8.7	32.9±5.7	2.58
H-ATLAS 133543.0+300402	HerBS-35	13:35:43.0	+30:04:02	136.6±5.4	145.7±5.8	125.0±6.9	35.8±4.7	2.68
H-ATLAS 125125.8+254930	HerBS-52	12:51:25.8	+25:49:30	57.4±5.8	96.8±5.9	109.4±7.2	39.4±5.8	3.44
H-ATLAS 091304.9-005344	HerBS-59	09:13:04.9	-00:53:44	118.2±6.4	136.8±7.4	104.3±7.7	42.0±5.4	2.63
H-ATLAS 130118.0+253708	HerBS-64	13:01:18.0	+25:37:08	60.2±4.8	101.1±5.3	101.5±6.4	54.3±4.6	4.04
H-ATLAS 115820.1-013752	HerBS-66	11:58:20.1	-01:37:52	119.8±6.8	123.7±7.7	101.5±7.9	25.8±4.2	2.19
H-ATLAS 113243.0-005108	HerBS-71	11:32:43.0	-00:51:08	67.8±7.3	105.8±8.2	99.8±8.8	<i>10.8±5.1</i>	2.58
H-ATLAS 083344.9+000109	HerBS-88	08:33:44.9	+00:01:09	71.0±7.6	96.0±8.1	95.9±8.8	16.3±5.8	<i>3.10</i>
H-ATLAS 113803.6-011737	HerBS-96	11:38:03.6	-01:17:37	85.1±7.3	98.4±8.2	94.8±8.8	11.3±5.1	<i>3.15</i>
H-ATLAS 113833.3+004909	HerBS-100	11:38:33.3	+00:49:09	96.8±7.3	106.4±8.1	93.4±8.7	6.3±5.9	2.22
HeLMS and HeRS Sources								
HERMES J233255.5-031134	HeLMS-2	23:32:55.4	-03:11:34	271±6	336±6	263±8		2.6895
HERMES J004410.2+011821	HeLMS-4	00:44:10.2	+01:18:21	113±7	177±6	209±8		4.1625
HERMES J234051.3-041937	HeLMS-5	23:40:51.3	-04:19:37	151±6	209±6	205±8		3.5027
HERMES J233620.7-060826	HeLMS-6	23:40:51.5	-04:19:38	151±6	209±6	205±8		3.4346
HERMES J232439.4-043934	HeLMS-7	23:24:39.4	-04:39:36	214±7	218±7	172±9		2.4726
HERMES J004714.1+032453	HeLMS-8	00:47:23.6	+01:57:51	398±6	320±6	164±8		2.2919
HERMES J004723.3+015749	HeLMS-9	00:47:23.6	+01:57:51	398±6	230±6	164±8		1.441
HERMES J005258.4+061319	HeLMS-10	00:52:58.6	+06:13:19	88±6	129±6	155±7		4.3726
HERMES J001615.8+032433	HeLMS-13	00:16:15.7	+03:24:35	195±6	221±6	149±7		2.765
HERMES J233255.7-053424	HeLMS-15	23:32:55.7	-05:34:26	148±6	187±6	147±9		2.4022
HERMES J005159.4+062240	HeLMS-18	00:51:59.5	+06:22:41	166±6	195±6	135±7		2.392
HERMES J001626.0+042613	HeLMS-22	00:16:26.0	+04:26:13	117±7	151±6	127±7		2.5093
HERMES J002220.9-015520	HeLMS-29	00:22:20.9	-01:55:20	66±6	102±6	121±7		5.1614
HERS J020941.1+001557	HerS-1	02:09:41.1	+00:15:57	826±7	912±7	718±8		2.5534
HERS 011640.1-000454	HerS-4	01:16:40.1	-00:04:54	137±7	196±7	190±8		3.791 ^a
HERS J010301.2-003300	HerS-6	01:03:01.2	-00:33:00	121±7	147±6	130±8		2.2153

Notes. The source names and *Herschel* flux densities are from Bakx et al. (2018) and Nayyeri et al. (2016). The SCUBA-2 flux densities of the HerBS sources are the revised values from Bakx et al. (2020c). The SCUBA-2 flux densities indicated in italics are classified as non-detections, as discussed in Sect. 3 of Bakx et al. (2018). The sources are listed in order of decreasing $S_{500\mu\text{m}}$ flux density. The references to the spectroscopic redshifts measurements are given in Bakx et al. (2018) and Nayyeri et al. (2016); for the HeLMS and HerS sources, the z_{spec} values have been updated for the sources included in Riechers et al. (2021) and Maresca et al. (2022) (and references therein); for HerBS-52 and HerBS-64, the spectroscopic redshifts are from Bakx et al. (2020a). Spectroscopic redshifts in italics are derived from on a single line detection and are therefore considered tentative. (a) The spectroscopic redshift of HerS-4 was measured with the LMT (A. Baker, priv. comm.).

Table A.6. Spectroscopic redshifts and emission line properties: HeLMS and HerS sources.

Source	Observed Position		z_{spec}	ΔV (km s^{-1})	Line	Line Flux (Jy km s^{-1})	Emission Lines			
	RA (J2000)	Dec.					Line	Line Flux (Jy km s^{-1})	Line	Line Flux (Jy km s^{-1})
HeLMS Sources										
HeLMS-1	23:34:40.97	-06:52:20.2	1.9047(1)	281±10	CO(2-1)	4.11±0.25	CO(4-3)	12.19±0.57		
HeLMS-3	00:02:15.96	-01:28:30.5	1.4202(1)	512±12*	CO(2-1)	11.24±0.86	CO(3-2)	20.72±1.55		
HeLMS-11	00:39:29.45	00:24:25.9	2.4829(1)	751±18*	CO(3-2)	4.71±0.41	CO(4-3)	6.81±0.55		
HeLMS-12	23:56:01.47	-07:11:43.1	2.3699(2)	551±25	CO(3-2)	10.70±0.75	CO(4-3)	17.81±1.15		
HeLMS-14	00:36:19.77	00:24:17.7	1.6169(1)	537±11*	CO(2-1)	8.62±0.52	CO(3-2)	13.99±0.86		
HeLMS-16	23:18:57.18	-05:30:34.7	2.8187(1)	1152±57*	CO(3-2)	5.31±0.76			[Cl](1-0)	6.47±1.38
HeLMS-17 E	23:25:58.56	-04:45:25.6	2.2983(1)	881±140	CO(4-3)	0.72±0.16	CO(6-5)	1.27±0.31		
HeLMS-17 W	23:25:57.92	-04:45:25.5	2.2972(1)	744±27	CO(4-3)	3.12±0.24	CO(6-5)	6.81±0.38	H ₂ O(2 ₁₁ -2 ₀₂)	1.80±0.19
HeLMS-19 E	23:22:10.20	-03:35:59.8	4.6871(1)	326±45	CO(4-3)	1.49±0.40	CO(5-4)	1.01±0.25	CO(7-6)	1.67±0.40
HeLMS-19 W	23:22:09.96	-03:36:01.4	4.6882(1)	250±24	CO(4-3)	1.76±0.38	CO(5-4)	1.01±0.25	CO(7-6)	2.71±0.45
HeLMS-19 EW	23:22:10.05	-03:36:00.8	4.6885(1)	332±31	CO(4-3)	3.25±0.40	CO(5-4)	1.76±0.16	CO(7-6)	4.38±0.42
HeLMS-19 EW							[Cl](1-0)	0.78±0.35	[Cl](2-1)	2.00±0.88
HeLMS-20	23:37:28.83	-04:51:06.3	2.1947(1)	298±20	CO(4-3)	3.83±0.85			[Cl](1-0)	2.05±0.49
HeLMS-21	00:18:00.21	-06:02:35.3	2.7710(2)	346±20	CO(3-2)	2.02±0.18	CO(5-4)	6.41±0.52		
HeLMS-23	00:58:41.09	-01:11:49.0	1.4888(1)	403±25	CO(2-1)	2.35±0.29	CO(3-2)	5.16±0.42		
HeLMS-24	00:38:14.00	-00:22:52.5	4.9841(1)	706±42	CO(4-3)	6.08±0.67	CO(5-4)	6.33±0.58	[Cl](1-0)	1.49±0.45
HeLMS-25	00:41:24.12	-01:03:07.8	2.1404(2)	932±39*	CO(2-1)	0.94±0.17	CO(4-3)	3.36±0.46		
HeLMS-26 E	00:47:47.56	06:14:39.4	2.6899(1)	328±53	CO(3-2)	0.35±0.14	CO(5-4)	1.65±0.44	[Cl](1-0)	1.19±0.36
HeLMS-26 W	00:47:46.76	06:14:45.3	2.6875(1)	815±159	CO(3-2)	0.89±0.30	CO(5-4)	2.82±0.75		
HeLMS-27	00:37:58.07	-01:06:20.1	3.7652(1)	658±67	CO(4-3)	1.61±0.28	CO(6-5)	5.63±0.78		
HeLMS-28	00:30:09.28	-02:06:25.1	2.5322(3)	749±27*	CO(3-2)	1.20±0.13			[Cl](1-0)	1.64±0.13
HeLMS-30	00:10:27.16	-02:46:26.4	1.8197(1)	662±15*	CO(2-1)	3.26±0.26	CO(4-3)	8.58±0.67		
HeLMS-31	00:13:53.53	-06:02:00.2	1.9494(1)	389±20*	CO(2-1)	4.96±0.48	CO(4-3)	10.46±0.99		
HeLMS-32	00:03:37.03	01:40:12.2	1.7149(3)	742±53*	CO(2-1)	1.63±0.38	CO(3-2)	1.63±0.43		
HeLMS-34	00:27:19.60	00:12:02.5	2.2714(1)	436±7*	CO(3-2)	2.61±0.15	CO(4-3)	4.10±0.22		
HeLMS-35	23:24:59.96	-00:56:44.1	1.6684(2)	857±70	CO(2-1)	2.25±0.37	CO(3-2)	7.91±0.86		
HeLMS-36	23:43:14.11	01:21:55.4	3.9802(1)	388±42	CO(4-3)	1.09±0.24	CO(5-4)	1.68±0.36	CO(6-5)	1.69±0.31
HeLMS-37	01:08:01.83	05:32:01.2	2.7576(2)	422±20*	CO(3-2)	0.78±0.13	CO(5-4)	2.00±0.31		
HeLMS-38	00:22:08.10	03:40:42.0	2.1898(1)	480±28*	CO(2-1)	1.95±0.28	CO(3-2)	4.27±0.66	CO(4-3)	4.35±0.68
HeLMS-39	00:29:36.02	02:07:13.1	2.7659(2)	478±20*	CO(3-2)	2.65±0.36	CO(5-4)	3.98±0.55		
HeLMS-40 W1	23:53:31.89	03:17:19.9	3.1445(1)	766±42*	CO(3-2)	1.26±0.25	CO(5-4)	3.03±0.42		
HeLMS-40 W2	23:53:31.58	03:17:21.0	3.1395(1)	681±51	CO(3-2)	1.11±0.30	CO(5-4)	1.86±0.40		
HeLMS-41	23:36:33.72	-03:21:20.2	2.3353(1)	557±61	CO(3-2)	2.09±0.35	CO(4-3)	3.22±0.50		
HeLMS-42	23:40:14.29	-07:07:36.5	1.9558(7)	463±92	CO(2-1)	1.90±0.59	CO(4-3)	2.61±0.72		
HeLMS-43	23:34:20.08	-00:34:58.1	2.2912(8)	529±101	CO(3-2)	3.07±0.99	CO(4-3)	2.17±0.57		
HeLMS-44	23:14:47.61	-04:56:56.0	1.3700(1)	385±16*	CO(2-1)	1.56±0.29	CO(3-2)	3.27±0.40		
HeLMS-45	00:12:27.07	02:08:06.4	5.3994(1)	725±75*	CO(5-4)	6.78±1.27	CO(8-7)	4.22±0.68	[Cl](1-0)	1.15±0.44
HeLMS-46	00:46:22.27	07:35:18.2	2.5765(2)	685±31*	CO(3-2)	2.42±0.37	CO(4-3)	3.52±0.56		
HeLMS-47	23:49:51.73	-03:00:17.5	2.2232(1)	694±46	CO(2-1)	2.15±0.24	CO(4-3)	6.84±0.68	[Cl](1-0)	2.08±0.59
HeLMS-48	23:28:33.52	-03:14:18.9	3.3514(1)	621±51	CO(3-2)	1.58±0.32	CO(4-3)	2.94±0.41	CO(5-4)	2.94±0.41
HeLMS-49	23:37:21.72	-06:47:42.0	2.2154(5)	531±111			H ₂ O(4 ₂₃ -3 ₃₀)	2.11±0.63	[Cl](1-0)	1.89±0.60
HeLMS-50	23:51:01.86	-02:44:23.5	2.0531(5)	1251±195	CO(3-2)	4.41±0.83	CO(4-3)	3.86±0.36		
HeLMS-51	23:26:17.61	-02:53:19.6	2.1567(5)	1064±59*	CO(4-3)	5.60±0.94			[Cl](1-0)	2.99±0.77
HeLMS-52	23:37:27.01	-00:23:41.1	2.2092(1)	344±25	CO(3-2)	3.23±0.37	CO(4-3)	4.26±0.43		
HeLMS-54	00:27:17.79	02:39:46.5	2.7070(8)	299±69	CO(3-2)	0.72±0.23			[Cl](1-0)	0.85±0.30
HeLMS-55	23:28:31.77	-00:40:36.6	2.2834(2)	471±29*	CO(3-2)	4.46±1.17	CO(4-3)	4.74±0.77		
HeLMS-56	00:13:25.82	04:25:07.3	3.3896(3)	814±38*	CO(3-2)	0.95±0.18	CO(4-3)	1.94±0.27		
HeLMS-57	00:35:19.56	07:28:03.7	1.9817(3)	762±106	CO(2-1)	2.75±0.59	CO(4-3)	5.15±1.01		
HerS Sources										
HerS-2	01:20:41.65	-00:27:05.8	2.0151(1)	748±23*	CO(2-1)	6.50±0.64	CO(4-3)	12.78±1.37		
HerS-3 NE	01:27:54.06	00:49:38.9	3.0608(1)	382±25	CO(3-2)	3.80±0.33	CO(5-4)	4.32±0.49		
HerS-3 SW	01:27:54.30	00:49:42.6	3.0607(3)	427±31	CO(3-2)	1.45±0.15	CO(5-4)	3.95±0.48		
HerS-5	01:26:20.60	01:29:49.9	1.4493(1)	742±22*	CO(2-1)	2.81±0.16	CO(3-2)	6.00±0.39		
HerS-7	01:01:33.69	00:31:53.7	1.9836(5)	693±83	CO(2-1)	3.91±0.86	CO(4-3)	5.68±0.91		
HerS-8	01:09:38.94	-01:48:22.9	2.2431(2)	718±108	CO(3-2)	1.14±0.30	CO(4-3)	2.75±0.56		
HerS-9	01:09:11.86	-01:17:33.4	0.8530(1)	134±26			HCN(3-2)	0.48±0.13	HCO ⁺ (3-2)	0.46±0.14
HerS-10	01:17:22.40	00:56:22.8	2.4688(1)	411±12*	CO(3-2)	2.85±0.24			[Cl](1-0)	1.61±0.35
HerS-11	00:58:47.20	-01:00:16.2	4.6618(3)	531±86	CO(4-3)	1.73±0.45	CO(7-6)	5.08±1.14		
HerS-12	01:25:46.18	-00:11:42.2	2.2706(3)	814±38*	CO(3-2)	5.17±0.78	CO(4-3)	7.45±1.27		
HerS-13	01:25:20.95	01:17:24.3	2.4759(4)	564±60	CO(3-2)	4.11±0.58			[Cl](1-0)	2.87±0.60
HerS-14	01:40:57.35	-01:05:46.8	3.3441(1)	895±37	CO(3-2)	3.33±0.50	CO(4-3)	6.17±0.71	CO(5-4)	8.32±0.50
HerS-15	01:21:06.83	00:34:55.8	2.3019(1)	672±19*	CO(3-2)	2.50±0.22	CO(4-3)	3.85±0.38		
HerS-16	02:14:34.48	00:59:23.9	2.1981(2)	786±27*	CO(4-3)	6.61±0.47			[Cl](1-0)	3.17±0.50
HerS-17	02:14:02.54	-00:46:11.0	3.0182(1)	793±18*	CO(3-2)	3.94±0.38	CO(5-4)	6.97±0.46		
HerS-18 E	01:32:12.31	00:17:55.6	1.6926(1)	733±36	CO(2-1)	4.72±0.45	CO(3-2)	7.16±0.47		
HerS-18 W	01:32:11.78	00:17:58.2	<i>0.5636(1)</i>	364±75	<i>CO(2-1)</i>	0.95±0.25				
HerS-19 SE	02:05:29.26	00:04:57.0	<i>3.1678(1)</i>	630±110	<i>CO(3-2)</i>	1.62±0.41				
HerS-19 W	02:05:28.82	00:05:01.2	<i>3.1394(1)</i>	662±190	<i>CO(3-2)</i>	1.22±0.51				
HerS-20	01:02:46.14	01:05:40.1	2.0792(1)	341±23	CO(2-1)	1.95±0.18	CO(4-3)	11.71±1.91		

Notes. The uncertainties in the spectroscopic redshifts, z_{spec} , given in brackets, correspond to the errors derived from the simultaneous Gaussian fits to the line profiles. Emission lines fitted with double Gaussians are highlighted with an asterisk after the value of ΔV . The redshifts and line identifications that are not robust are in italics (for HerS-18 W and HerS-19 SE & W). Further details about the line fitting procedure and the derivation of z_{spec} are provided in Appendix A.1.

Table A.7. Spectroscopic redshifts and emission line properties: HerBS sources.

Source	Observed Position		z_{spec}	ΔV (km s ⁻¹)	Line	Emission Lines				
	RA (J2000)	Dec.				Line Flux (Jy km s ⁻¹)	Line	Line Flux (Jy km s ⁻¹)	Line	Line Flux (Jy km s ⁻¹)
HerBS-38 SE	14:46:09.08	02:19:19.4	2.4775(1)	674±19*	CO(3-2)	2.19±0.22			[CI](1-0)	1.60±0.19
HerBS-38 W	14:46:08.33	02:19:29.8	2.4158(1)	873±67*	CO(3-2)	1.66±0.28	CO(4-3)	<1.16		
HerBS-38 NE	14:46:09.10	02:19:34.1	6.5678(1)	998±107			CO(7-6)	0.74±0.24	H ₂ O(2 ₁₁ -2 ₀₂)	0.82±0.22
HerBS-46	14:45:56.30	-00:48:51.8	1.8349(1)	994±41*	CO(2-1)	1.96±0.30	CO(4-3)	2.94±0.39		
HerBS-48	12:13:01.53	-00:49:23.2	3.1438(2)	773±83	CO(3-2)	3.59±0.69	CO(5-4)	5.16±0.74		
HerBS-50	12:03:19.14	-01:12:54.6	2.9283(2)	508±16*	CO(3-2)	2.56±0.25	CO(5-4)	4.02±0.30		
HerBS-51	12:07:09.12	-02:47:01.8	2.1827(1)	296±10	CO(3-2)	1.33±0.12	CO(4-3)	2.12±0.10	[CI](1-0)	1.06±0.27
HerBS-53 E	11:51:11.94	-01:26:38.7	1.4219(2)	362±60	CO(2-1)	0.98±0.23	CO(3-2)	1.44±0.36		
HerBS-53 W	11:51:12.28	-01:26:37.3	1.4236(3)	543±87	CO(2-1)	1.16±0.30	CO(3-2)	2.99±0.66		
HerBS-61	12:01:27.59	-01:40:45.8	3.7293(1)	722±40*	CO(3-2)	1.50±0.27	CO(4-3)	2.07±0.34	CO(6-5)	2.95±0.42
HerBS-62	12:15:42.81	-00:52:20.8	2.5738(1)	378±22	CO(3-2)	5.77±0.60	CO(4-3)	7.95±0.64		
HerBS-65	13:44:22.58	23:19:50.1	2.6858(5)	417±68	CO(3-2)	1.60±0.48	CO(5-4)	4.84±1.08		
HerBS-72	14:45:12.16	-00:15:10.5	3.6380(1)	541±124*	CO(4-3)	2.12±0.22	CO(6-5)	1.69±0.19	[CI](1-0)	0.72±0.52
HerBS-74	12:06:00.46	00:35:01.1	2.5596(3)	761±47*	CO(3-2)	1.79±0.38	CO(4-3)	1.01±0.22		
HerBS-76 E	13:35:34.09	34:18:34.8	2.3302(1)	200±15	CO(3-2)	3.01±0.30	CO(4-3)	4.01±0.45		
HerBS-76 W	13:35:33.83	34:18:33.4	2.3319(2)	742±149	CO(3-2)	0.75±0.12	CO(4-3)	1.26±0.38		
HerBS-78	14:33:52.56	02:04:17.2	3.7344(1)	489±30	CO(3-2)	1.09±0.17	CO(4-3)	2.04±0.20	CO(6-5)	2.05±0.23
HerBS-82	12:11:44.88	01:06:38.1	2.0583(1)	181±15*	CO(4-3)	4.07±0.38				
HerBS-83	12:18:13.06	01:18:43.2	3.9438(1)	742±114	CO(4-3)	1.23±0.34	CO(6-5)	1.62±0.40	H ₂ O(2 ₁₁ -2 ₀₂)	0.93±0.35
HerBS-85	11:47:52.85	-00:58:32.0	2.8169(1)	411±71	CO(3-2)	2.25±0.38	CO(5-4)	2.29±0.47	[CI](1-0)	1.86±0.55
HerBS-91 E	09:21:35.82	00:01:30.9	2.4048(1)	422±41	CO(3-2)	1.73±0.27	CO(4-3)	1.97±0.27		
HerBS-91 C	09:21:35.55	00:01:30.4	2.4047(1)	573±62	CO(3-2)	0.96±0.19	CO(4-3)	1.32±0.24	[CI](1-0)	1.05±0.27
HerBS-91 E&C	09:21:35.72	00:01:30.6	2.4048(1)	573±62					[CI](1-0)	1.05±0.27
HerBS-92 E&W	13:38:08.80	25:51:52.9	3.2644(2)	534±102	CO(3-2)	3.06±0.86	CO(5-4)	3.44±0.86		
HerBS-105	08:39:31.97	-01:18:00.0	2.6684(2)	520±29*	CO(3-2)	1.28±0.21	CO(5-4)	1.76±0.34		
HerBS-108	08:38:17.43	-00:41:34.4	3.7168(1)	541±53*	CO(3-2)	0.29±0.10	CO(4-3)	2.82±0.33	CO(6-5)	3.28±0.39
HerBS-109 NW	13:29:00.35	28:19:18.5	1.5850(1)	662±84	CO(2-1)	1.64±0.29	CO(3-2)	2.88±0.57		
HerBS-109 S	13:29:00.31	28:19:07.5	1.5843(3)	569±81	CO(2-1)	0.93±0.20	CO(3-2)	2.10±0.42		
HerBS-109 NE	13:29:00.77	28:19:16.3	2.8385(2)	330±62	CO(3-2)	0.70±0.18	CO(5-4)	1.21±0.37		
HerBS-110	14:18:33.24	01:02:10.9	2.6810(2)	915±25*	CO(3-2)	2.28±0.24	CO(5-4)	2.95±0.26		
HerBS-115	13:35:38.20	26:57:40.1	2.3706(1)	1020±72	CO(3-2)	2.45±0.29	CO(4-3)	3.27±0.37	[CI](1-0)	1.19±0.36
HerBS-116 E&W	12:13:48.05	01:08:11.0	3.1547(1)	800±107	CO(3-2)	3.02±0.57	CO(5-4)	4.14±0.87		
HerBS-124 W	12:21:58.48	00:33:24.9	2.2772(1)	449±35	CO(3-2)	2.26±0.25	CO(4-3)	3.62±0.45		
HerBS-124 E	12:21:58.65	00:33:25.0	2.2781(1)	630±175	CO(3-2)	0.55±0.22	CO(4-3)	0.85±0.35		
HerBS-125	13:04:32.18	29:53:39.6	2.5739(1)	615±55	CO(3-2)	2.29±0.30	CO(4-3)	3.79±0.50		
HerBS-126	14:51:35.27	-01:14:17.3	2.5875(3)	1083±128	CO(3-2)	2.05±0.33	CO(4-3)	0.80±0.16		
HerBS-127	13:21:28.76	28:20:23.2	3.1958(1)	839±98*	CO(3-2)	2.42±0.89	CO(5-4)	2.53±1.01		
HerBS-128	13:04:14.48	30:35:38.2	2.0681(2)	643±27*	CO(3-2)	2.54±0.32	CO(4-3)	2.26±0.28		
HerBS-129	13:00:53.80	26:03:00.3	3.3074(2)	645±19*	CO(3-2)	3.13±0.33	CO(5-4)	2.82±0.28		
HerBS-134	13:34:40.42	35:31:39.1	3.1725(3)	793±173	CO(3-2)	2.46±0.83	CO(5-4)	2.40±0.74		
HerBS-136	08:53:08.45	-00:57:28.8	3.2884(5)	1042±88	CO(3-2)	1.02±0.16	CO(5-4)	1.77±0.20		
HerBS-137	14:53:37.15	00:04:10.2	3.0408(3)	450±86	CO(3-2)	1.45±0.38	CO(5-4)	1.29±0.43		
HerBS-140	14:21:40.42	00:04:46.3	2.7799(2)	770±26*	CO(3-2)	0.69±0.09	CO(5-4)	1.53±0.16		
HerBS-143	14:18:10.10	-00:37:45.5	2.2406(1)	540±43	CO(3-2)	1.43±0.20	CO(4-3)	1.94±0.27	[CI](1-0)	1.00±0.16
HerBS-147	14:34:03.66	00:02:30.1	3.1150(1)	1789±300	CO(3-2)	2.00±0.60	CO(5-4)	2.98±0.68		
HerBS-149	13:38:27.61	31:39:55.6	2.6650(1)	243±21	CO(3-2)	2.29±0.30	CO(5-4)	5.45±0.65		
HerBS-150E	12:24:59.27	-00:56:51.9	3.6732(1)	780±128	CO(4-3)	1.92±0.51	CO(6-5)	1.03±0.23		
HerBS-150C	12:24:58.98	-00:56:47.5	3.6682(1)	212±64	CO(4-3)	0.26±0.10				
HerBS-150W	12:24:58.70	-00:56:48.7	3.6787(1)	1036±178	CO(4-3)	1.47±0.32				
HerBS-153	14:42:43.41	01:55:04.3	3.1501(2)	552±16*	CO(3-2)	0.89±0.19	CO(5-4)	3.72±0.33		
HerBS-157	08:49:57.77	01:07:10.8	1.8964(1)	849±45*	CO(2-1)	1.53±0.33	CO(4-3)	1.39±0.22		
HerBS-162 NE	14:43:34.73	-00:30:30.5	2.4742(4)	483±79*	CO(3-2)	1.18±0.27	CO(4-3)	0.55±0.14		
HerBS-162 SW	14:43:34.35	-00:30:35.7	2.4735(2)	288±41*	CO(3-2)	1.44±0.25	CO(4-3)	1.27±0.25		
HerBS-164	12:14:16.30	-01:37:03.7	2.0126(1)	644±97	CO(2-1)	1.04±0.27	CO(4-3)	3.25±0.27		
HerBS-165	09:06:13.91	-01:00:41.1	2.2251(1)	445±67	CO(3-2)	0.97±0.25	CO(4-3)	1.49±0.30		
HerBS-167	13:03:41.80	31:37:57.9	2.2144(7)	733±129*	CO(3-2)	1.05±0.27	CO(4-3)	0.97±0.25		
HerBS-169	08:38:59.47	02:13:27.4	2.6977(1)	617±48	CO(3-2)	1.73±0.18			[CI](1-0)	0.89±0.15
HerBS-171	08:39:45.21	02:10:18.1	2.4793(1)	382±56	CO(3-2)	0.95±0.20	CO(4-3)	1.17±0.26		
HerBS-172	14:50:40.47	00:33:35.8	2.9246(4)	386±41	CO(3-2)	1.34±0.22	CO(5-4)	0.74±0.11		
HerBS-175	12:19:00.83	00:33:26.9	3.1575(2)	660±77	CO(3-2)	3.48±0.55	CO(5-4)	0.43±0.09		
HerBS-176	13:12:22.15	27:02:17.8	2.9805(4)	733±76	CO(3-2)	3.63±0.53	CO(5-4)	5.22±0.84		
HerBS-177	11:54:33.72	00:50:41.8	3.9625(1)	606±53	CO(4-3)	3.36±0.49	CO(6-5)	4.97±0.75	H ₂ O(2 ₁₁ -2 ₀₂)	2.28±0.51
HerBS-179	11:55:20.97	-02:13:29.4	3.9423(1)	391±39	CO(4-3)	2.41±0.45	CO(6-5)	3.71±0.56	H ₂ O(2 ₁₁ -2 ₀₂)	0.52±0.31

Table A.7. continued - Spectroscopic redshifts and emission line properties: HerBS sources.

Source	Observed Position		z_{spec}	ΔV (km s ⁻¹)	Line	Line Flux (Jy km s ⁻¹)	Emission Lines			
	RA (J2000)	Dec.					Line	Line Flux (Jy km s ⁻¹)	Line	Line Flux (Jy km s ⁻¹)
HerBS-180 S	13:15:39.55	29:22:21.6	1.4527(1)	308±34	CO(2-1)	1.77±0.26	CO(3-2)	1.75±0.38		
HerBS-183	09:04:53.06	02:20:16.9	1.8910(1)	687±21*	CO(2-1)	0.89±0.09	CO(4-3)	3.21±0.32		
HerBS-185	09:24:08.92	-00:50:18.1	4.3238(1)	310±24	CO(7-6)	1.95±0.40	H ₂ O(2 ₁₁ -2 ₀₂)	0.85±0.30		
-							[CI](1-0)	0.51±0.34	[CI](2-1)	1.42±0.30
HerBS-187 E	08:37:05.58	02:00:31.5	1.8285(1)	591±76	CO(2-1)	2.09±0.40	CO(4-3)	3.40±0.63		
HerBS-187 W	08:37:05.35	02:00:32.7	1.8274(1)	403±78	CO(2-1)	0.60±0.20	CO(4-3)	1.82±0.47		
HerBS-188	08:42:59.98	02:50:01.7	2.7675(1)	1069±147*	CO(3-2)	1.94±0.38	CO(5-4)	2.54±0.53		
HerBS-190	09:04:05.28	-00:33:33.5	2.5890(2)	784±77	CO(3-2)	3.17±0.42	CO(4-3)	3.92±0.67		
HerBS-191	12:47:53.28	32:24:45.8	3.4428(3)	761±36*	CO(4-3)	3.23±0.55	CO(5-4)	2.52±0.39		
HerBS-193	08:53:51.96	-00:08:05.4	3.6951(1)	680±30	CO(3-2)	0.71±0.14	CO(4-3)	1.29±0.10	CO(6-5)	1.54±0.10
HerBS-194 N	08:55:21.30	-00:35:56.6	2.3335(1)	467±84	CO(3-2)	0.72±0.21	CO(4-3)	1.86±0.47		
HerBS-194 S	08:55:21.04	-00:36:11.9	2.3316(1)	538±140	CO(3-2)	0.38±0.14	CO(4-3)	0.34±0.13		
HerBS-197	12:20:34.05	-00:38:04.4	2.4170(1)	195±10	CO(3-2)	2.59±0.35	CO(4-3)	3.75±0.28	[CI](1-0)	1.11±0.16
HerBS-199 E	13:33:52.30	33:49:13.5	1.9248(2)	438±66	CO(2-1)	1.98±0.58	CO(4-3)	3.05±0.61		
HerBS-199 W	13:33:51.45	33:49:18.9	1.9197(1)	420±97	CO(2-1)	0.50±0.45	CO(4-3)	2.75±0.83		
HerBS-201	14:11:18.06	-01:06:52.6	4.1408(1)	553±69	CO(4-3)	1.17±0.25	CO(7-6)	1.15±0.22		
-							[CI](1-0)	0.71±0.21	[CI](2-1)	0.50±0.10
HerBS-202	14:33:28.37	02:08:09.8	2.0224(1)	411±21	CO(2-1)	1.51±0.31	CO(4-3)	2.66±0.18		
HerBS-204 W	13:29:09.21	30:09:59.3	3.4937(4)	378±34	CO(4-3)	1.32±0.16	CO(5-4)	2.07±0.31		
HerBS-204 E	13:29:09.71	30:09:57.1	3.4933(2)	365±44	CO(4-3)	1.01±0.18	CO(5-4)	1.27±0.23		
HerBS-205 NE	14:51:32.92	02:41:02.8	2.9600(1)	575±58	CO(3-2)	1.26±0.21	CO(5-4)	0.92±0.13		
HerBS-205 SE	14:51:32.91	02:40:58.9	2.9599(4)	450±53	CO(3-2)	0.66±0.13	CO(5-4)	0.50±0.08		
HerBS-205 W	14:51:32.38	02:41:00.5	2.9630(1)	566±117	CO(3-2)	0.27±0.12	CO(5-4)	0.57±0.16		
HerBS-206	14:04:21.76	-00:12:17.1	2.8122(4)	410±41	CO(3-2)	0.54±0.08	CO(5-4)	0.92±0.13		

Notes. The uncertainties in the spectroscopic redshifts, z_{spec} , given in brackets, correspond to the errors derived from the simultaneous Gaussian fits to the line profiles. Emission lines fitted with double Gaussians are highlighted with an asterisk after the value of ΔV . The redshift and line identifications that are not robust are in italics (in the case of HerBS-82). Further details about the line fitting procedure and the derivation of z_{spec} are provided in Appendix A.1.

Table A.8. Spectroscopic redshifts and emission line properties: HerBS sources from Neri et al. (2020).

Source	Observed Position		z_{spec}	ΔV (km s ⁻¹)	Line	Line Flux (Jy km s ⁻¹)	Emission Lines			
	RA (J2000)	Dec.					Line	Line Flux (Jy km s ⁻¹)	Line	Line Flux (Jy km s ⁻¹)
HerBS-34	13:34:13.87	26:04:57.5	2.6637(2)	330±10*	CO(3-2)	2.8±0.4	CO(5-4)	8.4±0.8		
HerBS-43a	13:24:18.79	32:07:54.4	3.2121(1)	1070±90*	CO(4-3)	5.5±0.8	CO(5-4)	6.7±0.8		
HerBS-43b	13:24:19.24	32:07:49.2	4.0543(7)	800±50*	CO(4-3)	1.7±0.3	CO(5-4)	1.5±0.3	CO(6-5)	4.8±0.9
HerBS-44	13:32:55.85	34:22:08.4	2.9268(2)	520±50*	CO(3-2)	4.9±0.5	CO(5-4)	12.5±1.2		
HerBS-54	13:15:40.72	26:23:19.6	2.4417(3)	1020±190	CO(3-2)	4.3±0.4	CO(4-3)	8.5±0.8		
HerBS-58	13:03:33.17	24:46:42.3	2.0842(1)	970±50*	CO(3-2)	5.3±1.0	CO(4-3)	5.2±1.5	[CI](1-0)	4.7±0.5
HerBS-70 E	13:01:40.33	29:29:16.2	2.3077(4)	770±50*	CO(3-2)	1.8±0.5	CO(4-3)	3.4±0.3	[CI](1-0)	3.5±0.7
HerBS-70 W	13:01:39.31	29:29:25.2	2.3115(1)	140±20	CO(3-2)	1.7±0.3	CO(4-3)	2.0±0.3		
HerBS-79	13:14:34.08	33:52:20.1	2.0782(8)	870±70*	CO(3-2)	4.1±0.8	CO(4-3)	5.5±0.5		
HerBS-89a	13:16:11.52	28:12:17.7	2.9497(1)	1102±83*	CO(3-2)	4.00±0.60	CO(5-4)	8.40±0.80		
HerBS-95 E	13:43:42.73	26:39:18.0	2.9718(3)	870±50*	CO(3-2)	1.0±0.1	CO(5-4)	3.6±0.3		
HerBS-95 W	13:43:41.55	26:39:22.7	2.9729(2)	540±30*	CO(3-2)	2.4±0.4	CO(5-4)	3.5±0.3		
HerBS-113	13:12:11.35	32:38:37.8	2.7870(8)	900±200*	CO(3-2)	6.1±1.2	CO(5-4)	13.5±1.4		
HerBS-154	13:22:58.11	32:50:51.7	3.7070(5)	310±40	CO(6-5)	7.6±0.7	[CI](1-0)	1.3±0.4	H ₂ O(2 ₁₁ -2 ₀₂)	1.5±0.3

Notes. The uncertainties in the spectroscopic redshifts, z_{spec} , given in brackets, correspond to the errors derived from Gaussian fits to the line profiles. ΔV values are the mean linewidths (FWHM) weighted by the peak intensities of the detected CO emission lines. Emission lines fitted with double Gaussians are highlighted with an asterisk after the value of ΔV - the reader is referred to Neri et al. (2020) for further details.

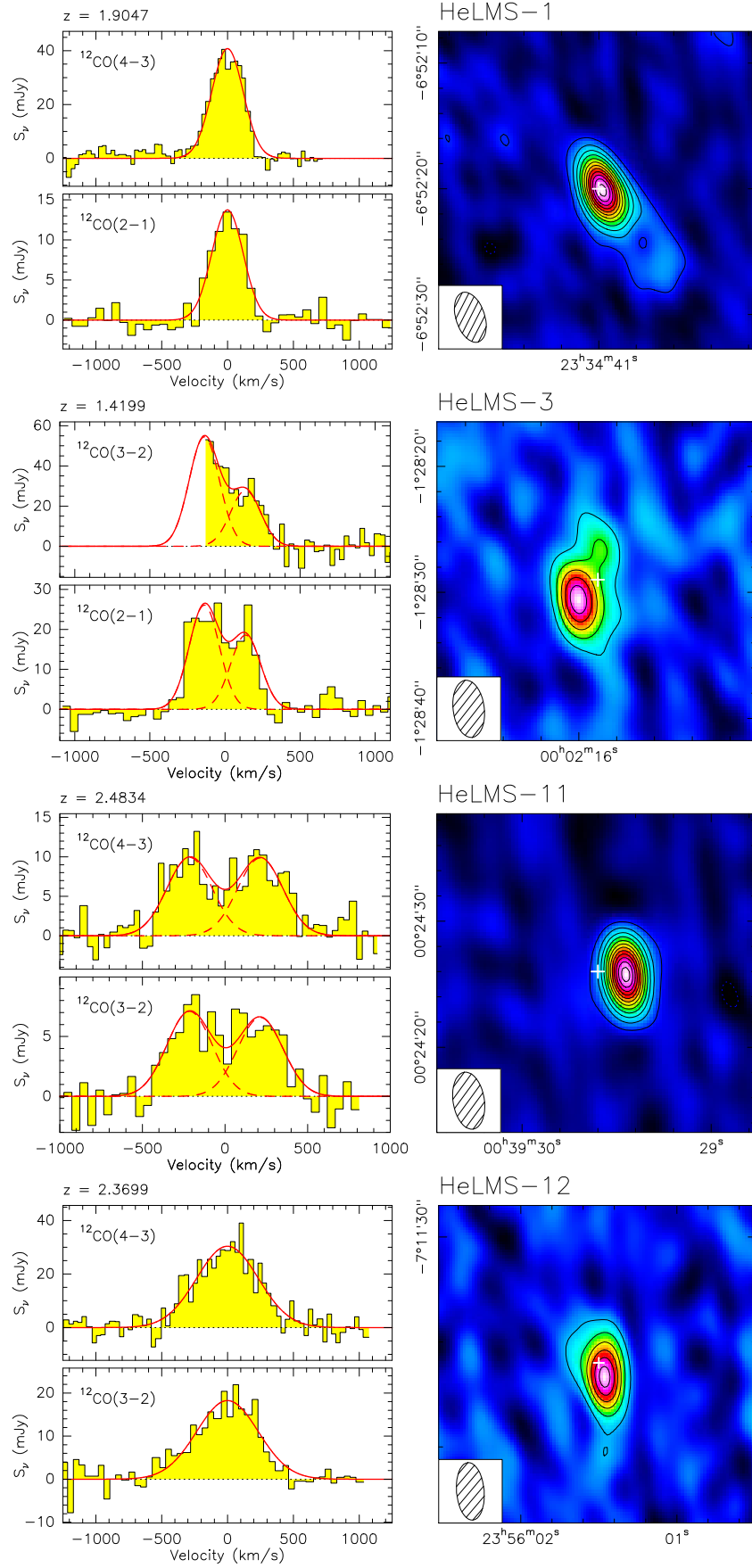


Fig. A.1. Combined continuum and emission lines maps of the z -GAL sources selected from the HeLMS sample and spectra of each of the detected emission lines. Further details are provided in Appendix A.1.

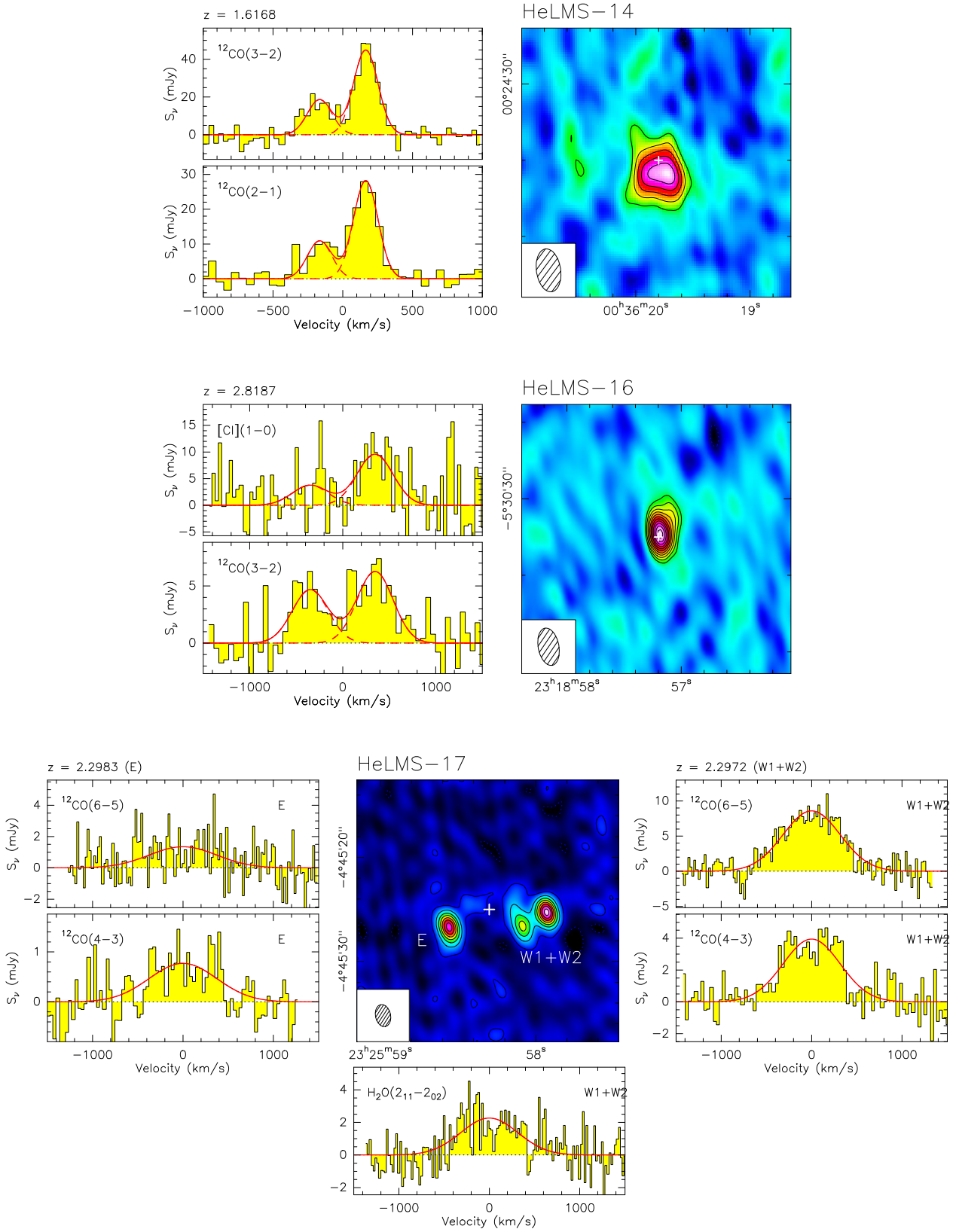


Fig. A.1. continued.

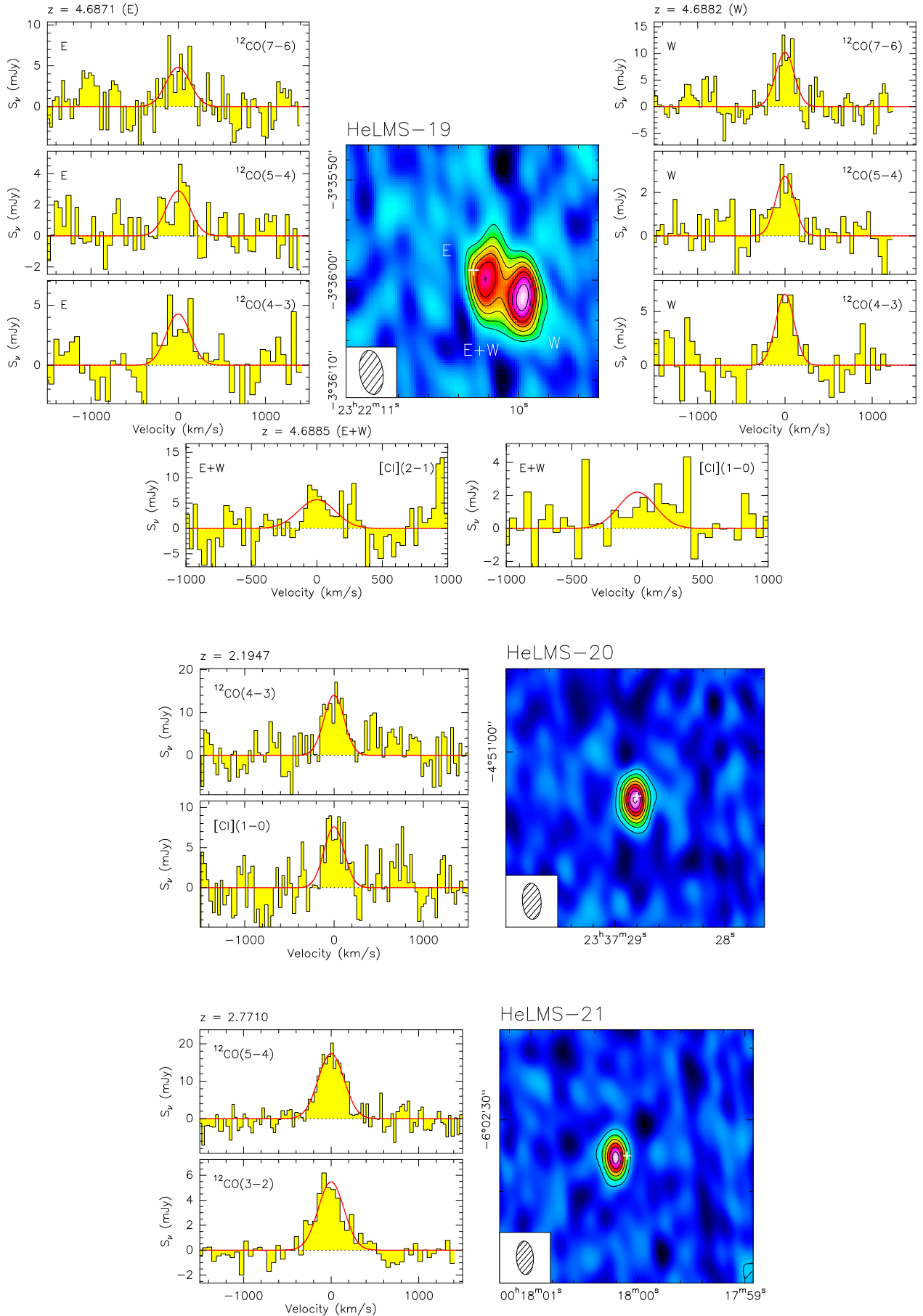


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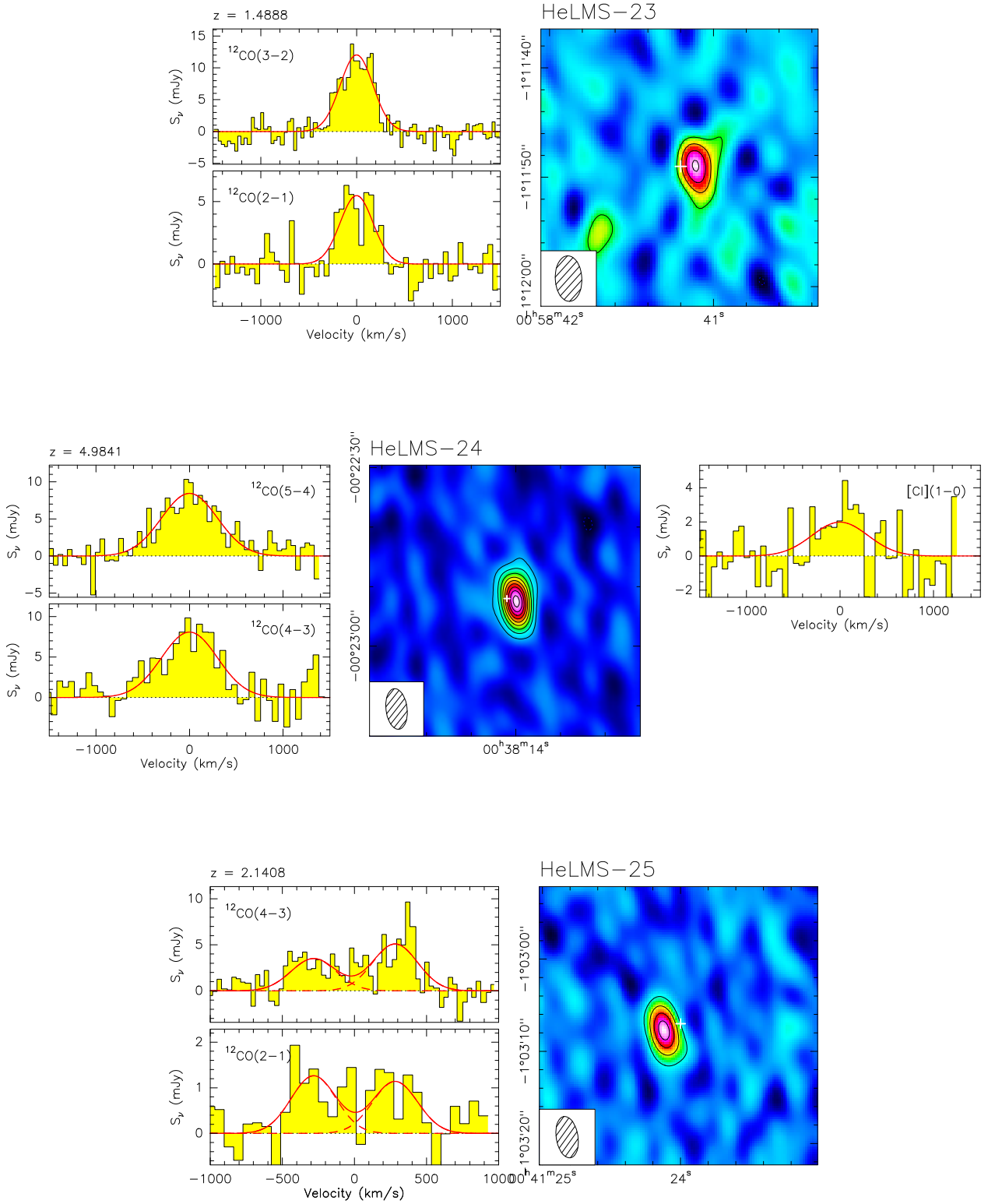


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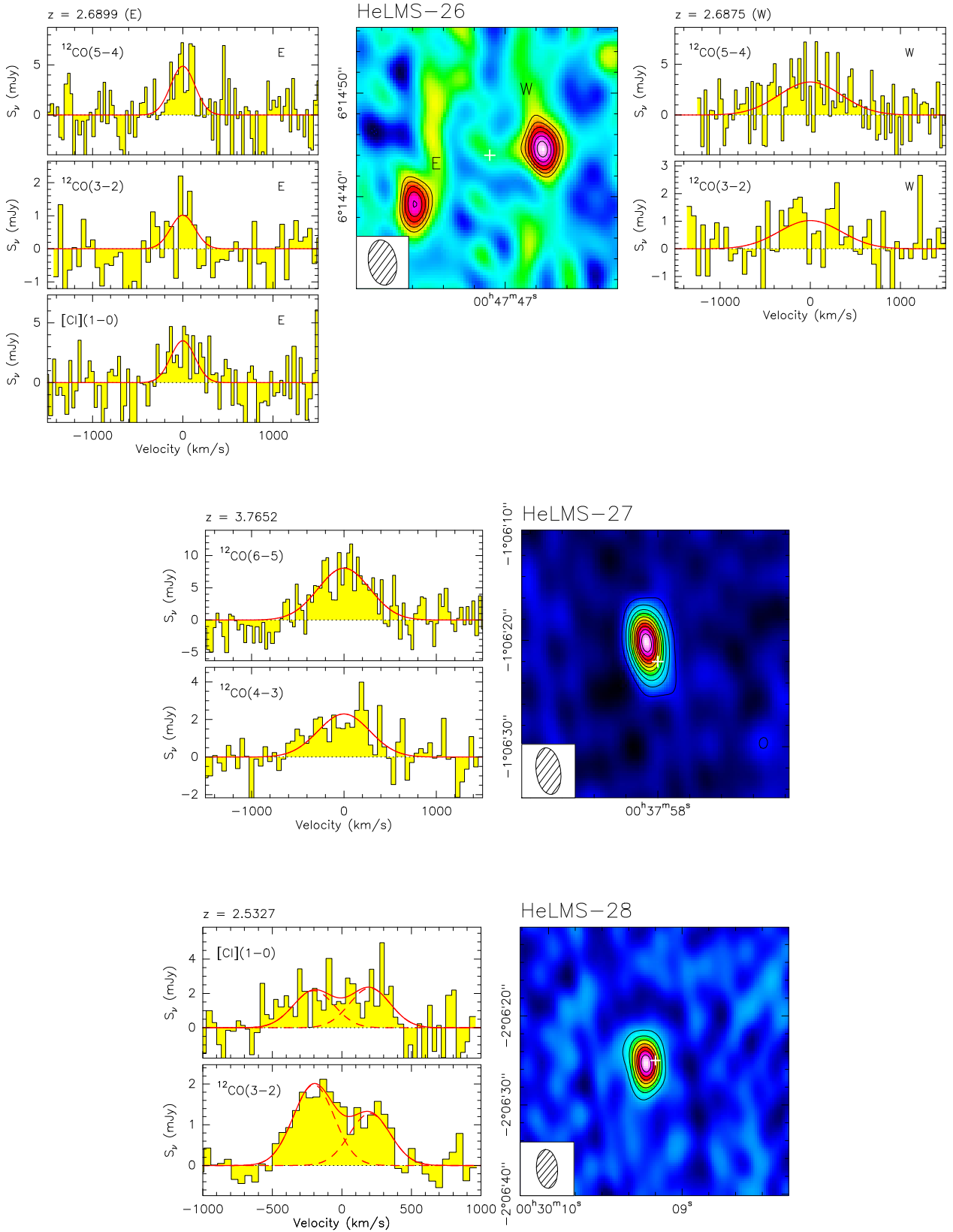


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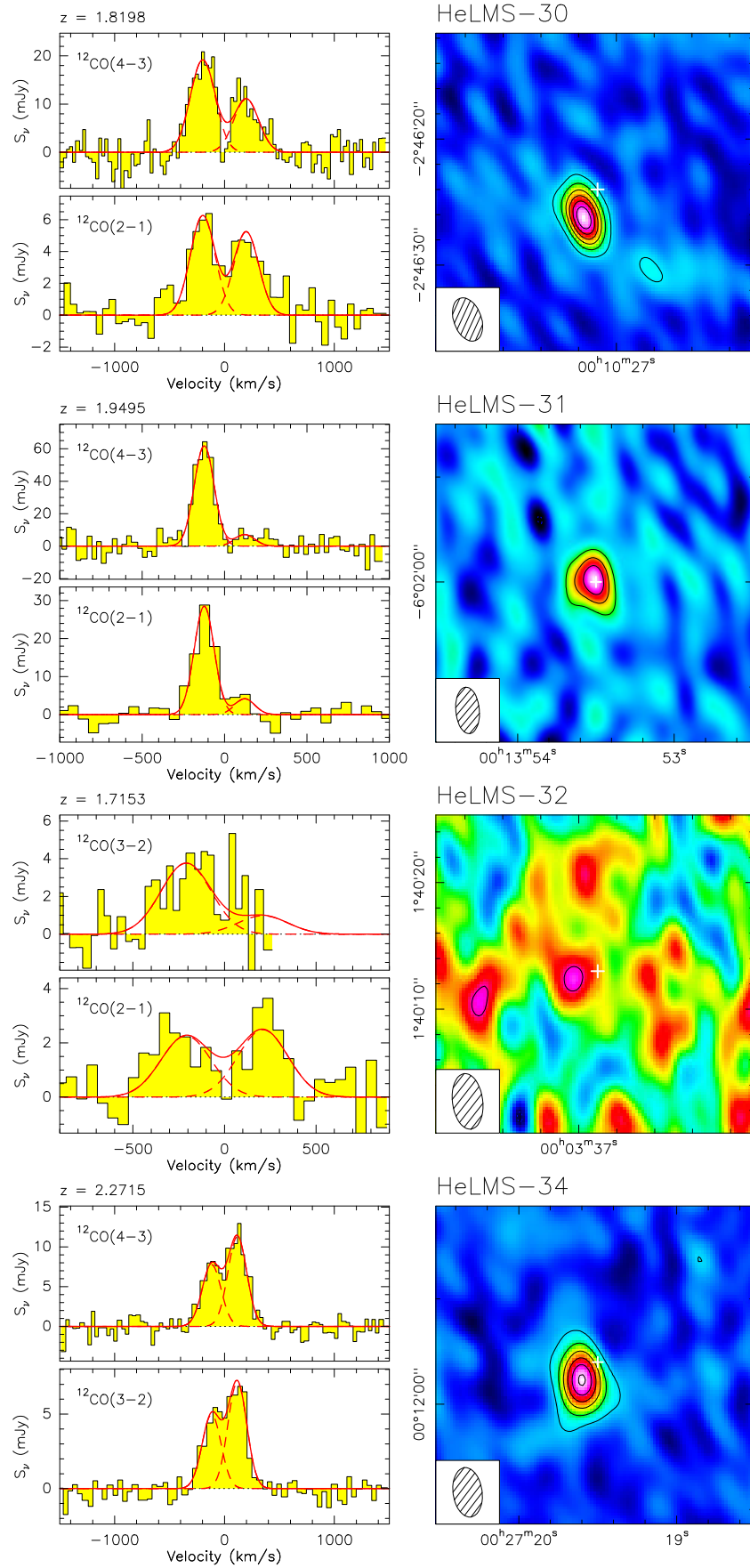


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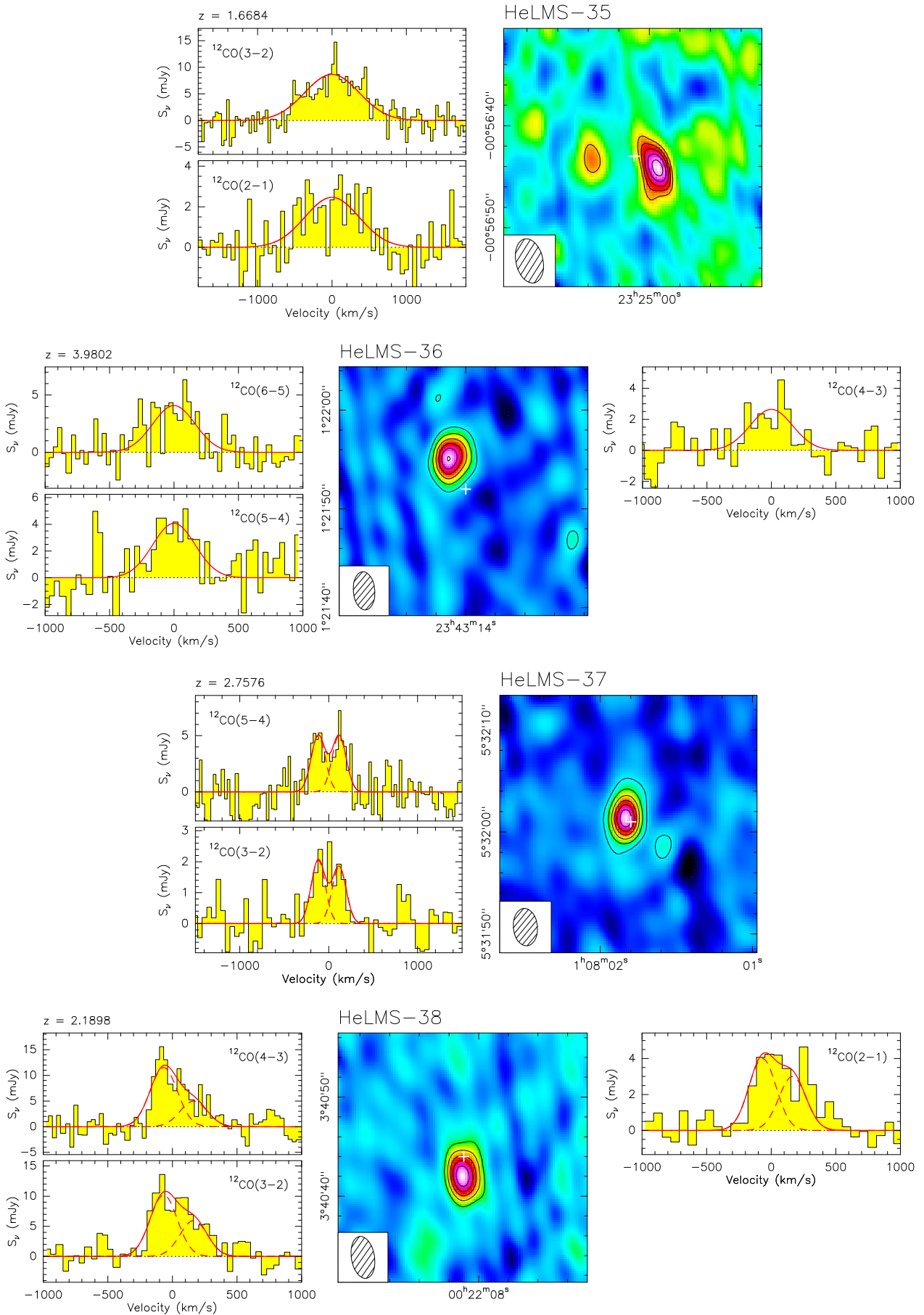


Fig. A.1. continued.

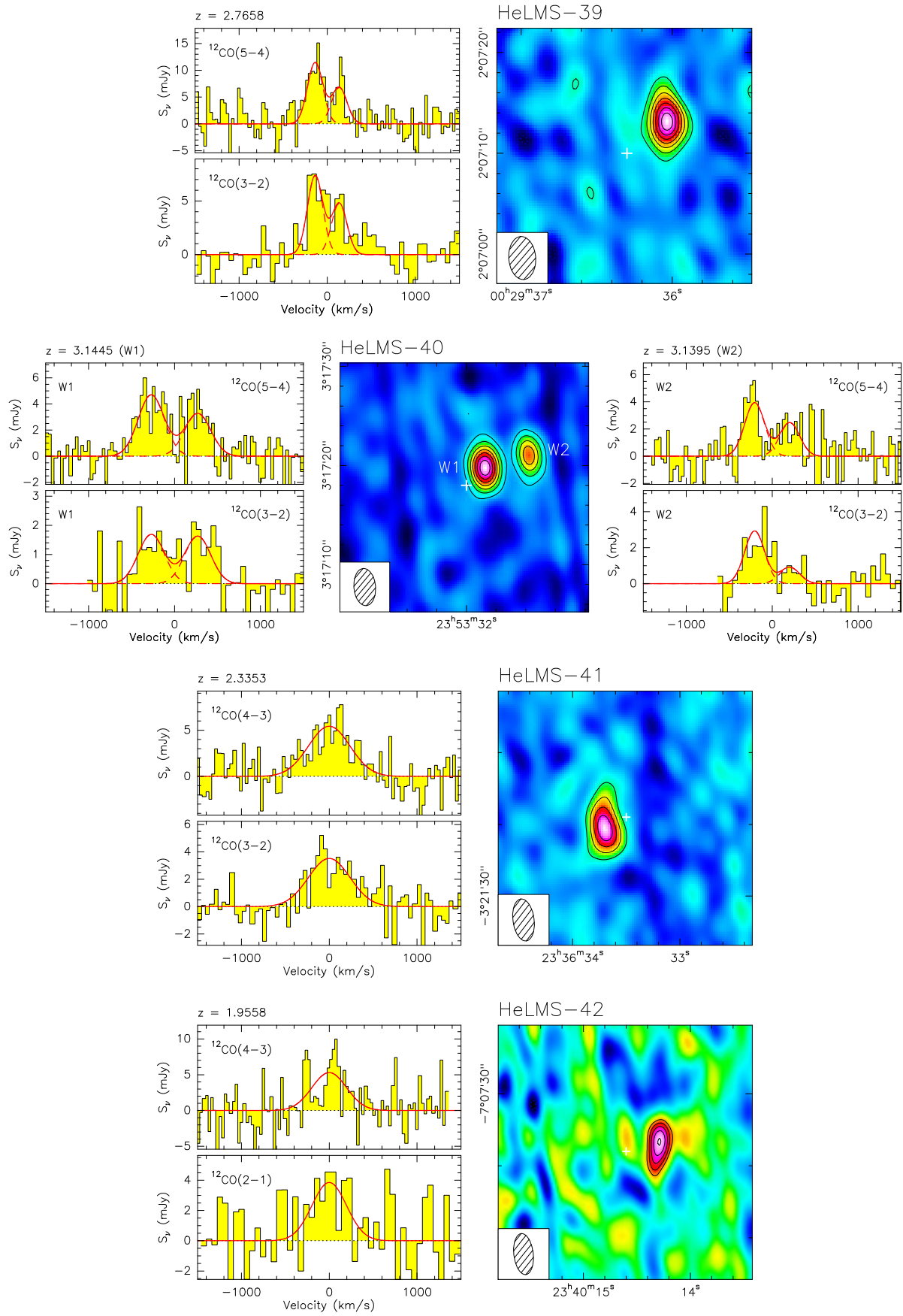


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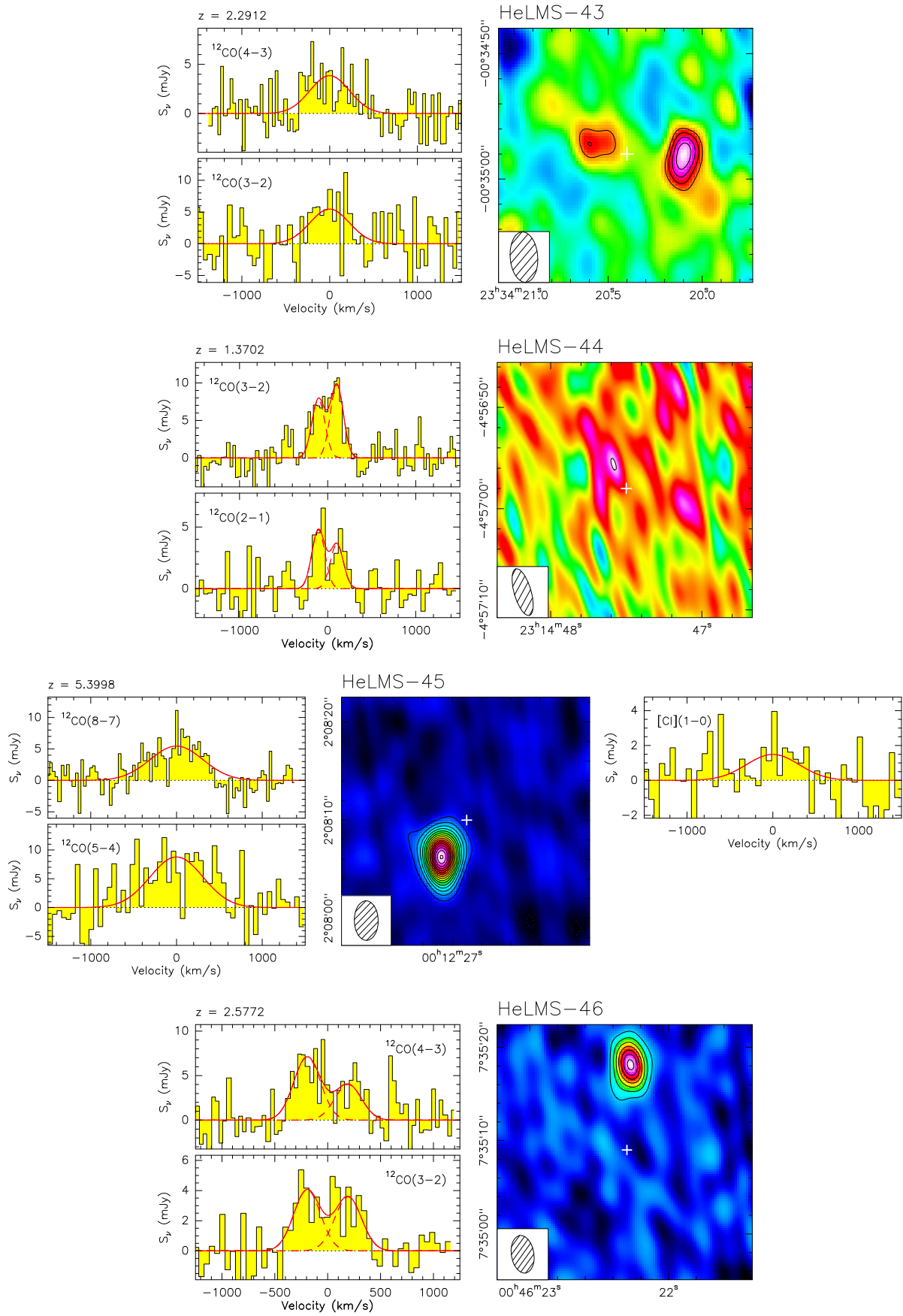


Fig. A.1. continued.

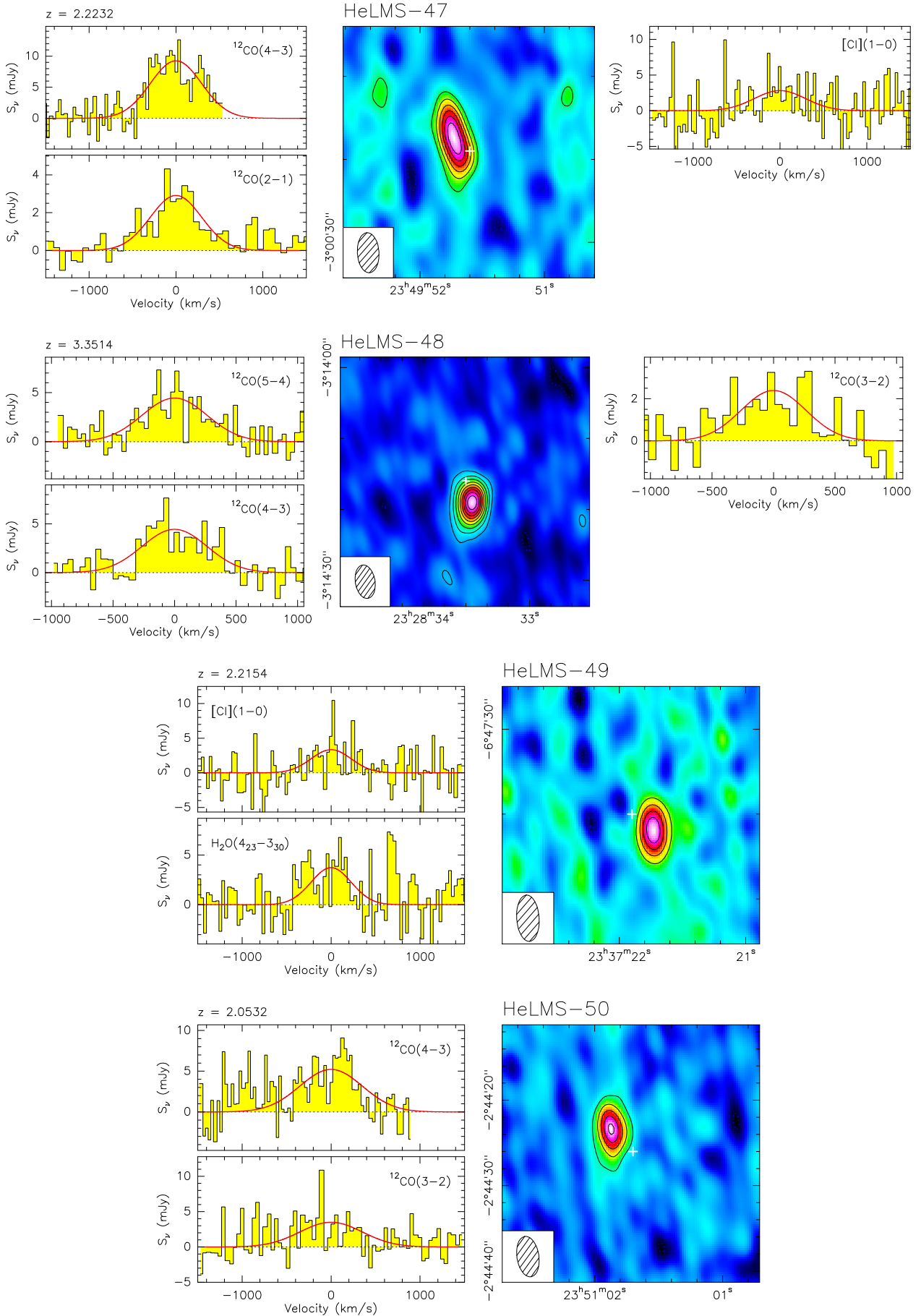


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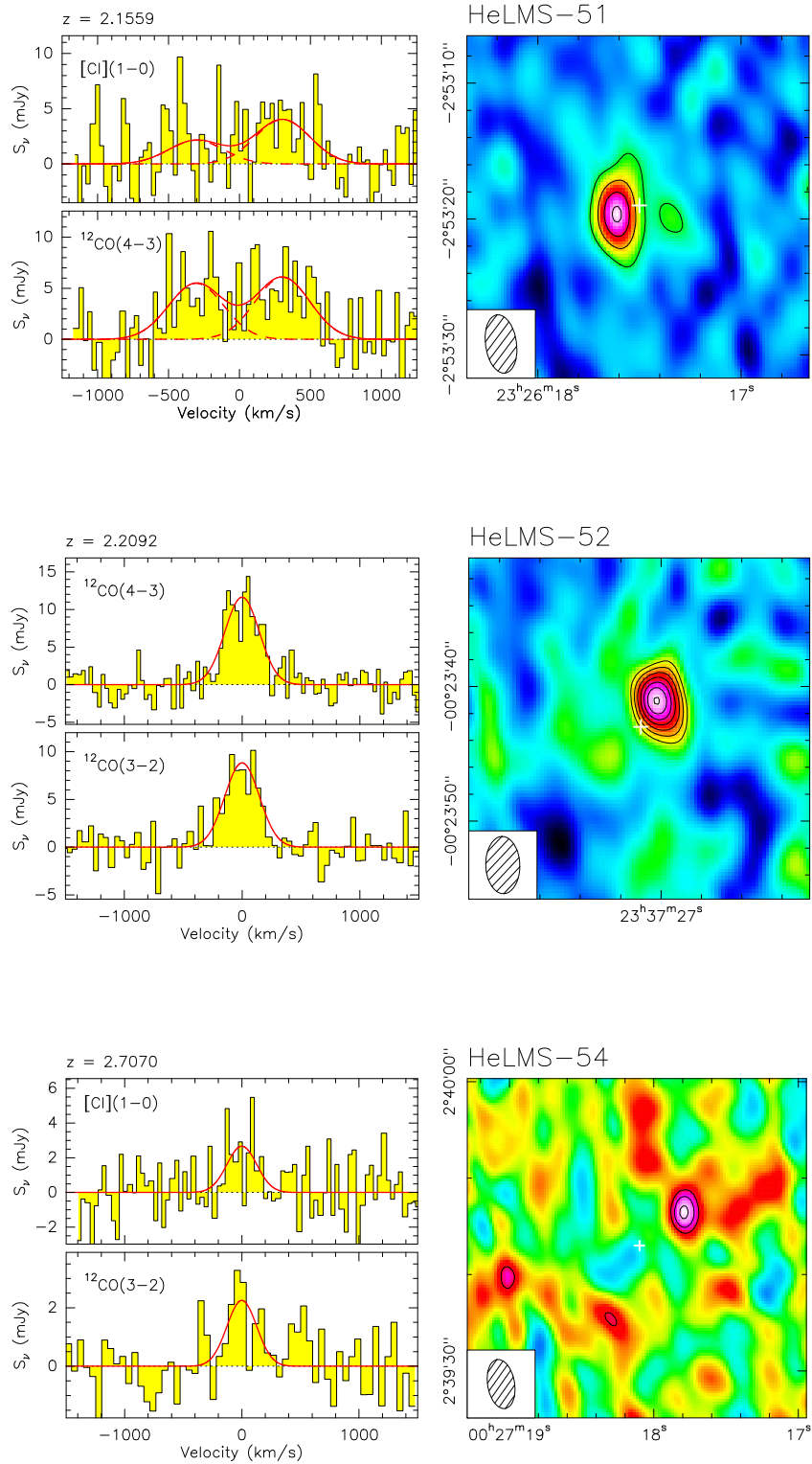


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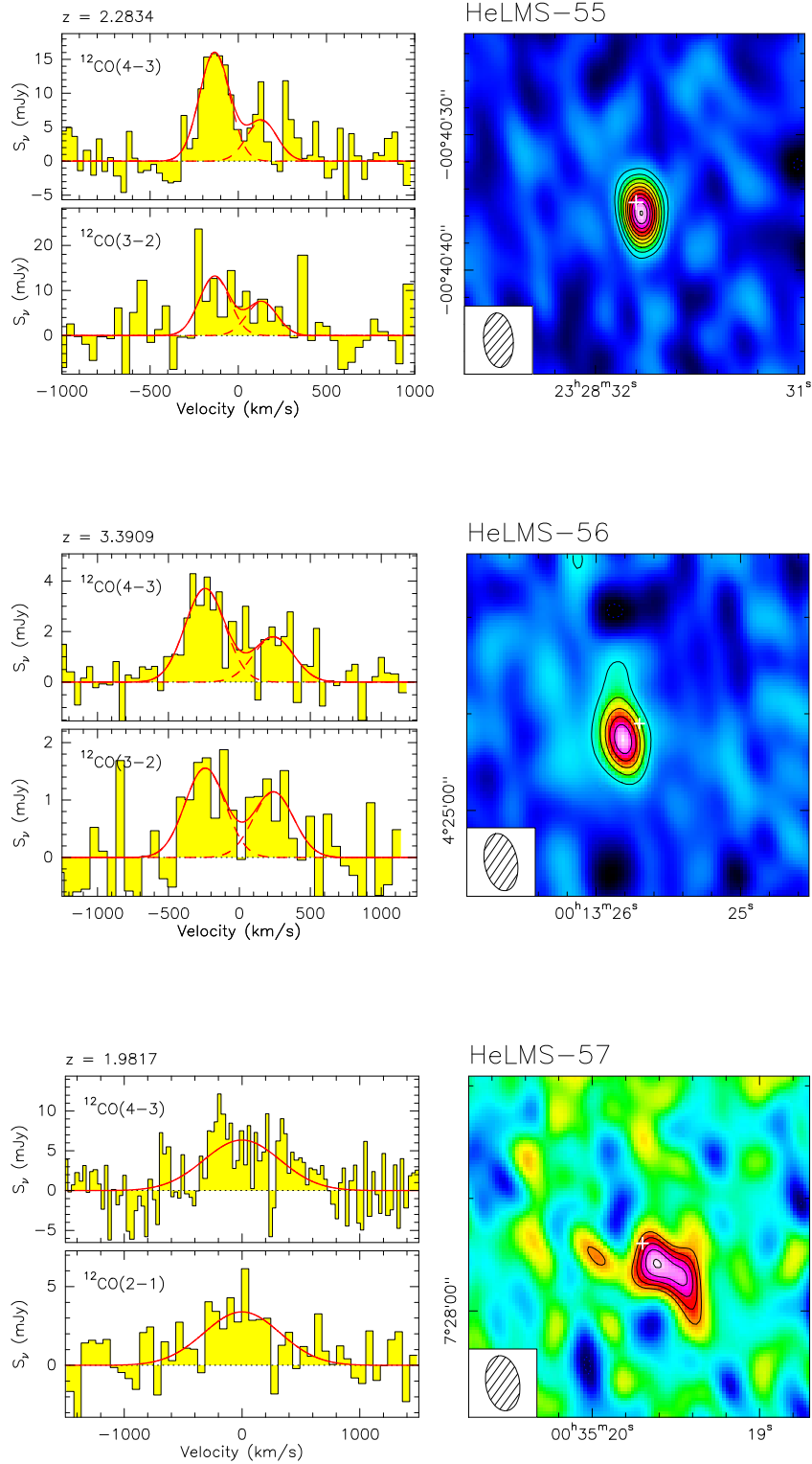


Fig. A.2. Combined continuum and emission lines maps of the z -GAL sources selected from the HerS sample and spectra of each of the detected emission lines. Further details are provided in Appendix A.1.

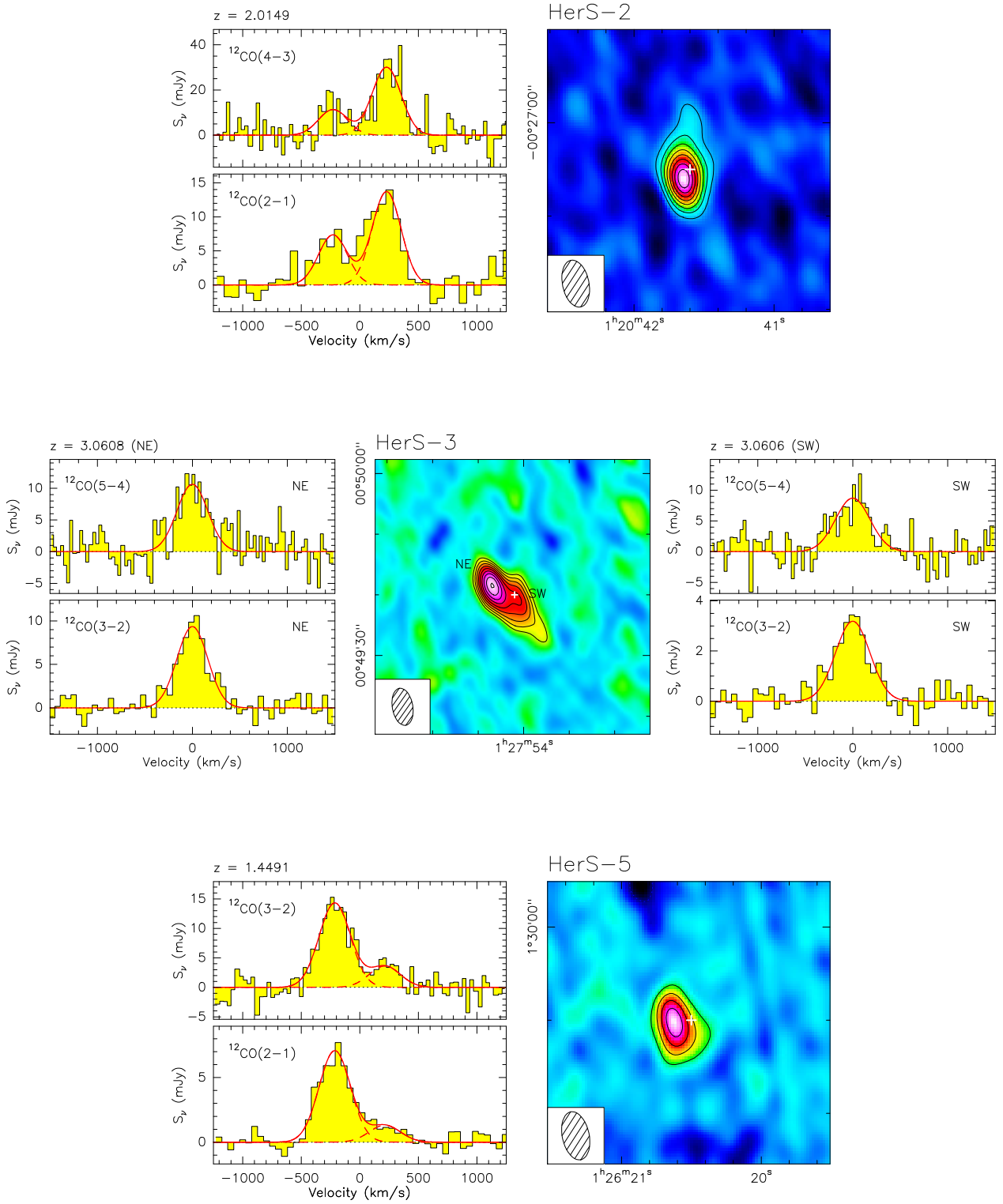


Fig. A.2. continued.

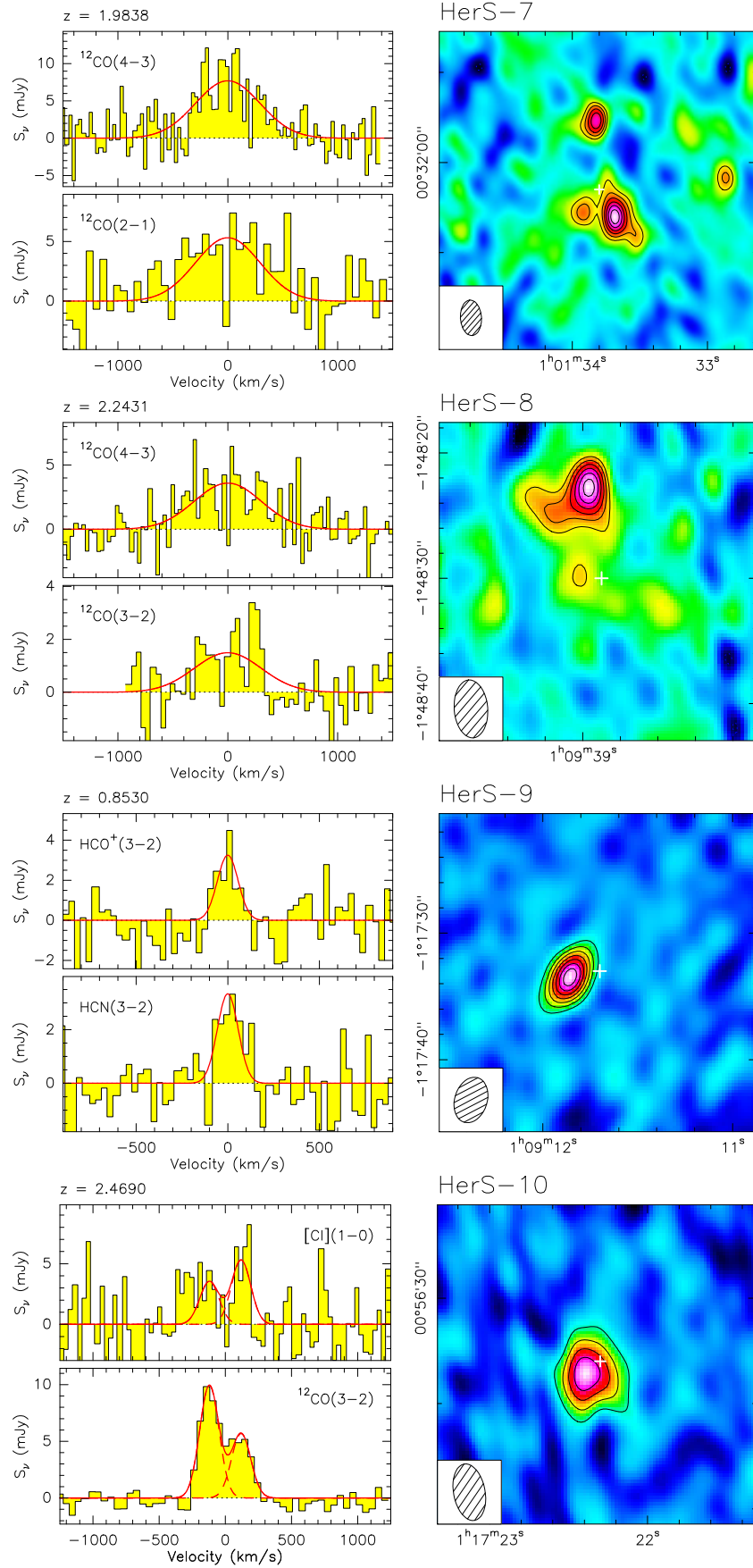


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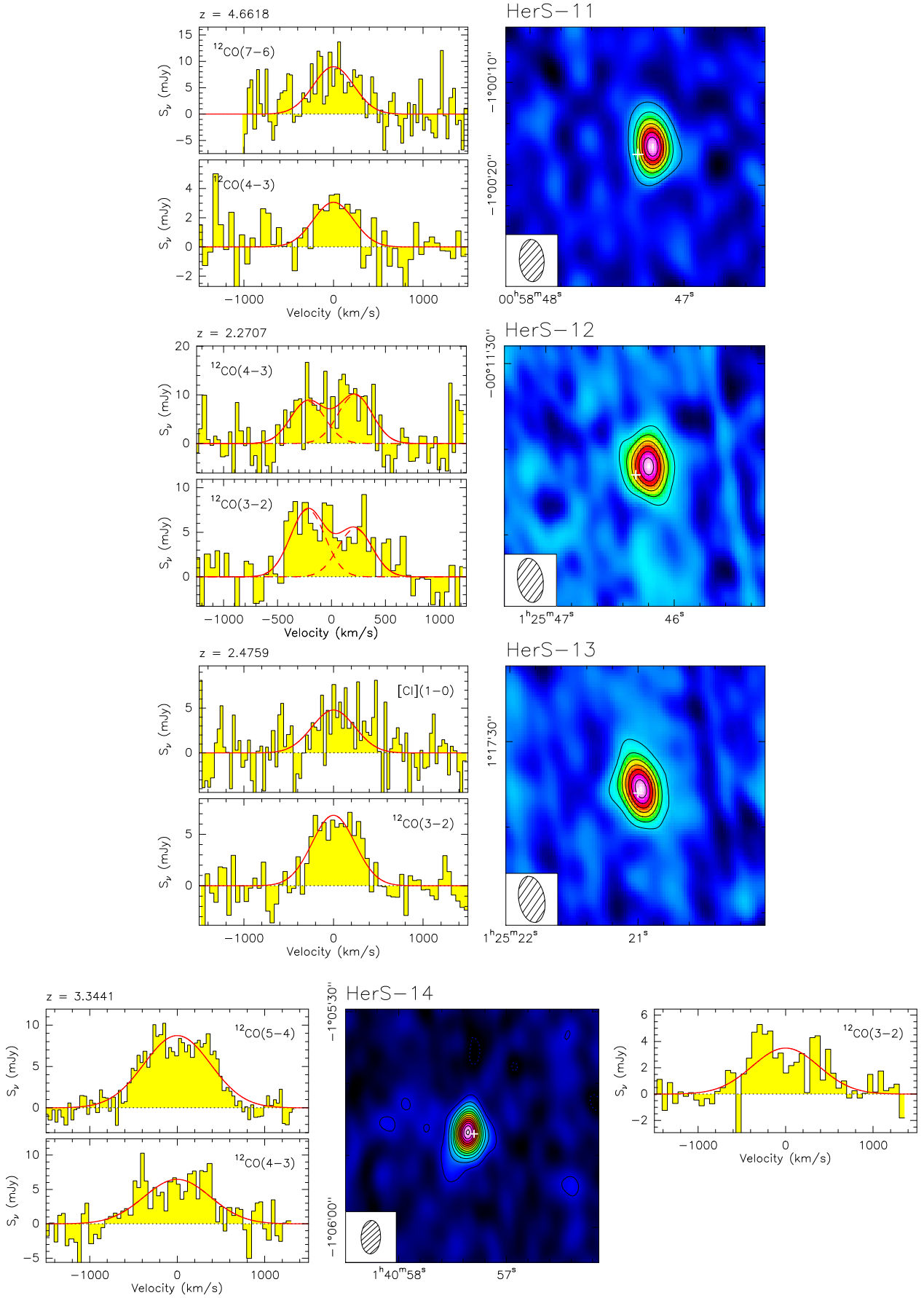


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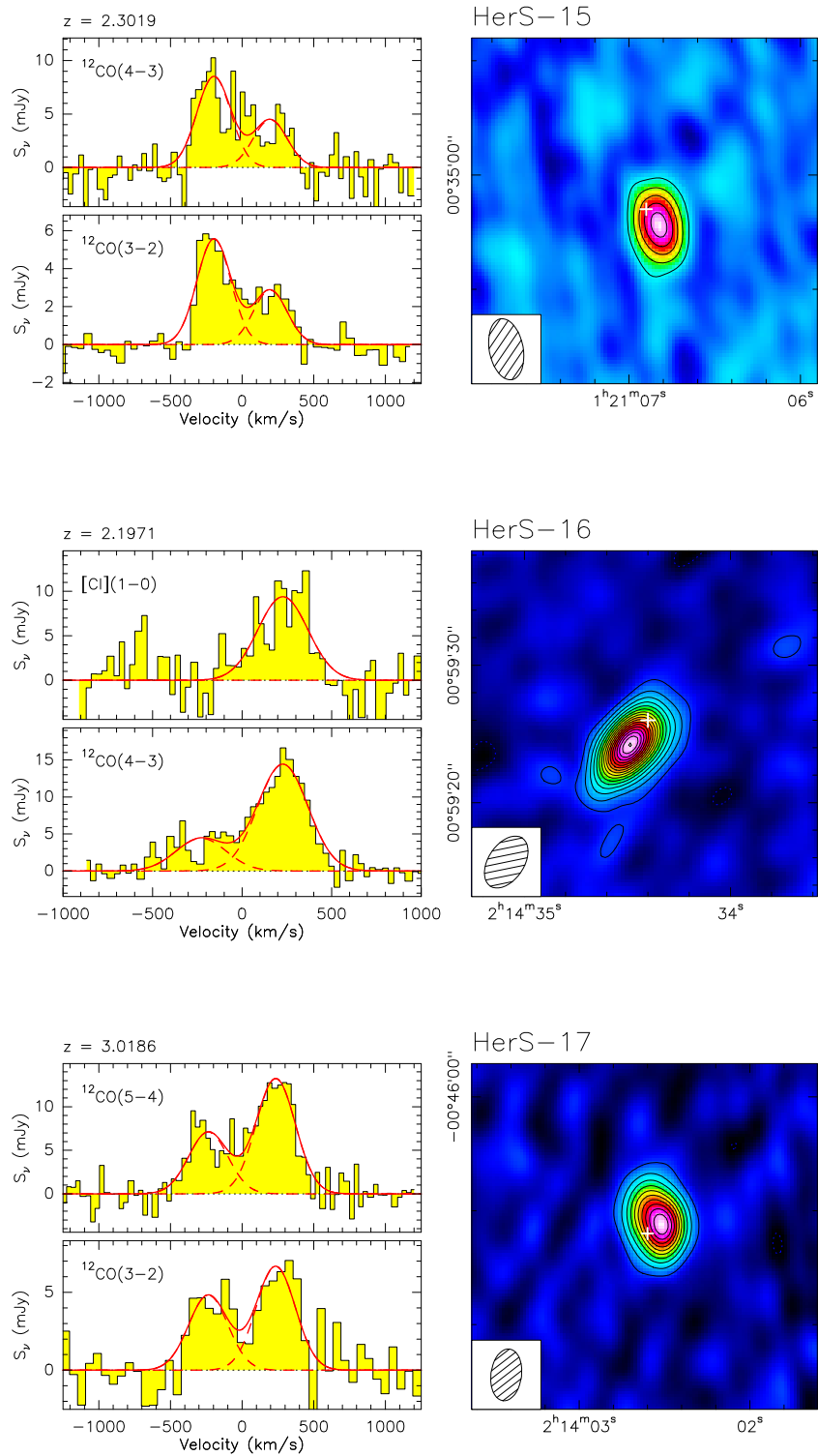


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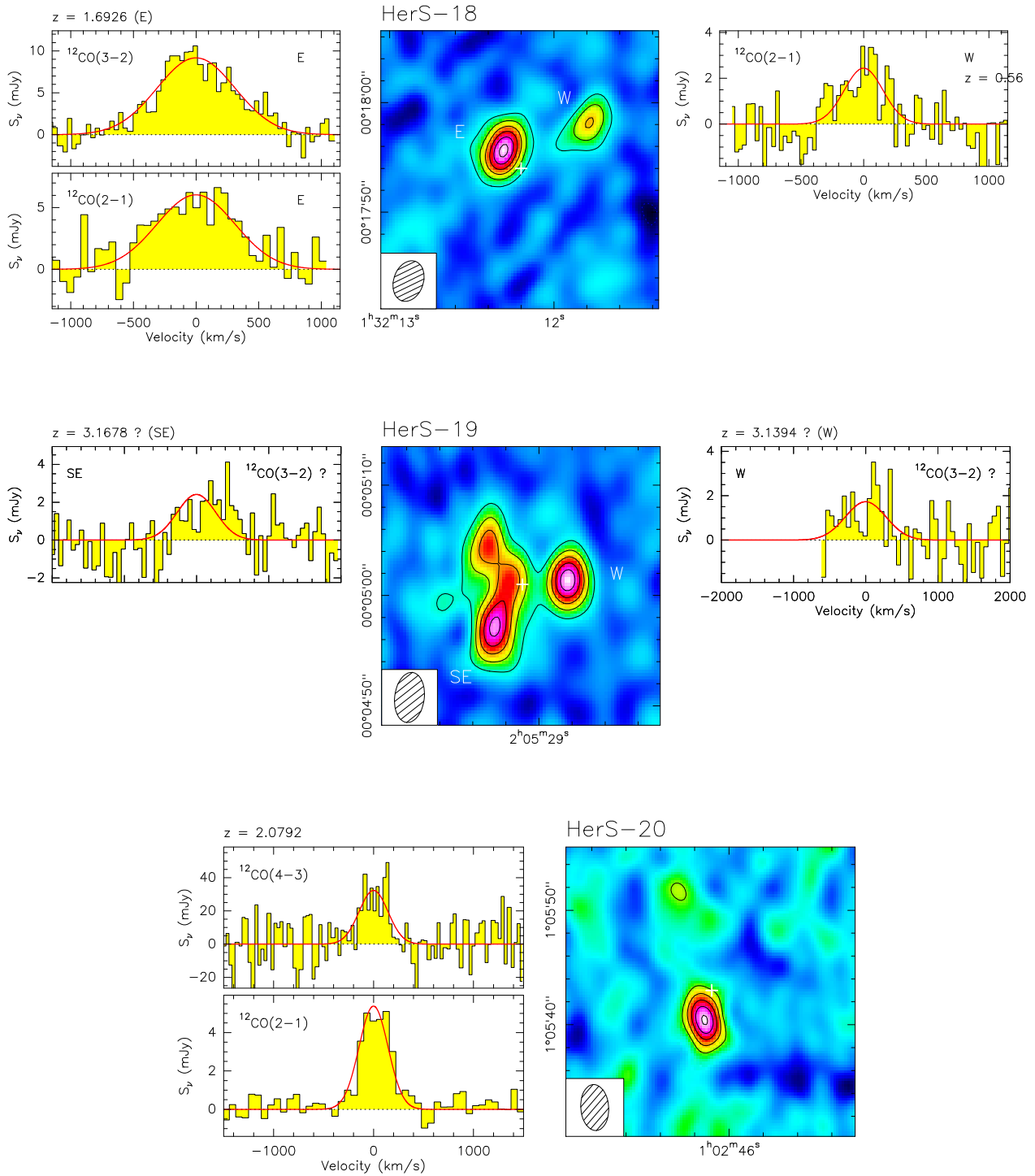


Fig. A.2. continued.

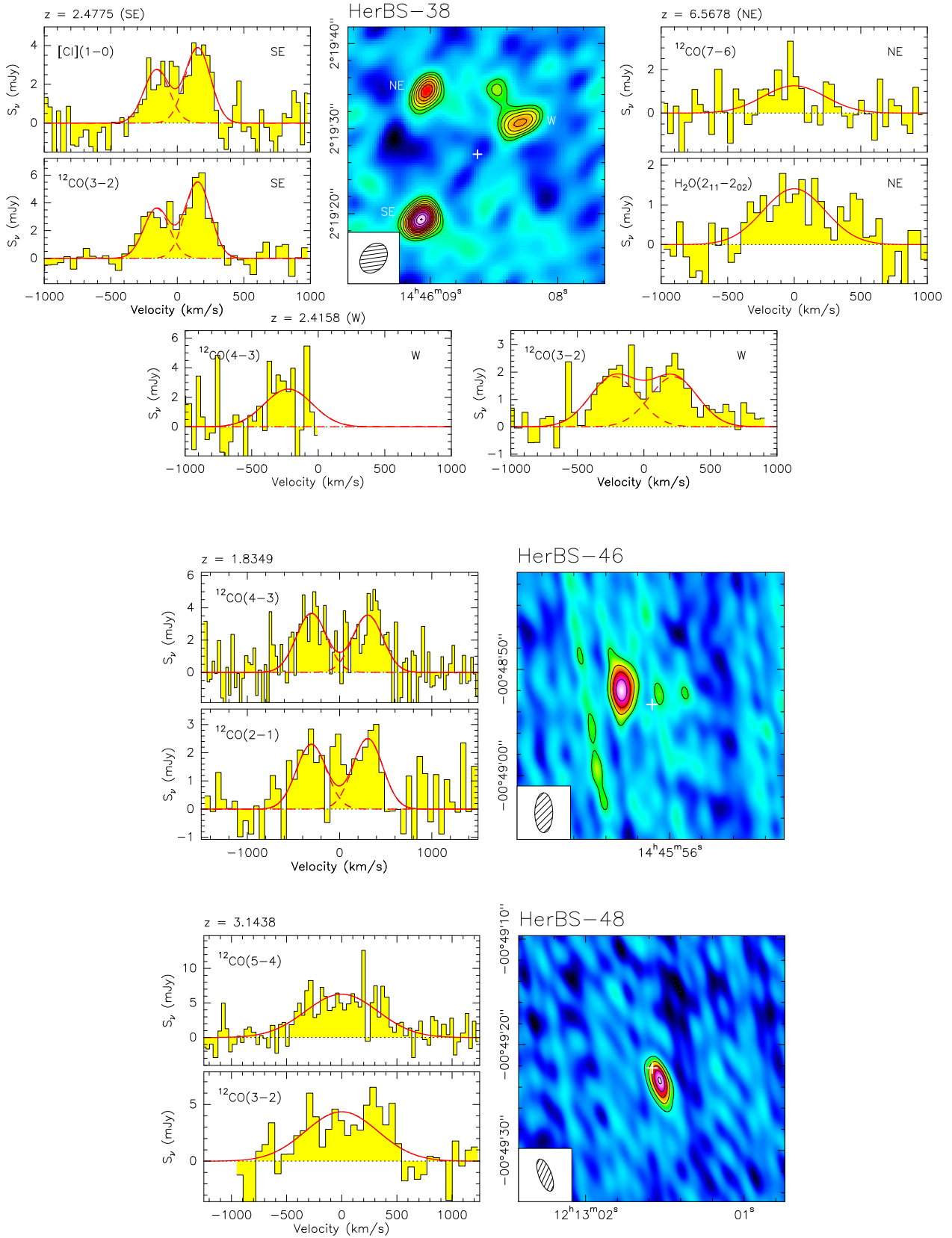


Fig. A.3. Combined continuum and emission lines maps of the z -GAL sources selected from the HerBS sample and spectra of each of the detected emission lines. Further details are provided in Appendix A.1.

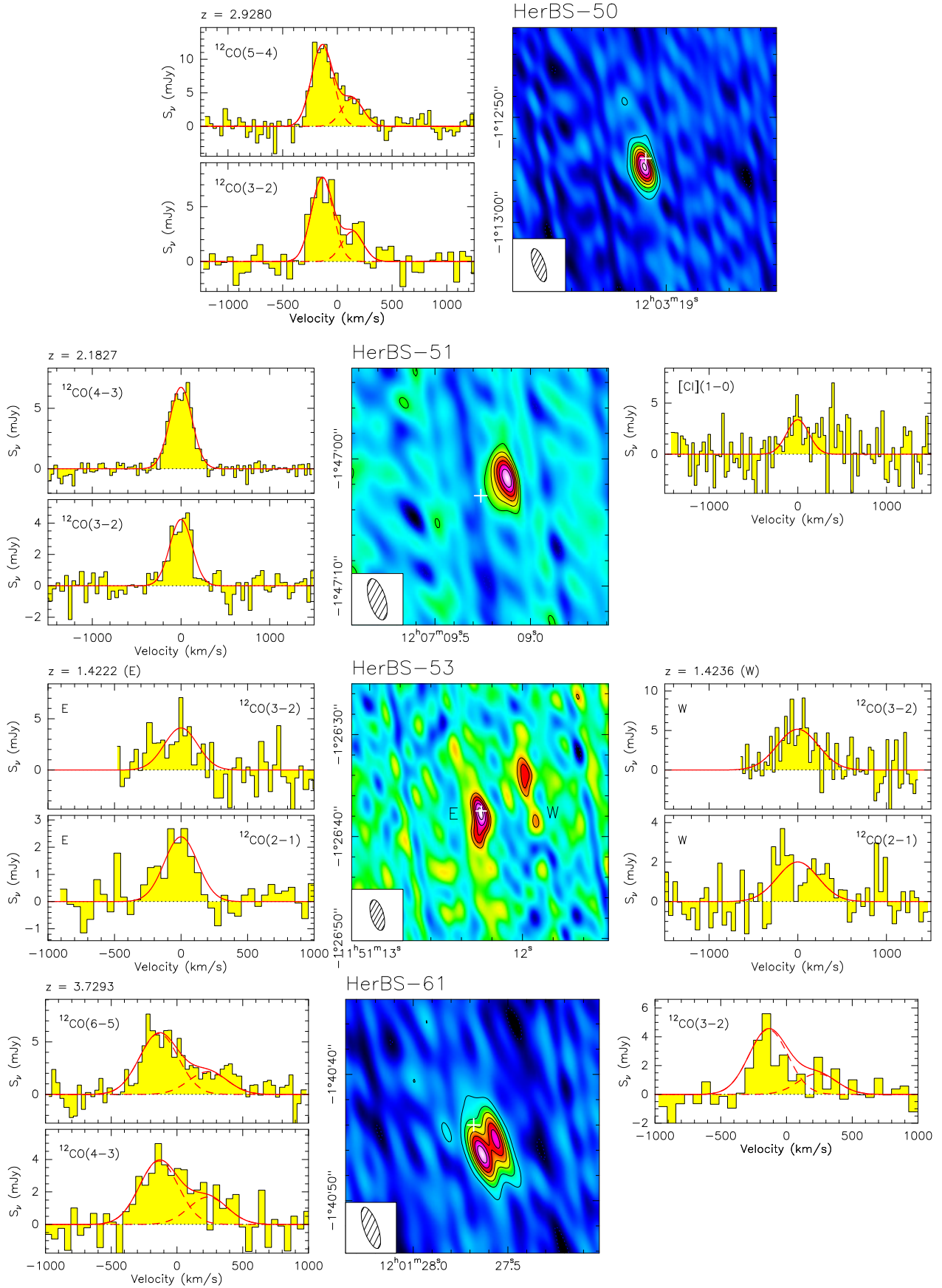


Fig. A.3. continued.

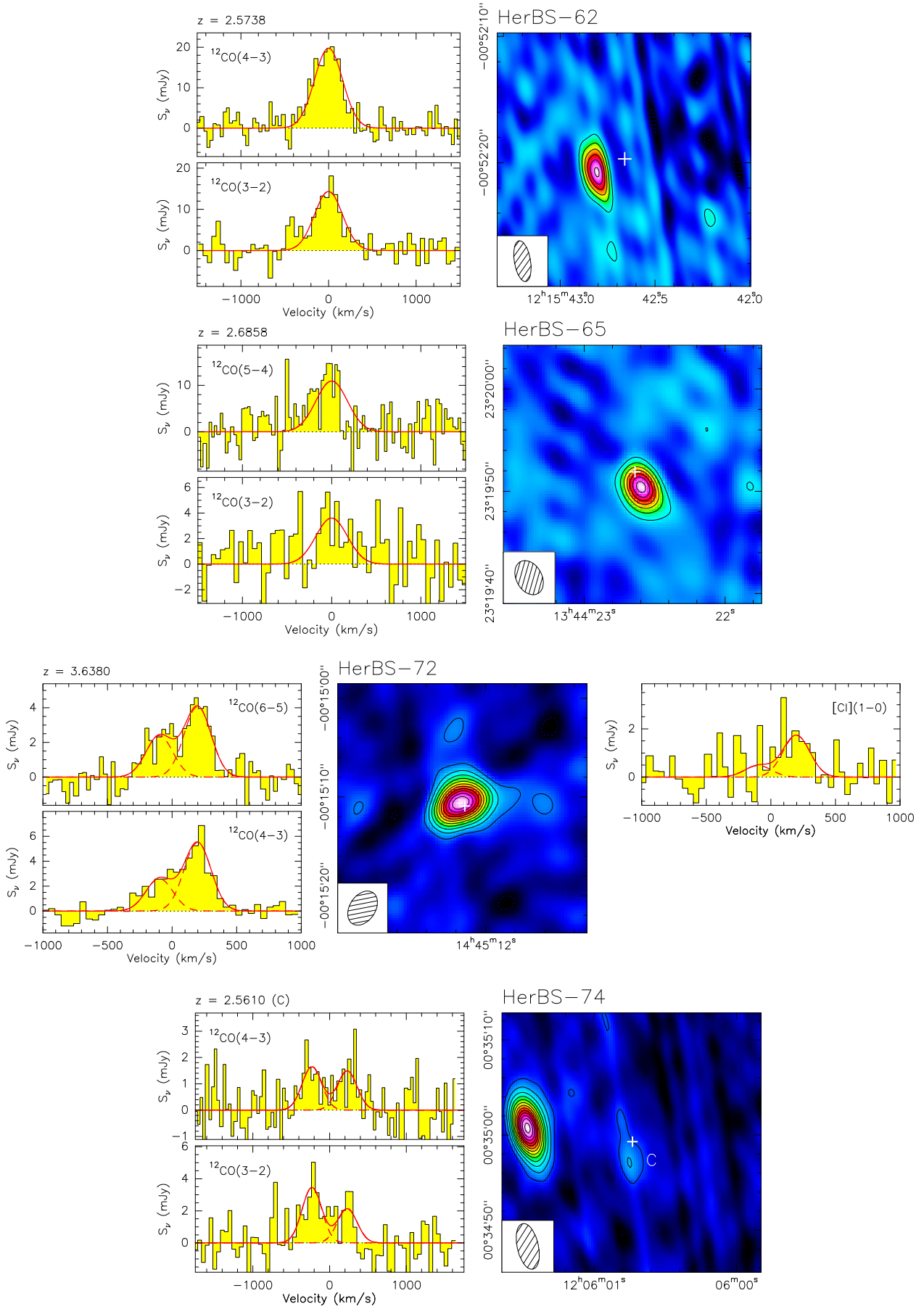


Fig. A.3. continued.

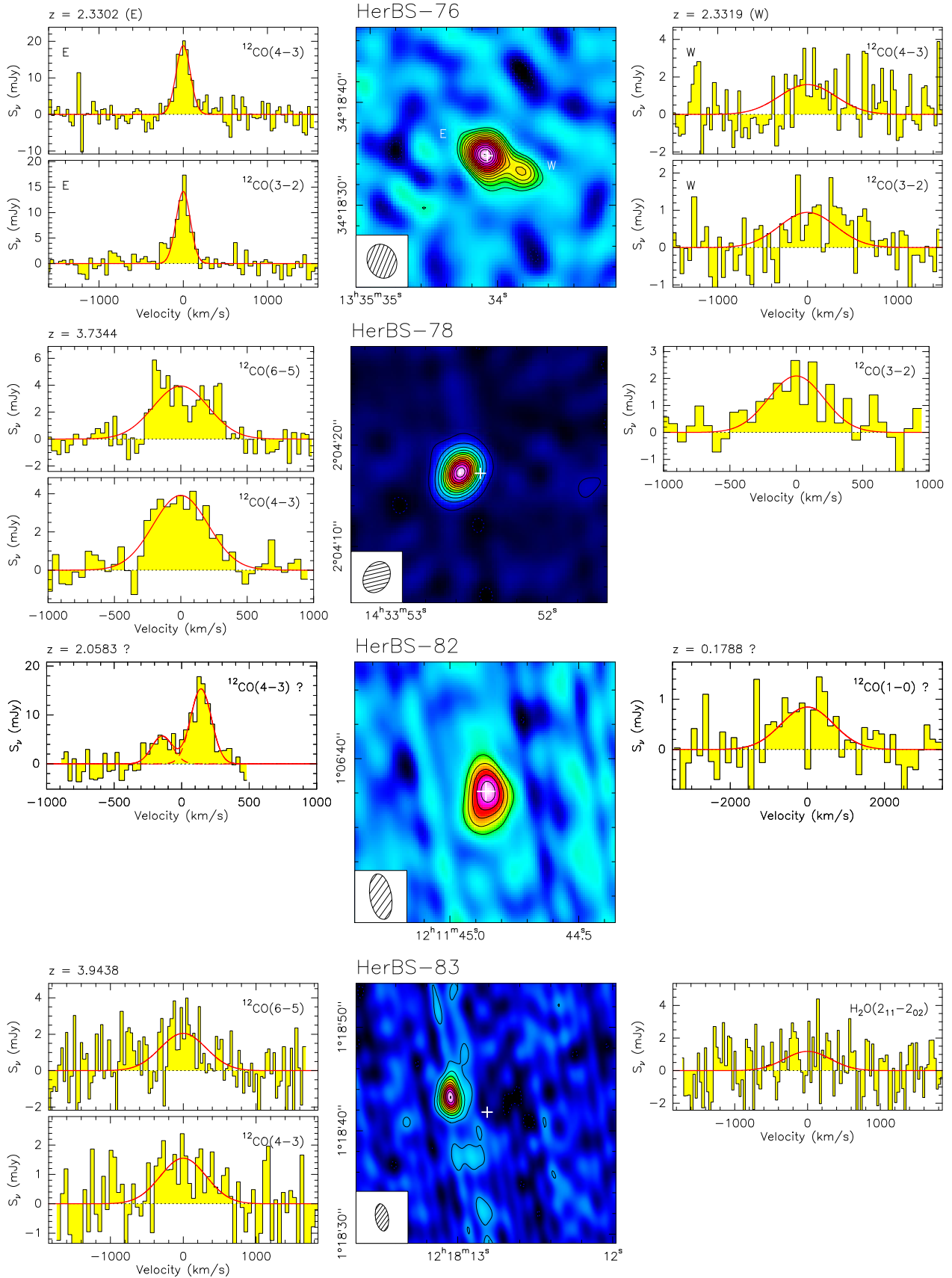


Fig. A.3. continued.

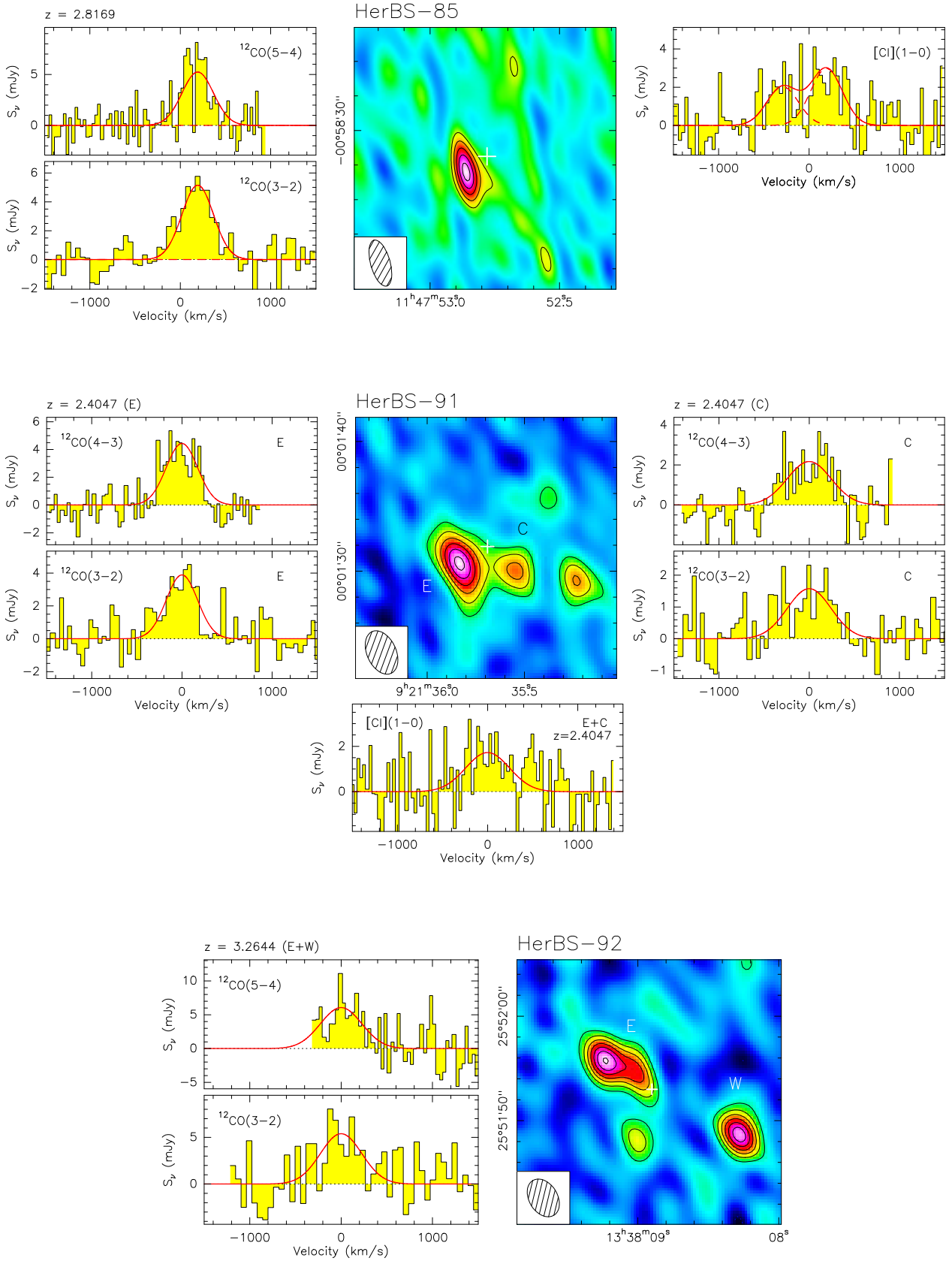


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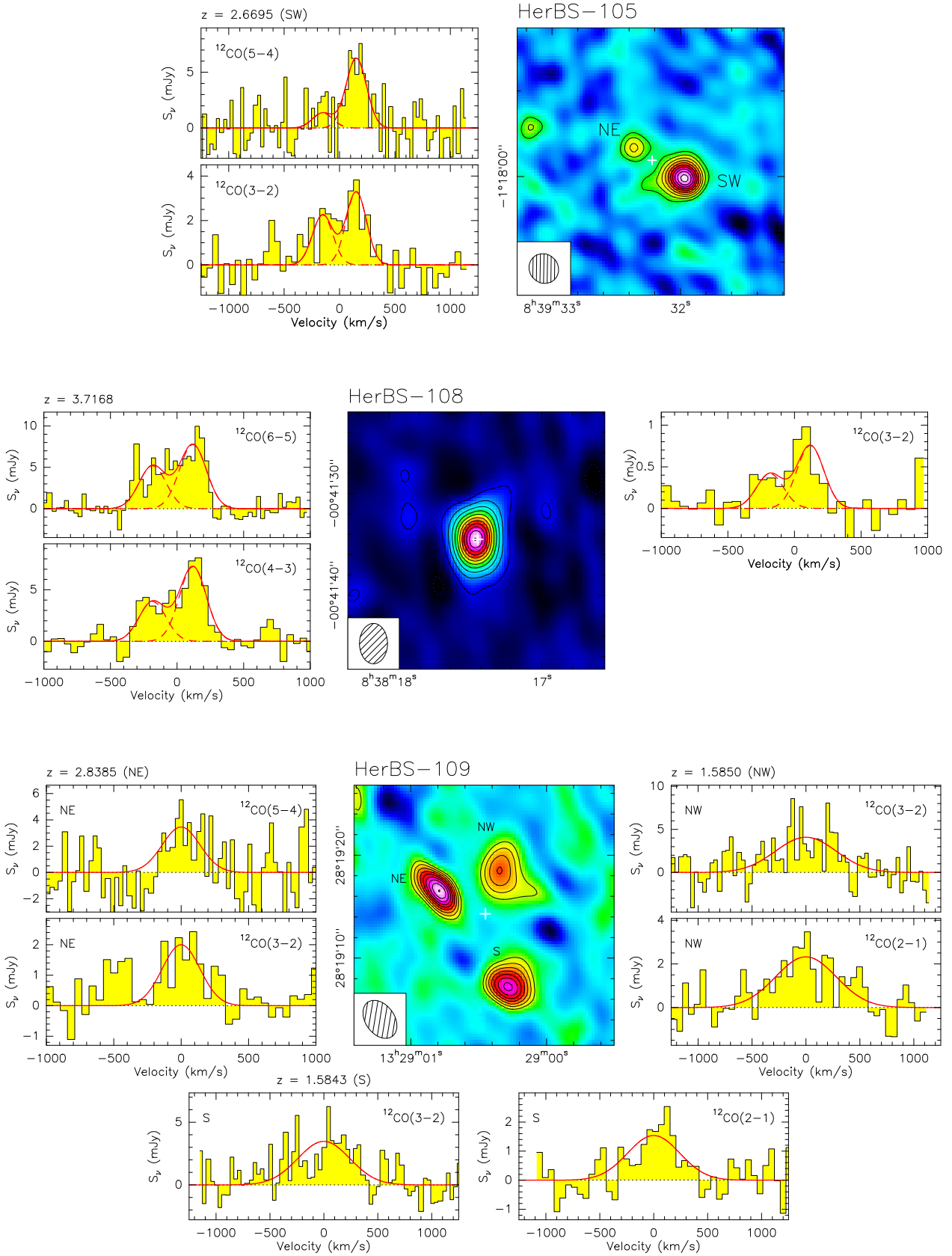


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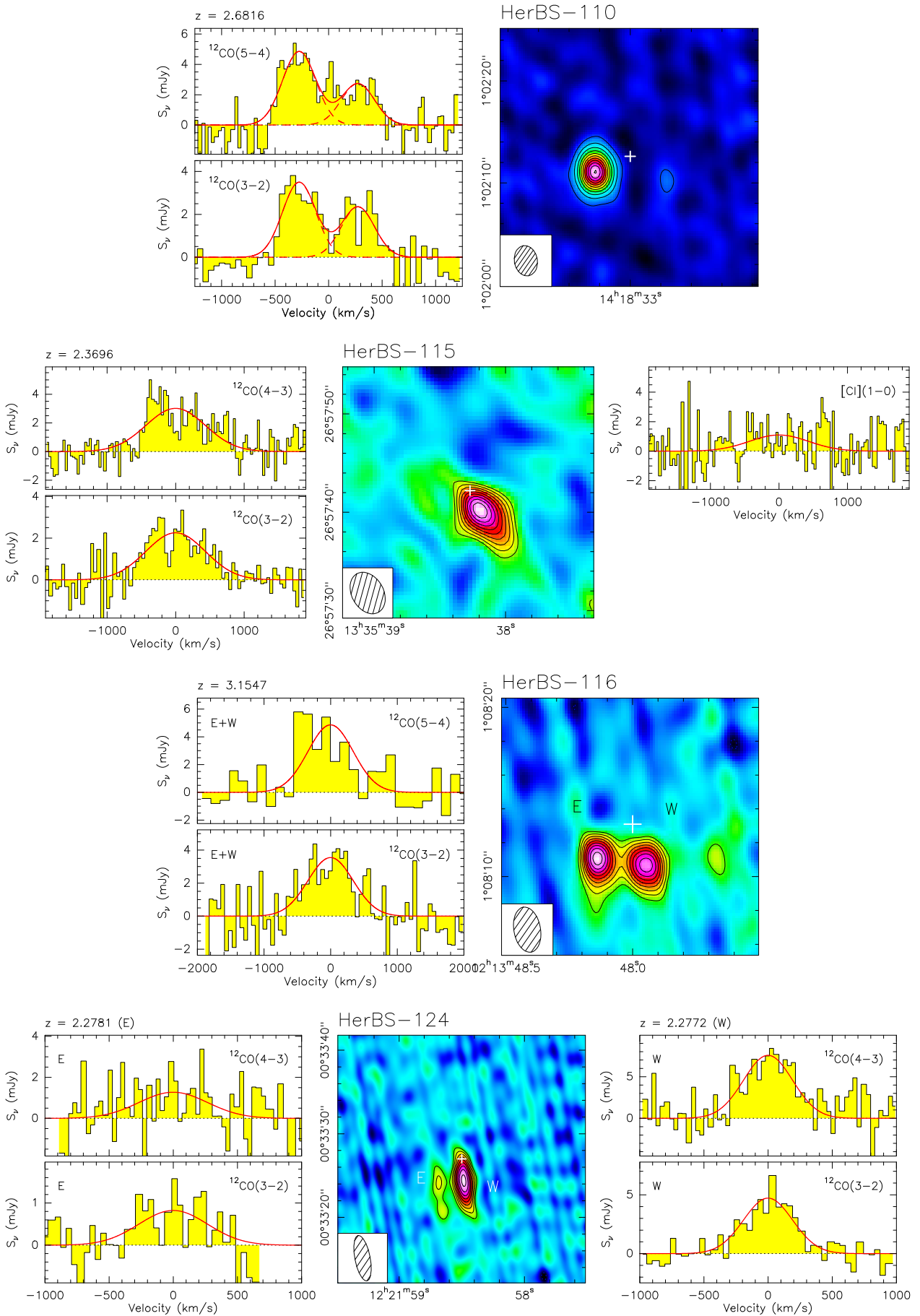


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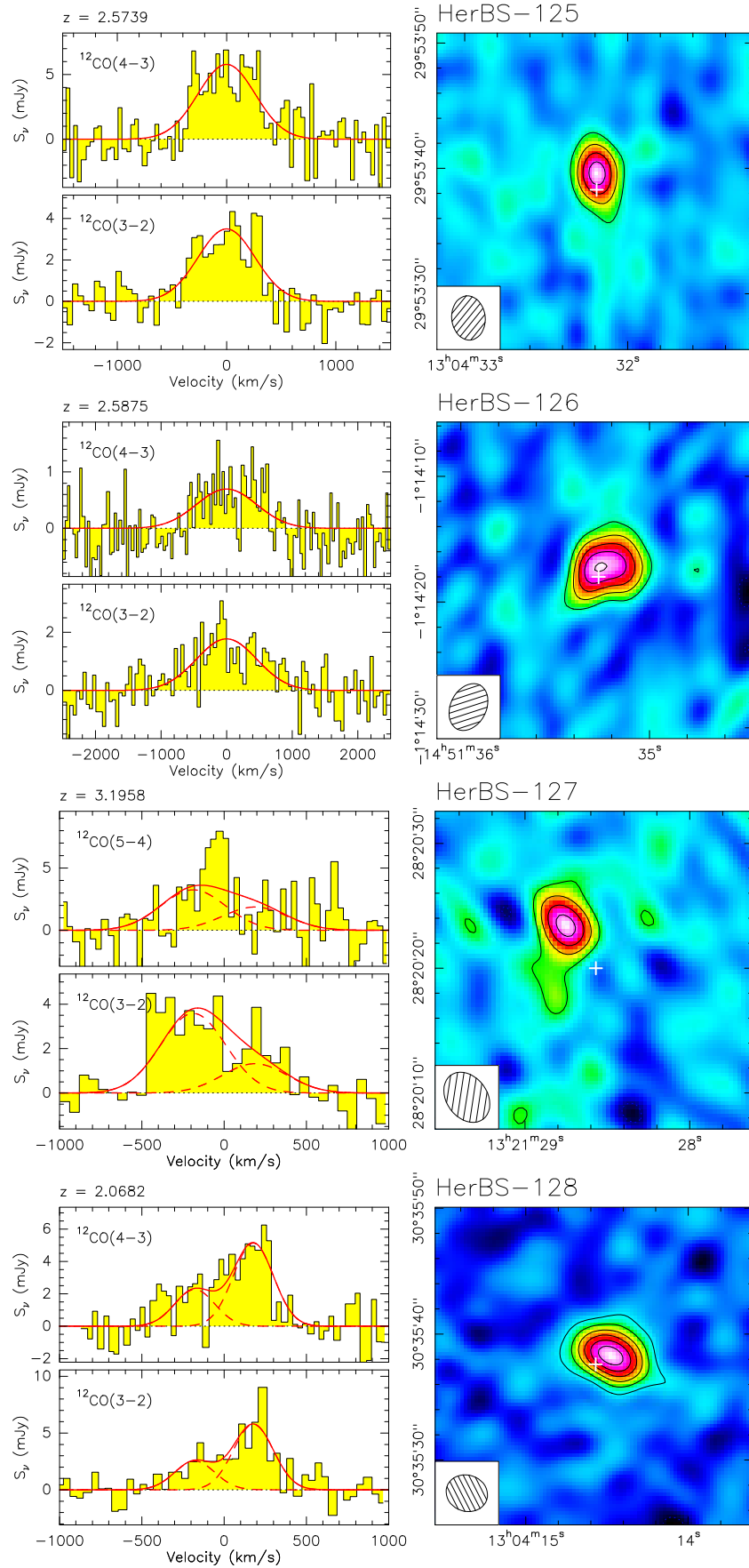


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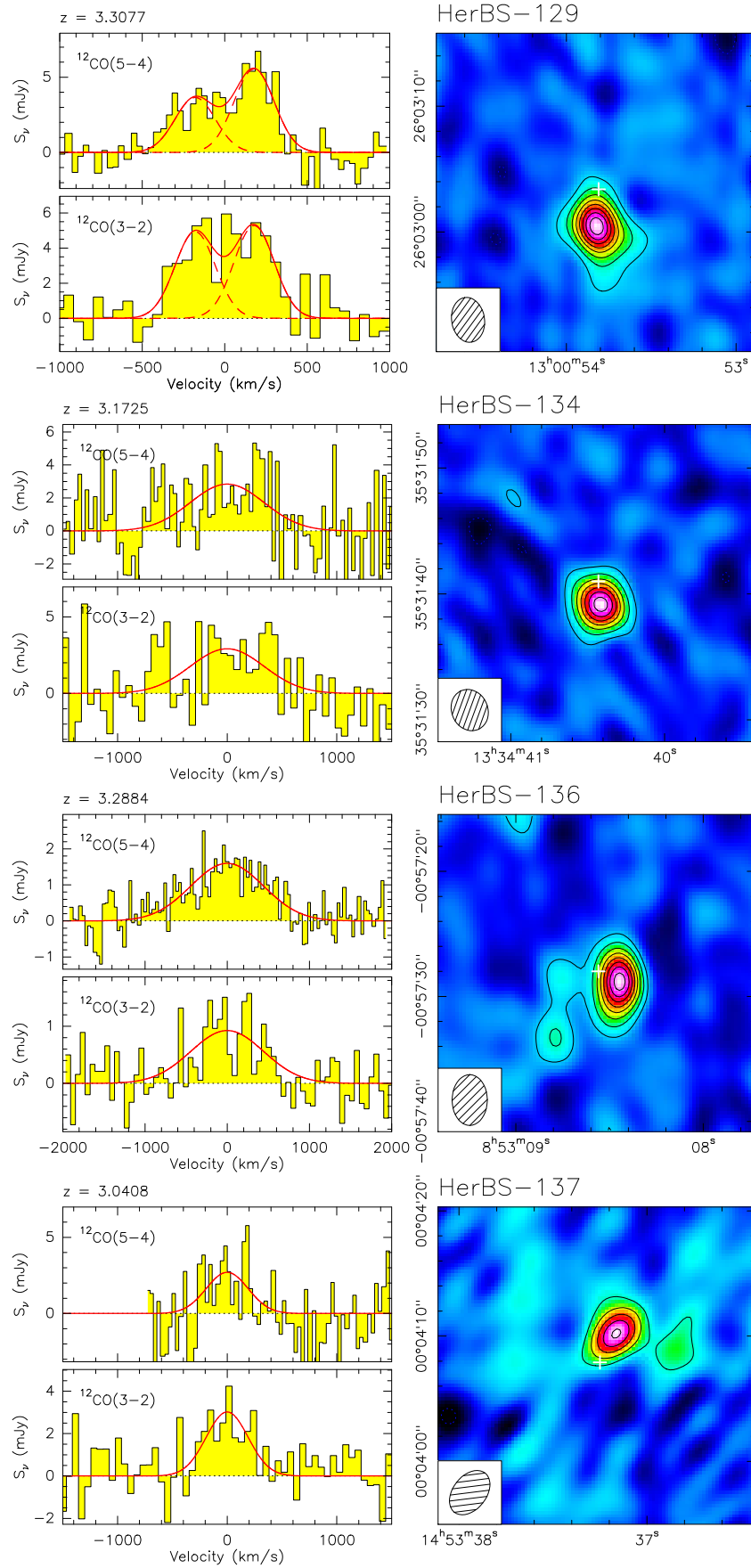


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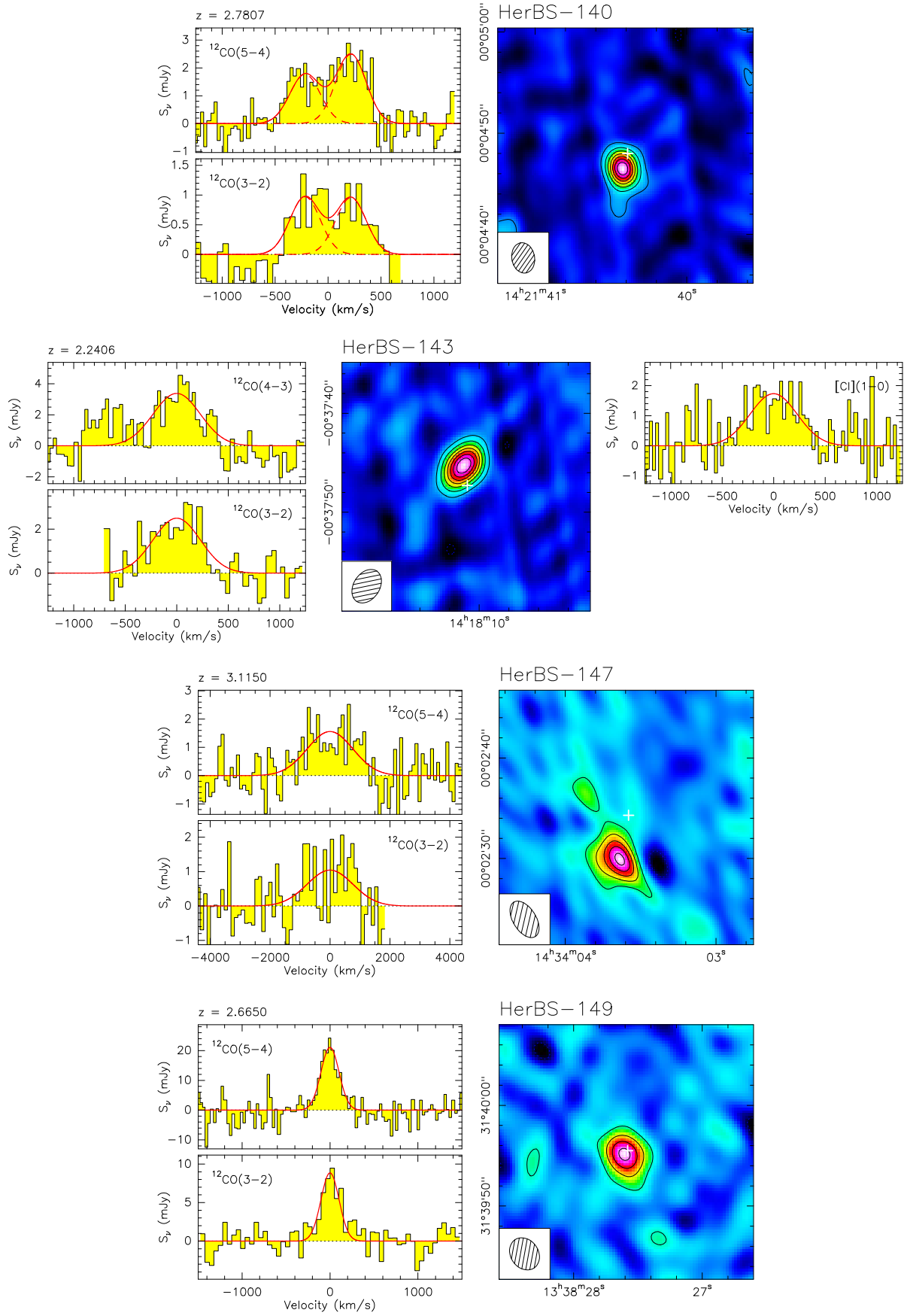


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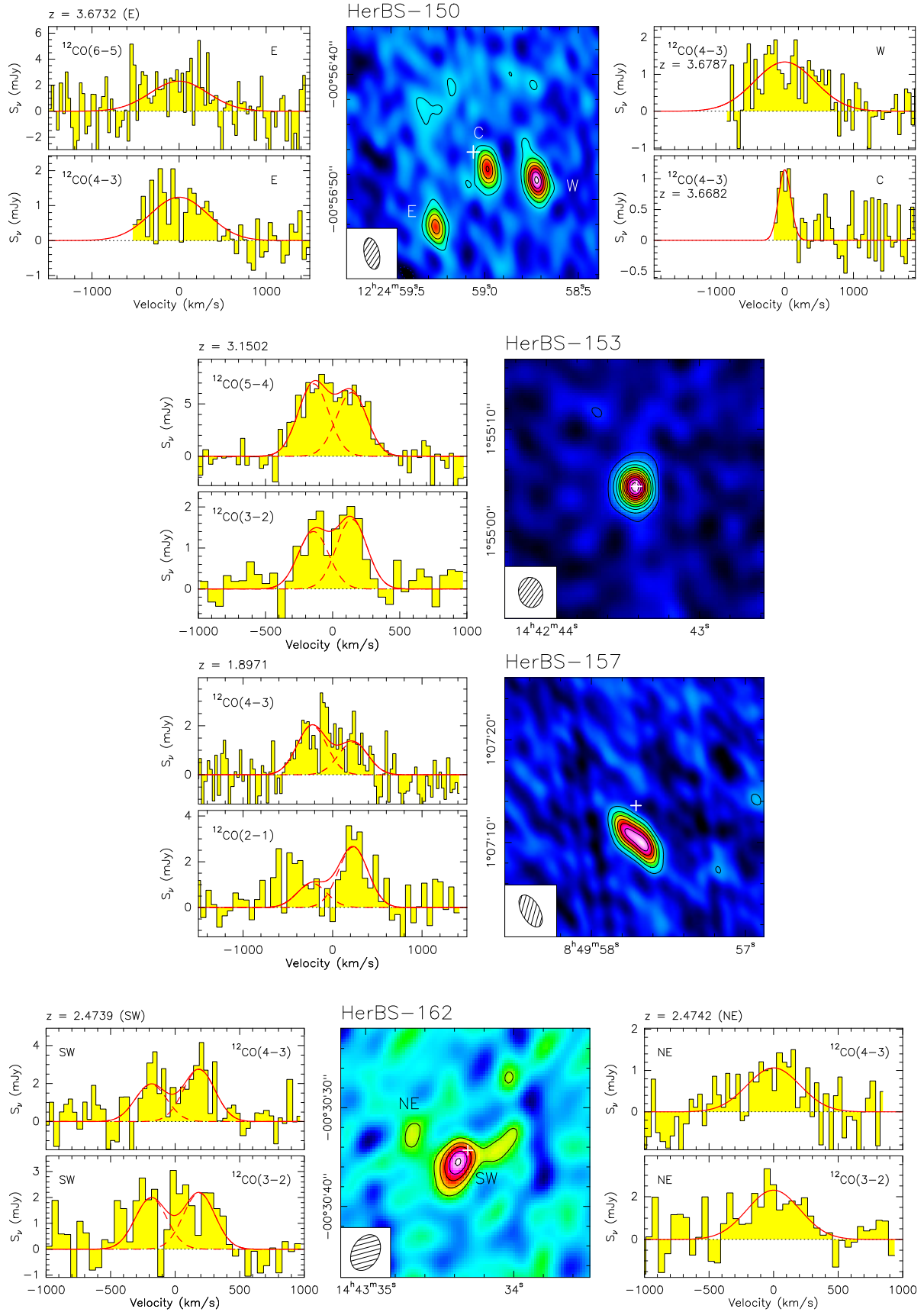


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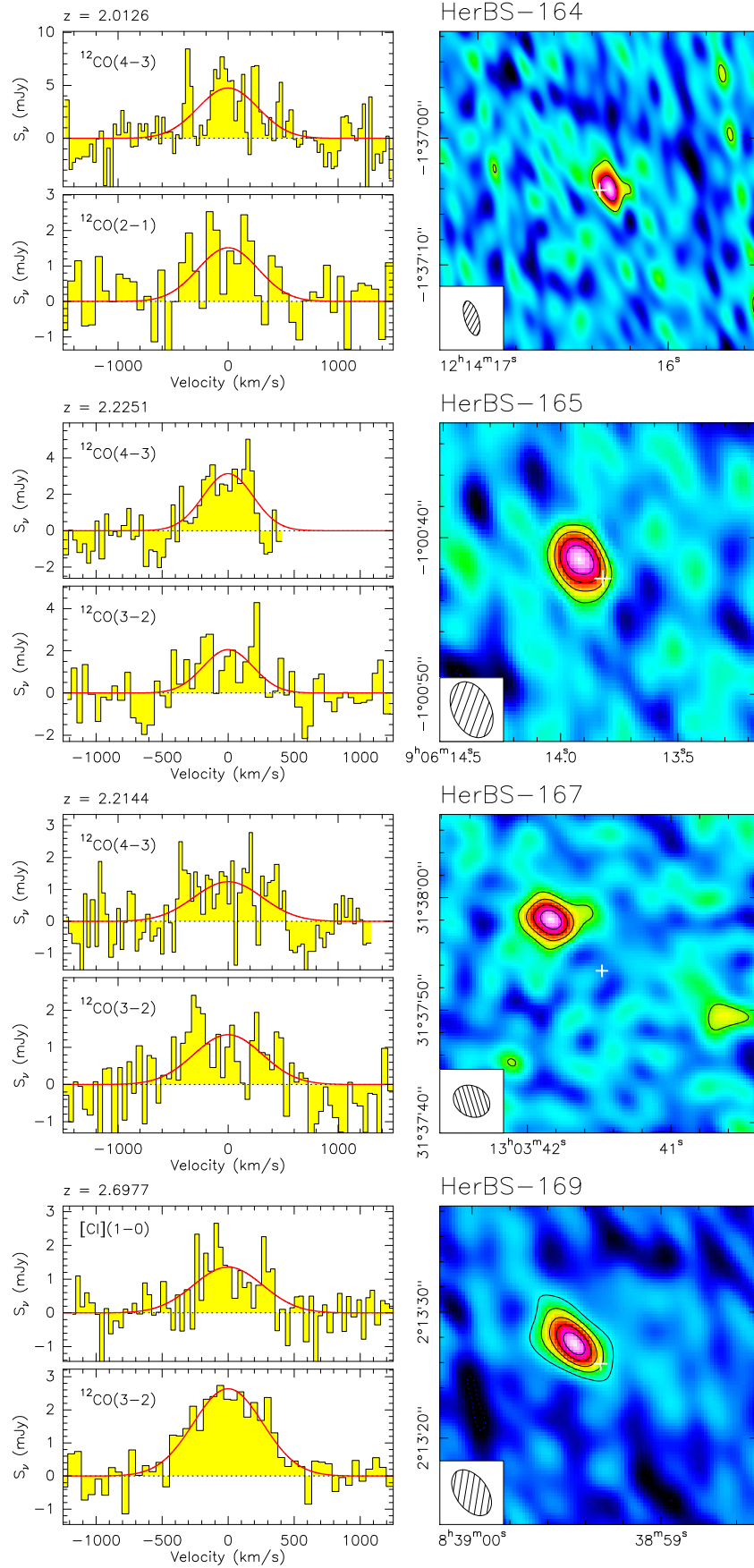


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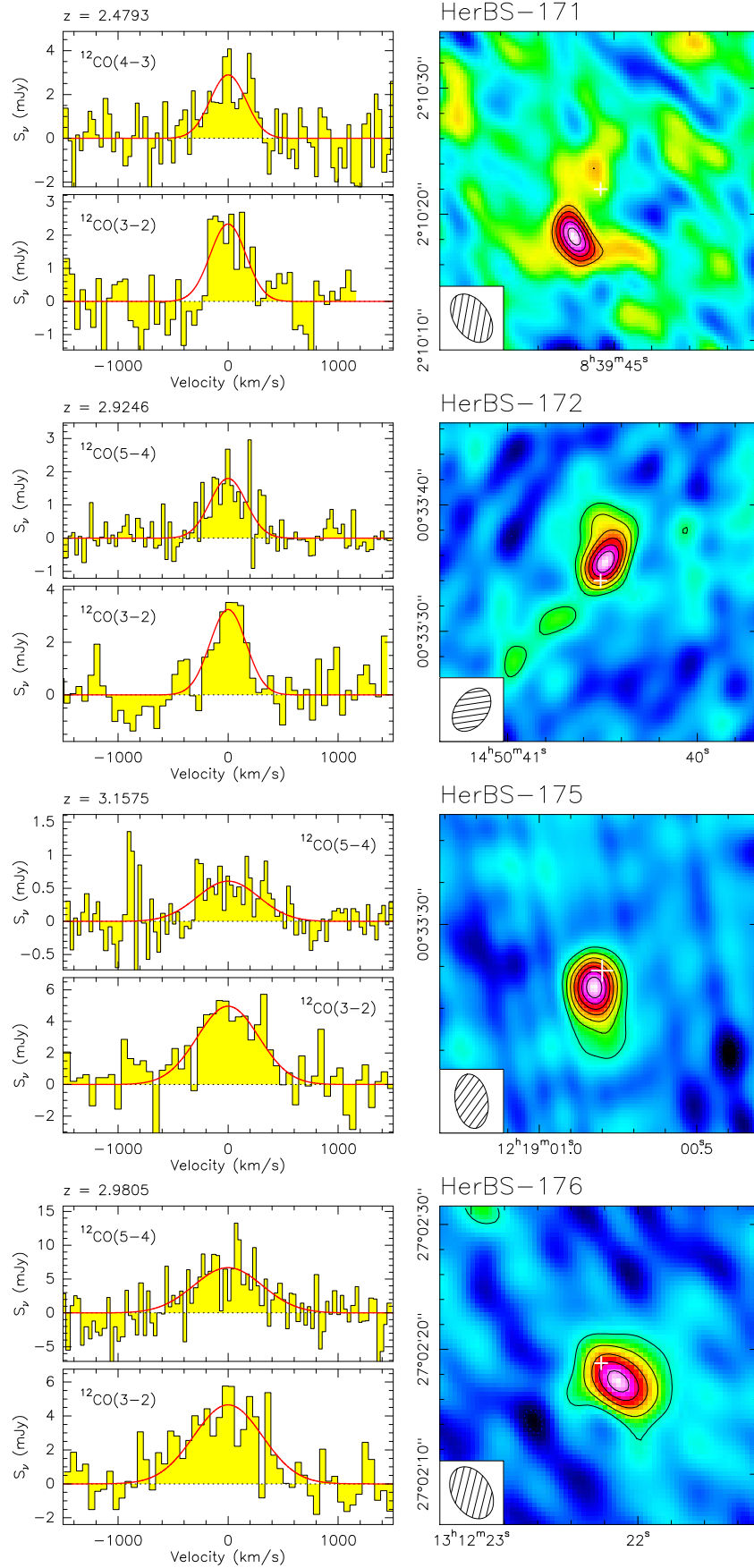


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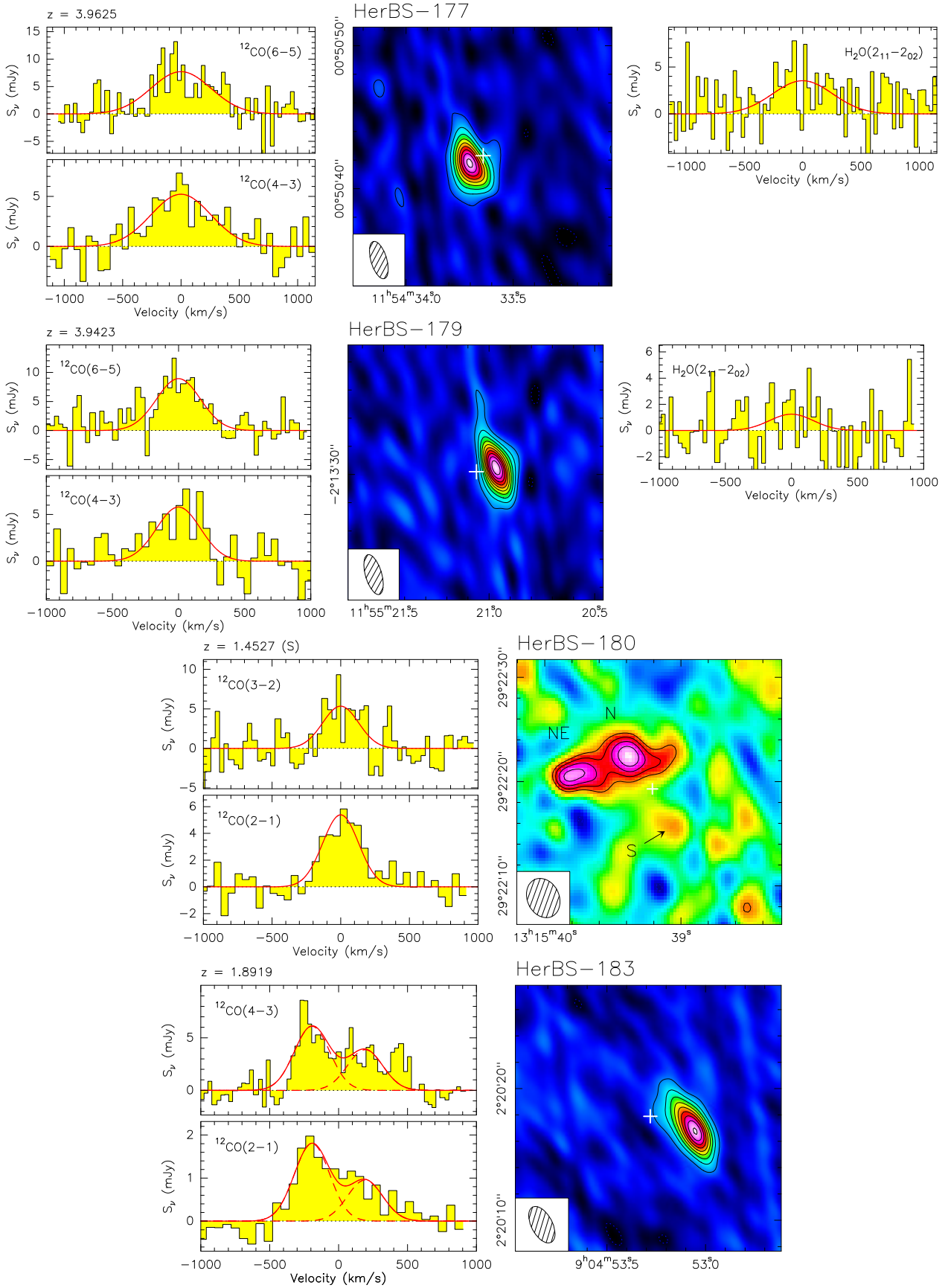


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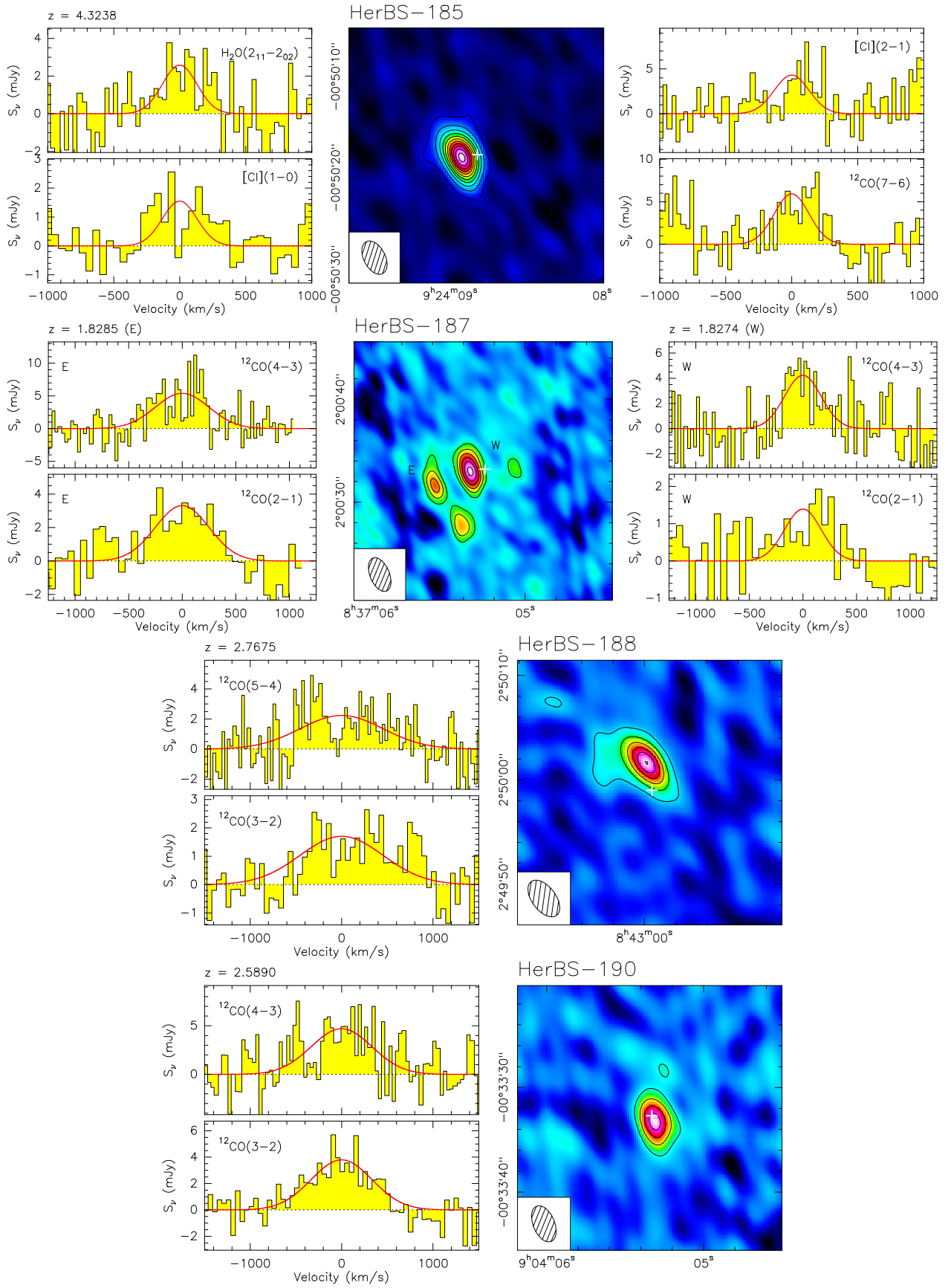


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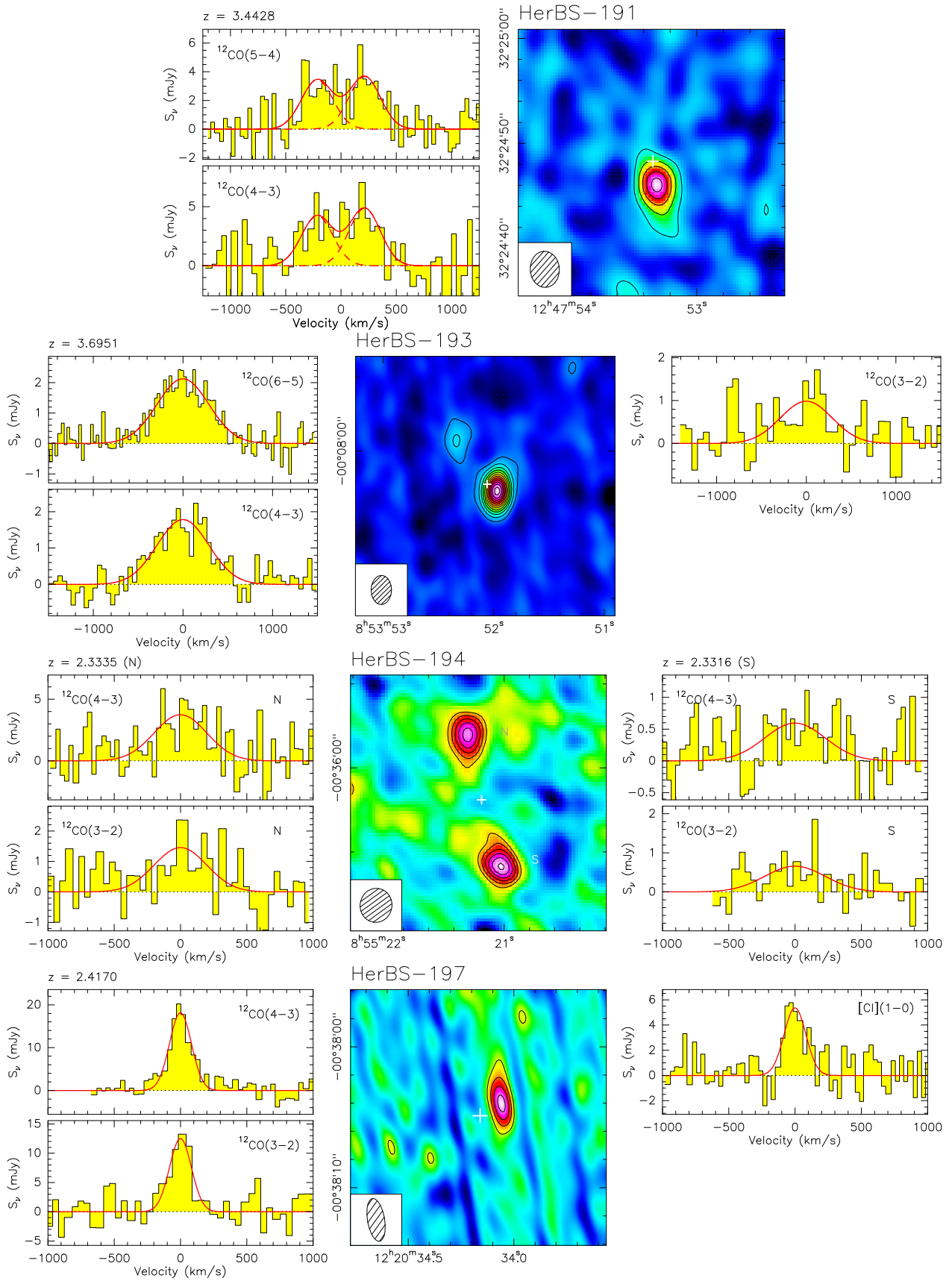


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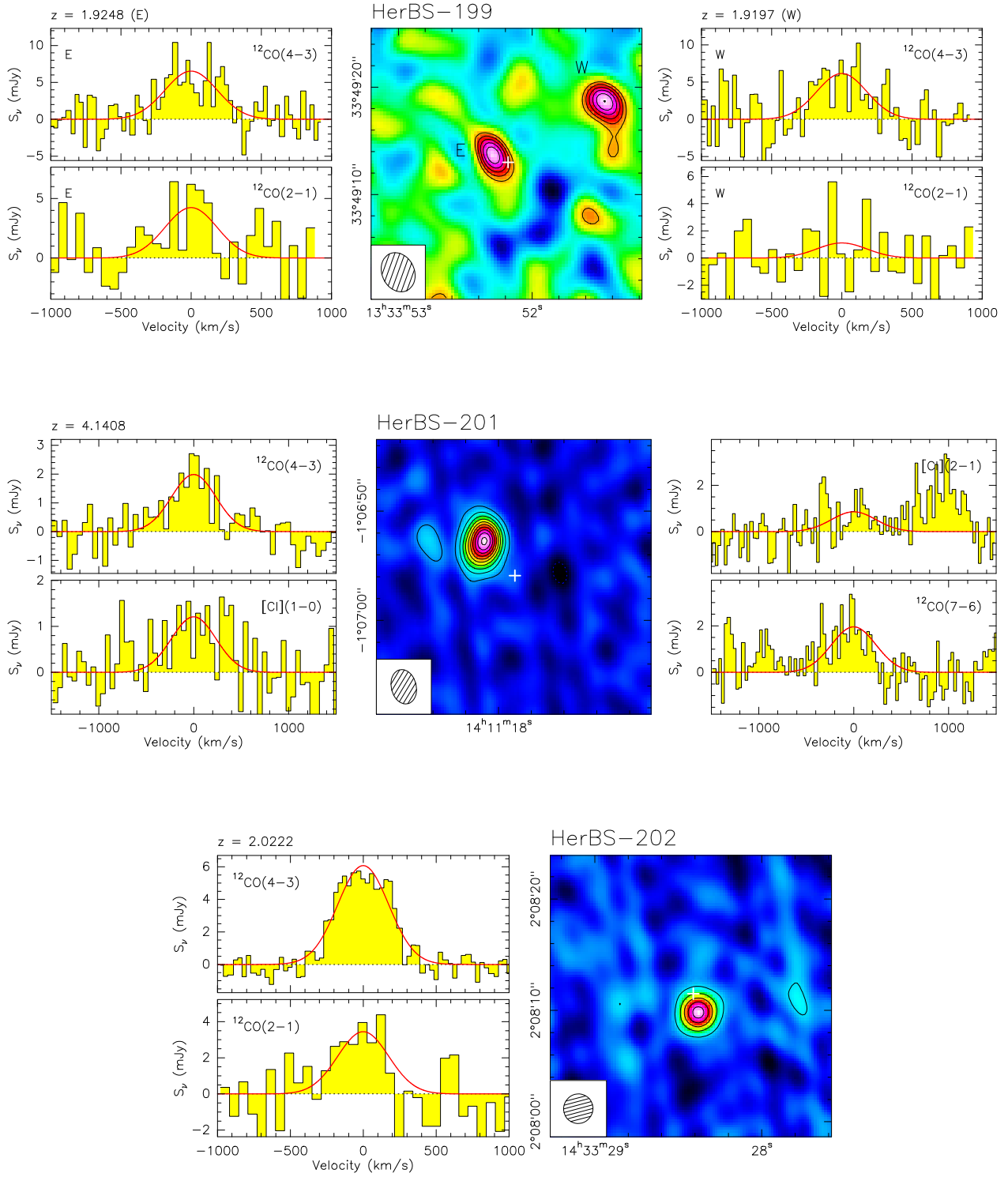


Fig. A.3. continued.

