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A three-component giant radio halo: The puzzling case of the galaxy cluster Abell 2142

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ABSTRACT

Context. Turbulence introduced into the intracluster medium (ICM) through cluster-merger events transfers energy to non-thermal components, and can trigger the formation of diffuse synchrotron radio sources. Typical diffuse sources in the form of giant radio halos and mini-halos are found in merging and relaxed cool-core galaxy clusters, respectively. On the other hand, recent observations reveal an increasing complexity to the non-thermal phenomenology.

Aims. Abell 2142 (A2142) is a mildly disturbed cluster that exhibits uncommon thermal and non-thermal properties. It is known to host a hybrid halo consisting of two components (H1 and H2), namely a mini-halo-like and an enigmatic elongated radio halo-like structure. We aim to investigate the properties, origin, and connections of each component.

Methods. We present deep LOFAR observations of A2142 in the frequency ranges 30–78 MHz and 120–168 MHz. With complementary multi-frequency radio and X-ray data, we analysed the radio spectral properties of the halo and assessed the connection between the non-thermal and thermal components of the ICM.

Results. We detect a third radio component (H3), which extends over the cluster volume on scales of ~ 2 Mpc, embeds H1 and H2, and has a morphology that roughly follows the thermal ICM distribution. The radio spectral index is moderately steep in H1 ($\alpha = 1.09 \pm 0.02$) and H2 ($\alpha = 1.15 \pm 0.02$), but is steeper ($\alpha = 1.57 \pm 0.20$) in H3. Our analysis of the thermal and non-thermal properties allowed us to discuss possible formation scenarios for each radio component. Turbulence from sloshing motions of low-entropy gas on different scales may be responsible for the origin of H1 and H2. We classified H3 as a giant ultra-steep spectrum radio halo, and find that it may trace the residual activity from an old energetic merger and/or inefficient turbulent reacceleration induced by ongoing minor mergers.

Key words. radiation mechanisms: thermal – radiation mechanisms: non-thermal – acceleration of particles – large-scale structure of Universe – galaxies: clusters: intracluster medium – galaxies: clusters: individual: Abell 2142

1. Introduction

Galaxy clusters are the largest gravitationally bound structures in the Universe and accrete mass through mergers and infalling matter from the filaments of the cosmic web (e.g. Kravtsov & Borgani 2012). The baryonic content is dominated by the intracluster medium (ICM), which is the hot ($T \sim 10^7$ – 10^8 K) and rarefied ($n_e \sim 10^{-2}$ – 10^{-4} cm⁻³) plasma that fills the space in between the galaxy cluster members, and emits through thermal bremsstrahlung in the X-ray band (Sarazin 1986). Magnetic fields and relativistic particles are present in the ICM. These can emit synchrotron radiation that generates a variety of diffuse radio sources, such as giant radio halos and mini-halos that exhibit typically steep spectra¹ in the range $\alpha \sim 1$ – 1.3 (see van Weeren et al. 2019 for a review).

Statistical studies show that the mass and the dynamical state of the host cluster are key to understanding the processes that trigger the formation of a diffuse radio source (e.g. Cassano et al. 2010a, 2013, 2023; Kale et al. 2013; Cuciti et al. 2015, 2021). Giant radio halos extending on megaparsec scales are preferentially found in massive and disturbed systems, whereas mini-halos are found to be confined within the cool core (on scales ~ 100 – 500 kpc) of relaxed clusters. According to reacceleration models, a cluster merger dissipates a fraction of its energy in the ICM through turbulence, which can reaccelerate electrons and amplify magnetic fields, thus generating a radio halo (e.g. Brunetti et al. 2001; Petrosian 2001; Brunetti & Lazarian 2007, 2016; Beresnyak et al. 2013; Miniati 2015). These models predict a wide range of spectral properties, including a large number of halos with very steep spectra, namely $\alpha \gtrsim 1.5$ (also referred to as ultra-steep spectrum radio halos), which are rare at high frequencies (~ 1 GHz), but should be more common at lower

¹ We define the spectral index from the flux density as $S_\nu \propto \nu^{-\alpha}$.

frequencies (hundreds of MHz; e.g. [Cassano et al. 2010b, 2023](#)). These can be old radio halos, in which the bulk of turbulence has already been dissipated, and/or may trace inefficient turbulent reacceleration associated with a minor merger and/or low-mass clusters. In line with the key predictions of the turbulent reacceleration scenario, low-frequency observations reveal very steep spectrum radio halos in a number of clusters ([Brunetti et al. 2008; Dallacasa et al. 2009; Macario et al. 2010; Bonafede et al. 2012; Venturi et al. 2017; Bruno et al. 2021; Di Gennaro et al. 2021; Duchesne et al. 2021; Edler et al. 2022; Pasini et al. 2022; Rajpurohit et al. 2023](#)).

Current data support turbulent reacceleration models, but the complex energy transfer mechanisms from megaparsec scales down to smaller scales are still poorly understood (e.g. [Brunetti & Jones 2014](#), for a review). Furthermore, the origin of the seed electrons that are reaccelerated by turbulence is unclear. It is possible that these are old, mildly relativistic populations of primary cosmic ray electrons (CRE) injected by AGN (e.g. [Fujita et al. 2007; Bonafede et al. 2014a; van Weeren et al. 2017](#)), or secondary CRE produced by hadronic collisions between protons in the ICM (e.g. [Brunetti & Blasi 2005; Brunetti & Lazarian 2011; Brunetti et al. 2017; Pinzke et al. 2017](#)). A similar scenario to that of giant halos was invoked for mini-halos, but in this case turbulence may be induced by sloshing of cold gas in the cluster core following perturbation by a minor merger (e.g. [Mazzotta & Giacintucci 2008; Fujita & Ohira 2013; ZuHone et al. 2013](#)). In the context of hadronic models (e.g. [Dennison 1980; Blasi & Colafrancesco 1999; Dolag & EnBlin 2000; Pfrommer et al. 2008](#)), mini-halos may also trace emission from fresh populations of secondary CRE, without requiring any reacceleration process ([Pfrommer & EnBlin 2004](#)).

Sensitive radio observations are revealing an increasing complexity to the phenomenology of diffuse sources. The distinction between halos and mini-halos is being challenged by the discovery of giant halos in relaxed systems ([Bonafede et al. 2014b; Kale et al. 2019; Raja et al. 2020](#)) and ‘hybrid’ halos with an inner mini-halo-like component and an outer radio halo-like component ([Farnsworth et al. 2013; Venturi et al. 2017; Savini et al. 2018, 2019; Biava et al. 2021](#)). Furthermore, diffuse radio emission was recently detected up to the canonical cluster outskirts (e.g. [Shweta et al. 2020; Rajpurohit et al. 2021a; Botteon et al. 2022; Cuciti et al. 2022](#)) and between pairs of clusters in pre-merger phases ([Botteon et al. 2018, 2020; Govoni et al. 2019; de Jong et al. 2022](#)), confirming the presence of non-thermal components and reacceleration processes taking place on such large scales. In particular, by means of Low Frequency ARray (LOFAR) observations below ~ 200 MHz, [Cuciti et al. \(2022\)](#) reported the discovery of a new type of diffuse source in four merging clusters (ZwCl 0634.1+4750, A665, A697, and A2218) that surrounds classical giant radio halos, which are referred to as megahalos. Megahalos extend up to scales of $\sim R_{500}^2$, exhibit a shallower surface brightness radial profile than that of their embedded radio halo, and are characterised by a very steep spectrum ($\alpha \gtrsim 1.6$ between ~ 50 – 150 MHz).

Studying targets with uncommon features can shed light on the physical properties that underlie standard classification schemes (e.g. relaxed and merger) for galaxy clusters; we study them here because of their potential to provide information about the origin of hybrid sources and the physical and/or evolutionary

connection between its radio components. The nearby and massive galaxy cluster Abell 2142 (hereafter A2142) is a favourable target for investigation of the role of intricate environmental and dynamical conditions at different scales in the formation and shaping of multi-component radio sources. A2142 is the main member of the A2142 supercluster; has an intermediate dynamical state that challenges the usual classification as a relaxed or merging system; is characterised by extremely complex dynamics, possibly triggered by many minor merger events; and is known to host a hybrid halo consisting of two radio components. In this work, we present our study of A2142 by means of new deep LOFAR observations that allowed us to discover an additional, ultrasteepest spectrum radio component of the hybrid halo on a large scale. We used complementary multi-frequency radio and X-ray data to probe the uncommon non-thermal and thermal properties of A2142. We discuss the origin of the only three-component radio halo discovered so far, and present the results of our search for evidence of possible emission in the form of a megahalo.

The paper is organised as follows. In Sect. 2, we describe the A2142 galaxy cluster. In Sect. 3, we present the radio and X-ray data and summarise their processing. In Sect. 4, we show the results of our analysis. In Sect. 5, we compare the properties of the hybrid halo in A2142 with those of the four recently discovered megahalos, and discuss the origin of its radio components. In Sect. 6, we summarise our work. Throughout this paper we adopted a standard Λ CDM cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$. At the cluster redshift $z = 0.0894$, the luminosity distance is $D_L = 408.6 \text{ Mpc}$ and $1'' = 1.669 \text{ kpc}$ (or $1' \sim 100 \text{ kpc}$).

2. The galaxy cluster Abell 2142

A2142 (RA_{J2000} = 15^h58^m20^s, Dec_{J2000} = 27°14′00″) is a nearby ($z = 0.0894$) galaxy cluster of mass $M_{500} = (8.8 \pm 0.2) \times 10^{14} M_\odot$ within a radius $R_{500} = 14.07 \pm 0.70 \text{ arcmin}$ ($1408.5 \pm 70.4 \text{ kpc}$ at the cluster redshift; [Planck Collaboration XXIV 2016; Tchernin et al. 2016](#)). The galaxy cluster A2142 is located at the centre of the A2142 supercluster, to which it gives the name ([Einasto et al. 2015; Gramann et al. 2015](#)). The A2142 supercluster is in turn part of an interacting system with the Corona Borealis supercluster ([Pillastrini 2019](#)). In Fig. 1 we overlay optical, X-ray, and radio images of the target from archival observations (see details below).

As derived from photometric and spectroscopical optical studies, approximately 900 galaxies within a radius of $\sim 3.5 \text{ Mpc}$ are confirmed members of A2142, and are hierarchically organised in many structures and substructures typically consisting of small groups ([Owers et al. 2011; Einasto et al. 2018; Liu et al. 2018](#)). The main and richest structure hosts the primary brightest cluster galaxy ‘BCG1’, which is located at the centre of the potential well of the cluster ([Okabe & Umetsu 2008; Wang & Markevitch 2018](#)). The dynamics of A2142 is extremely complex owing to several ongoing minor mergers, as galaxy groups of about ten members belonging to various substructures are infalling towards the cluster centre (e.g. [Owers et al. 2011; Eckert et al. 2017a; Liu et al. 2018](#)). At a projected distance of $\sim 180 \text{ kpc}$ from BCG1, the secondary brightest cluster galaxy, BCG2, has a high peculiar velocity (e.g. [Oegerle et al. 1995](#)) and is likely the main member of another merging group. Even though the cluster is still accreting through these minor mergers, [Einasto et al. \(2018\)](#) compared the present mass of A2142 and the evolution of simulated cosmic large structures and derived an estimate of the half-mass epoch;

² R_{500} is the radius enclosing $500\rho_c(z)$, where $\rho_c(z)$ is the critical density of the Universe at a given redshift; the corresponding mass is M_{500} .

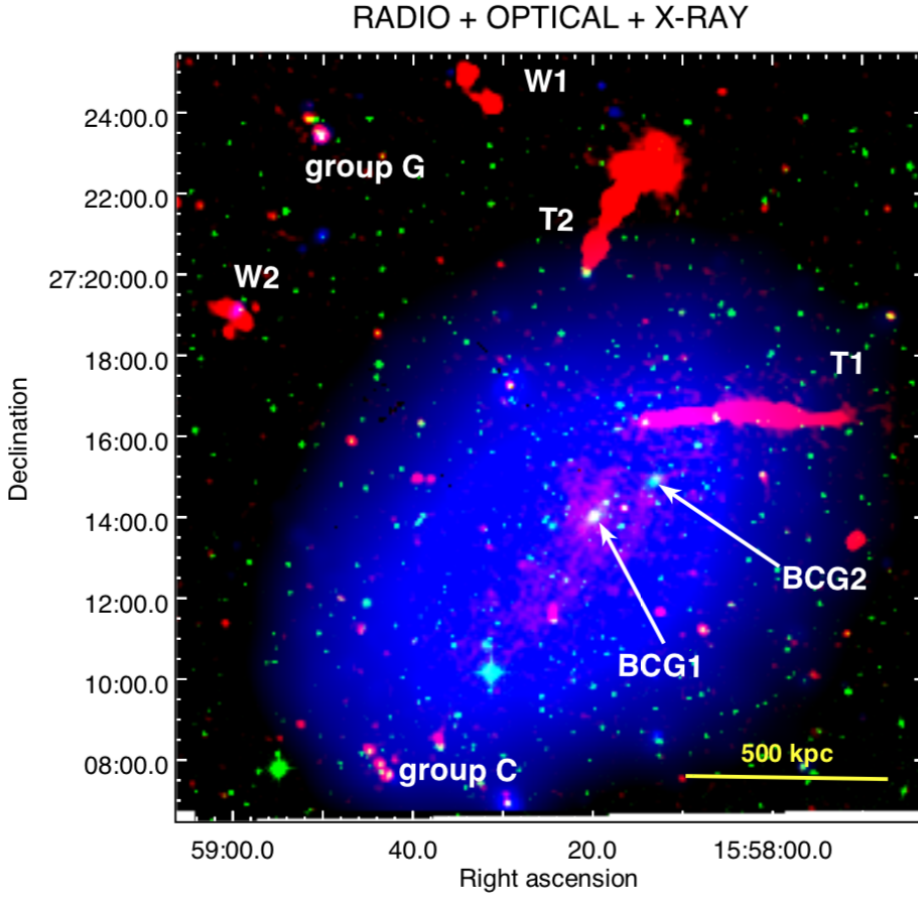


Fig. 1. Composite RGB image of A2142: radio (GMRT, 323 MHz) in red, optical (DSS-2, red filter) in green, and X-rays (*XMM-Newton*) in blue. The discrete sources discussed in the text are labelled as in Venturi et al. (2017).

according to these authors, half of the current mass of A2142 was likely accreted through merger events occurring $\gtrsim 4$ Gyr ago, thus forming the main structure of A2142.

In the X-ray band, the ICM has an elongated morphology along the NW–SE axis, and is aligned with the filamentary structure of the supercluster. The global temperature is $kT \gtrsim 9$ –10 keV, but it moderately decreases towards the central regions, where $kT \lesssim 7$ keV (e.g. Henry & Briel 1996). A2142 has morphological and thermodynamical properties intermediate between those of a relaxed cool core and an unrelaxed merging cluster (e.g. Cavagnolo et al. 2009; Rossetti et al. 2013; Tchernin et al. 2016; Cuciti et al. 2021). Even though a well-defined X-ray peak is present as in relaxed clusters, the density, temperature, and entropy of the ICM derived from its spectrum are not consistent with typical values found for cool cores, and therefore A2142 was classified by Wang & Markevitch (2018) as a rare case of a ‘warm’ core cluster (see also Rossetti & Molendi 2010; Molendi et al. 2023, for discussion on targets with similar properties).

Cold fronts are contact discontinuities detected in the X-rays as surface brightness edges, which were first discovered in A2142 by Chandra (Markevitch et al. 2000). In particular, A2142 exhibits a system of three roughly concentric cold fronts that follow a spiral-like path and are located close to the two BCGs (Markevitch et al. 2000; Markevitch & Vikhlinin 2007; Wang & Markevitch 2018). The X-ray peak of the warm core shows a prominent offset of 30 kpc with respect to the centre of the cluster (Wang & Markevitch 2018). It was suggested that an intermediate mass-ratio merger did not completely disrupt the core, but displaced a large fraction of gas from the minimum of the gravitational potential; this event caused extreme sloshing

motions of the cool gas, which then induced the formation of the cold fronts (Rossetti et al. 2013). A further cold front was discovered by *XMM-Newton* in the SE of the cluster at a distance of ~ 1 Mpc from the centre (Rossetti et al. 2013). The SE cold front may result from the long-term evolution of the central sloshing, which would indicate that it is a phenomenon able to move gas outwards from the core to larger radii (Rossetti et al. 2013).

In the radio band, diffuse synchrotron emission was first detected with the Very Large Array (VLA) in the central regions of the cluster (Giovannini & Feretti 2000), and then with the Green Bank Telescope (GBT) up to ~ 2 Mpc (Farnsworth et al. 2013). Deep follow-up observations at low frequencies with the Giant Metrewave Radio Telescope (GMRT) allowed Venturi et al. (2017) to classify the diffuse emission as a giant radio halo with two components, which are characterised by different morphological and spectral properties. The most compact, roundish, and brightest component (the ‘core’, ‘H1’) is spatially confined by the inner cold fronts. The faintest, most elongated, and largest component (the ‘ridge’, ‘H2’) extends in the direction of the SE cold front. Venturi et al. (2017) reported spectral indices of $\alpha \sim 1.3$ and $\alpha \sim 1.5$ between 118 and 1780 MHz, for the core and the ridge, respectively. The outermost emission revealed by the GBT at 2 Mpc is not detected with the GMRT, and therefore its properties are unconstrained. Venturi et al. (2017) suggested that both H1 and H2 may trace regions of turbulent particle reacceleration. In this case, turbulence in H1 would be generated by the dissipation of the kinetic energy of gas sloshing in the core, which is similar to the scenario invoked for the formation of mini-halos (e.g. Mazzotta & Giacintucci 2008; Fujita & Ohira 2013; ZuHone et al. 2013); turbulence in the ridge could have been

Table 1. Details of the LOFAR LBA (PI: L. Bruno, project code: LC17_012), LOFAR HBA (pointing name: P239+27 & PI: F. Vazza, project code: LC14_018), GMRT (PI: T. Venturi, project code: 23_017), uGMRT (PI: T. Venturi, project code: 33_052), and VLA (PI: D. Farnsworth, project code: 11B-156) radio data analysed in this work.

Instrument	Band name	Frequency coverage (MHz)	Observation date	On-source time (h)
LOFAR	LBA	30–78	08, 17, 23-Dec.-2021	16.0
LOFAR	HBA	120–168	15-Sept.-2018; 25, 31-Oct.-2020; 13-Nov.-2020	32.0
GMRT		307–339	27-Mar.-2013	5.0
uGMRT	Band-3	300–500	15-Mar.-2018	3.0
VLA _[C-array]	L	1000–2000	27-Apr.-2012	0.5
VLA _[D-array]	L	1000–2000	9-Oct.-2011	1.5

induced by less energetic mergers (including the same intermediate mass-ratio merger likely associated with the origin of core sloshing), or may trace the evolution of the central perturbations at larger scales. Alternatively, A2142 could be a hybrid hadronic/reacceleration radio halo, in which secondary CRE dominate the emission in H1, whereas the turbulent reaccelerated electrons dominate the emission in H2.

3. Observations and data reduction

In this section, we present the data analysed in this work and the corresponding data reduction. The details of the radio data are summarised in Table 1, while the details of *Chandra* X-ray data are summarised in Table 2.

3.1. LOFAR HBA radio data

A2142 was observed for 8 h in September 2018 in the context of the LOFAR Two Meter Sky Survey (LoTSS, pointing P239+27; Shimwell et al. 2017, 2019, 2022). Furthermore, 24 additional hours were spent on A2142 between October and November 2020. The High Band Antenna (HBA) Dutch array operating in the 120–168 MHz frequency range was employed for all the observations, with 23 core stations, 14 remote stations, and 14 international stations (the latter were not included in this work). Data were recorded with an integration time of 1 s and 64 channels (of width 3 kHz each) per sub-band, and then averaged to 16 channels per sub-band after removal of radio frequency interference (RFI). The source 3C 295 was used as a flux density scale calibrator.

All the data were processed together by means of the LOFAR Surveys KSP reduction pipelines, which perform direction-independent and direction-dependent calibration using PREFACTOR³ (van Weeren et al. 2016; Williams et al. 2016; de Gasperin et al. 2019) and ddf-pipeline⁴ v. 2.4, which makes use of DDFacet (Tasse et al. 2018) and KillMS (Tasse 2014a,b; Smirnov & Tasse 2015) to compute direction-dependent calibration of the whole LOFAR field of view (see also Tasse et al. 2021; Shimwell et al. 2022).

Following the ‘extraction and recalibration’ technique described in van Weeren et al. (2021), sources outside a square region of $25.8' \times 25.8'$ centred on the target were subtracted from the *uv*-data; we then performed eight additional cycles of amplitude and phase direction-independent self-calibration on the extracted datasets to improve the quality of the images towards A2142.

³ <https://github.com/lofar-astron/prefactor>

⁴ <https://github.com/mhardcastle/ddf-pipeline>

Table 2. Details of the *Chandra* X-ray data analysed in this work (PI of ObsID 5005: L. van Speybroeck; PI of ObsIDs 15186, 16564, 16565: M. Markevitch).

ObsID	CCDs	Observation date	Clean time (ks)
5005	S2, I0, I1, I2, I3	13-Apr.-2005	41.5
15186	S1, S2, S3, I2, I3	19-Jan.-2014	82.7
16564	S1, S2, S3, I2, I3	22-Jan.-2014	43.2
16565	S1, S2, S3, I2, I3	24-Jan.-2014	19.5

Uncertainties in the beam model of LOFAR HBA and calibration errors can introduce offsets in the flux density scale when amplitude solutions are transferred from the primary calibrator to the target (e.g. Hardcastle et al. 2021). As in LoTSS (Shimwell et al. 2022), the flux density scale was therefore set by cross-matching the LOFAR image with the NRAO VLA Sky Survey (NVSS; Condon et al. 1998), and by assuming a ratio of $S_{6C}/S_{NVSS} = 5.9124$ between the flux density of the 6C radio catalogue (Hales et al. 1988, 1990) and that of NVSS, at 150 MHz and 1.4 GHz, respectively. Due to this procedure, each of our LOFAR HBA images has to be multiplied by a flux scale correction factor of 0.8793.

3.2. LOFAR LBA radio data

A2142 was observed by LOFAR with the Low Band Antenna (LBA) in the frequency range 30–78 MHz on December 2021 for a total on-source time of 16 h. Observations were carried out in LBA Sparse Even mode by 24 core stations and 14 remote stations, with two separate beams simultaneously pointing on the target and flux density calibrator 3C 295, respectively. This observing setup both provides a larger field of view and reduces the impact of interfering sources from the first sidelobe with respect to the LBA_OUTER mode. Data were first acquired with an integration time of 1 s and 64 channels (of width 3 kHz each) per sub-band, and then averaged to 4 s and 8 channels per sub-band after removal of RFI and demixing, using a strategy similar to that used by de Gasperin et al. (2020a).

Calibration was performed by means of the Library for Low-Frequencies (LiLF⁵) pipelines. Following the process described in de Gasperin et al. (2019), gain solutions are obtained for the flux density calibrator to correct for polarisation alignment, bandpass, and clock drift, and then transferred to the target. Ionospheric effects on the target are corrected through steps of

⁵ <https://github.com/revoltek/LiLF>

direction-independent and direction-dependent calibration (see details in Tasse et al. 2018; de Gasperin et al. 2020b). This procedure provides an image of the whole field of view ($10^\circ \times 10^\circ$) at a resolution of $15''$.

Similarly to the LOFAR HBA data, the extraction and recalibration technique was successfully applied to LBA datasets as well (e.g. Edler et al. 2022; Pasini et al. 2022). We carried out the extraction technique on our data by testing several extraction regions to maximise the signal-to-noise ratio (S/N) of our images after the full self-calibration process; the best results were obtained with a circular extraction region of $23'$ in diameter. We then performed four rounds of amplitude and phase direction-independent self-calibration.

Compared to LOFAR HBA, the beam model of LOFAR LBA is expected to be more accurate (de Gasperin et al. 2023). Therefore, the flux density scale of LOFAR LBA does not require further corrections.

3.3. GMRT and uGMRT radio data

We retrieved archival GMRT observations of A2142 at 307–339 MHz for 5 h on-source, which were first presented by Venturi et al. (2017). The total bandwidth is 32 MHz (split into 256 channels). The source 3C 286 was used as an absolute flux density scale calibrator. We reprocessed the data by means of the Source Peeling and Atmospheric Modeling (SPAM) automated pipeline (Intema et al. 2009), which corrects for ionospheric effects and removes direction-dependent gain errors. Bright sources in the field are used to derive directional-dependent gains and fit a phase-screen over the field of view. Finally, images are corrected for the system temperature variations between the calibrators and the target. These procedures allowed us to reach a noise level of $\sim 40 \mu\text{Jy beam}^{-1}$ at $9''$, which is noticeably lower (by a factor ~ 3) than that reached by Venturi et al. (2017), mainly due to a better calibration of the short baselines.

Additional, unpublished uGMRT observations of A2142 are available in the archive. We retrieved a three-hour observation in the frequency range 300–500 MHz (band-3). The total bandwidth of 200 MHz is split into 4000 channels of 50 kHz each. The sources 3C 286 and 1602+334 were used as an absolute flux density scale and phase calibrators, respectively. We adopted the CASA Pipeline-cum-Toolkit for Upgraded GMRT data REduction (CAPTURE⁶; Kale & Ishwara-Chandra 2021) v.2.0 to perform a standard, fully automated calibration of the wide-band data and derive delay, bandpass, phase, and amplitude corrections. We then performed three cycles of phase-only and two cycles of phase plus amplitude self-calibration, which allowed us to reach a noise level of $\sim 30 \mu\text{Jy beam}^{-1}$.

3.4. Very Large Array radio data

We retrieved three VLA pointings on A2142 at 1–2 GHz (L band) in C and D configurations, for 0.5 and 1.5 h, respectively. The three pointings are slightly offset from the target, allowing us to mosaic the region with a more uniform sensitivity. The source 3C 286 was used as flux density calibrator, whereas J1609+2641 was used as phase calibrator. The data were recorded with 16 spectral windows, each divided into 64 channels.

The data reduction was carried out with the National Radio Astronomy Observatory (NRAO) Common Astronomy

Table 3. Summary of the parameters for LOFAR images discussed in Sect. 4.1.

ν (MHz)	$B_{\min}$ (λ)	Robust	Taper ($''$)	θ ($'' \times ''$)	PA (deg)	σ (mJy beam^{-1})
50	30	−1.0	–	14×10	86	1.9
50	30	−0.5	15	22×13	84	1.6
50	30	−0.5	30	44×32	84	2.7
50	30	−0.5	60	69×63	50	4.4
143	80	−0.5	–	9×6	88	0.075
143	80	−0.5	15	21×20	84	0.16
143	80	−0.5	30	38×35	16	0.25
143	80	−0.5	60	73×66	288	0.40
143	50	−0.5	120	128×117	310	1.0

Notes. Column 1: central frequency (ν). Columns 2–4: minimum baseline ($B_{\min}$), robust parameter of the Briggs weighting, and Gaussian tapering. Columns 5–7: restoring beam (θ), beam position angle (PA), and reached noise (σ).

Software Applications (CASA; McMullin et al. 2007) version 5.1. We performed standard initial phase, bandpass, and gain calibrations. Due to the usual high amount of RFI in this band, $\sim 50\%$ of the bandwidth was flagged. Further rounds of self-calibration were not required. In order to obtain deeper images, we combined the visibilities of the C and D configuration datasets. We finally split the remaining bandwidth into two datasets of ~ 250 MHz width each, centred on 1380 and 1810 MHz.

3.5. Radio imaging and source subtraction

The imaging process was carried out for GMRT, uGMRT, and LOFAR datasets with WSClean version 2.10 (Offringa et al. 2014; Offringa & Smirnov 2017), which can account for wide-field, multi-frequency, and multi-scale synthesis. To properly combine the VLA pointings, we imaged these datasets with tclean in CASA making use of mosaic gridding, as well as multi-frequency and multi-scale synthesis.

We produced images by varying the relative weights of baselines through the robust parameter of the Briggs weighting (Briggs 1995) and Gaussian tapering of the outer uv -coverage to study the radio emission on different spatial scales. The weighting schemes adopted for the images discussed in the next sections are summarised in Table 3.

To accurately measure the flux densities of the diffuse emission, the contribution of the embedded discrete sources needs to be removed. Following Venturi et al. (2017), we first obtained models of the discrete sources by imaging the data at high resolution, excluding baselines $< 2 \text{ k}\lambda$, corresponding to maximum recoverable scales of $\sim 100''$ (i.e. $\sim 170 \text{ kpc}$ at the cluster redshift). This procedure was carried out with WSClean for all the datasets, including the VLA datasets, whose pointings were handled separately. We then subtracted the clean components in the model images from the uv -data. Even though this process allowed us to remove the majority of the sources, residual artefacts associated with the two extended head-tail radio galaxies are still present; moreover, because of calibration imperfections, when decreasing the resolution, very faint residuals from compact sources can also be enhanced and severely contaminate the diffuse emission (e.g. Bruno et al. 2023). Owing to the complexity of the discrete sources and different quality of the subtraction at each frequency, here we assume no systematic subtraction errors, but consider regions of the target with as few discrete

⁶ <https://github.com/ruta-k/CAPTURE-CASA6>

sources as possible when measuring the flux density (see details in Sect. 4.2).

By ignoring the subtraction error contribution, uncertainties ΔS on the reported flux densities are given by

$$\Delta S = \sqrt{(\sigma \cdot \sqrt{N_{\text{beam}}})^2 + (\xi_{\text{cal}} \cdot S)^2}, \quad (1)$$

where σ is the RMS noise of the image, N_{beam} is the number of independent beams within the considered region, and ξ_{cal} is the calibration error. We adopted $\xi_{\text{cal}} = 10\%$ for LOFAR HBA (Shimwell et al. 2022) and LBA (de Gasperin et al. 2021), $\xi_{\text{cal}} = 6\%$ for GMRT and uGMRT band-3 (Chandra et al. 2004), and $\xi_{\text{cal}} = 5\%$ for VLA L-band (Perley & Butler 2013). In the following, position angles (PA) of radio beams and regions are measured north to east.

3.6. Chandra X-ray data

We analysed the archival *Chandra* X-ray data of A2142 that are summarised in Table 2. Observations were carried out in 2005 and 2014 in VFAINT mode with both ACIS-I and ACIS-S CCDs to recover the full extension of the target. We reprocessed the four observations using CIAO v.4.13, with CALDB v.4.9.4. We extracted light curves in source-free regions to filter out soft proton flares with the `lc_clean` algorithm. This procedure left a total clean time of 186.9 ks.

After correcting each pointing for the corresponding point spread function and exposure map, we combined them to produced a single flux image in the 0.5–2 keV band with the task `mergeobs`. Candidate point sources were identified through the `wavdetect` task; after visual inspection, we subtracted the confirmed ones.

Background event files were obtained through the `blanksky` tool, which reprojects blansky pointings in the direction of our observations, and calculates normalisation factors to match the count rates in the 9–12 keV band (with the parameter `weight_method=particle-rate`). These event files were used as input for `blanksky_image` tool to produce the corresponding background images.

X-ray spectral analyses were carried out with XSPEC (Arnaud 1996) v. 12.10.1. It is worth mentioning that present releases of XSPEC are not able to properly deal with spectra extracted from rescaled background events produced by `blanksky`. To this aim, we therefore manually rescaled the exposure time of the reprojected blansky pointings to match the count rates of our observations in the 9–12 keV band.

3.7. XMM-Newton X-ray data

A2142 was first observed by *XMM-Newton* in July 2011 (ObsID 0674560201, PI: M. Rossetti), and then followed up in July 2012 (ObsIDs 0694440101, 0694440501, 0694440601, 0694440201, PI: D. Eckert) as part of the XMM Cluster Outskirts Project (X-COP; Eckert et al. 2017b) to map both the centre and periphery of the cluster. These data consist of a mosaic of five pointings for a total exposure time of 195 ks. We retrieved the final products of these data (count, exposure, background, and background-subtracted flux maps in the 0.7–1.2 keV band) from the public website of the project⁷ (see Rossetti et al. 2013; Tchernin et al. 2016; Ghirardini et al. 2018, for details on observations and data processing).

⁷ <https://dominiqueeckert.wixsite.com/xcop>

4. Results

4.1. Radio images

In this section, we present the new LOFAR HBA and LBA images of A2142 shown in Figs. 2 and 3, respectively. A2142 hosts a number of compact and extended radio galaxies that were discussed in Venturi et al. (2017). Here we summarise the most interesting sources, which are labelled in the composite radio (GMRT), X-ray (*XMM-Newton*), and optical (DSS-2) image reported in Fig. 1 as in Venturi et al. (2017). The diffuse radio components of the halo are labelled in Fig. 4, where LOFAR HBA images at high and low resolutions are overlaid.

The most spectacular extended radio galaxies are ‘T1’ and ‘T2’, two head-tail galaxies. The head of T1 (i.e. the core of the radio galaxy) is coincident with the NW cold front, and the tail has a projected length of $\sim 5.5'$, corresponding to ~ 550 kpc, if measured from the HBA at the highest resolution ($9'' \times 6''$). The morphology of the tail is not straight and its width is not constant. Additional radio emission is detected by LOFAR HBA, which shows that the tail ends with a thin filament directly connected to a perpendicular arc-shape structure (see also Fig. 4), extending for $\sim 2'$ (~ 200 kpc) in the NS direction. The head-tail galaxy T2 is located at a projected distance of ~ 650 kpc from the cluster centre, and extends northwestwards for $\sim 4'$ (~ 400 kpc) in the highest resolution image. The complex morphology and features of T1 and T2 suggest an interplay with ICM motions, which will be the subject of a dedicated analysis in the future. The locations of the two BCGs are reported in Fig. 1. As mentioned earlier, the primary BCG is coincident with the cluster centre, whereas the secondary one is likely the main member of a merging group. Only BCG1 is radio active, hosting a compact radio galaxy. We refer to Venturi et al. (2017) for information on ‘W1’ and ‘W2’ (two wide-angle tails that are not cluster members), and galaxy groups ‘G’ (in the northeast; see also the X-ray counterpart in Eckert et al. 2017a) and ‘C’ (in the south).

The diffuse components H1 (core) and H2 (ridge) are visible at 143 MHz at high resolution (Fig. 2, top left panel), whereas they are undetected at the 3σ level in the $14'' \times 10''$ LBA image (Fig. 3, top left panel). With an intermediate resolution ($22'' \times 13''$), the core and the ridge are detected at 50 MHz as well (Fig. 3, top right panel). By decreasing the resolution and increasing the sensitivity to the diffuse emission, further emission is revealed, as suggested by Farnsworth et al. (2013) with GBT observations. Here we unambiguously confirm the existence of a third radio halo component, ‘H3’, which has an elliptical morphology elongated in the NW–SE direction, and embeds the core, the ridge, and the two head-tail galaxies T1 and T2, as shown in the $73'' \times 66''$ and $69'' \times 63''$ images at 143 and 50 MHz, respectively (see also Fig. 4).

In Fig. 5, the radio contours of the $21'' \times 20''$ and $73'' \times 66''$ resolution images at 143 MHz are overlaid on the *XMM-Newton* image of the cluster, with arrows indicating the position of the cold fronts. The core is confined by the southern central cold front, whereas the northern boundary is not clear owing to the presence of T1. The ridge extends along the NW–SE axis towards the outermost southern cold front. The region in between the core and the ridge, southeast of the cold front, shows a depletion of radio emission, which forms a ‘bay’-like structure that is labelled in Fig. 4. The bulk of the volume of the ICM is occupied by H3 and its radio emission follows the spatial distribution of the thermal X-ray emission, as typical of giant radio halos, reaching the outermost cold front from the south. We measured the maximum extension of H3 by considering the 2σ level of a LOFAR HBA image at resolution $128'' \times 117''$; we

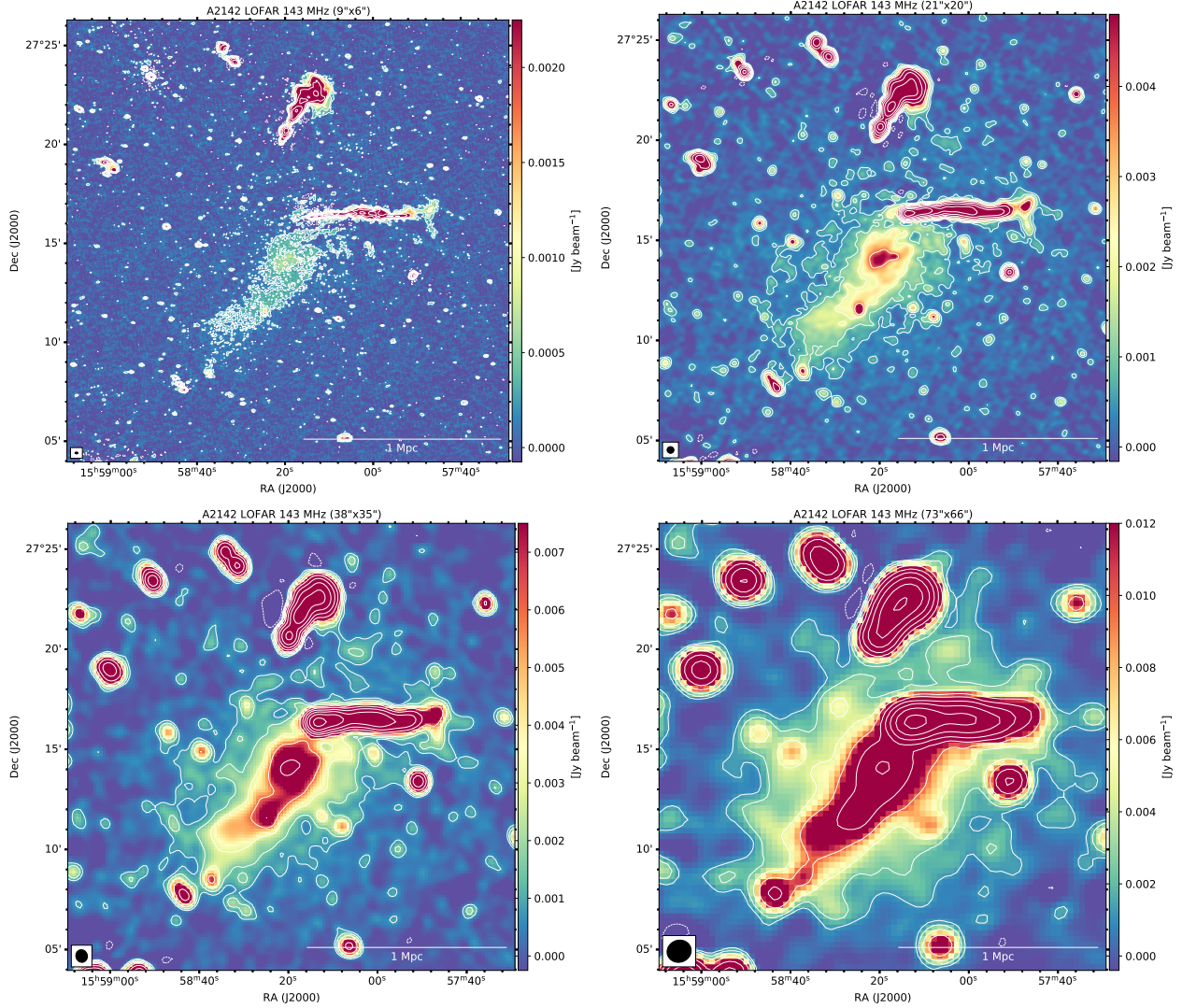


Fig. 2. A2142 LOFAR HBA radio images at 143 MHz at different resolutions. *Top left:* $9'' \times 6''$ resolution ($\sigma = 0.075 \text{ mJy beam}^{-1}$). *Top right:* $21'' \times 20''$ resolution ($\sigma = 0.16 \text{ mJy beam}^{-1}$). *Bottom left:* $38'' \times 35''$ resolution ($\sigma = 0.25 \text{ mJy beam}^{-1}$). *Bottom right:* $73'' \times 66''$ resolution ($\sigma = 0.40 \text{ mJy beam}^{-1}$). In all the panels, the contour levels are $[\pm 3, 6, 12, \dots] \times \sigma$.

can approximate its morphology as an ellipse of projected axis lengths of $2.4 \text{ Mpc} \times 2.0 \text{ Mpc}$.

The detection of H3 means that A2142 is the only hybrid radio halo with three distinct components discovered so far. In the following sections, we analyse the properties of the radio halo in detail with multi-frequency data to investigate the origins and possible connections between its components.

4.2. Spectral indices

The spectral index is key to probing the mechanisms responsible for the origin of radio halos. Owing to the extremely complex structure of A2142, a straightforward measurement of α is not trivial, and requires different methods to verify the self-consistency of our results. First, to accurately determine the spectral index of the various components, we produced sets of images at different resolutions with common uv -ranges that allow us to consistently measure the flux densities and minimise systematic uncertainties introduced by inhomogeneous uv -coverage of the different instruments (see details on the imaging parameters in Table 4).

The first set of images at $25''$ spans a frequency range from 50 to 1810 MHz. These provide sufficient resolution and sensitivity to measure the flux density of H1. As the northern limit of H1 cannot be easily determined, we considered the core as a sphere centred on BCG1 and confined to the south by the inner cold fronts, and therefore we measured the flux density in a circle (red circle in Fig. 6) of radius $R_{H1} = 1.1' \sim 110 \text{ kpc}$.

The uv -coverage of VLA is less dense at short spacings than those of LOFAR and GMRT. To avoid flux-density losses due to missing short baselines that would bias the spectral index towards steeper values, we excluded the VLA data from the measures in H2. As a compromise between resolution and sensitivity, we again considered images convolved to $25''$. Similarly to Venturi et al. (2017), we considered a box (blue box in Fig. 6) of $3.8'' \times 2.7''$ in size (corresponding to $380 \text{ kpc} \times 270 \text{ kpc}$) and $PA = 33^\circ$ that roughly encompasses the 3σ level of the ridge.

A third set of images from 50 to 407 MHz was obtained with a resolution of $85''$. At such low resolution, residuals from the discrete source subtraction are enhanced and have a non-negligible contribution when compared with the faint diffuse emission. As explained in Sect. 3.5, obtaining quantitative

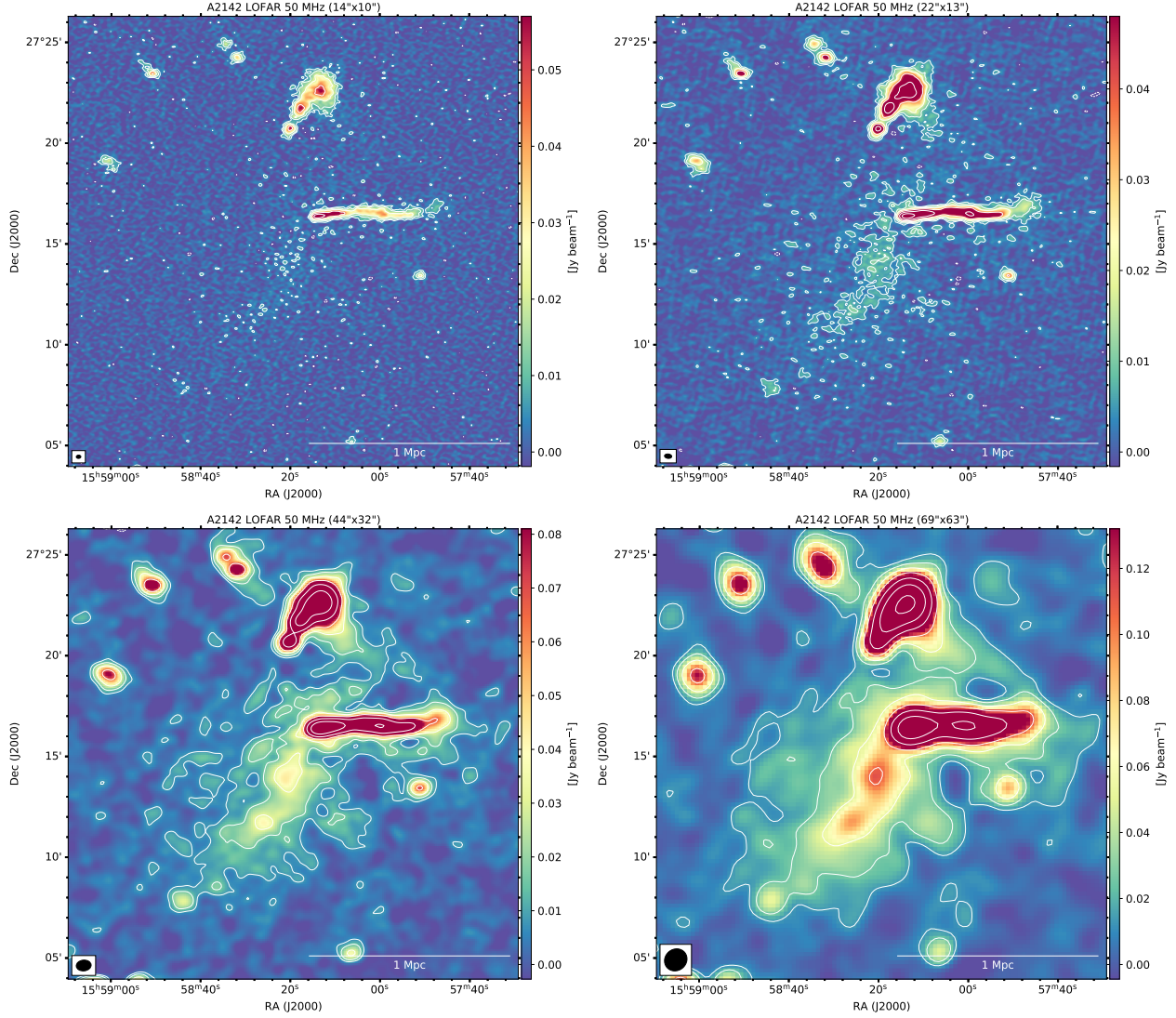


Fig. 3. A2142 LOFAR LBA radio images at 50 MHz at different resolutions. *Top left:* $14'' \times 10''$ resolution ($\sigma = 1.9 \text{ mJy beam}^{-1}$). *Top right:* $22'' \times 13''$ resolution ($\sigma = 1.6 \text{ mJy beam}^{-1}$). *Bottom left:* $44'' \times 32''$ resolution ($\sigma = 2.7 \text{ mJy beam}^{-1}$). *Bottom right:* $69'' \times 63''$ resolution ($\sigma = 4.4 \text{ mJy beam}^{-1}$). In all the panels, the contour levels are $[\pm 3, 6, 12, \dots] \times \sigma$.

estimates of the quality of the subtraction of each dataset is not trivial. Therefore, to minimise contamination by the residuals of discrete sources, we measured the local flux densities of H3 in two boxes (green boxes in Fig. 6) of sides $6' \times 3'$ ($600 \text{ kpc} \times 300 \text{ kpc}$) and $3.5' \times 3'$ ($350 \text{ kpc} \times 300 \text{ kpc}$), where the halo is detected at 3σ at each frequency.

Due to insufficient sensitivity at 323 and 407 MHz, we are only able to recover the outer emission of H3 with LOFAR. We thus obtained images at 50 and 143 MHz, and convolved them to $134''$ to further enhance the S/N. We considered two boxes (purple boxes in Fig. 6) of $7.7' \times 2.8'$ ($770 \text{ kpc} \times 280 \text{ kpc}$) and $7.4' \times 2.5'$ ($740 \text{ kpc} \times 250 \text{ kpc}$) in size that are free from source contamination, where the halo is detected at the 2σ level of both images.

All the regions defined as above are overlaid on the contours of the $134'' \times 134''$ HBA image of Fig. 6. We used the measured flux densities reported in Table 5 to fit the radio spectra with power laws that well describe the data points, as shown in Fig. 6. The fitted spectral indices are $\alpha_{\text{H1}}^{[50-1810]} = 1.09 \pm 0.02$ for H1, $\alpha_{\text{H2}}^{[50-323]} = 1.15 \pm 0.02$ for H2, $\alpha_{\text{H3}}^{[50-323]} = 1.36 \pm 0.05$ in the

inner parts (green boxes) of H3, and $\alpha_{\text{H3}}^{[50-143]} = 1.57 \pm 0.20$ in the outer parts (purple boxes) of H3 (see also the derivation of α for H3 from the total flux density in Sect. 4.3).

Figure 7 presents the spectral index maps (the corresponding error maps are shown in Fig. A.1) at $11'' \times 11''$ between 143 and 323 MHz (with and without discrete sources). The spectral index maps between 50 and 143 MHz at resolutions of $24''$, $75''$, and $136''$ are shown in Fig. 8 (the corresponding error maps are shown in Fig. A.2). The spectral indices inferred from these maps are globally consistent with the fitted spectra. The high-resolution 143–323 MHz spectral index maps (Fig. 7) show that the spectral index distribution is non-uniform, especially along the ridge; interestingly, a significant flattening ($\alpha \sim 0.9$) is observed in the region coincident with the radio bay. Outside H2 (Fig. 8), the spectrum further steepens up to ~ 1.6 . In the northern part, unreliable spectral indices $\alpha \gtrsim 2$ are the result of subtraction artefacts from T1 and T2. On the other hand, it is plausible that the tails of these galaxies release very old populations of electrons that are then reaccelerated by turbulence, and contribute to the steepest regions of the halo to the north, as suggested by the non-source-subtracted spectral index maps.

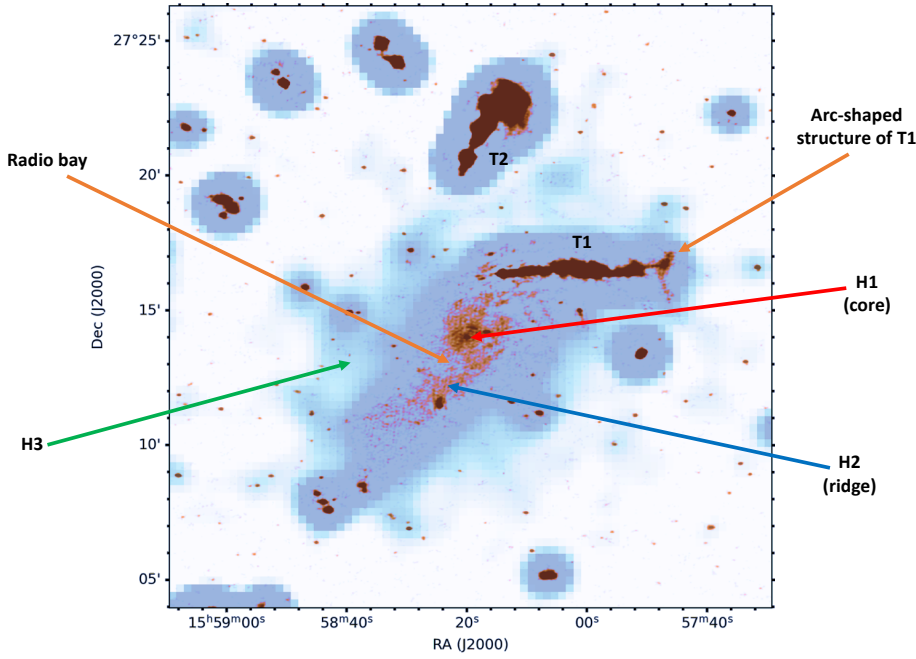


Fig. 4. Composite LOFAR HBA image of A2142 at high ($9'' \times 6''$, reddish colours) and low ($73'' \times 66''$, blueish colours) resolution. The radio components of the halo (H1, H2, H3), radio galaxies (T1, T2), and additional features discussed in the text are labelled.

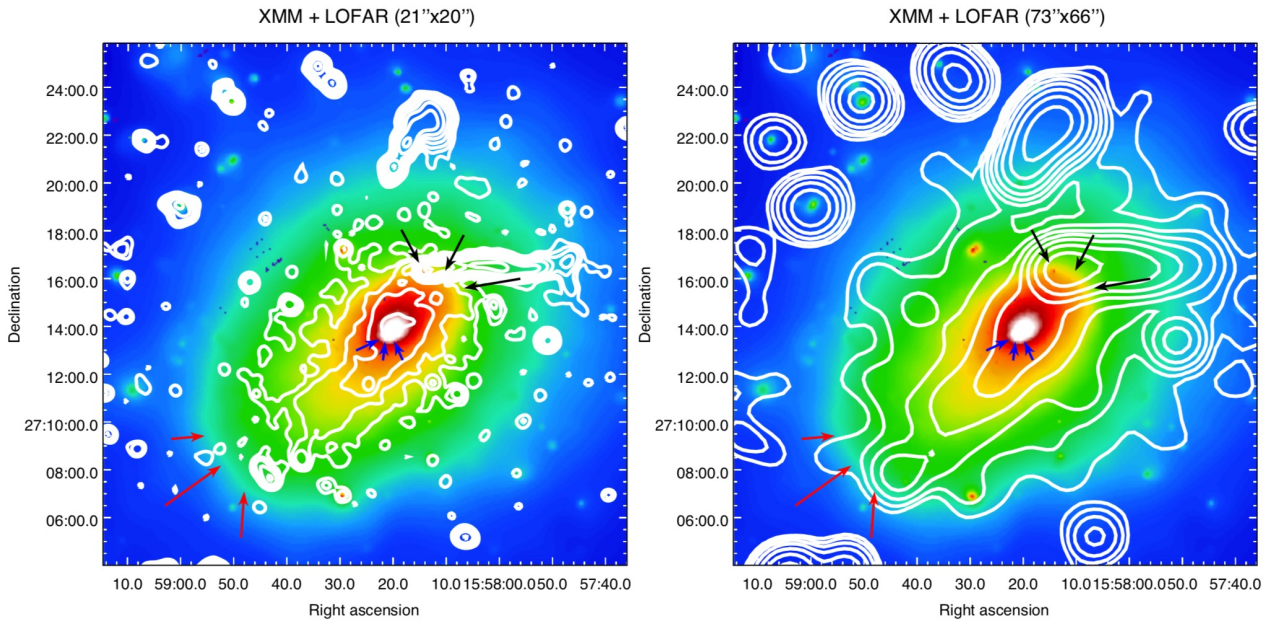


Fig. 5. A2142 LOFAR HBA radio contours from images in Fig. 2 (upper right and lower right panels) overlaid on *XMM-Newton* X-ray image. The locations of the cold fronts are indicated by arrows.

Table 4. Sets of images used for flux density measurements.

ν (MHz)	uv -range	Taper (arcsec)	θ_{conv} (arcsec)
50, 143, 323, 407, 1380, 1810	[250 λ –15 k λ]	–	25
50, 143, 323, 407	[60 λ –18 k λ]	15	25
50, 143, 323, 407	[60 λ –18 k λ]	60	85
50, 143	[50 λ –18 k λ]	120	134

Notes. Column 1: frequency of considered datasets. Columns 2 and 3: adopted uv -range and Gaussian taper. Column 4: circular beam adopted to convolve the images.

We point out that our integrated spectra are flatter than $\alpha \sim 1.3$ and $\alpha \sim 1.5$, as reported by Venturi et al. (2017) for H1 and H2, respectively. Discrepancies are likely caused by a combination of several factors. Indeed, we considered a smaller region for H1 to avoid possible contamination from H2 and/or residuals of subtraction from T1; improved the calibration of the GMRT data at 323 MHz; have a more accurate LOFAR HBA flux density value (e.g. uncertainties in Venturi et al. 2017 are $\sim 35\%$ for the LOFAR measurement); and considered only the datasets with the densest uv -coverage. Through our procedures, we obtained less scattered data points to be fitted than those reported by Venturi et al. (2017).

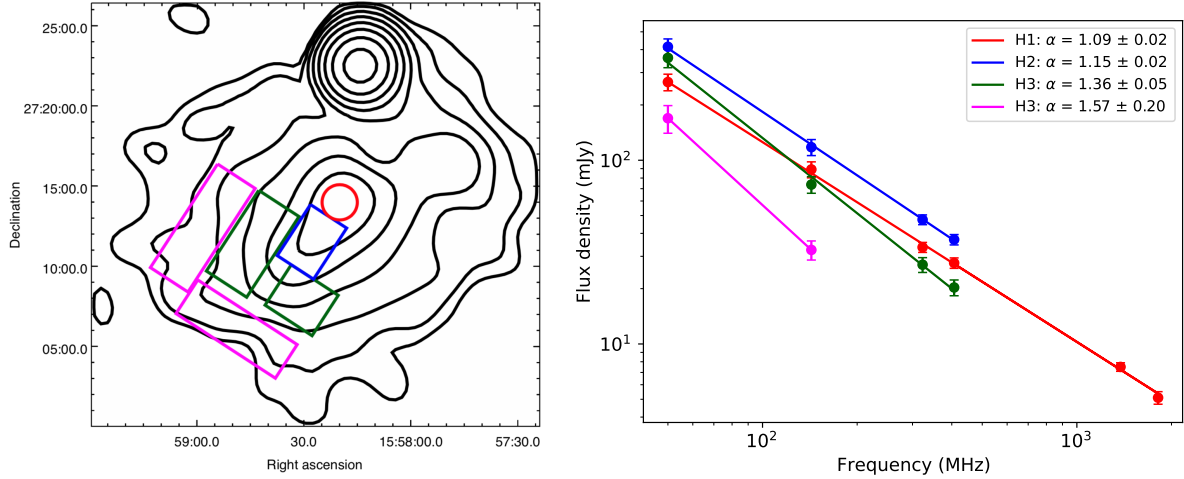


Fig. 6. Radio spectrum of A2142. *Left panel:* regions used to measure the flux densities of H1, H2, and H3 from various images overlaid on the 143 MHz contours ($[2, 4, 8, \dots] \times \sigma$) of the $134'' \times 134''$ image (see details in the text). *Right panel:* radio spectra of H1, H2, H3 fitted with power laws.

Table 5. Flux densities of H1, H2, and H3 as derived from the regions shown in Fig. 6.

Source	Region	S_{50} (mJy)	S_{143} (mJy)	S_{323} (mJy)	S_{407} (mJy)	S_{1380} (mJy)	S_{1810} (mJy)	α
H1	Red	266.4 ± 27.6	89.1 ± 8.9	33.6 ± 2.1	27.6 ± 1.8	7.5 ± 0.4	5.1 ± 0.4	1.09 ± 0.02
H2	Blue	414.4 ± 43.0	117.8 ± 11.8	47.5 ± 2.9	37.0 ± 2.4	—	—	1.15 ± 0.02
H3	Green	360.8 ± 41.9	73.7 ± 7.6	27.0 ± 2.5	20.3 ± 2.0	—	—	1.36 ± 0.05
H3	Purple	169.3 ± 29.0	32.5 ± 3.9	—	—	—	—	1.57 ± 0.20

Notes. Column 8 reports the fitted spectral index. The regions defined for H1 and H2 roughly cover their total extensions and flux densities are therefore total values. Flux densities of H3 are instead local values, obtained as the sum of the measurements within the considered boxes. Uncertainties on α for H1, H2, and inner H3 are fitting errors; the uncertainty on α for outer H3 is derived from the standard formula for the error propagation as in Eq. (A.1).

4.3. Radio surface brightness profile

In Sect. 4.2, we obtained local measures of the flux density and spectral index for H3. We now aim to determine global values by means of the average surface brightness and its radial profile.

As a first check, we masked the regions of H1, H2, and residuals from subtraction, and then obtained an estimate of the average surface brightness of H3 over an ellipse of $2.4 \text{ Mpc} \times 2.0 \text{ Mpc}$. By assuming that H3 uniformly fills the masked regions with the same average brightness, we estimated total flux densities of $S_{50} = 3116.6 \pm 316.0 \text{ mJy}$ and $S_{143} = 628.3 \pm 63.0 \text{ mJy}$.

More accurate values can be derived by fitting the surface brightness profile of the halo. Under simple assumptions of spherical symmetry, the surface brightness profile of both mini-halos and radio halos can be typically reproduced by an exponential law (Murgia et al. 2009):

$$I(r) = I_0 e^{-\frac{r}{r_e}}, \quad (2)$$

where I_0 is the central brightness and r_e is the e -folding radius. However, Eq. (2) is not able to properly fit the surface brightness of A2142 owing to its complex multi-component structure. We can assume that both H1 and H3 have spherical symmetry⁸, and that H3 fills a fraction of the volume of H1 and H2. We therefore

⁸ Even though an ellipsoidal geometry would be more reliable, we do not expect that our assumption introduces significant biases in the fitted parameters, since the minor to major axis length ratio (1.2) is relatively small.

considered a double exponential function:

$$I(r) = I_{0,\text{in}} e^{-\frac{r}{r_{e,\text{in}}}} + I_{0,\text{out}} e^{-\frac{r}{r_{e,\text{out}}}}, \quad (3)$$

where the contributions of the inner (H1) and outer (H3) components are summed. As we did not attempt to model the emission from the ridge because of its complex morphology, its contribution is therefore included in the fitted flux density of the outer component, below.

As shown in the examples in Fig. 9 for LOFAR HBA images at $40'' \times 40''$ and $134'' \times 134''$ resolution, the double exponential model can well reproduce the profile of A2142. In Table 6, we report the central brightness and e -folding radius for the outer component H2+H3, which we are interested in. The flux density is computed by integrating the surface brightness in annuli up to a radial distance of $\hat{r} = 3r_e$, which typically sets the finite extension of radio halos; this radial cut-off provides the 80% of the flux density that would be obtained by integrating the surface brightness up to $\hat{r} = \infty$, thus yielding $S_{\text{fit}} = 0.8 \times 2\pi I_0 r_e^2$. We find discrepancies in the fitted I_0 and r_e , depending on the input image; this is likely due to the (unmodelled) emission of the ridge, which contaminates the fits with different weights. On the other hand, the corresponding flux densities are well in agreement within the errors, and with the rough estimate provided above by considering the elliptical geometry and assuming a uniform distribution of the average brightness, which confirms the reliability of our approach.

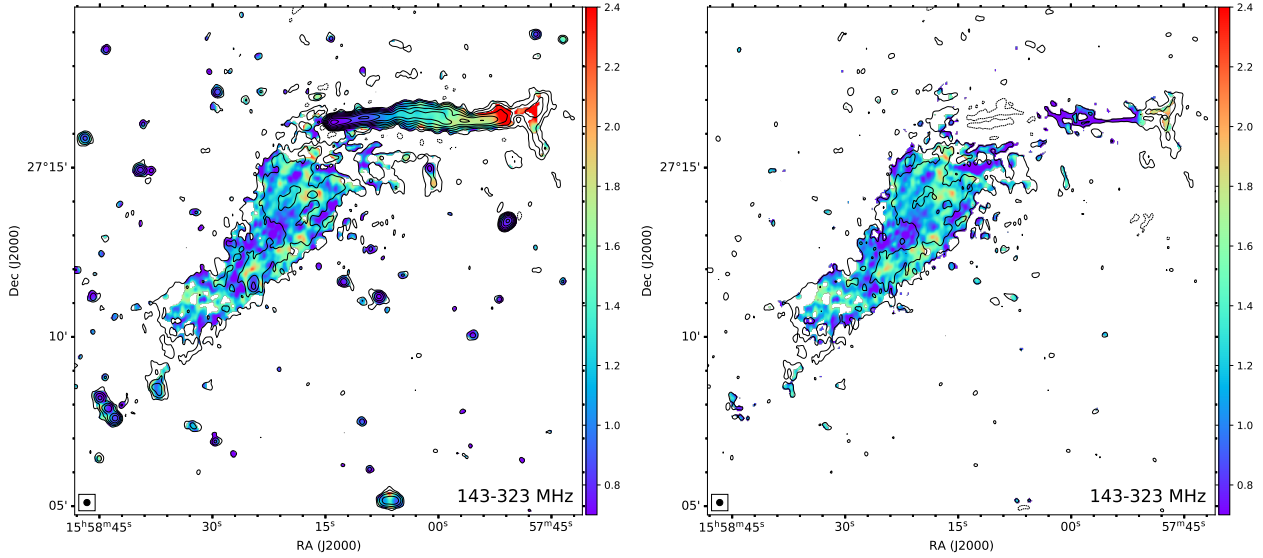


Fig. 7. A2142 spectral index maps between 143 and 323 MHz at $11'' \times 11''$ resolution, before (*left*) and after (*right*) the subtraction of the discrete sources. LOFAR contours are drawn in black.

By considering $S_{50} = 3271.2 \pm 528.3$ mJy and $S_{143} = 628.4 \pm 85.0$ mJy, we obtain a global $\alpha_{\text{H3}}^{[50-143]} = 1.57 \pm 0.20$, which is perfectly consistent with local values inferred in Sect. 4.2. This confirms that H3 is an ultrasteepest-spectrum radio halo.

4.4. Radio power

By considering the flux densities and spectral indices of H1, H2, and H3, we computed the corresponding radio powers as

$$P_{\nu} = 4\pi D_L^2 S_{\nu} (1+z)^{\alpha-1}. \quad (4)$$

If a flux density measurement was not provided for H2 or H3 at a given frequency, we extrapolated it from the fitted spectrum by assuming S_{143} as reference. The radio powers are reported in Table 7.

As mentioned, the flux density derived from the surface brightness profile of H3 in Sect. 4.3 also includes the contribution from H2. We scaled the corresponding radio power P_{143} to 150 MHz ($P_{150} = (1.2 \pm 0.2) \times 10^{25}$ W Hz $^{-1}$). As typical of ultrasteepest-spectrum radio halos (e.g. Cassano et al. 2013; Donnert et al. 2013; Cuciti et al. 2021), A2142 is located well below (by a factor of ~ 3) the scatter of the correlation between radio power and host mass reported in Cuciti et al. (2023).

4.5. Thermal properties and non-thermal emission

In this section, we compare the thermal and non-thermal properties derived from our X-ray and radio data, respectively. Radio halos are generally observed to have smooth surface brightness profiles, as discussed in Sect. 4.3. Nonetheless, Botteon et al. (2023) recently showed that substructures of the diffuse radio emission, such as surface brightness discontinuities, are detectable when observing radio halos with sufficiently high resolution and sensitivity, and that these can be co-spatial with substructures and discontinuities in the X-ray emission of the thermal ICM. Following a similar approach, we applied the Gaussian gradient magnitude (GGM; Sanders et al. 2016a,b) filtering to enhance surface-brightness discontinuities and substructures in A2142. We adopted $\sigma_{\text{GGM}} = 2$ pix and $\sigma_{\text{GGM}} = 8$ pix for our *Chandra* and LOFAR HBA images (where

1 pix = 1.5 arcsec in both cases) as the width of the derivative of the Gaussian filter function. The resulting GGM-filtered images are shown in Fig. 10. In the left panel, we observe some gas that departs from the BCG⁹ (black cross) and produces the inner cold fronts following a clear spiral path. In the radio counterpart (right panel), the brightest spots spatially correspond to the boundaries of the cold fronts. The western filamentary structure enhanced in the GGM-filtered radio image might suggest that the ridge is directly connected with the core, suggesting a common origin of the two radio components; on the other hand, even by filtering the *Chandra* image with wider σ_{GGM} , we do not find a corresponding X-ray structure that could confirm this hypothesis.

By using *XMM-Newton* X-ray images in different energy bands (see details of this procedure in Rossetti et al. 2007), Rossetti et al. (2013) produced projected entropy maps of A2142, and found that the ridge follows a trail of low-entropy ICM phase. We investigated this trend with *Chandra* data as well. Through CONTBIN¹⁰ (Sanders 2006), we adaptively binned our 0.5–2 keV exposure-corrected image in regions with a high signal-to-noise ratio of $S/N = 100$. We extracted the spectra in each region of the event and blanksky files, subtracted the corresponding background contribution from the ICM emission, and simultaneously fitted the resulting spectra in XSPEC with an absorbed thermal plasma component (phabs $\times$ apec) by keeping the hydrogen column density and metal abundance fixed at values of $N_{\text{H}} = 3.8 \times 10^{20}$ cm $^{-2}$ and $Z = 0.28 Z_{\odot}$ (Markevitch et al. 2000). The Cash statistics (Cstat; Cash 1979) was adopted for the fits; the chosen high S/N ensured Cstat/d.o.f. in the range 0.90–1.14 for each region, indicating the goodness of our fits. The procedure and systematic errors related to the derivation of the entropy map using the fitted temperature and apec normalisation are described in Appendix B. Figure 11 shows both our projected entropy map (left panel) and that produced by Rossetti et al. (2013; multiplied by the pixel area to match our units). Despite the different procedures, resolution, and S/N , the two maps are in close agreement in the common area that covers a radius of ~ 350 kpc. The LOFAR HBA

⁹ The X-ray peak is found at a distance of $21'' = 35$ kpc from BCG1, in line with Wang & Markevitch (2018).

¹⁰ <https://github.com/jeremysanders/contbin>

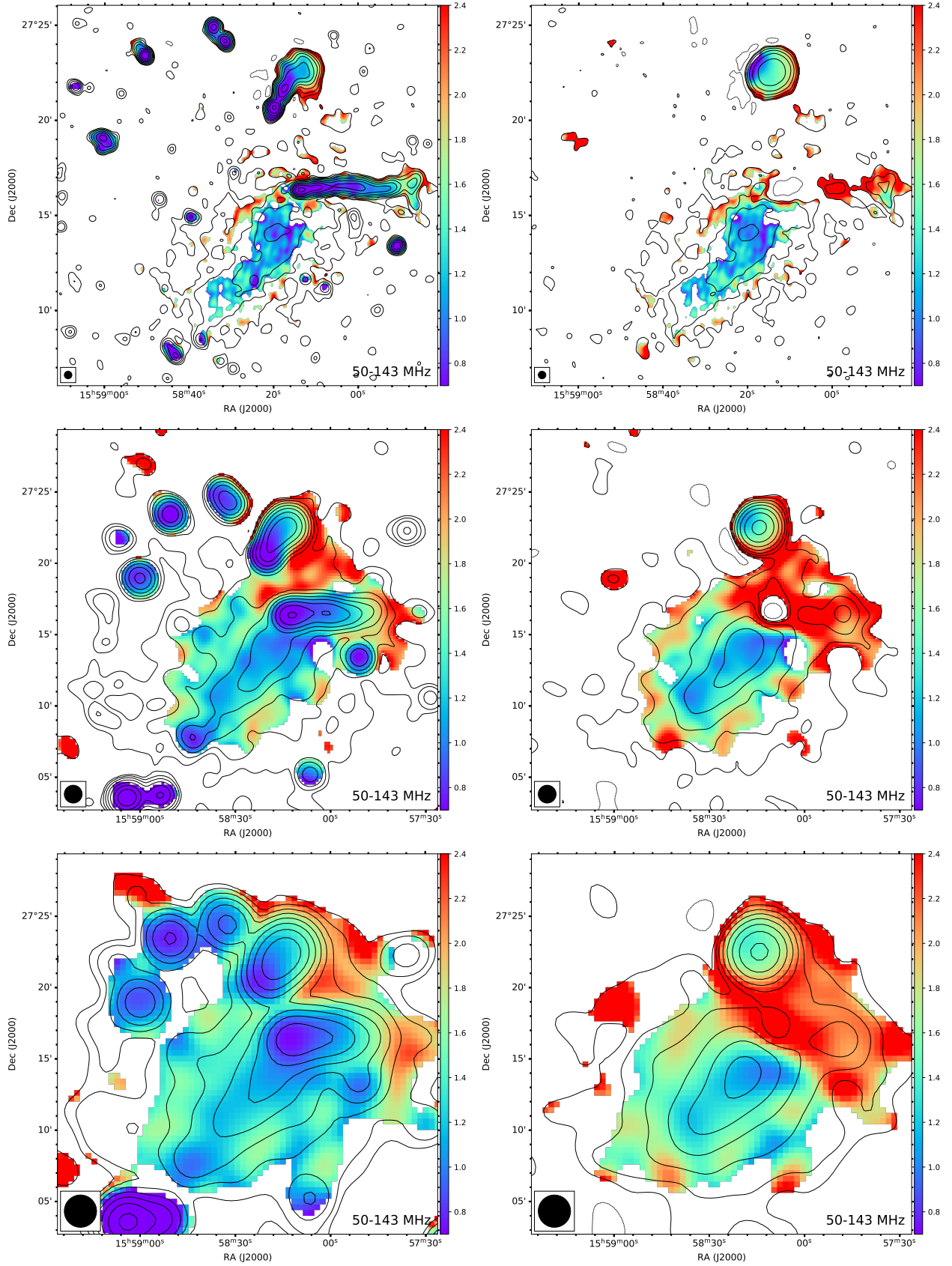


Fig. 8. A2142 spectral index maps between 50 and 143 MHz at various resolutions before (*left panels*) and after (*right panels*) the subtraction of the discrete sources. *Top:* maps at $24'' \times 24''$ resolution. *Middle:* maps at $75'' \times 75''$ resolution. *Bottom:* maps at $136'' \times 136''$ resolution. In all the panels, the corresponding contours at 143 MHz are drawn in black.

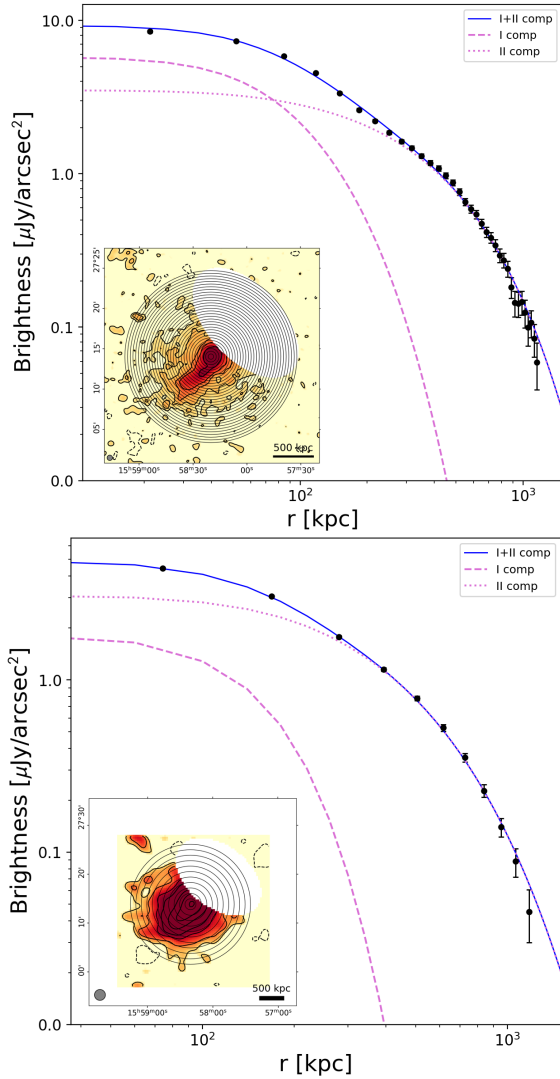


Fig. 9. Fit of the surface brightness of LOFAR HBA images at $40'' \times 40''$ (top, see inset) and $134'' \times 134''$ (bottom, see inset) with a double spherically symmetric exponential model (Eq. (3)). In both cases, residuals from T1 and T2 were masked and the width of the sampling annuli is half of the beam FWHM.

contours at $21'' \times 20''$ resolution (Fig. 2, upper right panel) are overlaid on the maps. In the radio core, the ICM has entropy values $s \lesssim 60 \text{ keV cm}^{5/3} \text{ arcmin}^{-2/3}$. As claimed by Rossetti et al. (2013), we confirm that the ridge is co-spatial with a trail of gas that extends in the NW–SE direction and is characterised by lower entropy values ($s \sim 100 \text{ keV cm}^{5/3} \text{ arcmin}^{-2/3}$) than those found along different directions at the same radial distance.

4.6. Point-to-point radio versus X-ray analysis

The radio surface brightness of both halos and mini-halos and the X-ray surface brightness of the thermal ICM are found to be spatially correlated as $I_R \propto I_X^k$ (e.g. Govoni et al. 2001). The slope k of this correlation indicates the relative spatial distributions of thermal and non-thermal components, and can therefore be used to probe the origin of the diffuse radio emission. Super-linear slopes $k > 1$ are associated with a narrower distribution of the emitting electrons than that of the thermal particles; conversely, sublinear $k < 1$ slopes are associated with narrower thermal and broader non-thermal distributions, respectively, sup-

porting a scenario where particles are accelerated and transported by turbulence generated on large scales. Typically, radio halos have been found to exhibit sublinear slopes (e.g. Hoang et al. 2019; Bruno et al. 2021; Duchesne et al. 2021; Rajpurohit et al. 2021b; Pasini et al. 2022), while superlinear slopes have been found in mini-halos (e.g. Govoni et al. 2009; Ignesti et al. 2020; Biava et al. 2021; Riseley et al. 2022).

We investigated the radio versus X-ray correlation in A2142 through the Point-to-point TReND EXtractor code¹¹ (PT-REX; Ignesti 2022), by gridding the regions of interest with beam-size square boxes. The extracted surface-brightness values were fitted in a logarithmic plane as $\log_{10} I_R = k \log_{10} I_X + c$ (where c is the intercept) with `linmix` (Kelly 2007), which can account for the errors of the measurements and intrinsic scatter of the linear regression through a Bayesian statistical approach. We considered LOFAR HBA and *XMM-Newton* images at a resolution of $128'' \times 117''$ as inputs, which allow us to investigate the faintest regions of the radio halo and the outskirts of the ICM, and considered a radio surface-brightness threshold of 2σ . As shown in Fig. 12, the data points are well correlated and fitted by a sublinear power law of slope $k = 0.82 \pm 0.03$ (Spearman and Pearson ranks are $\rho_S = 0.92$ and $\rho_P = 0.93$, respectively), as typical of usual radio halos.

In spite of the solid sublinear correlation we find, the slope might not be constant over the whole radio halo due to the presence of multiple radio components (see also the case of Coma in Bonafede et al. 2022), whose origin may be associated with different physical mechanisms. To investigate the radial behaviour of the slope $k = k(r)$, we considered concentric elliptical sectors that approximately follow the distribution of the ICM to measure I_R and I_X . The slope as a function of the radial distance is therefore computed as

$$k(r) = \frac{\Delta(\ln I_R)}{\Delta(\ln I_X)}, \quad (5)$$

where $\Delta(\ln I)$ is the difference of the logarithms of the radio or X-ray brightness evaluated in two consecutive sectors. Owing to its complex geometry, we ignored the region of H2. We considered radio images at various resolutions; the width of the annuli was set depending on the resolution and required S/N. An example of the sampling sectors is reported in the left panel of Fig. 13.

In the right panel of Fig. 13 we report the derived radial profiles $k(r)$ for different resolutions and S/N. We find a possible slightly superlinear correlation within the innermost regions (~ 50 – 100 kpc) of H1; on the other hand, the rapid drop in k with distance more likely suggests that this trend is caused by the offset between the radio and X-ray peaks, and not associated with a specific physical process. For $r \gtrsim 100$ kpc, $k(r)$ is on average sublinear, indicating that the distribution of the non-thermal components is broader than that of thermal components, and the slope varies in the range ~ 0.7 – 0.8 . Moreover, we observe two narrow peaks of $k(r)$, for $r \sim 150$ kpc and $r \sim 650$ kpc. These peaks might be related to the transitions H1–H2 and H2–H3, or trace edges in the radio and/or X-ray surface brightness. Even though the width of the sampling annuli is an important factor to ensure sufficient S/N on the different scales, the discussed trends are barely dependent on it, which means our conclusions remain valid. In Sect. 5.1, we further discuss the slope and its radial behaviour in order to constrain the possible formation scenario of each radio component.

¹¹ <https://github.com/AIgnesti/PT-REX>

Table 6. Results of the outer component in the double exponential fit of Eq. (3) (see Fig. 9).

ν (MHz)	θ	$I_{0,\text{out}}$ ($\mu\text{Jy arcsec}^{-2}$)	$r_{\text{e,out}}$ (kpc)	$S_{\text{out,fit}}$ (mJy)	ξ_{fit} (%)
50	$46'' \times 46''$	19.4 ± 1.1	299.0 ± 11.8	3129.7 ± 628.6	10
50	$132'' \times 132''$	19.5 ± 1.1	304.9 ± 15.5	3271.2 ± 528.3	12
143	$40'' \times 40''$	3.9 ± 0.2	303.8 ± 7.3	649.5 ± 110.6	7
143	$134'' \times 134''$	4.7 ± 0.3	272.2 ± 8.7	628.4 ± 85.0	9

Notes. Columns 1 and 2: frequency and resolution of the input image. Columns 3 and 4: fitted central surface brightness and e -folding radius. Columns 5 and 6: corresponding flux density ($S_{\text{fit}} = 0.8 \times 2\pi I_{0,\text{e}} r_{\text{e}}^2$) integrated up to $3r_{\text{e,out}}$ and error contribution (in percentage) from the fitting uncertainties (see details in the table notes). Errors on the fitted flux density are computed as $\Delta S_{\text{fit}} = \sqrt{(\sigma \cdot \sqrt{N_{\text{beam}}})^2 + (\xi_{\text{cal}} \cdot S_{\text{fit}})^2 + (\xi_{\text{fit}} \cdot S_{\text{fit}})^2}$ to include the uncertainties on the fitted parameters; ξ_{fit} is obtained from the standard formula for the error propagation as $\xi_{\text{fit}} = \sqrt{\left(\frac{\Delta I_0}{I_0}\right)^2 + 4\left(\frac{\Delta r_{\text{e}}}{r_{\text{e}}}\right)^2}$.

Table 7. Radio powers of H1, H2, and H3 at 50, 143, 323, 407, 1380, and 1810 MHz from measurements reported in Sects. 4.2 (for H1 and H2) and 4.3 (for H2+H3).

Source	P_{50} ($10^{24} \text{ W Hz}^{-1}$)	P_{143} ($10^{24} \text{ W Hz}^{-1}$)	P_{323} ($10^{23} \text{ W Hz}^{-1}$)	P_{407} ($10^{23} \text{ W Hz}^{-1}$)	P_{1380} ($10^{23} \text{ W Hz}^{-1}$)	P_{1810} ($10^{22} \text{ W Hz}^{-1}$)
H1	5.1 ± 0.5	1.7 ± 0.2	6.4 ± 0.4	5.3 ± 0.3	1.4 ± 0.1	9.7 ± 0.8
H2	7.9 ± 0.8	2.3 ± 0.2	9.1 ± 0.6	7.1 ± 0.5	$1.7 \pm 0.2^*$	$12.1 \pm 1.4^*$
H2+H3	64.8 ± 10.5	12.5 ± 1.7	$34.7 \pm 7.4^{**}$	$24.1 \pm 6.0^{**}$	$3.5 \pm 1.7^{**}$	$23.2 \pm 12.2^{**}$

Notes. Values marked with a single asterisk or a double asterisk are obtained by extrapolation with $\alpha = 1.15$ and $\alpha = 1.57$, respectively.

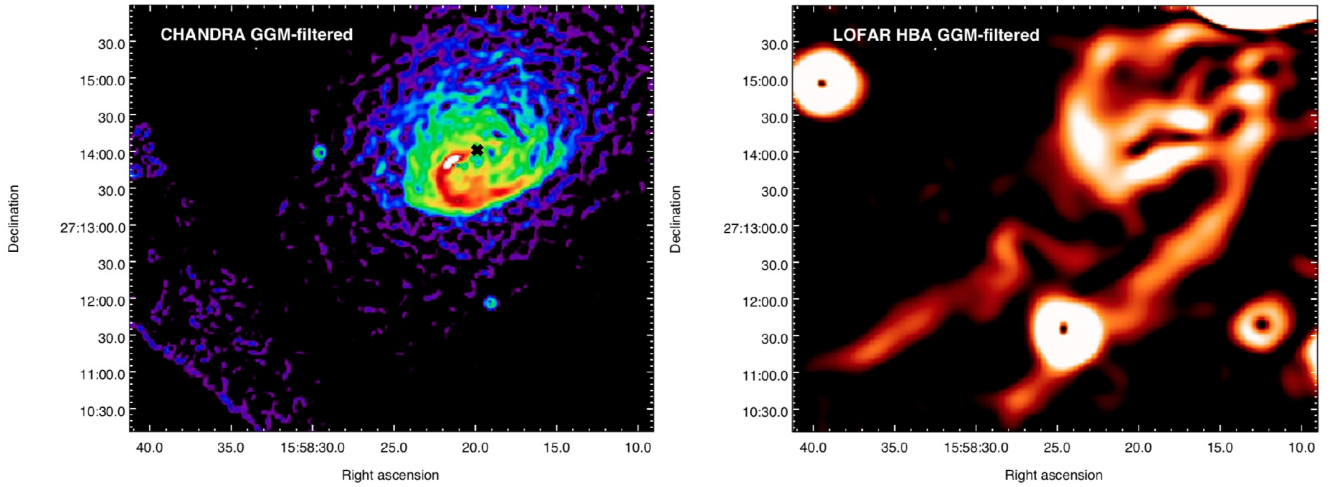


Fig. 10. GGM-filtered images of A2142. Discrete sources were not subtracted and appear as ring structures due to GGM-filtering. *Left panel:* Chandra X-ray flux image filtered with $\sigma_{\text{GGM}} = 3''$. The black cross marks the location of the BCG. The emission visible in the lower left corner is artificial and associated with the ACIS chip boundaries. *Right panel:* LOFAR HBA radio image filtered with $\sigma_{\text{GGM}} = 12''$. The radio emission of H1 follows the spiral pattern of the X-ray cold fronts.

5. Discussion

In this section, we discuss the origin of the three components of the halo and the possible existence of a megahalo in A2142.

5.1. Probing the hybrid halo origin

The results collected in the previous sections can be exploited to discuss the origin and connection of each component of the puzzling hybrid radio source. The core has a roundish morphology, is confined by multiple cold fronts on scales of ~ 200 kpc, has a moderately steep spectrum $\alpha \sim 1.1$, and the slope of the radio versus X-ray correlation is sublinear (at least on scales

~ 100 – 200 kpc, see Fig. 13). These properties may support the hypothesis that H1 is a mini-halo originating from turbulence induced by sloshing of cold gas (e.g. ZuHone et al. 2013) or from hadronic collisions between protons confined within the cold fronts (e.g. ZuHone et al. 2015). Nevertheless, the event that triggered the central sloshing is unclear. Rossetti et al. (2013) and Venturi et al. (2017) proposed the hypothesis of an intermediate mass-ratio merger (possibly involving the group of BCG2). Such an event is less disruptive than a major merger and would likely not affect the global properties and morphology of the ICM. On the other hand, strong sloshing motions could be induced by this kind of merger that could be able to move cold gas away from the centre, where mixing with the hotter gas phase

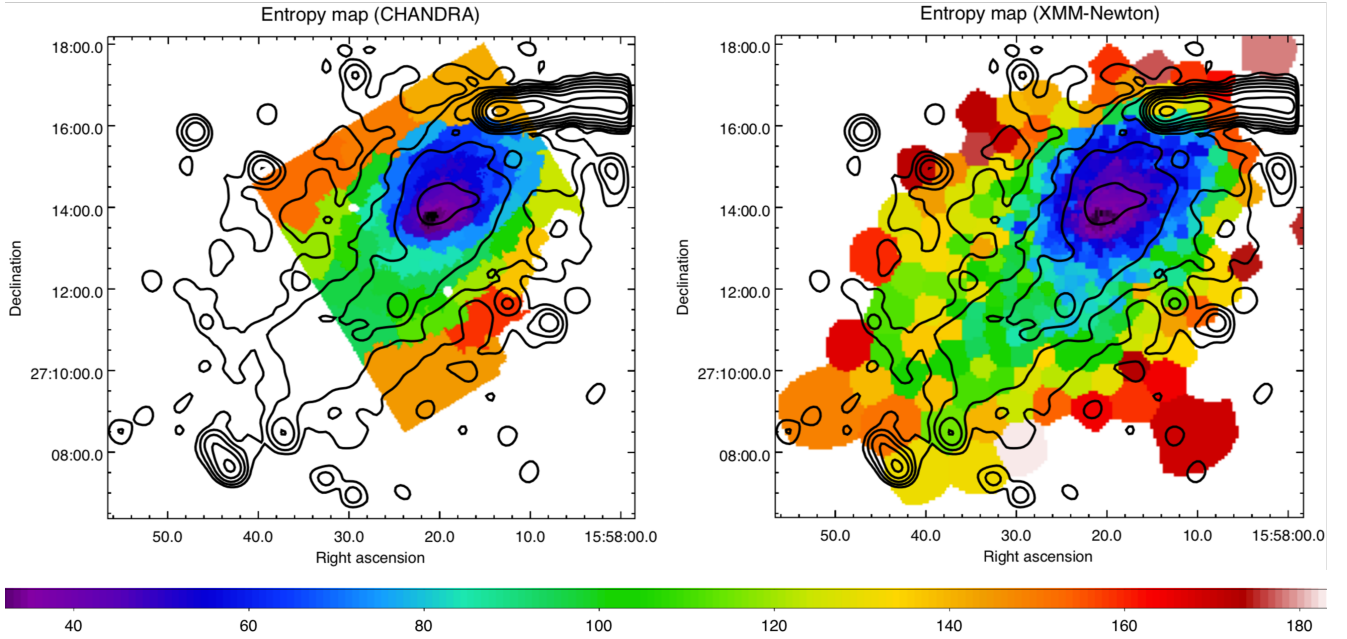


Fig. 11. Projected entropy maps of A2142 derived from *Chandra* (left) and *XMM-Newton* (right) in units of $\text{keV cm}^{5/3} \text{ arcmin}^{-2/3}$. Black contours are those of Fig. 2, upper right panel.

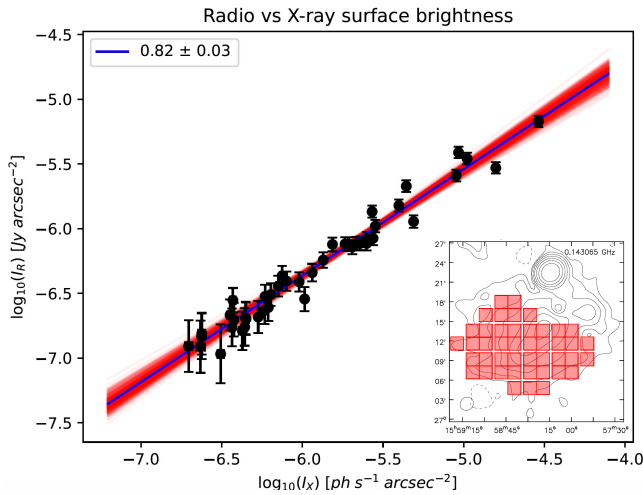


Fig. 12. Point-to-point radio (LOFAR HBA) vs. X-ray (*XMM-Newton*) analysis. The linear best fit provides a slope of $k = 0.82 \pm 0.03$. The sampling square boxes of $128''$ on a side are overlaid on the source-subtracted LOFAR HBA contours at $128'' \times 117''$, which are shown in the inset.

is then favoured (Wang & Markevitch 2018). This could lead to the formation of the prominently offset warm core from a former cool core; moreover, BCG1 would be deprived of most of the available fuel, thus explaining the absence of a powerful central AGN as in classical relaxed clusters.

The extreme sloshing motions in A2142 likely involve not only the inner regions of the cluster, but also larger scales, as witnessed by the presence of the southernmost cold front at ~ 1 Mpc from the centre (Rossetti et al. 2013) and the trail of low-entropy gas. In this case, sloshing could be partly responsible for the formation of the ridge, which extends for ~ 400 kpc along the NW–SE axis. The spectral index of the ridge is similar ($\alpha \sim 1.2$) to that of the core, and the projected entropy distribution of the ICM suggests a common origin for H1 and H2.

Turbulent motions induced by sloshing may trigger the radio emission of the ridge by reaccelerating particles and amplifying magnetic fields along the low-entropy trail. It is worth noting that turbulence in the ridge might also be supplied by the infalling groups along the NW–SE axis, which is the direction of the main accretion filament of the A2142 supercluster (Einasto et al. 2015, 2018, 2020); numerical simulations are needed to investigate this scenario.

The observed properties of H3 (large size, $\alpha \sim 1.6$, exponential radial profile, co-spatial thermal and non-thermal emission, sub-linear k) allow us to classify it as a giant ultrastep-spectrum radio halo. Concerning the origin of H3, we discuss two possible scenarios connected to either (i) an old merger or (ii) multiple minor mergers.

The first scenario for the origin of H3 is related to the half-mass epoch (estimated to be ~ 4 Gyr ago; Einasto et al. 2018), when A2142 formed at the centre of the supercluster through energetic mergers. At the present epoch, the turbulence injected by this event may not have been completely dissipated (see also Donnert et al. 2013). Therefore, the origin of H3 could be associated with this old merger, and its ultrastep spectrum could be the consequence of reacceleration in regions where turbulence is significantly dissipated. Furthermore, the hypothesis of an extremely old evolutionary stage of the halo may suggest an intriguing scenario to explain the peculiar thermal properties of A2142. Indeed, we might speculate that a former cool core was first disrupted by the old merger, and in the last few gigayears after this event began to reform again; nevertheless, this process is hampered by the perturbing minor mergers, which explains the presence of a warm core rather than a classical cool core.

As an alternative to the old merger scenario, the ultrastep spectrum of H3 could be originated by minor mergers (e.g. Cassano et al. 2006). It is unlikely that a single intermediate mass-ratio merger (as discussed for H1 and H2) is able to inject sufficient turbulent energy via sloshing to uniformly reaccelerate electrons on the large scales ($\gtrsim 2$ Mpc) of H3. However, a high number of minor mergers occurring over the cluster volume may be able to continuously supply turbulence and power

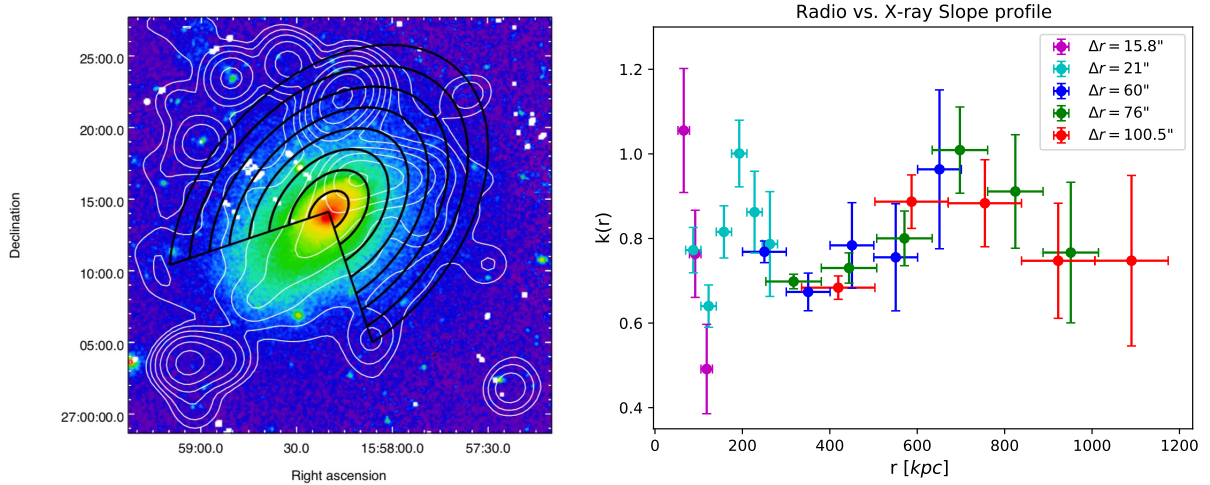


Fig. 13. Radial trends of the radio vs. X-ray correlation. *Left:* example of images (LOFAR HBA contours at $134''$ are overlaid on the *XMM-Newton* image) and elliptical sectors used to measure the surface brightnesses; for ease of inspection, the shown images include the discrete sources, whereas these were subtracted (or masked in case of residuals) during analysis. *Right:* radio vs. X-ray correlation slope as a function of the radial distance from BCG1. The widths of the concentric elliptical sampling sectors depend on the beam of the radio image and the required S/N, and are reported in the legend. The global trend $k(r)$ is sublinear, albeit with two narrow peaks.

the radio emission of H3 by reaccelerating electrons with a lower efficiency than turbulence induced by major mergers.

The uncommon thermal and non-thermal properties of A2142 offer the possibility to discuss the hypothesis of an evolutionary connection between H1, H2, and H3. By considering the model proposed by Zandanel et al. (2014), H1 and H3 might have a hadronic and turbulent reacceleration origin, respectively. Secondary CRe might significantly contribute to the emission of H1, but in this case the sublinear value of k on scales of $\sim 100\text{--}200\text{ kpc}$ would indicate particular conditions, namely a flat CRp distribution in the core and a strong ($B > B_{\text{CMB}}$, where $B_{\text{CMB}} \sim 4\text{ }\mu\text{G}$ is the equivalent magnetic field of the Cosmic Microwave Background at the cluster redshift) magnetic field (Brunetti & Jones 2014; Ignesti et al. 2020). Conversely, the observed properties (size, correlation slope, and spectral index) of H2 and H3 indicate that the secondary CRe contribute less of the emission in these regions, and the reacceleration scenario is therefore strongly favoured. In addition, H2 might have been the inner part of H3 in the past, hence implying the same origin of the two components; in this case, ICM turbulence could have then reshaped the inner emission into the present ridge.

Finally, A2142 is an interesting target to search for the origin of seed relativistic particles. Indeed, the possibility that (aged) particles originating in the core are transported and reaccelerated by sloshing turbulence along the ridge (Brunetti & Jones 2014) is a reasonable scenario that would be in line with thermal features, such as the low-entropy trail and the southernmost cold front. Furthermore, the seed particles to be reaccelerated in H2 and H3 may be primary CRe injected into the ICM by AGN and galaxy members of the groups infalling in A2142, as suggested by the prominent tails of T1 and T2.

5.2. A megahalo in A2142?

Megahalos are diffuse sources exhibiting a radial decline in the radio brightness that is shallower than that of their embedded radio halos (Cuciti et al. 2022). In addition, the four megahalos discovered so far share a large extension ($\sim R_{500}$) and an ultra-steep spectrum, and are found in disturbed clusters. Although the origin of megahalos is still poorly constrained, preliminary

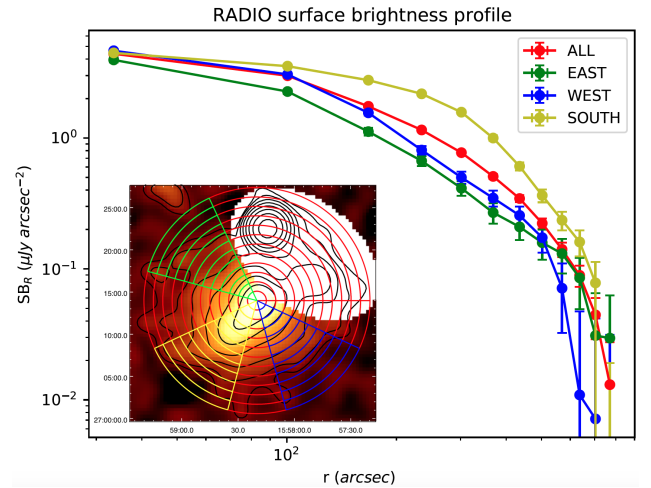


Fig. 14. Azimuthally averaged surface brightness profile of A2142 at $134'' \times 134''$ resolution, up to $0.95 R_{500}$. The inset reports the four sectors considered to extract the profiles: $0^\circ\text{--}360^\circ$ (red), $115^\circ\text{--}165^\circ$ (green), $205^\circ\text{--}255^\circ$ (yellow), $285^\circ\text{--}335^\circ$ (blue).

numerical simulations suggest that a baseline level of large-scale turbulence induced by the continuous accretion of matter onto the cluster may be common to all clusters, regardless of their dynamical state. As all megahalos were detected in massive ($M_{500} \sim 6\text{--}11 \times 10^{14} M_\odot$) clusters at intermediate to low redshift ($z \sim 0.17\text{--}0.28$), we searched for the presence of a megahalo in A2142, which is a very massive and nearby cluster as well.

To investigate the presence of a megahalo in A2142, throughout this section, we consider the 143 MHz, source-subtracted image (residuals from T1 and T2 were masked) at $134''$ resolution. We extracted the surface brightness profile in circular annuli of $67''$ in width (i.e. half the FWHM of the restoring beam) up to the 2σ contour level in different directions to check for possible biases associated with geometrical effects and/or residuals from subtraction. The extraction sectors and computed profiles are reported in Fig. 14. These do not show the typical shallower trend associated with megahalos.

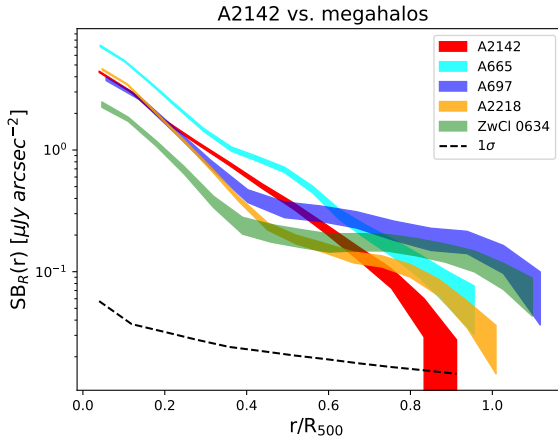


Fig. 15. Surface brightness profiles of A2142 and known megahalos (from Cuciti et al. 2022) as a function of the radial distance (normalised by R_{500}). The dashed line indicates the 1σ detection level of each bin for our observations. Different shapes are observed for A2142 and the other clusters.

In Fig. 15, we compare the surface brightness as a function of the radial distance (normalised by R_{500}) between A2142 and the four known megahalos. This plot shows that the shallower component arises on scales $r \sim 0.4\text{--}0.5 R_{500}$ for megahalos. In contrast, the profile of A2142 can be reproduced by two exponential laws (see Sect. 4.3) describing H1 and H3; the observed discontinuities in the profile at $r \sim 0.12 R_{500}$ and $r \sim 0.20 R_{500}$ mark the transitions from H1 to H2 and from H2 to H3, respectively. A possible megahalo in A2142 may have been detected by our deep observations beyond the exponential cutoff ($r \sim 0.9 R_{500}$) with a significance of $\sim 2\sigma$. In summary, Figs. 14 and 15 provide no evidence for a megahalo in A2142. As a sanity check, we also measured the surface brightness profile of A2142 from elliptical sectors (not shown); the shapes of the profiles from circular and elliptical annuli are consistent, which further supports our conclusions.

It is worth discussing whether the undetection of a megahalo is the result of calibration issues, or a combination of inadequate uv -coverage at short spacings and noise sensitivity. As shown by simulations in Shimwell et al. (2022), faint and extended emission may be partially lost if unmodelled during calibration steps. On the other hand, faint radio emission in nearby clusters – extending on even larger angular scales than those of A2142 – has been well recovered (e.g. Rajpurohit et al. 2021a; Bonafede et al. 2022; Botteon et al. 2022). Therefore, it is unlikely that possible emission from our target in the form of a megahalo was completely lost during calibration. For imaging, we adopted a minimum baseline of 50λ , which can recover sources in the sky of $\sim 1.15^\circ$ in size, corresponding to ~ 7 Mpc at the cluster redshift. Based on the results derived in Bruno et al. (2023) with a standard minimum baseline of 80λ , we should safely observe radio emission up to ~ 4 Mpc with losses of $\ll 20\%$ at most. Therefore, we can exclude the possibility that a megahalo on scales of $\gg R_{500}$ is not detectable due to missing short spacings.

6. Summary and conclusions

The main member of the A2142 supercluster is the nearby and massive galaxy cluster A2142, whose dynamical state is intermediate between a classical major merger and a relaxed cool core. A2142 was known to host a hybrid radio halo with two

distinct components (Venturi et al. 2017), namely the core (H1) and the ridge (H2). The core is roundish and confined by a system of cold fronts within the region of the warm X-ray core, whereas the ridge is elongated and roughly co-spatial with a trail of low-entropy thermal gas that ends in an extremely distant cold front (~ 1 Mpc from the centre). Furthermore, the ridge is aligned with the main accretion filament of A2142. An intermediate mass-ratio merger was invoked to explain the uncommon thermal features of A2142 (Rossetti et al. 2013). This event could lead to sloshing involving regions beyond the cluster core, without affecting the global morphology of the ICM.

In this work, we present new deep LOFAR HBA and LOFAR LBA observations allowing us to detect a new radio component (H3) whose emission follows the X-ray thermal distribution of the ICM up to scales of $\gtrsim 2$ Mpc. We investigated the properties, origin, and physical connection of the three components with complementary archival GMRT/uGMRT, VLA, *Chandra*, and *XMM-Newton* data. Our main results can be summarised as follows.

1. The properties of H1 (including the measured spectral index $\alpha = 1.09 \pm 0.02$) support the possibility that radio emission in this region is powered by sloshing in the core, which can generate turbulence and magnetic fields that confine CRp and/or reaccelerate relativistic electrons, as is thought to be the case for typical mini-halos.
2. The spectral index ($\alpha = 1.15 \pm 0.02$), morphology, and connection with the low-entropy gas trail suggest that the ridge can be powered by turbulence from large-scale sloshing, which could also be able to move and shape gas beyond the core. Additional supplies of turbulence in H2 might come from merging groups along the main accretion filament of A2142.
3. The size ($2.4 \text{ Mpc} \times 2.0 \text{ Mpc}$), spectral index ($\alpha = 1.57 \pm 0.20$), exponential surface brightness profile, and sublinear radio/X-ray spatial correlation ($I_R \propto I_X^k$, where $k \sim 0.7\text{--}0.8$) allow us to classify H3 as a giant ultrastep-spectrum radio halo. We propose two different scenarios for the origin of H3. In the first case, old mergers (~ 4 Gyr ago) from which the main structure of A2142 originate could have disrupted the former cool core and triggered the radio halo, which at present is observed in an advanced evolutionary phase. In this scenario, minor mergers may contribute by hampering the restoration of a classical cool core. As an alternative possibility, turbulence could be injected over the cluster volume by continuous minor mergers with small groups of galaxies. These mergers would inefficiently reaccelerate particles, leading to ultrastep spectra as predicted by turbulent reacceleration models. In both these scenarios, H2 could have been the inner part of H3, which was then reshaped by turbulence into the present ridge.
4. We investigated the possibility that A2142 hosts a megahalo, owing to a favourable combination of mass and redshift that is similar to those of the four confirmed cases reported by Cuciti et al. (2022). The analysis of the surface-brightness profiles in A2142 and in the literature megahalos does not support the presence of a megahalo in our target.

Our work primarily falls within the context of complex non-thermal phenomena occurring on the very large scales of the Universe, but could be used to probe the evolution of galaxy clusters. Indeed, we show that the hybrid halo emission of A2142 can be interpreted in terms of a particular stage of the development of the A2142 supercluster system. Our analysis suggests that various dynamical processes operating not only on

different spatial scales but also on different timescales could be responsible for the formation and evolution of multi-component halos. In this respect, radio observations can offer a valid perspective to investigate the complex dynamical history of galaxy clusters, from their innermost regions up to their outskirts.

Thanks to the high sensitivity provided by present and next-generation radio facilities, the number of hybrid halos is increasing (Lusetti et al. 2023; Biava et al., in prep.). By means of radio observations of clusters in an intermediate dynamical state, the relative roles of major and minor mergers and primary and secondary CRE in shaping thermal and non-thermal properties can be further constrained, and possibly tested by numerical simulations that follow the evolution of large-scale structures.

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References

Arnaud, K. A. 1996, in *Astronomical Data Analysis Software and Systems V*, eds. G. H. Jacoby, & J. Barnes, *ASP Conf. Ser.*, **101**, 17
 Astropy Collaboration (Robitaille, T. P., et al.) 2013, *A&A*, **558**, A33
 Astropy Collaboration (Price-Whelan, A. M., et al.) 2018, *AJ*, **156**, 123
 Beresnyak, A., Xu, H., Li, H., & Schlickeiser, R. 2013, *ApJ*, **771**, L31
 Bertocco, S., Goz, D., Tornatore, L., et al. 2020, in *Astronomical Data Analysis Software and Systems XXIX*, eds. R. Pizzo, E. R. Deul, J. D. Mol, J. de Plaa, & H. Verkouter, *ASP Conf. Ser.*, **527**, 303
 Biava, N., de Gasperin, F., Bonafede, A., et al. 2021, *MNRAS*, **508**, 3995
 Blasi, P., & Colafrancesco, S. 1999, *Astrophys. J.*, **12**, 169
 Bonafede, A., Brügggen, M., van Weeren, R., et al. 2012, *MNRAS*, **426**, 40

Bonafede, A., Intema, H. T., Brügggen, M., et al. 2014a, *ApJ*, **785**, 1
 Bonafede, A., Intema, H. T., Brügggen, M., et al. 2014b, *MNRAS*, **444**, L44
 Bonafede, A., Brunetti, G., Rudnick, L., et al. 2022, *ApJ*, **933**, 218
 Botteon, A., Shimwell, T. W., Bonafede, A., et al. 2018, *MNRAS*, **478**, 885
 Botteon, A., van Weeren, R. J., Brunetti, G., et al. 2020, *MNRAS*, **499**, L11
 Botteon, A., van Weeren, R. J., Brunetti, G., et al. 2022, *Sci. Adv.*, **8**, eabq7623
 Botteon, A., Markevitch, M., van Weeren, R. J., Brunetti, G., & Shimwell, T. W. 2023, *A&A*, **674**, A53
 Briggs, D. S. 1995, *Am. Astron. Soc. Meet. Abstr.*, **187**, 112.02
 Brunetti, G., & Blasi, P. 2005, *MNRAS*, **363**, 1173
 Brunetti, G., & Jones, T. W. 2014, *Int. J. Mod. Phys. D*, **23**, 1430007
 Brunetti, G., & Lazarian, A. 2007, *MNRAS*, **378**, 245
 Brunetti, G., & Lazarian, A. 2011, *MNRAS*, **410**, 127
 Brunetti, G., & Lazarian, A. 2016, *MNRAS*, **458**, 2584
 Brunetti, G., Setti, G., Feretti, L., & Giovannini, G. 2001, *MNRAS*, **320**, 365
 Brunetti, G., Giacintucci, S., Cassano, R., et al. 2008, *Nature*, **455**, 944
 Brunetti, G., Zimmer, S., & Zandanel, F. 2017, *MNRAS*, **472**, 1506
 Bruno, L., Rajpurohit, K., Brunetti, G., et al. 2021, *A&A*, **650**, A44
 Bruno, L., Brunetti, G., Botteon, A., et al. 2023, *A&A*, **672**, A41
 Cash, W. 1979, *ApJ*, **228**, 939
 Cassano, R., Brunetti, G., & Setti, G. 2006, *MNRAS*, **369**, 1577
 Cassano, R., Ettori, S., Giacintucci, S., et al. 2010a, *ApJ*, **721**, L82
 Cassano, R., Brunetti, G., Röttgering, H. J. A., & Brügggen, M. 2010b, *A&A*, **509**, A68
 Cassano, R., Ettori, S., Brunetti, G., et al. 2013, *ApJ*, **777**, 141
 Cassano, R., Cuciti, V., Brunetti, G., et al. 2023, *A&A*, **672**, A43
 Cavagnolo, K. W., Donahue, M., Voit, G. M., & Sun, M. 2009, *ApJS*, **182**, 12
 Chandra, P., Ray, A., & Bhatnagar, S. 2004, *ApJ*, **612**, 974
 Condon, J. J., Cotton, W. D., Greisen, E. W., et al. 1998, *AJ*, **115**, 1693
 Cuciti, V., Cassano, R., Brunetti, G., et al. 2015, *A&A*, **580**, A97
 Cuciti, V., Cassano, R., Brunetti, G., et al. 2021, *A&A*, **647**, A51
 Cuciti, V., de Gasperin, F., Brügggen, M., et al. 2022, *Nature*, **609**, 911
 Cuciti, V., Cassano, R., Sereno, M., et al. 2023, *A&A*, submitted [arXiv:2305.04564]
 Dallacasa, D., Brunetti, G., Giacintucci, S., et al. 2009, *ApJ*, **699**, 1288
 de Gasperin, F., Dijkema, T. J., Drabent, A., et al. 2019, *A&A*, **622**, A5
 de Gasperin, F., Vink, J., McKean, J. P., et al. 2020a, *A&A*, **635**, A150
 de Gasperin, F., Brunetti, G., Brügggen, M., et al. 2020b, *A&A*, **642**, A85
 de Gasperin, F., Williams, W. L., Best, P., et al. 2021, *A&A*, **648**, A104
 de Gasperin, F., Edler, H. W., Williams, W. L., et al. 2023, *A&A*, **673**, A165
 de Jong, J. M. G. H. J., van Weeren, R. J., Botteon, A., et al. 2022, *A&A*, **668**, A107
 Dennison, B. 1980, *ApJ*, **239**, L93
 Di Gennaro, G., van Weeren, R. J., Cassano, R., et al. 2021, *A&A*, **654**, A166
 Dolag, K., & EnBlin, T. A. 2000, *A&A*, **362**, 151
 Donnert, J., Dolag, K., Brunetti, G., & Cassano, R. 2013, *MNRAS*, **429**, 3564
 Duchesne, S. W., Johnston-Hollitt, M., & Wilber, A. G. 2021, *PASA*, **38**, e031
 Eckert, D., Gaspari, M., Vazza, F., et al. 2017a, *ApJ*, **843**, L29
 Eckert, D., Ettori, S., Pointecouteau, E., et al. 2017b, *Astron. Nachr.*, **338**, 293
 Edler, H. W., de Gasperin, F., Brunetti, G., et al. 2022, *A&A*, **666**, A3
 Einasto, M., Gramann, M., Saar, E., et al. 2015, *A&A*, **580**, A69
 Einasto, M., Deshev, B., Lietzen, H., et al. 2018, *A&A*, **610**, A82
 Einasto, M., Deshev, B., Tenjes, P., et al. 2020, *A&A*, **641**, A172
 Farnsworth, D., Rudnick, L., Brown, S., & Brunetti, G. 2013, *ApJ*, **779**, 189
 Fujita, Y., & Ohira, Y. 2013, *MNRAS*, **428**, 599
 Fujita, Y., Kohri, K., Yamazaki, R., & Kino, M. 2007, *ApJ*, **663**, L61
 Ghirardini, V., Ettori, S., Eckert, D., et al. 2018, *A&A*, **614**, A7
 Giovannini, G., & Feretti, L. 2000, *New Astron.*, **5**, 335
 Govoni, F., EnBlin, T. A., Feretti, L., & Giovannini, G. 2001, *A&A*, **369**, 441
 Govoni, F., Murgia, M., Markevitch, M., et al. 2009, *A&A*, **499**, 371
 Govoni, F., Orrù, E., Bonafede, A., et al. 2019, *Science*, **364**, 981
 Gramann, M., Einasto, M., Heinämäki, P., et al. 2015, *A&A*, **581**, A135
 Grant, C. E., Bautz, M. W., Ford, P. G., & Plucinsky, P. P. 2014, in *Space Telescopes and Instrumentation 2014: Ultraviolet to Gamma Ray*, eds. T. Takahashi, J. W. A. den Herder, & M. Bautz, *SPIE Conf. Ser.*, **9144**, 91443Q
 Hales, S. E. G., Baldwin, J. E., & Warner, P. J. 1988, *MNRAS*, **234**, 919
 Hales, S. E. G., Masson, C. R., Warner, P. J., & Baldwin, J. E. 1990, *MNRAS*, **246**, 256
 Hardcastle, M. J., Shimwell, T. W., Tasse, C., et al. 2021, *A&A*, **648**, A10
 Harris, C. R., Millman, K. J., van der Walt, S. J., et al. 2020, *Nature*, **585**, 357
 Henry, J. P., & Briel, U. G. 1996, *ApJ*, **472**, 137
 Hoang, D. N., Shimwell, T. W., van Weeren, R. J., et al. 2019, *A&A*, **622**, A20
 Hunter, J. D. 2007, *Comput. Sci. Eng.*, **9**, 90
 Ignesti, A. 2022, *New Astron.*, **92**, 101732
 Ignesti, A., Brunetti, G., Gitti, M., & Giacintucci, S. 2020, *A&A*, **640**, A37
 Intema, H. T., van der Tol, S., Cotton, W. D., et al. 2009, *A&A*, **501**, 1185
 Kale, R., & Ishwara-Chandra, C. H. 2021, *Exp. Astron.*, **51**, 95
 Kale, R., Venturi, T., Giacintucci, S., et al. 2013, *A&A*, **557**, A99

- Kale, R., Shende, K. M., & Parekh, V. 2019, *MNRAS*, **486**, L80
- Kelly, B. C. 2007, *ApJ*, **665**, 1489
- Kravtsov, A. V., & Borgani, S. 2012, *ARA&A*, **50**, 353
- Liu, A., Yu, H., Diaferio, A., et al. 2018, *ApJ*, **863**, 102
- Lusetti, G., Bonafede, A., Lovisari, L., et al. 2023, *A&A*, submitted [arXiv:2308.01884]
- Macario, G., Venturi, T., Brunetti, G., et al. 2010, *A&A*, **517**, A43
- Markevitch, M., & Vikhlinin, A. 2007, *Phys. Rep.*, **443**, 1
- Markevitch, M., Ponman, T. J., Nulsen, P. E. J., et al. 2000, *ApJ*, **541**, 542
- Mazzotta, P., & Giacintucci, S. 2008, *ApJ*, **675**, L9
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, in *Astronomical Data Analysis Software and Systems XVI*, eds. R. A. Shaw, F. Hill, & D. J. Bell, *ASP Conf. Ser.*, **376**, 127
- Miniati, F. 2015, *ApJ*, **800**, 60
- Molendi, S., De Grandi, S., Rossetti, M., et al. 2023, *A&A*, **670**, A104
- Murgia, M., Govoni, F., Markevitch, M., et al. 2009, *A&A*, **499**, 679
- Oegerle, W. R., Hill, J. M., & Fitchett, M. J. 1995, *AJ*, **110**, 32
- Offringa, A. R., & Smirnov, O. 2017, *MNRAS*, **471**, 301
- Offringa, A. R., McKinley, B., Hurley-Walker, N., et al. 2014, *MNRAS*, **444**, 606
- Okabe, N., & Umetsu, K. 2008, *PASJ*, **60**, 345
- Owers, M. S., Nulsen, P. E. J., & Couch, W. J. 2011, *ApJ*, **741**, 122
- Pasini, T., Edler, H. W., Brüggen, M., et al. 2022, *A&A*, **663**, A105
- Perley, R. A., & Butler, B. J. 2013, *ApJS*, **204**, 19
- Petrosian, V. 2001, *ApJ*, **557**, 560
- Pfrommer, C., & Enßlin, T. A. 2004, *A&A*, **413**, 17
- Pfrommer, C., Enßlin, T. A., & Springel, V. 2008, *MNRAS*, **385**, 1211
- Pillastrini, G. C. B. 2019, *New Astron.*, **69**, 1
- Pinzke, A., Oh, S. P., & Pfrommer, C. 2017, *MNRAS*, **465**, 4800
- Planck Collaboration XXIV. 2016, *A&A*, **594**, A24
- Raja, R., Rahaman, M., Datta, A., et al. 2020, *MNRAS*, **493**, L28
- Rajpurohit, K., Brunetti, G., Bonafede, A., et al. 2021a, *A&A*, **646**, A135
- Rajpurohit, K., Vazza, F., van Weeren, R. J., et al. 2021b, *A&A*, **654**, A41
- Rajpurohit, K., Osinga, E., Brienza, M., et al. 2023, *A&A*, **669**, A1
- Riseley, C. J., Rajpurohit, K., Loi, F., et al. 2022, *MNRAS*, **512**, 4210
- Robitaille, T., & Bressert, E. 2012, *Astrophysics Source Code Library* [record ascl:1208.017]
- Rossetti, M., & Molendi, S. 2010, *A&A*, **510**, A83
- Rossetti, M., Ghizzardi, S., Molendi, S., & Finoguenov, A. 2007, *A&A*, **463**, 839
- Rossetti, M., Eckert, D., De Grandi, S., et al. 2013, *A&A*, **556**, A44
- Russo, S. A., Bertocco, S., Gheller, C., & Taffoni, G. 2022, *Astron. Comput.*, **41**, 100648
- Sanders, J. S. 2006, *MNRAS*, **371**, 829
- Sanders, J. S., Fabian, A. C., Taylor, G. B., et al. 2016a, *MNRAS*, **457**, 82
- Sanders, J. S., Fabian, A. C., Russell, H. R., Walker, S. A., & Blundell, K. M. 2016b, *MNRAS*, **460**, 1898
- Sarazin, C. L. 1986, *Rev. Mod. Phys.*, **58**, 1
- Savini, F., Bonafede, A., Brüggen, M., et al. 2018, *MNRAS*, **478**, 2234
- Savini, F., Bonafede, A., Brüggen, M., et al. 2019, *A&A*, **622**, A24
- Shimwell, T. W., Röttgering, H. J. A., Best, P. N., et al. 2017, *A&A*, **598**, A104
- Shimwell, T. W., Tasse, C., Hardcastle, M. J., et al. 2019, *A&A*, **622**, A1
- Shimwell, T. W., Hardcastle, M. J., Tasse, C., et al. 2022, *A&A*, **659**, A1
- Shweta, A., Athreya, R., & Sekhar, S. 2020, *ApJ*, **897**, 115
- Smirnov, O. M., & Tasse, C. 2015, *MNRAS*, **449**, 2668
- Taffoni, G., Becciani, U., Garilli, B., et al. 2020, in *Astronomical Data Analysis Software and Systems XXIX*, eds. R. Pizzo, E. R. Deul, J. D. Mol, J. de Plaa, & H. Verkouter, *ASP Conf. Ser.*, **527**, 307
- Taffoni, G., Becciani, U., Bonafede, A., et al. 2022, arXiv e-prints [arXiv:2201.11526]
- Tasse, C. 2014a, arXiv e-prints [arXiv:1410.8706]
- Tasse, C. 2014b, *A&A*, **566**, A127
- Tasse, C., Hugo, B., Mirmont, M., et al. 2018, *A&A*, **611**, A87
- Tasse, C., Shimwell, T., Hardcastle, M. J., et al. 2021, *A&A*, **648**, A1
- Tchernin, C., Eckert, D., Ettori, S., et al. 2016, *A&A*, **595**, A42
- van Haarlem, M. P., Wise, M. W., Gunst, A. W., et al. 2013, *A&A*, **556**, A2
- van Weeren, R. J., Williams, W. L., Hardcastle, M. J., et al. 2016, *ApJS*, **223**, 2
- van Weeren, R. J., Andrade-Santos, F., Dawson, W. A., et al. 2017, *Nat. Astron.*, **1**, 0005
- van Weeren, R. J., de Gasperin, F., Akamatsu, H., et al. 2019, *Space Sci. Rev.*, **215**, 16
- van Weeren, R. J., Shimwell, T. W., Botteon, A., et al. 2021, *A&A*, **651**, A115
- Venturi, T., Rossetti, M., Brunetti, G., et al. 2017, *A&A*, **603**, A125
- Wang, Q. H. S., & Markevitch, M. 2018, *ApJ*, **868**, 45
- Williams, W. L., van Weeren, R. J., Röttgering, H. J. A., et al. 2016, *MNRAS*, **460**, 2385
- Zandanel, F., Pfrommer, C., & Prada, F. 2014, *MNRAS*, **438**, 124
- ZuHone, J. A., Markevitch, M., Brunetti, G., & Giacintucci, S. 2013, *ApJ*, **762**, 78
- ZuHone, J. A., Brunetti, G., Giacintucci, S., & Markevitch, M. 2015, *ApJ*, **801**, 146

Appendix A: Spectral index error maps

where ΔS is computed as in Eq. 1.

Figures A.1 and A.2 report the error maps associated with the spectral index maps shown in Figs. 7, 8. Errors are obtained as

$$\Delta\alpha = \left| \frac{1}{\ln\left(\frac{\nu_1}{\nu_2}\right)} \right| \sqrt{\left(\frac{\Delta S_1}{S_1}\right)^2 + \left(\frac{\Delta S_2}{S_2}\right)^2}, \quad (\text{A.1})$$

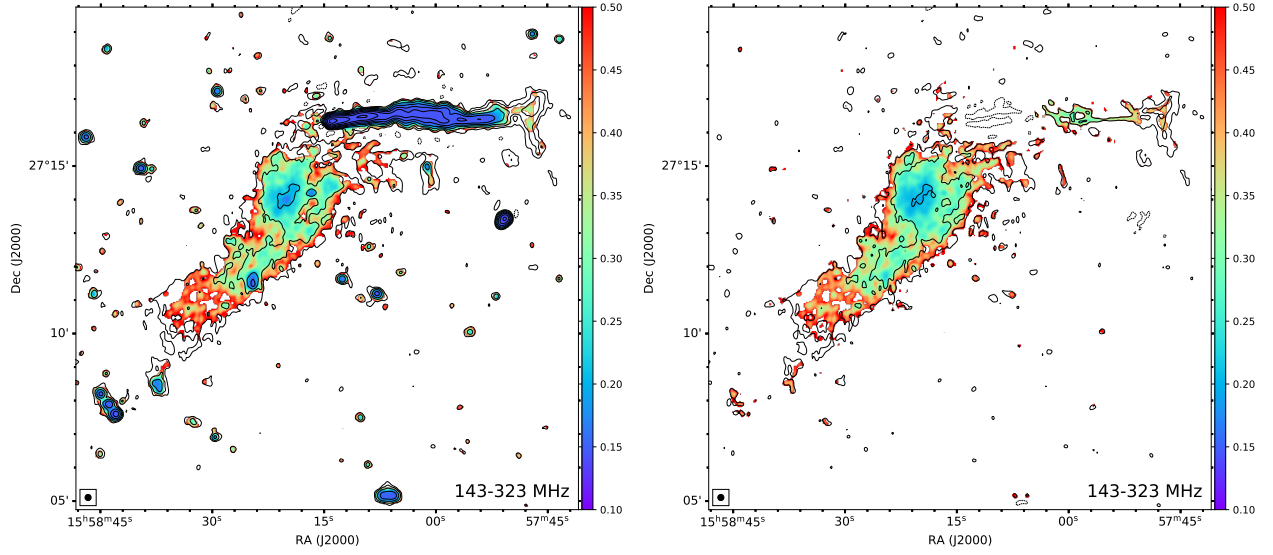


Fig. A.1. Spectral index error maps between 143 and 323 MHz corresponding to the maps reported in Fig. 7.

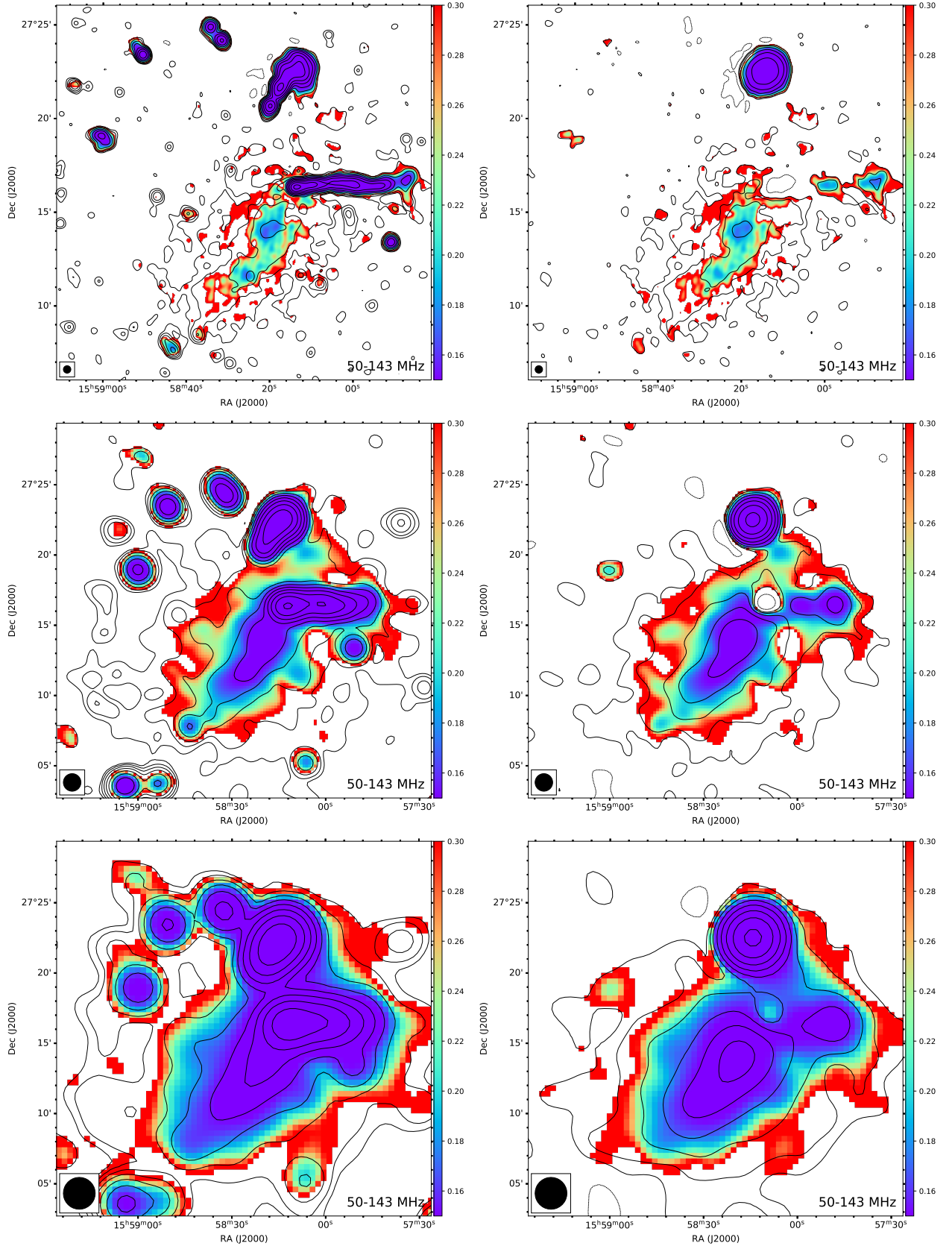


Fig. A.2. Spectral index error maps between 50 and 143 MHz corresponding to the maps reported in Fig. 8.

Appendix B: Thermodynamical maps

In the following, we describe the procedure adopted to derive the entropy map reported in Fig. 11, and associated systematic uncertainties. The fit of each spectrum provides a value of temperature kT (in units of keV) and normalisation $\mathcal{N}$, where the latter is a proxy for the density. Indeed, the normalisation of the *apec* model is defined as the integral of the squared numerical density ($n_e \times n_H$) over the volume:

$$\mathcal{N} = \frac{10^{-14}}{4\pi D_A^2 (1+z)^2} \int n_e n_H dV \quad [\text{cm}^{-5}], \quad (\text{B.1})$$

where D_A is the angular distance at the cluster redshift. We define the entropy as

$$s = kT \left(\frac{\mathcal{N}}{A} \right)^{-\frac{1}{3}} \quad [\text{keV cm}^{5/3} \text{ arcmin}^{-2/3}], \quad (\text{B.2})$$

where A is the area of each extracting region (in units of arcmin^2). As both the temperature and density are projected (or ‘pseudo’) quantities, the entropy has to be considered as a projected quantity as well.

As shown in Fig. B.1, we produced (pseudo-) temperature, normalisation, and entropy maps for both single and combined ACIS-I and ACIS-S chips to check for possible biases. The associated errors maps (not shown) are computed directly from the fitting errors for kT and $\mathcal{N}$, whereas uncertainties on s are

derived by assuming the standard error propagation formula as

$$\Delta s = s \sqrt{\left(\frac{k\Delta T}{kT} \right)^2 + \frac{1}{9} \left(\frac{\Delta \mathcal{N}}{\mathcal{N}} \right)^2}. \quad (\text{B.3})$$

In Fig. B.2 we compare the values of the fitted temperature (upper panels) and normalisation (lower panels) obtained separately from ACIS-I and ACIS-S. The fitted normalisations are in close agreement, i.e. within $\sim 5\%$; on the other hand, the fitted temperatures show a larger scatter, with typical differences of $\sim 10\% - 20\%$, and up to $\sim 30\%$. Some intrinsic scatter is expected given the different levels of degradation of the Chandra chips as a function of time (Grant et al. 2014). In principle, spectra from the earlier ACIS-I observation are expected to provide more accurate fitted temperatures than spectra from the more recent ACIS-S observations; nevertheless, the ACIS-S observations are deeper than the ACIS-I observation by a factor 3.5. Therefore, when the spectra of all the chips are fitted simultaneously, the temperatures are primarily driven by the ACIS-S observations, and statistical errors reduce. Despite this bias unavoidably affects the entropy, we recover consistent trends for both single and combined cases in the inner regions of interest, where $s \lesssim 120 \text{ keV cm}^{5/3} \text{ arcmin}^{-2/3}$; fitting errors for the combined case are $\Delta s \lesssim 5 \text{ keV cm}^{5/3} \text{ arcmin}^{-2/3}$ and $\Delta s \lesssim 10 \text{ keV cm}^{5/3} \text{ arcmin}^{-2/3}$ in the inner and outer regions, respectively. Therefore, we can genuinely rely on our results, which moreover are supported by the agreement with those of Rossetti et al. (2013) derived from *XMM-Newton* data.

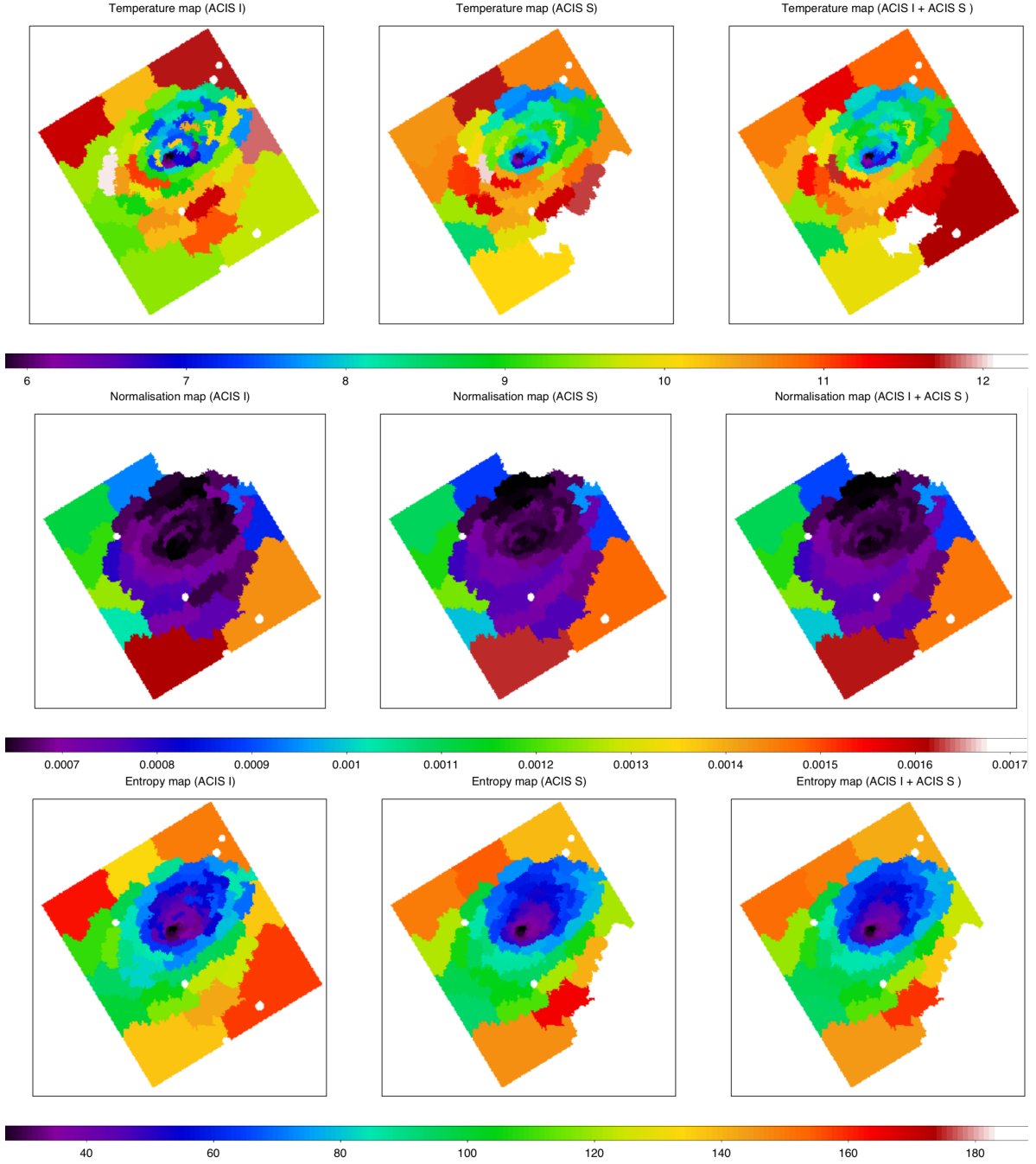


Fig. B.1. Projected thermodynamical maps of A2142 obtained with Chandra ACIS-I only (left columns), ACIS-S only (middle columns), and combined ACIS-I plus ACIS-S (right columns): temperature (top panels, in units of keV), apec normalisation (middle panels, in units of cm⁻⁵), and entropy (bottom panels, in units of keV cm^{5/3} arcmin^{-2/3}).

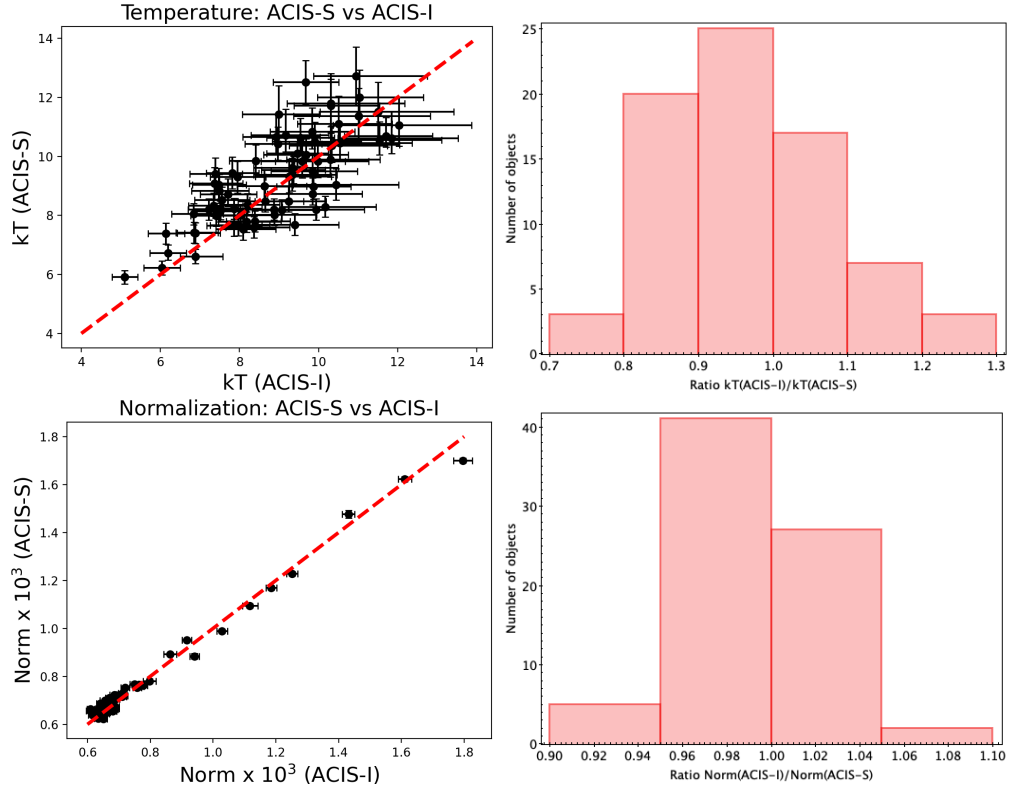


Fig. B.2. Comparison of fitted temperatures (upper panels) and normalisations (lower panels) derived from ACIS-I and ACIS-S separately. Left panels report the values, and the bisector line is plotted in red. Right panels report the distribution of the ratios of ACIS-I to ACIS-S values.