



Universiteit
Leiden
The Netherlands

The surface composition of six newly discovered chemically peculiar stars. Comparison to the HgMn stars mu Lep and beta Scl and the superficially normal B star nu cap

Monier, R.; Niemczura, E.; Kurtz, D.W.; Rappaport, S.; Bowman, D.M.; Murphy, S.J.; ... ; Le Ster, E.

Citation

Monier, R., Niemczura, E., Kurtz, D. W., Rappaport, S., Bowman, D. M., Murphy, S. J., ... Le Ster, E. (2023). The surface composition of six newly discovered chemically peculiar stars. Comparison to the HgMn stars mu Lep and beta Scl and the superficially normal B star nu cap. *The Astronomical Journal*, 166(2). doi:10.3847/1538-3881/acdb50

Version: Publisher's Version

License: [Creative Commons CC BY 4.0 license](https://creativecommons.org/licenses/by/4.0/)

Downloaded from: <https://hdl.handle.net/1887/3716672>

Note: To cite this publication please use the final published version (if applicable).



The Surface Composition of Six Newly Discovered Chemically Peculiar Stars. Comparison to the HgMn Stars μ Lep and β Scl and the Superficially Normal B Star ν Cap

Richard Monier¹, E. Niemczura², D. W. Kurtz^{3,4}, S. Rappaport⁵, D. M. Bowman⁶, Simon J. Murphy⁷, Yveline Lebreton^{1,8}, Remko Stuik⁹, Morgan Deal¹⁰, Thibault Merle¹¹, T. Kılıçoğlu¹², M. Gebran¹³, and Ewen Le Ster⁸

¹ LESIA, UMR 8109, Observatoire de Paris et Université Pierre et Marie Curie Sorbonne Universités, place J. Janssen, Meudon, France; Richard.Monier@obspm.fr

² University of Wrocław, Astronomical Institute, Kopernika 11, PL-51-622, Wrocław, Poland

³ Department of Physics, North-West University, Dr Albert Luthuli Drive, Mahikeng 2735, South Africa

⁴ Jeremiah Horrocks Institute, University of Central Lancashire, Preston, PR1 2HE, UK

⁵ Department of Physics, Kavli Institute for Astrophysics and Space Research, M.I.T., Cambridge, MA 02139, USA

⁶ Institute of Astronomy, KU Leuven, Celestijnenlaan 200D, B-3001 Leuven, Belgium

⁷ University of Southern Queensland, Centre for Astrophysics, West Street, Toowoomba, QLD 4350 Australia

⁸ Univ Rennes, CNRS, IPR (Institut de Physique de Rennes) - UMR 6251, F-35000 Rennes, France

⁹ Leiden Observatory, Leiden University, Niels Bohrweg 2, 2300 RA, Leiden, The Netherlands

¹⁰ Instituto de Astrofísica e Ciências do Espaço, Universidade do Porto, CAUP, Rua das Estrelas, PT4150-762 Porto, Portugal

¹¹ Institut d'astronomie et d'astrophysique, Université Libre de Bruxelles, CP. 226, Boulevard du Triomphe, B-1050 Brussels, Belgium

¹² Department of Astronomy and Space Sciences, Faculty of Science, Ankara University, 06100, Turkey

¹³ Department of Chemistry and Physics, Saint Mary's College, Notre Dame, IN 46556, USA

Received 2022 August 23; revised 2023 May 30; accepted 2023 May 30; published 2023 July 7

Abstract

We report on a detailed abundance study of six bright, mostly southern, slowly rotating late B-stars: HD 1279 (B8III), HD 99803 (B9V), HD 123445 (B9V), HD 147550 (B9V), HD 171961 (B8III), and HD 202671 (B5II/III), hitherto reported as normal stars. We compare them to the two classical HgMn stars μ Lep and β Scl, and to the superficially normal star ν Cap. In the spectra of the six stars, the Hg II line at 3984 Å line is clearly seen and numerous lines of P, Ti, Mn, Fe, Ga, Sr, Y, and Zr appear to be strong absorbers. A comparison of newly acquired and archival spectra of these objects with a grid of synthetic spectra for selected unblended lines reveals large overabundances of P, Ti, Cr, Mn, Sr, Y, Zr, Ba, Pt, and Hg, and underabundances of He, Mg, Sc, and Ni. The effective temperatures, surface gravities, low projected rotational velocities, and the peculiar abundance patterns of the six investigated stars show that they are new chemically peculiar stars, mostly new HgMn stars, and are reclassified as such. The evolutionary status of these stars has been inferred, and their ages and masses estimated. The two most massive objects, HD 1279 and HD 202671, might have evolved away from the main sequence recently, the other stars are main-sequence objects. HD 99803A is a sharp lined HgMn star with grazing eclipses. From TESS and MASCARA photometry, we determine an orbital period of $P_{\text{orb}} = 26.12022 \pm 0.00004$ day.

Unified Astronomy Thesaurus concepts: [Eclipsing binary stars \(444\)](#); [Chemically peculiar stars \(226\)](#)

Supporting material: machine-readable table

1. Introduction

The bright southern stars HD 99803 (HR 4423, B9V, $V = 5.14$), HD 123445 (HR 5294, B9V, $V = 6.17$), HD 147550 (HR 6096, B9V, $V = 6.24$), HD 171961 (HR 6990, B8III, $V = 5.77$), and HD 202671 (HR 8137, B5II/III, $V = 5.38$) have received little attention for their brightnesses: only 38, 57, 93, 35, and 79 references, respectively, can be found in SIMBAD for these stars.

We have recently undertaken a spectroscopic survey of all apparently slowly rotating late B-stars (B8–B9V) brighter than magnitude 7.0 and observable in the southern hemisphere which have $v_e \sin i \leq 60 \text{ km s}^{-1}$ and $\delta \leq -10^\circ$. We have added to these objects the northern star HD 1279 (HR 62, B8III, 53 references in SIMBAD) for which we obtained discretionary time at Observatoire de Haute Provence in January 2019.

This project addresses some fundamental questions of the physics of late B-stars. (i) Is our census of Chemically Peculiar (CP) stars complete up to the magnitude limits that we adopted? If not, then what are the physical properties of the newly found CP stars? (ii) Can we find new instances of rapid rotators seen pole-on (other than Vega) and study their physical properties (gradient of temperature across the disk, limb, and gravity darkening)?

Royer et al. (2014) carried out a similar exercise on a sample of 47 apparently slowly rotating A0–A1V stars in the northern hemisphere which satisfy $v_e \sin i \leq 60 \text{ km s}^{-1}$ and $\delta \geq -10^\circ$. Abundance analyses of selected lines allowed them to sort these stars into three groups: 13 chemically peculiar stars (among which four were new CPs), 17 superficially normal stars, and 17 spectroscopic binaries (Royer et al. 2014). Monier et al. (2015) reported on a first analysis of the abundances of Fe, Mn, and Hg, and showed that four members of that sample, i.e., HD 30085, HD 18104, HD 32867 and HD 53588, are newly discovered HgMn stars. Monier et al. (2018) have shown that HR 8844 is a new CP star, and is possibly an intermediate object between the Am and HgMn stars. Monier et al. (2019)



Original content from this work may be used under the terms of the [Creative Commons Attribution 4.0 licence](#). Any further distribution of this work must maintain attribution to the author(s) and the title of the work, journal citation and DOI.

provide detailed abundance analyses of HD 30085 and HD 30963, which highlight the HgMn status of these two stars.

In this paper, we report on abundance analyses from high resolution, high signal-to-noise, optical range échelle spectra of HD 1279 and five stars of the southern sample, i.e., HD 99803, HD 123445, HD 147550, HD 171961, and HD 202671. We also derive elemental abundances of two classical HgMn stars, μ Lep (HR 1702, HD 33904) and β Scl (HR 8937, HD 221507), which we use as comparison stars. We also use ν Cap (HR 7773, HD 193432), a superficially normal B9 V star, as a comparison.

Using model atmospheres and line synthesis, we derive for the first time the abundances of up to 40 chemical elements in these six stars and find that they depart strongly from solar. The overabundances of Mn, Sr, Y, Zr, Pt, and Hg, and the underabundances of He, Mg, Sc, and Ni establish that these stars are newly discovered CP stars, mostly new HgMn or PGa stars, whose atmospheres respectively harbor overabundances of manganese and mercury, and overabundances of phosphorus and gallium.

This paper is divided into seven sections. The first section recapitulates results from previous spectroscopic work on HD 1279 (HR 62), HD 99803 (HR 4423), HD 123445 (HR 5294), HD 147550 (HR 6096), HD 171961 (HR 6990), and HD 202671 (HR 8137), the second section describes new or archival spectroscopic observations of the eight stars. In the third section, we present the derivation of fundamental parameters, the modeling of the evolution of these new CP stars, including radiative diffusion and rotational mixing, and we provide estimates of their masses and ages. In the fourth section, we present frequency analyses of the recent light curves of these stars obtained with the Transiting Exoplanet Survey Satellite (Ricker et al. 2015) and we derive a rotational period when feasible. In Section 5, we discuss the model atmosphere analyses and the computations of synthetic spectra. In Section 6, we discuss the determination of elemental abundances for each star. In the conclusion, we discuss the chemical peculiarity of the six newly discovered chemically peculiar stars in the light of what is known of other HgMn stars.

2. Previous Spectroscopic Work on HD 1279, HD 99803, HD 123445, HD 147550, HD 171961, HD 202671, μ Lep, and β Scl

HD 1279 (HR 62) was ascribed a spectral type B7II by Garrison & Gray (1994). This star does not have a Renson nor a CCDM number: it was not previously known to be a chemically peculiar star or a spectroscopic binary. HD 99803 (HR 4423) is a visual binary with 13'' separation. Corbally (1984) classified HD 99803A as B9IV and HD 99803B as Am: kA3hA5VmA7 (meaning a marginal Am star); Monier et al. (2020) found HD 99803A to be a new HgMn star. HD 99803B is fainter ($V=7.77$) and cooler than HD 99803A; the lines of the two components are clearly separated in the HARPS spectrum that we have used. HD 123445 (HR 5294) is CCDM J14089-4328A, the other member of the system is much fainter and cooler ($V=12.52$). We see no signs of this companion in the spectra that we analyzed. HD 123445A (HR 5294) does not have any Renson member. HD 171961 (HR 6990) has neither a CCDM nor a Renson number. HD 147550 (HR 6096) is Renson 41660 and is classified as a B9 He-weak star. HD 202671 (HR 8137) is Renson 56480 and is classified as a B7 He-weak Mn star. Takeda et al. (2014) derived the

Table 1
Previous Abundance Studies for μ Lep (HR 1702) and β Scl

Chemical Element(s)	Reference	Star
Sr, Y	Dworetzky et al. (2008)	μ Lep and β Scl
Xe	Dworetzky et al. (2008)	μ Lep and β Scl
Ca	Cowley et al. (2007)	μ Lep
	Castelli & Hubrig (2004)	β Scl
Ga	Nielsen et al. (2005)	μ Lep
	Dworetzky et al. (1998)	μ Lep and β Scl
	Smith (1996)	μ Lep and β Scl
	Lanz et al. (1993)	β Scl
Hg	Dolk et al. (2003)	μ Lep and β Scl
	Woolf & Lambert (1999)	μ Lep
	Smith (1997)	β Scl
Cr	Savanov & Hubrig (2003)	μ Lep
Ne	Dworetzky & Budaj (2000)	μ Lep
N	Roby et al. (1999)	μ Lep
Mn	Jomaron et al. (1998)	μ Lep and β Scl
O	Takeda et al. (1999)	μ Lep
Cu, Zn	Smith (1994)	μ Lep and β Scl
Fe-peak	Smith & Dworetzky (1993)	μ Lep and β Scl
	Adelman et al. (1993)	μ Lep
Mg, Al, Si	Smith & Dworetzky (1993)	μ Lep and β Scl
C, N, O	Roby & Lambert (1990)	μ Lep
C	Catanzaro & Leone (2002)	β Scl

abundance of sodium and iron for this star, and Leone & Manfre (1997) derived the effective temperature. No extensive abundance study (i.e., for several elements) has been published for any of these six stars.

In contrast, the abundances of μ Lep (HR 1702) and β Scl (HR 8937) have been well studied. General abundance analyses of μ Lep and β Scl can be found in Adelman & Pintado (1997) and Thiam et al. (2010), respectively. Other works focus on the determination of the abundance of one element in several HgMn stars, and are collected in Table 1.

3. Observations

For the comparison stars μ Lep (HR 1702) and β Scl (HR 8937), we used archival FEROS spectra (Fiber-fed Extended Range Optical Spectrograph, $R=48,000$ ¹⁴) with good signal-to-noise (S/N) ratios. The spectra that we use were retrieved from the ESO Spectral archive.¹⁵

For the northern star HD 1279 (HR 62), one high resolution spectrum was obtained in service observing at Observatoire de Haute Provence with SOPHIE.¹⁶ The échelle spectrograph in its high resolution mode ($R=75,000$) yields a full spectral coverage from 3820 Å to 6930 Å in 39 orders. A detailed technical description of SOPHIE is given in Perruchot et al. (2008). SOPHIE is a cross-dispersed, environmentally stabilized échelle spectrograph dedicated to high precision radial velocity measurements. The spectra are extracted online from the detector images using a specific pipeline adapted from that of HARPS (High Accuracy Radial Velocity Planet Searcher¹⁷).

HD 171961 (HR 6990) was observed five times using the High Resolution Spectrograph (HRS;¹⁸ Crause et al. 2014)

¹⁴ <https://www.eso.org/sci/facilities/lasilla/instruments/feros.html>

¹⁵ http://archive.eso.org/wdb/wdb/adp/phase3_spectral/form

¹⁶ <http://www.obs-hp.fr/guide/sophie/sophie-eng.shtml>

¹⁷ <https://www.eso.org/sci/facilities/lasilla/instruments/harps.html>

¹⁸ <https://astronomers.salt.ac.za/instruments/hrs/>

Table 2
Log of Observed and Archival Spectra

Star ID	Observation Date	Instrument	R	Exposition Time (s)	S/N
μ Lep	08/12/2011	FEROS	48,000	60	517
	21/12/2019	HRS	65,000	60(V) +50(NIR)	450
β Scl	21/10/2005	FEROS	48,000	60	387
	21/12/2019	HRS	65,000	120(V) +80(NIR)	400
HD 1279	14/1/2019	SOPHIE	75,000	1200	209
HD 99803	29/4/2019	HRS	65,000	144(V) +118(NIR)	200
	07/4/2009	HARPS	115,000	200	305
HD 123445	02/2/2010	FEROS	48,000	450	277
HD 147550	24/3/2006	UVES	40,970	25	278
	24/3/2006	UVES	40,970	25	278
HD 171961	26/7/2019	HRS	65,000	206(V) +155(NIR)	220
	27/7/2019	HRS	65,000	206(V) +155(NIR)	320
	28/7/2019	HRS	65,000	206(V) +155(NIR)	190
	18/8/2019	HRS	65,000	206(V) +155(NIR)	230
	19/8/2019	HRS	65,000	206(V) +155(NIR)	240
HD 202671	21/08/2013	FEROS	48,000	600	457

behind the 11-m Southern African Large Telescope (SALT¹⁹). We have coadded the individual spectra. SALT HRS is a dual-beam (370–550 nm and 550–890 nm) fiber-fed, white-pupil, échelle spectrograph, employing VPH gratings as cross dispersers. The cameras are all-refractive. SALT HRS is an efficient single-object spectrograph using pairs of large (350 μ m and 500 μ m; 1''6 and 2''2) diameter optical fibers: one for the star and one for background sky. Three of the pairs feed image slicers before injection into the spectrograph, which delivers in its highest resolution mode a resolving power of $R \simeq 65,000$ (sliced 350 μ m fibers). A single 2k $\times$ 4k CCD captures all of the blue orders, while a 4k $\times$ 4k detector, using a fringe-suppressing deep-depletion CCD, is used for the red camera.

For the four remaining stars, i.e., HD 99803 (HR 4423), HD 123445 (HR 294), HD 147550 (HR 6096), and HD 202671 (HR 8137), we have used FEROS or HARPS archival spectra.

The observing dates, exposure times, and S/N at 5000 Å of the spectra are collected in Table 2. When several spectra were available for one star, we have used the spectrum with the highest S/N. In all cases, we have used the one-dimensional spectrum provided by the project.

4. Evolutionary Status and Rotational Periods Determination

4.1. Fundamental Parameters

The fundamental data for HD 1279 (HR 62), HD 99803A (HR 4423), HD 123445 (HR 5294), HD 147550 (HR 6096), HD 171961 (HR 6990), HD 202671 (HR 8137), β Scl (HR 8937), and μ Lep (HR 1702) are collected in Table 3. The spectral type retrieved from SIMBAD appears in row 2;

the apparent magnitudes appear in row 3; and the Strömgren indexes $b - y$, m_1 , and c_1 appear in rows 4, 5, 6. The photometric data were taken from Hauck & Mermilliod (1998).

For all stars, the effective temperature (T_{eff}) and surface gravity ($\log g$) were determined using the UVBYBETA code developed by Napiwotzki et al. (1993), and are collected in rows 7 and 8. This code is based on the Moon & Dworetsky (1985) grid, which calibrates the $uvby\beta$ photometry in terms of T_{eff} and $\log g$. The estimated errors on T_{eff} and $\log g$, are ± 125 K and ± 0.20 dex, respectively (see Section 4.2 in Napiwotzki et al. 1993).

The radial velocities were derived by cross-correlating Fe II lines with very accurate wavelengths from NIST in the range 4500 – 4600 Å with a synthetic spectrum computed for the derived effective temperature and surface gravity of the star and solar abundances. A parabolic fit to the upper part of the resulting cross-correlation function yields the Doppler shift, which is then used to shift the spectra to rest wavelengths. The projected rotational velocities were derived from the position of the first zero of the Fourier transform of individual lines, which are taken from Royer et al. (2014). The radial velocity, the microturbulent velocity, and projected equatorial velocity are collected in rows 8, 9, and 10 in Table 3.

4.2. Evolutionary Status and Mass and Age Determinations

To estimate the mass and age of the stars, we used the SPInS stellar model optimization tool (Lebreton & Reese 2020). SPInS uses a Bayesian approach to find the probability distribution function of stellar parameters from a set of constraints. At the heart of the code is a Markov Chain Monte Carlo solver coupled with interpolation within a precomputed stellar model grid. Here, we used the BaSTI-IAC grid of stellar models (Hidalgo et al. 2018). This grid is for a solar-scaled heavy element distribution with the solar mixture taken from Caffau et al. (2011) complemented by Lodders (2010), which corresponds to $(Z/X)_{\odot} = 0.0209$. The grid considers convective core overshooting included as an instantaneous mixing between Schwarzschild’s convective limit up to layers at a distance $\alpha_{\text{ov}} = 0.2H_P$ from it, where H_P is the pressure scale height at the Schwarzschild limit. Neither microscopic diffusion nor rotation are taken into account. We considered as a prior, Salpeter’s (1955) initial mass function.

We estimate the age and mass of each star on the basis of the Kiel diagram observational constraints, i.e., effective temperature T_{eff} and surface gravity $\log g$. We assume that the initial metallicity of the stars is close to solar, i.e., $[\text{Fe}/\text{H}] = 0.0 \pm 0.10$ dex. We provide the inferred mean ages A_{SPInS} and masses M_{SPInS} , and their 1σ errors in Table 3. The positions of the stars in the Kiel diagram are plotted in Figure 1, together with isochrones of age and metallicity inferred by SPInS. The figure also shows BaSTI-IAC stellar evolutionary tracks of solar metallicity with masses increasing, from right- to left-hand, from 2.0 to 7.5 $M_{\odot}$, by steps of 0.5 $M_{\odot}$.

The six stars are found to be young with ages between 100 and 261 Myr and masses in the range 2.8 – 4.9 $M_{\odot}$. For each couple of stars (HD 99803A (HR 4423), HD 147550 (HR 6096)) and (HD 123445 (HR 5294), β Scl (HR 8937)), stars have very close Kiel diagram parameters so that their ages and masses are very close. These stars lie on the main sequence, as so do HD 171961 (HR 6990) and μ Lep (HR 1702). For stars HD 1279 (HR 62) and HD 202671 (HR 8137), the situation is less clear. Considering the observed parameters, the stars are expected to be on the subgiant

¹⁹ <https://astronomers.salt.ac.za/>

Table 3
Adopted Fundamental Parameters for HD 1279, HD 99803A, HD 123445, HD 147550, HD 171961, HD 202671, β Scl, and μ Lep

		HD 1279	HD 99803	HD 123445	HD 147550	HD 171961	HD 202671	β Scl	μ Lep
Sp.T.		B8III	B9V	B9V	B9V	B8III	B5II/III	B9.5III	B9IV
V		5.852	5.136	6.170	6.241	5.772	5.383	4.370	3.290
$b - y$		−0.020	−0.012	−0.016	0.059	0.032	−0.048	−0.046	−0.043
m_1		0.092	0.122	0.109	0.109	0.074	0.096	0.125	0.106
c_1		0.588	0.944	0.746	1.035	0.558	0.547	0.677	0.653
T_{eff}	[K]	13300	10700	12000	10700	13800	13700	12500	12800
$\log g$	(cgs)	3.16	3.85	4.19	3.83	3.85	3.30	4.13	3.77
$v_e \sin i$	[km s ^{−1}]	28.0	8.0	66.0	4.8	55.0	25.0	26.0	18.0
$v_{\text{micr.}}$	[km s ^{−1}]	0.0	0.1	0.2	0.70	0.0	0.9	0.0	0.0
v_{rad}	[km s ^{−1}]	−9.20	3.0	−2.0	−24.0	−33.0	−12.50	0.40	27.10
[Fe/H] _s		0.18	−0.44	−0.52	0.05	0.00	0.40	0.23	0.23
M_{SPInS}	[$M_{\odot}$]	4.8 ± 0.4	2.8 ± 0.3	3.0 ± 0.2	2.8 ± 0.3	4.0 ± 0.4	4.9 ± 0.4	3.2 ± 0.2	3.7 ± 0.4
A_{SPInS}	(Myr)	106 ± 14	260 ± 63	119 ± 60	261 ± 62	100 ± 26	100 ± 12	110 ± 51	138 ± 28
Phase		SG/MS	MS	MS	MS	MS	SG/MS	MS	MS

Note. Uncertainties are 125 K on T_{eff} , 0.20 dex on $\log g$ and 0.10 dex on the present surface metallicity $[\text{Fe}/\text{H}]_s$. MS stands for main sequence and SG for subgiant.

branch but the SPInS inference places them close to the terminal age main sequence (TAMS) with surface gravities lower than the observed ones and error bars that do not overlap. More generally, it is interesting to note that for all stars but two (HD 123445 (HR 5294) and β Scl (HR 8937)), SPInS can find a solution corresponding to the observed effective temperature but for a $\log g$ -value higher than the observed one. It is difficult to identify the exact reason for this disagreement. On the theoretical side, these rather massive stars have very probably experienced rotation in the previous phases of their evolution, as well as microscopic diffusion and radiative accelerations on chemical elements. Rotation and gravitational settling bring fresh fuel to the H-burning core, which impacts the age-dating and enlarges the main sequence. Concerning rotation, it has been shown by Ekström et al. (2012) that rotation combined with a moderate amount of overshooting ($\alpha_{\text{ov}} = 0.1H_p$) have similar effects on the main-sequence width when compared with overshooting alone but with $\alpha_{\text{ov}} = 0.2H_p$, which is the amount considered in the BaSTI grids. Concerning diffusion, we calculated the evolution on the MS of a star of mass $M = 5.0 M_{\odot}$ with the cestam code including microscopic diffusion and radiative accelerations (Morel & Lebreton 2008; Marques et al. 2013; Deal et al. 2018), and checked that it has no impact for these stars. A precise modeling of each star is very demanding, and is therefore beyond the scope of this work, in particular because, on the observational side, the error on the values of $\log g$ is quite high, which hampers a precise characterization of these stars.

5. Analysis of Photometric Data

5.1. Analysis of the Photometric TESS and MASCARA Data of HD 99803 (HR 4423)

5.1.1. Frequency Analysis of the TESS Data- HD 99803 (HR 4423)

HD 99803 (TIC 163204899; HR 4423) was observed by TESS in Sector 10 in 2 minutes cadence.²⁰ The data are available in both simple aperture photometry (SAP) and presearch-data conditioning SAP (PDCSAP); we used the PDCSAP data for our analysis, converting intensity to magnitudes. The data have a time span of 24.96 days with a

center point in time of $t_0 = \text{BJD } 2458583.20612$, and comprise 15,992 data points, after clipping 51 outliers, which are probably the result of cosmic ray strikes (including solar radiation).

There is also a 2.6 days gap in the data at the time of perigee. The TESS orbit is a lunar-synchronous eccentric orbit with an orbital period of half the sidereal orbital period of the moon. This keeps the satellite from being perturbed significantly at apogee and allows for higher bandwidth data downloads at perigee. The gap in the data thus comes during the data transfers.

Figure 2 shows the light curve of the Sector 10 data where an eclipse is obvious. This eclipse was also noted by Kochukhov et al. (2021). The bottom panel of the figure shows a higher time resolution of the eclipse. HD 99803 (HR 4423) was observed again by TESS in Sector 36 two years later, in 10 minutes cadence, and another eclipse was observed. The S36 data are from Full Frame Images and the star was close to an edge in this sector, making extraction of the data problematic. Nevertheless, with assistance from Tibor Mitnyan (private communication) who used Andras Pál's FITSH package (Pál 2012), we were able to extract a time for the eclipse. We discuss the eclipses and the orbital period inferred from them in the next subsection.

The S10 data were analyzed using a fast Discrete Fourier Transform (Kurtz 1985) to produce an amplitude spectrum (Figure 3). There is no evidence of any p -mode pulsations out to the Nyquist frequency near 360 day^{-1} (not shown). In the low frequency range where g -modes occur there is no peak that cannot be ascribed to instrumental variations. Hence, there is also no evidence of g -modes. Comparison with the instability strips of Zdravkov & Pamyatnykh (2008) show that HD 99803A (HR 4423) is too cool to be either a β Cep star or an SPB star. There is a small peak at 0.6 day^{-1} that may appear to be of interest but when dividing the data into two subsets, this peak is most evident in the second half of the data set, and so it is not convincing as a pulsation frequency.

HD 99803A and B fall within one pixel in TESS observations. The problem of HD 99803A lying near to a chip edge in the S36 FFI data may indicate that HD 99803B actually contributed more of the combined light. This could have resulted in an apparently shallower eclipse, which we judge to be an artifact. There are also some low-amplitude δ Sct

²⁰ Available on the Mikulski Archive for Space Telescopes (MAST) via: doi: 10.17909/6jv0-4x21

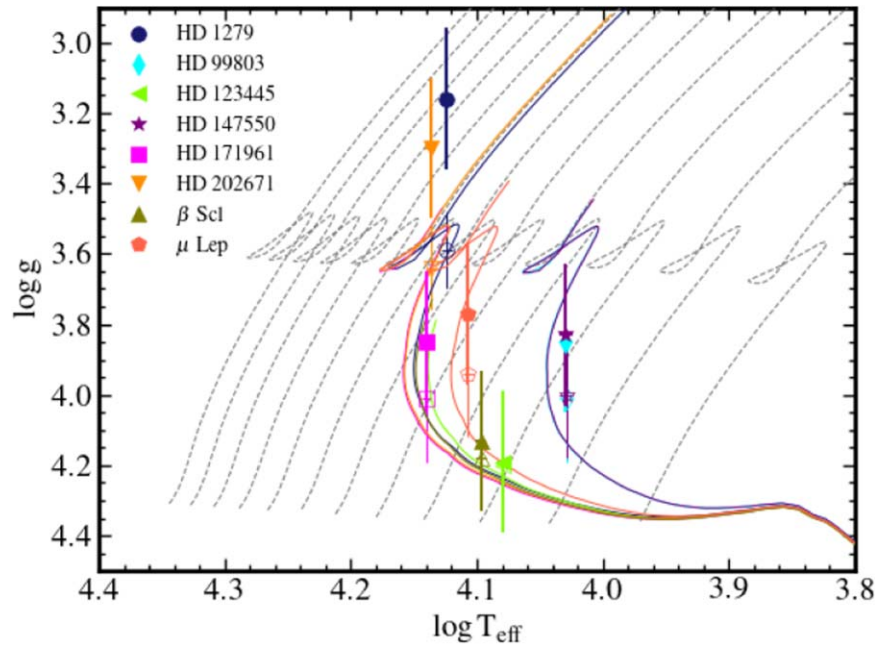


Figure 1. Positions of the six newly discovered chemically peculiar stars, and μ Lep and β Scl in an effective temperature-surface gravity Kiel diagram. Full symbols show the positions of the stars corresponding to the observed values of $\log g$ and T_{eff} (see Table 3), while open symbols are for the best stellar model values inferred by SPInS. Colored lines are isochrones calculated for the age A_{SPInS} inferred by SPInS. Gray-dashed lines are BaSTI-IAC evolutionary tracks of solar metallicity with masses increasing, from right- to left-hand, from 2.0 to 7.5 $M_{\odot}$, by steps of 0.5 $M_{\odot}$.

pulsations visible in the FT of the S36 data in the frequency range 17.7–32 day^{-1} . These very probably arise from the marginal Am star HD 99803B, which lies within the δ Sct instability strip. Many marginal Am stars are known to pulsate (Smalley et al. 2017). These data do not warrant further analysis at present. HD 99803 (HR 4423) will be observed again during TESS S63at, at that time the eclipse depth and p -mode pulsations can be examined further.

5.1.2. Eclipses Observed with TESS

The TESS Sector 10 eclipse occurred at BJD = 2, 458, 580.6328 ± 0.0002 , while a second eclipse was detected in Sector 36 at BJD = 2, 459, 285.8792 ± 0.0006 . There are no TESS observations between these two sectors, which are separated by two years. The difference between these two times of eclipse is 705.246 ± 0.001 day. Therefore, in the absence of any perturbing effects on the binary orbit, the orbital period is $P_{\text{orb}} = 705.246/N \pm 0.001/N$ days.

The eclipse depths are about 4% deep. For an equatorial occultation by the companion star, the inferred size of the companion star is approximately 20% of the radius of HD 99803A. Given that the latter is about $3 R_{\odot}$, this implies a companion radius of roughly $0.6 R_{\odot}$. If so, then this close companion would be orders of magnitude fainter than HD 99803A, and would not be expected to contribute noticeably to its spectrum. There would also be no readily detectable secondary eclipse.

Another consideration is that HD 99803A and B are co-moving Gaia stars where component B is 2.6 mag fainter than component A in G , with a $13''$ separation, which is some 2500 au on the plane of the sky. This fainter star is contained within the photometric aperture of HD 99803. However, if the eclipses were actually hosted by this co-moving companion, then the actual eclipse depth would be close to 50%. This is highly implausible.

5.1.3. Analysis of the MASCARA Photometry

We are also fortunate to have MASCARA coverage of this star over a nearly 6 yr interval. MASCARA (Multi-site All-Sky CAmERA; Talens et al. (2017)) and bRing (Stuik et al. 2017) have observed HD 99803 (HR 4423) since early-2017, providing $\sim 163,000$ photometric data points, each representing a set of 50 exposures of 6.4 s. After filtering the data, $\sim 115,000$ data points were used for further analysis. MASCARA and bRing have a pixel scale of approximately $1'$, while aperture photometry has been performed using a 2.5 pixel radius, with local background subtraction in an annulus of 6–21 pixels radius. The aperture photometry data have been calibrated using Talens et al. (2018) to remove most instrumental artifacts, mainly due to vignetting, PSF variations, and inter-pixel variations. Due to the use of fixed cameras, further strong systematics are known to exist at frequencies of 1 cycle per day, 1 cycle per sidereal day (and its harmonics), and synodic and sidereal lunar periods.

The first panel in Figure 4 shows the Box Least-Squares (BLS) transform of the MASCARA data. The largest peaks are associated with the synodic and sidereal periods of the lunar cycle, which affects the observations. There are also a number of peaks associated with subharmonics of the daily observing cycle. In addition to these non-astrophysical peaks, there is one significant peak at a period of 26.128 days, and its three higher harmonics, labeled 1–4. The second panel of Figure 4 shows a zoom-in of the BLS transform around the range of 25–31 days. As one can readily see, the peak at 26.128 days is much narrower than either peak associated with the lunar activity. Furthermore, the four harmonics of this period indicate that it represents a sharp feature in the light curve. The third panel in this figure shows a fold of the MASCARA data about the 26.128 days period. It is gratifying to see that it is a narrow eclipse that is about 4% deep, just as the TESS Sector 10 eclipse indicates.

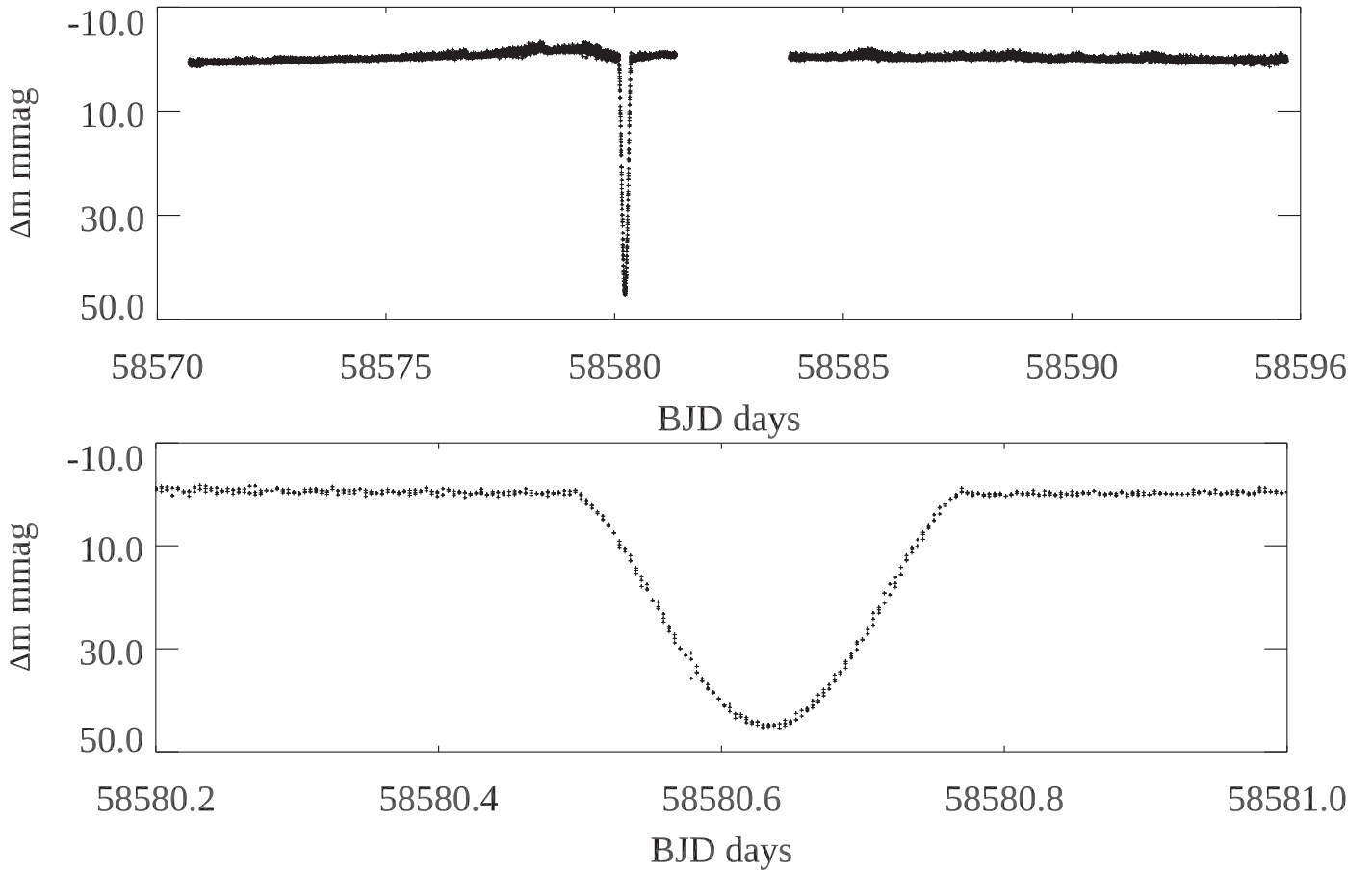


Figure 2. The light curve of HD 99803 obtained in 120 s cadence in Sector 10. The ordinate scale is Barycentric Julian Date—240 0000.0. The top panel is the full light curve; the bottom panel shows a higher time resolution view of the eclipse.

Finally, we note that the phase of the MASCARA eclipse agrees with the TESS eclipses to within $\lesssim 1\%$ of an orbital cycle. Therefore, the MASCARA period represents the same eclipsing behavior that we see in the TESS data. This, in turn, implies that the “missing” integer N in the expression above for the orbital period is 27. The accurate period determined from the time difference between the two TESS eclipses and the integer number of cycles between them is then 26.12022 ± 0.00004 days.

5.2. Frequency Analysis of TESS Data: HD 1279 (HR 62)

HD 1279 (TIC 440076466; HR 62) was observed by TESS in Sector 17 in 2 minutes cadence.²¹ The data are available in both SAP (simple aperture photometry) and PDCSAP (presearch-data conditioning SAP); we used the PDCSAP data for our analysis, converting intensity to magnitudes. The data have a time span of 22.85 days with a center point in time of $t_0 = \text{BJD } 2458776.11438$, and comprise 12969 data points, after clipping nine outliers, which are probably the result of cosmic ray strikes (including solar radiation). The TESS light curve of HD 1279 (HR 62) is shown in Figure 5.

The S17 data were analyzed using a fast Discrete Fourier Transform (Kurtz 1985) to produce an amplitude spectrum (Figure 6). There is no evidence of any p -mode pulsations out to the Nyquist frequency near 360 day^{-1} (not shown). In the low frequency range where g -modes occur there is no peak that

cannot be ascribed to instrumental variations. Hence, there is also no evidence of g -modes. A comparison with the instability strips of Zdravkov & Pamyatnykh (2008) shows that HD 1279 is too cool to be a β Cep star and too luminous to be an SPB star.

5.3. Frequency Analysis of TESS Data: HD 123445 (HR 5294)

HD 123445 (TIC 242288894; HR 5294) was observed by TESS in Sector 11 in 30 minutes cadence.²² The Fourier transform of the light curve from sector 38 (2 minutes cadence) reveals that HD 123445 (HR 5294) has a rotation period of roughly 3 days. It also pulsates with a frequency of about $f_1 = 3.8 \text{ c/d}$, and the harmonics $2f_1$ and f_1 show up clearly (Figure 7). The frequencies are collected in Table 4.

5.4. Frequency Analysis: HD 147550 (HR 6096), HD 171961 (HR 6990) and HD 202671 (HR 8137)

HD 147550 (TIC 71983021; HR 6096), HD 171961 (TIC 1286643; HR 6990), and HD 202671 (TIC 39364873; HR 8137) were not observed by TESS

6. Model Atmosphere Analysis and Synthetic Spectra Computation

For the six stars rotating with a $v_e \sin i$ lower than 30 km s^{-1} , the abundances or their upper limits for 40 chemical elements

²¹ Available on MAST via: doi: [10.17909/8rkx-c322](https://doi.org/10.17909/8rkx-c322).

²² Available on MAST via: doi: [10.17909/0shy-2f28](https://doi.org/10.17909/0shy-2f28).

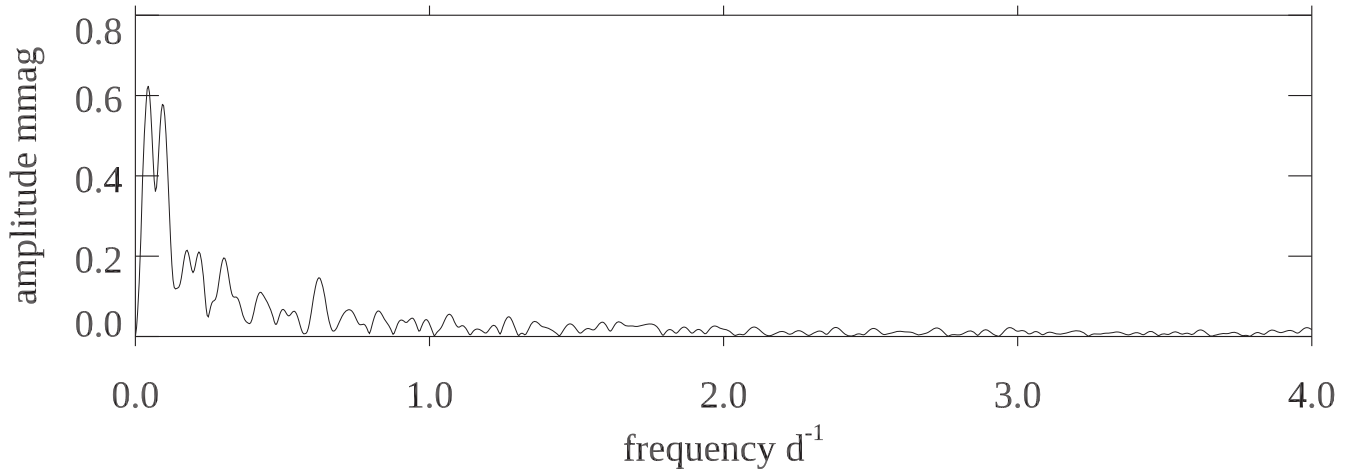


Figure 3. An amplitude spectrum of the Sector 10 data for HD 99803 (HR 4423) after the eclipse was excised from the time string. There is no evidence of any pulsation out to the Nyquist frequency near 360 day^{-1} . This plot shows on the low frequency range where there is a possibility of g -mode frequencies in a B9V star. No peaks that cannot be ascribed to instrumental variations are seen.

could be derived by iteratively adjusting synthetic spectra to the normalized spectra and looking for the best fit to carefully selected unblended lines. For the two stars rotating faster than 30 km s^{-1} , HD 123445 (HR 5294) and HD 171961 (HR 6990), the lines of elements with intrinsically low abundances (e.g., the lanthanides and heavy elements) are blended and fewer abundances could be derived. Specifically, synthetic spectra were computed assuming LTE using Hubeny & Lanz (1992) SYNSPEC49 code, which computes line profiles for elements up to $Z = 99$.

6.0.1. Microturbulent Velocity Determinations

To derive the microturbulent velocities, we have derived the iron abundance $[\text{Fe}/\text{H}]$ from unblended Fe II lines for a set of microturbulent velocities ranging from 0.0 to 2.0 km s^{-1} for each star. Figure 8 shows the standard deviation of $[\text{Fe}/\text{H}]$ as a function of the microturbulent velocity. The adopted microturbulent velocities are the values that minimize the standard deviations, i.e., for that value, all Fe II lines yield the same iron abundance. The microturbulent velocities of the eight stars, which are accurate to within $\pm 0.1 \text{ km s}^{-1}$, are collected in Table 3. Only two stars, HD 147550 (HR 6096) and HD 202671 (HR 8137), have microturbulent velocities that differ from 0 km s^{-1} .

6.1. Model Atmospheres

The ATLAS9 code (Kurucz 1992) was used to compute a first model atmosphere for the effective temperature and surface gravity of each star assuming a plane parallel geometry, a gas in hydrostatic and radiative equilibrium, and local thermodynamical equilibrium. The ATLAS9 model atmosphere contains 72 layers with a regular increase in $\log \tau_{\text{Ross}} = 0.125$ and was calculated assuming a solar chemical composition (Grevesse & Sauval 1998). It was converged up to $\log \tau = -5.00$ to attempt reproduce the cores of the Balmer lines. This ATLAS9 version uses the new opacity distribution function (ODF) of Castelli & Kurucz (2003) computed for that solar chemical composition. Once a first set of elemental abundances was derived using the ATLAS9 model atmosphere,

the atmospheric structure was recomputed for these abundances using the Opacity Sampling ATLAS12 code (Kurucz 2005, 2013). New slightly different abundances were then derived and a new ATLAS12 model was recomputed until the abundances of iteration $(n-1)$ differed of those of iteration (n) by less than $\pm 0.10 \text{ dex}$.

6.2. The Linelist

Atomic linelists from Kurucz’s database have provided the basis to construct our linelist.²³ These lists collect data mostly from the literature for light and heavy elements (usually critically evaluated transition probabilities from Martin et al. 1988 and Fuhr et al. 1988) and computed data by Kurucz (1992) for the iron group elements. A first linelist was built from Kurucz’s file gfall21oct16.dat²⁴ which includes hyperfine splitting levels. We then updated several oscillator strengths and damping parameters with more recent and accurate determinations when necessary. As a rule, we have preferred NIST²⁵ and Wiese et al. (1996) oscillator strengths for CNO and also Nilsson et al. (2006). The H I lines are calculated using Vidal et al. (1973) tables upgraded by Schoening & Butler (1989) up to H 10. The He I lines are computed in SYNSPEC49 by using specific tables, either from Shamey (1969) or Dimitrijevic & Sahal-Brechot (1984).

We introduced hyperfine components for a few lines of Mn II (Holt et al. 1999), isotopic and hyperfine transitions for the Ga II line at $6334.07 \text{ \AA}$ (Nielsen et al. 2000). To model the Hg II line at $3983.93 \text{ \AA}$, we have included nine hyperfine transitions from all isotopes from Hg^{196} up to Hg^{204} from Dolk et al. (2003). For Pt II, we have used the atomic data provided in Engleman (1989), Bohlender et al. (1998), and Jomaron et al. (1998).

For the Rare Earths, we retrieved all relevant transitions from the DREAM database.²⁶ We also used specific publications reporting on laboratory measurements: Sm II (Lawler et al. 2006), Nd II (Den Hartog et al. 2003), Eu II (Lawler et al. 2001b), and Tb II (Lawler et al. 2001a). In Kurucz’s linelists,

²³ <http://kurucz.harvard.edu/linelists/>

²⁴ <http://kurucz.harvard.edu/linelists/gfnew/gfall21oct16.dat>

²⁵ <http://physics.nist.gov/cgi-bin/AtData/linesform>

²⁶ <http://www.umh.ac.be/astro/dream.shtml>

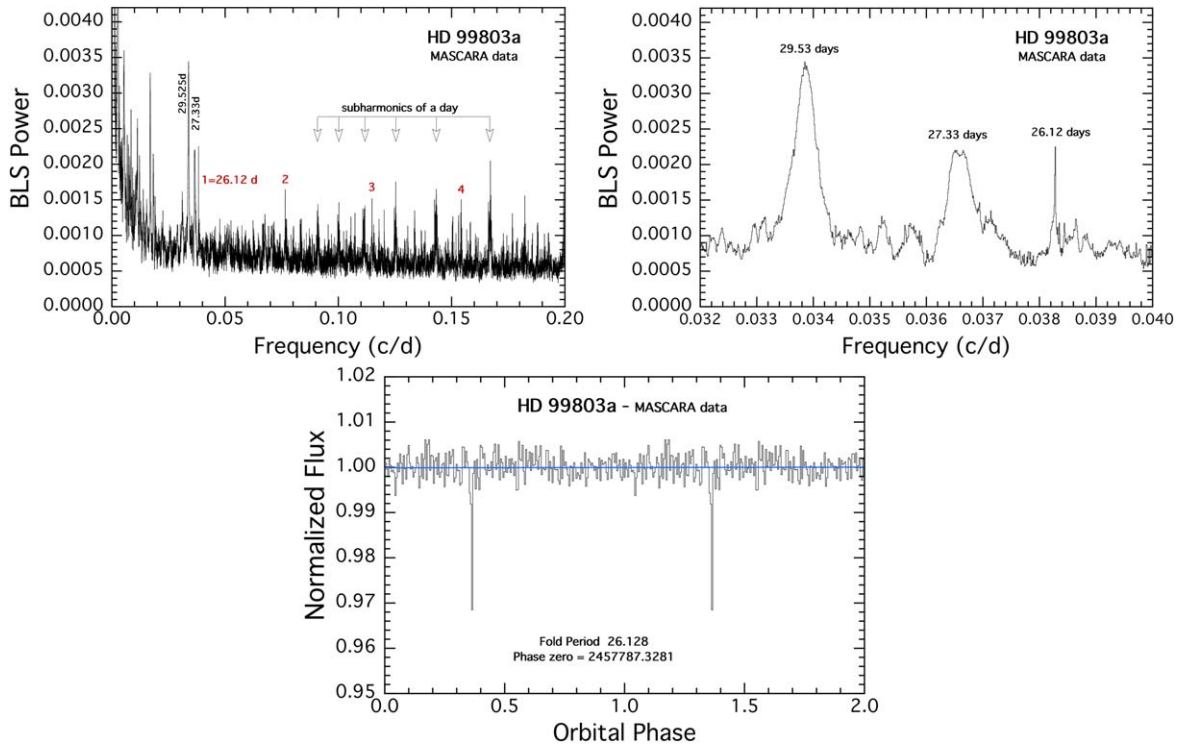


Figure 4. Detection of the binary period of HD 99803A with MASCARA data. Upper left-hand panel: BLS transform of 5.4 yr of MASCARA photometric data on HD 99803A. The 26.12 days detected period and its next three higher harmonics are marked in red font. The synodic and sidereal periods of the Moon are indicated, as are some of the subharmonics of the daily observation cycle. Upper right-hand panel: zoom-in around the two lunar periods and the eclipse period. Note that the former are extremely broad, while the eclipse period is appropriately narrow. Bottom panel: folded, binned, and averaged orbital light curve about a period of 26.128 days. The eclipse is narrow ($\lesssim 1\%$ of the orbital period) and $\sim 4\%$ deep, in agreement with the eclipse seen in the Sector 10 TESS data.

the damping constants are taken from the literature when available. For the iron group elements, they come from Kurucz’s (1992) computations. Additional damping constants for a few Si II lines were taken from Lanz et al. (1988). When they are not available from the linelist, damping constants were computed in SYNSPEC49 using an approximate formula (Kurucz & Avrett 1981).

6.3. Spectrum Synthesis

We have used only unblended lines to derive the abundances. For the two fastest rotators, HD 123445 (HR 5294) and HD 171961 (HR 6990), fewer abundances could be derived. For a given element, the final abundance is a weighted mean of the abundances derived for each transition. The weights are derived from the NIST grade assigned to that particular transition when available. The same weight was assigned to all transitions when no critical evaluation from NIST was available.

For several elements, in particular the heaviest elements, only one unblended line was available and the abundance should be regarded as an estimate only. For each modeled transition, the adopted abundance is that which provides the best fit to the observed normalized profile. A grid of synthetic spectra was computed with SYNSPEC49 (Hubeny & Lanz 1992) to model each selected unblended line of the 40 elements. Computations were iterated while varying the unknown abundance until minimization of the chi-square between the observed and synthetic spectrum was achieved.

7. Abundance Determinations

We discuss here the identifications of chemical elements, the lines selected for abundance determinations and the abundance determinations for each element. We also provide non-LTE abundance corrections for three light elements, namely Na I, Mg I, and Ca II. We use the Formato code (Merle et al. 2011) to build simple model atoms, using atomic data from NIST (Kramida et al. 2018) for energy levels, from VALD (Ryabchikova et al. 2015) for radiative bound-bound transitions, and from TOPbase (Cunto & Mendoza 1992) for radiative bound-free transitions. For inelastic collisions with electrons, we use the semi-empirical formula from Seaton (1962) for which a radiative counterpart exist and a collision strength of one for transitions with a forbidden radiative counterpart. For inelastic collisions with hydrogen, we use the formula from Drawin (1969) without scaling factor, except for Mg I where we implemented the mechanical quantum data from Guitou et al. (2015). The statistical equilibrium and the radiative transfer equation for each level and line in each model atom are consistently solved using the non-LTE 1D radiative transfer code MULTI (Carlsson 1986, 1992).

7.1. Helium

The strongest helium lines that are unambiguously detected are: λ 4026.19 Å, λ 4471.48 Å, and λ 5875.621 Å. The others are either weak or embedded into blends. The least blended, λ 4471.48 Å, has been synthesized to derive the helium LTE abundance. Helium is underabundant in all stars, the deficiency

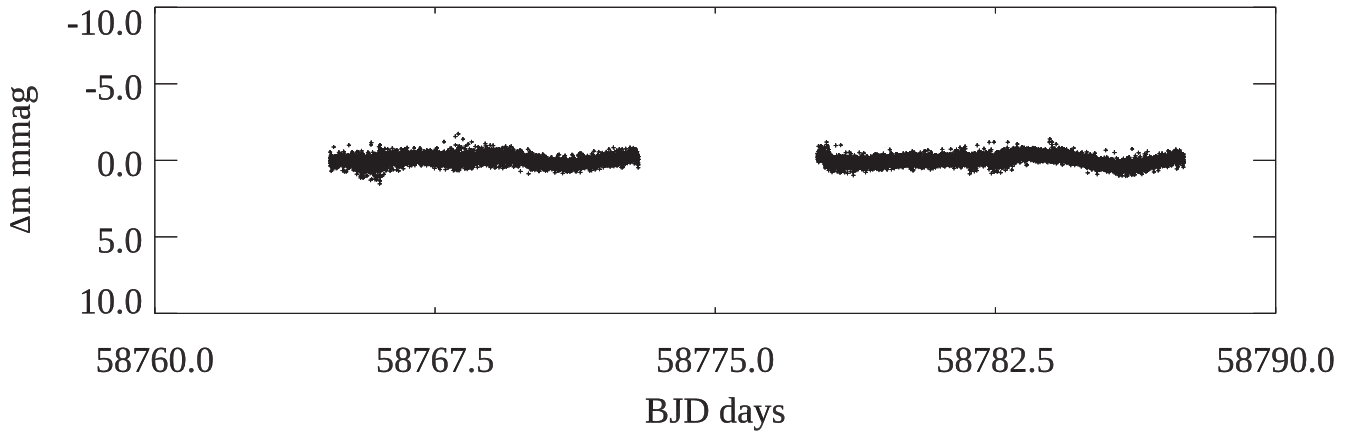


Figure 5. The light curve of HD 1279 (HR 62) obtained in 120-s cadence in Sector 17. The ordinate scale is Barycentric Julian Date—240 0000.0.

is the most important in HD 171961 (HR 6990), for which the abundance is only 4% of the solar abundance (-1.07).

7.2. The Light Elements (C to Ti)

7.2.1. Carbon

Most of the expected C II lines are blended with more abundant species. The least blended is the high excitation C II triplet at 4267.26 Å. Carbon is underabundant in HD 1279 (HR 62), HD 99803 (HR 4423), HD 123445 (HR 5294), HD 147550 (HR 6096), HD 202671 (HR 8137), β Scl (HR 8937), and μ Lep (HR 1702) with similar deficiencies, from 0.30 up to 0.60 times the solar abundance (-1.07).

Carbon is very underabundant in HD 171961 (HR 6990). Sakhibullin (1987) has shown that the C II line at 4267.26 Å is prone to NLTE effects for B stars hotter than 15,000 K. He suggests to use the C II lines at 3919.97 and 3921.69 Å to derive carbon abundances because they are less influenced by departures from LTE. These lines yield the same carbon abundances as 4267.26 Å for the eight stars.

7.2.2. Oxygen

For O I, the strongest expected allowed lines are those of multiplet 7 (NIST quality B+). The nine transitions of multiplet 7 dominate the opacity from 6155 Å to 6159 Å. Oxygen is underabundant in HD 99803 (HR 4423), HD 123445 (HR 5294), HD 147550 (HR 6096), HD 171961 (HR 6990), HD 202671 (HR 8137), and β Scl (HR 8937) with depletions of about 0.50 to 0.80 times the solar abundance (-3.31). In μ Lep (HR 1702), oxygen is about solar and marginally overabundant in HD 1279 (HR 62).

7.2.3. Neon

The Ne I line at 6402.25 Å (NIST quality AAA) was used to derive neon abundances. Neon is overabundant in HD 202671 (HR 8137), and quite underabundant in β Scl (HR 8937) and μ Lep (HR 1702). In all other stars, λ 6402.25 Å is not observed, which indicates an upper limit of one solar abundance (-4.07).

7.2.4. Sodium

We have used the 2 Na I lines at 4497.66 Å (NIST quality B+) and the resonance Na I lines at 5889.95 and 5895.92 Å (quality AA). The LTE overabundances of sodium derived from these lines are -5.67 (i.e., solar) for HD 1279 (HR 62)

and HD 171961 (HR 6990), half the solar value for HD 99803 (HR 4423), and overabundant by a factor of four in β Scl (HR 8937). No conclusion could be reached for HD 123445 (HR 5294), because of telluric lines, and for HD 202671 (HR 8137) and μ Lep (HR 1702), because of blending.

The NLTE corrections for the resonance Na I lines at 5889.95 and 5895.92 Å are about -0.52 and -0.60 dex, respectively, which yields Non-LTE abundances of Na of -6.27 in HD 1279 (HR 62) and HD 171961 (HR 6990), -5.97 for HD 99803A (HR 4423), and -5.37 for β Scl (HR 8937), which are well below the LTE determinations.

7.2.5. Magnesium

The unblended Mg II lines at 4390.514 Å and 4390.572 Å yield LTE abundances of -4.94 (HD 1279, HR 62), -5.0 (HD 99803A, HR 4423), -4.72 (HD 123445 (HR 5294), HD 171961 (HR 6990), β Scl (HR 8937)), -4.57 (HD 147550, HR 6096), -4.77 (HD 202671, HR 8137), and -4.54 (μ Lep, HR 1702), respectively, the solar abundance is -4.42 . The unblended Mg I line at 5172.68 Å yields similar abundances after a NLTE correction, which amounts to -0.33 dex. Magnesium therefore appears to be underabundant in each star.

7.2.6. Aluminum

The synthesis of the unblended Al II line at 4663.056 Å shows that aluminum is severely depleted in all stars except in HD 147550 (HR 6096), where it is overabundant (-5.23). The underabundance is the most pronounced in β Scl (HR 8937, -6.83). The solar abundance is -5.53 .

7.2.7. Silicon

The synthesis of strong and unblended lines of Si II reveals a large spread of the silicon abundances: from very deficient (-5.75) in HD 1279 (HR 62) to overabundant in HD 147550 (HR 6096, -4.25), β Scl (HR 8937, -4.09), and μ Lep (HR 1702, -3.91). The solar silicon abundance is -4.45 .

7.2.8. Phosphorus

The lines of P II expected strongest are λ 6024.18 Å and 6043.12 Å of multiplet 2. Phosphorus is overabundant in HD 1279 (HR 62, -5.07), HD 147550 (HR 6096, -5.67), HD 171961 (HR 6990, -5.00), HD 202671 (HR 8137, -4.32), μ Lep (HR 1702, -5.27), and β Scl (HR 8937, -5.14). None of

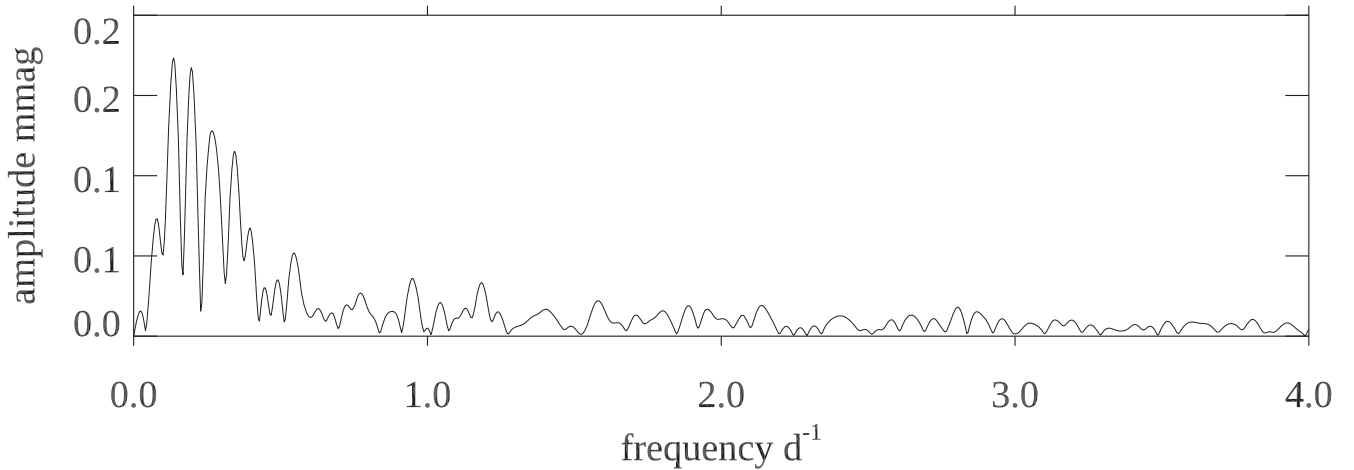


Figure 6. An amplitude spectrum of the Sector 17 data for HD 1279 (HR 62). There is no evidence of any pulsation out to the Nyquist frequency near 360 day^{-1} . This plot shows on the low frequency range where there is a possibility of g -mode frequencies in a B8III star. No peaks that cannot be ascribed to instrumental variations are seen.

the lines mentioned above are detected in HD 99803A (HR 4423) and HD 123445 (HR 5294), which suggests a one solar upper limit. The solar phosphorus abundance is -6.59 .

7.2.9. Sulfur

The lines of S II at $4162.67 \text{ \AA}$, $4142.25 \text{ \AA}$ and $4153.06 \text{ \AA}$ of multiplet 4 are clearly present and unblended in the spectra of all of these stars. Sulfur is solar in HD 123445 (HR 5294), HD 147550 (HR 6096), and μ Lep (HR 1702, -4.88) and underabundant in all other stars. The solar sulfur abundance is -4.67 .

7.2.10. Argon

The high excitation line of Ar II at $5194.53 \text{ \AA}$ was used to derive the abundance of argon. Argon is solar in HD 1279 (HR 62, -5.60). The Ar II line is blended in HD 202671 (HR 8137). The line of Ar II is not detected in HD 99803 (HR 4423), HD 123445 (HR 5294), HD 147550 (HR 6096), HD 171961 (HR 6990), μ Lep (HR 1702), and β Scl (HR 8937), which suggests a one solar upper limit. The solar argon abundance is -5.82 .

7.2.11. Calcium

The synthesis of the unblended Ca II lines at $3933.66 \text{ \AA}$ and $5019.97 \text{ \AA}$ yields the calcium LTE abundances, respectively -5.64 (HD 1279, HR 62), -6.34 (HD 99803A, HR 4423), -5.86 (HD 123445, HR 5294), -5.86 (HD 147550, HR 6096), -6.64 (HD 171961, HR 6990), -5.34 (HD 202671, HR 8137), -5.21 (μ Lep, HR 1702), and -5.24 (β Scl, HR 8937). The solar calcium abundance is -5.64 . The non-LTE correction for the $3933.66 \text{ \AA}$ amounts to $+0.37$ dex, Calcium therefore appears to be overabundant in HD 1279 (HR 62), HD 202671 (HR 8137), μ Lep (HR 1702), and β Scl (HR 8937).

7.2.12. Scandium

The line of multiplet 1 of Sc II at $4246.820 \text{ \AA}$ is the only unblended line (quality A' in NIST). Scandium appears to be overabundant in HD 1239 (HR 62, -8.13), HD 99803A (HR 4423, -8.35), and μ Lep (HR 1702, -8.13), and underabundant in HD 147550 (HR 6096, -9.13) and β Scl (HR 8937,

-9.83). The line is not detected in HD 123445 (HR 5294), HD 171961 (HR 6990), and HD 202671 (HR 8137), for which we adopt a one solar upper limit (the solar scandium abundance is -8.85).

7.2.13. Titanium

Ten lines of Ti II were used to derive the mean titanium abundances. Titanium is found to be overabundant in all stars, from marginally overabundant in HD 147550 (HR 6096, -6.82) to a sizable overabundance (-6.10) in β Scl (HR 8937). The solar titanium abundance is -6.98 .

7.3. The Iron-peak Elements (V to Zn) and Gallium

7.3.1. Vanadium

The vanadium abundance has been estimated using three lines of V II at $4005.705 \text{ \AA}$, $4023.378 \text{ \AA}$, and $4036.78 \text{ \AA}$.

The analysis of these lines yields an overabundance in HD 147550 (HR 6096); a solar abundance in HD 1239 (HR 62), HD 123445 (HR 5294), HD 171961 (HR 6990), and β Scl (HR 8937, -8.07); and a one solar upper limit for HD 202671 (HR 8137) and μ Lep (HR 1702). The solar vanadium abundance is -8.00 . The selected lines of V II are not observed in these two stars.

7.3.2. Chromium

Several strong and unblended lines of Cr II could easily be found in the spectra of all these stars. Six unblended lines of Cr II were synthesized. Chromium is solar in HD 171961 (HR 6990, -6.36) and overabundant in all stars from mildly overabundant in HD 1239 (HR 62), HD 123445 (HR 5294), HD 147550 (HR 6096), and HD 202671 (HR 8137) to overabundances close to 5 times the solar chromium abundance in HD 99803A (HR 4423), μ Lep (HR 1702), and β Scl (HR 8937). The solar chromium abundance is -6.33 .

7.3.3. Manganese

The two Mn II lines have hyperfine structures published by Holt et al. (1999). They yield overabundances in all stars but for HD 171961 (HR 6990), which has a solar manganese abundance (-6.57). The overabundances are very pronounced

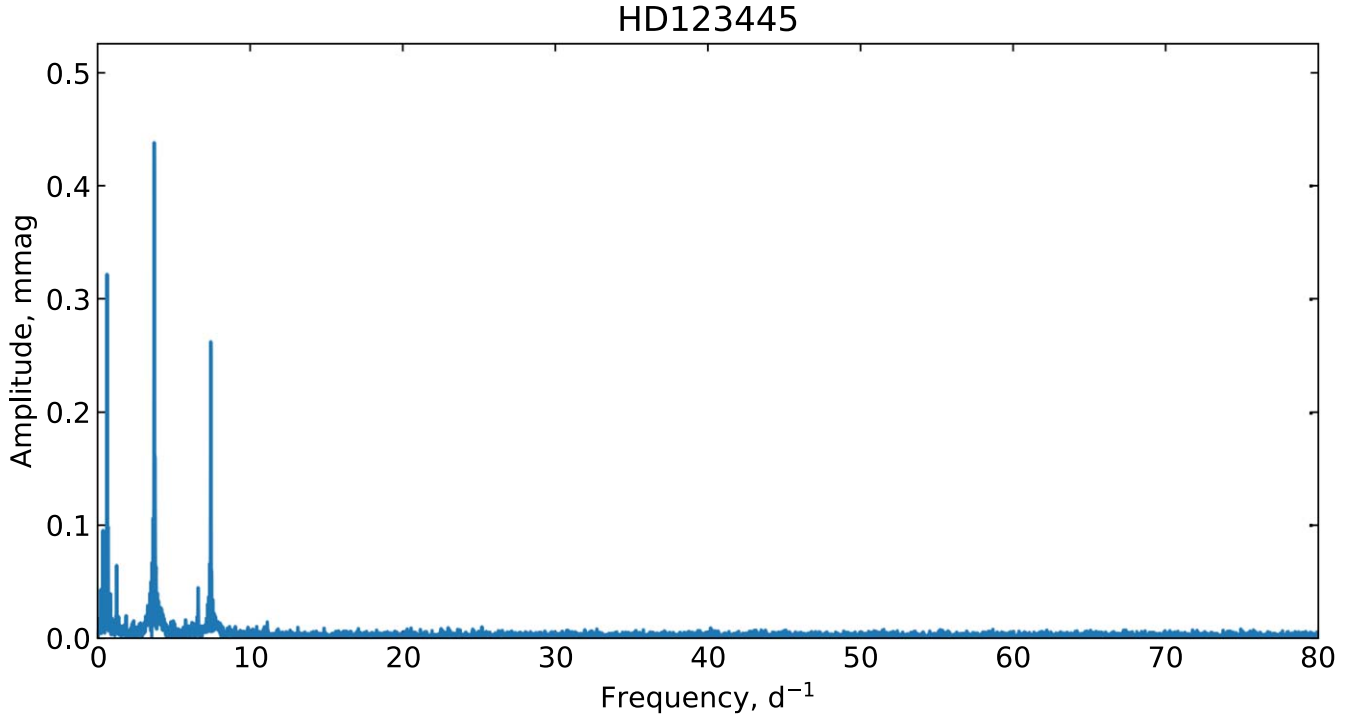


Figure 7. An amplitude spectrum of the Sector 17 data for HD 123445 (HR 5294).

Table 4
The Frequencies of HD 123445 (HR 5294)

Designation	Frequency day ⁻¹	Amplitude (mag)	Comment
F5	0.35658876	0.000083	
F4	0.62403082	0.000319	
F6	1.25185483	0.000059	
F1	3.71004238	0.000448	
F7	3.75556579	0.000149	
F8	6.59310162	0.000043	
F2	7.42008475	0.000260	2f1
F3	11.1301271	0.000013	3f1

in HD 1279 (HR 62, -5.21), HD 202671 (HR 8137, -5.10), μ Lep (HR 1702, -4.13), and β Scl (HR 8937, -4.21), and mild in HD 99803A (HR 4423, -6.00), HD 123445 (HR 5294, -6.21), and HD 147550 (HR 6096, -6.21). The solar manganese abundance is -6.61 .

7.3.4. Iron

The abundance of iron was derived from 11 Fe II lines, all of which are sensitive to small changes in $[\frac{Fe}{H}]$. Iron is mildly overabundant in HD 202671 (HR 8137, -4.10), about solar in HD 1279 (HR 62), HD 147550 (HR 6096), HR 6990, μ Lep (HR 1702), and β Scl (HR 8937, -4.50), and very underabundant in HD 99803A (HR 4423, -4.90) and HD 147550 (HR 6096, -5.02). The solar iron abundance is -4.50 .

7.3.5. Nickel

The nickel abundances was derived using one line at $4067.031 \text{ \AA}$ only and should be taken with caution. Nickel is found to be underabundant in HD 1279 (HR 62), HD 99803A (HR 4423), HD 171961 (HR 6990), HD 202671 (HR 8137),

and β Scl (HR 8937, -6.40 for these stars); very underabundant in μ Lep (HR 1702, -7.16) and HD 123445 (HR 5294, -7.00); and close to solar in HD 147550 (HR 6096, -5.78). The solar nickel abundance is -5.75 .

7.3.6. Gallium

We have used the high excitation Ga II line at $\lambda 6334.069 \text{ \AA}$ with hyperfine splitting of the gallium isotopes. It is the strongest expected line of Ga II in this temperature regime. The eight transitions retrieved and modified from Nielsen et al. (2000) are collected in Table 5.

The line of Ga II at $6334.069 \text{ \AA}$ is usually clearly present. Other lines of Ga II are present at $4251.108 \text{ \AA}$, $4254.032 \text{ \AA}$, $4255.64 \text{ \AA}$, and $4261.995 \text{ \AA}$; and $5360.313 \text{ \AA}$, $5363.353 \text{ \AA}$, and $5421.122 \text{ \AA}$ in the red part of the spectrum. Gallium is overabundant in HD 1279 (HR 62, -7.18), HD 123445 (HR 5294, -5.96), HD 171961 (HR 6990, -6.86), μ Lep (HR 1702, -5.38), and β Scl (HR 8937, -5.96), and solar in HD 99803A (HR 4423) and HD 147550 (HR 6096, -8.96), which is consistent with an upper limit of one solar in HD 202671 (HR 8137, the solar gallium abundance is -8.96 , no line observed).

7.4. The Sr-Y-Zr Triad

7.4.1. Strontium

The abundance of strontium was derived from the fit to the only unblended line of Sr II at $4215.52 \text{ \AA}$. Strontium is overabundant in HD 1279 (HR 62, -7.35), HD 147550 (HR 6096, -8.13), μ Lep (HR 1702, -6.43), and β Scl (HR 8937, -5.95); and solar in HD 99803A (HR 4423, -9.13), HD 123445 (HR 5294), HD 202671 (HR 8137), and HD 171961 (HR 6990). The solar strontium abundance is -9.03 .

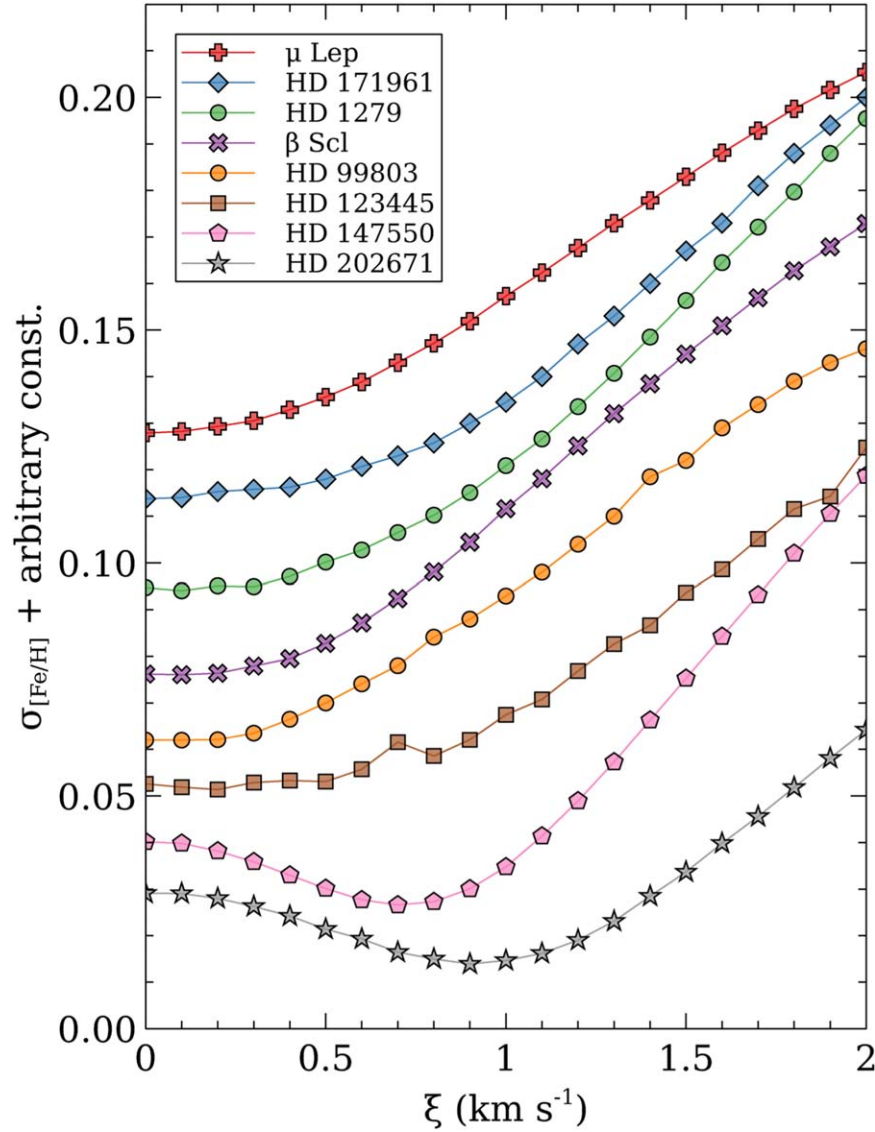


Figure 8. Microturbulent velocity determinations for the eight stars.

Table 5
The Ga II λ 6334.069 Å Linelist

Wavelength (Å)	ion	log gf	Lower Energy Level (cm ⁻¹)	EW (mÅ)
6333.930	Ga II	0.08	102944.595	14.7
6333.980	Ga II	0.21	102944.595	20.5
6333.990	Ga II	0.08	102944.595	14.6
6334.069	Ga II	0.10	102944.595	15.3
6334.080	Ga II	0.40	102944.595	26.5
6334.083	Ti II	-2.08	66521.008	0.0
6334.120	Ga II	0.02	102944.595	12.8
6334.200	Ga II	0.09	102944.595	20.9
6334.208	Cr II	-1.66	89812.422	0.0
6334.290	Ga II	0.01	102944.595	16.8
6334.368	Ni II	-2.36	120271.975	0.0

7.4.2. Yttrium

The yttrium lines are conspicuous in the spectra of HD 1279 (HR 62), HD 99803A (HR 4423), HD 123445 (HR 5294), HD 147550 (HR 6096), HD 202671 (HR 8137), μ Lep (HR 1702), β Scl (HR 8937), and HD 147550 (HR 6096).

These lines absorb from a few percent up to 10%. The only unblended lines are the low excitation lines at 3982.592 Å and 5662.925 Å. Yttrium is very overabundant in all stars from a factor of 5 times (-9.09) up to 12500 times (-5.69) the solar abundance, except in HD 171961 (HR 6990) where it is solar (the solar yttrium abundance is -9.79).

7.4.3. Zirconium

The only two Zr II unblended lines are λ 4443.008 Å and λ 4457.431 Å. Zirconium is solar in HD 1279 (HR 62), HD 123445 (HR 5294), and HD 171961 (HR 6990, the solar zirconium abundance is -9.42); and overabundant in HD 99803A (HR 4423, -8.12), HD 147550 (HR 6096, -8.64), HD 202671 (HR 8137, -7.42), μ Lep (HR 1702, -7.64), and β Scl (HR 8937, -7.32).

7.5. Barium

The two resonance lines of Ba II (Multiplet 1) at 4554.029 Å and 4934.076 Å and the low excitation line 6141.59 Å are present. The

Table 6
The Ba II Linelists

Wavelength (Å)	ion	log gf	Lower Energy Level (cm ⁻¹)	EW (mÅ)
4553.934	Zr II	-0.57	19514.840	0.1
4553.995	Ba II	-1.57	0.000	0.1
4553.997	Ba II	-1.57	0.000	0.1
4553.998	Ba II	-1.99	0.000	0.0
4553.999	Ba II	-1.82	0.000	0.1
4554.001	Ba II	-1.82	0.000	0.1
4554.001	Ba II	-2.22	0.000	0.0
4554.011	Cr I	-0.73	33040.093	0.0
4554.029	Ba II	0.02	0.000	16.6
4554.029	Ba II	-1.45	0.000	1.3
4554.029	Ba II	-0.94	0.000	3.8
4554.046	Ba II	-1.38	0.000	0.3
4554.049	Ba II	-1.82	0.000	0.1
4554.049	Ba II	-1.15	0.000	0.6
4554.050	Ba II	-2.52	0.000	0.0
4554.051	Ba II	-1.57	0.000	0.1
4554.052	Ba II	-2.29	0.000	0.0
4934.005	Fe I	-0.61	33507.120	0.4
4934.074	Ba II	-3.12	0.000	0.0
4934.074	Ba II	-1.33	0.000	0.8
4934.075	Ba II	-3.15	0.000	0.0
4934.075	Ba II	-1.10	0.000	1.3
4934.076	Ba II	-1.77	0.000	0.3
4934.076	Ba II	-1.25	0.000	0.9
4934.077	Ba II	-0.29	0.000	7.0
4934.084	Fe I	-2.30	26627.608	0.1
6141.597	Fe I	-3.12	33765.304	0.0
6141.713	Ba II	-0.22	5674.824	2.4
6141.713	Ba II	-1.03	5674.824	0.4
6141.714	Ba II	-1.18	5674.824	0.3
6141.714	Ba II	-1.26	5674.824	0.2
6141.715	Ba II	-1.69	5674.824	0.1
6141.717	Ba II	-3.07	5674.824	0.0
6141.718	Ba II	-3.05	5674.824	0.0
6141.731	Fe I	-1.61	29056.322	0.1

hyperfine structure of the various isotopes of barium that we have used is collected in Table 6. Barium is solar in HD 1279 (HR 62), HD 123445 (HR 5294), HD 202671 (HR 8137), β Scl (HR 8937), and HD 171961 (HR 6990, the solar barium abundance is -9.82); and overabundant in HD 99803A (HR 4423, -8.30), HD 147550 (HR 6096, -8.74), and μ Lep (HR 1702, -8.82).

7.6. The Rare Earths

We have searched for once ionized rare earths elements in the spectra of the eight stars, namely La II, Ce II, Pr II, Nd II, Sm II, Eu II, Gd II, Tb II, Dy II, Ho II, and Er II. Since the twice ionized rare earths often are the dominant stages in these atmospheres of late B-stars, we also looked for the twice ionized species, Pr III and Nd III using the lines listed in either NIST or in DREAM and the linelists published in Ryabchikova et al. (2006) and Ryabchikova et al. (2007). The lines of these ions are weak, a few of these lines could be found in most stars only be found in the slowest rotators. HD 147550 (HR 6096) is the star that has the most detections of rare earths.

7.6.1. Lanthanum

The spectrum synthesis of the La II line at 4042.91 Å feature is consistent with one solar abundance upper limit for HD 171961 (HR 6990) and β Scl (HR 8937). The line is detected in HD 1279 (HR 62), HD 99803A (HR 4423), and HD 202671 (HR 8137), and is properly reproduced with a solar abundance (the solar lanthanum abundance is -10.90). Lanthanum is overabundant in HD 147550 (HR 6096, -9.90) and very abundant in μ Lep (HR 1702, -8.00).

7.6.2. Cerium

The synthesis of the unblended Ce II lines at 4133.80 Å and 4460.21 Å yields a mild overabundance of cerium in HD 147550 (HR 6096, -9.24) and a one solar upper limit in all of the other stars (the solar cerium abundance is -10.42).

7.6.3. Praseodymium

The lines of Pr III at 5284.69 and 5299.99 Å are unblended and provide large overabundances in HD 147550 (HR 6096, -10.38), HD 99803A (HR 4423, -9.98), μ Lep (HR 1702, -8.98), and HD 123445 (HR 5294, -8.58). The synthesis of these lines is consistent with a solar abundance upper value for HD 1279 (HR 62), HD 202671 (HR 8137), HD 171961 (HR 6990), and β Scl (HR 8937, the solar praseodymium abundance is -11.28).

7.6.4. Neodymium

For Nd III, the strongest expected lines in DREAM are the resonance lines λ 5265.019 Å and λ 5294.099 Å. The synthesis of these lines yields moderate to large overabundances of Nd III in HD 99803A (HR 4423, -9.98), HD 1279 (HR 62, -9.28), HD 123445 (HR 5294, -8.18), and HD 147550 (HR 6096, -9.73). In HD 202671 (HR 8137), HD 171961 (HR 6990), μ Lep (HR 1702), and β Scl (HR 8937), the abundance of neodymium has a one solar upper value, i.e., -10.50 (the lines are not detected).

7.6.5. Samarium

The Sm II lines at 4280.79 Å and 4424.32 Å are unblended and are consistent with a one solar upper limit abundance of samarium (the solar samarium abundance is -11.04) in all stars except for HD 147550 (HR 6096) for which samarium is overabundant (-9.44).

7.6.6. Europium

The Eu II resonance line at 4129.73 Å is unblended and consistent with a one solar upper limit abundance of europium (the solar europium abundance being -11.48) in all stars except in HD 147550 (HR 6096), where europium is overabundant (-10.48).

7.6.7. Gadolinium

The Gd II line at 4037.32 Å is unblended and properly reproduced with large overabundances in HD 147550 (HR 6096, -9.63) and μ Lep (HR 1702, -7.45). In the other stars, gadolinium has a one solar upper limit (the solar gadolinium abundance is -10.88).

7.6.8. Terbium

The Tb II line at 4005.47 Å provides a one solar upper limit abundance (the solar terbium abundance is -11.70) for all stars except HD 147550 (HR 6096), where terbium is overabundant (-10.70).

7.6.9. Dysprosium

The Dy II line at 4000.45 Å provides a one solar upper limit (the solar dysprosium abundance is -10.90) in all stars except for HD 147550 (HR 6096), where dysprosium is overabundant (-9.30).

7.6.10. Holmium

The low excitation line of Ho II at 4152.62 Å provides a one solar upper limit (the solar holmium abundance is -11.52) in all stars except for HD 147550 (HR 6096), where holmium is overabundant (-9.04).

7.6.11. Erbium

For Er II, the line at 4142.91 Å is consistent with a one solar upper limit (the solar erbium abundance is -11.08) for all stars except HD 147550 (HR 6096), where erbium is very overabundant (-8.78).

7.6.12. Thulium

The line at 4242.15 Å of Tm II is consistent with a one solar upper limit (the solar thulium abundance is -11.90) in all eight stars.

7.6.13. Ytterbium

The Yb II line at 4135.095 Å provides a one solar upper limit for all eight stars (the solar ytterbium abundance is -11.16).

7.7. The Very Heavy Elements Hafnium, Osmium, Platinum, Gold, and Mercury

7.7.1. Hafnium

The Hf II line at 3918.08 Å yields a one solar upper limit (the solar hafnium abundance being -11.15) in all eight stars.

7.7.2. Osmium

The Os II line at 4158.44 Å is consistent with a one solar upper limit (the solar osmium abundance being -10.60) in all eight stars.

7.7.3. Platinum

The 4514.47 Å Pt II line is the only clearly detected line of platinum in the spectra of HD 99803A (HR 4423), μ Lep (HR 1702) and β Scl (HR 8937). In the other five stars, this line is not detected and the abundance of platinum has a one solar upper limit (the solar platinum abundance is -10.38). Our referee, Prof. M. Dworetsky, has kindly provided the hyperfine structure and isotopic shifts for the 4514.47 Å Pt II line (Engleman 1989; Bohlender et al. 1998; Jomaron et al. 1998). Inclusion of these hyperfine transitions significantly reduces the overabundances of platinum. The new abundances are -8.32 in HR 4423, -7.20 in HR 1702, and -6.20 in HR 8937. There are no other species at this wavelength, which

Table 7
The Hg II 3983.93 Å Linelist

Wavelength (Å)	ion	log gf	Lower Energy Level (cm $^{-1}$)	EW (mÅ)
3983.771	Hg II	-4.50	35514.999	1.2
3983.827	Ti I	-1.46	17075.258	0.0
3983.832	Fe I	-4.81	24338.766	0.0
3983.839	Hg II	-3.00	35514.000	14.1
3983.844	Hg II	-3.13	35514.000	12.5
3983.851	Cr II	-4.31	81962.291	0.0
3983.853	Hg II	-3.00	35514.000	14.1
3983.874	Fe I	-2.70	34039.515	0.0
3983.899	Cr I	0.35	20520.904	3.2
3983.912	Hg II	-2.50	35514.000	19.2
3983.932	Hg II	-3.10	35514.000	12.9
3983.941	Hg II	-2.90	35514.000	15.3
3983.956	Fe I	-1.02	21999.129	2.5
3983.986	Cr II	-2.32	90512.561	0.0
3983.991	Cr II	-2.88	82362.188	0.0
3983.993	Hg II	-3.00	35514.000	14.2
3984.022	Cr I	-2.47	35397.971	0.0
3984.072	Hg II	-3.00	35514.000	14.4

could account for the observed absorption. The wavelength scale around the Pt II line was checked using the Fe II neighboring control lines whose NIST wavelengths are accurate to within ± 0.001 Å.

7.7.4. Gold

The Au II line at 4052.790 Å is not clearly seen in the spectra of the eight stars, and is therefore reproduced by a one solar upper limit for all stars (-10.99).

7.7.5. Mercury

The abundance of mercury has been derived from the low excitation Hg II line at 3983.93 Å only, which is clearly present in all eight stars. This feature absorbs from 2% in HD 147550 (HR 6096) and HD 202671 (HR 8137) to about 12% of the continuum in β Scl (HR 8937). The other Hg II lines are all high excitation lines either weak or blended with more abundant species and were not synthesized. In particular, Hg II 6149.4749 Å is blended with Fe II 6149.258 Å in all stars.

To model the Hg II line at 3983.93 Å, we have included nine transitions, i.e., all hyperfine structure from the various isotopes from Hg¹⁹⁶ and Hg²⁰⁴ listed in Dolk et al. (2003). These transitions are collected together with blending lines from Ti I, Fe I, Cr I, and Cr II in Table 7, where the equivalent widths are computed for the final abundances of HR 62.

The wavelength scale was checked by using four control lines on each side of the Hg II line: shortwards the Zr II line at 3982.0250 Å, the Y II line at 3982.59 Å, and longwards the Zr II lines at 3984.718 Å and 3991.15 Å. Once corrected for the radial velocities of each star, the centers of these lines are found at their expected laboratory locations to within ± 0.02 Å, which we will adopt as the accuracy of the wavelength scale in this spectral region.

First, we checked the influence of possible contaminant species to the Hg II line from 3983.50 Å up to 3984.50 Å, which we estimate to be the maximum extension of the line wings. The only possible contaminants are three weak lines: Fe II λ 3983.737 Å, Cr I λ 3983.897 Å, and Fe I λ 3983.956 Å.

Table 8
The Final Abundances for the Eight Stars

Element	HD 1279 HR 62	HD 99803A HR 4423	HD 123445 HR 5294	HD 147550 HR 6096	HD 171961 HR 6990	HD 202671 HR 8137	μ Lep HR 1702	β Scl HR 8937	Sun
He (2)	10.33	10.63	10.23	10.53	9.53	9.93	9.93	9.93	10.93
C (6)	7.97	8.03	8.21	8.03	7.13	7.91	8.21	8.21	8.43
N (7)	7.83			7.83					7.83
O (8)	8.87	8.41	8.39	8.47	8.39	8.47	8.73	8.59	8.69
Ne (10)	7.93			7.93		8.41	8.23	7.23	7.93
Na (11)	6.24	5.94		6.24	6.24	6.24	6.94		6.24
Mg (12)	7.08	6.98	7.25	7.45	6.90	7.25	7.48	7.30	7.60
Al (13)	5.45	5.93	5.85	6.75	5.45	5.15	5.45	5.15	6.45
Si (14)	6.21	7.00	7.29	7.72	6.99	7.45	8.03	7.87	7.51
P (15)	6.89	5.41	5.41	6.29	6.95	7.71	6.82	6.68	5.41
S (16)	6.82	7.12	7.12	7.33	7.12	6.52	7.12	6.82	7.12
Cl (17)	6.50								5.50
Ar (18)	6.40					6.40			6.40
Ca (20)	7.04	5.64	6.12	6.22	5.34	6.64	6.74	6.74	6.34
Sc (21)	3.85	3.63		2.85			3.85	2.15	3.15
Ti (22)	5.65	5.36	5.46	5.11	5.53	5.63	5.83	5.75	4.95
V (23)	3.93	3.53	3.93	4.40				3.93	3.93
Cr (24)	5.88	6.27	6.04	5.78	5.64	5.94	6.36	6.29	5.64
Mn (25)	6.83	6.03	7.28	5.77	5.43	6.94	7.91	7.83	5.43
Fe (26)	7.68	7.06	6.98	7.55	7.50	7.90	7.73	7.73	7.50
Ni (28)	5.52	5.52	4.92	6.48	5.52	5.46	4.76	5.52	6.22
Ga (31)	4.82	3.04	6.04	3.04	5.14		6.62	6.04	3.04
Sr (38)	3.47	2.97	2.87	3.87	2.87	2.87	5.57	6.05	2.87
Y (39)	4.75	3.99	4.91	2.91	2.21	4.21	5.51	6.31	2.21
Zr (40)	2.58	3.88	2.58	3.36	2.58	2.58	4.36	4.68	2.58
Xe (54)	5.84			2.24		4.24	6.98	6.94	2.24
Ba (56)	2.18	3.69	2.18	3.26	2.18	2.18	3.18	2.18	2.18
La (57)		1.10		2.10		1.10	4.00		1.10
Ce (58)				2.76		1.58	1.58		1.58
Pr (59)	0.72	2.02	3.42	1.72			3.02		0.72
Nd (60)	2.72	2.02	3.82	2.27			1.42		1.42
Sm (62)				2.56			0.96	0.96	0.96
Eu (63)				1.52			0.52	0.52	0.52
Gd (64)				2.37			4.55		1.07
Tb (65)				1.30	0.30		0.30		0.30
Dy (66)		1.10		2.70	1.10		1.10		1.10
Ho (67)				2.96				0.48	0.48
Er (68)				3.22					0.92
Hf (72)							0.85		0.85
Os (76)					1.40				1.40
Pt (78)		1.90			0.02		2.72	4.02	0.02
Hg (80)	5.35	5.17	6.60	4.57	5.77	4.47	6.77	7.17	1.17
Bi (83)	4.95								0.65

The equivalent widths of these lines were computed for the derived Fe and Cr abundances, and the sum of their contributions is by far insufficient to reproduce the equivalent width of the observed feature at 3983.93 Å (about 64 to 70 mÅ) of the Hg II blend. Another test computed the flux without the Hg II transitions. This consistently resulted into a very weak absorption feature from 3983.50 to 3984.50 Å in each star. We can therefore conclude that the observed features at 3983.93 Å are indeed mostly due to the Hg II line and are almost free of blends.

Including the hyperfine structure of the seven isotopes from Table 2 of Dolk et al. (2003) significantly usually reduces the mercury abundance. The final combination of oscillator strengths is recorded in Table 7.

The final mercury abundances derived from the synthesis of the 3983.93 Å line yield overabundances of -7.43 for

HD 147550 (HR 6096) and HD 202671 (HR 8137); of about -7.43 for HD 1219 (HR 62) and HD 99803A (HR 4423); -6.23 in HD 171961 (HR 6990) and of -5.40 for HD 123445 (HR 5294); and of -5.23 for μ Lep (HR 1702) and -4.83 for β Scl (HR 8937). The solar mercury abundance is -10.91 .

The final abundances²⁷ are collected in Table 8. The error budget for these abundances is explained in the Appendix and uncertainties are collected in Table 9.

8. Discussion

The abundance analysis carried out here clearly establishes that the six slowly rotating southern stars that we analyzed are newly discovered chemically peculiar stars with large mercury

²⁷ We refer here to the absolute abundance in the star: $\log_{10}\left(\frac{X}{H}\right)_*$, where $\log_{10}(H) = 12.0$.

Table 9
Abundance Uncertainties for the Elements Analyzed in the Eight Stars

		He I	C II	O I	Mg II	Al II	Si II	P II	S II	Ca II
ΔT_{eff}	+200 K	-0.30	0.00	0.05	0.047	0.075	-0.017	0.00	0.00	0.109
$\Delta \log g$	0.15 dex	0.079	0.24	-0.06	-0.052	0.118	-0.017	-0.12	-0.12	-0.032
$\Delta \xi_r$	+0.20 km s ⁻¹	0.00	0.00	0.00	0.00	0.075	-0.035	-0.05	-0.046	0.00
$\Delta \log gf$	+0.10	-0.30	-0.12	-0.16	-0.25	-0.028	-0.053	-0.07	-0.071	-0.105
$\Delta \text{continuum}$		-0.12	0.21	0.075	0.13	0.25	0.062	0.06	0.061	0.058
$2\sigma_{[X/H]}$		0.45	0.34	0.19	0.29	0.20	0.092	0.160	0.159	0.165
		Sc II	Ti II	V II	Cr II	Mn II	Fe II	Ni II	Sr II	Y II
ΔT_{eff}	+200 K	0.22	0.041	0.028	0.057	0.00	0.079	0.176	0.138	0.146
$\Delta \log g$	0.15 dex	-0.079	0.015	-0.014	0.03	0.007	0.041	0.114	-0.058	
$\Delta \xi_r$	+0.20 km s ⁻¹	-0.038	-0.016	0.00	0.00	0.00	-0.095	0.00	-0.058	0.00
$\Delta \log gf$	+0.10	-0.13	-0.23	-0.046	-0.15	-0.015	-0.095	-0.155	-0.058	-0.05
$\Delta \text{continuum}$		0.067	0.03	0.014	0.046	0.014	0.0126	0.079	0.051	0.041
$2\sigma_{[X/H]}$		0.276	0.064	0.23	0.169	0.076	0.162	0.273	0.178	0.165
		Zr II	Ba II	La II	Ce II	Pr III	Nd III	Sm II	Eu II	Gd II
ΔT_{eff}	+200 K	0.058	0.114	0.11	0.11	0.11	0.11	0.11	0.11	0.11
$\Delta \log g$	0.15 dex	0.046	-0.032	-0.125	-0.13	-0.13	-0.13	-0.13	-0.13	-0.13
$\Delta \xi_r$	+0.20 km s ⁻¹	0.00	-0.097	-0.09	-0.09	-0.09	-0.09	-0.09	-0.09	-0.09
$\Delta \log gf$	+0.10	-0.066	-0.125	-0.13	-0.13	-0.13	-0.13	-0.13	-0.13	-0.13
$\Delta \text{continuum}$		0.146	-0.097	-0.09	-0.09	-0.09	-0.09	-0.09	-0.09	-0.09
$2\sigma_{[X/H]}$		0.173	0.251	0.25	0.25	0.25	0.25	0.25	0.25	0.25
		Dy II	Tb II	Ho II	Er II	Tm II	Yb II	Hf II	Os II	Pt II
ΔT_{eff}	+200 K	0.11	0.11	0.11	0.11	-0.155	0.11	0.11	0.11	0.11
$\Delta \log g$	0.15 dex	-0.13	-0.13	-0.13	-0.13	0.041	-0.13	-0.13	-0.13	-0.13
$\Delta \xi_r$	+0.20 km s ⁻¹	-0.09	-0.09	-0.09	-0.09	-0.097	-0.09	-0.09	-0.09	-0.09
$\Delta \log gf$	+0.10	-0.13	-0.13	-0.13	-0.13	-0.108	-0.13	-0.13	-0.13	-0.13
$\Delta \text{continuum}$		-0.09	-0.09	-0.09	-0.09	0.079	-0.09	-0.09	-0.09	-0.09
$2\sigma_{[X/H]}$		0.25	0.25	0.25	0.25	0.23	0.25	0.25	0.25	0.25
		Au II	Hg II							
ΔT_{eff}	+200 K	0.11	-0.155							
$\Delta \log g$	0.15 dex	-0.13	0.041							
$\Delta \xi_r$	+0.20 km s ⁻¹	-0.09	-0.097							
$\Delta \log gf$	+0.10	-0.13	-0.108							
$\Delta \text{continuum}$		-0.09	0.079							
$2\sigma_{[X/H]}$		0.25	0.23							

overabundances. The abundance patterns for these objects are shown in Figure 9. None of these objects appears to be a rapid rotator seen pole-on, such as Vega. We fail to find the characteristic flat-bottomed Fe II lines of Vega in the spectra that we analyzed. The inference of the age and mass of the stars with the SPInS tool and BaSTI stellar models shows that six stars are still on the main sequence, while for the two most massive ones, HD 1279 (HR 62) and HD 202671 (HR 8137) the evolutionary status cannot be firmly assessed. All these stars are young, as expected, and their estimated ages range from 100 to 261 Myr. The analysis of the TESS light curve of HD 99803A (HR 4423) shows that this star is an eclipsing binary star with an orbital period longer than the 24.96 days time span of the Sector 10 data. We estimate the radius of the transiting body to be about 0.2 the radius of the primary B9V star from the transit depth. This suggests that a secondary that is a lower main-sequence star is a plausible conjecture. There is no evidence for p -modes nor g -modes in the periodogram of

the TESS light curve. HD 99803 (HR 4423) is too cool to be a SPB star or a β Cephei star. There is no evidence either in this star for signatures of p -modes or g -modes. HD 123445 (HR 5294) is a pulsational variable with a principal period of 6 hours and 28 minutes, and two harmonics. HD 99803A (HR 4423) also stands out as a particularly interesting object because its spectrum has very sharp lines, similar to HD 72660 (HR 3383). There are very few very sharp lined early A/late B-stars. From the list of identifications provided for the final composition of HD 99803A (HR 4423) in Table 10 in the Appendix, we can conclude that the following species are present in the spectra of HD 99803A (and in the other stars as well): C II, N II, O I, Na I, Mg II, Mg I, Mg II, Si II, P II, S II, Ca II, Sc II, Ti II, V II, Cr II, Cr I, Mn II, Fe II, Ni II, Ga II, Sr II, Y II, Zr II, Ba II, Pt II, and Hg II. Among these species, Si II, Ti II, Cr II, Mn II, Fe II, Y II, and Zr II have numerous and strong lines, and are important opacity sources. Lines from neutrals: Cr I, Mn I, Fe I are also observed and yield lower abundances

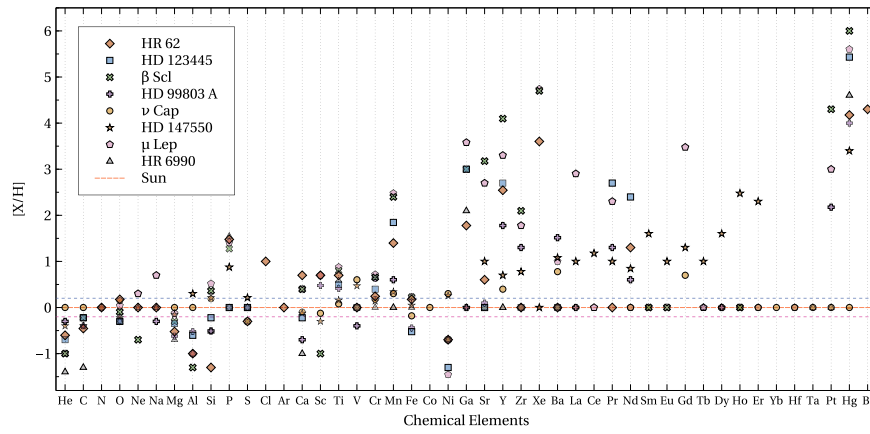


Figure 9. Comparison of the abundance patterns of the eight stars to that of the superficially normal star ν Cap (HR 7773).

Table 10
Line Identifications for HD 99803A (HR 4423)

λ_{obs} (Å)	λ_{lab} (Å)	Identification	$\log gf$	E_{low}	Relative Flux
3901.22	3900.515	Fe I	-0.92	26140.176	0.41
	3900.550	Ti II	-0.45	9118.260	
3903.61	3902.945	Fe I	-0.47	12560.933	0.87
3904.42	3903.73	Fe II	-1.49	60807.228	0.90
3906.31	3905.644	Cr II	-0.90	42986.619	0.56
3906.71	3906.035	Fe II	-1.83	44929.549	0.74
3911.99	3911.30	Cr II	-2.06	39824.381	0.82
3914.14	3913.45	Ti II	-0.53	8997.710	0.43
3914.83	3914.14	Ti II	-2.24	15257.430	0.92
3915.18	3914.49	Fe II	-4.05	13474.491	0.85
3916.61	3915.92	Zr II	-0.82	4248.300	0.90
3918.01	3917.32	Mn II	-1.15	55759.270	0.93
3919.20	3918.51	Fe II	-2.10	47674.718	0.91
3921.31	3920.62	Fe II	-1.20	60625.451	0.84
		C II	-0.21	131735.525	
...	...	...	...	...	...

(This table is available in its entirety in machine-readable form.)

than Cr II, Mn II and Fe II, indicating possibly departures from LTE.

We find evidence for overabundances of four rare earths from unblended lines of Ce II, Pr III, and Nd III and Gd II in the slowest rotating stars. For the other rare earths, we only find upper limits: the abundances must be solar or lower. The presence of once ionized rare earths is difficult to assess at temperatures above 10,000 K because they are not the dominant ionization stage. In general, the abundances for the rare earths should be regarded as the least reliable ones in this study because they were usually inferred from the synthesis of one weak line of the once ionized species only.

The overabundances found in the six newly discovered CP stars run from very mild ($2.0 \odot$ for Fe) up to quite large ($10^5 \odot$ for Hg). The underabundances run from mild $-0.80 \odot$ for carbon to pronounced $-0.02 \odot$ for nickel. Helium is quite underabundant in HD 171961 (HR 6990), for which the abundance is only 4% of the solar abundance. The comparison late B star, ν Cap (HR 7773, HD 193432), has a nearly solar composition for all elements, as already found by Monier et al. (2018) and Adelman (1991). For the other eight stars, the overall trend is that the light elements (He through Mg) are

most often deficient, whereas elements heavier than manganese tend to be increasingly overabundant as the atomic mass increases (with the exception of nickel). This pattern of abundances is most likely caused by radiative diffusion.

9. Conclusion

With their characteristic underabundances of most light elements up to calcium, and overabundances of manganese and elements heavier than strontium, in particular mercury, the six stars studied here are definitely new HgMn stars or new PGa stars. Their abundances differ from those of the classical HgMn stars μ Lep (HR 1702) and β Scl (HR 8937), possibly because their initial abundances differed slightly and because radiative diffusion operated in a slightly different manner in stars having different masses. For HD 147550 (HR 6096), the mild deficiencies for light elements and overabundances for strontium, yttrium, and zirconium, and several rare earths and mercury suggest that this object should be reclassified as a mild Chemically Peculiar star. We find that HD 123445 (HR 5294) pulsates with a main period of 6 hours 28 minutes, and two harmonics. More observations of this interesting star will help to elucidate its nature. A photometric monitoring of HD 99803A (HR 4423) over a period of time longer than 25 days span of the TESS observations should help measure the orbital period of this new eclipsing spectroscopic binary.

Acknowledgments

We thank Prof. Mike Dworetzky, our referee, for providing the hyperfine structure and isotopic shifts for the Pt II line at 4514 Å. We also thank Prof. Charles Cowley for his insightful comments during the analysis of HD 99803A. We are grateful to Tamas Borkovits, Tibor Mitnyan, and Veselin Kostov for their help with some of the TESS data. This work has made use of the VALD database, operated at Uppsala University, the Institute of Astronomy RAS in Moscow, and the University of Vienna. We have also used the NIST Atomic Spectra Database (version 5.4) available <http://physics.nist.gov/asd>. This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France. D.M.B. gratefully acknowledges funding from the Research Foundation Flanders (FWO) by means of a senior postdoctoral fellowship (grant agreement No. 1286521N). This work was supported by FCT/MCTES through the research grants UIDB/04434/2020, UIDP/04434/2020 and PTDC/FIS-AST/30389/2017. M.D. is

supported by national funds through FCT in the form of a work contract. Some of the observations reported in this paper were obtained with the Southern African Large Telescope (SALT) under programs: 2019-1-MLT-006, 2019-2-SCI-011, 2020-1-SCI-014, 2020-2-SCI-007, and 2021-1-MLT-004 (PI: Ewa Niemczura). Polish participation in SALT is funded by grant No. MEiN nr 2021/WK/01. E.N. is supported by the IDUB project held at the University of Wrocław.

Appendix A

Determination of Uncertainties

For a representative line of a given element, six major sources are included in the uncertainty determinations: the uncertainty on the effective temperature ($\sigma_{T_{\text{eff}}}$), on the surface gravity ($\sigma_{\log g}$), on the microturbulent velocity (σ_{ξ_r}), on the apparent rotational velocity ($\sigma_{v_e \sin i}$), the oscillator strength ($\sigma_{\log gf}$), and the continuum placement (σ_{cont}). These uncertainties are supposed to be independent, so that the total uncertainty σ_{tot_i} for a given transition (i) verifies:

$$\sigma_{\text{tot}_i}^2 = \sigma_{T_{\text{eff}}}^2 + \sigma_{\log g}^2 + \sigma_{\xi_r}^2 + \sigma_{v_e \sin i}^2 + \sigma_{\log gf}^2 + \sigma_{\text{cont}}^2 \quad (\text{A1})$$

The mean abundance $\langle \left[\frac{X}{H} \right] \rangle$ is then computed as a weighted mean of the individual abundances $[X/H]_i$ derived for each transition (i):

$$\left\langle \left[\frac{X}{H} \right] \right\rangle = \frac{\sum_i \left(\left[\frac{X}{H} \right]_i / \sigma_{\text{tot}_i}^2 \right)}{\sum_i (1 / \sigma_{\text{tot}_i}^2)} \quad (\text{A2})$$

and the total error, σ is given by:

$$\frac{1}{\sigma^2} = \sum_{i=1}^N (1 / \sigma_{\text{tot}_i}^2) \quad (\text{A3})$$

where N is the number of lines per element.

Appendix B

Identification of the Lines for HD 99803A (HR 4423)

The lines in the HARPS spectrum of HD 99803A (HR 4423) absorbing more than 2% of the continuum have been identified using the synthetic spectrum which best fits the observed spectrum. They are collected in the Table 10.

ORCID iDs

Richard Monier  <https://orcid.org/0000-0002-0735-5512>
D. W. Kurtz  <https://orcid.org/0000-0002-1015-3268>
S. Rappaport  <https://orcid.org/0000-0003-3182-5569>
D. M. Bowman  <https://orcid.org/0000-0001-7402-3852>
Simon J. Murphy  <https://orcid.org/0000-0002-5648-3107>
Yveline Lebreton  <https://orcid.org/0000-0002-4834-2144>
Remko Stuik  <https://orcid.org/0000-0001-7797-3749>
Morgan Deal  <https://orcid.org/0000-0001-6774-3587>
T. Kılıçoğlu  <https://orcid.org/0000-0001-6947-822X>
M. Gebran  <https://orcid.org/0000-0002-8675-4000>

References

Adelman, S. J. 1991, *MNRAS*, **252**, 116
Adelman, S. J., Cowley, C. R., Leckrone, D. S., Roby, S. W., & Wahlgren, G. M. 1993, *ApJ*, **419**, 276
Adelman, S. J., & Pintado, O. I. 1997, *A&AS*, **125**, 219

Bohlender, D. A., Dworetsky, M. M., & Jomaron, C. M. 1998, *ApJ*, **504**, 533
Caffau, E., Ludwig, H. G., Steffen, M., Freytag, B., & Bonifacio, P. 2011, *SoPh*, **268**, 255
Carlsson, M. 1986, *UppOR*, **33**, 117
Carlsson, M. 1992, in ASP Conf. Ser. 26, Cool Stars, Stellar Systems, and the Sun, ed. M. S. Giampapa & J. A. Bookbinder (San Francisco, CA: ASP), 499
Castelli, F., & Hubrig, S. 2004, *A&A*, **425**, 263
Castelli, F., & Kurucz, R. L. 2003, in IAU Symp. 210, Modelling of Stellar Atmospheres, ed. N. Piskunov, W. W. Weiss, & D. F. Gray (San Francisco, CA: ASP), A20
Catanzaro, G., & Leone, F. 2002, *NewA*, **7**, 495
Corbally, C. J. 1984, *ApJS*, **55**, 657
Cowley, C. R., Hubrig, S., Castelli, F., González, J. F., & Wolff, B. 2007, *MNRAS*, **377**, 1579
Crause, L. A., Sharples, R. M., Bramall, D. G., et al. 2014, *Proc. SPIE*, **9147**, 91476T
Cunto, W., & Mendoza, C. 1992, *RMxAA*, **23**
Deal, M., Alecian, G., Lebreton, Y., et al. 2018, *A&A*, **618**, A10
Den Hartog, E. A., Lawler, J. E., Sneden, C., & Cowan, J. J. 2003, *ApJS*, **148**, 543
Dimitrijevic, M. S., & Sahal-Brechot, S. 1984, *A&A*, **136**, 289
Dolk, L., Wahlgren, G. M., & Hubrig, S. 2003, *A&A*, **402**, 299
Drawin, H. W. 1969, *ZPhy*, **225**, 483
Dworetsky, M. M., & Budaj, J. 2000, *MNRAS*, **318**, 1264
Dworetsky, M. M., Jomaron, C. M., & Smith, C. A. 1998, *A&A*, **333**, 665
Dworetsky, M. M., Persaud, J. L., & Patel, K. 2008, *MNRAS*, **385**, 1523
Ekström, S., Georgy, C., Eggenberger, P., et al. 2012, *A&A*, **537**, A146
Engleman, R. J. 1989, *ApJ*, **340**, 1140
Fuhr, J. R., Martin, G. A., & Wiese, W. L. 1988, *JPCRD*, **17**, 4
Garrison, R. F., & Gray, R. O. 1994, *AJ*, **107**, 1556
Grevesse, N., & Sauval, A. J. 1998, *SSRv*, **85**, 161
Guitou, M., Spielfiedel, A., Rodionov, D., et al. 2015, *CP*, **462**, 94
Hauck, B., & Mermilliod, M. 1998, *A&AS*, **129**, 431
Hidalgo, S. L., Pietrinferni, A., Cassisi, S., et al. 2018, *ApJ*, **856**, 125
Holt, R. A., Scholl, T. J., & Rosner, S. D. 1999, *MNRAS*, **306**, 107
Hubeny, I., & Lanz, T. 1992, *A&A*, **262**, 501
Jomaron, C. M., Dworetsky, M. M., & Bohlender, D. A. 1998, *CoSka*, **27**, 324
Kochukhov, O., Khalack, V., Kobzar, O., et al. 2021, *MNRAS*, **506**, 5328
Kramida, A., Ralchenko, Y., Reader, J., & NIST ASD Team 2018, 5, <http://physics.nist.gov/asd>
Kurtz, D. W. 1985, *MNRAS*, **213**, 773
Kurucz, R. L. 1992, *RMxAA*, **23**, 187
Kurucz, R. L. 2005, *MSAIS*, **8**, 14
Kurucz, R. L. 2013, ATLAS12: Opacity sampling model atmosphere program, Astrophysics Source Code Library, ascl:1303.024
Kurucz, R. L., & Avrett, E. H. 1981, *SAO Special Report*, **391**
Lanz, T., Artru, M. C., Didelon, P., & Mathys, G. 1993, *A&A*, **272**, 465
Lanz, T., Dimitrijevic, M. S., & Artru, M.-C. 1988, *A&A*, **192**, 249
Lawler, J. E., Den Hartog, E. A., Sneden, C., & Cowan, J. J. 2006, *ApJS*, **162**, 227
Lawler, J. E., Wickliffe, M. E., Cowley, C. R., & Sneden, C. 2001a, *ApJS*, **137**, 341
Lawler, J. E., Wickliffe, M. E., den Hartog, E. A., & Sneden, C. 2001b, *ApJ*, **563**, 1075
Lebreton, Y., & Reese, D. R. 2020, *A&A*, **642**, A88
Leone, F., & Manfre, M. 1997, *A&A*, **320**, 257
Lodders, K. 2010, *Astrophysics and Space Science Proceedings, Principles and Perspectives in Cosmochemistry*, Vol. 16 (Berlin: Springer), 379
Marques, J. P., Goupil, M. J., Lebreton, Y., et al. 2013, *A&A*, **549**, A74
Martin, G. A., Fuhr, J. R., & Wiese, W. L. 1988, *Atomic Transition Probabilities. Scandium Through Manganese* (New York: AIP)
Merle, T., Thévenin, F., Pichon, B., & Bigot, L. 2011, *MNRAS*, **418**, 863
Monier, R., Gebran, M., & Royer, F. 2015, *A&A*, **577**, A96
Monier, R., Gebran, M., Royer, F., Kilicoglu, T., & Frémat, Y. 2018, *ApJ*, **854**, 50
Monier, R., Griffin, E., Gebran, M., et al. 2019, *AJ*, **158**, 157
Monier, R., Niemczura, E., & Kılıçoğlu, T. 2020, *RNAAS*, **4**, 234
Moon, T. T., & Dworetsky, M. M. 1985, *MNRAS*, **217**, 305
Morel, P., & Lebreton, Y. 2008, *Ap&SS*, **316**, 61
Napiwotzki, R., Schoenberner, D., & Wenske, V. 1993, *A&A*, **268**, 653
Nielsen, K., Karlsson, H., & Wahlgren, G. M. 2000, *A&A*, **363**, 815
Nielsen, K. E., Wahlgren, G. M., Proffitt, C. R., Leckrone, D. S., & Adelman, S. J. 2005, *AJ*, **130**, 2312
Nilsson, H., Ljung, G., Lundberg, H., & Nielsen, K. E. 2006, *A&A*, **445**, 1165
Pál, A. 2012, *MNRAS*, **421**, 1825

- Perruchot, S., Kohler, D., Bouchy, F., et al. 2008, *Proc. SPIE*, **7014**, 70140J
- Ricker, G. R., Winn, J. N., Vanderspek, R., et al. 2015, *JATIS*, **1**, 014003
- Roby, S. W., & Lambert, D. L. 1990, *ApJS*, **73**, 67
- Roby, S. W., Leckrone, D. S., & Adelman, S. J. 1999, *ApJ*, **524**, 974
- Royer, F., Gebran, M., Monier, R., et al. 2014, *A&A*, **562**, A84
- Ryabchikova, T., Piskunov, N., Kurucz, R. L., et al. 2015, *PhyS*, **90**, 054005
- Ryabchikova, T., Ryabtsev, A., Kochukhov, O., & Bagnulo, S. 2006, *A&A*, **456**, 329
- Ryabchikova, T., Sachkov, M., Kochukhov, O., & Lyashko, D. 2007, *A&A*, **473**, 907
- Sakhibullin, N. A. 1987, *SvA*, **31**, 151
- Salpeter, E. E. 1955, *ApJ*, **121**, 161
- Savanov, I., & Hubrig, S. 2003, in ASP Conf. Ser. 305, Magnetic Fields in O, B and A Stars: Origin and Connection to Pulsation, Rotation and Mass Loss, ed. L. A. Balona, H. F. Henrichs, & R. Medupe (San Francisco, CA: ASP), **230**
- Schoening, T., & Butler, K. 1989, *A&A*, **78**, 51
- Seaton, M. J. 1962, *PPS*, **79**, 1105
- Shamey, L. J. 1969, PhD thesis, Univ. of Colorado at Boulder
- Smalley, B., Antoci, V., Holdsworth, D. L., et al. 2017, *MNRAS*, **465**, 2662
- Smith, K. C. 1994, *A&A*, **291**, 521
- Smith, K. C. 1996, *Ap&SS*, **237**, 77
- Smith, K. C. 1997, *A&A*, **319**, 928
- Smith, K. C., & Dworetzky, M. M. 1993, *A&A*, **274**, 335
- Stuik, R., Bailey, J. I., Dorval, P., et al. 2017, *A&A*, **607**, A45
- Takeda, Y., Kawanomoto, S., & Ohishi, N. 2014, *PASJ*, **66**, 23
- Takeda, Y., Takada-Hidai, M., Jugaku, J., Sakaue, A., & Sadakane, K. 1999, *PASJ*, **51**, 961
- Talens, G. J. J., Spronck, J. F. P., Lesage, A. L., et al. 2017, *A&A*, **601**, A11
- Talens, G. J. J., Deul, E. R., Stuik, R., et al. 2018, *A&A*, **619**, A154
- Thiam, M., LeBlanc, F., Khalack, V., & Wade, G. A. 2010, *MNRAS*, **405**, 1384
- Vidal, C. R., Cooper, J., & Smith, E. W. 1973, *ApJS*, **25**, 37
- Wiese, W. L., Fuhr, J. R., & Deters, T. M. 1996, Atomic Transition Probabilities of Carbon, Nitrogen, and Oxygen : a Critical Data Compilation (Gaithersburg, MD: NIST)
- Woolf, V. M., & Lambert, D. L. 1999, *ApJ*, **521**, 414
- Zdravkov, T., & Pamyatnykh, A. A. 2008, *JPhCS*, **118**, 012079