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# Diffuse radio source candidate in CIZA J1358.9–4750

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## Abstract

We report on results of our upgraded Giant Metrewave Radio Telescope (uGMRT) observations for an early-stage merging galaxy cluster, CIZA J1358.9–4750 (CIZA1359), in

Band-3 (300–500 MHz). We achieved the image dynamic range of  $\sim 38000$  using the direction dependent calibration and found a candidate of diffuse radio emission at  $4\sigma_{\text{rms}}$  significance. The flux density of the candidate at 400 MHz,  $24.04 \pm 2.48$  mJy, is significantly positive compared to noise, where its radio power,  $2.40 \times 10^{24}$  W Hz $^{-1}$ , is consistent with those of typical diffuse radio sources of galaxy clusters. The candidate is associated with a part of the X-ray shock front at which the Mach number reaches its maximum value of  $\mathcal{M} \sim 1.7$ . The spectral index ( $F_\nu \propto \nu^\alpha$ ) of the candidate,  $\alpha = -1.22 \pm 0.33$ , is in agreement with an expected value derived from the standard diffusive shock acceleration (DSA) model. However, such a low Mach number with a short acceleration time would require seed cosmic rays supplied from active galactic nucleus (AGN) activities of member galaxies, as suggested in some other clusters. Indeed, we found seven AGN candidates inside the diffuse source candidate. Assuming the energy equipartition between magnetic fields and cosmic rays, the magnetic field strength of the candidate was estimated to be  $2.1 \mu\text{G}$ . We also find head–tail galaxies and radio phoenixes or fossils near CIZA1359.

**Key words:** galaxies: clusters: individual (CIZA J1358.9–4750)—radio continuum: galaxies—X-rays: galaxies: clusters

## 1 Introduction

The largest self-gravitating systems in the universe, galaxy clusters, are several Mpc in size and  $10^{14}$ – $10^{15} M_\odot$  in mass and are known to contain a hot ( $10^7$ – $8$  K) intracluster medium (ICM). This thermal energy is thought to be converted from the huge gravitational energy of the large-scale structure through the bottom-up structure formation. Pairs of sub-clusters in close proximity are thought to be colliding with each other (e.g., see Markevitch & Vikhlinin 2007 for a review); they are called merging clusters and are the sites of this energy conversion. A major energy-conversion mechanism is believed to be a shock wave formed during the merger. However, detailed physical mechanisms of shock waves such as particle acceleration, magnetic-field amplification, and turbulence generation are longstanding questions in astrophysics.

A shock wave in the ICM is often identified from X-ray observations of the ICM's density and/or temperature jumps. The shock is also found from radio observations of synchrotron radiation, called radio relics, emitted from the shock-accelerated cosmic-ray electrons (e.g., see Feretti et al. 2012; Akahori et al. 2018b; van Weeren et al. 2019 for reviews). The Fermi first-order acceleration, namely the diffusive shock acceleration (DSA; Blandford & Eichler 1987), is thought to be one of the plausible theories for the particle acceleration. Meanwhile, there are other classes of diffuse radio emission in galaxy clusters, radio halos, mini-halos, and radio bridges. Some of them are thought to be formed by turbulence based on the Fermi second-order acceleration (Brunetti et al. 2001; Fujita et al. 2003). A radio bridge is

a relatively new class of diffuse radio emission, which was found at the linked region of early-stage merging clusters (Govoni et al. 2019; Brunetti & Vazza 2020; Botteon et al. 2020). Since turbulence acceleration is affected by acceleration efficiency, seed cosmic-rays supplied from active galactic nucleus (AGN) jets of member galaxies were proposed for the radio bridges (Brunetti & Vazza 2020).

Another important factor in the thermal evolution of the ICM is AGN jets launched from the supermassive black holes of member galaxies. In the last decades, AGN jets are thought to be a promising source in the solution of the so-called cooling flow problem (see Fabian 1994 for a review), while the co-existence of cooling gas and AGN jets in the Phoenix galaxy cluster (Kitayama et al. 2020; Akahori et al. 2020) raises a new question on AGN feedback. A bent AGN jet in Abell 3376 indicated a tight connection between the jet and the coherent magnetic field at the cold front of the cluster (Chibueze et al. 2021). The spectral index distribution exhibits a plateau near the bending point, suggesting re-acceleration of cosmic rays likely by magnetic reconnection. Recently, AGN jets have become more recognized as a source of cosmic rays in the ICM. AGN jets connecting to radio relics are found in, for example, Abell 3411 (van Weeren et al. 2017) and Abell 3376 (Chibueze et al. 2023). In Abell 3411, the spectral index changes continuously along the radio structure, indicating spectral aging caused by cosmic ray electron cooling. Therefore, a detailed study of radio sources in galaxy clusters can provide new knowledge, such as particle acceleration, in addition to understanding the evolution of the sources themselves.

**Table 1.** Basic parameters of CIZA1359.

| Parameter                     | Value   | Reference* |
|-------------------------------|---|------------|
| RA (J2000.0)                  | 13 <sup>h</sup> 58 <sup>m</sup> 40 <sup>s</sup> | [1]        |
| Dec. (J2000.0)                | −47°46′00″                                      | [1]        |
| Redshift                      | 0.0740  | [1]        |
| $f_X[0.1 \sim 2.4\text{keV}]$ | $20.89 \times 10^{-12} \text{ erg cm}^{-1}$     | [1]        |
| $L_X[0.1 \sim 2.4\text{keV}]$ | $4.88 \times 10^{44} \text{ erg s}^{-1}$        | [1]        |
| kT (keV)                      | $5.6 \pm 0.2 \text{ keV}$ (south-east)          |            |
|                               | $4.6 \pm 0.2 \text{ keV}$ (north-west)          | [2]        |

\*References: [1] Kocevski et al. 2007; [2] Kato et al. 2015.

CIZA J1358.9–4750 (CIZA1359) is thought to be one of the only several known early-stage merging galaxy clusters. This object is a Clusters in the Zone of Avoidance (CIZA) survey target, thus CIZA1359 is found relatively close to the Galactic plane (Kocevski et al. 2007). These basic information is summarized in table 1. Kato et al. (2015) performed a detailed analysis of the Suzaku X-ray observation and found a discontinuous high-temperature region in the linked region between the two X-ray peaks of subclusters. They suggested that the high-temperature region was formed by the merger shock wave passing along the merger axis. Omiya et al. (2023) further studied X-ray properties of CIZA1359 and found another shock front at the northern edge of the hot region. These studies also suggest that the merger axis is off the plane of the sky.

CIZA1359 is relatively nearby at redshift  $z = 0.074$ , which is convenient for studying it in detail, and is expected to be investigated more closely. In this paper, we report on the results of the upgraded Giant Metrewave Radio Telescope (uGMRT) observation of CIZA1359, with the aim of detecting any diffuse emission of CIZA1359. In section 2 we describe the details of the uGMRT observations and the data reduction, and in section 3 we present some obtained radio images and spectral index maps. In section 4, we discuss the relic candidate of CIZA1359. We have used cosmological parameters  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_M = 0.3$ , and  $\Omega_\Lambda = 0.7$  in this work.

## 2 Observation and data reduction

### 2.1 The observation

We conducted uGMRT Band 3 (300–500 MHz) observations of CIZA1359 (the project code 39\_045). Both narrow- and wide-band modes were adopted. The center frequency and the bandwidth of the narrow-band mode are 317 MHz and 33 MHz, respectively, and those of the wide-band mode are 400 MHz and 200 MHz, respectively. The field of view and the angular resolution are 75′ and 8″.3, respectively, both in diameter at 400 MHz.

The observations were carried out in Cycle 39 and were split into two separate observations on 2021 January 13–14 (day 1) and 2021 February 24–25 (day 2) in International Atomic Time (TAI). The observing time was 5 hr and 15 min on day 1, and 4 hr and 16 min on day 2. In each day, we observed a flux density, bandpass, and polarization calibrator, 3C286, for 10 min at the beginning and the end of the observation, and observed a phase calibrator, 1349–393, for 5 min every 25 min. Thus, the target (CIZA1359) on-source time was 337 min in total; 196 min on day 1 and 141 min on day 2. According to the observation log, two 45-m-diameter antennas, C03 and C11, were not used on day 2. Therefore, only day 1 data were used in this study, as day 2 data tend to be a little noisy.

### 2.2 The data reduction

The data were analysed using SPAM (Source Peeling and Atmospheric Modeling; Intema 2014) which is based on the AIPS (Astronomical Image Processing System), produced and maintained by NRAO (the National Radio Astronomy Observatory). SPAM employed the AIPS 31DEC13 and was controlled by python 2.7. There is a bright compact source of about  $2 \times 10^3 \text{ mJy}$  in the field of view, and the beam pattern of this bright source is clearly visible in the dirty image, showing that this source contaminates the image by causing strong sidelobes. Therefore, we applied SPAM’s direction-dependent calibration (DDC; Intema et al. 2017) to improve the dynamic range of the final image. Self-calibration is also applied to the direction-independent calibration (DIC) before the DDC. The CLEAN algorithm is used for imaging in SPAM.

In the analysis of the narrow-band data, we aim to create a catalog of sources for use in the DDC for the wide-band data. To select peeling sources, a list of radio sources in the field-of-view was compiled using a source catalog from the Tata Institute of Fundamental Research (TIFR) GMRT Sky Survey (TGSS; Intema et al. 2017) in this analysis. Imaging in SPAM was performed with a Briggs robustness parameter of  $-1.0$ . The Python Blob Detector and Source Finder (PyBDSF; Mohan & Rafferty 2015) was used to catalog the compact sources in the field-of-view.

In the analysis of the wide-band data, the data were split into six sub-bands and the DDC was applied to each sub-band data using the source catalog obtained from the analysis of the narrow-band data. The final output from the SPAM was the outlier-removed  $uv$  data. After that, a full-band radio image was derived by combining the  $uv$  data of sub-bands using WSClean (w-stacking clean; Offringa et al. 2014). In the imaging with WSClean, we first employ uniform weighting of the robustness to identify compact sources from the data. The modeling and subtraction of

**Table 2.** Radio maps in this paper.

| Label              | Frequency<br>(MHz) | BW<br>(MHz) | RMS<br>(mJy beam <sup>-1</sup> ) | Beam size<br>(" × ") | Beam PA<br>(°) | Figure         |
|--------------------|--------------------|-------------|----------------------------------|----------------------|----------------|----------------|
| Narrow-band        | 317                | 33          | $8.2 \times 10^{-2}$             | $21.0 \times 6.9$    | -0.1           | —              |
| Wide-band          | 400                | 200         | $3.7 \times 10^{-2}$             | $14.8 \times 5.2$    | -6.2           | 1              |
| Sub-band 01        | 317                | 33          | $8.6 \times 10^{-2}$             | $22.8 \times 5.7$    | 1.6            | 7              |
| Sub-band 02        | 350                | 33          | $1.7 \times 10^{-1}$             | $15.1 \times 5.0$    | -0.5           | 7              |
| Sub-band 03        | 385                | 33          | $8.1 \times 10^{-2}$             | $17.8 \times 5.6$    | -0.3           | 7              |
| Sub-band 04        | 417                | 33          | $5.7 \times 10^{-2}$             | $16.2 \times 5.4$    | 1.6            | 7              |
| Sub-band 05        | 450                | 33          | $5.0 \times 10^{-2}$             | $14.7 \times 4.9$    | -1.1           | 7              |
| Sub-band 06        | 481                | 33          | $1.1 \times 10^{-1}$             | $12.7 \times 4.6$    | -1.4           | 7              |
| Smoothed Wide-band | 400                | 200         | $1.0 \times 10^{-1}$             | $25 \times 25$       | 0.0            | 1, 2, 3, and 7 |

the compact sources can reduce their sidelobes, particularly from bright ones. Next, a multi-scale CLEAN (Cornwell 2008; Rich et al. 2008) was performed with the robustness closer to natural weighting to derive the diffuse emission from the data.

The primary beam effect was corrected by using the AIPS task *pbcor*. The function applied to the primary beam model is  $f(x) = 1.0 - \frac{2.939x}{10^3} + \frac{33.312x^2}{10^7} - \frac{16.659x^3}{10^{10}} + \frac{3.066x^4}{10^{13}}$ , where  $x$  is the distance parameter (see the AIPS Cookbook for details).<sup>1</sup> AIPS was also used to edit the data at each frequency to the same pixel size and spatial resolution, if necessary.

When calculating the flux density  $F_\nu$ , we adopted an empirical 10% error ( $\sigma_{\text{abs}} = 0.1$ ) of the absolute flux density (see also section 3.1, according to the DDC flux decay). For a diffuse source, the flux density error,  $\sigma_{F_\nu}$ , was given by  $\sigma_{F_\nu} = \sqrt{(\sigma_{\text{rms}}\sqrt{N_b})^2 + (\sigma_{\text{abs}}F_\nu)^2}$ , where  $\sigma_{\text{rms}}$  is the rms noise in an image, and  $N_b$  is the number of beams in the diffuse source (e.g., Kale et al. 2022).

### 2.3 Other data

In this study, X-ray data from Suzaku are used to confirm the spatial correlation with the ICM (Kato et al. 2015). In addition, the ICM temperature inferred from the XMM-Newton data was used as an indicator of the shock region (Omiya et al. 2023). The MeerKAT Galaxy Cluster Legacy Survey Data Release 1 (MGCLS DR1; Knowles et al. 2022) were also combined to determine the spectral index. The used MGCLS data were *Enhanced imaging products*, which was corrected the primary beam effect.

## 3 Results

### 3.1 Total intensity maps

We derived the total intensity map of the narrow-band data at 317 MHz. The DDC was applied to the imaging and

the robustness parameter of  $-1.0$  was set. The rms noise level of the map is  $8.2 \times 10^{-2}$  mJy beam<sup>-1</sup> with the DDC (table 2), while it is 1.73 mJy beam<sup>-1</sup> with the DIC. Therefore, the DDC improved the sensitivity and dynamic range by more than one order of magnitude. We found and cataloged 423 radio sources using the PyBDSF. The brightest source is PMN J1401–4733 at the north-east of the field of view, with a total flux density of  $2.01 \times 10^3$  mJy and a peak intensity of  $1.41 \times 10^3$  mJy beam<sup>-1</sup>. Thus, the achieved image dynamic range is 17195.

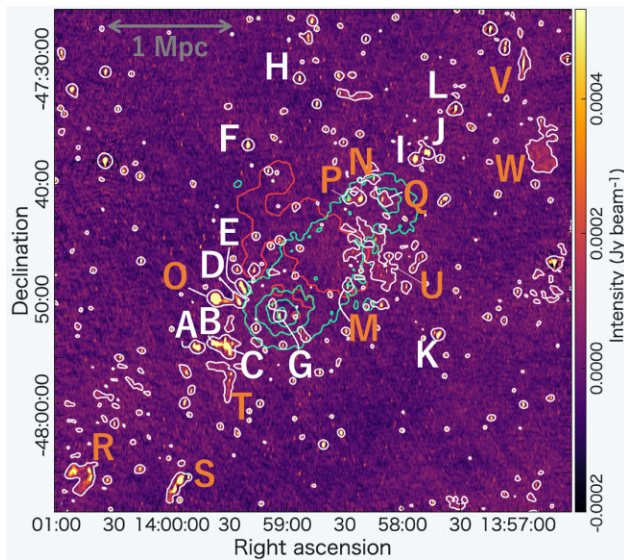
Next, we derived the total intensity map of the wide-band data at 400 MHz. We used the above source catalog as a prior sky model for the DDC, then obtained the rms noise level of the map,  $3.7 \times 10^{-2}$  mJy beam<sup>-1</sup> (table 2), or the achieved image dynamic range of 38108 for the robustness parameter of  $-1.0$ . The wide-band image of CIZA1359 is shown in figure 1. The background colour is the radio intensity distribution, where the white contour indicates the 0.4 mJy beam<sup>-1</sup> level ( $4\sigma_{\text{rms}}$ ) for the uGMRT smoothed at 25". The alphabetic labels indicate 23 distinguished sources. The white labels from A to L are known sources, while the orange labels from M to W are newly detected extended sources.<sup>2</sup>

One of the known issues on the DDC analysis is that a lot of DDC solutions result in global decay of the fluxes across the field (e.g., Patil et al. 2016). To assess this decay, we checked the visibility amplitude with respect to the DDC and DIC results. The maximum angular scale of Source U, which is the largest feature among the structures detected in this paper, is 6', which is about 0.6 kilo wavelength at 300 MHz. The medians of the visibility amplitude over a range of 0.6 kilo wavelength are 1.168 Jy for the DDC and 1.247 Jy for the DIC, respectively. Therefore, we measured an offset of about 0.94 in the amplitude ratio around the angular scale of Source U. The error can be compared to an empirical 10% error of the absolute flux we display in this paper. The flux accuracy checks are summarized in appendix 1.

<sup>1</sup> (<http://www.aips.nrao.edu/cook.html>).

<sup>2</sup> After we submitted this paper, Kale et al. (2022) reported Sources M to Q.





**Fig. 1.** Wide-band image of CIZA1359. The background image is the total intensity at 400 MHz with the bandwidth of 200 MHz. The beam pattern with the size,  $14''.8 \times 5''.2$ , is shown in the gray filled elliptical at the lower left-hand corner of the figure. The white contours show the intensity level of  $0.4 \text{ mJy beam}^{-1}$  which corresponds to  $4\sigma_{\text{rms}}$  with the smoothed  $25''$  resolution. The green contours show the X-ray surface brightness distribution of Suzaku (Kato et al. 2015) in arbital units at intervals of 1.81, 3.64, 5.46, and 7.28. The red contour shows the region with a temperature of 6 keV or above inferred from the recent X-ray observations (Omiya et al. 2023). The diffuse radio sources, including relic candidate, are labeled with letters A to W.

**Table 3.** Imaging parameters performed to focus on the diffuse emission.

| Weighting | RMS<br>( $\text{mJy beam}^{-1}$ ) | Beam size<br>( $'' \times ''$ ) | Beam PA<br>( $^\circ$ ) | Taper<br>( $''$ ) | Figure |
|-----------|-----------------------------------|---------------------------------|-------------------------|-------------------|--------|
| Uniform   | 0.18                              | $10.5 \times 3.2$               | -2.4                    | 22.5              | —      |
| Briggs -1 | 1.23                              | $25.2 \times 21.2$              | 24.7                    | 22.5              | —      |
| Briggs 0  | 1.56                              | $36.8 \times 22.7$              | 29.6                    | 22.5              | 2a     |
| Briggs 1  | 6.24                              | $296.3 \times 40.6$             | 24.7                    | 22.5              | —      |

To facilitate our discussion, a compact-source-subtracted image for Source U was derived. The compact sources were subtracted from the image by fitting with a point-source model with a Gaussian function using PyBDSF. We explored the robustness closer to natural weighting to image the diffuse emission. Table 3 summarizes the noise level and resolution from CLEAN with different weightings; we employed a robustness of 0.0. The results are shown in figure 2a. The hot ICM region is shown as red contours, which would indicate the approximate location of the merger shock front. The white contours are the same as that in figure 1. The black contours correspond to  $3\sigma_{\text{rms}}$  in the compact-source-subtracted image, which is shown by the background color.

### 3.2 Spectral index map

We convolved all images to a  $25''$  square beam using the AIPS task *convl*, where the pixel size and the number of pixels were fixed using the AIPS task *hgeom*. We then adopted the least-square fit to derive the best-fitting spectral index assuming a power-law and calculated the index pixel by pixel. We performed the fitting in a linear space to account for the negative flux value caused by the noise. To derive the spectral index, we added the MGCLS DR1 data (see subsection 2.3) to our uGMRT data. Details on the calculation of the spectral index are summarized in appendix 2. Figures 2b and 3 show the spectral index maps of newly detected extended sources in CIZA1359.

The background colour indicates the spectral index,  $\alpha$ , such that  $F_\nu \propto \nu^\alpha$ , and the white contour is the same as that in figure 1.

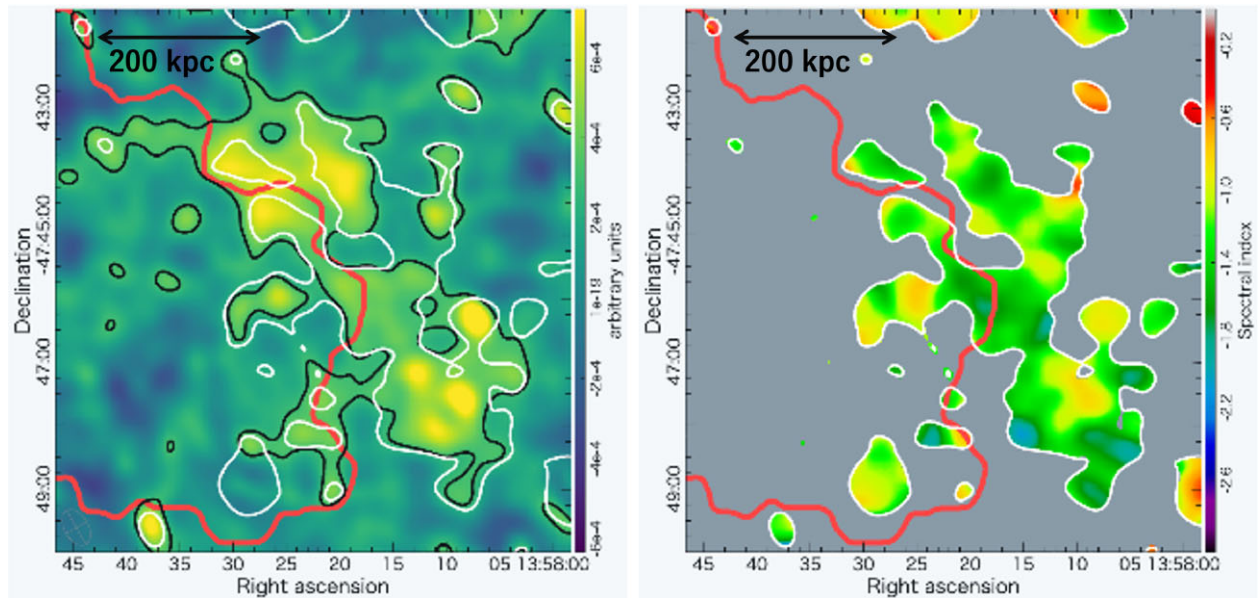
We also attempted to calculate in-band spectral indices for bright compact sources using sub-band images, for each GMRT and MeerKAT data. We obtained approximately the same index no matter which data is used for fitting for the sources, such as Source A. On the other hand, because faint sources including many diffuse-emission features have a low SN in each pixel, the sub-band spectral fitting results in a spectral index close to the slope of the noise floor. Therefore, for the faint sources we used the combined data in each band, where the center frequencies are 400 MHz and 1280 MHz, to derive the spectral index  $\alpha_{400-1280}$ .

In addition to the spectral index of each pixel, we derived the mean spectral index using the total flux densities shown in columns 4 and 5 of table 4. The total flux density was estimated by assuming the size of each source. The resultant mean spectral index is listed in column 6 of table 4. Here, the spectral index error was derived from the error propagation term  $\{\sqrt{[F_{\nu_1} \log(\nu_1/\nu_2)]^{-2} \sigma_{F_{\nu_1}}^2 + [-F_{\nu_2} \log(\nu_1/\nu_2)]^{-2} \sigma_{F_{\nu_2}}^2}\}$ .

### 3.3 Source catalog

We found 23 distinguished radio features in the image. We labeled them as Sources A to W (figure 1). The parameters for each source are summarized in table 4.

Sources A to L were reported in the previous ATCA observation (Akaishi et al. 2018a), where several sources are closely concentrated on the south-western rim of the southern subcluster of CIZA1359. The spectral indices are comparable to those estimated with ATCA and those of typical AGN. We found that some of the sources were clearly larger in spatial scale than the synthesized beam and were resolved into multiple components. Source G is located near the center of the south subcluster and has an apparent size of about  $25'' \times 22''$  ( $=38 \text{ kpc} \times 33 \text{ kpc}$ ) at the  $3\sigma_{\text{rms}}$



**Fig. 2.** Radio and spectral maps of Source U. White and red contours are the same as in figure 1. Left (a): The background color shows the point source subtracted intensity. This unit is equivalent to  $\text{Jy beam}^{-1}$ . Because the point source is subtracted with a Gaussian fit, we use the expression “arbitrary unit” in the sense that it is different from the color unit in figure 1. The black contours have a point source subtracted intensity of  $3.0 \times 10^{-4}$ , which corresponds  $3\sigma_{\text{rms}}$ . Right (b): Spectral index map of Source U.

signal-to-noise level. This is one order of magnitude smaller than the typical size of mini-halos ( $\sim 500$  kpc), so it is more likely to be an AGN radio lobe. Source G is cataloged in SIMBAD as the galaxy 2MASX J13590381–4751311. It has the redshift of  $z = 0.074$ , which is consistent with the redshift of CIZA1359.

Sources M to Q were reported in the previous GMRT observation (Kale et al. 2022). Similarly to Source G, Source Q is more likely to be an AGN, although it is located near the center of the northern subcluster with a size of  $60'' \times 30''$  ( $=90 \text{ kpc} \times 45 \text{ kpc}$ ). Source Q is cataloged in SIMBAD as the galaxy 2MASX J13581085–4741243. Its redshift,  $z = 0.074$ , is consistent with the redshift of CIZA1359. Source O, which is unresolved in TGSS and the Sydney University Molonglo Sky Survey (SUMSS; Bock et al. 1999) images, consists of a round structure and a narrow east–west linear structure. A high-resolution image such as MGCLS at 1.28 GHz shows a head–tail galaxy-like structure, while there is no corresponding source in the ATCA image at 2.1 GHz (Akahori et al. 2018a). Indeed, a steep spectral index of  $-2.01$  and the 400 MHz total flux density,  $48.73 \pm 4.88$  mJy, predicts a 2 GHz flux density of 1.9 mJy (with a size of  $90'' \times 90''$ ), which is almost the same as the sensitivity limit of the ATCA observation. Source O can be classified as an ultra-steep spectrum (USS) source found in galaxy clusters (Mandal et al. 2019), so that Source O may be a fossil plasma source. We could not find any corresponding source to Source O in SIMBAD.

Source R is about 2 Mpc away from the southern subcluster center to the south-east and is cataloged in SIMBAD

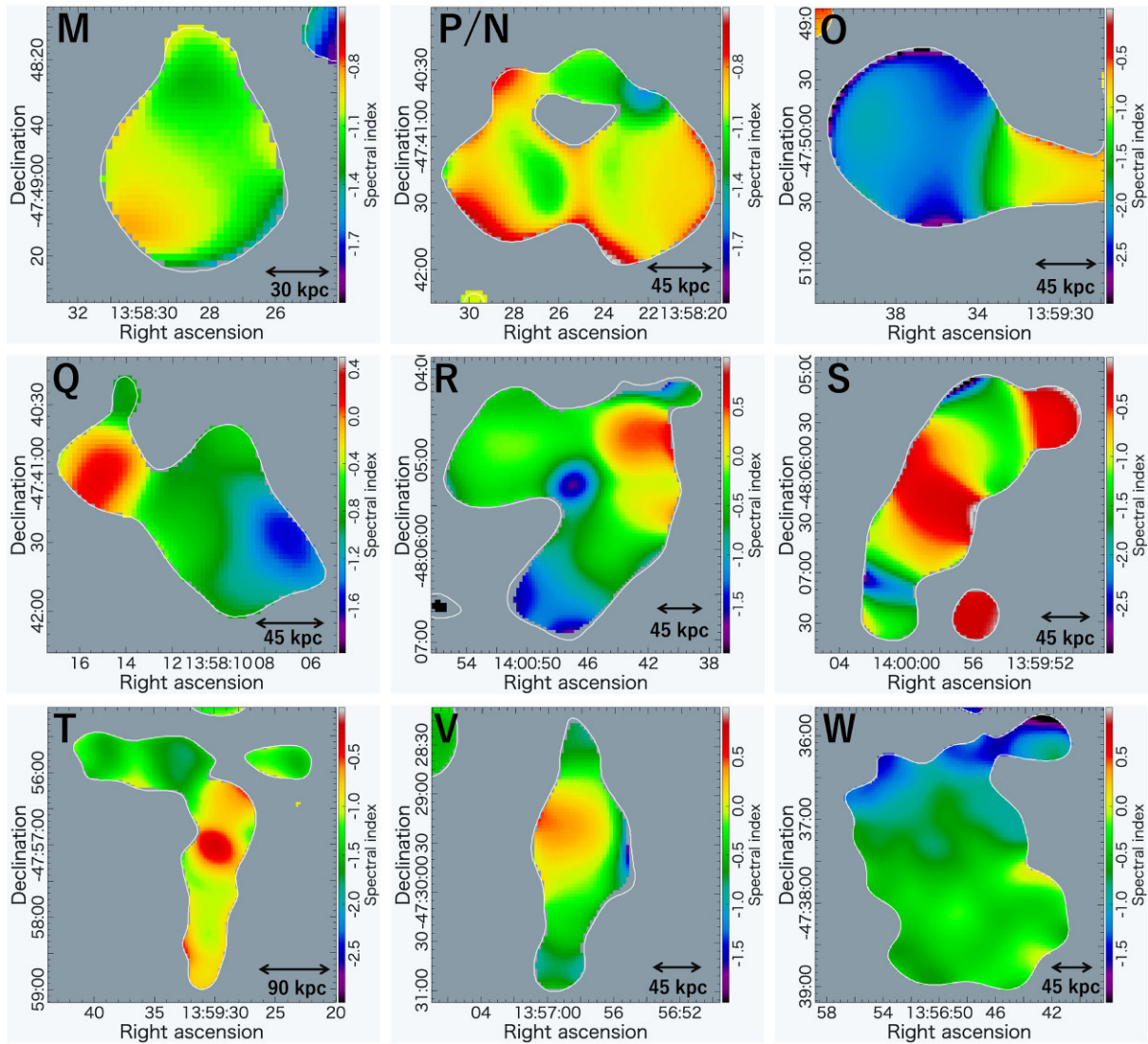
as the galaxy 2MASX J14004272–4804474. We found that it has a radio structure like a head–tail galaxy. There is also a spectral index gradient, with aging from the head to the tail. It has a redshift of  $z = 0.075$ , which is comparable to that of CIZA1359.

Source S has an FR II-like radio structure about 1.5 Mpc away from the southern subcluster center. SIMBAD cataloged a galaxy at  $z = 0.054$  as 2MASX J13595922–4805486 in the neighbourhood. However, 2MASX J13595922–4805486 may be associated with a faint radio structure seen at  $11''$  away to the northeast from Source S.

Source T is located at the southeast of the CIZA1359 and has a head–tail galaxy-like radio structure. It has an elongated structure to the south and a faint structure to the east. There is no associated source within  $30''$  in SIMBAD. At the peak flux position of the south component of Source T, there is a galaxy, 2MASX J13592976–4757043, in SIMBAD. The redshift of 2MASX J13592976–4757043 is  $z = 0.081$ , which is located at the far side of CIZA1359.

Source U is a candidate of diffuse cluster emission. It is located at the south of the northern subcluster and in between the two subclusters. Four galaxies were found by SIMBAD within a  $4'$  radius centred on Source U; 2MASX J13580947–4745213 is located at the north-west of Source U and has redshift  $z = 0.078$ , 2MASX J13581294–4748183 is located at the south of Source U with the redshift  $z = 0.069$ , 6dFGS gJ135839.1–474723 is located at the east of





**Fig. 3.** Spectral index maps of newly detected extended sources in CIZA1359. The spectral index  $\alpha$  is calculated using  $F_\nu \propto \nu^\alpha$ . The white contours are the same as figure 1. Note that the color range is different for each.

Source U with the redshift  $z = 0.067$ , and LEDA 184317 is located at the north-east of Source U with an unknown redshift. Source U is the largest diffuse source in the uGMRT image (see figure 1). The signal-to-noise ratio of Source U is around  $4\sigma_{\text{rms}}$ , where  $\sigma_{\text{rms}} = 0.10 \text{ mJy beam}^{-1}$ . With  $N_b = 23.0$ , the point source subtracted flux density is  $24.04 \pm 2.48 \text{ mJy}$ , which significantly deviates from the null. We explore Source U in detail in section 4.

Source V is cataloged in SIMBAD as a galaxy 2MASX J13565832–4729231, with the redshift  $z = 0.078$ , similar to that of CIZA1359. It has an elongated structure extending north–south. The extended structure appears to be connected to Source W; Source V is like a bipolar radio jet.

Source W is quadrangle in shape and is the second-largest radio structure in the image. No corresponding sources

were found in SIMBAD in the vicinity of this source. Sources V and W are similar to the structure known as a radio phoenix.

## 4 Discussion

We explore Source U in detail in subsection 4.1 next, followed by the discussion of its origin (likely a radio relic) supposing its detection, in subsection 4.2.

### 4.1 Source U: Diffuse radio structure candidate

#### 4.1.1 Location

First, we focus on the spatial location of Source U. As described in the Introduction, the recent X-ray observation found a pair of shock fronts in the linked region (Omiya



**Table 4.** Radio source catalog of the CIZA1359 field.\*

| Label<br>(1) | Mean<br>intensity<br>400 MHz <sup>†</sup><br>(2) | Peak flux<br>400 MHz <sup>†</sup><br>(3) | Flux                        |                              | $\alpha_{400-1280}$<br>(6) | Identification<br>(7)   |
|--------------|--|--|-----------------------------|------------------------------|----------------------------|---|
|              |  |  | 400 MHz <sup>†</sup><br>(4) | 1280 MHz <sup>†</sup><br>(5) |                            |   |
| A            | 4.74   | 17.85 ± 0.21                             | 25.90 ± 2.59                | 13.10 ± 1.32                 | -0.59 ± 0.28               | —   |
| B            | 4.86   | 17.41 ± 0.11                             | 29.62 ± 2.97                | 13.50 ± 1.37                 | -0.68 ± 0.28               | 2MASX J13593870-4753472   |
| C            | 3.72   | 17.54 ± 0.12                             | 43.54 ± 4.36                | 17.49 ± 1.77                 | -0.78 ± 0.28               | 2MASX J13593065-4754053   |
| D            | 2.47   | 10.47 ± 0.16                             | 14.04 ± 1.42                | 6.35 ± 0.67                  | -0.68 ± 0.29               | 2MASX J13592419-4750253   |
| E            | 3.26   | 11.29 ± 0.23                             | 21.98 ± 2.21                | 10.84 ± 1.11                 | -0.61 ± 0.28               | 2MASX J13592518-4749333   |
| F            | 2.56   | 9.28 ± 0.08                              | 20.30 ± 2.04                | 9.58 ± 0.97                  | -0.65 ± 0.28               | —   |
| G            | 1.18   | 4.63 ± 0.10                              | 10.93 ± 1.11                | 4.88 ± 0.52                  | -0.69 ± 0.29               | 2MASX J13590381-4751311   |
| H            | 2.48   | 9.63 ± 0.09                              | 15.31 ± 1.54                | 9.37 ± 0.95                  | -0.42 ± 0.28               | —   |
| I            | 2.91   | 13.71 ± 0.14                             | 31.22 ± 3.13                | 15.12 ± 1.52                 | -0.62 ± 0.28               | 2MASX J13575383-4737543   |
| J            | 5.79   | 28.21 ± 0.26                             | 38.69 ± 3.88                | 17.56 ± 1.77                 | -0.68 ± 0.28               | —   |
| K            | 3.12   | 11.54 ± 0.18                             | 15.18 ± 1.53                | 6.16 ± 0.64                  | -0.78 ± 0.29               | —   |
| L            | 5.47   | 30.51 ± 0.07                             | 32.66 ± 3.27                | 13.51 ± 1.37                 | -0.76 ± 0.28               | —   |
| M            | 2.20   | 8.77 ± 0.13                              | 9.11 ± 0.93                 | 2.73 ± 0.32                  | -1.04 ± 0.31               | —   |
| N            | 1.81   | 8.20 ± 0.16                              | 11.13 ± 1.13                | 3.63 ± 0.42                  | -0.96 ± 0.30               | 1RXS J135821.7-474126   |
| O            | 4.70   | 19.57 ± 0.63                             | 50.57 ± 5.06                | 4.87 ± 0.56                  | -2.01 ± 0.30               | —   |
| P            | 1.30   | 4.12 ± 0.14                              | 5.55 ± 0.58                 | 1.83 ± 0.25                  | -0.95 ± 0.34               | —   |
| Q            | 0.72   | 1.99 ± 0.19                              | 5.52 ± 0.60                 | 2.59 ± 0.35                  | -0.65 ± 0.34               | 2MASX J13581085-4741243   |
| R            | 1.39   | 5.60 ± 0.42                              | 31.06 ± 3.13                | 27.22 ± 2.75                 | -0.11 ± 0.28               | 2MASX J14004272-4804474   |
| S            | 2.41   | 9.79 ± 0.29                              | 31.32 ± 3.15                | 13.29 ± 1.36                 | -0.74 ± 0.28               | —   |
| T            | 0.65   | 2.07 ± 0.32                              | 10.06 ± 1.06                | 3.43 ± 0.48                  | -0.93 ± 0.35               | 2MASX J13592976-4757043   |
| U            | 0.54   | 1.29 ± 0.15                              | 28.93 ± 2.96                | 7.03 ± 0.94                  | -1.22 ± 0.33               | 2MASX J13580947-4745213<br>2MASX J13581294-4748183<br>6dFGS gJ135839.1-474723 |
| V            | 1.57   | 7.34 ± 0.34                              | 12.75 ± 1.30                | 14.39 ± 1.46                 | 0.10 ± 0.28                | 2MASX J13565832-4729231   |
| W            | 0.74   | 1.24 ± 0.20                              | 21.34 ± 2.18                | 11.60 ± 1.25                 | -0.52 ± 0.29               | —   |

\*Columns: (1) Source name. (2) Average intensity in mJy beam<sup>-1</sup>, which adopts the white contours of figure 1 as the size of the source. The error can be used  $3.7 \times 10^{-2}$  mJy beam<sup>-1</sup>. (3) Peak intensity in mJy beam<sup>-1</sup>; the size of the source is the same as column 2. The error used is an rms noise around the source in the point source subtracted image. (4) Integrated flux which is not subtracting point sources in mJy at 400 MHz; the size of the source is the same as column 2. (5) Same as column 4 but frequency is 1280 MHz. (6) Spectral index, which is calculated using the fluxes in columns 4 and 5. (7) Corresponding galaxies nearby in sky-plane.

<sup>†</sup>Data sources: Columns (2), (3), and (4) This work; Column (5) MGCLS (Knowles et al. 2022).

et al. 2023); the north shock at the northern edge of the hot region, and the south shock at the southern edge of the hot region, where the hot region is shown as the red solid line in figure 2a. The pair of shocks seem to have emerged from the interface of the subclusters and be propagating toward each subcluster core.

Such a merger shock has been considered as a site of cluster diffuse radio emission, based on an expectation that the shock accelerates cosmic-ray electrons emitting synchrotron radiation. In fact, the location of Source U is broadly consistent with that of the western part of the north shock indicated by the red contours in figure 2a. In fact, the shape of Source U and the shock plane do not exactly coincide with each other. If Source U was excited by a shock wave, it would flow down and age at the downstream-side of the shock front. However, it is located at the slightly upstream side. This would be interpreted as being due to the misalignment of the merger axis

with the sky plane, i.e., a projection effect of the viewing angle.

#### 4.1.2 Structure

The shock-associated diffuse radio emission is often seen in late-stage merging clusters and they are called “radio relics.” Although CIZA1359 is known as an early-stage merging cluster, Source U extends about  $5' \times 6'$  (=450 kpc  $\times$  540 kpc), which is comparable in size to radio relics (Ferretti et al. 2012).

In the above shock scenario, one may also expect radio emission from the south shock, although our observation did not find any candidate. Interestingly, Omiya et al. (2023) estimated the Mach number of the shocks and found that the north shock has a higher Mach number,  $\mathcal{M} = 1.7$ , while the other part has a lower value of  $\mathcal{M} = 1.4$ . Therefore, Source U is consistent with the theoretical expectation that a shock wave with a higher Mach number forms a

brighter radio emission so that it accelerates cosmic rays more efficiently (Ha et al. 2018). However, DSA does not work well at such low Mach numbers. We discuss the need for re-acceleration in subsection 4.2.

#### 4.1.3 Radio power

The radio power of diffuse radio emission in galaxy clusters has been studied in the literature and thus the radio power is also useful to examine whether Source U is a real emission or not. We calculate the monochromatic radio power using the following equation:

$$P_\nu = 4\pi D_L^2 \int I_\nu d\Omega, \quad (1)$$

where  $D_L(z = 0.07) \sim 9.74 \times 10^{24}$  m is the luminosity distance (Wright 2006),  $I_\nu$  is the radio intensity, and  $\Omega$  is the area of diffuse emission.

We integrate  $I_\nu$ , which is the point source subtracted intensity, within the area of the white contours in figure 2a. To compare the monochromatic radio power in this work with those of 1.4 GHz in the literature, it was converted to the flux at 1.4 GHz from flux at 400 MHz using the spectral index of  $-1.22$ . The derived radio power is  $P_{1.4\text{ GHz}} = 2.40 \times 10^{24}$  W Hz $^{-1}$ . This is consistent with the known radio relics, halos, shown in figure 4. Therefore, there is no immediate problem in considering Source U as a diffuse radio source of a galaxy cluster.

Note that if we adopt the formula of equation (2) in Akahori et al. (2018a) to derive the upper limit of the radio power, we obtain a value that is broadly consistent with that derived in their work. The reason for the non-detection in ATCA could be due to the steep spectrum of Source U.

#### 4.1.4 Magnetic field

Finally, although there are uncertainties caused by theoretical assumptions, it is possible to derive the magnetic field strength, and to discuss the reality of the candidate from comparison with previous estimations. Assuming the energy equipartition between magnetic fields and cosmic rays, the magnetic field strength can be estimated from synchrotron radiation as follows: (Beck & Krause 2005),

$$B_{\text{eq}} = \left[ \frac{4\pi(-2\alpha + 1)(K_0 + 1)I_\nu E_p^{1+2\alpha}(\nu/2c_1)^{-\alpha}}{(-2\alpha - 1)c_2(-\alpha)L c_4(i)} \right]^{1/(-\alpha+3)}, \quad (2)$$

where  $\alpha$  denotes the spectral index of synchrotron radiation,  $K_0$  is the number density ratio of cosmic ray nuclei to that of the electrons,  $L$  is the path length of the synchrotron emitting media,  $I_\nu$  is the intensity at frequency  $\nu$ , and  $E_p$  is the proton rest energy. The coefficients are

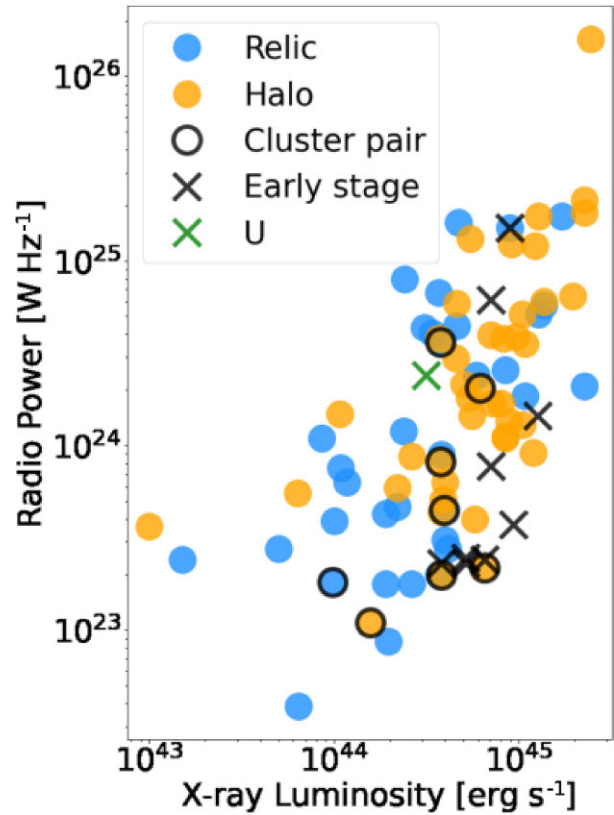


Fig. 4. Relationship between radio power at 1.4 GHz and X-ray luminosity at 0.1–2.4 keV of the diffuse radio sources. The green cross is data of Source U obtained from this study. The blue and orange dots show the data of the radio halo and relic, respectively, taken from Ferretti et al. (2012). Some cluster pairs (table 5) and early stage mergers (table 6) are shown by the black ring and black cross, respectively.

$c_1 = 3e/(4\pi m_e^3 c^5) = 6.3 \times 10^{18}$  erg $^{-2}$  s $^{-1}$  G $^{-1}$ ,  $c_2 = 4.56 \times 10^{-24}$  erg G $^{-1}$  sr $^{-1}$ ,  $c_4 = 1$ , and  $E_p = 1.5 \times 10^{-3}$  erg. We adopted our best-fitting value of  $\alpha = -1.22$  and  $\nu = 400$  MHz, and we assumed the typical values of  $L = 500$  kpc  $\sim 1.5 \times 10^{24}$  cm and  $K_0 = 100$  (Kierdorf et al. 2017). We then obtained the field strength of 2.1  $\mu$ G from the intensity of Source U, 0.4 mJy beam $^{-1}$ . Such a  $\mu$ G magnetic field is commonly found in galaxy clusters (e.g., Kierdorf et al. 2017). We note that this strength is insensitive to the parameters of equation (2). One  $\mu$ G-order field strength is derived even when the parameters are changed. However, figure 2 suggests that a patchy structure exists within Source U. The structure can cause an error in the estimation of the average strength in equation (2).

We can check the field strength from an empirical radial dependence of magnetic field. Using the following equation (Bonafede et al. 2010),

$$B(r) = B_0 \times \left[ \frac{n_e(r)}{n_0} \right]^\eta, \quad n_e = n_0 \left( 1 + \frac{r^2}{r_c^2} \right)^{-\frac{3}{2}\beta}, \quad (3)$$

**Table 5.** Radio parameters of the cluster pairs.\*

| Name  | RA                    | Dec                   | Structure | Log $P$ (1.4)          | $L_X(10^{44})$          | Pair  | Merger phase         |
|-------|-----------------------|-----------------------|-----------|------------------------|-------------------------|-------|----------------------|
| (1)   | ( $^{\circ}$ )<br>(2) | ( $^{\circ}$ )<br>(3) | (4)       | [W Hz $^{-1}$ ]<br>(5) | [erg s $^{-1}$ ]<br>(6) | (7)   | (8)                  |
| A209  | 22.990                | −13.576               | Halo      | 24.31                  | 6.17                    | A222  | —                    |
| A399  | 44.485                | 13.016                | Halo      | 23.30                  | 3.80                    | A401  | Early <sup>†</sup>   |
| A401  | 44.737                | 13.582                | Halo      | 23.34                  | 6.52                    | A399  | Early <sup>†</sup>   |
| A2061 | 230.336               | 30.671                | Relic     | 23.65                  | 3.95                    | A2067 | —                    |
| A2063 | 230.758               | 8.639                 | Relic     | 23.26                  | 0.98                    | MKW3s | —                    |
| A2256 | 255.931               | 78.718                | Halo      | 23.91                  | 3.75                    | A2271 | Complex <sup>‡</sup> |
| A2256 | 255.931               | 78.718                | Relic     | 24.56                  | 3.75                    | A2271 | Complex <sup>‡</sup> |
| A3562 | 203.383               | −31.673               | Halo      | 23.04                  | 1.57                    | A3558 | Complex <sup>§</sup> |

\*Columns: (1) Name of radio associated cluster. (2) Right ascension of the cluster. (3) Declination of the cluster. (4) Structure with radio information in columns (5) and (6). (5) Logarithm of radio power at 1.4 GHz from Feretti et al. (2012). (6) X-ray luminosity in the 0.1–2.4 keV band in  $10^{44}$  units from Feretti et al. (2012). (7) Names of cluster that pair with the one in column (1). (8) Expected merging phase; “early” and “complex” mean the early stage which has not completed a core-crossing and the multiple merger, respectively.

<sup>†</sup>Govoni et al. (2019).

<sup>‡</sup>Breuer et al. (2020).

<sup>§</sup>Haines et al. (2018).

the field strength is estimated to be  $2.7 \mu\text{G}$  at a relic position, 2.5 (225 kpc) away from the cluster centre in the north, assuming the central magnetic strength to be  $B_0 = 4.7 \mu\text{G}$  and the radial power-law slope  $\eta = 0.5$  (Bonafede et al. 2010). We also adopt the central gas density  $n_0 = 2.54 \times 10^{-3} \text{ cm}^{-3}$ , the core radius  $r_c = 165 \text{ kpc}$ , and the  $\beta$ -model parameter  $\beta = 0.67$  from (Omiya et al. 2023). Note that  $B_0$  and  $\eta$  are from the Coma cluster and are expected to vary with cluster. Indeed, in Abell 2382,  $B_0 = 3.5 \mu\text{G}$  and  $\eta = 0.5$  are obtained (Guidetti et al. 2008), and in Abell 2255,  $B_0 = 2 \mu\text{G}$  and  $\eta = 0.5$  are obtained (Govoni et al. 2006). Since these values for CIZA1359 are unknown, we adopted those of the Coma cluster as representative values.

The magnetic field strengths of CIZA1359 as estimated by these two independent methods are in good agreement although they include large theoretical assumptions. This means that the magnetic field strength of CIZA1359 is consistent with other galaxy clusters.

## 4.2 Origin of Source U

Our assessment of Source U based on its location, structure, spectrum, power, and magnetic field do not suggest that Source U is noise. In this subsection, we discuss the origin of Source U, supposing that Source U is real.

### 4.2.1 Comparison with other early-stage merging clusters

To understand the origin of Source U, it is useful to compare it with diffuse radio sources in early-stage merging clusters. This is because the cluster’s physical properties that would be related to the origin are very different between early- and late-stage merging clusters. Here, since it is rather

difficult to identify early-stage merging clusters, and there is no catalog of early-stage merging clusters yet, we look for them using the following two methods. The first is to find radio-associated cluster pairs, and the second is to make a list of well-known early-stage merging clusters.

Radio-associated cluster pairs were searched by catalog matching. We (1) checked the coordinates of the radio-associated clusters listed in tables 1 and 3 of Feretti et al. (2012) using SIMBAD, and (2) cataloged if they are within the Planck’s beam FWHM (7.18) from the coordinates listed in table 1 of Planck Collaboration (2013) using TOPCAT (Taylor 2005). As a result, we found seven radio-associated cluster pairs (table 5). It should be noted that this catalog may contain not only early-stage merging clusters but also late-stage merging clusters which have close separations of subclusters. Furthermore, some may be random pairs which are only close in projection on the plane of the sky and have significantly different redshift, although such a case is rare because galaxy clusters sparsely exist in the Universe.

Some clusters, including new discoveries, are well-known as early-stage merging clusters (e.g., 1E 2216.0–0401 and 1E 2215.7–0404; Gu et al. 2019). Table 6 summarizes the information of the clusters that are believed to be early-stage merging clusters. 1E2216.0–0401 indicates a temperature jump in the ICM between the cluster pair (Gu et al. 2019). Gu et al. (2019) suggested that the system is an early-stage merging cluster. They also reported diffuse radio sources between the cluster pairs. They concluded that they are bright AGNs affected in part by the merger shock. Abell 141 (Duchesne et al. 2021) also has a temperature jump and was reported to be an early-stage

**Table 6.** List of well-known early-stage merging clusters.\*

| Name           | RA<br>( $^{\circ}$ ) | Dec<br>( $^{\circ}$ ) | Radio source | Log $P(1.4)$<br>[W Hz $^{-1}$ ] | $L_X(10^{44})$<br>[erg s $^{-1}$ ] | Pair |
|----------------|----------------------|-----------------------|--------------|---------------------------------|------------------------------------|------|
| (1)            | (2)                  | (3)                   | (4)          | (5)                             | (6)                                | (7)  |
| A98            | 11.608               | 20.490                | Galaxy       | —                               | —                                  | 1    |
| A115           | 13.998               | 26.321                | Relic        | 25.18                           | 8.9                                | 2    |
| A141           | 16.376               | −24.655               | Unknown      | 24.16                           | 12.6 <sup>†</sup>                  | 3    |
| A399           | 44.485               | 13.016                | Halo         | 23.36                           | 3.8                                | 4    |
| A399/401 d     | 44.737               | 13.582                | Relic        | 23.39                           | 5.2 <sup>‡</sup>                   |      |
| A399/401 f     | 44.737               | 13.582                | Relic        | 23.36                           | 5.2 <sup>‡</sup>                   |      |
| A401           | 44.737               | 13.582                | Halo         | 23.38                           | 6.5                                |      |
| A1758          | 203.134              | 50.510                | Relic        | 23.57                           | 9.4 <sup>§</sup>                   | 5    |
| A1758N         | 203.134              | 50.510                | Halo         | 24.79                           | 7.1                                |      |
| A1758S         | 203.134              | 50.510                | Halo         | 23.89                           | 7.1                                |      |
| A1775          | 205.482              | 26.365                | Halo         | —                               | 1.5 <sup>§</sup>                   | 6    |
| A3391          | 96.564               | −53.681               | Galaxy       | —                               | 1.3 <sup>  </sup>                  | 7    |
| A3395          | 96.880               | −54.399               | Galaxy       | —                               | 1.3 <sup>  </sup>                  |      |
| 1E 2215.7−0404 | 334.585              | −3.828                | Galaxy       | —                               | 0.8 <sup>‡</sup>                   | 8    |
| 1E 2216.0−0401 | 334.673              | −3.766                | Galaxy       | —                               | 0.8 <sup>‡</sup>                   |      |
| CIZA1359       | 209.667              | −47.767               | Relic        | 24.38                           | 3.1 <sup>**</sup>                  | 9    |

\*Columns (1) to (6) are the same as in table 5. Column (7) is cluster pair number.

<sup>†</sup>Ebeling et al. (1996).

<sup>‡</sup>Average value of A399 and A401.

<sup>§</sup>Ebeling et al. (1998).

<sup>||</sup>Vikhlinin et al. (2009).

<sup>‡</sup>Gioia et al. (1990).

\*\*Kato et al. (2015).

merging cluster. Radio structures were also found between the subclusters, but it is not possible to isolate whether they are radio bridge, relic, or halo due to lack of spatial resolution. Abell 1775 has a similar X-ray morphology to early-stage merging clusters (Botteon et al. 2021). Slushing or slingshot effects have been reported. Diffuse radio sources were detected and reported; their structures seem to be slingshot radio halos associated with the X-ray structure. Abell 115 clearly shows two subclusters in its X-ray morphology (Lee et al. 2020). No radio structure was detected between them, while a radio relic is present and implies rather a late-type merging cluster. A3391-A3395 (Brüggen et al. 2021) and A98 have radio structures that seem to associate with the head–tail galaxy (Paterno-Mahler et al. 2014).

Figure 4 plots the radio-associated cluster pairs (table 5) and the early-stage merging clusters (table 6) as black circles and black crosses, respectively. We see that their distributions are very scattered in the relation between radio power and X-ray luminosity. CIZA1359, the green cross, appears to be inside the dispersed distribution, supporting that there is no contradiction in considering CIZA1359 to be a member of these families.

#### 4.2.2 Is Source U a radio relic?

We next discuss how Source U can be classified. As broadly described in the Introduction, there are some known classifications of diffuse radio emission of galaxy clusters (see, e.g., Feretti et al. 2012; van Weeren et al. 2019 for reviews). We look into the acceleration mechanism of cosmic ray electrons which produce the observed radio emission of Source U. Discussion about the acceleration mechanism is helpful for the source classification.

As discussed above, Source U seems to be associated with the northern shock with a Mach number of 1.7. Based on the standard DSA and test-particle regime, the energy spectrum of the relativistic electrons,  $p$ , of  $n(E)dE \propto E^{-p}dE$ , depends on the shock compression,  $C$ , as  $p = (C + 2)/(C - 1)$ . We obtain  $C \sim 1.96 \{ = [(3/4M^2) + 0.25]^{-1} \}$  for the Mach number 1.7, which gives  $p \approx 4.1$ . If the magnetic field is roughly constant over the radio source, such a power-law electron distribution will lead to synchrotron emission with a spectral index  $\alpha_M = -(p - 1)/2 \approx -1.56$  for  $F_\nu \propto \nu^{\alpha_M}$ . The observed spectral index and its error  $\alpha \pm \sigma_{\text{ind}} = -1.22 \pm 0.33$  is consistent within  $\sim 1\sigma_{\text{ind}}$  error with the estimated spectral index. It could be the case that the DSA is working for the particle acceleration at cluster shocks.



It is, on the other hand, already known that the acceleration efficiency at a weak shock is far too low to reproduce the observed radio luminosity (e.g., Kang & Ryu 2011; Kang et al. 2012; Pinzke et al. 2013; Vink & Yamazaki 2014). To explain this issue, several possibilities are proposed, which include (1) re-acceleration of pre-accelerated electrons (e.g., Markevitch et al. 2005; Kang et al. 2012; van Weeren et al. 2016), (2) shock drift accelerations (e.g., Matsukiyo et al. 2011; Guo et al. 2014b, 2014a), (3) other mechanisms, for instance turbulence accelerations (e.g., Fujita et al. 2015, 2016; Kang et al. 2017). All observational systematic on both X-ray and radio side is also the cause of the discrepancy between the observed value and the DSA theory (e.g., Akamatsu et al. 2017; Stroe et al. 2017; Hoang et al. 2018).

Indeed, there are seven compact sources in the region of Source U, with spectral indices of  $-1.5$  to  $-0.5$ , which are comparable to the typical value of radio jets (Sabater et al. 2019). Moreover, Source U is well aligned with the temperature jump (figure 2). Therefore, a possible origin of Source U would be the result of merger shock re-acceleration of pre-seeded cosmic-ray electrons by AGNs. Note that the spectral index map (figure 2b) shows patchy distribution and a less-clear global gradient due to the aging with respect to the shock propagation direction, implying multiple seeding from these AGN.

Moreover, our lack of detection of diffuse radio emission associated with the southern shock with the Mach number of 1.4 may indicate that there is a threshold of efficient (re-)acceleration at the shock age of  $\sim 50$  Myr (Kato et al. 2016). It is thus important to examine whether theoretical models of particle acceleration can explain this marginal detection and non-detection simultaneously. There are also compact radio sources in the other shock regions, and, for example, the spectral indices range from  $-1.5$  to  $-0.5$  for six compact sources at the eastern part of the south shock. Thus, there is also a possibility of seeding cosmic-ray electrons from AGN. It is necessary to clarify whether they are member galaxies associated with CIZA1359. Those are future works that will follow on from this paper.

Finally, we discuss which of the traditional classifications for cluster diffuse emission this fits into. Source U indicates that (I) it is found in between two subclusters of an early-stage merging cluster, (II) it has a structure along the shock, (III) it possesses a relatively flat spectral index of  $-1.22$ , and (IV) multiple non-bright radio point sources are located within the diffuse radio emission and there are no bright AGNs. These facts imply that Source U is a radio relic. Feature (I) is also seen in the radio bridges of early-stage merging galaxy clusters, such as Abell 399 and 401 (Govoni et al. 2019) and Abell 1758 (Botteon et al. 2020). However, these clusters also have radio halos and relics (Botteon et al. 2018), while CIZA1359 has no any other diffuse sources.

Radio relics tend to be brighter than radio bridges, so that we expect this applies to CIZA1359 as well. As no other diffuse emission has been detected in CIZA1359, it is more reasonable to consider Source U as a relic than a bridge.

Govoni et al. (2019) found the formation scenario of the radio bridges between Abell 399 and 401. That is, the contact of two clusters generates a shock and turbulence is excited at the post-shock region. Seed cosmic-ray electrons are re-accelerated through turbulence by the Fermi second-order acceleration mechanism. Such turbulence (re-)acceleration could be realized in CIZA1359 as well, though the short age of  $\sim 50$  Myr prefers the direct shock acceleration by the Fermi first-order acceleration mechanism. In other words, the radio relic candidate of CIZA1359 is a precursor of a radio bridge. Even in this case, a low Mach number such as  $\mathcal{M} = 1.7$  would require seeding of cosmic rays to achieve efficient acceleration and radio emission (e.g., Gu et al. 2019).

## 5 Summary

We reported the results on a SPAM-based analysis of uGMRT observations at 300–500 MHz for the early-stage merging galaxy cluster, CIZA J1358.9–4750 (CIZA1359). We found many radio sources such as a head–tail galaxies, FR II type radio lobes, AGNs, and so-called radio phoenixes or fossils. We found a diffuse radio source candidate named Source U with a flux density of  $24.04 \pm 2.48$  mJy roughly along a part of the shock front found in the previous X-ray observations.

We discussed whether Source U is real or noise from several aspects of its properties. First, the location of Source U is consistent with that of the shock front. Such an association is often seen in radio relics. Secondly, the size is comparable to known radio relics. Interestingly, the structure of Source U coincides with the shock structure where the Mach number of the shock wave reaches its maximum value of  $\mathcal{M} \sim 1.7$ . Thirdly, the relation between the radio power and the X-ray luminosity is in good agreement with that of other radio relics. Finally, the energy-equipartition magnetic-field strength,  $2.1 \mu\text{G}$ , is a typical value seen in galaxy clusters and relics. The above facts favor that Source U is a real radio relic.

If Source U is a real diffuse radio source, this study confirmed that even a very weak ( $\mathcal{M} \sim 1.7$ ) shock can accelerate cosmic-rays and emit observable radio emission. Moreover, we did not find any radio candidate at the shock with  $\mathcal{M} \sim 1.4$ , suggesting the existence of an acceleration-efficiency threshold around the Mach number. We suspect that seed cosmic rays were supplied by some compact radio sources (AGNs) associated with Source U and the re-acceleration took place at the shock.

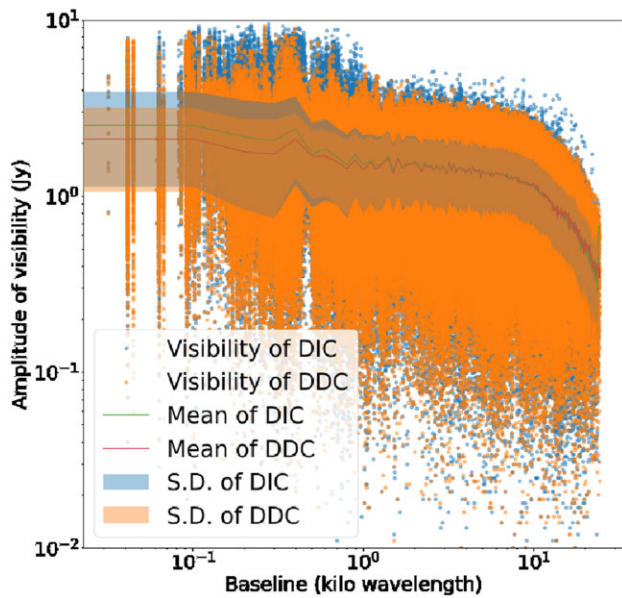


Fig. 5. Power spectrum of uGMRT data. Blue and orange points show the data from DIC and DDC, respectively. The solid lines show the data averaged over 0.1 kilo wavelength bins, and the shadows show their standard deviations.

It is important to identify the redshifts of the radio sources in order to elucidate the origin of the relic candidates. The identification is also necessary to examine whether head–tail galaxies seen in this observation interact with CIZA1359 or not. In addition, the relic candidate has a relatively steep spectrum. Therefore, future observations at lower frequencies such as 144 MHz would be promising to detect the candidate. Finally, ultimate sensitive observations with the Square Kilometre Array (SKA) will enable us to understand the details of radio sources as well as its polarization.

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## Appendix 1 Flux accuracy

We confirmed the amplitude of visibility with respect to the DDC and DIC results as described in subsection 3.1. Figure 5 shows the amplitude of visibility at each baseline length.

The DDC calibration employs data over 1 kilo wavelength. The amplitudes of DDC and DIC are different in each region of the figure 5. At scales below the most diffuse component (Source U), DDC has an amplitude that is about 6% lower than DIC.

We further compared our results to the flux of the TGSS (Intema et al. 2017) to test the validity of the 10% error in absolute flux. First, we convolved TGSS and our

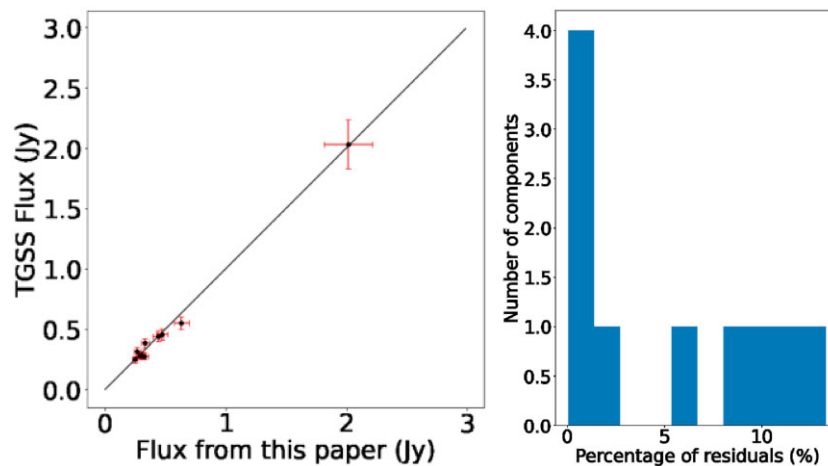


Fig. 6. Left (a): The red bars show a 10% error relative to flux. The black line is the expected line ( $y = x$ ). Right (b): The distance between the black line and the data divided by the flux from this paper, i.e., the percentage of error relative to flux.

results to the same resolution ( $65''$ ). Secondly, we performed Gaussian fitting on each component with  $SN > 20$  using pybdsf. Finally, we matched for each component by position using TOPCAT (Taylor 2005). The separation used for the matching was the same value as the resolution, and since TGSS is a 150 MHz band, it was re-scaled using  $\alpha = -0.8$ . The resulting flux relationship is shown in the figure 6. These results mean that the error in flux of this paper is usually well within 10%. The minimum flux in the figure 6 depends on the noise level of the TGSS.

## Appendix 2 Spectral index

An example of a spectral fit is shown in figure 7. We examine in this section whether the spectral index of Source U is due to noise behavior.

Figure 8a show the histogram of the spectral indices for the pixels within the Source U. We obtained the average spectral index of  $-1.23 \pm 0.27$ , the error means the standard deviation of the histogram. However, we need a careful assessment of the spectral index for marginal sources like Source U, because of the significant contribution of noise for the spectrum index fitting.

We have checked the spectral indices of all pixels in the image and found that they are distributed with three peaks at 0,  $-0.7$ , and  $-6.8$  as shown in figure 8b. The 0-centered peak is thought to be derived from spectral fitting by bad pixels, while the initial value of 0 remains as it is because the fitting does not converge; indeed the 0-centered peak disappears in the histograms of the pixels brighter than  $0.5 \text{ mJy beam}^{-1}$ , i.e., the pixels possessing high signal-to-noise ratio ( $SN \gtrsim 5$ ), which has a peak at  $-0.71 \pm 0.47$  (figure 8c). The other,  $-0.7$ - and  $-6.8$ -centered, peaks are thought to be derived from spectral fitting between noise floors in the data of uGMRT and MeerKAT.<sup>3</sup> The  $-0.7$ -centered peak is consistent with the values estimated from the noise distribution such as is shown in figure 7. The  $-6.8$ -centered peak occurs frequently when fitting negative and

<sup>3</sup> The image rms noise of uGMRT is larger than that of MeerKAT. When given error values of uGMRT and MeerKAT with the same sign, they make up an artificial negative spectral index. In our data, this artificial index is  $-0.7$ .

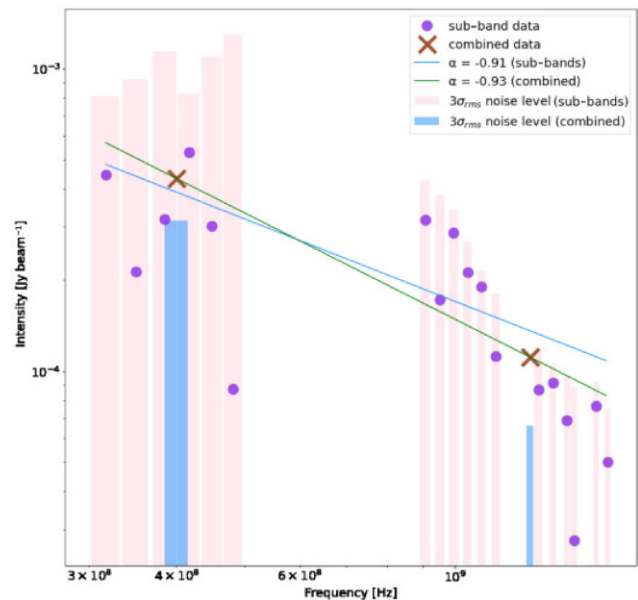
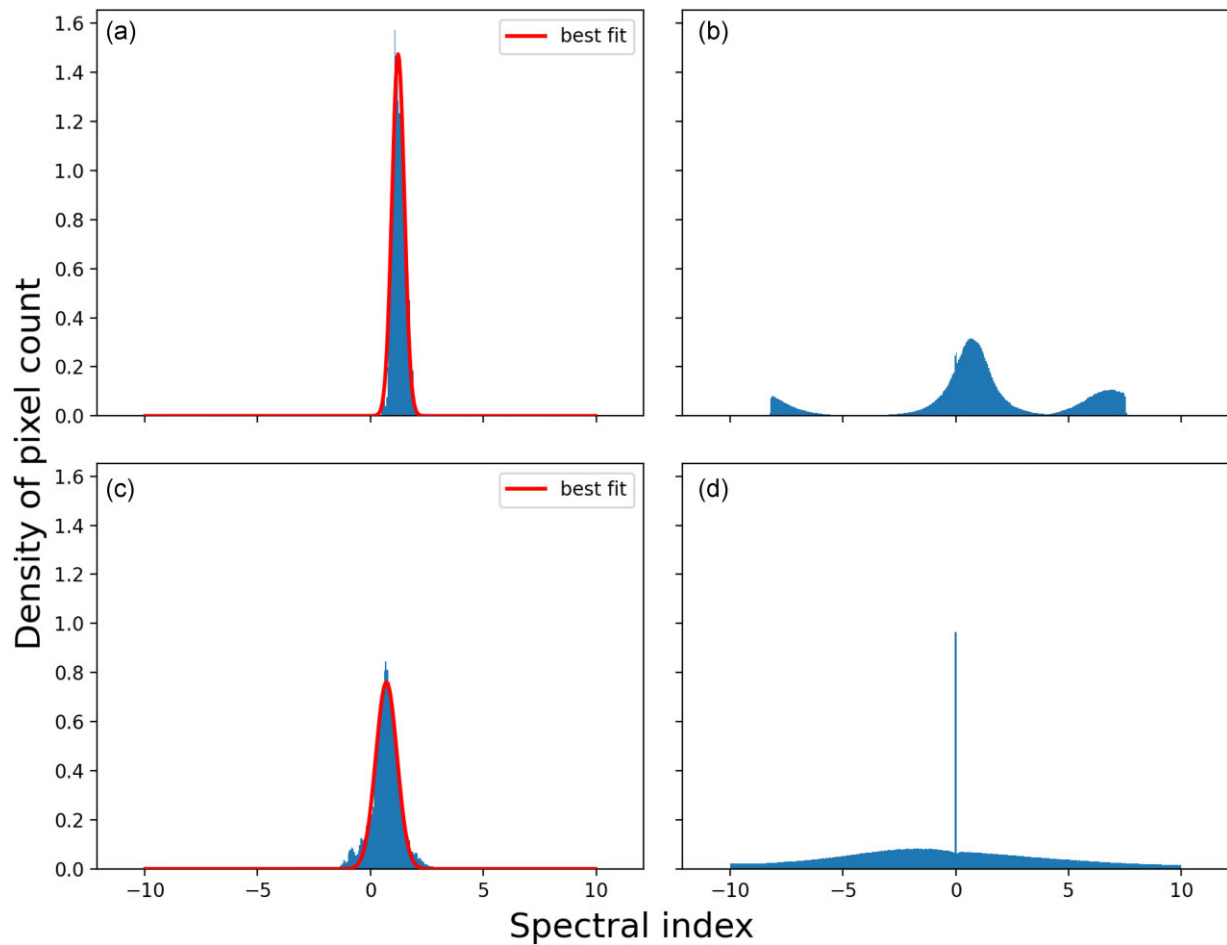


Fig. 7. Example of radio spectrum in Source U. The points marker shows the data for each frequency channel. The crosses are combined data from uGMRT and MeerKAT. The solid lines show the best-fitting power-law models, where the blue line considers all data for each frequency channel. The green line is the result of fitting using combined data from uGMRT and MeerKAT. The pink and blue bars indicate the  $3\sigma_{rms}$  of each channel and the combined data, respectively.

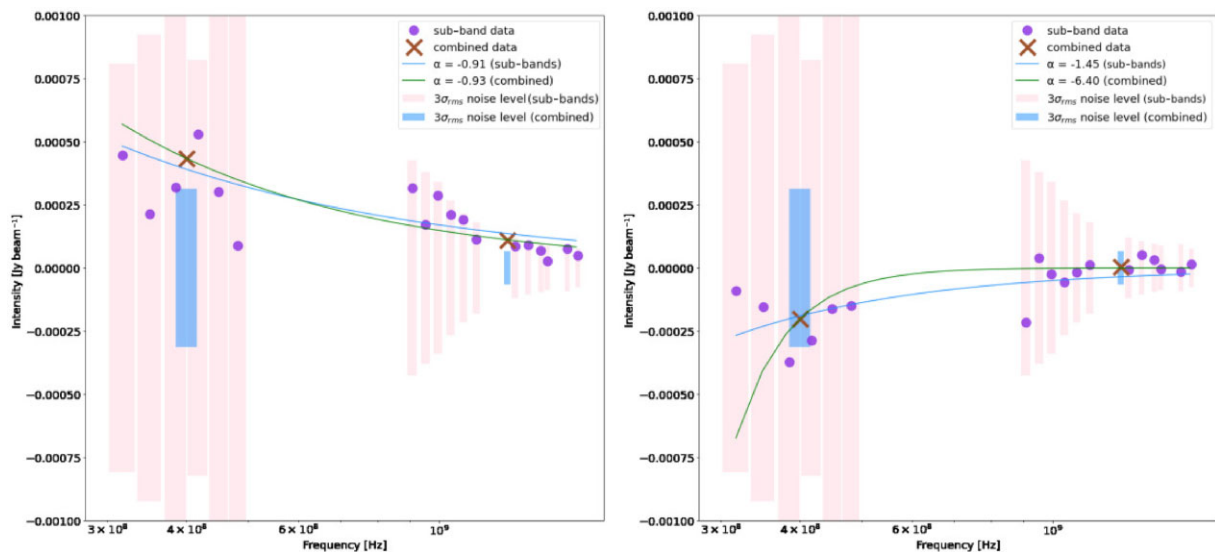
positive noise (see figure 9b). Indeed, the index histogram calculated from only uGMRT data (figure 8d) does not show the  $-0.7$ - and  $-6.8$ -centered peaks; it has no clear peak other than a 0-centered one.

We also show the spectra in linear space in figure 9. Panel (a) shows the spectrum of the same pixel as in figure 7, and panel (b) shows the spectrum of the noise pixel at the outer edge of the field of view. Panel (a) tends to have predominantly positive data compared to noise pixel [panel (b)] which suggests that Source U is different from the noise trend. Also, the green line in panel (b) shows that the spectral index is very steep because of the fitting with negative 400 MHz data and positive 1280 MHz data.

The above results support that it is more favorable to consider that Source U is a real radio source. The value  $-1.23 \pm 0.27$  is consistent with the spectral index of Source U inferred from the flux density.



**Fig. 8.** Histograms of spectral indices. This is normalized so that the total area is unity. The histograms were sampled between  $-10$  and  $10$  with even intervals of  $0.05$  in the spectral index axis. (a): Spectral index histogram of Source U. The blue bars show the data. The red line shows the best Gaussian fit. (b): Spectral index histogram of all the pixels within the image. The data of uGMRT and MeerKAT are both taken for the power-law fit. (c): Same as panel (b) but for the pixels with an intensity greater than  $0.4 \text{ mJy beam}^{-1}$  at the smoothed uGMRT image which correspond to the white contours in figure 1. (d): Same as panel (b) but the spectral index was calculated only with uGMRT data.



**Fig. 9.** Left: Same as figure 7, but the vertical axis is linear. Right: Same as the left-hand panel, but it is the spectrum of pixels where noise is dominant.



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