



Universiteit
Leiden
The Netherlands

Now you see it, now you don't: star formation truncation precedes the loss of molecular gas by 100 Myr in massive poststarburst galaxies at $z \sim 0.6$

Bezanson, R.; Spilker, J.S.; Suess, K.A.; Setton, D.J.; Feldmann, R.; Greene, J.E.; ... ; Verrico, M.

Citation

Bezanson, R., Spilker, J. S., Suess, K. A., Setton, D. J., Feldmann, R., Greene, J. E., ... Verrico, M. (2022). Now you see it, now you don't: star formation truncation precedes the loss of molecular gas by 100 Myr in massive poststarburst galaxies at $z \sim 0.6$. *The Astrophysical Journal*, 925(2). doi:10.3847/1538-4357/ac3dfa

Version: Publisher's Version
License: [Creative Commons CC BY 4.0 license](#)
Downloaded from: <https://hdl.handle.net/1887/3562536>

Note: To cite this publication please use the final published version (if applicable).



Now You See It, Now You Don't: Star Formation Truncation Precedes the Loss of Molecular Gas by ~ 100 Myr in Massive Poststarburst Galaxies at $z \sim 0.6$

Rachel Bezanson¹ , Justin S. Spilker^{2,3,13} , Katherine A. Suess^{4,5,6} , David J. Setton¹ , Robert Feldmann⁷ ,
Jenny E. Greene⁸ , Mariska Kriek^{4,9} , Desika Narayanan^{10,11,12} , and Margaret Verrico¹

¹ Department of Physics and Astronomy and PITT PACS, University of Pittsburgh, Pittsburgh, PA 15260, USA; rachel.bezanson@pitt.edu

² Department of Astronomy, University of Texas at Austin, 2515 Speedway, Stop C1400, Austin, TX 78712, USA

³ Department of Physics and Astronomy and George P. and Cynthia Woods Mitchell Institute for Fundamental Physics and Astronomy, Texas A&M University, 4242 TAMU, College Station, TX 77843-4242, USA

⁴ Astronomy Department, University of California, Berkeley, CA 94720, USA

⁵ Department of Astronomy and Astrophysics, University of California, Santa Cruz, 1156 High Street, Santa Cruz, CA 95064 USA

⁶ Kavli Institute for Particle Astrophysics and Cosmology and Department of Physics, Stanford University, Stanford, CA 94305, USA

⁷ Institute for Computational Science, University of Zurich, CH-8057 Zurich, Switzerland

⁸ Department of Astrophysical Sciences, Princeton University, Princeton, NJ 08544, USA

⁹ Leiden Observatory, Leiden University, P.O.Box 9513, NL-2300 AA Leiden, The Netherlands

¹⁰ Department of Astronomy, University of Florida, 211 Bryant Space Science Center, Gainesville, FL 32611, USA

¹¹ University of Florida Informatics Institute, 432 Newell Drive, CISE Bldg E251 Gainesville, FL 32611, US

¹² Cosmic Dawn Centre at the Niels Bohr Institute, University of Copenhagen and DTU-Space, Technical University of Denmark, Denmark

Received 2021 October 4; revised 2021 November 17; accepted 2021 November 26; published 2022 February 2

Abstract

We use ALMA observations of CO(2–1) in 13 massive ($M_* \gtrsim 10^{11} M_\odot$) poststarburst galaxies at $z \sim 0.6$ to constrain the molecular gas content in galaxies shortly after they quench their major star-forming episode. The poststarburst galaxies in this study are selected from the Sloan Digital Sky Survey spectroscopic samples (Data Release 14) based on their spectral shapes, as part of the Studying QUenching at Intermediate- z Galaxies: Gas, angular momentum, and Evolution (*SQUIGGLE*) program. Early results showed that two poststarburst galaxies host large H_2 reservoirs despite their low inferred star formation rates (SFRs). Here we expand this analysis to a larger statistical sample of 13 galaxies. Six of the primary targets (45%) are detected, with $M_{H_2} \gtrsim 10^9 M_\odot$. Given their high stellar masses, this mass limit corresponds to an average gas fraction of $\langle f_{H_2} \equiv M_{H_2}/M_* \rangle \sim 7\%$ or $\sim 14\%$ using lower stellar masses estimates derived from analytic, exponentially declining star formation histories. The gas fraction correlates with the D_n4000 spectral index, suggesting that the cold gas reservoirs decrease with time since burst, as found in local K+A galaxies. Star formation histories derived from flexible stellar population synthesis modeling support this empirical finding: galaxies that quenched $\lesssim 150$ Myr prior to observation host detectable CO(2–1) emission, while older poststarburst galaxies are undetected. The large H_2 reservoirs and low SFRs in the sample imply that the quenching of star formation precedes the disappearance of the cold gas reservoirs. However, within the following 100–200 Myr, the *SQUIGGLE* galaxies require the additional and efficient heating or removal of cold gas to bring their low SFRs in line with standard H_2 scaling relations.

Unified Astronomy Thesaurus concepts: Post-starburst galaxies (2176); Galaxy quenching (2040); Galaxy evolution (594); Quenched galaxies (2016); Galaxies (573); High-redshift galaxies (734)

1. Introduction

The process by which star-forming massive disk galaxies shut off their ongoing star formation and join the older population of quiescent elliptical galaxies is one of the most poorly understood aspects of galaxy evolution. The dominance of ellipticals and the relative inefficiency of star formation as measured by abundance matching of the dark matter halo and galaxy mass functions above M^* has led to the need to introduce additional “feedback” into the galaxy formation process. At the massive end, this feedback is generally attributed to active galactic nuclei (AGNs; e.g., Croton et al. 2006). Most modern cosmological simulations that form realistic populations of massive galaxies introduce some mode of energy injection that is attributed to supermassive black holes

(e.g., Crain et al. 2015; Schaye et al. 2015; Anglés-Alcázar et al. 2017; Weinberger et al. 2017; Pillepich et al. 2018; Davé et al. 2019; Rodríguez Montero et al. 2019), but in other cases, the inefficiency to accrete and replenish the cold gas supplies of the most massive systems is tied to the halo mass (e.g., Feldmann & Mayer 2015; Davé et al. 2017; Feldmann et al. 2017). In all cases, the link between cold H_2 and ongoing star formation is a built-in assumption. These expectations are empirically well motivated at large scales by the relatively gas-rich nature of galaxies with ongoing star formation (e.g., Saintonge et al. 2011a, 2011b, 2012; Tacconi et al. 2013, 2018) and at small scales by the strong correlation between the surface density of active star formation and the density of molecular hydrogen (e.g., Kennicutt 1998; Schrubba et al. 2011). Although H_2 reservoirs in dynamically hot elliptical galaxies appear to be less efficient at fueling their low-level star formation, this effect is secondary; in the local universe quiescent galaxies are extremely gas depleted (e.g., Davis et al. 2011; Young et al. 2011; Davis et al. 2013).

Empirical studies suggest that the majority of massive elliptical galaxies formed their stars in early, short-lived

¹³ NHFP Hubble Fellow.



episodes—indicating the importance of a rapid mode of quenching (e.g., Thomas et al. 2005; Pacifici et al. 2016; Tacchella et al. 2021). It is therefore interesting to investigate the properties of poststarburst galaxies, sometimes referred to as E+A or K+A galaxies, which are selected to be the direct products of a fast track of quenching that shuts off a dramatic episode of star formation within $\lesssim 1$ Gyr (Dressler & Gunn 1983; Zabludoff et al. 1996). This truncation produces characteristic spectral signatures originating from A stars¹⁴ that dominate after more massive stars have died combined with a lack of signatures of instantaneous star formation (e.g., [O II] or H α emission lines). If indeed these galaxies are in transition and star formation has already been shut down, a reasonable expectation would be that their gas reservoirs already resemble those of older quiescent galaxies.

In this context, it is surprising that poststarburst galaxies have been demonstrated to host enigmatic and significant H₂ reservoirs ($f_{\text{H}_2} \equiv M_{\text{H}_2}/M_*$ up to $\sim 50\%$) even after their star formation stops abruptly. The majority of studies of H₂ in poststarburst galaxies have been limited to the local universe, where such galaxies are extremely rare, but detailed studies are relatively accessible (e.g., Alatalo 2015; French et al. 2015; Rowlands et al. 2015; Alatalo et al. 2016). Given that the most massive galaxies have the oldest stellar populations and therefore quenched at the earliest epochs (e.g., Thomas et al. 2005; McDermid et al. 2015), these low-redshift relics of the late-time quenching process do not necessarily reflect the processes that shut off star formation at early times. Intriguingly, H₂ in local poststarburst galaxies has been demonstrated to disappear on short (~ 100 Myr) timescales (French et al. 2018). However, no similar tests have been conducted beyond the local universe.

Pushing observations of the possible link—or lack thereof—between quenching and the availability of H₂ to an epoch when galaxies are shutting down their primary episodes of star formation is a critical test of the theoretical models that could form the most massive galaxies in the universe. Only a handful of quiescent galaxies at $z > 0.1$ have been targeted for H₂ using CO lines, yielding only four detections and/but deep nondetections of (older) quiescent galaxies at “cosmic noon” ($1 \lesssim z \lesssim 2$) (Sargent et al. 2015; Spilker et al. 2018; Bezanson et al. 2019; Belli et al. 2021; Williams et al. 2021). In a pilot study of two massive poststarburst galaxies at $z \sim 0.6$ that were spectroscopically selected to have recently shut down a major episode of star formation, Suess et al. (2017) reported both retained significant molecular gas reservoirs ($f_{\text{H}_2} \sim 10\%–30\%$).

In this paper we expand upon that work with a larger sample of 13 galaxies from the Studying QUenching at Intermediate- z Galaxies: Gas, angular momentum, and Evolution (*SQuIGGLE*) survey. This enables an investigation of trends between H₂ and the spectroscopic properties of massive poststarburst galaxies at $z \sim 0.6$, immediately after they quench their dominant episode of star formation. In Section 2, we briefly describe the *SQuIGGLE* sample and the ALMA CO(2–1) observations. Section 3 discusses the H₂ reservoirs implied by the CO(2–1) fluxes and puts the sample in the context of other studies of star-forming and poststarburst galaxies. Finally Section 4 provides a discussion of

the results and avenues for future study. Throughout this paper we assume a concordance Λ CDM cosmology with $\Omega_\Lambda = 0.7$, $\Omega_m = 0.3$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, a Chabrier (2003) initial mass function, and quote AB magnitudes.

2. Data

2.1. The *SQuIGGLE* Sample

The *SQuIGGLE* sample is selected from the Sloan Digital Sky Survey (SDSS) Data Release 14 (DR14) spectroscopic database (Abolfathi et al. 2018) to have strong Balmer breaks and blue slopes redward of the break using rest-frame filters, following Kriek et al. (2010). For a detailed description of the spectroscopic identification and stellar populations of *SQuIGGLE* galaxies, we refer the reader to the survey paper by Suess et al. (2021). In summary, for all objects in the database with $z > 0.5$, each spectrum is integrated within three medium-width synthetic rest-frame filters (U_m , B_m , and V_m), which span the Balmer/4000 Å break and the spectral shape just redward of the break. We only include galaxies with signal-to-noise ratio (S/N) > 6 in the B_m and V_m fluxes and colors similar to those of A-type stars ($U_m - B_m > 0.975$ and $-0.25 < B_m - V_m < 0.45$). This selection yields a sample of 1318 unique objects, which span $0.50 < z < 0.94$. Figure 1 shows the distribution of the full *SQuIGGLE* sample in i magnitude (left) and stellar mass (right) versus redshift.

For this work, we rely on two sets of stellar population synthesis modeling, which we use in different contexts. The first set of models uses *Prospector* (Johnson & Leja 2017; Leja et al. 2017; Johnson et al. 2021) to fit the SDSS *ugriz* and WISE ($3.4 \mu\text{m}$ and $4.5 \mu\text{m}$) photometry and spectra (Abolfathi et al. 2018; Schlafly et al. 2019) with a custom set of “nonparametric” star formation histories (SFHs), assuming a Kriek & Conroy (2013) dust law. These SFHs are similar to default continuity prior nonparametric models from Leja et al. (2019), with three fixed duration, variable star formation rate (SFR) bins prior to 2 Gyr of lookback time, five equal-mass bins with variable edges that follow, and a unique final bin with flexibility in timing and SFR normalization. Without this final bin, the default continuity assumptions could be too strict to allow for dramatic bursts or quick truncation in star formation as one might expect for poststarburst galaxies, effectively blurring out SFHs and biasing instantaneous SFRs to higher values. While conducting extensive recovery testing of stellar population properties using *Prospector*, we found that our adopted flexible SFHs provide excellent recovery of instantaneous SFR (interpolated over the last 1 Myr) for sufficient SFR ($\gtrsim 1 M_\odot \text{ yr}^{-1}$), below which the measured star formation rates were poorly constrained by the existing spectrophotometric data set (Suess et al. 2021). Therefore, when making comparisons to scaling relations, we set the SFRs to a floor value of $1 M_\odot \text{ yr}^{-1}$ and label those points as upper limits.

We note that the low SFRs are consistent with follow-up Keck/LRIS spectroscopy targeting H α (Suess et al. 2021). Although SFRs derived from H α luminosity are less uncertain than, e.g., the [O II] luminosity-based SFRs used in Suess et al. (2017), due to dust and other contaminating ionizing sources, H α -based SFRs are insensitive to heavily dust-obscured star formation. We see no strong evidence for such extreme obscuration e.g., in the 2 mm continuum data presented in this paper but will return to this in Section 4.

¹⁴ We note that strong Balmer absorption features are also apparent in BV and later-type stars, but adopt the convention of referring to these signatures as A-type signatures. This becomes important for some galaxies in the *SQuIGGLE* survey for which inferred postquenching ages are shorter than the lifetimes of A stars.

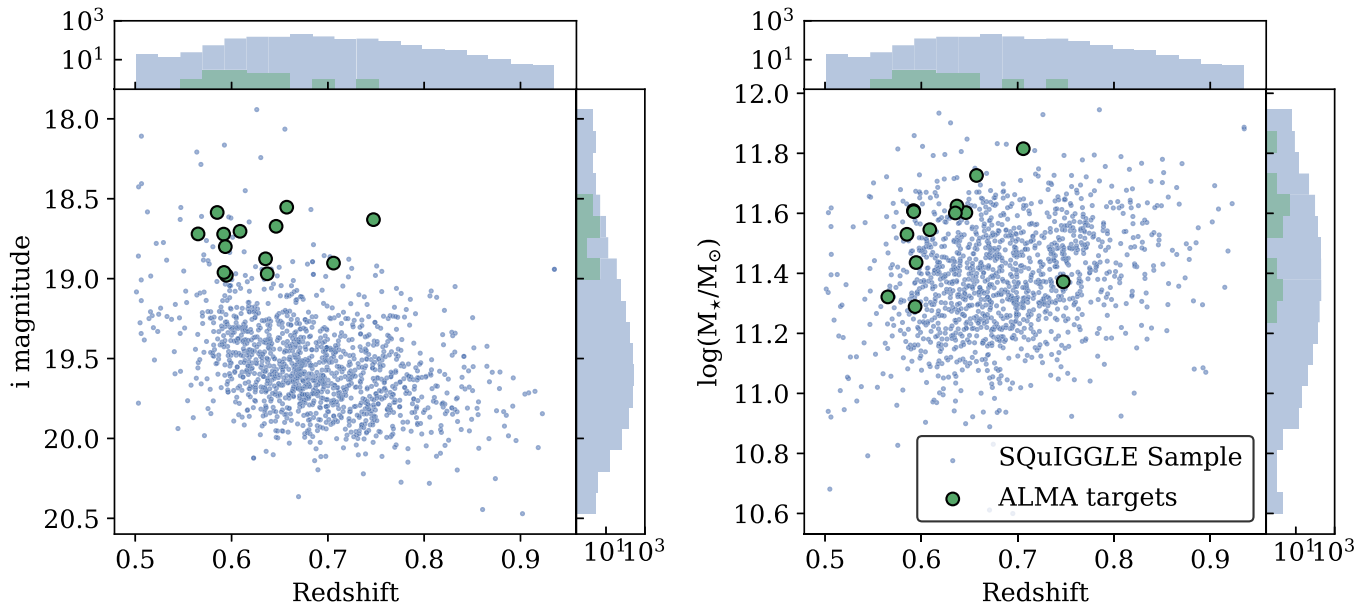


Figure 1. The distribution of *SQuIGGLE* massive poststarburst galaxies in *i* magnitude (left) and stellar mass (right) vs. spectroscopic redshift. In all panels the full sample is indicated by the blue symbols and histograms, and the 13 galaxies targeted by our ALMA observations are highlighted in green. The ALMA targets are skewed toward slightly lower redshifts than the full sample ($z < 0.75$) to catch the CO(2–1) line in ALMA band 4. ALMA targets were selected to have higher stellar masses and brighter magnitudes to allow for efficient multiwavelength follow-up.

This Prospector spectral energy distribution (SED) modeling is designed to accurately recover SFHs and SFRs, particularly immediately before quenching. However, often these histories are more extended than a more standard exponentially declining or delayed exponential analytic SFH, yielding significantly higher stellar masses. This 0.1–0.3 dex offset is a generic consequence of “nonparametric” SED modeling (see, e.g., Leja et al. 2019; Lower et al. 2020) and is perhaps more extreme for the poststarburst galaxies in this sample. While we expect that the higher stellar masses likely reflect the intrinsic properties of the galaxies, we also fit the SDSS spectra and photometry with delayed exponential SFHs (as described in Setton et al. 2020) assuming similar Chabrier (2003) IMF, Bruzual & Charlot (2003) libraries, and a Calzetti (1997) dust law using FAST++, an implementation of the FAST (Fitting and Assessment of Synthetic Templates) software (Kriek et al. 2009). The stellar masses derived from these fits ($M_{*,\text{FAST}}$) are an average of 0.38 dex lower than the stellar masses derived in the default fits. We plan to expand upon these differences in an upcoming paper (Suess et al. 2021) but in this work adopt the lower values to place *SQuIGGLE* galaxies on scaling relations for consistency. Regardless of the technique used to calculate stellar masses, galaxies in the *SQuIGGLE* Survey are generally bright and massive ($\langle \log M_*/M_\odot \rangle = 11.4$), which is primarily driven by the spectroscopic signal-to-noise cut.

Although these selection criteria are designed to identify galaxies immediately following the truncation of a significant episode of star formation based on the shape of their SEDs, it almost exclusively includes galaxies that would also be designated as poststarbursts based on strong $H\delta$ absorption (98% of the sample has Lick $H\delta_A \geq 4.0\text{\AA}$). For a comprehensive review of the range of poststarburst galaxy identification methods, we refer the reader to French (2021). Figure 2 shows the *SQuIGGLE* sample as small blue points in Lick $H\delta_A$ versus D_n4000 parameter space, which is commonly used to study the demographics of broad galaxy populations (see, e.g., Kauffmann et al. 2003). All indications from weak [O II] emission and full spectral modeling indicate that galaxies in the *SQuIGGLE* sample would pass cuts

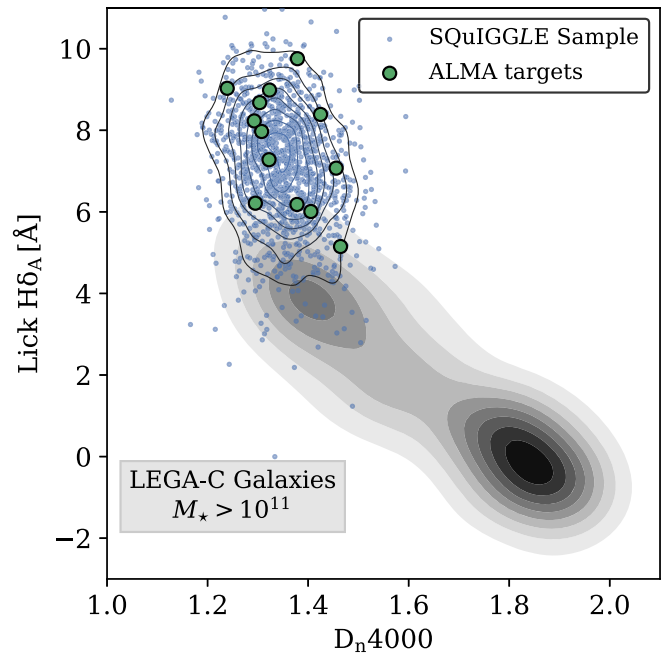


Figure 2. The distribution of the *SQuIGGLE* sample in Lick $H\delta_A$ and D_n4000 spectral indices. Blue symbols indicate the full data set and green circles highlight those targeted by our ALMA CO(2–1) survey. The distribution of the population of similarly massive ($\log M_*/M_\odot > 11$) galaxies at $0.6 < z < 0.8$ from the LEGA-C survey is indicated by the gray shaded region. Although the selection of poststarburst galaxies in *SQuIGGLE* is based on their broad spectral shapes, the sample exhibits strong $H\delta_A$ absorption features and very blue D_n4000 spectral indices relative to the full population of massive galaxies at similar redshifts. The galaxies targeted in our ALMA survey span the range of $H\delta_A$ and D_n4000 parameter space spanned by the full *SQuIGGLE* survey.

designed to identify objects without significant ongoing star formation, but the traditional $H\alpha$ SFR indicator is redshifted out of the spectral wavelength coverage for the full data set. For comparison, we show the distribution of similarly massive ($\log M_*/M_\odot > 11$) galaxies at a similar epoch ($0.6 < z < 0.8$)

Table 1
Properties of the ALMA Observations

ID	R.A. (Degrees)	Decl. (Degrees)	ALMA Program	Observation Date	Integration Time (s)	Angular Resolution ($''$)
SDSS_J0912+1523	138.17821	15.38479	2016.1.01126.S	2017-01-08	5866.56	1.57
SDSS_J2202-0033	330.60121	-0.55955	2016.1.01126.S	2017-03-07	5685.12	2.00
SDSS_J1448+1010	222.19133	10.16960	2017.1.01109.S	2018-03-12	5999.616	0.74
SDSS_J0753+2403	118.43406	24.06005	2017.1.01109.S	2018-03-20	5987.52	0.72
SDSS_J1203+1807	180.98548	18.13016	2017.1.01109.S	2018-03-14	5927.04	0.70
SDSS_J1007+2330	151.80432	23.51530	2017.1.01109.S	2018-03-15	5927.04	0.75
SDSS_J1053+2342	163.44737	23.70956	2017.1.01109.S	2018-03-16	5987.52	0.77
SDSS_J0233+0052	38.49722	0.87734	2017.1.01109.S	2018-04-04	5987.52	1.35
SDSS_J1302+1043	195.70387	10.71748	2017.1.01109.S	2018-03-21	5987.52	0.67
SDSS_J1109-0040	167.38393	-0.66774	2017.1.01109.S	2018-04-10	5927.04	1.11
SDSS_J0046-0147	11.66247	-1.78856	2017.1.01109.S	2018-04-15	5987.52	1.02
SDSS_J0027+0129 ^a	6.85600	1.49942	2017.1.01109.S	2018-04-10	5927.04	1.11
SDSS_J2258+2313	344.52365	23.22115	2017.1.01109.S	2018-05-01	5987.52	1.17

Note.

^a This galaxy was also observed as part of 2016.1.01126.S, but those observations did not pass quality assessment (QA).

from the third data release of the Large Early Galaxy Astrophysics Census (LEGA-C; van der Wel et al. 2016; Straatman et al. 2017; van der Wel et al. 2021) in gray contours. As demonstrated in Wu et al. (2018), massive galaxies at this epoch tend to have relatively old stellar populations as evidenced by weak Balmer absorption and strong 4000Å breaks, corresponding to high D_n4000 spectral indices. This is in stark contrast with the recently quenched poststarburst galaxies in the *SQuIGGLE* survey; the LEGA-C galaxy distribution only barely reaches the low end of the *SQuIGGLE* range in $H\delta_A$ indices. Although these samples probe similar redshifts, this lack of overlap is primarily an effect of survey volume. The ~ 1.6 deg² section of the COSMOS field that was targeted by LEGA-C is relatively large for an extragalactic field, but it probes a vastly smaller volume than the BOSS Survey (10,000 deg²; Dawson et al. 2013).

2.2. ALMA CO(2–1) Observations

We have targeted a subset of the *SQuIGGLE* sample for extensive multiwavelength follow-up studies, preliminary results of which have been published in several articles. In Suess et al. (2017) we published ALMA CO(2–1)-based detections of vast H_2 reservoirs ($\log M_{H_2}/M_\odot \gtrsim 10.0$, or $f_{H_2} \sim 20\%$ and 4%)¹⁵ in two galaxies. These results suggest that the cold molecular gas is common in massive, recently quenched galaxies at $z \sim 0.6$, but are far from conclusive. In the current paper, we present the demographics of H_2 (as probed by CO(2–1)) in a larger sample of 13 galaxies, adding 11 galaxies to those presented in Suess et al. (2017). The targeted subset is indicated by the green circles and histograms in Figures 1 and 2. We note specifically that the subset of galaxies selected for ALMA follow-up (green points) span the range of stellar populations of the full *SQuIGGLE* sample in this parameter space. Properties of the observations including program numbers, observation dates, integration times, and spatial resolution of the data are included in Table 1 and physical properties of the sample are also enumerated in Table 2. Because we chose this subsample for follow-up observations, our selection is biased toward brighter galaxies (in the i band)

that can be observed by ALMA in the southern hemisphere. These targets are brighter due to a combination of slightly lower redshifts and higher masses than the full *SQuIGGLE* selection. However, we note that the spectral diversity of the full *SQuIGGLE* data set, e.g., as probed by $H\delta_A$ and D_n4000 in Figure 2, is well sampled by the ALMA targets.

Following the strategy of our pilot study (Suess et al. 2017, ALMA Program #2016.1.01126.S, PI: Bezanson), we targeted CO(2–1) in 11 additional galaxies using ALMA band 4 in Cycle 5 (PI: Bezanson, Program# 2017.1.01109.S). All targets were observed in two 80 minute observing blocks with the full 12 m array. Total on-source integration times were ~ 1.7 hr, and the angular resolution of the resulting data cubes ranges from $0''.7$ to $2''$ (see Table 1).

Continuum images at ~ 2 mm were created using the full bandwidth of the ALMA data, excluding ± 500 km s⁻¹ around the expected frequency of the redshifted CO(2–1) line. The data reach continuum sensitivities ~ 6 – 10 μ Jy. Continuum emission was detected in one target, SDSS_J0753+2403; the flux ratio between the upper and lower sidebands of the ALMA data and a detection at 1.4 GHz (Greene et al. 2020) make clear that this is nonthermal synchrotron emission from AGN activity. Following Suess et al. (2017), the nondetections of the other sources imply upper limits on the presence of any highly obscured star formation, $SFR \lesssim 50 M_\odot/\text{yr}$, assuming a standard modified blackbody function and dust temperature $T_{\text{dust}} = 30$ K. These limits are highly uncertain due to the unknown dust temperature because the observed-frame 2 mm is far from the peak of the dust SED but are sufficient to rule out very highly obscured SFRs.

All CO(2–1) spectra are extracted by fitting circular Gaussian models to the visibilities using *uvmultifit* (Martí-Vidal et al. 2014). We generate spectra at effective velocity resolutions of ≈ 50 , 100, and 200 km s⁻¹. In the sources where CO is clearly detected, we often find evidence of velocity gradients and/or marginally resolved source sizes in individual channels. We allow the position and size of the spatial Gaussian components to be free parameters across the line profiles and fix them to the phase center and the median size for line-free channels. In undetected sources, we simply fix the centroid to the phase center and the size to $2''$ in all cases, a value typical of the detected sources.

¹⁵ We note that in Suess et al. (2017) we used delayed tau SFHs; the stellar masses based on flexible SFHs yield higher M_* values and therefore lower f_{H_2} measurements of $\sim 14.5\%$ and $\sim 1\%$ for the same galaxies.

Table 2
ALMA Target Sample Properties

Galaxy ID	Plate-MJD-Fiber	z_{spec}	$\log M_*/M_\odot$	SFR ($M_\odot \text{ yr}^{-1}$)	Lick $H\delta_A$ ($\text{\AA}$)	D_n4000	SdvCO(2-1) (Jy kms $^{-1}$)	$L'_{\text{CO}(2-1)}$ ($10^9 \text{ K kms}^{-1} \text{ pc}^2$)	R_{CO} (kpc)	$\log M_{\text{H}_2}/M_\odot$
SDSS_J1448+1010	5475-56011-379	0.6462	$11.60^{+0.04}_{-0.07}$	$1.06^{+0.99}_{-0.94}$	7.28 ± 0.34	1.32 ± 0.02	0.82 ± 0.05	4.82 ± 0.30	1.5 ± 0.6	10.29 ± 0.03
SDSS_J0753+2403	4466-55857-198	0.5652	$11.32^{+0.03}_{-0.05}$	$0.10^{+0.25}_{-0.09}$	8.99 ± 0.21	1.32 ± 0.01	<0.09	<0.42	...	<9.22
SDSS_J1053+2342	6417-56308-55	0.6370	$11.62^{+0.03}_{-0.08}$	$0.29^{+0.67}_{-0.28}$	6.01 ± 0.44	1.41 ± 0.02	<0.15	<0.86	...	<9.54
SDSS_J0027+0129	4302-55531-482	0.5851	$11.53^{+0.03}_{-0.03}$	$1.44^{+0.66}_{-1.15}$	8.23 ± 0.34	1.29 ± 0.02	<0.12	<0.59	...	<9.38
SDSS_J2202-0033	1105-52937-311	0.6573	$11.73^{+0.03}_{-0.02}$	$1.99^{+1.91}_{-1.70}$	8.68 ± 0.27	1.30 ± 0.01	0.27 ± 0.03	1.62 ± 0.21	<8.7	9.81 ± 0.06
SDSS_J2258+2313	6308-56215-977	0.7058	$11.82^{+0.03}_{-0.05}$	$0.94^{+1.85}_{-0.71}$	6.21 ± 0.53	1.30 ± 0.02	1.86 ± 0.08	13.10 ± 0.56	12.4 ± 0.9	10.72 ± 0.02
SDSS_J0233+0052	705-52200-614	0.5918	$11.61^{+0.02}_{-0.04}$	$0.01^{+0.22}_{-0.01}$	8.39 ± 0.41	1.42 ± 0.02	<0.11	<0.56	...	<9.35
SDSS_J0046-0147	4370-55534-762	0.6088	$11.55^{+0.03}_{-0.03}$	$0.14^{+1.11}_{-0.14}$	9.76 ± 0.30	1.38 ± 0.02	<0.12	<0.61	...	<9.39
SDSS_J1109-0040	278-51900-193	0.5935	$11.29^{+0.09}_{-0.03}$	$2.33^{+1.12}_{-1.62}$	7.97 ± 0.36	1.31 ± 0.02	0.78 ± 0.06	3.85 ± 0.30	5.7 ± 0.9	10.19 ± 0.03
SDSS_J1203+1807	2595-54207-459	0.5946	$11.44^{+0.02}_{-0.05}$	$0.02^{+0.19}_{-0.02}$	7.07 ± 0.42	1.46 ± 0.02	<0.09	<0.45	...	<9.26
SDSS_J1007+2330	6458-56274-501	0.6353	$11.60^{+0.03}_{-0.03}$	$0.89^{+0.99}_{-0.86}$	5.15 ± 0.54	1.46 ± 0.03	<0.15	<0.85	...	<9.53
SDSS_J0912+1523	2438-54056-396	0.7473	$11.37^{+0.03}_{-0.02}$	$0.81^{+1.33}_{-0.76}$	9.03 ± 0.33	1.24 ± 0.01	1.07 ± 0.05	8.47 ± 0.36	6.2 ± 0.8	10.53 ± 0.02
SDSS_J1302+1043 ^a	5421-55980-512	0.5921	$11.61^{+0.04}_{-0.06}$	$0.26^{+0.93}_{-0.26}$	6.18 ± 0.93	1.38 ± 0.04	0.78 ± 0.09	3.85 ± 0.42	3.3 ± 0.7	10.19 ± 0.05

Note. Upper limits for the undetected sources are 3σ and assume an 800 km s^{-1} line width.

^a This target was included in an early generation of the *SQUIGGLE* sample based on the SDSS DR12 spectra, but the spectral shape differs slightly in the DR14 spectrum, causing the rest-frame colors to fall out of the *SQUIGGLE* selection cuts. We include this galaxy in all further analysis.

^b The *Prospector* modeling for this galaxy is unable to reproduce the WISE $4.5 \mu\text{m}$ flux, driving the fit against the edges of the priors for dust. We expect that this is due to an asymptotic giant branch star or neighboring light that is not accounted for in the model parameter space. Therefore, we exclude WISE band 2 for this object and the fit converges well.

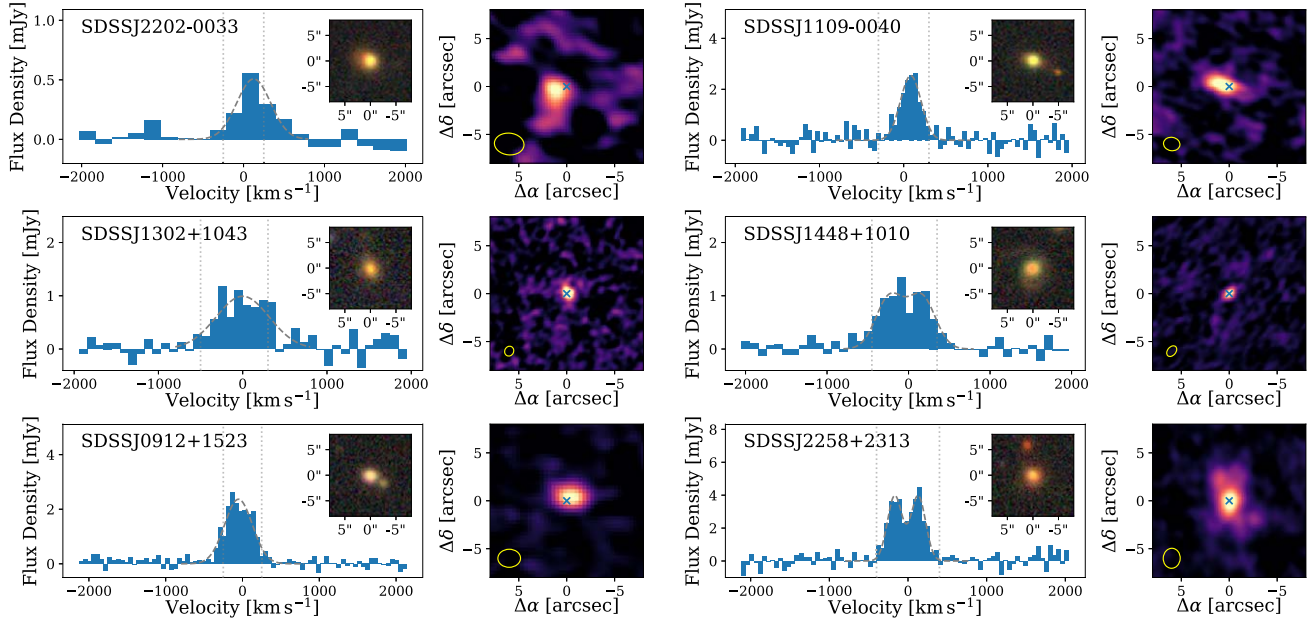


Figure 3. Spatially integrated CO(2–1) spectra (left) and line maps (right) for *SQuIGGLE* galaxies that are detected in our ALMA band 4 follow-up study. Gaussian spectral fits are included as dashed lines, and line maps are collapsed between vertical dashed lines. Galaxies are sorted from low to high $L_{\text{CO}(2-1)}$. Inset panels show optical *grz* images from the DESI Legacy Survey DR9 (Dey et al. 2019).

Integrated line fluxes are estimated by fitting extracted spectra, using one (spectral) Gaussian (SDSS_J2202-0033, SDSS_J1109-0400, SDSS_J13202+1043, and SDSS_J0912+1523) or two Gaussians (SDSS_J1488-1010 and SDSS_J2258-2312) to the CO(2–1) line profiles. For undetected galaxies, the upper limits on the CO(2–1) line flux are conservatively estimated using a single 800 km s^{-1} wide channel; upper limits on the line fluxes for these sources scale as $\sqrt{\Delta V/800 \text{ km s}^{-1}}$ for alternative choices of the velocity width ΔV . We stack the integrated spectra of the seven galaxies that are individually undetected using the 200 km s^{-1} spectra, finding a $\approx 3.7\sigma$ detection of CO(2–1) in the stacked spectrum. We verify this detection with a simple image-plane stack of the CO data cubes, but stress that this is merely for visualization purposes due to the varying spatial resolution of the input data cubes. We estimate the spatial extent of CO(2–1) emission in each detected galaxy by fitting 2D Gaussians in the image plane. Circularized half-width half-maximum values are quoted in Table 2 in physical units (kiloparsec).

Of the 13 targeted galaxies, CO(2–1) emission was detected in 6. The CO(2–1) spatially integrated spectra and line maps for the detected *SQuIGGLE* galaxies are included in Figure 3 and undetected galaxies are included in Figure 4. Spectral Gaussian fits are included as dashed lines, and vertical lines indicate the regions used to generate the line maps. Optical images (*grz*) are included from the DESI Legacy Survey (Dey et al. 2019) as insets to the CO(2–1) spectra. We note that of the seven galaxies that are undetected in CO(2–1), three data cubes include significantly detected lines in close physical and kinematic proximity to the *SQuIGGLE* galaxies. We interpret these as representing CO(2–1) emission from the cold gas reservoirs of neighboring galaxies. We note that due to the high stellar masses of the galaxies in this sample, we expect them to reside in dense environments but defer the analysis of these nearby sources to future work. We detect nonthermal continuum emission in SDSS_J0753+2403, which we subtract from the spectrum in Figure 4.

Of the six detected galaxies, CO(2–1) is largely spatially coincident with the optical centroid of the galaxies, as determined from the SDSS imaging, with the exception of SDSS_J1109-0040 and SDSS_J2202-0033 (also in Suess et al. 2017). In SDSS_J2202-0033, the $\sim 1''$ offset from the optical centroid is not significant given the resolution and S/N of the data; we note that the optical image of this galaxy appears to be slightly asymmetric. The CO(2–1) emission in SDSS_J1109-0040 is spatially offset by $\sim 1''\text{--}3''$; it is possible that this galaxy is in the process of a close, late-stage merger for which only one galaxy is detected in CO. A possible companion galaxy is visible in the optical image, on the opposite side from the direction of elongation. Alternatively, we may be detecting an especially strong CO outflow. In either case the emission is not especially broad ($\sigma = 123 \text{ km s}^{-1}$) in comparison to the other detected sources. In two additional sources, SDSS_J1448+1010 and SDSS_J2258+2313, the CO emission appears to extend beyond the optical extent of the galaxies, reaching distances up to $\sim 4''\text{--}5''$ from the center. These two objects may similarly be exhibiting late-stage mergers or molecular outflows in which a significant fraction of the molecular gas has been removed, or the gas may be associated with low-surface-brightness stellar light not apparent in the SDSS imaging. These offset and extended sources are targets of follow-up ALMA observations that will be presented in future work. For this work we assume that all CO(2–1) emission detected is physically associated with the *SQuIGGLE* poststarburst galaxies.

Figure 5 shows the CO(2–1) luminosity versus redshift (left), stellar mass (center), and *i* magnitude. To calculate H_2 masses from CO(2–1) we assume $r_{21} = 1.0$, making the conservative assumption of thermalized emission, (e.g., Combes et al. 2007; Dannerbauer et al. 2009; Young et al. 2011) and a Milky Way–like CO– H_2 conversion factor of $\alpha_{\text{CO}} = 4.0$ (Bolatto et al. 2013 and references within). We note that these assumptions are unlikely to lead to significant uncertainties in the implied H_2 properties. r_{21} could only be lower than assumed, which would

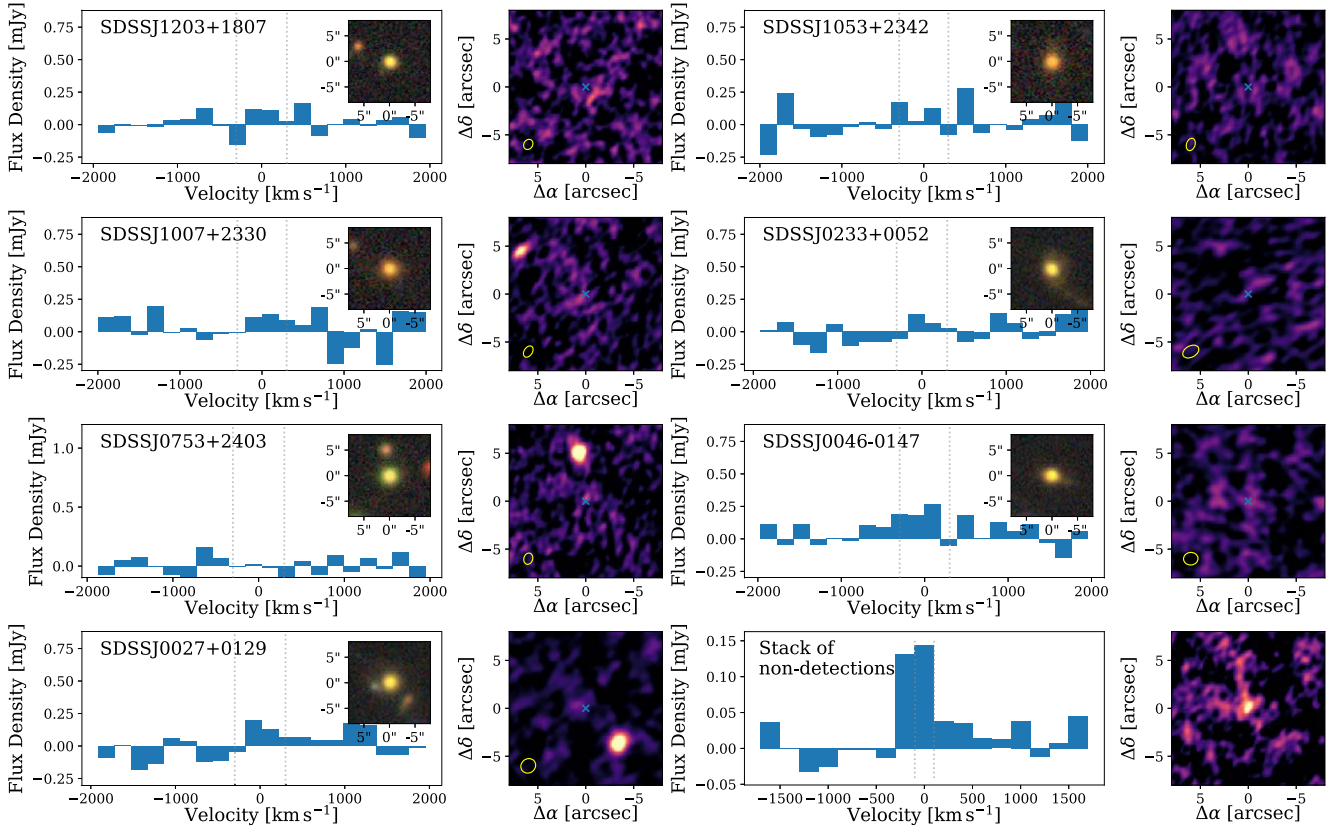


Figure 4. Spatially integrated CO(2–1) spectra (left) and line maps (right) and optical images (left inset) for undetected *SQuIGGLE* galaxies, sorted by decreasing i magnitude. Although individual galaxies are not detected, the stacked emission is detected (lower right). We note that we subtract a constant continuum offset from the spectrum of SDSS_J0753+2403, which we attribute to nonthermal continuum emission.

increase the M_{H_2} . Higher values of α_{CO} would require either cooler and lower-velocity-dispersion gas (increasing the CO line opacity) or lower metallicities (decreasing the CO abundance) than in the Milky Way, both of which are unlikely. Lower α_{CO} could arise from hotter gas or higher velocity dispersions than typical giant molecular clouds in the Milky Way, which may be realistic in this sample. However, a factor of ~ 2 – 3 reduction in α_{CO} would be counteracted by any decrease in r_{21} if the gas is not thermalized (e.g., Narayanan et al. 2012). Implied H_2 masses are indicated on the right axis in this figure. In the left panel there is a weak trend indicating that the higher-redshift galaxies in the *SQuIGGLE* sample are more likely to be detected in CO(2–1), which is possibly suggestive of an evolutionary sequence; however, we emphasize that the selection criteria are fairly narrow. CO(2–1) luminosity and M_{H_2} are independent of stellar mass or magnitude.

3. H_2 Reservoirs of *SQuIGGLE* Galaxies

3.1. H_2 and Scaling Relations

In this section we compare the enigmatic H_2 reservoirs of *SQuIGGLE* galaxies to the scaling relations defined by “normal” star-forming galaxies at high and low redshift and to other samples of sub-main-sequence galaxies. In Figure 6, we show the SFRs of the *SQuIGGLE* galaxies in blue (full sample) and green symbols for the ALMA targets. We include the “star-forming main sequence” scaling relation for star-forming galaxies at the average redshift of *SQuIGGLE* (Whitaker et al. 2012b). The *SQuIGGLE* sample lies significantly below this relation, although many have comparable gas reservoirs to coeval star-forming

galaxies. The SFRs for the CO-detected *SQuIGGLE* galaxies (green circles) are on average systematically higher by a factor of ~ 4 than the undetected counterparts (hatched circles). However, we emphasize that measuring very low SFRs ($\text{SFR} \lesssim 1 M_\odot \text{ yr}^{-1}$, dashed line) is extremely challenging.

Figure 7 shows the SFR versus H_2 gas mass for the *SQuIGGLE* poststarburst galaxies (green symbols). The empirical scaling relation for CO-based H_2 measurements is traced by colored contours. At lower SFR, massive galaxies at $z \sim 0$ from COLDGASS (Saintonge et al. 2011a, 2011b, 2012) are shown in black (for galaxies with detected CO(1–0) lines) and red (upper limits on CO(1–0)) contours. Star-forming galaxies at $\langle z \rangle = 1.2$ from PHIBSS/PHIBSS2 surveys (Tacconi et al. 2010; Genzel et al. 2015; Freundlich et al. 2019) are shown by blue contours. In the left panel, only individual *SQuIGGLE* galaxies and the *SQuIGGLE* stack are included. Detected galaxies are offset by $\gtrsim 1$ dex in M_{H_2} , while the stack of nondetections (hatched square) is consistent with the upper limits of galaxies with low SFRs in the local universe. In the right panel we include other samples of galaxies with SFRs that place them below the main sequence of star-forming galaxies (e.g., Noeske et al. 2007). This panel includes poststarburst galaxies at $z \sim 0$ from the SDSS (French et al. 2015; Rowlands et al. 2015). Much like the galaxies in *SQuIGGLE*, poststarburst galaxies from French et al. (2015; yellow diamonds) exhibit a range of H_2 reservoirs, which places some galaxies ~ 0.5 dex more gas-rich than the upper limits of local low-SFR galaxies. Poststarburst galaxies from Rowlands et al. (2015; purple diamonds) have higher quoted SFRs and lower M_{H_2} , lying well within the distribution of galaxies in the

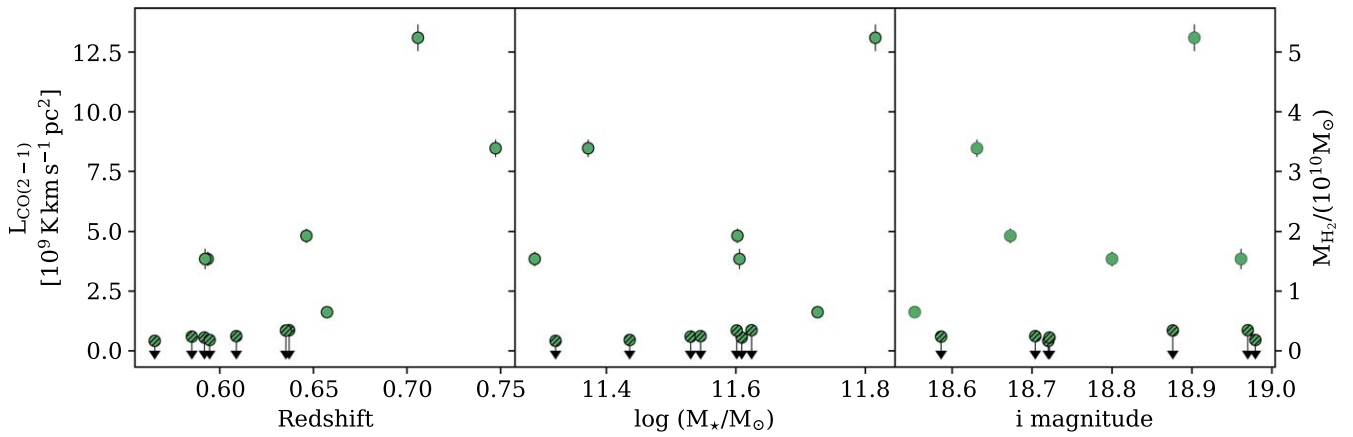


Figure 5. CO(2–1) luminosity (left axis) and M_{H_2} (right axis, assuming $\alpha_{\text{CO}} = 4.0$) vs. redshift (left panel), stellar mass (center panel), and i magnitude (right panel). Galaxies are primarily detected in the higher-redshift subsample ($z \gtrsim 0.65$) but are otherwise detected across the full range of stellar mass and magnitude.

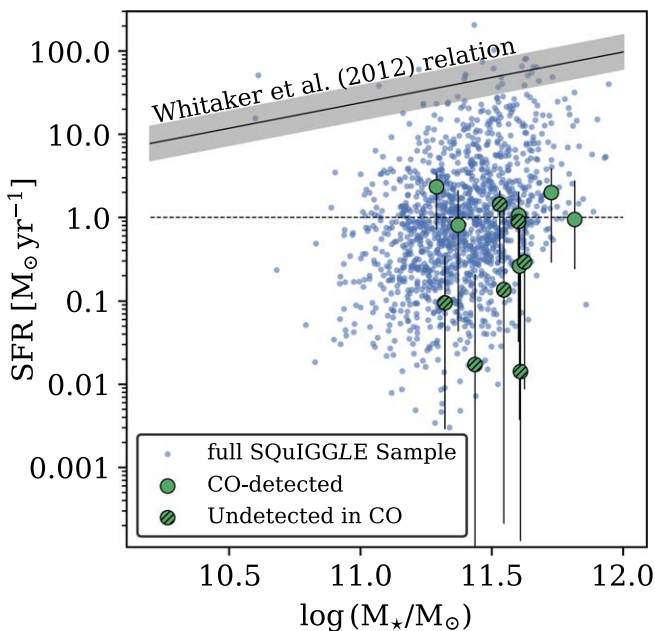


Figure 6. SFR vs. M_* measured from joint photometric and spectroscopic modeling using *Prospector*, for the full *SQuIGGLE* sample (blue) and ALMA targets, with the $1 M_\odot \text{ yr}^{-1}$ threshold below which SFRs are robustly recovered (Suess et al. 2021). Green circles indicate those that are detected in CO(2–1), hatched filling identifies undetected galaxies. We include the scaling relation for star-forming galaxies at the average redshift of the sample from Whitaker et al. (2012b). The *SQuIGGLE* galaxies span a wide range of SFRs, but CO-detected targets fall below the main sequence by at least an order of magnitude.

COLDGASS sample. At higher redshift, we also include galaxies at a roughly coeval $\langle z \rangle = 0.7$ from Spilker et al. (2018) as maroon hexagons, which are selected from the LEGA-C survey (van der Wel et al. 2016) to be sub-main sequence, but have optical spectra that suggest more gradual SFHs than poststarburst galaxies. Finally we include CO-based measurements of M_{H_2} for nine quiescent galaxies at $z \sim 1.5$ (four detections and five upper limits), including two with poststarburst spectral signatures (strong Balmer absorption lines; Bezanson et al. 2019; Belli et al. 2021; Williams et al. 2021). The H_2 reservoirs of the galaxies detected by Belli et al. (2021; blue squares) are comparable to those of the *SQuIGGLE* galaxies; however, their higher SFRs (likely due to the fact that

they are observed at higher redshift) place them closer to the coeval star-forming galaxies in the PHIBSS/PHIBSS2 survey (Tacconi et al. 2013, 2018).

It is useful to examine the SFR and H_2 masses in the more traditional projection of the Kennicutt–Schmidt (K-S) relation (e.g., Kennicutt 1998), as the density of cold gas is a more direct tracer of the fueling of star formation. As the galaxies are also unresolved in the existing ground-based imaging, we adopt circularized sizes derived from the CO(2–1) emission. Figure 8 shows the Σ_{SFR} versus Σ_{H_2} for *SQuIGGLE* galaxies that are detected in CO(2–1) and other samples with reliable size measurements, using the same symbols as Figure 7. We note that all other samples are included using sizes estimated from stellar effective radii. Given the uncertainty in spatial extent that impacts both surface densities, detected *SQuIGGLE* galaxies can only come closer to the tight K-S relation if their CO(2–1) emission is more extended than the spatial distribution of any residual star formation.

A number of groups have investigated additional integrated scaling relations between gas reservoirs and the fueling of star formation and star formation efficiency in galaxies (e.g., Lilly et al. 2013; Tacconi et al. 2013; Scoville et al. 2017; Tacconi et al. 2018, 2020). In part motivated by the time evolution of the characteristic $\text{SFR}(M_*)$ of galaxies, some papers advocate for redshift-evolving multivariate regressions that minimize the scatter, quantifying e.g., $M_{\text{H}_2}(\text{SFR}, M_*, z)$. The majority of the observational constraints on these scaling relations rely on galaxies that are near, on, or above the star-forming main sequence at any epoch. As a result, this sample provides a useful tool to probe the scatter about these relations at low SFRs. As motivated in Section 2, we adopt slightly different definitions for stellar mass and SFR in Figures 9 and 10. These definitions are more consistent with those used to derive the scaling relations. We use $M_{*,\text{FAST}}$, from delayed exponential SFHs, for the former and treat $\text{SFR} = 1 M_\odot \text{ yr}^{-1}$ as a floor for derived values, below which SFRs are plotted as upper limits.

Figure 9 shows the specific SFR (sSFR), normalized relative to the star-forming main sequence, versus H_2 gas fractions for the sample of poststarburst galaxies presented in this paper. We also include a coeval sample of massive and sub-main-sequence galaxies from Spilker et al. (2018). These individual measurements can be contrasted with extensions of two redshift-dependent scaling relations from Scoville et al. (2017; dark gray band) and

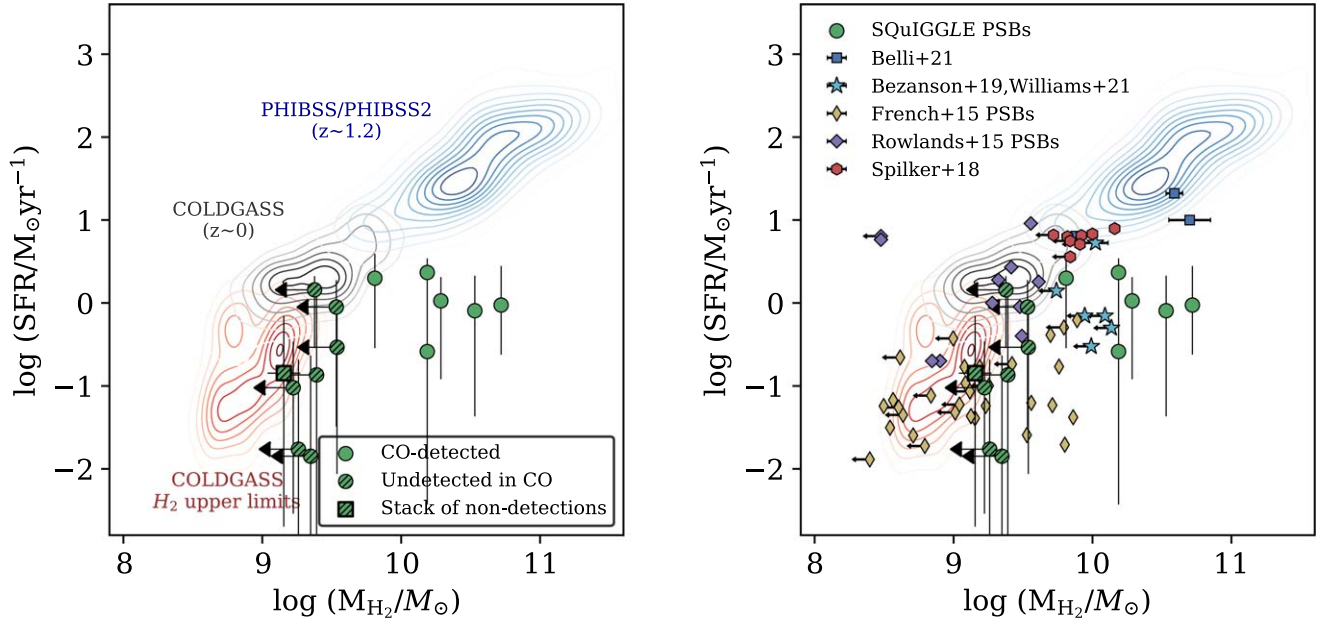


Figure 7. SFR vs. H_2 gas mass for *SQuIGGLE* galaxies and comparison samples of massive galaxies from COLDGASS at $z \sim 0$ (Saintonge et al. 2011a, 2011b) and star-forming galaxies at $z \sim 1.2$ from PHIBSS/PHIBSS2 (Tacconi et al. 2010, 2013). Green symbols in both panels indicate the *SQuIGGLE* poststarburst galaxies. Additional samples of $z \sim 0$ poststarburst galaxies (French et al. 2015; Rowlands et al. 2015) and galaxies below the star-forming SFR–stellar mass “main sequence” with CO-based H_2 mass estimates are included in the right panel (Spilker et al. 2018; Bezanson et al. 2019; Belli et al. 2021; Williams et al. 2021). Although most galaxies lie near the M_{H_2} –SFR relation, a number of *SQuIGGLE* and French et al. (2015) poststarburst galaxies harbor large H_2 reservoirs for their low SFRs. This is most dramatic for the CO(2–1)-detected *SQuIGGLE* galaxies, which are offset by over an order of magnitude in M_{H_2} .

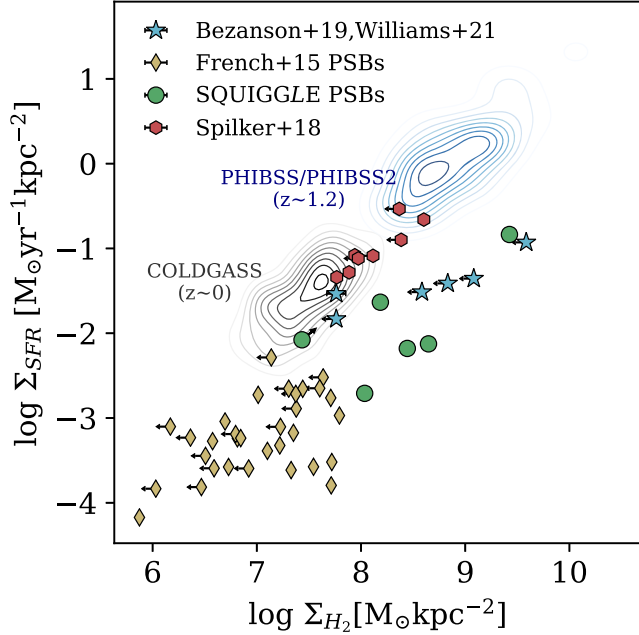


Figure 8. SFR surface density vs. H_2 surface density, or the Schmidt–Kennicutt relation for the detected *SQuIGGLE* galaxies. Sizes are estimated from the CO(2–1) emission, which is likely an overestimate of the stellar sizes. Even at the most extended limit, this sample lies at the highest-density edge of the local French et al. (2015) poststarburst galaxies and more than an order of magnitude offset from the K–S relation for “normal” star-forming galaxies.

Tacconi et al. (2020; light gray band), plotted assuming the average redshift and stellar mass of the *SQuIGGLE* ALMA sample. For each scaling relation, the quoted scatter is indicated by the gray band and the scaling relation is indicated by a solid gray or black line where the relation is calibrated and dotted lines

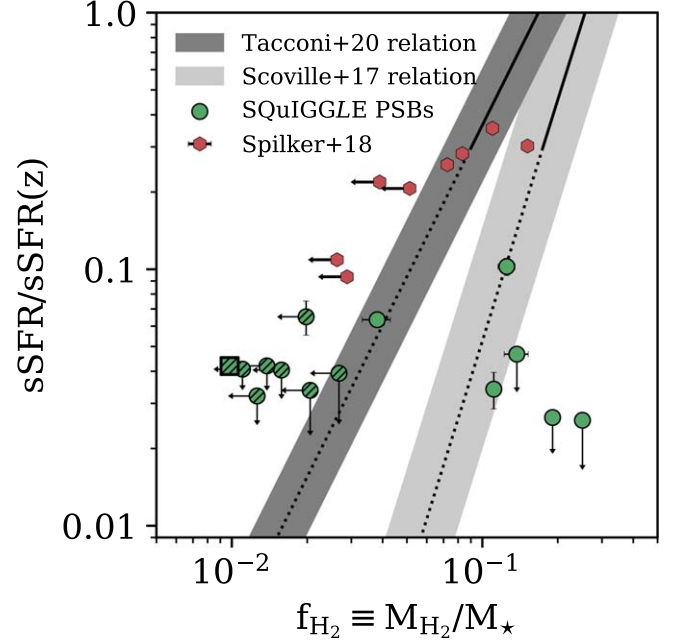


Figure 9. Specific SFR, normalized by the star-forming main sequence at the average redshift of the sample, vs. H_2 gas fraction for quiescent galaxies from this sample and a coeval sample of less extreme quiescent galaxies from Spilker et al. (2018). The combined sample reveals a clear increase in the scatter about integrated scaling relations in the quiescent regime, even when accounting for the lower efficiency of star formation in depleted galaxies.

to indicate where each is extrapolated. Although there is some agreement between the galaxies with higher sSFR and the scaling relations, at lower specific SFRs, where the two relations diverge, it is clear that there is a significant offset between the H_2 -detected

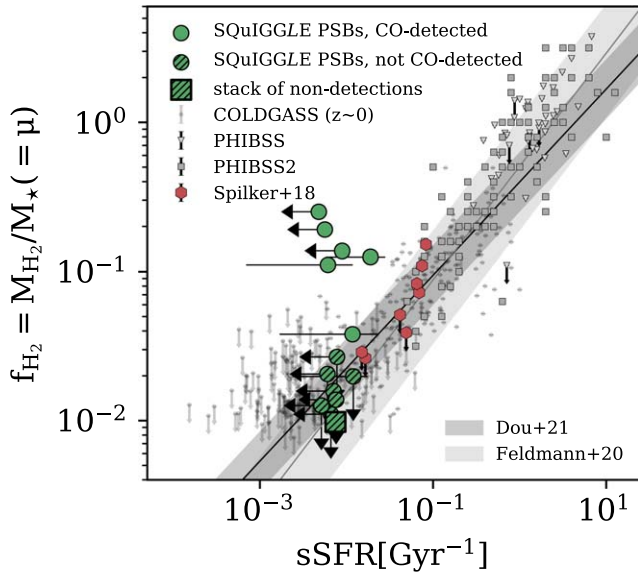


Figure 10. H_2 fraction vs. sSFR in absolute quantities, representing a suggested fundamental relation within which galaxies evolve (e.g., Feldmann 2020; Dou et al. 2021). We note that the gas-rich *SQuIGGLE* galaxies lie significantly outside the scatter in this relation, but the deviation may be temporary and the stack of older, nondetections lies on the extended relation.

SQuIGGLE galaxies (green circles) and the undetected counterparts (hatched green symbols)—individually or in the stack (square). Detected galaxies would need to have vastly underestimated SFRs (by at least an order of magnitude) to be consistent with either scaling relation. Such high SFRs are disfavored by, e.g., the nondetection of 2 mm dust continuum emission in any source.

It is also possible to define more inclusive scaling relations, along the lines of, e.g., the fundamental metallicity relation (e.g., Mannucci et al. 2010), that encompass the redshift evolution of the galaxy population by spanning the range of properties through which galaxies evolve. In addition to being easier to graphically depict, the simplicity of such scaling relations is appealing because it does not require relying on evolving physics. One such parameter space is the relation between H_2 gas fraction (typically denoted as f_{H_2} or μ) and specific SFR, as shown in Figure 10. In this figure we can include galaxies at a number of cosmic times, spanning from $z \sim 0$ (Saintonge et al. 2011a, 2011b, 2012) to $z \sim 2$ (Tacconi et al. 2013). At different redshifts, normal star-forming galaxies (all indicated by gray points) fall along a similar relation but occupy varying regimes in sSFR. We show two such parameterizations of this integrated scaling from Feldmann (2020) and Dou et al. (2021). In this projection, we note that quiescent galaxies from Spilker et al. (2018) and the stacked nondetections no longer fall outside of the scatter, rather they lie near lower-redshift star-forming counterparts. However, the H_2 -rich galaxies in this sample remain dramatic outliers in this projection of gas fueling as well. We note that $z \sim 0$ poststarburst galaxies from the French et al. (2018) sample lie in a similarly offset location, but we omit them from the diagram for clarity of presentation. This offset from all other samples of galaxies suggests that the deviation is temporary, but whether the H_2 -rich galaxies rejuvenate at SFRs that correspond to their gas reservoirs or lose, heat, or deplete H_2 and quench permanently cannot be determined.

3.2. Star Formation Histories and H_2 Reservoirs

Finally, we investigate the relationship between the stellar populations of the galaxies in this sample and the residual gas fractions in Figure 11. In Figure 11(a), we show *SQuIGGLE* galaxies on the $H\delta_A$ versus D_n4000 space, colored by f_{H_2} . Tracks were generated using Flexible Stellar Population Synthesis (FSPS; Conroy et al. 2009; Conroy & Gunn 2010), assuming two top-hat SFHs, $A_V = 0.5$, and solar metallicity. Lines of constant age are indicated by thin dotted lines, and time since burst in Myr is labeled in rectangular boxes. Galaxies with the highest gas fractions (circular symbols) are located toward the left, or low D_n4000 , portion of this panel, and galaxies that are undetected at the survey depths ($\lesssim 3 \times 10^9 M_\odot$, hatched circles) have higher D_n4000 values. This corresponds to roughly an inverse relationship between the existence of a gas reservoir and time since burst, with the majority of detectable H_2 gas in galaxies that quenched in the last ~ 100 –200 Myr.

Trends between molecular gas fraction and stellar continuum indices are shown in the center and right panels of Figure 11: f_{H_2} versus $H\delta_A$ (Figure 11(b)) and D_n4000 (Figure 11(c)). There is no significant trend with $H\delta_A$ (Figure 11(b)). This suggests that residual gas reservoirs are independent of the fraction of mass that was formed during the major burst, which primarily drives the maximum $H\delta_A$. The analysis of this trend is complicated by the fact that this parameter is double valued in its time evolution. The trend with D_n4000 is much more apparent; the majority (five of seven) of galaxies with $D_n4000 < 1.4$ have $f_{H_2} > 5\%$, while at higher values only one galaxy has detectable CO(2–1) emission. The average f_{H_2} implied by the stacked CO(2–1) flux from individually undetected galaxies is consistent with the decreasing trend in f_{H_2} with D_n4000 . Unlike $H\delta_A$, D_n4000 increases monotonically with age; the trend we observe corresponds to declining H_2 reservoirs after the end of the starburst event. After about a hundred Myr, the molecular gas fractions dwindle below $\sim 1\%$.

The *Prospector* stellar population synthesis modeling used to determine, e.g., the stellar masses and SFR of galaxies in the sample also provides flexible SFHs. From these SFHs, we compute the time between rapid quenching of star formation and the time of observation, which we refer to as the time since quenching (t_q ; see Suess et al. 2021). In Figure 12, we show that the H_2 gas masses (left) and fractions (right) are even more strongly correlated with t_q than with the empirical indices. This is likely driven by the fact that these derived SFHs are based on the full suite of spectral indices, essentially incorporating complex combinations of features that vary on different timescales (see, e.g., Figure 11(a)). It is immediately clear that the H_2 -rich versus H_2 -poor bimodality is statistically significant; only galaxies with $t_q \lesssim 200$ Myr retain $f_{H_2} > 1\%$. We note that the single galaxy that is detected in H_2 , but has a slightly longer t_q such that it overlaps with the H_2 -poor subset, is the single object that falls out of the DR14 spectroscopic S/N cuts. This may explain the relatively large uncertainty on t_q ; however, we include this potentially discrepant object in our analysis to avoid introducing a confirmation bias. We perform a student T-test on the full sample and verify that the detected H_2 and 3σ limits on M_{H_2} are not drawn from the same distribution, with $p = 0.02$. In addition to the bimodality, we quantify the implied exponential decay timescale by fitting $\ln M_{H_2}/M_*$ versus t_q , including the detected galaxies and the stack of undetected targets using orthogonal distance regression with `scipy.ODR`. The

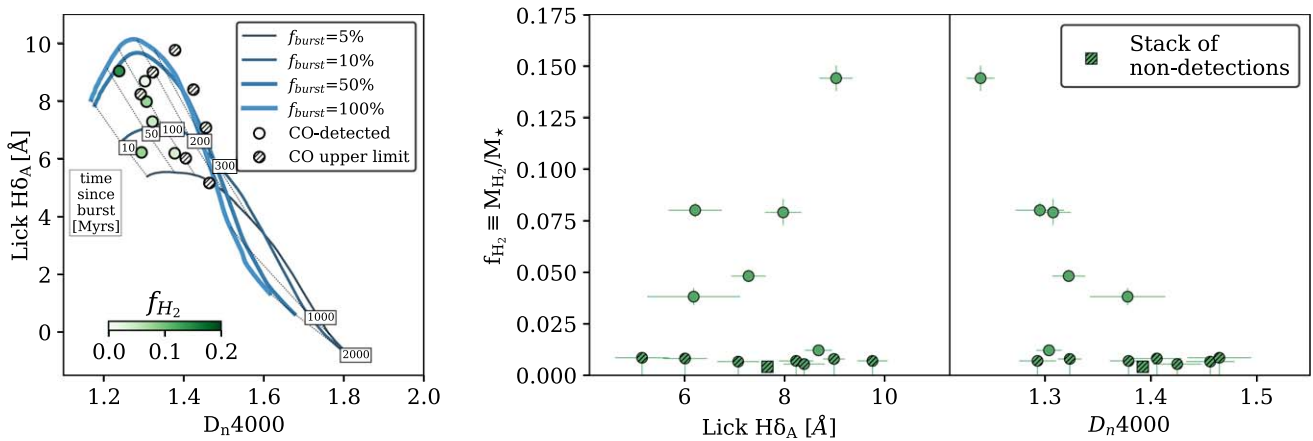


Figure 11. *SQuIGGLE* galaxies in $H\delta_A$ vs. D_n4000 space, colored by H_2 gas fraction, with upper limits indicated by triangles (left panel). Two-burst evolutionary tracks are indicated in blue, with lines of constant age in dotted black. Galaxies do not show a clear trend between H_2 reservoirs and $H\delta_A$ (center panel); however, they exhibit higher H_2 gas fractions ($\gtrsim 5\%$) at the lower D_n4000 (right panel). This suggests that molecular gas reservoirs diminish with time after ~ 100 Myr.

resulting fit and confidence interval are included in Figure 12(b) as a dashed line and gray band. This analysis yields an exponential depletion timescale of 84 ± 45 Myr, a remarkably rapid decline in the H_2 reservoirs postquenching, as we discuss further below.

4. Discussion

The primary result of this paper is that nearly half of the massive, poststarburst galaxies at $z \sim 0.6$ in the *SQuIGGLE* sample retain significant H_2 reservoirs ($\sim 1\text{--}5 \times 10^{10} M_\odot$), building upon our pilot study of two galaxies (Suess et al. 2017). Timing derived from stellar population synthesis modeling suggests that this H_2 disappears rapidly; no CO(2–1) emission is detected in galaxies observed $\gtrsim 150$ Myr after their star formation truncated. The existence of similar enigmatic molecular gas reservoirs has also been reported in a number of samples of poststarburst galaxies in the local universe (e.g., French et al. 2015; Rowlands et al. 2015; Alatalo et al. 2016). Moreover, we find a striking difference in the molecular gas properties of young versus old poststarbursts, suggesting that the cold gas rapidly disappears $\sim 100\text{--}200$ Myr postquenching. Intriguingly, a detailed study of the SFHs of the local sample found a similar anticorrelation between t_q and the H_2 reservoirs, finding an exponential depletion timescale of 117–230 Myr (French et al. 2018). Those local samples span a much larger range in stellar masses than *SQuIGGLE* ($9 \lesssim \log M_*/M_\odot \lesssim 11.5$). At low redshift, similarly massive ($\log M_*/M_\odot \sim 11$) $z \sim 0$ galaxies tend to be products of less extreme ($\sim 10\%$ mass fractions) bursts. Regardless, the similarity with the maximum t_q beyond which no *SQuIGGLE* galaxies are detected in CO(2–1) is striking, suggesting the two populations experience similar ties between quenching and cold gas depletion in the ~ 200 Myr after quenching.

We emphasize that the depletion times due to ongoing star formation for the *SQuIGGLE* galaxies are very long ($\gtrsim 1$ Gyr) given their large H_2 reservoirs and low unobscured SFRs. In Figure 13 we show the star formation depletion times ($t_{\text{dep}} \equiv M_{H_2}/\text{SFR}$) for *SQuIGGLE* targets (green symbols) and the stack of nondetections (square) versus t_q . It is immediately clear that the younger *SQuIGGLE* galaxies have insufficient residual star formation for the observed reservoirs to deplete, especially if we assume that this trend corresponds to an

evolutionary sequence. The exponential depletion time calculated for this sample ($\tau_{\text{dep}} \sim 84 \pm 45$ Myr) is comparable (within $\sim 1\sigma$) to the depletion timescale found by French et al. (2018) in low-redshift poststarburst galaxies. French et al. (2018) argue convincingly that such a rapid depletion timescale cannot be explained by ongoing minimal star formation, extreme variations in stellar populations (e.g., the initial mass function), stellar winds, or supernova feedback, pointing to this as possible smoking-gun evidence for AGN feedback. We emphasize that we cannot rule out the possibility that star formation is only temporarily halted in the H_2 -rich *SQuIGGLE* galaxies, which are caught in the opposite transition, immediately prior to rejuvenation. In that case, the two halves of the *SQuIGGLE* sample (gas-rich and gas-poor) may represent two different populations altogether depending on whether or not they will resume star formation in the future.

Another possible explanation for the large H_2 reservoirs that do not appear to fuel star formation could be that these galaxies harbor heavily dust-obscured star formation, causing depletion time estimates to be severely underestimated. We do not see evidence of continuum emission due to dust-obscured star formation in these galaxies in the 2 mm ALMA data presented in this work,¹⁶ placing a limit of $\text{SFR} \lesssim 50 M_\odot \text{ yr}^{-1}$. Although a number of galaxies in the *SQuIGGLE* data set are detected by the Very Large Array (VLA) Faint Images of the Sky at Twenty-Centimeters (FIRST) survey (Becker et al. 1995), those data are too shallow to be sensitive to realistic levels of radio emission due to ongoing star formation (such fluxes would correspond to $\text{SFR} \gtrsim 1000 M_\odot \text{ yr}^{-1}$). Instead, we interpret those detections as originating from AGN activity and explore that connection in a separate paper (Greene et al. 2020). In general, estimating the SFR for poststarburst galaxies is challenging. An analysis of local E+A galaxies found significant scatter among different indicators, concluding that total infrared luminosity provides overestimates of the intrinsic SFRs (Smercina et al. 2018). This implies that if $\text{SFR}_{\text{IR}} \lesssim 50 M_\odot \text{ yr}^{-1}$ under standard assumptions, the true upper limit would be even stronger. We note that even if the SFRs are underestimated by an order of magnitude, the depletion times for the most extreme, youngest *SQuIGGLE* would still be

¹⁶ As noted in Section 2.1, one galaxy, SDSS J0753+2403 has detected nonthermal synchrotron emission, which we attribute to AGN activity.

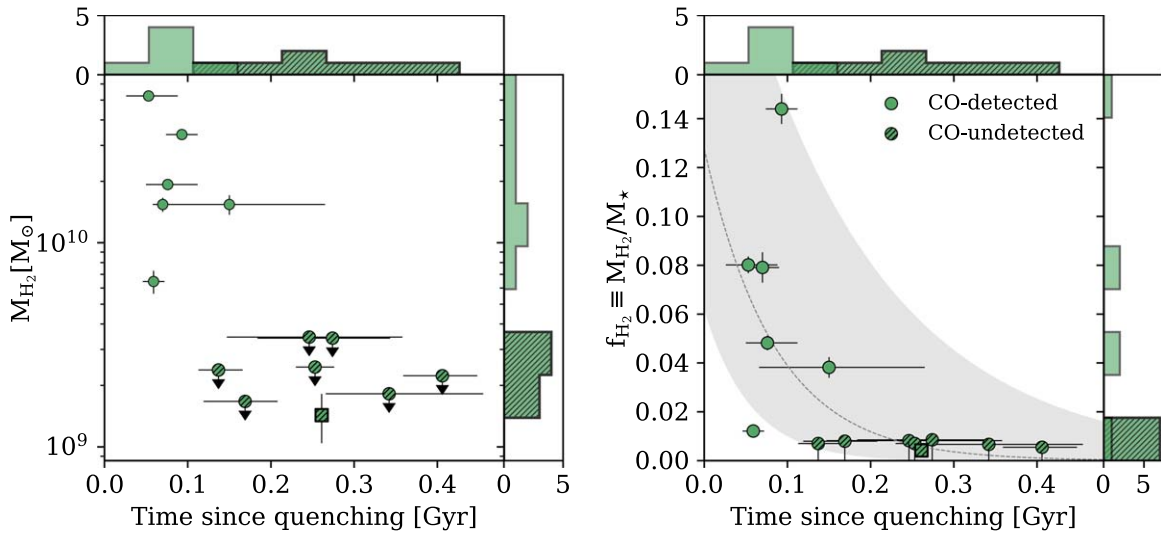


Figure 12. H_2 gas mass (left) and gas fraction (right) vs. time since quenching for *SQuIGGLE* targets (green circles) and average H_2 fraction derived from the stack of individually undetected galaxies (hatched square). Histograms of the two samples are included at the top and right of each panel. The correlation between whether a galaxy retains significant H_2 and the time since quenching is striking; all galaxies older than $t_q \gtrsim 200$ Myr are undetected in CO emission. The best-fit exponentially declining relation (with $\tau_{\text{dep}} \approx 84 \pm 45$ Myr) is included as a gray band in the right panel.

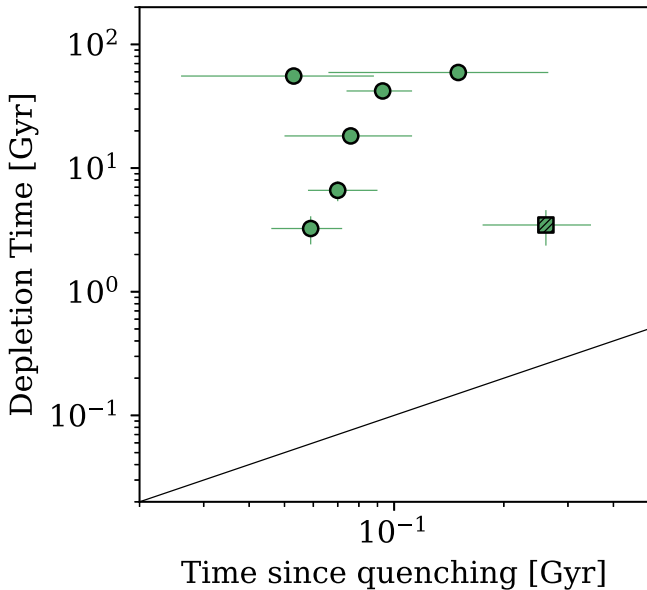


Figure 13. Depletion time vs. time since quenching for CO-detected *SQuIGGLE* targets (circles) and average depletion time derived from the stack of individually undetected galaxies (hatched square). Depletion times are much longer than the time since quenching (the one-to-one line is indicated by the black line), necessitating additional physical heating or removal of H_2 beyond residual star formation to connect the younger and older groups if this represents an evolutionary sequence.

closer to ~ 1 Gyr, which is much longer than the range in t_q probed by this sample. Therefore, if we assume that the trend in Figure 12 is a time sequence, ~ 2 dex of dust-obscured star formation would be needed to deplete the H_2 , which we would expect to see in continuum emission in the Band 4 ALMA data.

Interestingly, this empirical finding of large H_2 reservoirs in the youngest *SQuIGGLE* galaxies is suggestively similar to an observed trend in the AGN occurrence rates within the same sample (Greene et al. 2020). This is especially clear for AGNs identified via optical emission lines (high $[\text{O III}]/\text{H}\beta$ ratios), which are ~ 10 times more common in the youngest

SQuIGGLE galaxies. Although the parent sample is the same for both studies, only one galaxy (SDSS_J1448+1010) with strong $[\text{O III}]$ emission, indicating the presence of an AGN, is included in this ALMA sample. Therefore, we cannot make any robust claims about a possible causal correlation between the presence of an actively accreting supermassive black hole and the removal, heating, or destruction of H_2 within or surrounding these massive poststarburst galaxies.

While poststarburst galaxies remain a subdominant population of galaxies for the past ~ 7 Gyr, observational studies of massive galaxies of the high-redshift universe have begun to demonstrate that around $z \sim 2-3$ the general population of massive ($\log M_*/M_\odot \sim 11$) galaxies is in the process of rapidly quenching their primary episode of star formation (Whitaker et al. 2012a; Muzzin et al. 2013; Tomczak et al. 2014; Davidzon et al. 2017). Spectroscopic samples of such distant, massive galaxies indicate that poststarburst stellar populations are common (e.g., van de Sande et al. 2011; Bezanson et al. 2013; van de Sande 2013; Belli et al. 2015; Carnall et al. 2019; Kriek et al. 2019; Wild et al. 2020; Tacchella et al. 2021), suggesting that many massive galaxies undergo a dramatic truncation of star formation. Which physical processes are responsible for driving those quenching events that can also concurrently destroy, deplete, or heat their molecular gas reservoirs are poorly constrained empirically. Bezanson et al. (2019) placed stringent upper limits $f_{\text{H}_2} \lesssim 7\%$ in a $z = 1.522$ galaxy and similarly Williams et al. (2021) expanded the sample to include five additional galaxies. From the overall low H_2 fractions in that sample, Williams et al. (2021) conclude the need for rapid ($\tau_{\text{dep}} \sim 0.3$ Gyr) depletion. We note that this timescale cannot be compared directly to the t_q parameter used in this study, which starts specifically after star formation shuts down. The galaxies in those samples were significantly older, with a postburst age closer to $\gtrsim 1$ Gyr than the *SQuIGGLE* galaxies that are detected in CO(2–1). Therefore, although the galaxies at that epoch are generally more gas-rich than local galaxies (e.g., Tacconi et al. 2013, 2018; Freundlich et al. 2019; Belli et al. 2021), we would not have predicted that they would be detected based on the results presented in this paper.

Therefore, while the striking similarity between studies of poststarburst galaxies at $z \sim 0$ and the more extreme cases at $z \sim 0.6$ presented in this paper is suggestive that quenching does not necessarily coincide with the elimination of the H_2 reservoirs, direct studies of galaxies during the primary quenching epoch will be critical in strengthening our understanding of the quenching mechanisms at play.

While no model includes the existence of H_2 after star formation shuts down by construction, the rapid disappearance of these cold gas reservoirs in poststarburst galaxies is easily consistent with the gas-poor nature of older quiescent galaxies (e.g., Young et al. 2011, 2014). As they stand, these data present a challenge to galaxy formation models. Larger, more statistical studies of H_2 reservoirs in this and similar populations of recently quenched galaxies could precisely time the disappearance of H_2 . We anticipate that spatially resolved H_2 maps may provide clues as to the distribution and kinematics of the H_2 , which may help explain how it is stabilized against collapse. Initial studies of the stellar kinematics of this sample have revealed a range of rotational support (Setton et al. 2020) and, in one case, that the CO(2–1) kinematics follow the stellar motion (Hunt et al. 2018). French et al. (2018) found a smooth exponential depletion of gas over time to provide the best fit for local galaxies. Deeper observations of individually undetected galaxies could assess how smooth the transition from H_2 rich is for our intermediate-redshift galaxies, as the current observations seem to reveal a more discontinuous distribution.

Theoretical and observational studies have pointed out that dynamical support against H_2 collapse could also stem from turbulent pressure in the gas. For example, slightly lower star formation efficiency is observed in morphologically classified local early-type galaxies; the same molecular gas reservoirs form ~ 2.5 times fewer stars in ellipticals than in spiral galaxies (e.g., Davis et al. 2014). Similarly, quiescence below the main sequence is driven by both depleted gas reservoirs and diminished star formation efficiency (e.g., Piotrowska et al. 2020). This effect is often attributed to relatively deep gravitational wells and increased disordered orbits, “morphological quenching” (Martig et al. 2009), but could also be driven by streaming motions (Meidt et al. 2013). Gensior et al. (2020) found that the existence of compact, spheroidal structures can indeed drive turbulent pressure, pointing toward the feasibility of morphological quenching. This could be especially relevant given that the sizes of poststarburst galaxies are often compact, even with respect to older quiescent galaxies (e.g., Yano et al. 2016; Wu et al. 2020; D. Setton 2021, in preparation.). It has been demonstrated the turbulent motions in the interstellar medium, perhaps induced by a combination of shocks and magnetic fields, could support H_2 against collapse (e.g., Federrath 2015). Simulations of relativistic jets can drive shocks that diminish SFRs by a factor of ~ 2 (e.g., Mandal et al. 2021), but not all specific implementations of jet astrophysics can effectively quench star formation (e.g., Su et al. 2021). Furthermore, given that most of these processes can decrease star formation efficiency by only a modest factor of a few, it remains unclear whether even combining these models could explain the order-of-magnitude offsets in star formation efficiency exhibited by the youngest *SQuIGGLE* galaxies. Perhaps further analysis of systematic suites of simulations (e.g., Su et al. 2021) could use the immediate decrease in star formation efficiency and subsequent rapid (~ 100 Myr)

disappearance of H_2 to differentiate among feedback models. Such studies could be especially constrained by extracting observed quantities matched to the *SQuIGGLE* data set. In particular, Su et al. (2021) showed the promising efficiency of cosmic-ray jets, which rapidly shut off star formation on a similar timescale; a more careful comparison would be needed to assess whether the depletion time lag is also consistent.

Although the similarity of the French et al. (2018) results at $z \sim 0$ and the more extreme poststarbursts at intermediate redshifts presented in this paper suggests a fundamental challenge to galaxy formation models, the strongest test will come from earlier times, at the peak epoch of galaxy quenching and transformation (e.g., Wild et al. 2016). In this paper, we focus on CO-based measurements of H_2 ; however, recent studies have suggested elevated average cold gas reservoirs in quiescent galaxies at cosmic noon based on stacked far-infrared dust continuum emission (e.g., Gobat et al. 2018; Magdis et al. 2021), in apparent contradiction with low or absent M_{H_2} in individual galaxies (e.g., Caliendo et al. 2021; Whitaker et al. 2021; Williams et al. 2021). One possible interpretation is that these stacks include a subset of H_2 -rich recently quenched ($t_q \lesssim 200$ Myr) galaxies averaged with a depleted majority. Identifying these young galaxies requires spectroscopic data of sufficient quality to precisely measure SFHs. Although current spectroscopic samples of galaxies at cosmic noon are somewhat rare, in the coming years massively multiplexed spectrographs with NIR capabilities, like the Prime Focus Spectrograph on Subaru (Takada et al. 2014) or MOONS (Multi-Object Optical and Near-infrared Spectrograph; Maiolino et al. 2020), will produce hundreds of recently quenched targets at this critical epoch. Follow-up studies of their H_2 reservoirs, either with CO or dust-based estimates, will ultimately determine the timescale and simultaneity of quenching and the disappearance of H_2 in massive quiescent galaxies.

R.S.B., J.E.G., D.J.S., and D.N. gratefully acknowledge support from NSF-AAG#1907697, #1907723, and #1908137. J.S.S. acknowledges support provided by NASA through NASA Hubble fellowship grant #HF2-51446 awarded by the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., for NASA, under contract NAS5-26555. R.F. acknowledges financial support from the Swiss National Science Foundation (grant No. PP00P2_157591, PP00P2_194814, and 200021_188552). This paper makes use of the following ALMA data: ADS/JAO.ALMA #2016.1.01126.S and ADS/JAO.ALMA #2017.1.01109.S. ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada), MOST and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO and NAOJ. The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. K.A.S. gratefully acknowledges the UCSC Chancellors Fellowship.

This work was performed in part at the Aspen Center for Physics, which is supported by National Science Foundation grant No. PHY-1607611. We also thank the North American ALMA Science Center (NAASC) for their generous funding that helped support the participation of junior scientists (including D.J.S. and K.A.S.) at the Aspen Center for Physics workshop in 2020 February. The NAASC is part of the

National Radio Astronomy Observatory, a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

Funding for SDSS-III has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, and the U.S. Department of Energy Office of Science. The SDSS-III website is <http://www.sdss3.org/>.

SDSS-III is managed by the Astrophysical Research Consortium for the Participating Institutions of the SDSS-III Collaboration including the University of Arizona, the Brazilian Participation Group, Brookhaven National Laboratory, Carnegie Mellon University, University of Florida, the French Participation Group, the German Participation Group, Harvard University, the Instituto de Astrofísica de Canarias, the Michigan State/Notre Dame/JINA Participation Group, Johns Hopkins University, Lawrence Berkeley National Laboratory, Max Planck Institute for Astrophysics, Max Planck Institute for Extraterrestrial Physics, New Mexico State University, New York University, Ohio State University, Pennsylvania State University, University of Portsmouth, Princeton University, the Spanish Participation Group, University of Tokyo, University of Utah, Vanderbilt University, University of Virginia, University of Washington, and Yale University.

The Legacy Surveys consist of three individual and complementary projects: the Dark Energy Camera Legacy Survey (DECaLS; Proposal ID #2014B-0404; PIs: David Schlegel and Arjun Dey), the Beijing-Arizona Sky Survey (BASS; NOAO Prop. ID #2015A-0801; PIs: Zhou Xu and Xiaohui Fan), and the Mayall z-band Legacy Survey (MzLS; Prop. ID #2016A-0453; PI: Arjun Dey). DECaLS, BASS, and MzLS together include data obtained, respectively, at the Blanco telescope, Cerro Tololo Inter-American Observatory, NSF’s NOIRLab; the Bok telescope, Steward Observatory, University of Arizona; and the Mayall telescope, Kitt Peak National Observatory, NOIRLab. The Legacy Surveys project is honored to be permitted to conduct astronomical research on Iolkam Du’ag (Kitt Peak), a mountain with particular significance to the Tohono Oodham Nation.

NOIRLab is operated by the Association of Universities for Research in Astronomy (AURA) under a cooperative agreement with the National Science Foundation.

This project used data obtained with the Dark Energy Camera (DECam), which was constructed by the Dark Energy Survey (DES) collaboration. Funding for the DES Projects has been provided by the U.S. Department of Energy, the U.S. National Science Foundation, the Ministry of Science and Education of Spain, the Science and Technology Facilities Council of the United Kingdom, the Higher Education Funding Council for England, the National Center for Supercomputing Applications at the University of Illinois at Urbana-Champaign, the Kavli Institute of Cosmological Physics at the University of Chicago, Center for Cosmology and Astro-Particle Physics at the Ohio State University, the Mitchell Institute for Fundamental Physics and Astronomy at Texas A&M University, Financiadora de Estudos e Projetos, Fundação Carlos Chagas Filho de Amparo, Financiadora de Estudos e Projetos, Fundação Carlos Chagas Filho de Amparo a Pesquisa do Estado do Rio de Janeiro, Conselho Nacional de Desenvolvimento Científico e Tecnológico and the Ministerio da Ciencia, Tecnologia e Inovacao, the Deutsche Forschungsgemeinschaft and the Collaborating Institutions in the Dark Energy Survey. The Collaborating Institutions are Argonne National Laboratory, the University of California at

Santa Cruz, the University of Cambridge, Centro de Investigaciones Energeticas, Medioambientales y Tecnologicas-Madrid, the University of Chicago, University College London, the DES-Brazil Consortium, the University of Edinburgh, the Eidgenössische Technische Hochschule (ETH) Zurich, Fermi National Accelerator Laboratory, the University of Illinois at Urbana-Champaign, the Institut de Ciències de l’Espai (IEEC/CSIC), the Institut de Física d’Altes Energies, Lawrence Berkeley National Laboratory, the Ludwig Maximilians Universität München and the associated Excellence Cluster Universe, the University of Michigan, NSF’s NOIRLab, the University of Nottingham, the Ohio State University, the University of Pennsylvania, the University of Portsmouth, SLAC National Accelerator Laboratory, Stanford University, the University of Sussex, and Texas A&M University.

BASS is a key project of the Telescope Access Program (TAP), which has been funded by the National Astronomical Observatories of China, the Chinese Academy of Sciences (the Strategic Priority Research Program “The Emergence of Cosmological Structures” grant No. XDB09000000), and the Special Fund for Astronomy from the Ministry of Finance. The BASS is also supported by the External Cooperation Program of Chinese Academy of Sciences (Grant No. 114A11KYSB20160057), and the Chinese National Natural Science Foundation (grant No. 11433005).

The Legacy Survey team makes use of data products from the Near-Earth Object Wide-field Infrared Survey Explorer (NEOWISE), which is a project of the Jet Propulsion Laboratory/California Institute of Technology. NEOWISE is funded by the National Aeronautics and Space Administration.

The Legacy Surveys imaging of the DESI footprint is supported by the Director, Office of Science, Office of High Energy Physics of the U.S. Department of Energy under contract No. DE-AC02-05CH1123, by the National Energy Research Scientific Computing Center, a DOE Office of Science User Facility under the same contract; and by the U.S. National Science Foundation, Division of Astronomical Sciences under Contract No. AST-0950945 to NOAO.

Facility: ALMA.

Software: astropy (Astropy Collaboration et al. 2013), scipy (Virtanen et al. 2020), seaborn (Waskom 2021).

ORCID iDs

Rachel Bezanson  <https://orcid.org/0000-0001-5063-8254>
 Justin S. Spilker  <https://orcid.org/0000-0003-3256-5615>
 Katherine A. Suess  <https://orcid.org/0000-0002-1714-1905>
 David J. Setton  <https://orcid.org/0000-0003-4075-7393>
 Robert Feldmann  <https://orcid.org/0000-0002-1109-1919>
 Jenny E. Greene  <https://orcid.org/0000-0002-5612-3427>
 Mariska Kriek  <https://orcid.org/0000-0002-7613-9872>
 Desika Narayanan  <https://orcid.org/0000-0002-7064-4309>
 Margaret Verrico  <https://orcid.org/0000-0003-1535-4277>

References

- Abolfathi, B., Aguado, D. S., Aguilar, G., et al. 2018, *ApJS*, **235**, 42
- Alatalo, K. 2015, *ApJL*, **801**, L17
- Alatalo, K., Lisenfeld, U., Lanz, L., et al. 2016, *ApJ*, **827**, 106
- Anglés-Alcázar, D., Faucher-Giguère, C.-A., Quataert, E., et al. 2017, *MNRAS*, **472**, L109
- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, **558**, A33
- Becker, R. H., White, R. L., & Helfand, D. J. 1995, *ApJ*, **450**, 559
- Belli, S., Contursi, A., Genzel, R., et al. 2021, *ApJL*, **909**, L11

- Belli, S., Newman, A. B., & Ellis, R. S. 2015, *ApJ*, **799**, 206
- Bezanson, R., Spilker, J., Williams, C. C., et al. 2019, *ApJL*, **873**, L19
- Bezanson, R., van Dokkum, P., van de Sande, J., Franx, M., & Kriek, M. 2013, *ApJL*, **764**, L8
- Bolatto, A. D., Wolfire, M., & Leroy, A. K. 2013, *ARA&A*, **51**, 207
- Bruzual, G., & Charlot, S. 2003, *MNRAS*, **344**, 1000
- Caliendo, J. N., Whitaker, K. E., Akhshik, M., et al. 2021, *ApJL*, **910**, L7
- Calzetti, D. 1997, *AJ*, **113**, 162
- Carnall, A. C., McLure, R. J., Dunlop, J. S., et al. 2019, *MNRAS*, **490**, 417
- Chabrier, G. 2003, *PASP*, **115**, 763
- Combes, F., Young, L. M., & Bureau, M. 2007, *MNRAS*, **377**, 1795
- Conroy, C., & Gunn, J. E. 2010, FSPS: Flexible Stellar Population Synthesis, ascl:1010.043
- Conroy, C., Gunn, J. E., & White, M. 2009, *ApJ*, **699**, 486
- Crain, R. A., Schaye, J., Bower, R. G., et al. 2015, *MNRAS*, **450**, 1937
- Croton, D. J., Springel, V., White, S. D. M., et al. 2006, *MNRAS*, **365**, 11
- Dannerbauer, H., Daddi, E., Riechers, D. A., et al. 2009, *ApJL*, **698**, L178
- Davé, R., Anglés-Alcázar, D., Narayanan, D., et al. 2019, *MNRAS*, **486**, 2827
- Davé, R., Rafiee-Ashtiani, M. H., Thompson, R. J., & Hopkins, P. F. 2017, *MNRAS*, **467**, 115
- Davidzon, I., Ilbert, O., Laigle, C., et al. 2017, *A&A*, **605**, A70
- Davis, T. A., Alatalo, K., Bureau, M., et al. 2013, *MNRAS*, **429**, 534
- Davis, T. A., Alatalo, K., Sarzi, M., et al. 2011, *MNRAS*, **417**, 882
- Davis, T. A., Young, L. M., Crocker, A. F., et al. 2014, *MNRAS*, **444**, 3427
- Dawson, K. S., Schlegel, D. J., Ahn, C. P., et al. 2013, *AJ*, **145**, 10
- Dey, A., Schlegel, D. J., Lang, D., et al. 2019, *AJ*, **157**, 168
- Dou, J., Peng, Y., Renzini, A., et al. 2021, *ApJ*, **907**, 114
- Dressler, A., & Gunn, J. E. 1983, *ApJ*, **270**, 7
- Federrath, C. 2015, *MNRAS*, **450**, 4035
- Feldmann, R. 2020, *CmPhy*, **3**, 226
- Feldmann, R., & Mayer, L. 2015, *MNRAS*, **446**, 1939
- Feldmann, R., Quataert, E., Hopkins, P. F., Faucher-Giguère, C.-A., & Kereš, D. 2017, *MNRAS*, **470**, 1050
- French, K. D. 2021, *PASP*, **133**, 072001
- French, K. D., Yang, Y., Zabludoff, A., et al. 2015, *ApJ*, **801**, 1
- French, K. D., Yang, Y., Zabludoff, A. I., & Tremonti, C. A. 2018, *ApJ*, **862**, 2
- Freundlich, J., Combes, F., Tacconi, L. J., et al. 2019, *A&A*, **622**, A105
- Gensior, J., Kruijssen, J. M. D., & Keller, B. W. 2020, *MNRAS*, **495**, 199
- Genzel, R., Tacconi, L. J., Lutz, D., et al. 2015, *ApJ*, **800**, 20
- Gobat, R., Daddi, E., Magdis, G., et al. 2018, *NatAs*, **2**, 239
- Greene, J. E., Setton, D., Bezanson, R., et al. 2020, *ApJL*, **899**, L9
- Hunt, Q., Bezanson, R., Greene, J. E., et al. 2018, *ApJL*, **860**, L18
- Johnson, B., Foreman-Mackey, D., Sick, J., et al. 2021, dfm/python-fsps: python-fsps v0.4.0, vv0.4.0, Zenodo, doi:10.5281/zenodo.4577191
- Johnson, B., & Leja, J. 2017, Bd-J/Prospector: Initial Release, vv0.1, Zenodo, doi:10.5281/zenodo.1116491
- Kauffmann, G., Heckman, T. M., Tremonti, C., et al. 2003, *MNRAS*, **346**, 1055
- Kennicutt, R. C. 1998, *ARA&A*, **36**, 189
- Kriek, M., & Conroy, C. 2013, *ApJL*, **775**, L16
- Kriek, M., Labbé, I., Conroy, C., et al. 2010, *ApJL*, **722**, L64
- Kriek, M., Price, S. H., Conroy, C., et al. 2019, *ApJL*, **880**, L31
- Kriek, M., van Dokkum, P. G., Labbé, I., et al. 2009, *ApJ*, **700**, 221
- Leja, J., Johnson, B. D., Conroy, C., et al. 2019, *ApJ*, **877**, 140
- Leja, J., Johnson, B. D., Conroy, C., van Dokkum, P. G., & Byler, N. 2017, *ApJ*, **837**, 170
- Lilly, S. J., Carollo, C. M., Pipino, A., Renzini, A., & Peng, Y. 2013, *ApJ*, **772**, 119
- Lower, S., Narayanan, D., Leja, J., et al. 2020, *ApJ*, **904**, 33
- Magdis, G. E., Gobat, R., Valentino, F., et al. 2021, *A&A*, **647**, A33
- Maiolino, R., Cirasuolo, M., Afonso, J., et al. 2020, *Msngr*, **180**, 24
- Mandal, A., Mukherjee, D., Federrath, C., et al. 2021, *MNRAS*, **508**, 4738
- Mannucci, F., Cresci, G., Maiolino, R., Marconi, A., & Gnerucci, A. 2010, *MNRAS*, **408**, 2115
- Martig, M., Bournaud, F., Teyssier, R., & Dekel, A. 2009, *ApJ*, **707**, 250
- Martí-Vidal, I., Vlemmings, W. H. T., Muller, S., & Casey, S. 2014, *A&A*, **563**, A136
- McDermid, R. M., Alatalo, K., Blitz, L., et al. 2015, *MNRAS*, **448**, 3484
- Meidt, S. E., Schinnerer, E., García-Burillo, S., et al. 2013, *ApJ*, **779**, 45
- Muzzin, A., Marchesini, D., Stefanon, M., et al. 2013, *ApJ*, **777**, 18
- Narayanan, D., Krumholz, M. R., Ostriker, E. C., & Hernquist, L. 2012, *MNRAS*, **421**, 3127
- Noeske, K. G., Weiner, B. J., Faber, S. M., et al. 2007, *ApJL*, **660**, L43
- Pacifici, C., Kassin, S. A., Weiner, B. J., et al. 2016, *ApJ*, **832**, 79
- Pillepich, A., Springel, V., Nelson, D., et al. 2018, *MNRAS*, **473**, 4077
- Piotrowska, J. M., Bluck, A. F. L., Maiolino, R., Concas, A., & Peng, Y. 2020, *MNRAS*, **492**, L6
- Rodríguez Montero, F., Davé, R., Wild, V., Anglés-Alcázar, D., & Narayanan, D. 2019, *MNRAS*, **490**, 2139
- Rowlands, K., Wild, V., Nesvadba, N., et al. 2015, *MNRAS*, **448**, 258
- Saintonge, A., Kauffmann, G., Kramer, C., et al. 2011a, *MNRAS*, **415**, 32
- Saintonge, A., Kauffmann, G., Wang, J., et al. 2011b, *MNRAS*, **415**, 61
- Saintonge, A., Tacconi, L. J., Fabello, S., et al. 2012, *ApJ*, **758**, 73
- Sargent, M. T., Daddi, E., Bournaud, F., et al. 2015, *ApJL*, **806**, L20
- Schaye, J., Crain, R. A., Bower, R. G., et al. 2015, *MNRAS*, **446**, 521
- Schlaflly, E. F., Meisner, A. M., & Green, G. M. 2019, *ApJS*, **240**, 30
- Schruba, A., Leroy, A. K., Walter, F., et al. 2011, *AJ*, **142**, 37
- Scoville, N., Lee, N., Vanden Bout, P., et al. 2017, *ApJ*, **837**, 150
- Setton, D. J., Bezanson, R., Suess, K. A., et al. 2020, *ApJ*, **905**, 79
- Smercina, A., Smith, J. D. T., Dale, D. A., et al. 2018, *ApJ*, **855**, 51
- Spilker, J., Bezanson, R., Barišić, I., et al. 2018, *ApJ*, **860**, 103
- Straatman, C. M. S., Glazebrook, K., Kacprzak, G. G., et al. 2017, *ApJ*, **839**, 57
- Su, K.-Y., Hopkins, P. F., Bryan, G. L., et al. 2021, *MNRAS*, **507**, 175
- Suess, K. A., Bezanson, R., Spilker, J. S., et al. 2017, *ApJL*, **846**, L14
- Suess, K., Kriek, M., & Bezanson, R. 2021, arXiv:2111.14878
- Tacchella, S., Conroy, C., Faber, S. M., et al. 2021, arXiv:2102.12494
- Tacconi, L. J., Genzel, R., Neri, R., et al. 2010, *Natur*, **463**, 781
- Tacconi, L. J., Genzel, R., Saintonge, A., et al. 2018, *ApJ*, **853**, 179
- Tacconi, L. J., Genzel, R., & Sternberg, A. 2020, *ARA&A*, **58**, 157
- Tacconi, L. J., Neri, R., Genzel, R., et al. 2013, *ApJ*, **768**, 74
- Takada, M., Ellis, R. S., Chiba, M., et al. 2014, *PASJ*, **66**, R1
- Thomas, D., Maraston, C., Bender, R., & Mendes de Oliveira, C. 2005, *ApJ*, **621**, 673
- Tomczak, A. R., Quadri, R. F., Tran, K.-V. H., et al. 2014, *ApJ*, **783**, 85
- van de Sande, J., Kriek, M., Franx, M., et al. 2011, *ApJL*, **736**, L9
- van de Sande, J., Kriek, M., Franx, M., et al. 2013, *ApJ*, **771**, 85
- van der Wel, A., Bezanson, R., D'Eugenio, F., et al. 2021, *ApJS*, **256**, 44
- van der Wel, A., Noeske, K., Bezanson, R., et al. 2016, *ApJS*, **223**, 29
- Virtanen, P., Gommers, R., Oliphant, T. E., et al. 2020, *Nature Methods*, **17**, 261
- Waskom, M. L. 2021, *JOSS*, **6**, 3021
- Weinberger, R., Springel, V., Hernquist, L., et al. 2017, *MNRAS*, **465**, 3291
- Whitaker, K. E., Kriek, M., van Dokkum, P. G., et al. 2012a, *ApJ*, **745**, 179
- Whitaker, K. E., van Dokkum, P. G., Brammer, G., & Franx, M. 2012, *ApJL*, **754b**, L29
- Whitaker, K. E., Williams, C. C., Mowla, L., et al. 2021, *Natur*, **597**, 485
- Wild, V., Almaini, O., Dunlop, J., et al. 2016, *MNRAS*, **463**, 832
- Wild, V., Taj Aldeen, L., Carnall, A., et al. 2020, *MNRAS*, **494**, 529
- Williams, C. C., Spilker, J. S., Whitaker, K. E., et al. 2021, *ApJ*, **908**, 54
- Wu, P.-F., van der Wel, A., Bezanson, R., et al. 2020, *ApJ*, **888**, 77
- Wu, P.-F., van der Wel, A., Gallazzi, A., et al. 2018, *ApJ*, **855**, 85
- Yano, M., Kriek, M., van der Wel, A., & Whitaker, K. E. 2016, *ApJL*, **817**, L21
- Young, L. M., Bureau, M., Davis, T. A., et al. 2011, *MNRAS*, **414**, 940
- Young, L. M., Scott, N., Serra, P., et al. 2014, *MNRAS*, **444**, 3408
- Zabludoff, A. I., Zaritsky, D., Lin, H., et al. 1996, *ApJ*, **466**, 104