



Universiteit
Leiden
The Netherlands

Testing afterglow models of FRB 200428 with early post-burst observations of SGR 1935 + 2154

Cooper, A.J.; Rowlinson, A.; Wijers, R.A.M.J.; Bassa, C.; Gourdji, K.; Hessels, J.; ... ; Veen, S. ter

Citation

Cooper, A. J., Rowlinson, A., Wijers, R. A. M. J., Bassa, C., Gourdji, K., Hessels, J., ... Veen, S. ter. (2022). Testing afterglow models of FRB 200428 with early post-burst observations of SGR 1935 + 2154. *Monthly Notices Of The Royal Astronomical Society*, 517(4), 5483-5495. doi:10.1093/mnras/stac2951

Version: Publisher's Version
License: [Creative Commons CC BY 4.0 license](#)
Downloaded from: <https://hdl.handle.net/1887/3561856>

Note: To cite this publication please use the final published version (if applicable).

Testing afterglow models of FRB 200428 with early post-burst observations of SGR 1935 + 2154

A. J. Cooper^{1,2}★, A. Rowlinson^{1,2}, R. A. M. J. Wijers¹, C. Bassa², K. Gourdji³, J. Hessels^{1,2},
A. J. van der Horst⁴, V. Kondratiev², D. Michilli⁵, Z. Pleunis⁶, T. Shimwell^{2,7} and S. ter Veen²

¹*Anton Pannekoek Institute, University of Amsterdam, Postbus 94249, 1090 GE Amsterdam, the Netherlands*

²*ASTRON, The Netherlands Institute for Radio Astronomy, Oude Hoogeveensedijk 4, 7991 PD, Dwingeloo, the Netherlands*

³*Centre for Astrophysics and Supercomputing, Swinburne University of Technology, Hawthorn VIC 3122, Australia*

⁴*Department of Physics, George Washington University, 725 21st St NW, Washington, DC 20052, USA*

⁵*MIT Kavli Institute for Astrophysics and Space Research, 77 Massachusetts Ave, Massachusetts Institute of Technology, Cambridge, MA 02139, USA*

⁶*Dunlap Institute for Astronomy & Astrophysics, University of Toronto, 50 St. George Street, Toronto, ON M5S 3H4, Canada*

⁷*Leiden Observatory, Leiden University, P.O. Box 9513, 2300 RA Leiden, the Netherlands*

Accepted 2022 October 11. Received 2022 October 11; in original form 2022 September 19

ABSTRACT

We present LOFAR imaging observations from the April/May 2020 active episode of magnetar SGR 1935 + 2154. We place the earliest radio limits on persistent emission following the low-luminosity fast radio burst FRB 200428 from the magnetar. We also perform an image-plane search for transient emission and find no radio flares during our observations. We examine post-FRB radio upper limits in the literature and find that all are consistent with the multiwavelength afterglow predicted by the synchrotron maser shock model interpretation of FRB 200428. However, early optical observations appear to rule out the simple versions of the afterglow model with constant-density circumburst media. We show that these constraints may be mitigated by adapting the model for a wind-like environment, but only for a limited parameter range. In addition, we suggest that late-time non-thermal particle acceleration occurs within the afterglow model when the shock is no longer relativistic, which may prove vital for detecting afterglows from other Galactic FRBs. We also discuss future observing strategies for verifying either magnetospheric or maser shock FRB models via rapid radio observations of Galactic magnetars and nearby FRBs.

Key words: acceleration of particles – fast radio bursts – stars: individual: SGR1935 + 2154 – stars: magnetars – stars: neutron.

1 INTRODUCTION

Fast radio bursts (FRBs) are bright, millisecond duration flashes of coherent radio emission (Lorimer et al. 2007; Thornton et al. 2013; Petroff et al. 2016; The CHIME/FRB Collaboration 2021). Over 20 FRBs have now been localized¹ to their host galaxies (Chatterjee et al. 2017; Bannister et al. 2019; Ravi et al. 2019; Bhandari et al. 2020; Heintz et al. 2020; Marcote et al. 2020; Nimmo et al. 2022), confirming that these bursts are (usually) of extragalactic origin, as was already implied by large dispersion measures. Two varieties of FRB source have been identified, categorized by whether the source is observed to repeat or not. Bursts observed from these different categories of sources appear to have different properties (Pleunis et al. 2021), which may be suggestive of different progenitors or emission mechanisms.

Since the discovery of FRBs, many theoretical models have been presented to explain the bursts involving various astrophysical sources and radiation mechanisms (Platts et al. 2019). Many of these models focus on highly magnetized neutron stars as prime

progenitor candidates for a few reasons. First, the millisecond emission time-scale implies compact emission regions $R \approx ct_{\text{FRB}} \Gamma^2 \approx 10^6 t_{\text{FRB},-3} \Gamma_0^2$ cm, where Γ is the bulk Lorentz factor of the emission region towards the observer and we have used the convenient notation $X_n \equiv X/10^n$. In addition, the luminosity of FRBs (≈ 8 orders of magnitude larger than giant millisecond radio pulses in the Milky Way) might imply they are powered by the large magnetic energy reservoir in the magnetospheres of highly magnetized neutron stars known as magnetars. For decades, magnetars within our Galaxy have been observed to emit spontaneous X-ray bursts, which appear to follow a similar wait time distribution and luminosity function to FRBs (e.g. SGR 1900 + 14, Göğüş et al. 1999; SGR 1806-20, Göğüş et al. 2000; FRB 121102, Gourdji et al. 2019), as well as periods of activity and quiescence. Finally, other observed properties of FRBs can be explained within a magnetized neutron star progenitor model: large rotation measures (Michilli et al. 2018), downward drifting sub-bursts (e.g. Hessels et al. 2019), and polarization angle swings across bursts (Luo et al. 2020).

The exact radiation mechanism within a magnetar progenitor framework is also not understood, with the two primary classes of models distinguished by where the FRB is emitted: close to the surface of the neutron star in magnetospheric models (Katz 2016; Kumar, Lu & Bhattacharya 2017; Lu & Kumar 2018; Wadiasingh &

* E-mail: a.j.cooper@uva.nl

¹<http://frbhosts.org>

Timokhin 2019; Kumar & Bošnjak 2020; Lyutikov & Popov 2020), and far away from the surface in models which rely on maser emission in magnetized shocks (Beloborodov 2017; Metzger, Margalit & Sironi 2019; Beloborodov 2020). Both classes of radiation model appear to explain the basic properties of FRBs.

Magnetospheric models of FRBs require the coherent radiation of accelerated particles from close to the neutron star surface, through e.g. curvature emission (Kumar et al. 2017). It has been shown that many of the observed properties of FRBs discussed above can be reproduced within this model (e.g. Lu & Kumar 2018; Wang et al. 2019, 2022; Cooper & Wijers 2021; Yang & Zhang 2021). Coherent curvature radiation requires spatial inhomogeneities in the particle distribution (‘clumps’ or ‘bunches’), which may be created during the particle acceleration phase through the two-stream instability (Cheng & Ruderman 1977; Usov 1987; Lu & Kumar 2018; Kumar, Gill & Lu 2022). However, whether clumps form on the required time-scales is less clear (see Melrose, Rafat & Mastrano 2021 and references therein for critical discussion of curvature radiation as a pulsar mechanism), and it has further been suggested that the FRBs produced within the magnetosphere may fail to propagate to an observer (Beloborodov 2021).

The maser shock model appears to be a theoretically robust way in which coherent emission can be produced, as the maser has been shown to emit narrow-band radio emission in 1D and 3D particle-in-cell (PIC) simulations of magnetized, relativistic shocks (Plotnikov & Sironi 2019; Sironi et al. 2021). Furthermore, such emission could be ubiquitous in nature wherever such shocks occur, and non-magnetar progenitors have been invoked to explain observed periodicity in repeating sources (Sridhar et al. 2021). Such ubiquity could account for the high observed volumetric rate of FRBs (Ravi et al. 2019; Lu & Kumar 2019; Luo et al. 2020). However, the specific properties of recently reported FRBs are difficult to explain in a maser shock scenario. In particular, the microsecond variability of FRB 20180916B (Nimmo et al. 2022) and the 30 ms separation between distinct sub-bursts in Galactic FRB 200428 (CHIME/FRB Collaboration 2020) appear to be contradictory to the large length scales associated with the synchrotron maser emission region.

In magnetar models of FRBs, multiwavelength, or multimessenger counterparts to emission provide invaluable observing opportunities with which to distinguish between theoretical models. Both magnetospheric and maser shock models predict a high-energy counterpart to FRB emission (Metzger et al. 2019; Cooper & Wijers 2021; Yang & Zhang 2021), however the large distances to extragalactic FRBs mean predicted X-ray/gamma-ray fluxes are below the detection threshold of current instrumentation. The first version of the synchrotron maser shock model (Lyubarsky 2014), in which relativistic magnetar hyperflares ($U_B \approx 10^{48}$ erg) interact with pulsar wind nebula, predicted a very high-energy TeV component to FRBs. In Metzger et al. (2019) this model was adapted to consider interaction with a wind-like environment, or a mildly relativistic baryonic shell emitted from a previous flare. Using these models, Metzger et al. (2019) predict a multiwavelength afterglow due to thermal synchrotron emission, based on gamma-ray burst (GRB) afterglow models (Sari, Piran & Narayan 1998). Notably, as the shocks required for the maser are highly magnetized & relativistic, magnetic field lines are compressed such that shocks are quasi-perpendicular. For this reason, shock accelerated non-thermal electrons are not necessarily expected (Sironi & Spitkovsky 2009), in contrast to what is observed in other astrophysical shocks. Observationally, identifying a multiwavelength afterglow could confirm the synchrotron maser model of FRBs, since magnetospheric models are not predicted to produce them. Persistent radio counterparts have been identified in two localized

repeating FRB sources: FRB 121102 (Chatterjee et al. 2017) and FRB 20190520B (Niu et al. 2022). The origin of the persistent radio emission is not definitively known, the spectra and lack of variability may imply the compact source is an AGN that resides relatively close to the FRB source. Follow-up of all other localized FRBs have yielded only upper limits to afterglow counterparts (Bhandari et al. 2018), although these are generally not constraining for the maser shock model.

In this work, we present Low Frequency Array (LOFAR; van Haarlem et al. 2013) imaging results of SGR 1935 + 2154, and discuss more broadly the multiwavelength afterglow predicted after FRB 200428 in the maser shock model. In Section 2, we discuss the behaviour of magnetar SGR 1935 + 2154 before and during the 2020 active phase. In Section 3, we present the LOFAR imaging results, where we perform a search for bright radio bursts observable in the image-plane (see Bailes et al. 2021 for the time-domain search, as well as other radio observations) and provide limits for persistent low-frequency emission. In Section 4, we interpret the LOFAR results and suggest observing strategies for verification of FRB emission mechanisms. In Section 5, we discuss the synchrotron maser shock model and its application to FRB 200428. In particular, we find that early time upper limits appear to rule out simple versions of the maser shock model as applied to the Galactic FRB. We further discuss extensions to the model including non-thermal radiation at late times, and make afterglow predictions for Galactic and nearby, extragalactic FRBs. We conclude our findings in Section 6.

2 SGR 1935 + 2154

Galactic magnetar SGR 1935 + 2154 was discovered in July 2014 through a series of short X-ray bursts (Cummings et al. 2014; Lien et al. 2014; Israel et al. 2016), and past X-ray activity from 2011 was found in an archival search (Campana et al. 2014). Further observations in 2015 with *CHANDRA* & *XMM-Newton* found many more X-ray bursts (Israel et al. 2016), as well as the first intermediate flares (Kozlova et al. 2016). This led to the first measurements of a spin period and spin-down rate of $P = 3.245$ s (Israel et al. 2016) and $\dot{P} = 1.43 \times 10^{-11}$, relatively typical values for a Galactic magnetar. Using these measured values, Israel et al. (2016) derive the characteristic age and surface magnetic field

$$\text{of the magnetar as: } \tau_c = \frac{P}{2\dot{P}} \approx 3.6 \text{ kyr and } B_s = \sqrt{\frac{3c^3 I P \dot{P}}{8\pi^2 R_{\text{NS}}^6}} \approx$$

2.2×10^{14} G. Interestingly, \dot{P} shows a decreasing trend, with a measured negative $\ddot{P} = -3.5(7) \times 10^{-19} \text{ s}^{-1}$ observed. The source position was determined to be at R.A. = 19: 34: 55.5978, Dec. = 21: 53: 47.7864, with an accuracy of 0.6 arcsec (90 per cent confidence, Israel et al. 2016; see also the discovery of an infrared counterpart by Levan, Kouveliotou & Fruchter 2018 at Galactic coordinates $b = 293.73128$, $l = 21.896608$). The magnetar was associated with a supernova remnant SNR G57.2 + 0.8 by Gaensler (2014), with a full multiwavelength radio study of continuum and persistent emission presented in Kothes et al. (2018). The authors find that the age of the SNR is approximately 41 kyr, i.e. much older than the characteristic age of the magnetar derived from its spin properties (Israel et al. 2016). The distance to SNR G57.2 + 0.8 was revised by Zhou et al. (2020), finding that the SNR and thus the magnetar is likely closer than the 10 kpc previously assumed, at just $D = 6.6 \pm 0.7$ kpc.

In late 2019 (Ambrosi et al. 2019; Wood, Bissaldi & Fermi GBM Team 2019) and especially in early 2020 (Palmer 2020; Younes et al. 2020), the magnetar entered a new phase of extreme activity. On the 2020 April 27 at 14:34:24, an extremely bright radio burst

was emitted from SGR 1935 + 2154. The burst was observed by the Survey for Transient Astronomical Radio Emission 2 (*STARE-2*; Bochenek et al. 2020b) an all-sky radio telescope (Bochenek et al. 2020a). It was also detected in the side lobes of *CHIME*, the Canadian Hydrogen Intensity Mapping Experiment telescope (*CHIME/FRB* Collaboration 2020). *STARE-2* detected the burst with a 1.281–1.468 GHz lower limit fluence of $1.5 \pm 0.3 \times 10^6$ Jy ms and a full-width half-maximum duration of 0.61 ms. *CHIME/FRB* observed two bursts separated by 28.9 ms, with 400–800 MHz durations of 0.585 ± 0.014 ms and 0.335 ± 0.007 ms, respectively, and average fluences of $0.48 \pm_{0.24}^{0.48} \times 10^6$ Jy ms and $0.22 \pm_{0.11}^{0.22} \times 10^6$ Jy ms. The lower observed fluence could be attributed to a steep spectral rise with frequency such that the burst is brighter at the *STARE-2* observing frequency. The implied luminosity was lower than extragalactic FRBs by ≈ 4 orders of magnitude, but was brighter than any other coherent radio emission observed from within our Galaxy with the possible exception of the brightest 2 MJy nanoshot observed from the Crab pulsar (Hankins & Eilek 2007). On 30th April, a weak radio burst (0.06 Jy ms) was reported by Zhang et al. (2020), and two further radio bursts were reported later by Kirsten et al. (2021) with fluences of 112 Jy ms and 24 Jy ms.

FRB 200428 occurred during a very active period of high-energy bursts from SGR 1935 + 2154, and many hundreds of short X-ray bursts were observed by NICER, Fermi, Swift, AGILE, INTEGRAL, Insight-HXMT, Konus-Wind, and other X-ray satellite observatories (for catalogues of X-ray bursts see Younes et al. 2020; Cai et al. 2022a; Li et al. 2022). Remarkably, an X-ray burst temporally coincident with the FRB was observed by four X-ray instruments: Integral (Mereghetti et al. 2020), Insight-HXMT (Li et al. 2021), AGILE (Tavani et al. 2021), and Konus-Wind (Ridnaia et al. 2021).

The X-ray/radio energy ratio of the radio burst was $\frac{E_X}{E_r} \approx 10^5$.

The radio-coincident X-ray burst had a harder spectrum than other bursts seen from SGR 1935 + 2154 in the same window of activity, extending to 250 keV (Mereghetti et al. 2020; Li et al. 2021) despite a relatively typical total fluence. The stark contrast is clear when the X-ray burst is compared to the many Fermi-GBM and NICER observed bursts in Younes et al. (2021). Within magnetospheric models of FRBs, this can plausibly be explained by a non-thermal counterpart observable only when an FRB is observed (Cooper & Wijers 2021), in addition to a regular short magnetar burst; or due to rapid relaxation of perturbed magnetic field lines (Yang et al. 2020). In the synchrotron maser shock picture of FRBs, the hard X-ray burst represents the initial peak of the afterglow emission that is observable at lower frequencies at later times (Margalit et al. 2020b).

3 LOFAR OBSERVATIONS

SGR 1935 + 2154 entered an active phase in April 2020.² We obtained Directors Discretionary Time observations using LOFAR, comprising of 14 h spread over three dates in 2020. The observation details are provided in Table 1. Each observation was obtained using the LOFAR High Band Antennas (HBA) and the full Dutch array (24 core stations and 14 remote stations), covering a frequency range of 120–168 MHz and a central frequency of 144 MHz. The frequency range is divided into 244 sub-bands with bandwidths of 195.3 kHz. The recorded data have a time resolution of 1s, and a frequency resolution of 64 channels per sub-band. These data were pre-processed using the standard methods for LOFAR.

Table 1. The observations attained of SGR 1935 + 2154. The ID numbers correspond to the IDs for the preprocessed data in the LOFAR Long Term Archive.² The duration of each target observation was 2 h, separated by 10 min calibrator observations of 3C 295.

Observation ID	Start Date (UTC)	Start Time (UTC)	Calibrator ID
L780243	2020 April 29	03:08	L780239
L780251	2020 April 29	05:20	L780247
L780651	2020 May 11	01:49	L780655
L780659	2020 May 11	04:01	L780655
L780667	2020 May 11	06:13	L780671
L797090	2020 October 21	14:30	L797092
L797094	2020 October 21	16:42	L797096

Table 2. The parameters used for imaging each observation set in WSCLEAN. Deep images are created using two different minimum baselines.

Parameter	Deep Images	Snapshot Images
Size (pixels)	4096 × 4096	2048 × 2048
Pixel Scale (arcsec)	5	5
-mgain	10	10
-j	10	10
Automatic masking of sources (σ)	10	10
Maximum baseline length (λ)	8000	8000
Minimum baseline length (λ)	0/1000	700
Number of frequency channels	6	1
Automatic threshold (σ)	3	3
Briggs weighting robustness	−0.5	−0.5
Number of iterations	10 000	10 000
Primary beam applied	True	False
Weighting rank filter	3	3
Clean border	0	0
Reorder visibilities	True	False
Fit restoring beam	True	True
Number of time intervals	1	180

To calibrate these LOFAR data, we used the PREFACTOR pipeline (version 3.2)³ developed by the LOFAR Radio Observatory with the default parameters and version 3.1 of the LOFAR software. The calibrator source chosen for these observations was 3C 48. As the field of SGR 1935 + 2154 is close to the A-Team source Cygnus A, we used the PREFACTOR demixing options to subtract the contribution of Cygnus A from the target data. The data were averaged to a time step of 8 s and a frequency resolution of 48.82 kHz (4 channels per subband).

3.1 Epoch imaging

All imaging of these data was conducted using WSCLEAN version 2.7.0⁴ (Offringa et al. 2014). The imaging settings used for these observations are outlined in Table 2. Deep images of the field were created for the three observation days separately, as well as a fully combined image. Two versions of the deep images were created; one with the full data set (Fig. 1) and the second with a minimum baseline length of 1000λ to remove the extended emission from the supernova remnant. These images were used to search for faint persistent emission from the location of SGR 1935+2154. No persistent emission

²<https://lta.lofar.eu>

³<https://github.com/lofar-astron/prefactor>

⁴<http://wsclean.sourceforge.net>

Table 3. The flux density measurements at the location of SGR 1935 + 2154 for each observing epoch and a combined image. The flux density measurements were obtained via a constrained flux density measurement using PYSE. The 3σ upper limits are obtained by measuring the rms noise in a region surrounding SGR 1935 + 2154.

Observation date	Flux density (mJy)	3σ upper limit (mJy per beam)
2020 April 29	2.2 ± 5.8	9.5
2020 May 11	1.8 ± 4.3	6.8
2020 October 21	0.7 ± 4.6	7.9
All	1.5 ± 3.8	6.0

was detected at the location of SGR 1935+2154 and the limits obtained for persistent emission from SGR 1935 + 2154 are provided in Table 3. We provide a constrained flux density measurement at the position of SGR 1935 + 2154 obtained using PYSE (Carbone et al. 2018) and assuming the restoring beam parameters. Additionally, we measure the rms noise in a 30×30 pixel box centred on the position of SGR 1935 + 2154 and use the rms noise to calculate a 3σ upper limit.

3.2 Snapshot imaging

To search for bright pulses from SGR 1935+2154, we also imaged the observations on shorter snapshot time-scales. We integrated across the observed frequencies to create a single image. As SGR 1935 + 2154 has a high dispersion measure (332.7 pc cm^{-3} ; CHIME/FRB Collaboration 2020), it takes time for a dispersed pulse to traverse the bandwidth of the HBA observation. This time is calculated using

$$\tau_{\text{DM}} = 8.3 \Delta\nu DM \nu^{-3} \mu\text{s}, \quad (1)$$

where τ_{DM} is the dispersion delay in seconds, $\Delta\nu$ is the bandwidth in MHz, ν is the observing frequency in GHz, and DM is the dispersion measure in pc cm^{-3} (Taylor & Cordes 1993). For these LOFAR observations, we find that the dispersion delay corresponds to 44 s. Therefore, we create snapshot images of 40 s in duration (as the data are integrated to 8 s intervals) and obtained 1081 images in total. We note that we could conduct de-dispersion imaging (e.g. Anderson et al. 2021; Tian et al. 2022) to obtain deeper constraints on very short duration dispersed pulses such as FRBs. However, beam-formed observations were also attained for these observations and very short duration dispersed pulses would be optimally detected in those observations (Bailes et al. 2021). Our snapshot images targeted longer duration flares that may not have been detectable in the beam-formed searches. The snapshot images were created using the settings given in Table 2. A minimum baseline length of 700λ was used so that the extended emission from the supernova remnant was not visible in the snapshot images as this leads to a better constraint on emission from a point source at the location of SGR 1935 + 2154.

Following the method outlined in Rowlinson et al. (2022), we correct for any systematic flux density variations between the snapshot images produced by assuming that the majority of sources are not variable. As SGR 1935 + 2154 is in the centre of the field, we consider sources out to a radius of 1 degree (covering the inner 50 per cent of the full width half maximum of the LOFAR HBA beam at 150 MHz⁵). We extract all sources in the first 40 s snapshot

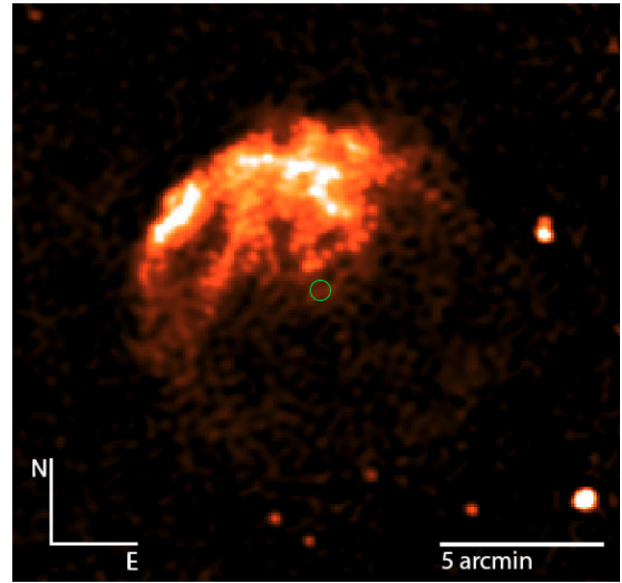


Figure 1. LOFAR image of SGR 1935+2154 (denoted by the green circle) and the surrounding remnant SNR G57.2 + 0.8. This image makes use of the full data sets from all epochs.

image within that radius using PYSE with a detection threshold of 8σ and confirm they are point sources by visual inspection (one source was rejected as it consisted of 2 blended point sources). These selection criteria resulted in 10 sources that were input into the LOFAR Transients Pipeline (TRAP; Swinbank et al. 2015) using the monitoring list capability. A flux density correction factor is then calculated and applied to each individual image.⁶ The average flux density correction is ~ 6 per cent.

Following the correction for systematic flux density variations, we conduct image quality control by measuring the rms noise in the inner eighth of each image and plot a histogram of these data in Fig. 2. We fit the typical rms noise distribution for the images with a Gaussian distribution to give an average rms value of $17.4^{+2.6}_{-2.3}$ mJy per beam. Any image with an rms noise that is more than 3σ deviant from the average rms value is rejected. This analysis rejected 110 images, corresponding to 10 per cent of the sample.

To search for emission at the location of SGR 1935 + 2154, we processed the images using TRAP and monitor the position of SGR 1935 + 2154 to provide a flux measurement in each epoch. We used the standard TRAP settings apart from using a 4σ detection threshold and setting the force beam parameter to True. As we are focusing on the location of SGR 1935+2154, which lies at the centre of the field, we use a small extraction radius of 120 pixels (corresponding to 10 arcmin). We detect no emission at the location of SGR 1935 + 2154. In Fig. 3, we plot a histogram of the flux density values measured at the location of SGR 1935+2154 and fit the distribution with a Gaussian distribution. The fitted Gaussian distribution gives an average flux density measurement per epoch of 13 ± 0.7 mJy with a standard deviation of 23 mJy, corresponding to a 3σ upper limit for persistent point source emission of 82 mJy. This suggests the flux at the location of SGR 1935 + 2154 is significantly non-zero, which by inspection of Fig. 1, can be attributed to extended emission from the surrounding SNR.

⁵<http://old.astron.nl/radio-observatory/astronomers/lofar-imaging-capabilities-sensitivity/lofar-imaging-capabilities/lofa>

⁶using the script https://github.com/transientskp/TrAP_tools/blob/master/exampleScripts/correctSystematicFluxOffset.py

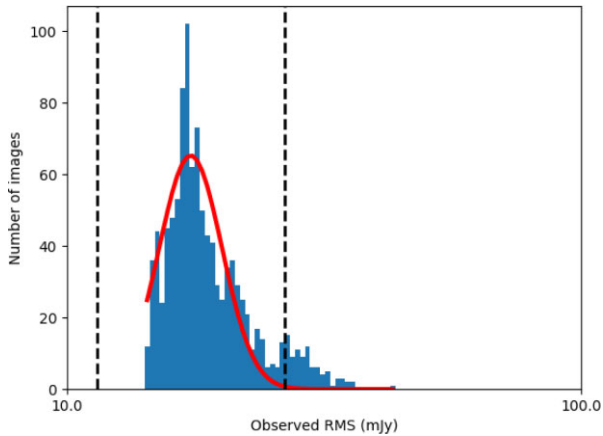


Figure 2. Histogram of the rms values of the inner eighth of all the 40-s snapshot images used. The values are fit with a Gaussian distribution shown by the solid red curve. All images that are $>3\sigma$ deviant from the average rms are rejected and the thresholds are shown with the black dashed lines.

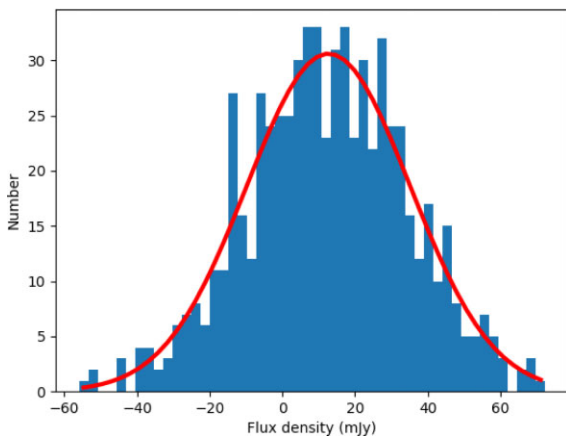


Figure 3. A histogram of the flux density values for SGR 1935 + 2154 obtained by conducting a constrained fit at the location in each snapshot image.

4 INTERPRETATION OF LOFAR OBSERVATIONS & DISCUSSION

4.1 Simultaneous radio & X-ray limits on transient flares

Low-frequency image-plane searches for radio transients have recently yielded detections of new types of coherent radio transient with long periods (Caleb et al. 2022; Hurley-Walker et al. 2022). These new sources have inferred surface magnetic fields consistent with magnetars. We have performed a search for transient low-frequency emission in snapshot images on time-scales of 40 s (i.e. the dispersion delay across the LOFAR band of a burst from SGR 1935 + 2154) and find no transient emission at the location of the magnetar. Given the non-detections in the LOFAR time-domain search presented in Bailes et al. (2021), coupled with the fact that second-hour time-scale transient radio emission has not been yet observed from Galactic magnetars, we did not necessarily expect a detection from this search.

One of the primary motivators of radio observations of magnetars in active phases, other than detecting period-modulated radio pulses from magnetars (see Section 4.2), is to obtain constraints on the X-ray/radio luminosity ratio during X-ray magnetar bursts. In Fig. 4 we show the times of LOFAR and AARTFAAC observations,

overplotted with high-energy bursts observed by various X-ray observatories. One high-energy burst observed by Insight-HXMT (burst number 56 in Cai et al. 2022a) occurred during our LOFAR observations on 11/5/2022. The burst had a reported 1–250 keV flux of $2.29^{+0.18}_{-0.17} \times 10^{-7} \text{ erg s}^{-1} \text{ cm}^{-2}$ (Cai et al. 2022b). The duration of this burst was 0.093, 0.06, and 0.076 s in the high-energy (HE), medium-energy (ME) and low-energy (LE) detectors respectively, resulting in a X-ray fluence of approximately $2 \times 10^{-8} \text{ erg cm}^{-2}$ across 1–250 keV. The time-domain search for LOFAR bursts in this observation provided upper limits to the spectral fluence of a 1 ms width radio burst of 0.075 Jy ms (Bailes et al. 2021). Assuming the bandwidth of a typical radio burst is $\delta\nu = 10^9 \text{ Hz}$ (i.e. similar to the observed bandwidth of FRB 200428), we find a fluence limit of $7.5 \times 10^{-18} \text{ erg cm}^{-2}$, resulting in a minimum X-ray/radio fluence ratio of $\frac{\mathcal{F}_X}{\mathcal{F}_{r,144\text{MHz}}} > 3 \times 10^9$. This ratio is very constraining, as shock models of FRBs typically predict luminosity ratios of approximately 10^5 (Margalit et al. 2020b).

The Amsterdam-ASTRON Radio Transient Facility And Analysis Centre (AARTFAAC), is an all-sky radio transient monitor operating in parallel with LOFAR (Prasad et al. 2016), sensitive to bright transient and variable emission at low-frequencies. AARTFAAC was operating with 16 subbands spanning 60.15–64.84 MHz during an X-ray burst observed from SGR 1935 + 2154 by Insight-HXMT (burst 19 in Cai et al. 2022a). The dispersion delay from SGR 1935 + 2154 from the Insight-HXMT X-ray band to the AARTFAAC band is $\tau_{\text{X-ray}} = 324 \text{ s}$, and the in-band dispersion delay across 60.15–64.84 MHz is $\tau_{\text{in-band}} = 54 \text{ s}$. No radio transient was detected at the time of the X-ray burst. We can derive a meaningful fluence sensitivity by comparing the weakest bursts observed in Kuiack et al. (2020) from PSR B0950 + 08 ($4 \times 10^4 \text{ Jy ms}$), and rescaling by a factor proportional to the square root of the ratio of in-band dispersion delay from the different sources to account for a loss of sensitivity due to pulse smearing: $\left(\frac{\tau_{\text{PSR B0950+08}}}{\tau_{\text{SGR 1935+2154}}}\right)^{1/2} \approx 11$. We therefore find a conservative fluence sensitivity of approximately 10^6 Jy ms for the dispersed burst from SGR 1935 + 2154, and we can constrain the fluence ratio of the X-ray burst observed by Insight-HXMT to $\frac{\mathcal{F}_X}{\mathcal{F}_r} \gtrsim 10^2$.

Some magnetospheric models of FRBs predict that FRBs are always produced with short magnetar X-ray bursts, but are only beamed towards the observer in a small number of cases $\approx \Gamma^{-1}$ (Lu, Kumar & Zhang 2020). Here, Γ is the Lorentz factor of the emitting set of particles, and is on the order of 50–1000 (Kumar et al. 2017). In the context of these models, coordinated radio and X-ray observations of active magnetars will either lead to the detection of another Galactic FRB, or build up a sample of radio-quiet magnetar bursts. If such a sample is significantly large (i.e. $N \gg \Gamma$) it can be used to statistically constrain Γ , and thus magnetospheric models of FRBs more broadly. Using single-dish 25-m class telescopes, Kirsten et al. (2021) were able to place strong limits on radio counterparts to 59 high-energy bursts from SGR 1935 + 2154, effectively ruling out values of $\Gamma \ll 50$. Furthermore, Lin et al. (2020a) found that 29 Fermi-GBM bursts were co-observed during one observing epoch with the FAST radio telescope, providing stringent upper limits on the radio/X-ray ratio of these bursts. The caveats with pursuing these kinds of simultaneous observation are two-fold. First, we know that (repeating) FRBs generally have a narrow bandwidth (e.g. The CHIME/FRB Collaboration 2021) and therefore observing at a one frequency does not necessarily preclude the emission of a radio burst at another. Secondly, magnetar X-ray bursts are stochastic and

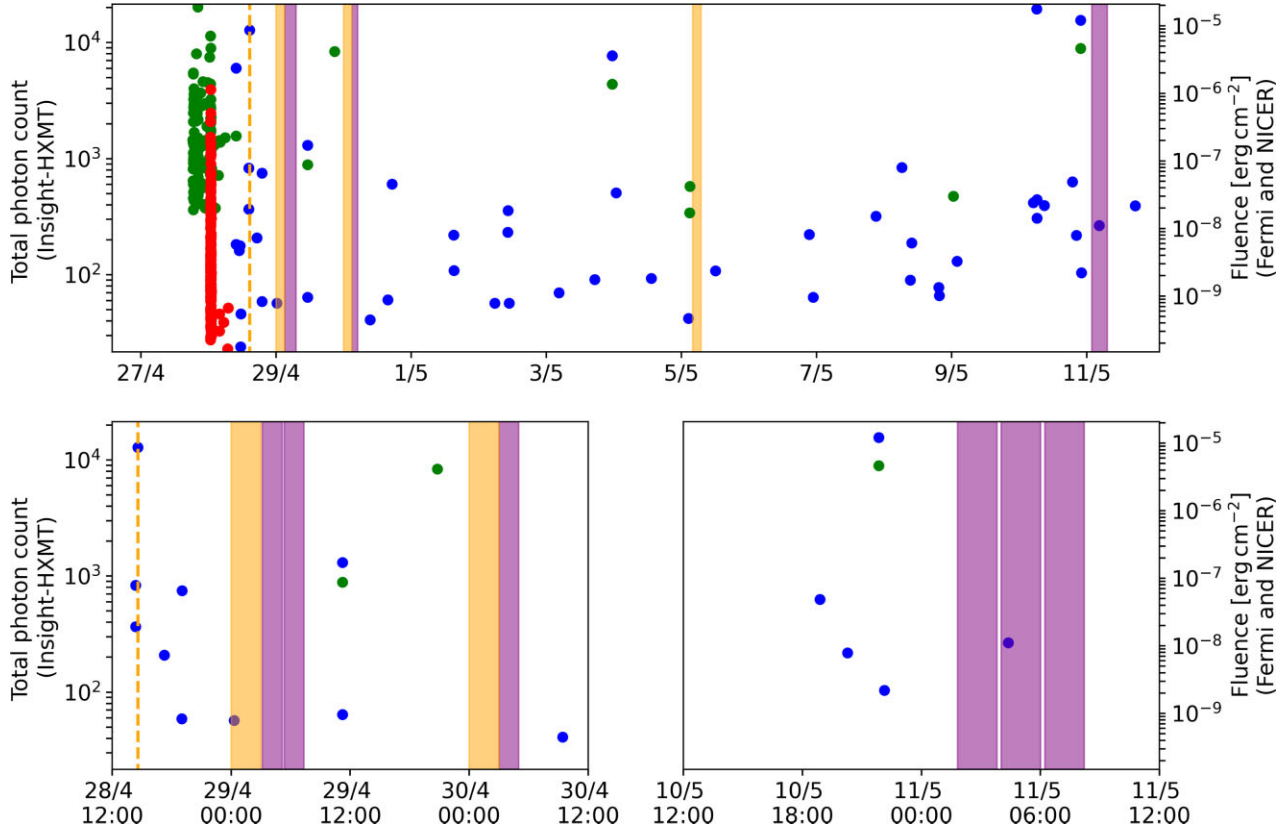


Figure 4. We show the times of LOFAR observations (denoted by orange background hue) and AARTFAAC (Prasad et al. 2016) observations (denoted by purple background hue) of SGR 1935 + 2154 during the April/May 2020 active episode. We note that observations on the April 30 were beamformed only, and therefore do not appear in our imaging analysis Section 3. The time of the FRB as observed by CHIME & STARE-2 is denoted by the vertical dashed orange line. We also show counts/fluence and times of X-ray bursts as observed by Insight-HXMT (blue dots; Cai et al. 2022a), NICER (red dots; Younes et al. 2020), and Fermi-GBM (green dots; Lin et al. 2020b). Insight-HXMT data points refer to the total number of photon counts in the burst across all detectors (left y-axis), whereas NICER and Fermi-GBM data points refer to the total fluence of the burst in the 0.5–10 keV and 8–200 keV bands, respectively. The second row of plots show timeline zoomed in on LOFAR observations, where 10 min calibrator scans are visible, as are the two X-ray bursts for which we were on source. Interpretation of these observations is discussed in Section 4.

difficult to predict. Building such a sample will require either a pipeline from X-ray detections to radio observatories that enable rapid target-of-opportunity observations⁷; or long-term observations of magnetars with underutilized radio telescopes. For the latter, we note that achieving a constraint of $\frac{L_X}{L_r} > 10^5$ for a typical Galactic magnetar bursts only requires fluence limits on the order of 100–1000 Jy ms, and therefore does not require time on extremely sensitive radio instruments.

4.2 Radio-quiet magnetars as FRB sources

Pulsed, persistent emission is observed from a sub-population of magnetars which are radio-loud (Camilo et al. 2006, 2007; Kramer et al. 2007; Levin et al. 2010; Anderson et al. 2012; Esposito et al. 2020; Lower et al. 2020). In Rea et al. (2012), the authors describe the Fundamental Plane of radio magnetars, suggesting that in general radio-loud magnetars have quiescent 2–10 keV X-ray luminosities

⁷A possible observation strategy would be to use low-frequency radio telescopes and utilize the dispersion delay of radio emission to probe radio bursts associated with rapidly reported X-ray bursts. Alternatively, one may attempt to identify X-ray burst storms to trigger observations when more X-ray bursts are expected within the observation duration.

below their rotational energy loss rate $L_{\text{rot}} = 4\pi^2 I P \dot{P}^{-3}$, where I is the neutron star moment of inertia. The conjecture that enhanced X-ray emission precludes radio emission (or vice versa) is supported by observations of PSR J1119-6127, a radio pulsar with a high magnetic field where radio pulses shut off when X-ray bursts are observed (Archibald et al. 2017). Similar X-ray and radio mode switches have also been observed in older (but non-recycled) pulsars with modest inferred magnetic fields, including PSR B0943 + 10 (Hermsen et al. 2013) and PSR B0823 + 26 (Hermsen et al. 2018).

SGR 1935 + 2154 is thus far the only Galactic magnetar to have produced a bright radio burst, and the radio limits presented here and in Bailes et al. (2021) suggest SGR 1935 + 2154 is not a radio-loud magnetar. This is consistent with the Fundamental Plane of Rea et al. (2012) given the observed persistent X-ray luminosity of SGR 1935 + 2154. Three individual radio bursts/pulses were reported after the FRB (Zhang et al. 2020; Kirsten et al. 2021), albeit at lower luminosities. The one-off radio bursts and pulses strongly suggest that high X-ray bursting activity does not disrupt a magnetar’s ability to produce *transient* radio bursts, and possibly that persistent pulsed radio emission precludes transient bursts. The association between radio-quiet magnetars and FRB-emitting magnetars is more intriguing when considered in the context of the recent discovery of highly magnetized neutron stars which exist beyond the canonical pulsar death lines (e.g. Ruderman & Sutherland

1975; Chen & Ruderman 1993) in $\dot{P} - P$ phase space (Caleb et al. 2022; see also Hurley-Walker et al. 2022). The implication that a neutron star's inability to produce typical pulsar-like emission could be a prerequisite for a FRB production may hint that pulsed radio emission observed from (ultra) long period magnetars is powered by a different emission mechanism than pulsars, possibly more similar to the radiation mechanism that powers FRBs. Understanding the particle acceleration and coherent radiation mechanism operating in these long period sources may advance our understanding of how FRBs are generated.

5 MASER SHOCK MODEL & AFTERGLOW CONSTRAINTS

In the synchrotron maser shock model of FRBs, coherent radiation is generated by the gyration of particles in relativistic, magnetized shocks (Lyubarsky 2014; Beloborodov 2017, 2020; Metzger et al. 2019; Sironi et al. 2021), as has been previously discussed in the literature before the discovery of FRBs (Langdon, Arons & Max 1988; Gallant et al. 1992; Hoshino et al. 1992; Usov & Katz 2000). The model suggests that flaring magnetars can produce relativistic ejecta, which provide conditions conducive to synchrotron maser emission upon interaction with external material in a surrounding nebula (Lyubarsky 2014), or from a previous flare (Metzger et al. 2019). As the shock propagates relativistically, a multiwavelength afterglow is expected (Metzger et al. 2019), with emission peaking at successively lower frequencies due to absorption and a decreasing bulk Lorentz factor Γ .

5.1 Afterglow model

We present relevant dynamical & radiation models, following the afterglow prescription detailed in Metzger et al. (2019) & Margalit et al. (2020b) in this subsection. We extend the model of Metzger et al. (2019) into the non-relativistic expansion phase relevant for late-time radio observations, remaining agnostic about the nature of the circumburst medium. Much of the early afterglow model is similar to models developed for GRB afterglows (e.g. Mészáros, Rees & Wijers 1998; Sari et al. 1998). The late-time dynamics of interest here after a deceleration time t_{dec} rely on the self-similar model of blast-waves described in Blandford & McKee (1976) and non-relativistic Sedov-Taylor model of an expanding blast-wave (Taylor 1950; Sedov 1959). Using these models we can make detailed predictions of the evolution of the shock and multiwavelength afterglow.

5.1.1 Deceleration phase

As shown in Metzger et al. (2019), the shock reaches the deceleration radius r_{dec} in a time t_{dec} , which is equal to the central engine activity time, i.e. $t_{\text{dec}} \approx t_{\text{FRB}} \approx 1$ ms, such that we use these subscripts interchangeably. All parameters denoted with these subscripts refer to the value at the deceleration time. All time-scales of interest for emission below the X-ray band occur at $t \gg t_{\text{dec}}$ therefore we are only interested in solutions to the dynamics and radiation in the deceleration and non-relativistic phases. The bulk Lorentz factor for an adiabatically evolving shock at times $t \gg t_{\text{FRB}}$ is

$$\begin{aligned} \Gamma(r > r_{\text{FRB}}) &= \left(\frac{E_{\text{flare}}(17 - 4k)}{16\pi m_p n_{\text{ext}} r^3 c^2} \right)^{1/2} \\ &= \Gamma_{\text{FRB}} \left(\frac{r}{r_{\text{FRB}}} \right)^{\frac{k-3}{2}}, \end{aligned} \quad (2)$$

where r and n_{ext} are the radius of the shock from the central engine and the external number density of the local environment that the shock ploughs into. n_{ext} depends on the type of environment considered and in general has the form $n_{\text{ext}} \propto r^{-k}$ where $0 \leq k < 3$. Wind-like medium ($k = 2$) and constant medium ($k = 0$) are commonly used, and either may be applicable to the environment surrounding a magnetar (Metzger et al. 2019). Assuming an external density at $n_{\text{ext}}(r_{\text{FRB}})$ at $r = r_{\text{FRB}}$,

$$n_{\text{ext}}(r) = n_{\text{ext}}(r_{\text{FRB}}) \left(\frac{r}{r_{\text{FRB}}} \right)^{-k}. \quad (3)$$

In the observer frame, the distance of the adiabatic shock front from the central engine as a function of time is

$$\begin{aligned} r(t > t_{\text{FRB}}) &= 2c\Gamma^2 t = 2c\Gamma_{\text{FRB}}^2 \left(\frac{r}{r_{\text{FRB}}} \right)^{k-3} \\ &\propto t^{\frac{1}{4-k}} \end{aligned} \quad (4)$$

Therefore, by equations (2) & (4) the shock Lorentz factor varies as a function of observer time as

$$\begin{aligned} \Gamma(r > r_{\text{FRB}}) &= \Gamma_{\text{FRB}} \left(\frac{1}{r_{\text{FRB}}} \right)^{\frac{k-3}{2}} \left((2c\Gamma_{\text{FRB}}^2)^{\frac{1}{4-k}} r_{\text{FRB}}^{\frac{3-k}{4-k}} \right)^{\frac{k-3}{2}} \\ &\propto t^{\frac{k-3}{8-2k}}. \end{aligned} \quad (5)$$

Given this, Metzger et al. (2019) compute the afterglow emission during this deceleration phase. Electrons in the shock are assumed to be in energetic equipartition with the ions such that their mean thermal Lorentz factor is: $\gamma_{\text{therm}} = \frac{m_p}{2m_e} \Gamma$. Particles are assumed not to undergo non-thermal acceleration by the shock (see Section 5.5). The magnetic field is parameterized in terms of the magnetization parameter σ as a fraction of the thermal energy density of the shock such that

$$B(t > t_{\text{FRB}}) = \sqrt{64\pi\sigma\Gamma^2 m_p c^2 n_{\text{ext}}} \propto t^{\frac{-3}{8-2k}}, \quad (6)$$

where we assume throughout $0.1 \leq \sigma \leq 1$, as required for coherent maser emission. As in GRB afterglows, the critical synchrotron frequency ν_{sync} and the cooling frequency ν_{cool} (Sari et al. 1998) in the observer frame is

$$\begin{aligned} \nu_{\text{sync}} &= \frac{qB\gamma_{\text{therm}}^2\Gamma}{2\pi m_e c} \propto t^{\frac{3k-12}{8-2k}} \\ \nu_{\text{cool}} &= \frac{qB}{m_e c} \gamma_{\text{c,cool}}^2 \Gamma \propto t^{\frac{3k-10}{8-2k}}, \end{aligned} \quad (7)$$

where $\gamma_{\text{c,cool}} = \frac{6\pi m_e c}{\sigma_T \Gamma B^2 t}$. Using the parameter values of Margalit et al. (2020b), we find that for the flare associated with FRB 200428 from SGR 1935 + 2154, $\nu_{\text{sync}} < \nu_{\text{cool}}$ after a time of just $t = 0.99 \sigma_1^2$ s. This means that for even the earliest optical data discussed in Section 5.3, the slow cooling regime holds. We note that for larger magnetization values (i.e. $\sigma = 1$) the earliest time predictions within a minute of the burst will be affected.

The spectral luminosity at the critical synchrotron frequency is

$$L_{\nu, \text{sync}} \propto N_{\text{therm}} \Gamma B, \quad (8)$$

where $N_{\text{therm}} \propto n_{\text{ext}} R_{\text{sh}}^3$ is the number of radiating thermal electrons. In the slow-cooling regime (Sari et al. 1998; Margalit et al. 2020b), the spectral luminosity is

$$L_\nu = \begin{cases} L_{\nu, \text{sync}} \left(\frac{\nu}{\nu_{\text{sync}}} \right)^{1/3} & \nu < \nu_{\text{sync}} \\ L_{\nu, \text{sync}} \exp \left(\left(-\frac{\nu}{\nu_{\text{sync}}} \right)^{1/3} \right) & \nu > \nu_{\text{sync}} \end{cases} \quad (9)$$

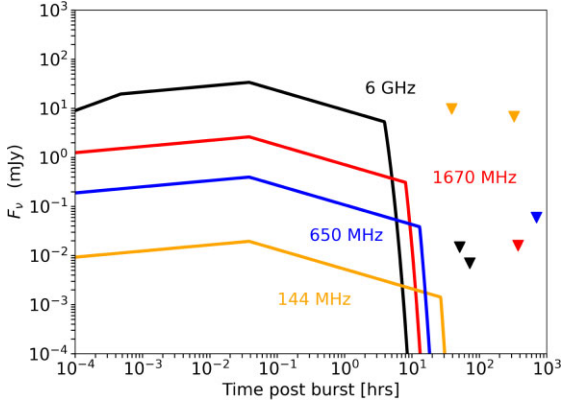


Figure 5. Radio afterglow lightcurves following the FRB from SGR 1935 + 2154, following the prescription of Margalit et al. (2020b) in their fig. 9 as discussed in the text. We show those frequencies with the most stringent and earliest radio observations (Bailes et al. 2021), including the LOFAR limits we present in this work. The lack of non-thermal particles leads to a steep decline in all lightcurves after just hours as the synchrotron frequency drops below the observing frequency.

where the $1/3$ exponent in the second line reflects recent theoretical work by Margalit & Quataert (2021). In practice, for FRB 200428 we compute the predicted afterglow by fitting the initial X-ray data point and using the temporal scaling relations developed for GRB afterglow models described below (Mészáros et al. 1998; Sari et al. 1998; Metzger et al. 2019). To approximate the absorption of lower frequency emission, we follow the method of Margalit et al. (2020b) used in their fig. 9, by ensuring the spectral luminosity is not larger than the expected synchrotron self-absorption value of

$$L_{\nu, \text{SSA}} \approx \frac{8}{3} \pi^2 m_e R_{\text{sh}}^2 v^2 \gamma_{\text{therm}}. \quad (10)$$

5.1.2 Non-relativistic phase

The low luminosity of FRB 200428 means it enters a non-relativistic expansion phase much earlier than a bright FRB (Margalit et al. 2020b). We can model this phase by comparison to the dynamics of supernova remnants, i.e. Sedov–Taylor expansion (Taylor 1950; Sedov 1959). We can find the time at which the shock enters this phase by considering when $\Gamma \approx 1$. Inspection of equations (4) and (5) tell us $\Gamma(t > t_{\text{FRB}}) \propto t^{\frac{k-3}{8-2k}}$ such that the approximate time and radius at which the shock becomes non-relativistic as

$$t_{\text{non-rel}} \approx t_{\text{FRB}} \left(\frac{1}{\Gamma_{\text{FRB}}} \right)^{\frac{8-2k}{k-3}}$$

$$r_{\text{non-rel}} \approx r_{\text{FRB}} \left(\frac{t_{\text{non-rel}}}{t_{\text{FRB}}} \right)^{\frac{1}{4-k}}. \quad (11)$$

For FRB 200428, we find that $t_{\text{non-rel}} \approx 45$ s. We can find the synchrotron frequency and spectral luminosity of the thermal electrons at the time the shock is non-relativistic by equation (7),

$$\nu_{\text{sync, non-rel}} = \nu_{\text{sync, dec}} \left(\frac{t_{\text{non-rel}}}{t_{\text{FRB}}} \right)^{\frac{3k-12}{8-2k}}. \quad (12)$$

In the slow-cooling regime, $L_{\nu, \text{sync}} \approx \text{constant}$ in the deceleration phase (Sari et al. 1998) such that

$$L_{\nu, \text{sync, non-rel}} = L_{\nu, \text{sync, dec}}. \quad (13)$$

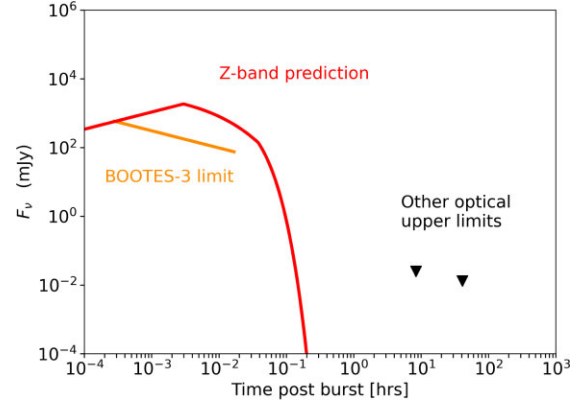


Figure 6. We show in red the z-band lightcurve following the FRB from SGR 1935 + 2154, following the prescription of Margalit et al. (2020b). In orange, the span of the observation taken by BOOTES-3 telescope (after 1 s on source), and in black we include additional optical upper limits by BOOTES-2 and LCOGT presented in Lin et al. (2020a). The peak of the predicted flux is approximately one order of magnitude above the extinction-corrected limit set by the BOOTES-3 observation.

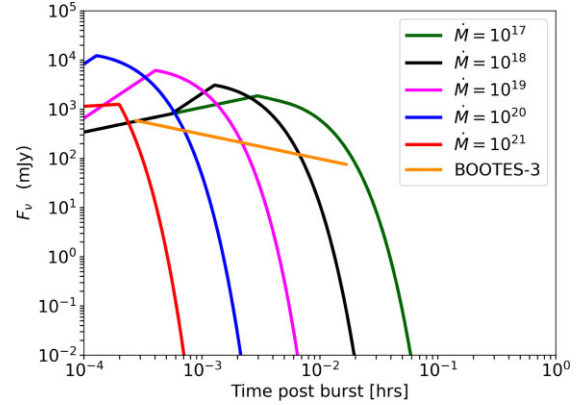


Figure 7. Afterglow z-band lightcurves the wind-like environment case for various values of the mass injection rate \dot{M} . The flux sensitivity of the BOOTES-3 limit is assumed to scale as $T_{\text{obs}}^{-1/2}$ and we show only limits from 1 s after the start of the observation. This appears to rule out the most values of the mass injection rate; however, extrapolation of the observations to the time of the FRB (where they are reported to begin) may completely rule out the maser shock model of FRB 200428.

After defining these quantities at the non-relativistic transition time, we can make predictions for this phase. Using the well-known Sedov–Taylor expansion solution in a constant medium, the shock radius in this regime is

$$r(r > r_{\text{non-rel}}) = r_{\text{non-rel}} \left(\frac{t}{t_{\text{non-rel}}} \right)^{2/5}. \quad (14)$$

In the slow-cooling non relativistic regime, the critical thermal synchrotron frequency and synchrotron luminosity are

$$\nu_{\text{sync}}(t > t_{\text{non-rel}}) = \nu_{\text{sync, non-rel}} \left(\frac{t}{t_{\text{non-rel}}} \right)^{-3}$$

$$L_{\nu, \text{sync}}(t > t_{\text{non-rel}}) = L_{\nu, \text{sync, non-rel}} \left(\frac{t}{t_{\text{non-rel}}} \right)^{3/5}. \quad (15)$$

The above description assumes the thermal particles are relativistic, and therefore is valid while $\gamma_{\text{therm}} \gg 1$. For FRB 200428, in the constant medium case the thermal electrons become non-relativistic

approximately 1 d post-burst due to the low initial Lorentz factor. By this time the synchrotron cut-off frequency has dropped below the observing frequency for radio wavelengths of interest (see Fig. 5). However, extrapolating afterglow lightcurves beyond this time will require additional consideration of the deep-Newtonian regime similarly to GRB afterglows (e.g. Barniol Duran & Giannios 2015). Finally, we note that similar to SNR evolution, the shock is expected to enter the radiative snow-plough phase at a threshold velocity $v_{\text{sh}} \approx \sqrt{\frac{kT}{m_e}} \approx 10^7 \text{ cm s}^{-1}$. For the constant medium case, this occurs for the FRB 200428 at approximately 1000 d post-burst affecting the predicted shock dynamics and lightcurves at very late times, which are not relevant for the observed upper limits.

Using the above, we can make lightcurve predictions through the deceleration and non-relativistic phases for a variety of FRB bursts, as in Margalit, Metzger & Sironi (2020a). Following the prescription of Margalit et al. (2020b), we use the observed properties of FRB 200428 & the coincident X-ray burst to normalize the initial values r_{FRB} , Γ_{FRB} , t_{FRB} given the requirements of the maser and the high-energy burst. Crucially, the peak of the X-ray afterglow at t_{FRB} is normalized to the fluence and peak frequency observed by Insight-HXMT ($\mathcal{F} \approx 7 \times 10^{-7} \text{ erg cm}^{-2}$; $\nu_{\text{peak}} \approx 50 \text{ keV}$; Li et al. 2021). These parameters are outlined for FRB 200428 from SGR 1935 + 2154 in Margalit et al. (2020b) (equations 10–13; used in Figs 5 and 9).

5.2 Radio afterglow

In Fig. 5, we show predicted afterglow lightcurves for four frequencies for which radio limits were obtained within a short time after the initial burst (Bailes et al. 2021), assuming a constant medium ($k = 0$). We find that that model-constraining observations could have been attained with the Karl G. Jansky Very Large Array (VLA; Thompson et al. 1980) & European VLBI Network (EVN)⁸ at 6 and 1.67 GHz respectively if these instruments were on source before the synchrotron frequency of thermal electrons falls below the observing frequency. However, this would necessitate observations beginning just hours after the initial burst, which is incompatible with typical target-of-opportunity delay for large-scale radio facilities. In this case the obtained LOFAR limits presented in Section 3 are not very constraining, primarily attributable to the brightness of the extended emission from SNR G57.2 + 0.8 at 144 MHz.

5.3 Optical afterglow

In Lin et al. (2020a), the authors report on *FAST* observations of SGR 1935 + 2154, as well as a multiwavelength campaign spanning X-ray, optical and radio observations taken after FRB 200428. Of particular note is a minute-long z-band BOOTES-3 (Castro-Tirado et al. 1999) observation that occurred simultaneously with the FRB 200428. The observation began on the 2020 April 28 at 14:34:24.03 (Extended Data Table 1; Lin et al. 2020a), concluding a minute later and setting an upper limit of 17.9 mag. The authors revise this to just 11.7 mag when corrected for dust extinction of 6.2 mag, corresponding to a flux density of approximately 75 mJy. The peak of the X-ray burst coincident with FRB 200428 occurred after the optical observations began (Mereghetti et al. 2020; Ridnaia et al. 2021), meaning the optical limit is constraining for the very early time afterglow of the FRB. Lin et al. (2020a) discuss the limit with

respect to fast optical bursts predicted by Yang, Zhang & Wei (2019) to be produced coincident with FRBs. In Fig. 6, we show z-band lightcurves using the above method (see also Fig. 9 in Margalit et al. 2020b), and note that the BOOTES-3 upper limit significantly constrains the FRB afterglow. We assume that the flux limit scales as $F_{\text{lim}} \propto (T_{\text{obs}})^{-1/2}$ and that limits could be placed 1 s after the FRB and the start of the observations. We note that there are uncertainties with extrapolating the upper limit to the start time of the observation in this manner. Even without such extrapolation, the reported data strongly suggests that the predicted afterglow of the maser shock model in a uniform medium presented in Margalit et al. (2020b) is not compatible with the observed upper limit.

5.3.1 Wind-like case

One way in which the maser shock model could circumvent the optical upper limit is to invoke a more complex, non-uniform environment. In this subsection, we consider the FRB 200428 afterglow shock propagating into a wind-like ($k = 2$) medium at early times. To correct the initial values of Γ_{FRB} & r_{FRB} we refer to the wind-like case in section 2.2.3 of Metzger et al. (2019).

$$\Gamma_{\text{FRB}} = 2.8 E_{\text{flare}}^{1/4} \dot{M}_{19}^{-1/4} \beta_{\text{W}}^{1/2} \delta t^{1/2} \quad (16)$$

$$r_{\text{FRB}} = 1.5 \times 10^9 \text{ cm } E_{\text{flare}}^{1/2} \dot{M}_{19}^{-1/2} \beta_{\text{W}}^{1/2} \delta t^{1/2}, \quad (17)$$

where E_{flare} is the total energy of the flare, \dot{M} is the mass injection rate in units of grams per second, and $\beta_{\text{W}} = v_{\text{w}}/c$ is the time-averaged magnetar wind velocity divided by the speed of light. We note that the Lorentz factor required to explain the low-luminosity X-ray flare is relatively low, challenging the implicit assumption required for maser emission that the flare is initially ultra-relativistic. As a lower limit to the flare energy E_{flare} , we adopt the X-ray luminosity of the coincident flare so that $E_{\text{flare}} = 7 \times 10^{39} \text{ erg}$. We also consider a range of values for the unknown mass injection rate for this source: $\dot{M} = 10^{17-21} \text{ g s}^{-1}$. Values of $\dot{M} = 10^{19-21} \text{ g s}^{-1}$ have been shown to be consistent with the persistent radio nebula and rotation measure of FRBs from FRB 121102 (Margalit & Metzger 2018). However, we also consider values $\dot{M} < 10^{19} \text{ g s}^{-1}$ due to the lack of any persistent radio nebula at the location of SGR 1935 + 2154 in Section 3, and the fact that this source does not appear to be a prolific FRB emitter. In the non-relativistic regime, we use the Sedov–Taylor expansion solution corrected for wind-like media such that $R \propto t^{2/3}$.

In Fig. 7 we show the z-band predictions for the wind-like medium case, considering a range of values for \dot{M} . We also show the flux limits obtained by BOOTES-3, again assuming that the flux limit sensitivity scales as $T_{\text{obs}}^{-1/2}$ from 1 s after the start of the observation. We find that the optical limits are consistent with scenarios in which $\dot{M} > 10^{19} \text{ g s}^{-1}$, but caution that a more detailed analysis of the earliest second of the optical data set may rule out all values of $\dot{M} < 10^{21} \text{ g s}^{-1}$. Finally, we note that $n \propto R^{-k}$ where $k > 2$ might be expected during the early-time deceleration phase at the time of the BOOTES-3 observation. This is because the charge density of the magnetar’s magnetosphere (which has a radial dependence of $k \approx 3$) likely contributes to the local density at $r \approx 10^9 \text{ cm}$. Estimated in terms of the Goldreich-Julian density (Goldreich & Julian 1969), the contribution is

$$n_{\text{GJ}}(r_{\text{FRB}}) \approx \frac{2B_s}{cqP} \left(\frac{R_{\text{NS}}}{R} \right)^{-3} \approx 10^4 \text{ cm}^{-3} B_{s,14} P_0^{-1}. \quad (18)$$

This is comparable to the required value of $n_{\text{ext}} = 4 \times 10^4 \text{ cm}^{-3}$ in the maser shock model of FRB 200428 in equation (12) of Margalit

⁸<https://www.evlbi.org/>

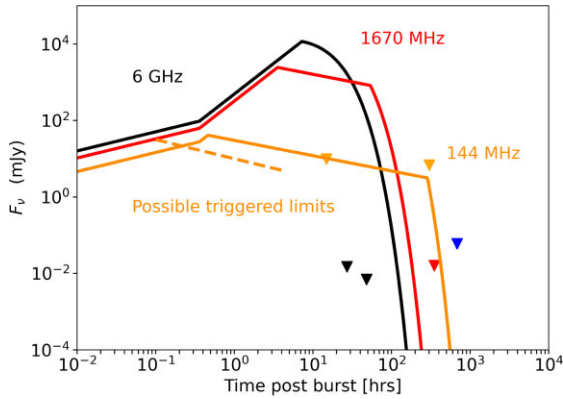


Figure 8. Radio afterglow predictions for a typical luminosity FRB from SGR 1935 + 2154, at frequencies for which early upper limits were obtained for FRB 200428. We include the expected sensitivity from triggered observations with low-frequency radio telescopes that have rapid response capabilities, such as LOFAR or the Murchison Wide Array (MWA; Tingay et al. 2013). We assume a conservative 5 mJy sensitivity for a 4 h integration at 144 MHz, and that sensitivity increases as \sqrt{T} where T is the observation time (dashed orange line). The observations are presumed to begin 6 min after the burst, longer than LOFAR’s current rapid response capability. Inverted triangles correspond to radio upper limits at different frequencies from Bailes et al. (2021) as in Fig 5.

et al. (2020b) therefore considering variable values of k as a function of r may be required for a more accurate description of the afterglow.

5.4 Prospects of radio afterglow detection for the next Galactic FRB

As aforementioned, the lower limit to the radio luminosity of the Galactic FRB 200428 from SGR 1935 + 2154 was approximately 4–5 orders of magnitude less than typical FRB luminosities. In this subsection, we discuss strategies with which to detect an afterglow for a future Galactic FRB.

A bright FRB occurring within the Galaxy could be detected by all-sky radio telescopes such as STARE-2⁹ or AARTFAAC (Prasad et al. 2016), or in the side-lobes of smaller field-of-view radio telescopes. Furthermore, the prompt X-ray component of the afterglow (or coincident magnetar X-ray burst) will be very bright and possibly detected by wide-field gamma-ray instruments such as *Fermi-GBM*, or other X-ray instruments as in the case of the burst from SGR 1935 + 2154.

In Fig. 8, we present 6 GHz, 1.67 GHz, and 144 MHz afterglow lightcurves for a ‘typical’ luminosity ($E_{\text{flare}} = 10^{43}$ erg) FRB from within our Galaxy, along with obtained radio limits for the SGR 1934 + 2154. The afterglow in this figure is calculated using the reference values of r_{FRB} , Γ_{FRB} , t_{FRB} & $v_{\text{sync, dec}}$ for a typical FRB described in equations (30)–(32) and (57) in Metzger et al. (2019). We scale the early X-ray afterglow fluence appropriately such that the $\mathcal{F}_{\nu, \text{sync, max}} \approx E_{\text{flare}} / (4\pi D^2)$. We find that of the radio limits attained after the 2020 burst, only the earliest LOFAR & VLA limits can constrain the model in its current form. We note that for this brighter burst, the condition that $\gamma_{\text{therm}} > 1$ as required by the afterglow model holds until 1000 h post-burst, i.e. the entire time of interest for the radio afterglow.

⁹STARE-2 is now decommissioned, but a successor instrument to the project is planned: Connor et al. (2021)

We also show in Fig. 8 plausible flux limits that could have been attained if radio telescopes capable of rapid automated response such as LOFAR (van Haarlem et al. 2013; Rowlinson et al. 2022) and MWA (Tremblay et al. 2015; Hancock et al. 2019; Anderson et al. 2021) were triggered on the initial burst and were on source within 10 min. We have assumed a 5 mJy flux limit after 4 h observation, where the flux limit scales as $F_{\text{lim}} \propto (T_{\text{obs}})^{-1/2}$ as generally expected, unless the flux threshold is limited by poor u-v coverage for very short integration times. If we are fortunate enough to observe a bonafide FRB from within our Galaxy, rapid optical and radio observations on minute to day time-scales will be crucial to observe the afterglow and verify or falsify the maser shock model of FRBs. Such observations are only possible if the FRB event is reported in a timely manner on networks that distribute astronomical alerts rapidly (such as GCN Barthelmy et al. 1998 or VEvent Williams & Seaman 2006 networks), and radio telescopes have programs in which observations can be interrupted for rapid or automatic repointing.

5.5 Non-thermal radio afterglows

Shock acceleration of particles resulting in non-thermal distributions and radiation occurs almost ubiquitously in high energy astrophysical transients such as supernovae remnants (Yuan et al. 2011; Ackermann et al. 2013), gamma-ray bursts (Waxman 1997), active galactic nuclei (Blandford & Königl 1979), and X-ray binary jets (Markoff, Falcke & Fender 2001). However, particle-in-cell (PIC) simulations of relativistic, magnetized shocks suggest that particle acceleration in shocks does not occur very efficiently (Sironi & Spitkovsky 2009; Sironi, Keshet & Lemoine 2015). As relativistic shocks sweep up organised magnetic field lines, they are compressed into the downstream shocked medium. As compression occurs, the angle between the magnetic field lines and the shock velocity θ increases, such that they are quasi-perpendicular $\theta \approx 90$ deg. PIC simulations show no significant self-generated turbulence or magnetic field from particles in relativistic magnetized shocks (Sironi & Spitkovsky 2011), and as such the particles are forced to slide along the background magnetic field lines. As these field lines run perpendicular to the shock velocity, particles do not escape the shock and therefore do not undergo repeated shock crossings required for Fermi-like shock acceleration. This is quantified in Sironi & Spitkovsky (2009), where the authors define a critical angle θ_{crit} , such that if $\theta > \theta_{\text{crit}}$, particles would have to propagate at greater than the speed of light in order to outrun the shock and undergo multiple shock crossings as required for non-thermal acceleration. For this reason, the FRB afterglow model presented in Metzger et al. (2019), Margalit et al. (2020b) assumes purely thermal radiation, nor are non-thermal particles included in Figs 5, 6, 7, or 8.

As the FRB-initiating shock wave progresses it decelerates to smaller bulk Lorentz factors and propagates into regions of lower magnetic field strength (equations 5 and 6). The typical magnetization and Lorentz factor values assumed in simulations of relativistic, magnetized shocks are $\sigma > 0.1$ and $\Gamma > 10$. For the SGR 1935 + 2154 burst, the afterglow model in Margalit et al. (2020b) suggests the shock enters a non-relativistic Sedov-Taylor phase at $t_{\text{non-rel}} \approx 45$ s via equation (11). This is exceptionally early due to the low-luminosity of this FRB: $t_{\text{non-rel}}$ will be larger by a factor of approximately 100 for a bright FRB, using the shock values from Margalit et al. (2020a), corresponding to a time $t_{\text{non-rel}} \approx 1.5$ h. We suggest that in this non-relativistic phase, particles are able to undergo multiple shock crossing as has been observed in magnetized shocks of supernovae remnants.

To model the non-thermal radiation in the Sedov-Taylor phase, we assume 10 per cent of the total (constant) energy of the shock $E_{\text{sh}} = \frac{4}{3}\pi R_{\text{sh}}^3 n_{\text{ext}} m_p v_{\text{sh}}^2 \propto t^0$ is available for non-thermal particle acceleration, as is canonically expected for cosmic ray acceleration in SNRs (e.g. Strong, Moskalenko & Ptuskin 2007). We further assume an equipartition in energy between hadronic and leptonic acceleration. For this calculation we assume a uniform density medium throughout such that $n_{\text{ext}} = n_{\text{FRB}}$, noting that if the constant density does not extend to such radii, the non-thermal radiation would be lower than predicted here. Non-thermal radiation strongly depends on the magnetic field at the shock. However, in the non-relativistic regime the shock's magnetic field strength is uncertain due to the unknown magnetization of the magnetosphere-ISM transition medium, and possible amplification of the compressed magnetic field due to non-resonant Bell instability (Vink & Laming 2003; Bell 2004). We parametrize the magnetic field in terms of a fraction of the thermal energy density (i.e. equation 6) and the shock-amplified ISM value, such that

$$B = \max\left(\sqrt{64\pi\sigma\beta^2 m_p c^2 n_{\text{ext}}}, \chi_{\text{sh}} B_{\text{ISM}}\right), \quad (19)$$

where $\chi_{\text{sh}} \approx 4$ is the shock compression ratio, the magnetization $\sigma = 0.1-1$ (Metzger et al. 2019), β is the velocity of the blast-wave in units of the speed of light, and we assume $B_{\text{ISM}} \approx 3 \mu\text{G}$. The minimum electron energy can be expressed in terms of the thermal particle energy

$$E_{\text{min}} = \epsilon_e \frac{(p-2)}{2(p-1)} \frac{m_p \beta^2 c^2}{m_e} + m_e c^2, \quad (20)$$

where $\epsilon_e = 0.05$ is the fraction of shock energy (i.e. half of 10 per cent of the total energy due to equipartition) that goes into accelerating electrons and p is the slope of the power law distribution of shock-acceleration electrons which we assume to be $p = 2.2$ (Sironi & Giannios 2013). The spectral luminosity in the slow cooling regime relevant for the radio afterglow is therefore

$$L_\nu \propto \frac{K}{p+1} \nu^{-(p-1)/2} B^{(p+1)/2} \quad (21)$$

where $K = E_{\text{sh}} \epsilon_e (p-2) E_{\text{min}}^{p-2}$ is the normalization factor of the non-thermal electron distribution. We note that exponential suppression of non-thermal particle radiation is assumed for times $t < t_{\text{non-rel}}$, since we do not expect shock acceleration at these times (see above). We treat synchrotron self-absorption using equation (10), replacing γ_{therm} with the Lorentz factor for which the critical synchrotron frequency is the observing frequency ν .

In Fig. 9, we show the predicted thermal (dashed) and non-thermal (solid) afterglow lightcurves for two observing frequencies. For FRB 200428, the weak burst means the afterglow is relatively dim, and does not challenge the radio upper limits previously discussed. However, we note that particularly at low frequencies the non-thermal radiation could contribute to the overall flux shortly before the thermal synchrotron cut-off. In Fig. 10, we show a similar plot for a bright FRB for two distances, corresponding to Galactic or nearby extragalactic FRBs. For a bright Galactic FRB, non-thermal radio emission could be detectable up to years after the initial burst. Long-term monitoring of such a source would provide opportunities for detailed modelling of the afterglow, contributing to our understanding of shocks more generally. However, non-thermal radio emission from even a nearby extragalactic FRB is not likely to be bright enough to probe.

The closest repeating FRB resides at a distance of $D \approx 3.6 \text{ Mpc}$ in M81 (Bhardwaj et al. 2021; Kirsten et al. 2022). Kirsten et al. (2022)

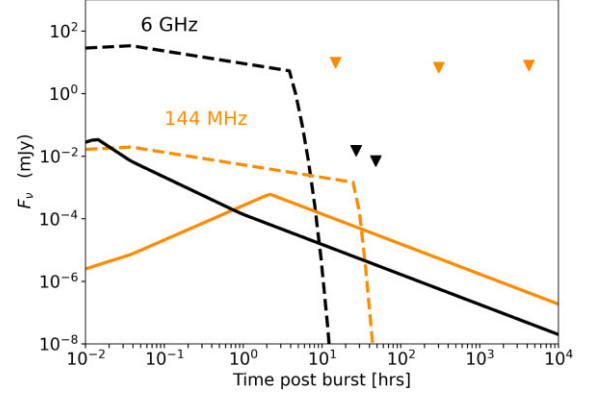


Figure 9. Dotted and solid lines denote the thermal and non-thermal components respectively to the afterglow associated with the FRB from SGR 1935 + 2154, as described in the text. The major break denotes the time at which the emission becomes optically thin. The non-thermal afterglow contributes most significantly for lower frequencies and at later times, near the thermal synchrotron cut-off. For all times of interest, the shock magnetic field strength is always greater than the amplified ISM value. Black and orange inverted triangles correspond radio upper limits taken after FRB 200428 by VLA and LOFAR at 6 GHz and 144 MHz, respectively.

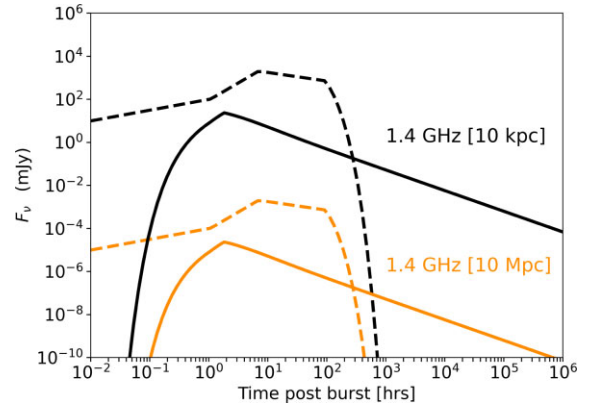


Figure 10. Dotted and solid lines denote the thermal and non-thermal components respectively to the afterglow after a typical luminosity FRB, where $\nu_{\text{obs}} = 1.4 \text{ GHz}$. We show two distances representative of a bright Galactic burst (black) and a nearby extragalactic FRB. The non-thermal emission is relatively weak, but allows for late-times monitoring of the lightcurve in the Galactic FRB case. The thermal radio afterglow is tentatively observable ($> \mu\text{Jy}$) in the hours following a close extragalactic FRB. We assume exponential suppression of non-thermal acceleration before the shock becomes non-relativistic.

obtained deep persistent flux constraints ($6.5 \mu\text{Jy}$ at 1.5 GHz) of the source before their reported FRBs, but 1–6 months after previously reported bursts by Bhardwaj et al. (2021). Unfortunately, the low luminosity of the observed FRBs mean that predicted afterglow emission is not observable.

6 CONCLUSIONS

In this work we have presented results from LOFAR imaging observations of SGR 1935 + 2154, following the Galactic magnetar burst FRB 200428. We discuss interpretations of the LOFAR results, stressing the importance of simultaneous X-ray/radio observations of active magnetars to constrain both magnetospheric and flare/shock models of FRBs. We also make recommendations of rapid obser-

vations on minute-day time-scales following Galactic or nearby extragalactic FRBs. We have analysed optical and radio early time upper limits in the context of models of FRB 200428 that predict a multiwavelength afterglow, namely the synchrotron maser shock model. We have found that early BOOTES-3 optical observations appear to rule out simple versions of this model, but invoking a wind-like environment close to the FRB emission zone may mitigate the constraints. We also suggest that shock accelerated particle populations should be considered at late times when the shock is non-relativistic; however, we find that such non-thermal emission is too faint to be observable, except for future Galactic FRBs.

ACKNOWLEDGEMENTS

AC would like to thank Ben Margalit for a thorough explanation of the shock-maser afterglow model for the specific case of SGR 1935 + 2154, and for describing the calculation underlying the afterglow light curve predictions in the non-relativistic regime in their work. We would also like to thank Jeremy Harwood, Oliver Boersma, and Rhaana Starling for useful discussions and suggestions, and the anonymous referee whose comments improved this work.

Sections of this paper are based on data obtained with the *International LOFAR Telescope* (ILT) under project codes DDT13_004, DDT13_006, DDT14_007. LOFAR is the Low Frequency Array designed and constructed by ASTRON. It has observing, data processing, and data storage facilities in several countries, that are owned by various parties (each with their own funding sources), and that are collectively operated by the ILT foundation under a joint scientific policy. The ILT resources have benefited from the following recent major funding sources: CNRS-INSU, Observatoire de Paris and Université d'Orléans, France; BMBF, MIWF-NRW, MPG, Germany; Science Foundation Ireland (SFI), Department of Business, Enterprise and Innovation (DBEI), Ireland; NWO, The Netherlands; The Science and Technology Facilities Council, UK; Ministry of Science and Higher Education, Poland.

DATA AVAILABILITY

The LOFAR visibility data are publicly available at the LOFAR Long Term Archive (<https://lta.lofar.eu/>) under the relevant project codes. A basic reproduction package providing the scripts and data required to reproduce the figures and tables of this paper will be available upon publication.

REFERENCES

- Ackermann M. et al., 2013, *Science*, 339, 807
 Ambrosi E., Barthelmy S. D., D’Ai A., D’Elia V., Lien A. Y., Page K. L., Neil Gehrels Swift Observatory Team, 2019, GRB Coordinates Netw., 26153, 1
 Anderson G. E. et al., 2012, *ApJ*, 751, 53
 Anderson G. E. et al., 2021, *PASA*, 38, e026
 Archibald R. F. et al., 2017, *ApJ*, 849, L20
 Bailes M. et al., 2021, *MNRAS*, 503, 5367
 Bannister K. W. et al., 2019, *Science*, 365, 565
 Barniol Duran R., Giannios D., 2015, *MNRAS*, 454, 1711
 Barthelmy S. D. et al., 1998, in Meegan C. A., Preece R. D., Koshut T. M., eds, AIP Conf. Ser. Vol. 428, Gamma-Ray Bursts, 4th Hunstville Symposium. Am. Inst. Phys., New York, p. 99
 Bell A. R., 2004, *MNRAS*, 353, 550
 Beloborodov A. M., 2017, *ApJ*, 843, L26
 Beloborodov A. M., 2020, *ApJ*, 896, 142
 Beloborodov A. M., 2021, *ApJ*, 922, L7
 Bhandari S. et al., 2018, *MNRAS*, 475, 1427
 Bhandari S. et al., 2020, *ApJ*, 895, L37
 Bhardwaj M. et al., 2021, *ApJ*, 910, L18
 Blandford R. D., Königl A., 1979, *ApJ*, 232, 34
 Blandford R. D., McKee C. F., 1976, *Phys. Fluids*, 19, 1130
 Bochenek C. D., McKenna D. L., Belov K. V., Kocz J., Kulkarni S. R., Lamb J., Ravi V., Woody D., 2020a, *PASP*, 132, 034202
 Bochenek C. D., Ravi V., Belov K. V., Hallinan G., Kocz J., Kulkarni S. R., McKenna D. L., 2020b, *Nature*, 587, 59
 Cai C. et al., 2022a, *ApJS*, 260, 24
 Cai C. et al., 2022b, *ApJS*, 260, 25
 Caleb M. et al., 2022, *Nature Astron.*, 6, 828
 Camilo F., Ransom S. M., Halpern J. P., Reynolds J., Helfand D. J., Zimmerman N., Sarkissian J., 2006, *Nature*, 442, 892
 Camilo F., Reynolds J., Johnston S., Halpern J. P., Ransom S. M., van Straten W., 2007, *ApJ*, 659, L37
 Campana S., de Ugarte Postigo A., Thoene C. C., Gorosabel J., Rea N., Coti Zelati F., 2014, GRB Coordinates Netw., 16535, 1
 Carbone D. et al., 2018, *Astron. Comput.*, 23, 92
 Castro-Tirado A. J. et al., 1999, *A&AS*, 138, 583
 Chatterjee S. et al., 2017, *Nature*, 541, 58
 Chen K., Ruderman M., 1993, *ApJ*, 402, 264
 Cheng A. F., Ruderman M. A., 1977, *ApJ*, 212, 800
 CHIME/FRB Collaboration, 2020, *Nature*, 587, 54
 Connor L. et al., 2021, *PASP*, 133, 075001
 Cooper A. J., Wijers R. A. M. J., 2021, *MNRAS*, 508, L32
 Cummings J. R., Barthelmy S. D., Chester M. M., Page K. L., 2014, *Astron. Telegram*, 6294, 1
 Esposito P. et al., 2020, *ApJ*, 896, L30
 Gaensler B. M., 2014, GRB Coordinates Netw., 16533, 1
 Gallant Y. A., Hoshino M., Langdon A. B., Arons J., Max C. E., 1992, *ApJ*, 391, 73
 Göğüş E., Woods P. M., Kouveliotou C., van Paradijs J., Briggs M. S., Duncan R. C., Thompson C., 1999, *ApJ*, 526, L93
 Göğüş E., Woods P. M., Kouveliotou C., van Paradijs J., Briggs M. S., Duncan R. C., Thompson C., 2000, *ApJ*, 532, L121
 Goldreich P., Julian W. H., 1969, *ApJ*, 157, 869
 Gourdji K., Michilli D., Spitler L. G., Hessels J. W. T., Seymour A., Cordes J. M., Chatterjee S., 2019, *ApJ*, 877, L19
 Hancock P. J. et al., 2019, *PASA*, 36, e046
 Hankins T. H., Eilek J. A., 2007, *ApJ*, 670, 693
 Heintz K. E. et al., 2020, *ApJ*, 903, 152
 Hermsen W. et al., 2013, *Science*, 339, 436
 Hermsen W. et al., 2018, *MNRAS*, 480, 3655
 Hessels J. W. T. et al., 2019, *ApJ*, 876, L23
 Hoshino M., Arons J., Gallant Y. A., Langdon A. B., 1992, *ApJ*, 390, 454
 Hurley-Walker N. et al., 2022, *Nature*, 601, 526
 Israel G. L. et al., 2016, *MNRAS*, 457, 3448
 Katz J. I., 2016, *ApJ*, 826, 226
 Kirsten F., Snelders M. P., Jenkins M., Nimmo K., van den Eijnden J., Hessels J. W. T., Gawroński M. P., Yang J., 2021, *Nature Astron.*, 5, 414
 Kirsten F. et al., 2022, *Nature*, 602, 585
 Kothes R., Sun X., Gaensler B., Reich W., 2018, *ApJ*, 852, 54
 Kozlova A. V. et al., 2016, *MNRAS*, 460, 2008
 Kramer M., Stappers B. W., Jessner A., Lyne A. G., Jordan C. A., 2007, *MNRAS*, 377, 107
 Kuiack M., Wijers R. A. M. J., Rowlinson A., Shulevski A., Huizinga F., Molenaar G., Prasad P., 2020, *MNRAS*, 497, 846
 Kumar P., Lu W., Bhattacharya M., 2017, *MNRAS*, 468, 2726
 Kumar P., Bošnjak Ž., 2020, *MNRAS*, 494, 2385
 Kumar P., Gill R., Lu W., 2022, *MNRAS*, 516, 2697
 Langdon A. B., Arons J., Max C. E., 1988, *Phys. Rev. Lett.*, 61, 779
 Levan A., Kouveliotou C., Fruchter A., 2018, *ApJ*, 854, 161
 Levin L. et al., 2010, *ApJ*, 721, L33
 Li C. K. et al., 2021, *Nature Astron.*, 5, 378
 Li X. et al., 2022, *ApJ*, 931, 56
 Lien A. Y. et al., 2014, GRB Coordinates Netw., 16522, 1

- Lin L. et al., 2020a, *Nature*, 587, 63
- Lin L., Göğüş E., Roberts O. J., Baring M. G., Kouveliotou C., Kaneko Y., van der Horst A. J., Younes G., 2020b, *ApJ*, 902, L43
- Lorimer D. R., Bailes M., McLaughlin M. A., Narkevic D. J., Crawford F., 2007, *Science*, 318, 777
- Lower M. E., Shannon R. M., Johnston S., Bailes M., 2020, *ApJ*, 896, L37
- Lu W., Kumar P., 2018, *MNRAS*, 477, 2470
- Lu W., Kumar P., 2019, *MNRAS*, 483, L93
- Lu W., Kumar P., Zhang B., 2020, *MNRAS*, 498, 1397
- Luo R. et al., 2020, *Nature*, 586, 693
- Lyubarsky Y., 2014, *MNRAS*, 442, L9
- Lyutikov M., Popov S., 2020, preprint ([arXiv:2005.05093](https://arxiv.org/abs/2005.05093))
- Marcote B. et al., 2020, *Nature*, 577, 190
- Margalit B., Metzger B. D., 2018, *ApJ*, 868, L4
- Margalit B., Quataert E., 2021, *ApJ*, 923, L14
- Margalit B., Metzger B. D., Sironi L., 2020a, *MNRAS*, 494, 4627
- Margalit B., Beniamini P., Sridhar N., Metzger B. D., 2020b, *ApJ*, 899, L27
- Markoff S., Falcke H., Fender R., 2001, *A&A*, 372, L25
- Melrose D. B., Rafat M. Z., Mastrano A., 2021, *MNRAS*, 500, 4530
- Mereghetti S. et al., 2020, *ApJ*, 898, L29
- Mészáros P., Rees M. J., Wijers R. A. M. J., 1998, *ApJ*, 499, 301
- Metzger B. D., Margalit B., Sironi L., 2019, *MNRAS*, 485, 4091
- Michilli D. et al., 2018, *Nature*, 553, 182
- Nimmo K. et al., 2022, *Nature Astron.*, 6, 393
- Niu C. H. et al., 2022, *Nature*, 606, 873
- Offringa A. R. et al., 2014, *MNRAS*, 444, 606
- Palmer D. M., 2020, *Astron. Telegram*, 13675, 1
- Petroff E. et al., 2016, *PASA*, 33, e045
- Platts E., Weltman A., Walters A., Tendulkar S. P., Gordin J. E. B., Kandhai S., 2019, *Phys. Rep.*, 821, 1
- Pleunis Z. et al., 2021, *ApJ*, 923, 1
- Plotnikov I., Sironi L., 2019, *MNRAS*, 485, 3816
- Prasad P. et al., 2016, *J. Astron. Instrum.*, 5, 1641008
- Ravi V. et al., 2019, *Nature*, 572, 352
- Rea N., Pons J. A., Torres D. F., Turolla R., 2012, *ApJ*, 748, L12
- Ridnaia A. et al., 2021, *Nature Astron.*, 5, 372
- Rowlinson A. et al., 2022, preprint ([arXiv:2203.16918](https://arxiv.org/abs/2203.16918))
- Ruderman M. A., Sutherland P. G., 1975, *ApJ*, 196, 51
- Sari R., Piran T., Narayan R., 1998, *ApJ*, 497, L17
- Sedov L. I., 1959, *Similarity and Dimensional Methods in Mechanics*. Academic Press. Cambridge, Massachusetts
- Sironi L., Giannios D., 2013, *ApJ*, 778, 107
- Sironi L., Spitkovsky A., 2009, *ApJ*, 698, 1523
- Sironi L., Spitkovsky A., 2011, *ApJ*, 726, 75
- Sironi L., Keshet U., Lemoine M., 2015, *Space Sci. Rev.*, 191, 519
- Sironi L., Plotnikov I., Näätäjä J., Beloborodov A. M., 2021, *Phys. Rev. Lett.*, 127, 035101
- Sridhar N., Metzger B. D., Beniamini P., Margalit B., Renzo M., Sironi L., Kowlakas K., 2021, *ApJ*, 917, 13
- Strong A. W., Moskalenko I. V., Ptuskin V. S., 2007, *Ann. Rev. Nucl. Part. Sci.*, 57, 285
- Swinbank J. D. et al., 2015, *Astron. Comput.*, 11, 25
- Tavani M. et al., 2021, *Nature Astron.*, 5, 401
- Taylor G., 1950, *Proc. R. Soc.*, 201, 159
- Taylor J. H., Cordes J. M., 1993, *ApJ*, 411, 674
- The CHIME/FRB Collaboration, 2021, *ApJS*, 257, 59
- Thompson A. R., Clark B. G., Wade C. M., Napier P. J., 1980, *ApJS*, 44, 151
- Thornton D. et al., 2013, *Science*, 341, 53
- Tian J. et al., 2022, *PASA*, 39, e003
- Tingay S. J. et al., 2013, *PASA*, 30, e007
- Tremblay S. E. et al., 2015, *PASA*, 32, e005
- Usov V. V., 1987, *ApJ*, 320, 333
- Usov V. V., Katz J. I., 2000, *A&A*, 364, 655
- van Haarlem M. P. et al., 2013, *A&A*, 556, A2
- Vink J., Laming J. M., 2003, *ApJ*, 584, 758
- Wadiasingh Z., Timokhin A., 2019, *ApJ*, 879, 4
- Wang W., Zhang B., Chen X., Xu R., 2019, *ApJ*, 876, L15
- Wang W.-Y., Yang Y.-P., Niu C.-H., Xu R., Zhang B., 2022, *ApJ*, 927, 105
- Waxman E., 1997, *ApJ*, 485, L5
- Williams R. D., Seaman R., 2006, in Gabriel C., Arviset C., Ponz D., Enrique S., eds, *ASP Conf. Ser. Vol. 351, Astronomical Data Analysis Software and Systems XV*. Astron. Soc. Pac., San Francisco. p. 637
- Wood J., Bissaldi E., Fermi GBM Team, 2019, *GRB Coordinates Netw.*, 25975, 1
- Yang Y.-P., Zhang B., 2021, *ApJ*, 919, 89
- Yang Y.-P., Zhang B., Wei J.-Y., 2019, *ApJ*, 878, 89
- Yang Y.-P., Zhu J.-P., Zhang B., Wu X.-F., 2020, *ApJ*, 901, L13
- Younes G. et al., 2020, *ApJ*, 904, L21
- Younes G. et al., 2021, *Nature Astron.*, 5, 408
- Yuan Q., Liu S., Fan Z., Bi X., Fryer C. L., 2011, *ApJ*, 735, 120
- Zhang C. F. et al., 2020, *Astron. Telegram*, 13699, 1
- Zhou P., Zhou X., Chen Y., Wang J.-S., Vink J., Wang Y., 2020, *ApJ*, 905, 99

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.