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The MUSE eXtremely deep field: first panoramic view of an Mg II emitting intragroup medium[★]

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ABSTRACT

Using the exquisite data from the MUSE eXtremely Deep Field (MXDF), we report the discovery of an Mg II emission nebula with an area above a 2σ significance level of 1000 proper kpc². This provides the first panoramic view of the spatial distribution of magnesium in the intragroup medium of a low-mass group of five star-forming galaxies at $z = 1.31$. The galaxy group members are separated by less than 50 physical kpc in projection and ≈ 120 km s⁻¹ in velocity space. The most massive galaxy has a stellar mass of $10^{9.35} M_{\odot}$ and shows an Mg II P-Cygni line profile, indicating the presence of an outflow, which is consistent with the spatially resolved spectral analysis showing $\approx +120$ km s⁻¹ shift of the Mg II emission lines with respect to the systemic redshift. The other galaxies are less massive and only show Mg II in emission. The detected Mg II nebula has a maximal projected extent of ≈ 70 kpc, including a low-surface-brightness ($\approx 2 \times 10^{-19}$ erg s⁻¹ cm⁻² arcsec⁻²) gaseous bridge between two subgroups of galaxies. The presence of absorption features in the spectrum of a background galaxy located at an impact parameter of 19 kpc from the closest galaxy of the group indicates the presence of gas enriched in magnesium even beyond the detected nebula seen in emission, which suggests that we are observing the tip of a larger intragroup medium. The observed Mg II velocity gradient suggests an overall rotation of the structure along the major axis of the most massive galaxy. Our MUSE data also reveal extended Fe II* emission in the vicinity of the most massive galaxy, aligned with its minor axis and pointing towards a neighboring galaxy. Extended [O II] emission is found around the galaxy group members and at the location of the Mg II bridge. Our results suggest that both tidal stripping effects from galaxy interactions and outflows are enriching the intragroup medium of this system.

Key words. galaxies: groups: general – Galaxy: formation – galaxies: evolution – galaxies: interactions – intergalactic medium

1. Introduction

In the Λ CDM framework, theories and simulations predict that galaxies grow in mass and size by accreting cold gas and by merging with other galaxies. The environment is known to play a crucial role in shaping galaxy growth. Indeed, theoretical work indicates that tidal interactions and merging events can dramatically impact the properties of the galaxies and their circumgalac-

tic medium (CGM, e.g., [Hani et al. 2018](#)). Studying galaxy systems is therefore crucial for understanding the effect of environment on galaxy formation and evolution.

A large number of galaxy groups found in spectroscopic redshift surveys have been reported in the literature (e.g., [Cucciati et al. 2010](#); [Iovino et al. 2016](#); [Abril-Melgarejo et al. 2021](#)). Galaxy groups are generally referred to as the lower mass end (dark matter halo mass $\leq 10^{13} M_{\odot}$) of galaxy clusters, although the definition of groups in terms of halo mass is vague. A commonly adopted definition of a galaxy group is an association of three or more galaxies with a physical distance of

[★] Based on observations made with ESO telescope at the La Silla Paranal Observatory under the large program 1101.A-0127.

$\Delta r \lesssim 500$ kpc and a velocity separation of $\Delta v \lesssim 500$ km s⁻¹ (e.g., [Knobel et al. 2009](#); [Diener et al. 2013](#)).

Gas exchanges between the group members and their CGM as well as galaxy interactions shape the intragroup medium (IGrM). It is, thus, a great laboratory for studying the effects of the local environment and gas flows on galaxy formation and evolution. The IGrM is a multi-phase medium. At very low redshifts, it is mainly detected and studied through the X-ray emission from its hot phase (e.g., [Lovisari et al. 2015](#); [Eckert et al. 2017](#); [Oppenheimer et al. 2021](#), and associated papers for a review). Those studies are however limited to massive galaxy groups (dark matter, hereafter DM, halo mass $>10^{13} M_{\odot}$) due to the low X-ray flux at lower masses ([Lovisari et al. 2021](#)). Observations of neutral hydrogen in nearby galaxy groups reveal that the cold phase of the IGrM also shows extended structures ([Michel-Dansac et al. 2010](#); [Borthakur et al. 2010](#); [Serra et al. 2013](#)). At higher redshifts, most constraints on the associated IGrM come from absorption line studies in lines of sight probed by bright background objects. Several possibilities are proposed and discussed in the literature to explain the origin of the absorbing gas in systems with three or more galaxies: the gas detected in absorption could be associated with the IGrM of the system ([Nielsen et al. 2018](#)), the CGM of the group members ([Bordoloi et al. 2011](#); [Fossati et al. 2019](#)), tidal tail remnants ([Kacprzak et al. 2010](#); [Dutta et al. 2020](#)), or orbiting gas clouds ([Bielby et al. 2017](#)) after galaxy interactions.

The advent of integral field unit (IFU) instruments is revolutionizing the study of the circum- and intragroup media by allowing their direct mapping in emission. The unprecedented sensitivity of VLT/MUSE ([Bacon et al. 2010](#)) recently allowed for the detection of ionized and enriched intragroup nebula at $z \lesssim 0.7$. [Epinat et al. \(2018\)](#) reported the detection of a 150 kpc large [O II] nebula at $z \approx 0.7$ surrounding a dozen galaxies with stellar masses around $10^{10} M_{\odot}$. [Chen et al. \(2019\)](#) reported a 100 kpc H α blob around a galaxy group with a dynamical mass of $\approx 3 \times 10^{12} M_{\odot}$ at $z \approx 0.3$. While [Epinat et al. \(2018\)](#) do not exclude a possible contribution from active galactic nucleus (AGN) outflows, both analyses favor a scenario where the observed metal enriched gas has been extracted from galaxies by tidal stripping forces. Moreover, [Johnson et al. \(2018\)](#) discovered six and [Helton et al. \(2021\)](#) three spatially extended nebulae emitting in [O III], [O II] and H β in $z \approx 0.5$ galaxy groups hosting a quasar. Both studies suggest that nebulae are a signature of galaxy interactions in quasar host groups.

At higher redshift, the IGrM of star-forming galaxies is mainly detected in H I Lyman α $\lambda 1215.67$ Å (Ly α) and generally referred to as Ly α blobs (LABs, e.g., [Steidel et al. 2010](#); [Caminha et al. 2016](#); [Vanzella et al. 2017](#); [Herenz et al. 2020](#)). The Ly α line traces the neutral phase of the IGrM at $z > 0.3$; at $z = 0$, Ly α observations are extremely limited. In order to connect the ionized intragroup media observed at lower redshifts to those of high redshift galaxy systems, a tracer of the neutral gas phase observable at low- z is needed. The Mg II $\lambda\lambda 2796, 2803$ (hereafter Mg II) doublet is a very good candidate. Indeed, because of the lower ionization potential of Mg I (7.6 eV) compared to H I (13.6 eV), hydrogen gas is neutral when Mg II photons are emitted, namely, Mg II traces the H I gas. Because Mg II is a resonant line, we expect to observe extended Mg II emission, similarly to Ly α (e.g., [Wisotzki et al. 2016](#); [Leclercq et al. 2017](#)).

The detection of extended Mg II emission was reported for the first time ten years ago by [Rubin et al. \(2011\)](#) and later by [Martin et al. \(2013\)](#), both using slit spectroscopy of galaxies at $z \approx 0.7$ and 0.9, respectively. Statistical evidence for such Mg II halos at slightly higher redshift ($z \approx 1.5$) was demonstrated by

[Erb et al. \(2012\)](#) via the stacking of long-slit spectra. Their analyses were, however, limited by the slit aperture, preventing us from getting a complete view of the galaxy surroundings. By reporting non-detections in narrow band (NB) images of five star-forming galaxies at $z \approx 0.7$, [Rickards Vaught et al. \(2019\)](#) reinforced the assumption that the detection of the extended Mg II gas is challenging. Recently, [Burchett et al. \(2021\)](#) re-observed the [Rubin et al. \(2011\) \$z = 0.7\$ galaxy using the Keck Cosmic Webb Imager \(Keck/KCWI, \[Martin et al. 2010\]\(#\)\) and confirmed the detection of significant extended Mg II emission spanning over \$\approx 40\$ kpc. They were able to perform the first detailed spatially resolved analysis of an Mg II halo and found that isotropic outflow models best fit the data. This detection has been followed up by two similar discoveries \(\[Zabl et al. 2021\]\(#\); \[Wisotzki et al., in prep.\]\(#\)\) in the MUSE GAs FLOW and Wind \(MEGAFLOW\) survey and MUSE-Deep fields \(\[Bacon et al. 2017\]\(#\)\), respectively. \[Zabl et al. \\(2021\\)\]\(#\) reported the first discovery of an Mg II emission halo probed by a quasar sightline around a \$z = 0.702\$ galaxy. Thanks to the 3D MUSE view on the CGM of this galaxy, the authors were able to compare the observations with toy models and found good consistency with a bi-conical outflow model. These very recent discoveries highlight the fact that IFU instruments truly are game changers for the study of the circum- and intragroup medium of galaxies.](#)

In this paper, we report the discovery of the first intragroup medium mapped in Mg II, detected in the MUSE eXtremely Deep Field (MXDF, [Bacon et al., in prep.](#) and see also [Bacon et al. 2021](#)). The emitting nebula encompasses a group of five low-mass galaxies ($M_* < 10^{9.4} M_{\odot}$) at $z \approx 1.31$ with a halo mass of $\approx 10^{11.7} M_{\odot}$. It extends over ≈ 70 kpc (physical) and unveils gaseous connections between galaxies. By taking advantage of the three-dimensional (3D) information provided by the MUSE data cubes, we study the kinematics of the nebula. Diffuse [O II] $\lambda\lambda 3726, 3729$ (hereafter [O II]) and Fe II* ($\lambda 2365, \lambda 2396, \lambda 2612, \lambda 2626$, hereafter Fe II*) emission are also detected and allow a comparison between the different gas phases of the IGrM. This first detection of an Mg II emitting intragroup nebula – until now, only locally probed along absorption lines of sight – allows us to shed light on the highly debated existence of an IGrM in low-mass galaxy groups and on its origin.

The paper is organized as follows. We describe the MXDF observations, data reduction, and catalog classification in Sect. 2. The group and the properties of the galaxy members studied in this paper are presented in Sect. 3. Section 4 describes the detection and analysis of the Mg II intragroup nebula. In Sect. 5, we investigate the [O II] and Fe II* properties of the group and compare them with Mg II. Then, in Sect. 6, we discuss the existence and origin of the intragroup medium and the mechanisms that make it shine. Section 6 also provides a comparison with the literature. Finally, we present our summary and conclusions in Sect. 7.

Throughout the paper, all magnitudes are expressed in the AB system and distances are in physical units that are not comoving. We assume a flat Λ CDM cosmology with $\Omega_m = 0.315$ and $H_0 = 67.4$ km s⁻¹ Mpc⁻¹ ([Planck Collaboration VI 2020](#)); in this framework, a 1'' angular separation corresponds to 8.6 kpc proper at the redshift of the group ($z \approx 1.31$).

2. MXDF observations

Here, we provide a summary of the data acquisition, data reduction, and catalog building processes. More details can be found in [Bacon et al. \(2021\)](#) and in the upcoming survey paper [Bacon et al. \(in prep.\)](#).

The MXDF data were taken as part of the MUSE Guaranteed Time Observations program between August 2018 and January 2019 under photometric conditions. The observations were performed with the VLT GALACSI/AOF ground-layer adaptive optics system (Kolb et al. 2016; Madec et al. 2018). A single 140-h pointing was completed after rejection of bad quality exposures. The MUSE field of view was rotated between each observation in order to reduce the systematics. The resulting quasi-circular field is more than 100 h deep in the inner 31'' radius and the exposure time decreases to ≈ 10 h at a radius of 41''. The MXDF data partially overlaps with the MUSE 10-h and 30-h fields from Bacon et al. (2017), denoted mosaic and udf-10, respectively, as well as with the 1-h MUSE-Wide fields (Urrutia et al. 2019, and in prep.).

The data reduction is similar to the one described in Bacon et al. (2017); it is based on the MUSE public pipeline (Weilbacher et al. 2020) and includes some improvements resulting in reduced systematics and a better sky subtraction. We estimated the point spread function (PSF) using the *muse-psfrec* software (Fusco et al. 2020). The final PSF is modeled using a Moffat function (Moffat 1969) with parameters of 0''.6 (0''.4) full width at half maximum (FWHM) and $\beta = 2.1$ (1.8) at the blue (red) end of the datacube. The datacube contains 157034 spatial pixels (spaxel) of 0''.2 \times 0''.2. The number of spectra matches the number of spaxels with a wavelength range of 4700 Å to 9350 Å divided into 3721 pixels of 1.25 Å. The datacube also contains the estimated variance for each pixel.

The MXDF source catalog (Bacon et al., in prep.) consists of (i) sources blindly detected and extracted using the ORIGIN software (Mary et al. 2020) which is designed to detect emission lines in 3D data sets and (ii) sources extracted using the ODHIN software (Bacher 2017) based on an HST catalog (Rafelski et al. 2015), which performs source deblending in MUSE using high resolution HST images. The ODHIN approach is similar to TDOSE (Schmidt et al. 2019) developed for the MUSE wide survey (Urrutia et al. 2019; Schmidt et al. 2021) with the differences that it is non-parametric, it uses multiple broadband HST images, and it implements a regularization process to avoid noise amplification for very close sources. The sources are then inspected by three experts and a final catalog of 733 sources with redshift and associated confidence is created. It includes 406 new spectroscopic redshift measurements with respect to the previous MUSE deep field catalog (Inami et al. 2017).

3. A group of galaxies at $z \approx 1.3$

The Mg II nebula was discovered during the search for extended Mg II emission around galaxies in the MUSE *Hubble* Ultra Deep Field (UDF, Bacon et al. 2017). A visual inspection, first in the udf-10 and then in the newly observed MXDF data (see Appendix A), revealed, at first glance, Mg II emission offset from a galaxy and extended in one direction. By increasing the size of the search area and looking for neighbors, we found this object to be surrounded by four close galaxies. In this section, we focus on the environment (Sect. 3.1), global properties and integrated spectral features of the group members (Sects. 3.2 and 3.3, respectively). The Mg II intragroup nebula detected in the five-member galaxy group is presented and analyzed in the next section (Sect. 4).

3.1. The galaxy group and its environment

The group we are studying here consists of five galaxies at redshift $z \approx 1.31$. Using a friends-of-friends algorithm

(Huchra & Geller 1982) connecting galaxies separated by less than 450 kpc and 500 km s^{-1} in the MXDF, UDF, and MUSE-Wide catalogs (Bacon et al., in prep.; Urrutia et al. 2019, and in prep.), we found that this group is at the center of a larger structure of 14 galaxies with spectroscopically confirmed redshifts (with high confidence, $Z_{\text{CONF}} > 1$ in the MUSE catalogs, except for one object, see below) and separated by less than 460 km s^{-1} in redshift space. This structure spans the MUSE UDF mosaic field over 1.9 Mpc in projection (see Fig. 1, left panel). Three galaxies aligned in projection with the five-galaxy group form a 0.9 Mpc long substructure extending in the MUSE-Wide fields, namely, beyond the MUSE deep fields.

For the rest of the paper, we only focus on the subgroup of five galaxies located at the center of the larger structure (area within the red square in the first panel of Fig. 1 and shown in the two other panels) because they are the galaxies embedded in the newly detected Mg II nebula. For the purposes of clarity, the five galaxies are designated by letters¹, where galaxy A is the most massive one (see Sect. 3.2). The other letters (from B to E) are attributed to galaxies depending on their projected distance to galaxy A (from the closest to the furthest, see middle panel of Fig. 1). We note that galaxy B has a lower redshift confidence level ($Z_{\text{CONF}} = 1$) than the other group members in the MUSE UDF catalog (Bacon et al., in prep.). Based on a careful analysis of the [O II] NB image (see Sect. 5.1) – where [O II] emission is clearly detected at the position of galaxy B when setting low flux cuts – and the rather well constrained photometric redshift ($1.25 < z_{\text{BPZ}} < 1.47$), we are confident about the redshift of this source and therefore about its group membership.

The group members are located within ≈ 50 kpc in projection and 120 km s^{-1} in redshift space (middle and right panels of Fig. 1). The projected distances and velocity offsets from galaxy A are (9 kpc, $+43 \text{ km s}^{-1}$), (13 kpc, -76 km s^{-1}), (37 kpc, $+2 \text{ km s}^{-1}$), and (41 kpc, $+43 \text{ km s}^{-1}$) for galaxies B, C, D, and E, respectively (see right panel of Fig. 1 and Table 1). Galaxies B and C are aligned with the major axis of galaxy A. Similarly, galaxy D is aligned with the major axis of galaxy E.

A virial mass of $\approx 10^{11} M_{\odot}$ was estimated from the velocity dispersion of the five members using the gapper method (Beers et al. 1990; Cucciati et al. 2010; Abril-Melgarejo et al. 2021). The dynamical halo mass becomes $\approx 10^{12.3} M_{\odot}$ when considering the 14 galaxies of the larger structure. The corresponding virial radius is ≈ 50 kpc (Eq. (1) of Lemaux et al. 2012) for the five-member group (≈ 150 kpc when considering the whole structure). Alternatively, using the stellar to halo mass relation from Behroozi et al. (2019) at the redshift of the group, we calculated a DM halo mass of $10^{11.5} M_{\odot}$ for the central galaxy A (with 0.3 dex of uncertainties) and a virial radius of ≈ 90 kpc. Those two methods therefore indicate that the five group members reside inside the virial radius of the group and galaxy A.

Galaxy A is also the brightest galaxy of the larger structure in the UDF (≈ 24.5 mag in the *F775W* filter) and therefore corresponds to the dominant “brightest group galaxy” (BGG) near the halo center where the Xray intensity peak tracing the IGrM is usually observed in Xray studies (Oppenheimer et al. 2021). The BGGs have been found to have disk-like morphologies (Moffett et al. 2016), which seems to be the case for galaxy A given its elongated shape (Sect. 3.2). We note that one galaxy located at ≈ 1 Mpc outside the MUSE deep field and at the

¹ The MUSE identifiers in Inami et al. (2017) are #32, #121, #77 and #65 for the galaxies A,C,D and E, respectively. The galaxy B is a newly detected source which is in the MXDF catalog only (Bacon et al., in prep.) and has the identifier #8493.

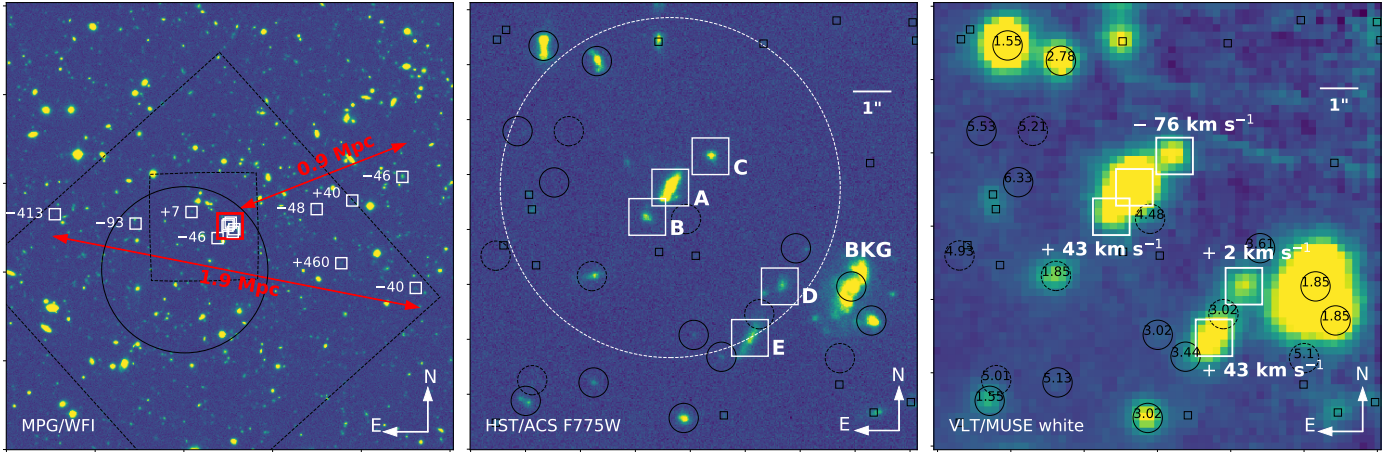


Fig. 1. Galaxy group and its environment. *Left:* galaxies belonging to the same structure (see Sect. 3.1) as the galaxies embedded in the $z \approx 1.31$ Mg II nebula (shown as the red square) are indicated by white squares. The velocity relative to the most massive galaxy of the group (galaxy A, see *middle panel*) is indicated for each galaxy in km s^{-1} . The structure spans ≈ 2 Mpc in projection and the galaxies appear well aligned. A ≈ 1 Mpc substructure extending beyond the MUSE Deep Field (big black dashed square) is also indicated. The MXDF and udf-10 fields are shown with the solid circle and small dashed square, respectively. The background image is an MPG(ESO2.2m)/WFI image (Hildebrandt et al. 2006). *Middle:* zoom-in window of $(12''.2 \times 12''.2)$ on the five galaxies around which the Mg II nebula has been discovered. The galaxies are designated by letters where galaxy A is the most massive one (see Sect. 3.2), and the other letters (from B to E) are attributed to galaxies depending on their projected distance to galaxy A (from the closest to the furthest). The white circle is centered on galaxy A and has a radius of 40 kpc ($\approx 5''$). The background image is the HST/ACS $F775W$ image. The galaxies in the Rafelski et al. (2015) HST catalog without a MUSE redshift are indicated by the small black squares. The spectroscopic MUSE redshifts with confidence 1 and >1 (Bacon et al., in prep.; see also Inami et al. 2017) are indicated as black dashed and solid circles, respectively. The redshift values are indicated on the *right panel*. The label “BKG” points to a bright background galaxy acting as a sightline for absorption line studies of the Mg II nebula (see Sect. 4.2). *Right:* same spatial window as the *middle panel* showing the MUSE white light image of the group (white squares) and its surrounding. The positions of the galaxies B, C, D and E relative to galaxy A in velocity space are indicated in white. The redshift measurement procedure of the group members is described in Sect. 3.3.

outskirt of the detected structure is brighter (≈ 23 mag in the $F775W$ filter), with a stellar mass of $\approx 10^{10.6} M_{\odot}$ (Urrutia et al. 2019).

The middle panel of Fig. 1 shows the $F775W$ HST/ACS image of the group and its close environment in a $\approx 100 \times 100 \text{ kpc}^2$ ($12''.2 \times 12''.2$) window. The known photometric redshift ranges (BPZ) from the Rafelski et al. (2015) catalog are indicated for objects without MUSE spectroscopic redshifts. Galaxies with MUSE redshifts (Bacon et al., in prep.; see also Inami et al. 2017) are shown as solid or dotted black circles, depending on the redshift confidence level (middle and right panels of Fig. 1). The systemic redshifts of the five group members have been measured for this analysis in Sect. 3.3.

A bright background galaxy ($m_{AB,F775W} = 23.4$), referred to as “BKG”, spectroscopically confirmed to be at $z \approx 1.8463$ (MUSE ID #18 in Inami et al. 2017) is located at an impact parameter of $2''.2$ (or 19 kpc) from galaxy D, i.e., within the virial radius of galaxy A. This galaxy produces a strong continuum in the MUSE data and thus appears like an interesting sight line for the analysis of the Mg II nebula (see Sect. 4.2) in absorption. No other background projected neighbors, inside a 200 kpc radius around galaxy A, have strong enough detected continua in the MUSE data to allow the detection of absorption features.

3.2. Global properties of the galaxies

The stellar masses and star formation rates (SFR) of the group members were estimated by fitting their spectral energy distribution (SED) from HST photometry ($F225W$, $F275W$, $F336W$, $F435W$, $F606W$, $F775W$, $F850LP$, $F105W$, $F125W$, $F140W$, $F160W$, Rafelski et al. 2015) using MAGPHYS (da Cunha et al. 2008, 2015). This version of MAGPHYS uses a Chabrier (2003)

initial mass function (IMF), Bruzual & Charlot (2003) stellar models and Charlot & Fall (2000) dust attenuation. Moreover, the star formation histories (SFH) are considered exponential with random bursts superimposed. The resulting properties of the group members are given in Table 1. With a highest stellar mass of $\log M_{*}/M_{\odot} = 9.35^{+0.16}_{-0.14}$, the galaxies of this group are rather low-mass objects. This is in contrast with previous studies where the systems embedded in ionized nebulae usually host several galaxies with stellar masses higher than $10^{10} M_{\odot}$ (e.g., Epinat et al. 2018; Helton et al. 2021). Despite the high uncertainties on the SFR values of the less massive galaxies as well as the few constraints on the so-called “star formation main sequence” for low-mass galaxies at $z = 1.3$, the five galaxies can be considered as normal star forming galaxies (Whitaker et al. 2014; Boogaard et al. 2018).

The two most massive galaxies (A and E) are seen edge-on with projected axis ratios of $q = 0.329 \pm 0.005$ and $q = 0.247 \pm 0.015$, respectively (corresponding to inclinations of 71 and 76°, respectively, when considering a flat disk model) and a similar position angle $\text{PA} \approx -27^{\circ}$ (measurements from van der Wel et al. 2012 in the WFC3/ $F105W$ band, see Table 1). The three other galaxies have rounder shapes with $q \gtrsim 0.7$ and smaller sizes.

Regarding possible AGN contamination, we refer to Feltre et al. (2018, their Sect. 3.3). By cross-matching their Mg II sample (which includes four out of five of our group members) with the 7 Ms Source Catalogs of the Chandra Deep Field South Survey (Luo et al. 2017), the authors found an X-ray counterpart for only one source, which is not one of the galaxies analyzed in this paper. We do not find any X-ray counterpart for galaxy B (not in Feltre et al. 2018). The faintness and low stellar masses of the galaxies (Table 1), the reasonable line widths, and the lack of typical AGN lines such as C II] $\lambda 3425$, also indicate that the

Table 1. Physical properties of the five group members.

ID	MID	z_{syst}	M_{UV} [mag]	$\log_{10}(M_*)$ [M_{\odot}]	$\log_{10}(\text{SFR})$ [$M_{\odot} \text{ yr}^{-1}$]	Δv_A [km s^{-1}]	Δp_A [kpc]	q	PA [$^{\circ}$]	EW ₂₇₉₆ [\AA]
A	32	1.30661	-19.09 ± 0.04	$9.35^{+0.16}_{-0.14}$	$0.5^{+0.2}_{-0.06}$	0	0	0.329 ± 0.005	-26.5 ± 0.4	1.5 ± 0.2
B	8493	1.30695	-16.71 ± 0.17	$8.49^{+0.08}_{-0.1}$	$-0.48^{+0.08}_{-0.06}$	+43	9.2	0.543 ± 0.038	52 ± 6	4.9 ± 0.9
C	121	1.30603	-17.09 ± 0.11	$8.55^{+0.1}_{-0.05}$	$-0.55^{+0.07}_{-0.02}$	-76	13.0	0.559 ± 0.032	54 ± 6	1.3 ± 0.2
D	77	1.30663	-17.49 ± 0.16	$8.76^{+0.14}_{-0.06}$	$0.24^{+0.18}_{-0.32}$	+2	37.3	0.644 ± 0.038	-12 ± 6	1.7 ± 0.4
E	65	1.30694	-17.85 ± 0.11	$8.94^{+0.14}_{-0.07}$	$-0.07^{+0.1}_{-0.01}$	+43	41.2	0.247 ± 0.015	-28 ± 1	1.7 ± 0.4

Notes. ID: label of the galaxy group members as shown in Fig. 1 (middle panel). MID: MUSE identifier (Inami et al. 2017; Bacon et al., in prep.). z_{syst} : systemic redshift as measured in Sect. 3.3. M_{UV} : absolute UV magnitude close to 1500 \AA restframe measured in the F336W HST band (Rafelski et al. 2015). $\log_{10}(M_*)$: logarithm of the stellar mass in M_{\odot} . SFR: star formation rate in $M_{\odot} \text{ yr}^{-1}$. The parameters M_* and SFR are based on SED fitting (Sect. 3.2). Δv_A : velocity offset from galaxy A in km s^{-1} . Δp_A projected distance from galaxy A in kpc. q : projected axis ratio. PA: position angle in degree. The galaxy morphological parameters q and PA are from van der Wel et al. (2012) and measured in the WFC3/F105W band. EW₂₇₉₆: rest frame equivalent width of the Mg II $\lambda 2796$ line in \AA as measured by pyplatefit (Bacon et al., in prep., Sect. 3.3).

presence of an AGN in those galaxies is unlikely. Nonetheless, the presence of low luminosity or heavily obscured AGN cannot be entirely excluded.

3.3. Integrated spectra

Figure 2 displays a zoom-in HST/ACS view in the F775W filter (first column) and the integrated spectrum (second column) of each group member. We display the ODHIN spectrum (see Sect. 2), which makes use of the higher spatial resolution of the HST images to perform a deblending of sources in the MUSE cubes. The deblending feature of the ODHIN software is particularly needed here as the emission from the [A,B,C] and [D,E] galaxies are overlapping at MUSE resolution (see Fig. 3a). The corresponding integrated error spectra are shown in gray in each panel and errors are directly indicated on the spectra (in gray) in the insets.

The systemic redshifts of the group members are calculated from those spectra by fitting the [O II] doublet, which is the line with the highest signal-to-noise ratio (S/N) for all sources. We fit the doublet using two Gaussian functions with fixed peak separation and considering the variance of the MUSE cube. The errors on the fit parameters are estimated using bootstrapping; 1000 realizations of the line doublet are generated using the error spectrum assuming that the errors are normally distributed around the observed flux values. The resulting errors on the systemic redshift measurements (given in Table 1) are on the order of 10^{-5} .

All the galaxies show a stellar continuum with clear [O II] and Mg II emission (see inset panels in Fig. 2). We note that both Mg II and [O II] doublets are impacted by significant sky-lines. The spectrum of the most massive galaxy (A) displays other emission lines like [Ne III], H8, H ϵ as well as C II $\lambda 2326$ and [O II] $\lambda \lambda 2470, 2471$. Moreover, the Balmer absorption is detected on the red side of the [O II] line. The Mg II line of galaxy A shows a P-Cygni profile and a detection of the Mg I $\lambda 2852$ absorption line. At the redshift of the group, the UV transitions Fe II* ($\lambda 2365, \lambda 2396, \lambda 2612, \lambda 2626$) are observable with MUSE; most of the corresponding emission and absorption lines are detected in the spectrum of galaxy A (top-right panel of Fig. 2). Using the shallower mosaic and udf-10 MUSE data-cubes, Feltre et al. (2018) and Finley et al. (2017) respectively studied the Mg II and Fe II* emission of star-forming galaxies. Both studies classified the Mg II doublet of galaxy A as a

P-Cygni profile. In Feltre et al. (2018), galaxy C is classified as an Mg II emitter and galaxy E as a non-detection. The redshift confidence level of galaxy D was not high enough to be included in their sample and, as mentioned before, galaxy B is a newly detected galaxy that was not in their sample either.

The Mg II rest frame equivalent widths (EW) of the Mg II $\lambda 2796$ line calculated using pyplatefit in Bacon et al. (in prep., see also Bacon et al. 2021) are given in Table 1. The Pyplatefit software performs a stellar continuum fit using a simple population model (Brinchmann et al. 2013) and fits the emission and absorption lines after subtracting the continuum. The intrinsic Mg II $\lambda \lambda 2796, 2803$ absorption in the spectra of O/B stars is expected to be very weak ($\approx 0.6 \text{\AA}$ for population ages up to a few 100 Myr, Henry et al. 2018) and therefore cannot explain the observed absorption. The Mg II absorption is most likely dominated by continuum photons pumping by Mg⁺ ions along the line of sight, followed by isotropic re-emission out of line of sight, as a result of resonant scattering. This leads to a net absorption in front of the continuum sources, and possibly large-scale diffuse emission that fills the absorption troughs. This effect is called emission infill (e.g., Prochaska et al. 2011; Scarlata & Panagia 2015; Zhu et al. 2015; Finley et al. 2017). The EW values reported in Table 1 are not corrected for this effect. We also did not attempt to correct our Mg II observations for emission infill, nor the spectra or the narrow bands (Sect. 4.1) because it would inevitably imply an astrophysical assumption that absorbing and emitting regions are spatially different and can be separated. Since the physical meaning of any simple correction is limited, we present the full complexity of these Mg II data and highlight the novelty of the detection of such diffuse large-scale structures. Our non-corrected measurements enable (i) direct and simple comparisons with other studies in terms of detection limits, observed flux, and spatial variations, as well as (ii) the reproducibility of our results and measurements. Finally, we note that the absorption correction has an effect at galaxy scale only, so our conclusions about the large scale intragroup medium are not impacted by the non-correction.

In summary, we are looking at a compact group of five low-mass members located at the center of a larger structure at high redshift. Taking advantage of the deep MXDF data, we analyzed the Mg II properties of the IGrM in the next section (Sect. 4) and its [O II] and Fe II* properties in Sect. 5.

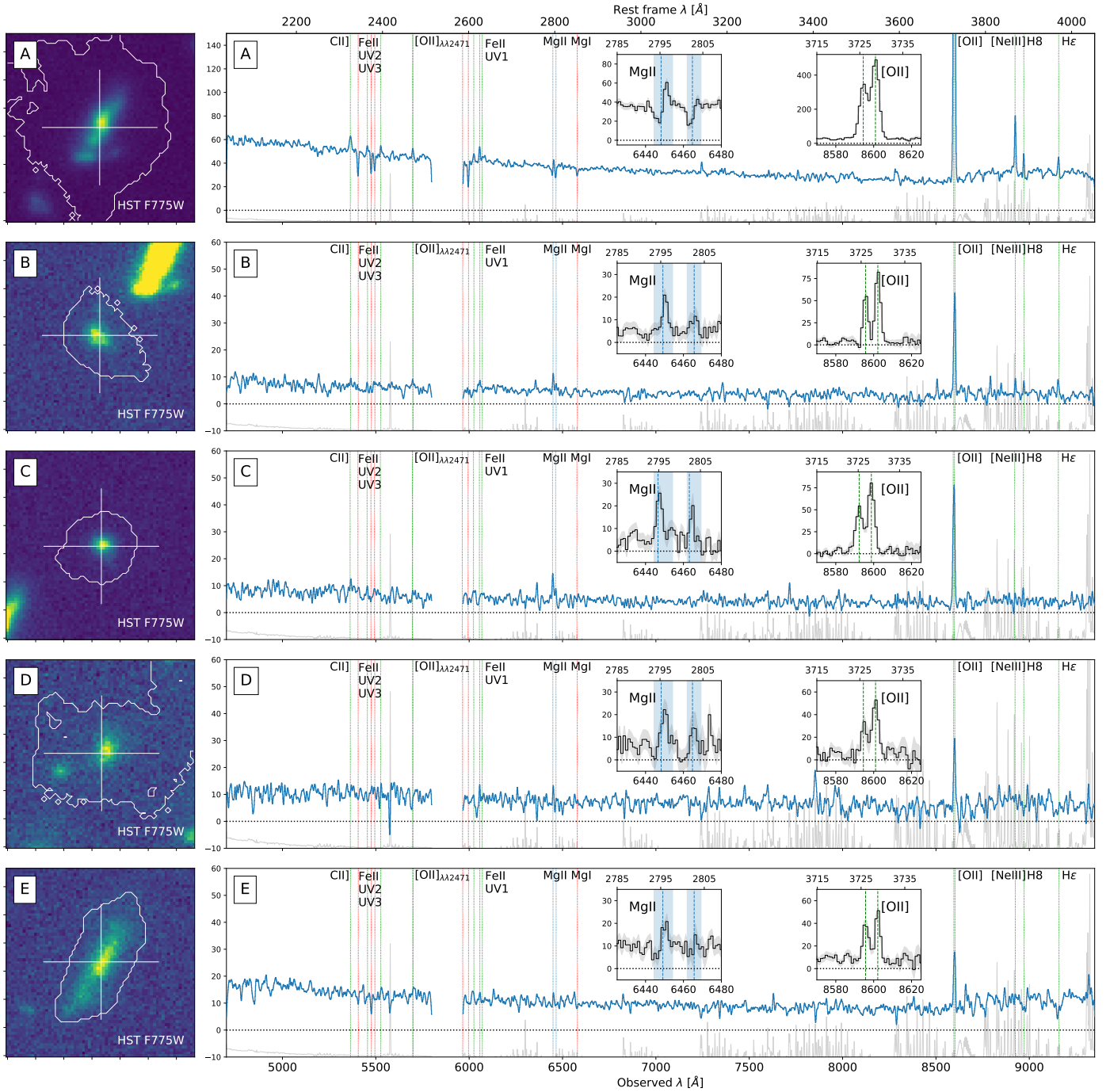


Fig. 2. Integrated spectra of the five group members. *Left:* $2'' \times 2''$ HST/ACS F775W images of the five group members. The central white cross and contour indicate the HST coordinates and segmentation map from Rafelski et al. (2015). *Right:* integrated spectrum (units of $10^{-20} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$) of the galaxy shown in the corresponding left panel extracted within the HST segmentation map contour using the ODHIN software (Bacher 2017, see Sect. 3.3) and smoothed with a 7 \AA FWHM Gaussian function. The positions of several emission (absorption) lines are indicated by vertical dotted green (red) lines. The gray unsmoothed spectra show the 1σ uncertainties which have been offset for readability. The unsmoothed Mg II and [O II] lines (vertical dashed lines) are shown in the insets. The vertical blue shaded areas indicate the spectral widths used to construct the Mg II NB image shown in Fig. 3a.

4. Mg II intragroup nebula

4.1. Mg II narrow-band image

The Mg II structure was extracted in a subcube of $12''.2 \times 12''.2$ ($\approx 100 \times 100 \text{ kpc}^2$ at $z \approx 1.31$) centered on the coordinates $\text{RA} = 53^\circ 09' 25''$, $\text{Dec} = -27^\circ 46' 44''$. We estimated the continuum by performing a spectral median filtering on the subcube

using a wide spectral window of 200 spectral pixels (250 \AA). After subtracting this continuum-only cube from the original one, we obtained an emission lines-only cube, from which we optimally created the Mg II NB image of the group as follows.

In order to encompass all the Mg II emitting flux from the five galaxies of the group while limiting the noise, we adopted the following method: (i) the Mg II line is extracted by integrating

inside a circular aperture including the five group members, centered on the continuum-subtracted subcube, which maximizes the integrated flux of the Mg II $\lambda 2796$ line ($r = 3.2''$), (ii) the NB image of the $\lambda 2796$ Mg II line is created by summing the continuum-subtracted subcube over the wavelength range 6445.0–6453.75 Å ($\approx 400 \text{ km s}^{-1}$) delimiting the line; the borders of the line are determined by wavelengths for which the flux reaches zero. The same procedure is applied to extract the $\lambda 2803$ line of the Mg II doublet (6462.5–6468.75 Å, $\approx 290 \text{ km s}^{-1}$). The total Mg II emission NB image is finally obtained by adding the $\lambda 2796$ and $\lambda 2803$ NB images. A broader band Mg II image is presented in Fig. C.1 and confirms that our method provides an S/N-optimized NB image without missing flux (see the residuals image in the right panel of Fig. C.1).

The NB image shown in Fig. 3a has been smoothed using a $0''.6$ (3 MUSE spaxels) FWHM Gaussian in order to increase the S/N while keeping the spatial resolution as good as possible (PSF of $0.52''$ FWHM, see Fig. B.1). The 1σ significance level corresponds to a surface brightness (SB) of $1 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. As shown in Fig. 3b, the exposure time ranges from ≈ 30 to 90 h through the Mg II nebula (white contour). This explains the noisier regions in the upper right corner of the NB image. As a consequence, we refrain from estimating the noise from empty regions around the nebula to compute the limiting SB contours and use the propagated variance of the MUSE cube in each pixel after smoothing. The total Mg II flux within 2σ significance level is $7.48 \pm 0.50 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ and corresponds to a total luminosity of $8.0 \pm 0.5 \times 10^{40} \text{ erg s}^{-1}$ without dust correction.

The eight 1.25 \AA slices (from -177 to $+286 \text{ km s}^{-1}$) of the MUSE subcube are summed to create the $\lambda 2796$ Mg II NB image are shown in Fig. 4. We note that because of the line spread function, those slice images are somewhat correlated. In spite of this, the spatial variations are still visible between adjacent slices (see Sect. 4.4).

4.2. Mg II morphology

Figure 3a shows the extended nature of the Mg II emission around this group of five galaxies. The observed nebula encompasses a projected area of 15 arcsec^2 , corresponding to 1000 kpc^2 at the redshift of the group, above the 2σ SB threshold of $2 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. The detected nebula resides within the virial radius of galaxy A (Sect. 3.1) and has an elongated shape linking the five group members with maximal projected extent reaching $\approx 70 \text{ kpc}$ along the A/E direction and $\approx 40 \text{ kpc}$ along the B/C direction.

The strongest Mg II emission coincides spatially with galaxy C. Another emission peak is observed in between galaxies A and B. While globally the nebula appears aligned with the minor axis of galaxy(ies) A (and E) as expected e.g., in a scenario where Mg II is emitted by biconical outflows launched perpendicular to the galaxy disk (Bouché et al. 2012), here the two intensity peaks are aligned with the major axis of the most massive galaxy A. This point is discussed in Sect. 6.2.

The 2D map reveals a lack of flux at the positions of galaxies A and BKG (background galaxy). Those features are artificially created by the continuum subtraction procedure (see Sect. 4.1): the absorption features appear negative once the continuum has been subtracted. The fact that Mg II in emission is observed at the position of galaxy A in the $[+54; +112]$ and $[+112; +170] \text{ km s}^{-1}$ channel maps of Mg II $\lambda 2796$ (Fig. 4) indicates that both absorption and emission coexist in this region. The lack of flux detected

at the position of galaxy A in the summed NB image (Fig. 3a) therefore corresponds to regions where the absorption is dominant over the emission. As indicated in Sect. 3.3, our study is not aimed at obtaining a correction for the emission inflly effects here because our study focuses on the larger scale Mg II nebula.

Interestingly, while Mg II emission lines are detected in the integrated spectra of the galaxies D and E, the Mg II map shows no flux at the location of galaxy E and offset emission near galaxy D. This can be caused by the galaxy's absorption being completely filled by the emission (e.g., Zabl et al. 2021). The spectra of the galaxies D and E indeed show hints of low S/N absorption lines, which are within the Mg II NB spectral bandwidth (see Fig. 2). We also note that the segmentation map of galaxy D extends beyond the stellar body of the galaxy (Fig. 2). This implies that the Mg II emission observed in the spectrum of galaxy D likely originates from circumgalactic regions (as seen in the Mg II map in Fig. 3a).

The ODHIN spectrum of BKG ($z \approx 1.8463$, see Fig. 1c), which is located outside the detected Mg II emission nebula and at a projected distance of 19 kpc from galaxy D, shows Mg II absorption at the same redshift as the nebula (see Fig. 3c). Such a detection indicates that the Mg II gas even extends beyond the detected Mg II emission nebula. Using the line fitting procedure described in Sect. 3.3, we measured an Mg II $\lambda 2796$ rest frame EW of $0.5^{+0.6}_{-0.4} \text{ \AA}$ for the absorption. This measurement falls well on the Mg II EW versus impact parameter relation found in several absorption studies (e.g., Lundgren et al. 2021). Using the Ménard & Chelouche (2009) relation between the HI column density and the rest frame Mg II $\lambda 2796$ EW we obtained the rough estimate of $10^{19} \text{ atoms cm}^{-2}$ for the extended gas column density in this line of sight.

The spatially resolved MUSE view of the nebula also reveals that the Mg II distribution is different from that of the continuum: the Mg II gas is more extended than the stellar content of the galaxies (see Sect. 4.3) and a projected bridge seems to link galaxy subgroups [A,B,C] and [D,E] (see Sect. 4.3). This bridge might not be a projection effect because of its tentative detection ($S/N \lesssim 2$) in the single 1.25 \AA (58 km s^{-1}) cube slice between $+54$ and $+112 \text{ km s}^{-1}$ relative to the systemic redshift of galaxy A (bottom left panel of Fig. 4). We can also mention that the Mg II spatial distribution is not homogeneous in between the galaxies: several intensity peaks are detected at $S/N > 3$. They suggest the presence of gas clumps or satellite galaxies (undetected in the HST and MUSE data) residing in the IGrM (see discussion in Sect. 6.2.1).

4.3. Radial Mg II surface brightness profile

In order to highlight the extended nature of the Mg II emission around and in between the galaxy group members, we computed azimuthally averaged radial surface brightness profiles of the Mg II and continuum emission centered on galaxy A (Fig. 5, first panel). The Mg II profile (blue) is measured on the unsmoothed Mg II NB image (Sect. 4.1) after masking pixels located outside an ellipse corresponding to the 1σ significance level in order to increase the S/N in the two-pixel-wide annuli (or truncated annuli) used for the aperture photometry (see inset panels). The width of the annuli ($0.4''$) is comparable to the spatial resolution element of the data (PSF FWHM of $0.52''$).

The Mg II continuum radial SB profile (black) is measured from a continuum image extracted from the continuum-only cube within a window of $\pm 20000 \text{ km s}^{-1}$ around the Mg II doublet. This broad spectral window does not include strong

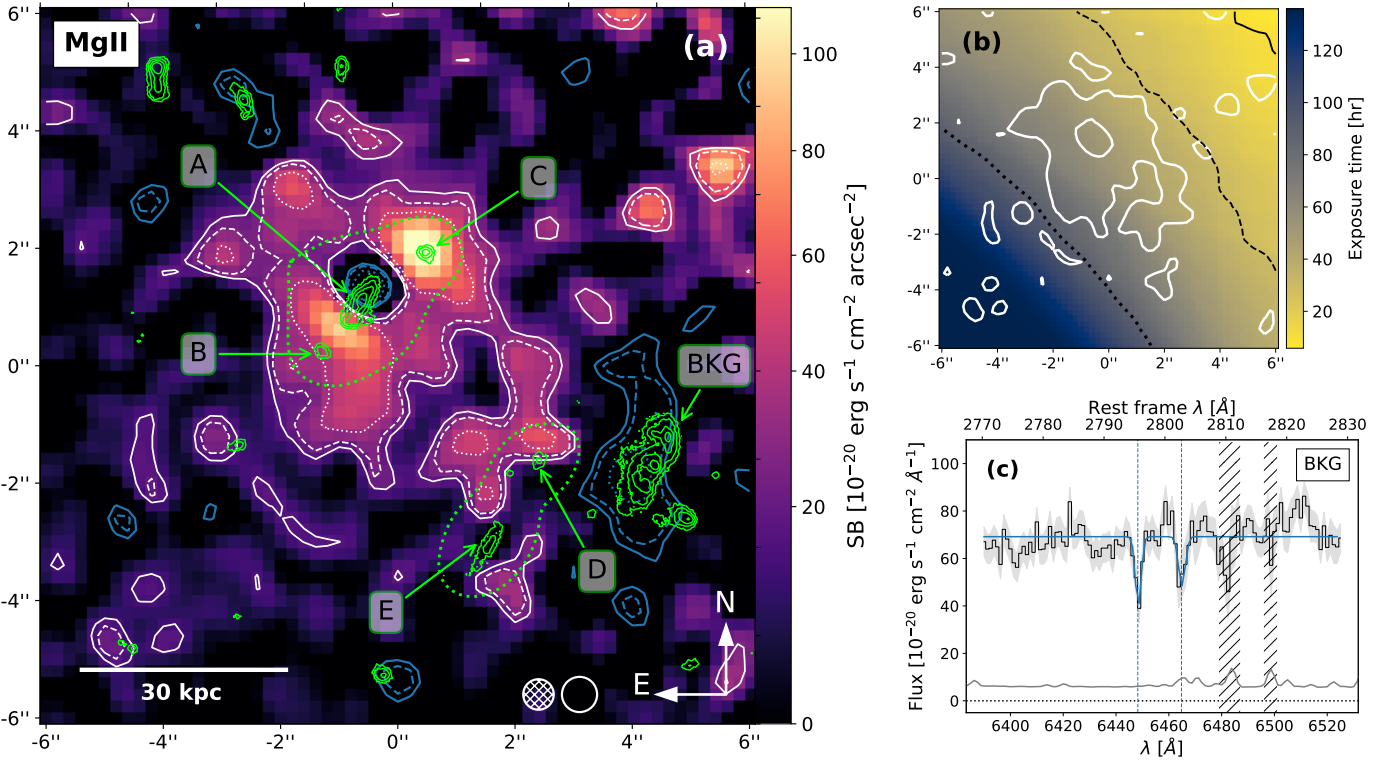


Fig. 3. Metal-enriched intragroup medium revealed by MUSE. (a) Mg II narrow-band image of the group (see Sect. 4.1) smoothed with a $0.6''$ FWHM Gaussian and plotted with a power-law stretch. The white contours correspond to Mg II significance levels of 1.5 , 2 and 3σ (solid, dashed, and dotted, respectively) where 1σ corresponds to a SB level of $1 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. The blue contours indicate the negative SB values (same coding as the positive contours) corresponding to the absorbing regions (see Sect. 4.2). The green contours trace the continuum of the group members close to Mg II and the neighboring galaxies as detected in the HST/ACS $F775W$ image. The dotted green contour corresponds to the outer HST continuum contours of the group members set at the MUSE resolution. The FWHM of the MUSE PSF and smoothing kernel are shown with a hatched and empty circle, respectively, at the bottom-right of the figure. (b) Exposure map centered on the group. The solid, dashed, and dotted black contours indicate 10, 30, and 90 h depth, respectively. The white contour corresponds to the Mg II significance level of 1.5σ . (c) Mg II absorption in the spectrum of the background galaxy BKG located at an impact parameter of 19 kpc from galaxy D in projection and at the nebula redshift. The blue line is the best-fit model of the Mg II doublet absorption (Sect. 4.2). The 1σ uncertainties are shown in gray and the hatched areas indicate the position of strong sky lines and noise peaks. The systemic redshift of galaxy D is indicated by vertical dashed lines.

absorption or emission lines and ensures a good enough S/N. The neighboring sources were rigorously masked in order to avoid contamination and the same elliptical mask used on the Mg II image was applied to compute the continuum SB profile. Errors were measured in each (truncated) annulus using the estimated variance from the MUSE data cube. To aid the visual comparison, the continuum profile has been rescaled to the Mg II SB at the position of galaxy C, which is where the brightest Mg II intensity peak is observed. Galaxy C is unlikely to be affected by Mg II absorption as its measured Mg II line ratio $\lambda 2796 \text{ \AA} / \lambda 2803 \text{ \AA}$ is very close to two, which is the expected value for optically thin emission (see Appendix D). Moreover, no fluorescent Fe II* lines are significantly detected for galaxy C which is in good agreement with no infilling effect and therefore no absorption (see Appendix D, Mauerhofer et al. 2021).

The Mg II SB profile appears more extended than the continuum, especially in between the galaxies [A,B,C] and [D,E] – radial positions are indicated by vertical dashed lines – where a low SB projected “bridge” at $\approx 2 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$ level can be identified. In order to differentiate between the extended emission around the galaxy subgroup [A,B,C] and in between the galaxy subgroups [A,B,C] and [D,E], we compute the radial SB profiles above and below the [A,B,C] axis (see insets for illustration) in the middle and right panels of Fig. 5,

respectively. The Mg II emission appears more extended than the rescaled continuum in both directions (above and below the [A,B,C] axis), confirming the detection of diffuse Mg II emission around the group members as well as a low SB Mg II bridge ($\approx 2 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$) detected with more than 1.5σ significance spanning $\approx 50 \text{ kpc}$ in projection between the galaxy subgroups [A,B,C] and [D,E] (see Fig. 3a). We go further into the analysis of the Mg II nebula in the next section by looking at the spatial variation of the spectral properties of Mg II.

4.4. Spatially resolved Mg II properties

The kinematics of the gas connecting the galaxy group can provide crucial information on its nature and origin. To reveal the spatial variations of the Mg II spectral profile within the nebula, we built a 2D binned map using the weighted Voronoi tessellation method (Cappellari & Copin 2003; Diehl & Statler 2006). This method allows us to increase the S/N in the surroundings of the galaxies where the surface brightness is the lowest. The binning was performed on the Mg II NB image constructed from the MXDF datacube (see Sect. 4.1), for which the pixels with $S/N < 1.5$ (corresponding to the outer SB contour in Fig. 3a) have been masked in order to remove the noise-dominated pixels. The resulting binned Mg II map consists of

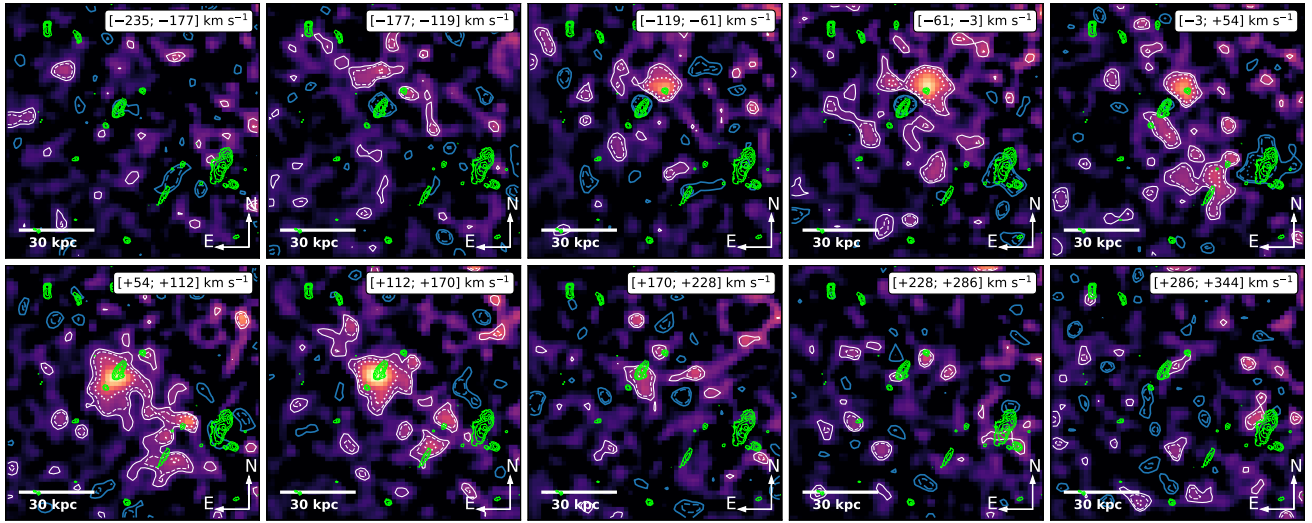


Fig. 4. Channel maps around the Mg II $\lambda 2796$ emission. Each map corresponds to one MUSE subcube slice (1.25 \AA or 58 km s^{-1} at the redshift of the group). The velocity window relative to the systemic redshift of galaxy A is indicated on each panel. Contours are the same as in Fig. 3a. The eight central slices (from -177 to $+286 \text{ km s}^{-1}$) are summed to create the optimized $\lambda 2796$ Mg II NB image shown in Fig. 3a (see Sect. 4.1).

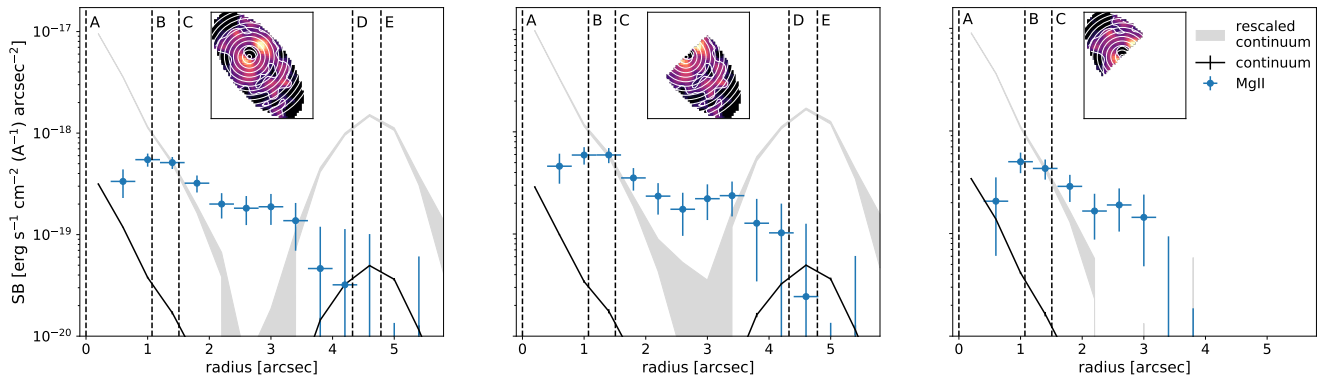


Fig. 5. Azimuthally averaged radial surface brightness profile of the Mg II nebula centered on galaxy A (blue data points). For comparison the profile of the Mg II continuum ($\pm 20000 \text{ km s}^{-1}$ around the Mg II doublet) shown in black is also rescaled (gray) to the Mg II SB at the location of galaxy C (where Mg II absorption is unlikely, see Appendix D). The positions of the group members are indicated with vertical dashed lines. The radial profiles have been calculated using unsmoothed Mg II and continuum images after masking the pixels outside the 1σ significance level approximated with an elliptical contour as shown in the inset (left). *Middle and right panels:* radial profiles after masking the pixels above and below the [A,B,C] axis, respectively. In the insets, we also show the two-pixel annuli used to construct the profiles as white circles.

20 bins with $S/N > 3$ and probes the detected Mg II nebula above $\approx 2 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$.

The spectral extraction is performed from the MUSE continuum-subtracted cube smoothed with a $0.6''$ FWHM Gaussian in order to increase the S/N while keeping a good spatial resolution (see Appendix B). We extracted the Mg II doublet in each resulting bin and measured the properties of the emission lines by modeling the doublet as a sum of two Gaussian profiles with a fixed peak separation.

The errors on the line parameters (peak position and FWHM) were estimated using bootstrapping. For each segmented area, we generated 1000 realizations of the extracted line where each pixel is randomly drawn from a normal distribution centered on the original pixel value and with standard deviation derived from the estimated noise value. The noise of each extracted spectrum corresponds to the standard deviation of the data measured in two 200 \AA wide spectral windows located on each side of the Mg II doublet. We note that due to smoothing as well as the PSF, the line parameters between neighboring pixels are somewhat correlated. However, the spatial variations are still visible.

Figure 6 (top left panel) shows the resolved map of the Mg II $\lambda 2796$ peak position relative to the systemic redshift of galaxy A. The observed Mg II velocity gradient suggests an overall rotation of the structure along the major axis of galaxy A consistent with an extension of the ISM rotation (as probed by the [O II] emission, see Sect. 5.1) at the CGM scale within which the satellite galaxies B and C are embedded. This ISM/CGM co-rotation is also observed in simulations and observations (e.g., Smith et al. 2019; Zabl et al. 2019, 2021; Ho et al. 2020; Tejos et al. 2021).

The top-right panels of Fig. 6 show the unsmoothed Mg II doublet extracted in the elliptical areas traced in purple and designated by the same number on the map. These apertures are independent of the Voronoi bins and aim to show the Mg II line profile in some diffuse regions of the nebula. The velocity variations are small ($< 50 \text{ km s}^{-1}$) in the diffuse regions.

The kinematics of the nebula appear to be overall shifted towards higher velocities (by $+50 \text{ km s}^{-1}$ on average) compared to the [O II] emission (bottom left panel, see Sect. 5.1). The highest velocity variations are observed along the minor axis of

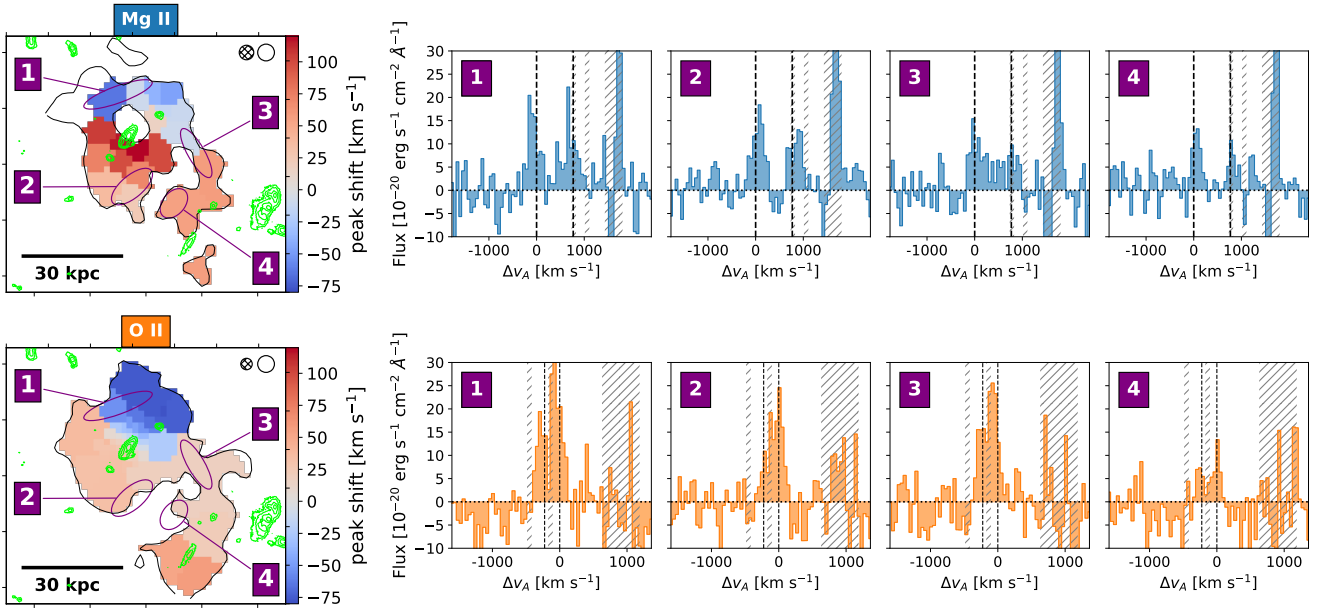


Fig. 6. Spatially resolved Mg II and [O II] properties. *Left:* binned line-of-sight velocity maps of the Mg II (*top*) and [O II] (*bottom*) emission relative to the systemic redshift of galaxy A (Sects. 4.4 and 5.1). The black contours correspond to the 1.5σ significance level. The green contours trace the galaxies detected in the HST/ACS $F775W$ image. The FWHM of the MUSE PSF and smoothing kernel are shown with a hatched and an empty circle, respectively, at the *top right of the panels*. The diverging Mg II and [O II] colormaps are centered on the systemic redshift of galaxy A and have the same dynamical range to ease the visual comparison (Sect. 5.1). *Right:* Mg II (*top*) and [O II] (*bottom panels*) doublet extracted in the elliptical areas traced in purple and designated by the same number on the map (*left panels*). The vertical dashed lines correspond to the systemic redshift of galaxy A. The vertical hatched bands indicate the presence of sky lines.

galaxy A and reach $\approx +120 \text{ km s}^{-1}$ and -80 km s^{-1} with respect to the systemic redshift of galaxy A. As discussed in Sect. 6.2.1), this supports the presence of an outflow emerging from galaxy A.

5. [O II] and Fe II* emission

At the redshift of the group, MUSE covers the non-resonant collisionally excited [O II] doublet and the fluorescent Fe II* lines. The corresponding NB images and radial SB profiles are shown in Fig. 7 and are created following the same procedure as for the Mg II emission (Sects. 4.1 and 4.3, respectively).

5.1. Low-surface-brightness [O II] nebula

The [O II] NB map (Fig. 7, top left) is obtained by summing the continuum-subtracted cube in the wavelength range [8590, 8605] Å ($\approx 500 \text{ km s}^{-1}$). In order to improve the S/N, the image was smoothed using a $0''.6$ FWHM Gaussian function. The white contours correspond to the 1.5 , 2 , and 3σ significance levels with 1σ corresponding to $2.5 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. The total [O II] flux and luminosity within the 2σ significance level contour (dashed white line) are $5.95 \pm 0.07 \times 10^{-17} \text{ erg s}^{-1} \text{ cm}^{-2}$ and $6.36 \pm 0.07 \times 10^{41} \text{ erg s}^{-1}$, respectively.

The [O II] azimuthally averaged radial SB profiles are shown in the lower left panels of Fig. 7. The profiles are constructed following the same procedure as for Mg II (see Sect. 4.3) with an additional step for the construction of the continuum image where the cube slices within $\pm 20\,000 \text{ km s}^{-1}$ polluted by skylines have been masked. This is particularly important because at the redshift of the nebula, the [O II] line falls next to many sky lines. The [O II] doublet is itself contaminated by a skyline (see vertical gray line in the lower right panels of Fig. 6) preventing us from in-depth analysis.

When comparing the [O II] and rescaled continuum radial SB profiles in Fig. 7, only one data point is above the rescaled continuum profile with more than 2σ significance. The inner profile has the same shape as the continuum one implying that the [O II] emission is not significantly more extended than the continuum inside a radius of $\approx 2''$ from galaxy A. Extended [O II] emission at the same position as the Mg II bridge, namely, in between the galaxies [A,B,C] and [D,E], is detected with 2σ significance level in the [O II] NB image (Fig. 7, top left). However, due to averaging effects, the [O II] bridge is less significant on the radial profiles computed below the [A,B,C] axis, namely, in the direction of the [D,E] subgroup (first column third row of Fig. 7). The radial SB profile computed above the galaxies A, B and C (bottom left panel) indicates extended [O II] emission at radii $> 2.5''$ from galaxy A, suggesting the presence of an ionized [O II] halo around this galaxy subgroup. Comparisons of the Mg II and [O II] spatial extent and radial SB profiles are presented in Fig. 7 (top two right panels). Since the MUSE PSF is narrower at the [O II] wavelength than at the Mg II wavelength, we degraded the [O II] NB image by convolving it by a Moffat function whose parameters ($a = 0.22$ and $b = 2.87$, see Bacon et al. 2017) have been empirically determined so that the [O II] and Mg II images have the same PSF. While the [O II] and Mg II spatial distributions are consistent, the Mg II radial SB profiles (azimuthally averaged and in directions, bottom right panels) are all flatter than the [O II] profiles. This can be attributed to the resonant nature of the Mg II transition (see Sect. 6.2.2).

The binned line-of-sight velocity map of the [O II] emission is shown in the bottom left panel of Fig. 6. The map is constructed following the same procedure as for Mg II (Sect. 4.4) and consists of more bins (57) because of the higher S/N of the [O II] emission near the galaxies. At the position of the galaxies, the observed spatial variations of the [O II] peak positions are consistent with their redshift relative to galaxy A (right panel of

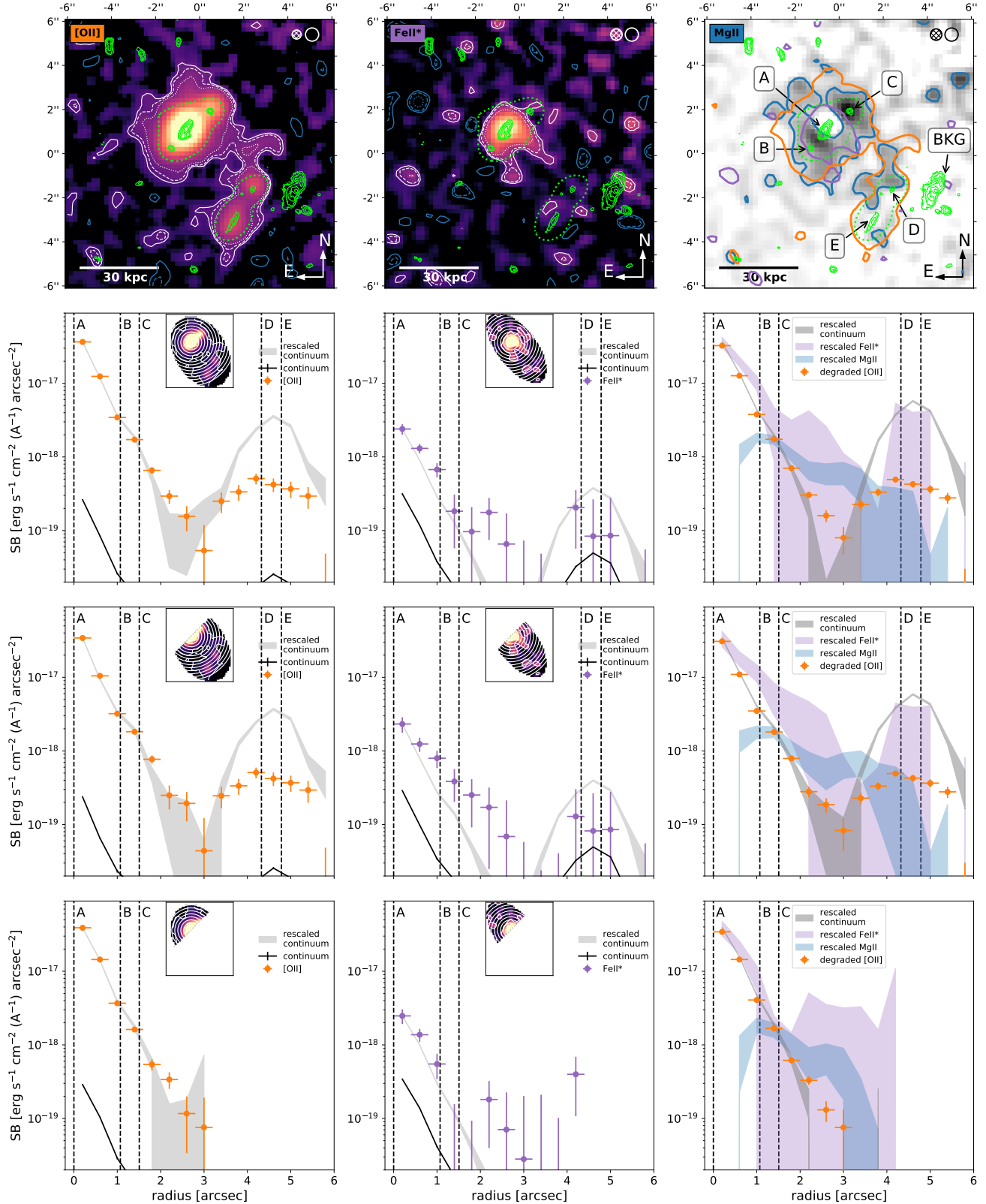


Fig. 7. [O II] and Fe II* spatial properties. *First column:* [O II] NB image (top, Sect. 5.1). Legend is the same as in Fig. 3a. The 1σ level corresponds to a SB of $2.5 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. *Bottom panels:* total and directional (below and above the [A,B,C] axis) azimuthally averaged radial SB profiles centered on galaxy A (from top to bottom, respectively, see also insets for illustration) of the [O II], [O II] continuum and rescaled [O II] continuum (orange, black and gray, respectively) emission. The radial positions of other group members with respect to galaxy A are indicated by the vertical dashed lines. *Second column:* same as first column but for Fe II* (purple). *Third column:* comparison of the Mg II, [O II] and Fe II* 2σ significance level contours (blue, orange, and purple, respectively) superimposed on the Mg II NB image (top). *Bottom panels:* show a comparison of the radial SB profiles where the Fe II* profile (shaded purple) has been rescaled to the [O II] profile at the SB peak and Mg II (shaded blue) at the position of galaxy C (Appendix D). For this plot the [O II] map was degraded to the resolution of the Mg II emission (Sect. 5.1).

Fig. 1). The lower right panels of Fig. 6 show the unsmoothed [O II] lines extracted in the same elliptical apertures as for Mg II (top panels). Interestingly, in region #4 the Mg II emission is stronger than [O II]. This can be explained by the resonant nature of the Mg II transition (see Sect. 6.2.2) resulting in Mg II being more extended than [O II]. The existence of the [O II] bridge is reinforced by the detection of the [O II] doublet in this area (region #3 in Fig. 6).

5.2. Extended Fe II* emission

The Fe II* NB image (top middle panel of Fig. 7) is obtained by summing the λ 2365, λ 2396, λ 2612 and λ 2626 NB images, of which the observed spectral windows are [5448.75, 5456.25], [5522.5, 5530.0], [6021.25, 6027.5], and [6041.25, 6066.25] Å, respectively. The total Fe II* flux and luminosity within the significance level contour of 2σ is $5.3 \pm 0.5 \times 10^{-18}$ erg s $^{-1}$ cm $^{-2}$ and $5.7 \pm 0.6 \times 10^{40}$ erg s $^{-1}$, respectively.

Given that the Mg II and Fe II* lines are close in wavelength, we consider the same stellar continuum (see Sect. 4.1). Therefore, the radial SB profiles of the continuum (black profile in the middle bottom panels) are the same as for Mg II. The comparison between the Fe II* emission and the rescaled continuum radial SB profile reveals that the Fe II* emission is more extended than the continuum (see central panel of Fig. 7) second row. The Fe II* extended emission appears more significant in the direction of the galaxies D and E (compare the central panels of the Fig. 7 third and fourth rows). This is confirmed in the NB image where the Fe II* emission indeed extends along the minor axis of galaxy A and towards galaxy D. We discuss the origin for the extended nature of the Fe II* emission in Sect. 6.2. Contrary to Mg II and [O II] emission, the Fe II* emission does not surround the five-galaxy group but is rather centered on the most massive galaxy A (see the Mg II, [O II] and Fe II* contours and radial SB profiles comparison in the last column of Fig. 7).

Similarly, Finley et al. (2017) using MUSE found the Fe II* emission of a $z = 1.29$ galaxy to be more extended than the stellar continuum and aligned with the minor axis of the galaxy. The velocity gradient and shape of the Fe II* emission suggest the presence of a conical outflow. The faintness of the Fe II* emission of galaxy A hampers a kinematic analysis.

6. Discussion

6.1. Existence of an Mg-enriched intragroup medium

Historically, the existence of an intragroup medium enriched in magnesium has been suggested by absorption studies in order to explain the fact that while the strong Mg II absorbers (rest frame equivalent width $EW > 0.8$ Å) are generally associated with a single galaxy (within 100 kpc, e.g., Bouché et al. 2007, 2012; Ho et al. 2017; Schroetter et al. 2016, 2019; Zabl et al. 2019; Lundgren et al. 2021), sometimes weaker Mg II absorbers ($EW < 1$ Å, e.g., Muzahid et al. 2018; Dutta et al. 2020) or H I absorbers (e.g., Péroux et al. 2017; Rahmani et al. 2018) can be matched to groups of multiple galaxies ($\Delta r < 500$ kpc and $\Delta v < 500$ km s $^{-1}$), which is in agreement with clustering studies (Bouché et al. 2006; Lundgren et al. 2009; Gauthier et al. 2009).

The question of whether observed absorption features (with $EW < 1$ Å) imprinted in the spectrum of a background source are tracing gas coupled to the group or circumgalactic gas of an individual galaxy is still a matter of debate. Indeed, absorption studies reported contradictory results, some suggesting that Mg II absorbing systems were associated with the CGM of the indi-

vidual group members (e.g., Bordoloi et al. 2011; Fossati et al. 2019) and others to a widespread IGrM (e.g., Bielby et al. 2017; Péroux et al. 2017; Nielsen et al. 2018) originating from a mixture of previous tidal interactions between group members and outflowing winds. Disentangling the two is very challenging, especially when there is only one-dimensional information for a given system.

Our discovery of the first extended Mg II nebula surrounding and connecting the members of a five-galaxy group provides a panoramic view of the neutral enriched gas residing in between galaxies. This group is rather small and compact (see Sect. 3.1), meaning that it is likely an interacting system where the CGM of individual galaxies are mixed. This idea is reinforced by the fact that the detected Mg II nebula resides inside the virial radius of galaxy A estimated at ≈ 100 kpc. Moreover, the detection of a gaseous bridge in between the two subgroups [A,B,C] and [D,E] suggests that we are observing a widespread low-surface-brightness structure embracing the five galaxies, also called the “intragroup” medium. We note that there is no precise definition of the IGrM in the literature. Our analysis brings the first observational evidence of a low SB diffuse component of neutral gas residing within a galaxy group at $z > 1$ that we define as the intragroup medium.

The detection of Mg II absorption in the spectrum of the background source BKG (see Sect. 4.2) indicates that the IGrM is actually even more extended than the detected nebula. Our EW measurement of the Mg II absorption (Sect. 4.2) is in good agreement with the EW versus impact parameter relation established by quasar absorption line studies (e.g., Lundgren et al. 2021). This result highlights the complementarity of the absorption and emission methods to study the gas around galaxies.

6.2. Origin of the intragroup nebula

In this section, we discuss two points: the presence of enriched gas beyond the galaxies and the mechanisms that make it shine.

6.2.1. Reasons for the presence of enriched gas so far from the stars

Several scenarios have been suggested to explain the presence of magnesium-enriched gas in between the galaxies of a group. Magnesium atoms are α elements released in the ISM and CGM by core collapse supernovae. One can therefore naturally imagine that the magnesium gas residing outside galaxies has been ejected through strong supernovae-driven outflows. The P Cygni shape of the Mg II doublet in the integrated spectrum of galaxy A (see Fig. 2 top) is in good agreement with this scenario. Indeed, according to radiative transfer (RT) models, this line profile is a signature of resonant scattering in an optically thick outflowing medium where the redshifted emission corresponds to back-scattered photons reflected by the receding medium (Haiman et al. 2000; Dijkstra et al. 2006; Verhamme et al. 2006; Kollmeier et al. 2010). The Mg II lines of galaxies D and E seem to be redshifted compared to their systemic redshifts (Fig. 2), however, the continuum is too faint to detect significant absorption lines and, thus, too faint to detect the P Cygni profiles, namely, the presence of outflows for those galaxies. We detected $>2\sigma$ Mg II emission aligned with the minor axis of galaxy E and near galaxy D (Fig. 3). These regions could have been enriched via outflows after supernovae feedback. However, the S/N of our data is too poor in these areas to carry out a resolved kinematics study. The resolved Mg II map (Fig. 6) shows that the two regions along the minor axis of galaxy A have different

velocities with opposite signs ($\approx +120$ and -80 km s^{-1} with respect to the systemic redshift of galaxy A), providing additional evidence of the presence of an outflow. Even if we do not eventually find clear evidence of an outflow in the four other galaxies, we cannot exclude the deposition of metals in the CGM by past feedback processes explaining the observed diffuse Mg II emission detected around each of the galaxy members. In that scenario, we could imagine that the low-surface-brightness gaseous bridge in-between the galaxy subgroups [A,B,C] and [D,E] could be due to galactic wind transfer between the group members as observed in the FIRE simulation (Anglés-Alcázar et al. 2017).

The proximity of the galaxies, as well as the detection of a bridge in-between the galaxy subgroups is also compatible with another scenario where the magnesium gas has been tidally stripped out of galaxies due to past gravitational interactions among the galaxy group members. In this framework, the detected enriched bridge would correspond to a tidal feature linking the galaxies. In this scenario, Bielby et al. (2017) suggested that the IGrM could consist of multiple cool gas systems orbiting around galaxies to form the IGrM. As mentioned in Sect. 4.2, we detect several $S/N > 3$ Mg II intensity peaks within the IGrM, which, if real, could be associated with the cool gas clouds described in Bielby et al. (2017). According to hydrodynamical simulations, such cold clumps could originate from gas density perturbations after tidal interactions (Nelson et al. 2020). As mentioned above, we could also argue that the gas in the bridge could be outflowing material from galaxy A undergoing an “intergalactic transfer” to galaxy D (Anglés-Alcázar et al. 2017). However, neither the fluorescent Fe II*, nor the non-resonant [O II] emission are symmetrically distributed around galaxy A; rather they extend preferentially towards galaxy D and E (see top panels of Fig. 7), suggesting that the enriched gas detected in the bridge has been stripped of galaxies by tidal or ram pressure forces. Moreover, we do not observe any strong velocity gradients in the bridge, indicating that the bridge dynamics are consistent with the rest of the structure (left panel of Fig. 6) and reinforcing the presence of tidally disrupted material. It is, however, more difficult to disentangle the different scenarios in the close surroundings of the galaxies where – given the very small physical separation between the sources – outflows, tidal interactions, and intergalactic transfer could play a role.

Given the proximity of the five group members, we suggest that past tidal interactions are likely the dominant mechanisms explaining the presence of magnesium-enriched gas in the IGrM. Plenty of cool stripped gas is actually observed in simulations in the form of tidal features (e.g., Rosdahl & Blaizot 2012; Nelson et al. 2020). In Rosdahl & Blaizot (2012), the $\approx 10^{11.5} M_{\odot}$ DM halo (similar to our system) at $z = 3$ indeed shows orbiting satellites and tidal tails. Such H I tidally disruptive material is also observable in the local Universe, namely, in between the galaxies of the M 81 system (Yun et al. 1994).

6.2.2. Possible mechanisms behind the shine of the intragroup magnesium gas

Several mechanisms can explain the emission of Mg II photons from the intragroup gas. The possible sources, which can arise at all scales (ISM, CGM, IGrM), are as follows:

Stellar continuum at $\lambda \sim 2800 \text{ \AA}$. By summing the continuum contribution of the five galaxies in the group, we estimate the Mg II photon budget from the stellar continuum to $9.0 \pm 1.0 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ corresponding to a total luminos-

ity of $9.6 \pm 1.1 \times 10^{40} \text{ erg s}^{-1}$ without dust correction and considering the sum of the two doublet lines (see Sect. 4.1).

Photo-ionization by Mg II ionizing radiation ($\lambda < 825 \text{ \AA}$), from the stellar continuum of galaxies in the group or from the UV background (UVB) and collisional ionization of Mg^+ ions into Mg^{2+} ions, followed by recombination cascades leading to the emission of nebular Mg II photons. Using Cloudy photo-ionization models, we estimate the nebular Mg II photon budget corresponding to collisions and photo-ionization from the stellar continuum (see Appendix E).

Mg II photons produced by shocks resulting from gravitational interactions or galactic outflows from galaxies.

The Mg II doublet is a resonant line meaning that the Mg II photons are likely to be scattered within an Mg^+ gas cloud. This property is usually invoked to explain the diffuse resonant emission around galaxies, namely, Ly α halos (e.g., Kusakabe et al. 2019; Leclercq et al. 2020). However, this process erases the information on the location where the photons have been emitted. In the following, we discuss the possible distributions of sources and mechanisms that can explain the intragroup Mg II luminosity, and light distribution.

Central source with scattering. Stellar continuum: the first possibility to explain extended Mg II emission is the scattering of stellar continuum photons, namely, the re-emission of continuum photons absorbed in the Mg II transition. While Wisotzki et al. (in prep.) found that this process is likely the dominant one in their system, Zabl et al. (2021) found that it is (under the assumption of a biconical outflow) not likely enough to explain the brightness of the Mg II halo. In the group studied here, the Mg II continuum luminosity ($9.6 \pm 1.1 \times 10^{40} \text{ erg s}^{-1}$ without dust correction and integrating over the same velocity range as the Mg II NB image, see Sect. 4.1) is enough to explain the observed Mg II nebula (total luminosity of $8.0 \pm 0.5 \times 10^{40} \text{ erg s}^{-1}$). This suggests that the continuum scattering scenario is an important process at play in this system.

Nebular emission: Diffuse Mg II emission can also be produced by the scattering of nebular Mg II photons produced in the ISM of galaxies. For this mechanism, the Mg II photons are produced by collisions and recombinations associated with stellar UV radiation in star-forming regions of galaxies. The escaping photons can then scatter into the surrounding Mg^+ gas and some are redirected towards the observer. We used Cloudy models (Ferland et al. 2017) to estimate the intrinsic Mg II flux budget produced by the stars of the five galaxy group members. We found an intrinsic Mg II/[O II] flux ratio varying between ~ 10 and 30% (see Appendix E for more details). The total observed Mg II/[O II] flux ratio of the nebula (12%) being lower than the intrinsic one, or similar depending on the models, we conclude that there is enough nebular Mg II emission to explain the observed Mg II extended emission through resonant scattering of photons originally produced in H II regions. This scenario is reinforced by the spectral shapes of the galaxies (i.e., P Cygni and redshifted lines with respect to the systemic redshift), suggesting that Mg II photons scattered in an outflowing medium (see Sect. 6.2.1).

In situ emission. Collisions + photo-ionization: another scenario to explain the extended Mg II emission is that enough Mg II ionizing photons ($\lambda_0 < 825 \text{ \AA}$) escape the ISM of galaxies – perhaps facilitated by galaxy-scale outflows – so that Mg II photons are produced through photo-ionization in situ, that is, directly in the CGM or IGrM. It is however well known that the escape of ionizing photons from galaxies is very low (e.g., Malkan et al. 2003; Bridge et al. 2010; Rutkowski et al. 2016;

Alavi et al. 2020). The spatial coincidence of the Mg II and the non-resonant collisionally excited [O II] emission lines support this “non-scattering” origin. We note that the resulting Mg II photons can then scatter in the surrounding magnesium located in the IGrM, as in the first scenario. This can explain the “butterfly” shape of the ionized Mg II nebula and the fact that the Mg II emission has flatter radial SB profiles than the [O II] profiles (right panels of Fig. 7 and Sect. 5.1).

UVB. The UVB photons emitted by all the galaxies and quasars of the Universe at the redshift of the group with energy higher than 15 eV (i.e., $\lambda_0 < 825 \text{ \AA}$) can also ionize Mg^+ in the IGrM and thus produce Mg II photons. We compare the contributions of the UVB at $z = 1.3$ (Haardt & Madau 2012) and the total observed (i.e., dust attenuated) stellar emission of the group members in Fig. E.2 (see also Appendix E). By considering a very low escape fraction of the Mg^+ ionizing flux (0.7%) we find that the UVB dominates the Mg^+ ionizing budget in the IGrM (i.e., $r > 10 \text{ kpc}$ from the stars). In the vicinity of the stars ($r \approx 10 \text{ kpc}$), the contribution of the stellar emission becomes dominant. Therefore the UVB appears to be a non-negligible source of ionization at large distances from the stars, namely, in the CGM/IGrM, if we assume no escape of Mg II ionizing photons from the galaxies. When considering that a non-zero fraction of the ionizing flux escapes (e.g., 14% in Fig. E.2), the contribution of the UVB to the powering of the Mg II nebula becomes negligible. While, as mentioned above, the ionizing escape fraction in galaxies is usually found to be very low, galaxy C shows some hints for a clear line of sight (see Appendix D), thus indicating the possibility of ionizing flux leakage (e.g., Chisholm et al. 2020). A precise measurement of the ionizing flux escape fraction from the group members is needed to go deeper in the analysis but is beyond the scope of this paper.

Satellite. In order to explain the Mg II emission at the position of the bridge, we estimated upper limits of $\approx 10^7 M_\odot$ and $\approx 10^{-3} M_\odot \text{ yr}^{-1}$ for the stellar mass and SFR, respectively (Madau et al. 1998; Whitaker et al. 2014), of an undetected galaxy at $z \approx 1.3$ in the extremely deep F775W HST image (limiting magnitude of 29.5 mag corresponding to a UV luminosity of $3 \times 10^{37} \text{ erg s}^{-1}$, Rafelski et al. 2015). In other words, if there are unseen galaxies powering the Mg II emission at the position of the bridge, they should correspond to ultra-low luminosity Mg II emitters at $z \approx 1$. Bacon et al. (2021) has recently proposed the existence of a population of very faint Ly α emitters which could contribute to the extended Ly α emission tracing overdense structures. We think that such a hypothesis is unlikely to hold for the Mg II group studied here since it would require an Mg II EW ($\gtrsim 20 \text{ \AA}$) above the highest Mg II EW (i.e., 19 \AA) measured in the MUSE deep fields (Feltre et al. 2018).

Shocks. In the compact and complex configuration of our group, a large fraction of Mg II photons are likely to be produced due to the shocks resulting from gravitational interactions among the group members or outflows from galaxies (Heckman et al. 1990; Monreal-Ibero et al. 2006).

In order to distinguish between these sources of ionization, we would need additional lines such as [O III] $\lambda 5007$ or H β $\lambda 4861$ in order to use line diagnostics for a comparison with photo-ionization and shock models (e.g., Alarie & Morisset 2019). Moreover, the presence of skylines in both the Mg II and [O II] doublet hampers any further investigations.

All in all, our experiments suggests that both the UV stellar emission (at $\lambda < 825 \text{ \AA}$ and $\lambda \approx 2800 \text{ \AA}$) and the UVB

contribute to make the Mg II intragroup medium shine. The photo-ionization of Mg^+ by the UVB appears like the dominant scenario only if the ionizing escape fractions from the group members are close to zero, which is likely for galaxy A but questionable for galaxy C. The spatial coincidence of the Mg II and the non-resonant and collisionally excited [O II] nebulae suggests that the in situ emission by collisions and photo-ionization processes should be favored. The reality is likely that a mixture of different mechanisms arise at all scales. Our analysis suggests that such nebulae can be commonly found around groups of star-forming galaxies, providing that data are robust enough. This will be statistically investigated in an upcoming paper (Leclercq et al., in prep.).

6.3. Comparison with the literature

We now compare the properties of this new intragroup Mg II nebula with (i) the known ionized gas structures found around galaxy groups and (ii) recently reported Mg II extended emission around individual galaxies.

6.3.1. Ionized nebula in galaxy groups

Today, only a few ionized nebulae around galaxy groups (DM halo mass $\lesssim 10^{13} M_\odot$) are known and most of them have been detected thanks to the arrival of IFU instruments. The first detection of a large ionized structure in a galaxy group was reported by Epinat et al. (2018). Using MUSE, they found a 150 kpc large [O II] nebula at $z \approx 0.7$ embracing a dozen of galaxies with maximum stellar masses of $\approx 10^{10.9} M_\odot$ (the DM halo mass of the group is $6.5 \times 10^{13} M_\odot$, see Abril-Melgarejo et al. 2021). By investigating the kinematics and ionization properties, the authors concluded that gas was stripped out of the galaxies by tidal forces after galaxy interactions and AGN outflow. A year later, Chen et al. (2019) mapped a 100 kpc large H α structure spanning over a 14-galaxy group (Kollmeier et al. 2010) at $z \approx 0.3$. Their analysis of the emission morphology and kinematics indicates that most of those nebulae are powered by shocks and turbulent gas motions associated with gas stripping after gravitational interactions of the group members. Two other recent studies reported the detection of ionized nebulae emitting in [O III], [O II] and H β around six and three galaxy groups hosting a quasar (Johnson et al. 2018; Helton et al. 2021, respectively). All those studies favor a scenario where the enriched gas in the IGrM originates from tidally stripped gas after galaxy interactions.

The galaxy group studied here (i.e., embedded in the Mg II nebula) is at higher redshift and consists of fewer galaxies which are also less massive than the group members embedded in the other ionized nebulae. It also has a more compact configuration ($< 50 \text{ kpc}$) in projection and in velocity space ($< 120 \text{ km s}^{-1}$). Although the galaxy group embedded in the Mg II nebula shows different properties compared to the previously published systems, we all favor a scenario where the enriched gas in the IGrM originates from tidally stripped gas after galaxy interactions.

Finally, we can also mention the discovery by Rupke et al. (2019) of an [O II] nebula around a massive galaxy ($10^{11.1} M_\odot$) at $z \approx 0.46$. This nebula has an hourglass shape which is similar to our nebula. According to the authors, this spatial distribution indicates an ionized bipolar outflow. However our discovery shows that, if we have deep enough data to detect the continuum of companion galaxies, this kind of morphology is also compatible with a gas stripping scenario. Further support

for this possibility comes from the two tidal tails visible in the HST image of the published [O II] nebula.

6.3.2. Mg II halos around individual galaxies

Thanks to the high sensitivity of MUSE and to a long exposure time, we discovered the first Mg II nebula around a galaxy group at any redshift. Extended Mg II emission around three individual galaxies at $z \simeq 0.7$ has very recently been reported in Burchett et al. (2021), Zabl et al. (2021) and Wisotzki et al. (in prep.) using the Keck/KCWI and MUSE instruments. These galaxies are more massive ($10^{9.75} < M_* [M_\odot] < 10^{10.05}$) than galaxy A ($10^{9.35} M_\odot$) and are above the main sequence of star forming galaxies (see Zabl et al. 2021 for a compilation), which is not the case for galaxy A (Sect. 3.2), although the uncertainties on the SFR value are large (Table 1). Moreover, the three studies of Mg II halos are consistent with the presence of outflows, which is also very likely in galaxy A.

At $z \simeq 0.7$, the [O II] emission is covered by MUSE. Both Zabl et al. (2021) and Wisotzki et al. (in prep.) found extended [O II] emission but with a steeper spatial profile compared to Mg II. The Mg II emission measured in these recent studies extends out from the center of the galaxies to a radius of 20 to 40 kpc. In particular, the Mg II halo analyzed by Zabl et al. (2021) spans a total area of 1000 kpc^2 above the 2σ significance level which is comparable to our nebula.

We can also imagine a scenario where our galaxy group would be a progenitor of the observed galaxies with Mg II halos at $z \simeq 0.7$. In this scenario we would be seeing the pre-merging stage leading to a more massive galaxy. Actually, both the galaxies studied in Wisotzki et al. (in prep.) and Zabl et al. (2021) are likely in a late stage merger as they both show significant asymmetric substructures in their central regions (see Appendix A.2 in Zabl et al. 2021). Using a cosmological simulation, Ventou et al. (2019) investigated the probability that interacting galaxies will merge in the future. According to this study, the three galaxies A, B, and C, and the two galaxies D and E, separated in projection by $\Delta r < 25 \text{ kpc}$ and $\Delta v < 100 \text{ km s}^{-1}$, have 70% chance to merge (see their Fig. 2b). The systems [A,B,C] and [D,E], separated by $\Delta r < 50 \text{ kpc}$ and $\Delta v < 300 \text{ km s}^{-1}$, have at least 30% change of merging by $z = 0$ (their Sect. 3.2.3).

Interestingly, both objects from Zabl et al. (2021) and Wisotzki et al. (in prep.) have a less massive companion located at $< 100 \text{ km s}^{-1}$ in velocity space and at $\lesssim 50 \text{ kpc}$ in projection. The authors do not exclude the presence of tidal stripping effects in those systems, reinforcing the idea that interactions might play a crucial role in the redistribution of metals within the circumgalactic and intragroup media. Finally, it appears that the signatures of outflows and tidal interactions are very difficult to disentangle. Moreover, according to simulations, we actually expect the two to co-exist, which seems to be the case in our system.

7. Summary and conclusions

Thanks to the extraordinarily deep MXDF data, we report the first detection of an intragroup medium shining in Mg II. The nebula surrounds and connects five neighboring galaxies at $z \simeq 1.31$, separated by less than 50 kpc in projection and $\approx 120 \text{ km s}^{-1}$ in velocity space (Figs. 1 and 2). With a DM halo mass of $\approx 10^{11.7} M_\odot$ and the highest galaxy mass of $\approx 10^{9.3} M_\odot$ (Table 1), this is a low-mass, compact, and high redshift system, as compared to previous groups that have been found to be surrounded by ionized nebulae. The detection of Mg II absorption features

in the spectrum of a background galaxy located at an impact parameter of 19 kpc from the group indicates that the intragroup medium studied here is even larger than the nebula seen in emission. Our ≈ 60 -h deep MUSE data (Fig. 3b) allowed us to spatially and spectrally map the extended gaseous nebula of the system. Our observations provide a novel and panoramic view on the IGrM, which has historically only been probed in 1D by studies employing absorption line techniques. The present study allowed us to deepen our understanding of the existence and origin of enriched gas outside of star-forming galaxies residing in groups. Our results can be summarized as follows:

1. We detected a 1000 kpc^2 Mg II emitting nebula with total Mg II flux of $7.48 \pm 0.50 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ within a 2σ level isophote, corresponding to a total luminosity of $8.0 \pm 0.5 \times 10^{40} \text{ erg s}^{-1}$. Our optimized NB image construction procedure (Sect. 4.1) allowed us to capture most of the Mg II flux while limiting the noise (Appendix C). Once slightly smoothed to increase the S/N while preserving a good spatial resolution (Appendix B), our panoramic Mg II map (Fig. 3a) reveals that the Mg II nebula has an elongated shape surrounding and connecting the five group members with a maximal projected extent of $\approx 70 \text{ kpc}$. A low SB ($\approx 2 \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$) bridge connects the galaxies [A,B,C] and [D,E] (Figs. 3a and 5). This bridge does not appear to be a projection effect because of its tentative detection ($S/N \lesssim 2$) in a single 1.25 \AA (58 km s^{-1}) cube slice (Fig. 4). While the highest intensity peak coincides with galaxy C (Appendix D), others are detected with $S/N > 3$ within the nebula. We see evident absorption at the position of galaxy A and the background galaxy BKG (Fig. 3c). BKG is located at an impact parameter of 19 kpc from galaxy D, outside of the 2σ detection edge of the nebula. The detection of absorption lines in its spectrum at the redshift of the group indicates that there is gas enriched in magnesium even beyond the detected nebula. This underlines the complementarity between emission and absorption studies.
2. Taking advantage of our IFU data, we studied the kinematics of the Mg II gas surrounding the group using the 2D weighted Voronoi binning to increase the S/N in the lower SB regions (Sect. 4.4). Our resolved Mg II velocity map (Fig. 6) suggests an overall rotation of the structure along the major axis of galaxy A consistent with an extension of the ISM rotation. A comparison with the [O II] kinematics shows that the kinematics of the Mg II nebula is overall redshifted with respect to the [O II] kinematics. The highest velocity variations are observed along the minor axis of galaxy A where the bulk of the gas reaches velocities of $+120$ and -80 km s^{-1} with respect to the systemic redshift, indicative of outflowing gas.
3. At the redshift of the group, MUSE covers the Fe II* and [O II] lines (Fig. 2). We measured extended [O II] emission as well as a $S/N \simeq 2$ [O II] bridge at the same location as the Mg II bridge. The Fe II* emission appears more extended than the continuum around galaxy A and extends preferentially towards the neighboring galaxy D, similarly to the Mg II and [O II] emission (Fig. 7).
4. This first panoramic and 3D view of an IGrM detected in Mg II revealed a diffuse low SB metal-enriched gaseous component surrounding and connecting the galaxy group members (Sect. 6.1). The evidence supporting outflows, at least from the most massive galaxy A as well as the proximity of the five group members, suggests that the IGrM has been enriched through both stellar outflows and tidal interactions between the group members (Sect. 6.2.1).

5. We listed the different mechanisms that can power the Mg II nebula in Sect. 6.2.2: stellar UV continuum at Mg II, nebular emission, collisional excitation, and discussed the possible distributions of sources that can explain the intragroup Mg II luminosity and light distribution. We found that the UV stellar continua, both at $\lambda < 825 \text{ \AA}$ and $\lambda \approx 2800 \text{ \AA}$, are possible sources powering the observed Mg II nebula (Appendix E.1). We compare the stellar emission to the UVB contribution (Appendix E.2) and found that the UVB is the dominant process only if the total ionizing escape fraction is close to zero which is likely the case for galaxy A, but questionable for galaxy C. The spatial coincidence of the Mg II and [O II] nebulae favors the in situ emission by collisional and photoionization processes. We conclude that we need additional lines to distinguish between the different sources of ionization/excitation.
6. When comparing our results with the literature (Sect. 6.3) and specifically to previously detected ionized nebulae in galaxy groups (DM halo mass $\lesssim 10^{13} M_{\odot}$), we found that our nebula is less extended and surrounds fewer and lower mass galaxies. Moreover, a comparison between our Mg II nebula and the three Mg II halos recently mapped by IFUs and detected around individual galaxies at lower redshift ($z \approx 0.7$), reveals similar spatial extents. For those three cases, the authors reported the presence of outflows. Interestingly, two of the three galaxies from the literature have a close companion galaxy implying that tidal stripping cannot be excluded. Finally, we suggest the possibility that our $z = 1.3$ system constitutes a pre-merging stage of the observed $z \approx 0.7$ galaxies surrounded by an Mg II halo. This is reinforced by the fact that all three $z \approx 0.7$ galaxies show hints for a late-stage merger.

This discovery paper sheds light on the existence and origin of the IGrM in one low-mass system of five galaxies. More observations are needed to generalize the presence of a low SB diffuse enriched gaseous component within low-mass galaxy groups. While facilitated by the advent of IFU instruments such as MUSE, such observations are, however, very expensive in telescope time as they require several tens of hours in exposure time. More cases of extended Mg II emission have been detected in the MXDF data and are the topic of an upcoming paper (Leclercq et al., in prep.).

Using the Mg II emission to map the circumgalactic and intragroup media is particularly interesting because it traces the same cool and neutral gas phase as the Ly α emission. By comparing the Ly α and Mg II halo properties we will be able to probe the spatial distribution of the enriched versus pristine gas around galaxies and better characterize the gas exchanges between the galaxies and their environments, which is crucial for improving our understanding of galaxy evolution.

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Appendix A: Nebula appearance at different depths

The Mg II nebula was first discovered in the 31-hour deep MUSE udf-10 data cube (second panel of Fig. A.1). Part of this field has been re-observed at greater depth (≈ 60 h) as part of the MXDF program (Sect. 2). Here, we compare the appearance of the detected Mg II IGrM in these two fields (left and middle panels) and in the combined datacube (udf-10 and MXDF) reaching a ≈ 90 h depth (right panel).

In the combined data, the shape of the Mg II nebula is reassuringly preserved, the S/N increased but the PSF is larger. For comparison we also show the appearance of the nebula in the ten-hour deep mosaic field, highlighting the fact that deep observations are crucial to study the CGM and IGrM of low-mass galaxies.

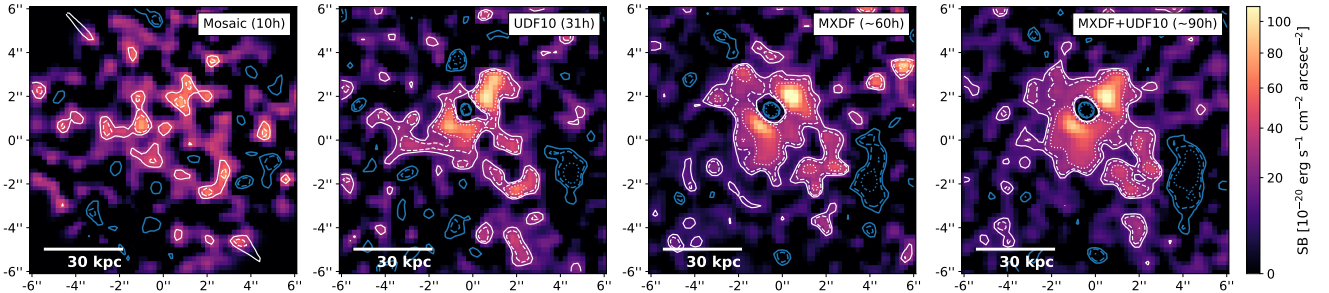


Fig. A.1. Mg II NB image of the nebula as constructed in Sect. 4.1 using the mosaic (10 h), udf-10 (31 h), MXDF (≈ 60 h) and a combination of the udf-10 and MXDF (≈ 90 h) MUSE data cubes. The legend is the same as in Fig. 3a.

Appendix B: Smoothing effects on the detected nebula shape

Although our MUSE data are very deep (≈ 60 h, see Fig. 3b), the diffuse Mg II emission of the group revealed by such deep observations is still very faint. In order to increase the S/N, while keeping a spatial resolution that is as good as possible, we smoothed our data with a Gaussian kernel whose FWHM (3 MUSE pixels or $0''.6$) is slightly higher than the MUSE PSF FWHM ($0''.52$).

Figure B.1 shows the dependence of the Mg II nebula S/N map on the smoothing kernel size. We choose to apply a smoothing as light as possible to keep an as good as possible spatial resolution.

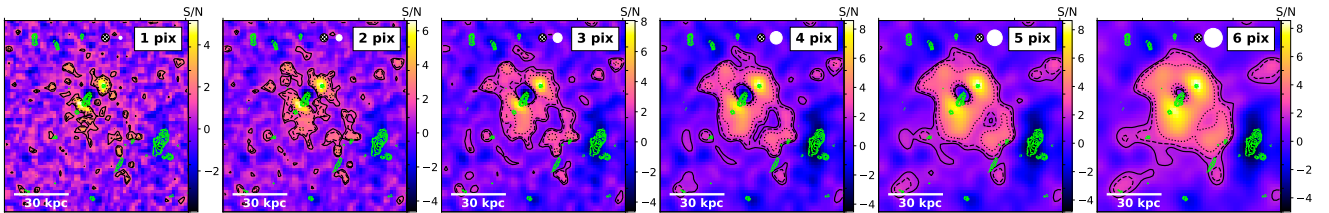


Fig. B.1. Mg II S/N map as constructed in Sect. 4.1 and smoothed using a 1, 2, 3, 4, 5, and 6 spaxel FWHM Gaussian from left to right, respectively. The FWHM values are indicated and shown as white circle (top right) in each panel. For comparison, the FWHM of the MUSE PSF is shown with a hatched white circle. The solid, dashed and dotted black contours correspond to a S/N of 1.5, 2 and 3, respectively. The adopted smoothing kernel has a FWHM of $0''.6$ (i.e., 3 spaxels or ≈ 5 kpc at $z = 1.31$) and allows us to improve the S/N while keeping a good spatial resolution.

Appendix C: Spectral bandwidth effect on the Mg II nebula detection

As described in Sect. 4.1, the final Mg II NB image (Fig. 3a) has been optimized to encompass all the Mg II flux in the $\lambda 2796$ and $\lambda 2803$ lines while limiting the noise. The selected wavelength bands are shown by the shaded blue areas in the first panel of Fig. C.1 and indeed cover both the Mg II lines extracted in an area above the 2σ significance level (first panel). In order to emphasize the importance of optimizing the NB image and to check for any missing flux, we created a broader Mg II NB image (middle panel) encompassing the whole Mg II spectral range ($2792\text{--}2806$ Å in rest frame, purple shaded area in left panel). This broader-band (BB) Mg II image is noisier and the overall shape (green contour) of the nebula is lost in the noise. Looking at the residual image (i.e., the difference between the NB and BB Mg II images, right panel), we can see that, while some absorbing flux at the position of galaxy A is missed, all the emitting flux is captured by our NB construction procedure.

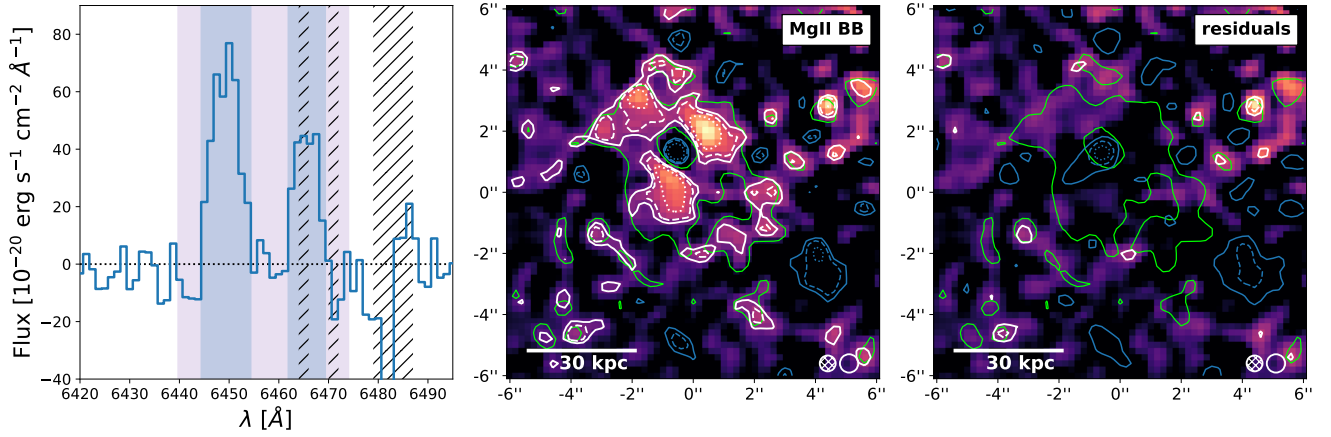


Fig. C.1. Spectral bandwidth effect on the Mg II nebula detection. *Left:* Mg II spectrum of the nebula integrated within the 2σ significance level contour (see Fig. 3a). The vertical blue shaded areas indicate the spectral widths used to construct the Mg II NB image shown in Fig. 3. The shaded purple area shows the wavelength band used to build the broader NB image shown in the middle panel. The hatched areas indicate the position of sky lines. *Middle:* Broader Mg II NB image extended by $\pm 500 \text{ km s}^{-1}$ on the $\lambda 2796$ and $\lambda 2803$ line outskirts (purple shaded area in the left panel) encompassing the whole Mg II spectral range (2792–2806 Å in rest frame). *Right:* Difference between the Mg II NB and broader band (blue and purple shaded areas in the left panel) images. The contours are the same as in Fig. 3 except that the green contour here shows the 1.5σ contour of the S/N-optimized Mg II NB image.

Appendix D: Optically thin emission in the galaxy C line of sight

In Sect. 4.3 we rescaled the continuum radial SB profile to the Mg II radial SB value matching the position of galaxy C. This choice is motivated on the following basis:

- the fact that the lines are well fitted by Gaussian functions (see Fig. D.1) indicates that the Mg II light emitted from galaxy C does not undergo strong radiative transfer (RT) effects or absorption.
- the best-fit Mg II line flux ratio F_{2796}/F_{2803} of 2.11 ± 0.69 for galaxy C corresponds to optically thin emission and thus is in good agreement with a lack of RT effects and absorption in this line of sight.
- as a resonant line, we can expect the so-called infilling effect to be significant for Mg II in this system. In other words, the absorbed photons can be re-emitted resonantly and escape the galaxy. As such, scattered Mg II photons can contribute to the spectrum of the galaxy by “filling” the Mg II absorption troughs. Mauerhofer et al. (2021) studied this effect for a simulated galaxy and found that it is more important for directions showing strong fluorescence lines. No fluorescent Fe II* lines are significantly detected in the galaxy C spectrum (Fig. 2, third line). The lack of infilling would reinforce the lack of scattering in the line of sight of galaxy C because the photons creating the infilling effect must have scattered at least once according to Mauerhofer et al. (2021).

To sum up, the lack of hints for absorption features and RT effects in the spectrum of galaxy C indicate that the Mg II emission is optically thin at the location of C. Those motivations justify our continuum rescaling procedure used to highlight the extended nature of the Mg II emission on radial SB profiles (Sect. 4.3).

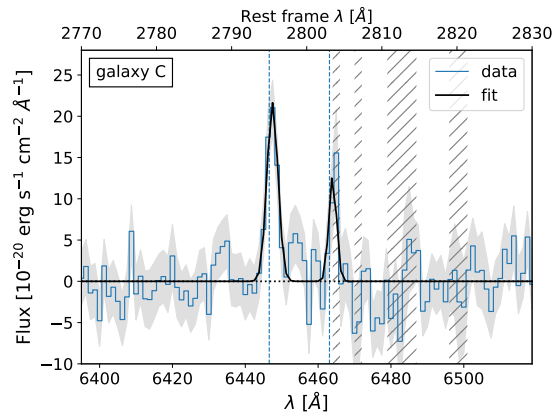


Fig. D.1. Mg II doublet line fitting of galaxy C. The best fit line (black) is shown along with the continuum subtracted ODHIN spectrum (blue). The uncertainties are shown in gray and the position of sky lines with gray hatches.

Appendix E: Estimating the intrinsic Mg II photons budget using Cloudy modeling

E.1. Intrinsic nebular emission in H II regions

In order to predict an Mg II intrinsic stellar photon budget, we produced a grid of Cloudy models with a closed spherical geometry (Ferland et al. 2017). The input SEDs were created by combining the MAGPHYS intrinsic spectra for the group galaxies (obtained by fitting the Rafelski et al. 2015 HST photometry, see Sect. 3.2) with constant star-forming model SEDs obtained from the binary population synthesis code BPASSv2.2.1 (Eldridge et al. 2017) as shown in Fig. E.2. The BPASS SED was used to extend the MAGPHYS spectra to ionizing wavelengths. The age of the stellar population models was set to $10^{8.75}$ years and the metallicity (Z) to two times solar, in accordance to the MAGPHYS best fit parameters. The stopping criteria of the calculations were set by the Hydrogen column density, with values between $\log_{10}(N_{\text{Hstop}}) = 17$ and 20 cm^{-2} . This parameter cuts the calculation at some column density values inside the Stömgren Sphere. Finally, the luminosity of the models was set by the amount of hydrogen ionizing photons ($\log(q(H)) \approx 51 \text{ photons s}^{-1}$), constraining the ionization parameter to $\log(U) \approx -2.6$.

We adopt two different prescriptions for the element abundances in the cloud: (1) solar abundance ratios from Grevesse et al. (2010) where $\frac{Mg}{H} \sim 10^{-4.5}$ and, (2) ISM abundances based on the mean warm and cold phases of the ISM ($\frac{Mg}{H} \sim 10^{-4.9}$). It is noteworthy to mention that (1) does not include grains, while (2) has a combination of silicates and graphites defined by Mathis et al. (1977). The results of these models can be seen in Figure E.1, in which we show how the intrinsic Mg II to [O II] ratio changes as a function of the stopping column density. We find that the models provide a higher intrinsic Mg II to [O II] ratio than the observed one (where dust attenuation and scattering might explain the discrepancy) when considering solar abundance ratios and ISM abundances with a stopping column density $> 10^{18} \text{ cm}^{-2}$. This suggests that the Mg II emission of the nebula can be explained by stellar processes alone through resonant scattering of photons originally produced in H II regions.

E.2. Comparison between the stellar versus UVB contribution to the Mg II nebulae

In order to assess the origin of the Mg^+ ionizing photons ($\lambda < 825 \text{ \AA}$), leading to the emission of Mg II photons, we compare the total SED of the group members (sum of the five SEDs, see Sect. E.1) with the UVB SED from Haardt & Madau (2012) at the redshift of the group (i.e., $z = 1.3$). Those SEDs are in rest frame and show the flux as it would be observed at a distance varying between 10 and 35 kpc of the stars. This range in radius allowed us to evaluate and compare the contributions of the stars and UVB at the edge of the detected nebula as well as closer to the galaxies. We show two models with arbitrary values of the escape fractions 0.7% and 14% ($\log_{10}(N_{\text{Hstop}}) = 18$ and 18.5 , respectively), calculated as the transmitted to intrinsic ratio of the flux integrated $< 912 \text{ \AA}$.

In Figure E.2 we show that the contribution of the UVB to the Mg^+ ionizing flux ($\lambda < 825 \text{ \AA}$ or $E > 15.04 \text{ eV}$) is dominant in the nebula ($10 < r [\text{kpc}] < 35$) only when considering an ionizing escape fraction close to zero. When getting closer to the stars ($r < 10 \text{ kpc}$), the contribution of the UV stellar emission becomes dominant. At higher escape fractions (e.g., 14% as shown on the figure), the contribution of the UVB quickly becomes negligible, even at the outskirts of the nebula.

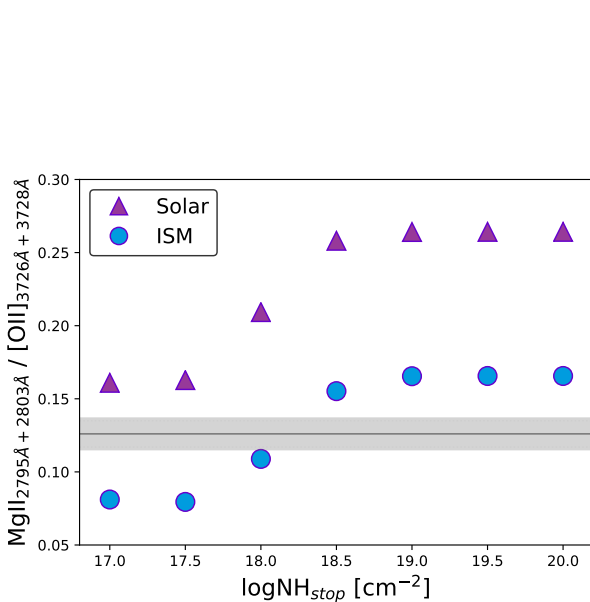


Fig. E.1. Predicted intrinsic nebular Mg II to [O II] ratio using Cloudy models constrained by the best fit parameters of the observations (see SEDs in the right panel). The symbols represent the chemical abundances used, triangles for solar and circles for ISM. The horizontal gray line and shaded area show the observed Mg II to [O II] ratio from the group and the 1σ uncertainties, respectively.

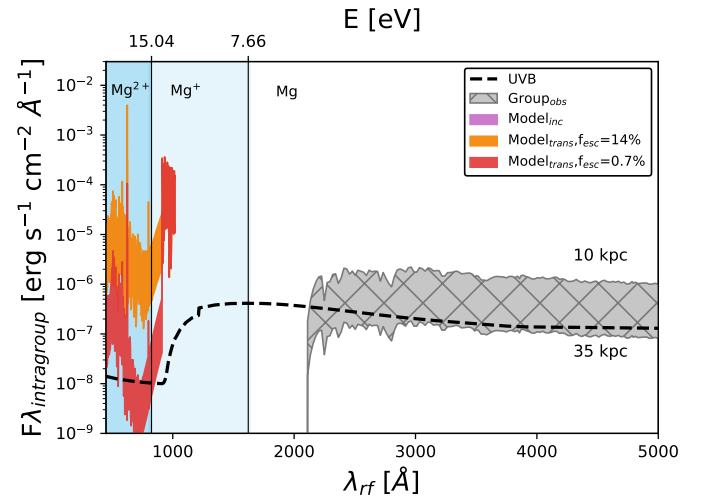


Fig. E.2. Comparison of the UVB vs stellar photons budget, ionizing Mg^+ in Mg^{2+} . In purple we show the intrinsic BPASS model for $Z = 2Z_{\odot}$ and age of $10^{8.75}$ years for a constant SF stellar population. The transmitted BPASS model considering 0.7% and 14% ionizing escape fraction is shown in red and orange, respectively. The gray spectrum shows the attenuated spectrum of the group (sum of five SEDs, Sect. E.1). The dispersion of all spectra depends on the radius at which we estimate the flux, the lower and upper boundaries of the areas are for 35 and 10 kpc, respectively. The black dashed line shows the UVB from Haardt & Madau (2012) at $z = 1.3$.