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Radial Velocities and Kinematic Ages of Nearby T Dwarfs from Keck/NIRSPEC High-Resolution Spectroscopy



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Introduction

Precise measurements of radial (RV) and rotational ($v_{\text{sin}i}$) velocities of stars are essential for studying stellar kinematics (space velocities and dispersions), binary orbits (mass measurements and formation), and rotational dynamics (angular momentum evolution). The high-resolution spectroscopic observations necessary to make these measurements are challenging for the intrinsically faint and low-temperature ultracool dwarfs, stellar and sub-stellar objects with masses below $0.1 M_{\odot}$. Previous local UCD kinematic studies indicated conflicting L dwarf kinematic ages^{5,6,10,15}, with little constraints on the T dwarf kinematic due to smaller sample size ($N = 9$)¹⁵. We present a radial and rotational velocity survey of 37 nearby ($d \leq 20$ pc) T dwarfs based on forward-modeling analysis of nearly 20 years of high-resolution spectra obtained with Keck/NIRSPEC.

Modeling the Spectral Data

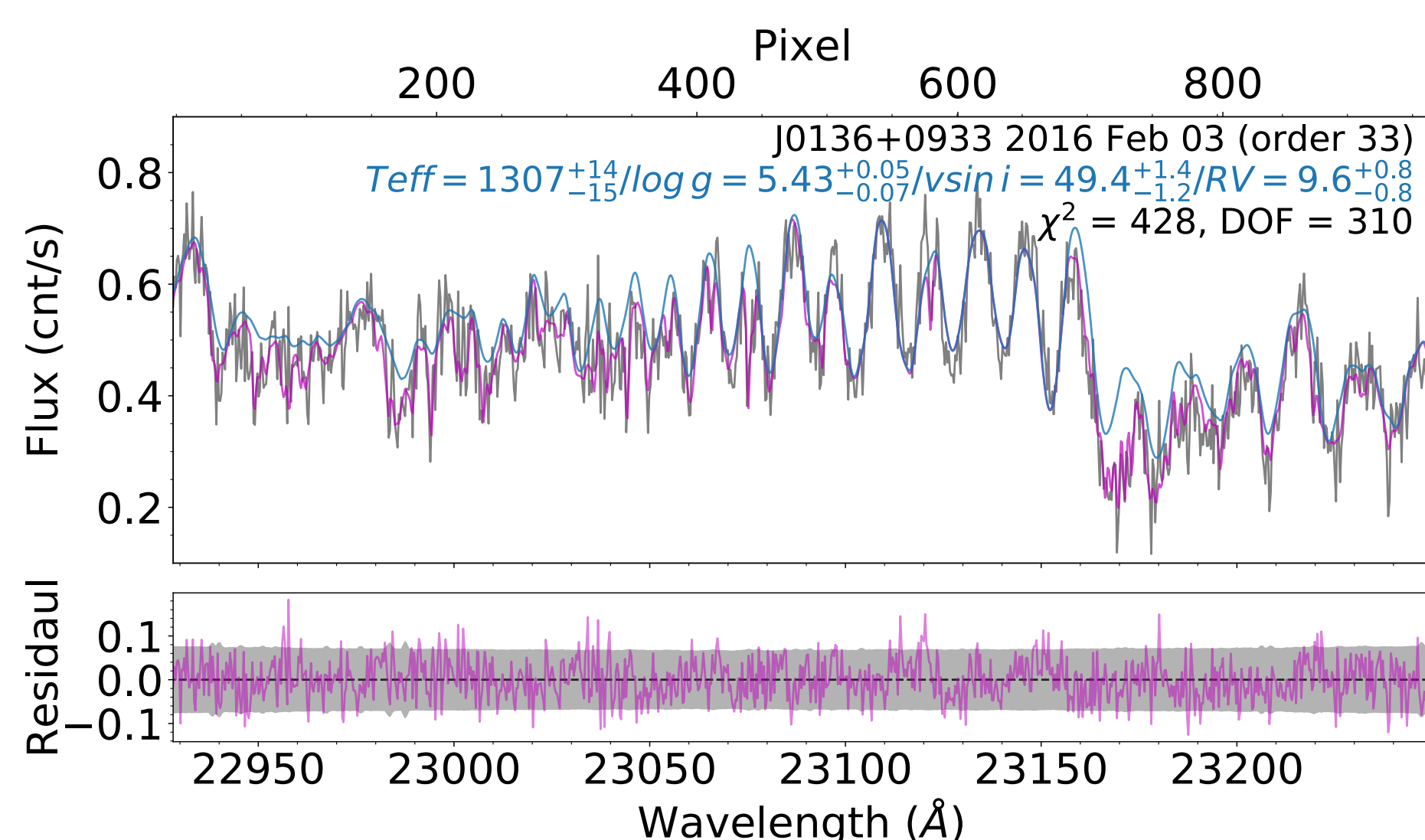


Figure 1. BT-Settl model fit of the order 33 spectrum of the T2.5 J0136+0933, observed on 2016 February 3 (UT). *Upper*: the grey line is the observed spectra; the magenta and blue lines are the stellar model with and without telluric absorption, respectively. *Lower*: difference of the data - model (magenta) with $\pm 1\sigma$ data uncertainty shaded in grey.

We built upon the forward modeling method^{5,7} and employed MCMC to extract the effective temperature, gravity, $v_{\text{sin}i}$, RV, telluric airmass and water vapor parameters (Figure 1). Based on our analysis of 37 T dwarfs, our RV and $v_{\text{sin}i}$ measurements are generally consistent with previous results, and we achieve **median precisions of 0.5 and 0.9 km/s**, respectively. RV precision is better for late-M/L dwarfs as they typically have higher S/N and smaller $v_{\text{sin}i}$ values.

UCD Kinematics & Ages

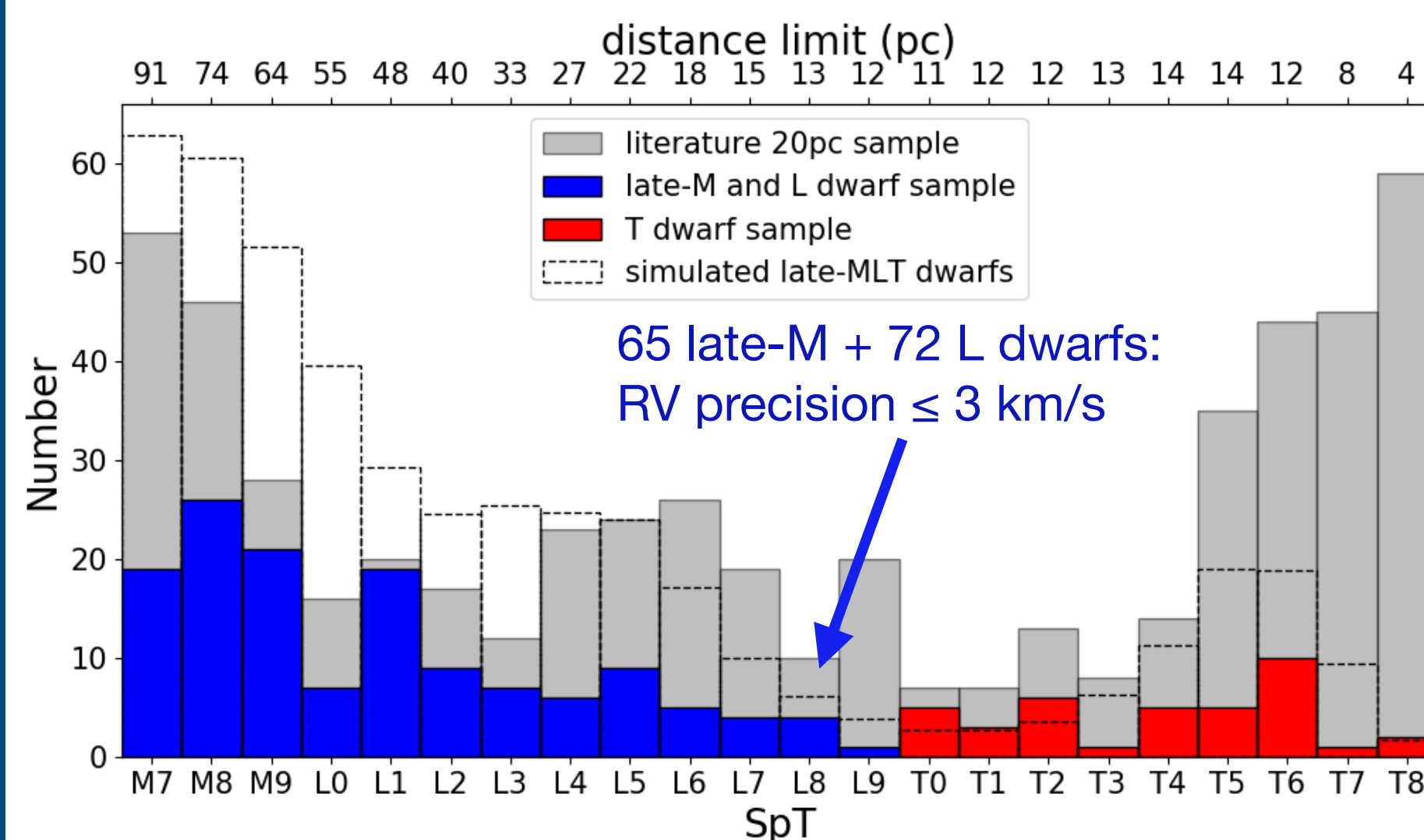


Figure 2. Spectral type distribution of our 20 pc late-M and L dwarf kinematic sample with RV uncertainty of $\leq 3 \text{ km s}^{-1}$ (blue histogram), and our NIRSPEC T dwarf sample (red histogram).

Kinematic ages were computed² from empirical age-velocity dispersions for a local sample of 173 UCDs with RV uncertainty $\leq 3 \text{ km s}^{-1}$ (Figure 2). The estimated ages for the late-M and T dwarfs are comparable ($4.1 \pm 0.3 \text{ Gyr}$ and $3.5 \pm 0.3 \text{ Gyr}$), while the L dwarf population appears too old ($5.7 \pm 0.3 \text{ Gyr}$). However, the local L dwarf sample has a higher fraction of thick disk sources, and **removing them brings the L dwarf age into alignment ($4.1 \pm 0.3 \text{ Gyr}$), resolving a decade-old mystery**. A population simulation assuming an exponential star formation rate² from 0.1 to 9 Gyr and a mass range from 0.01 to $0.15 M_{\odot}$ predicts ages consistent with the measurements (Figure 3).

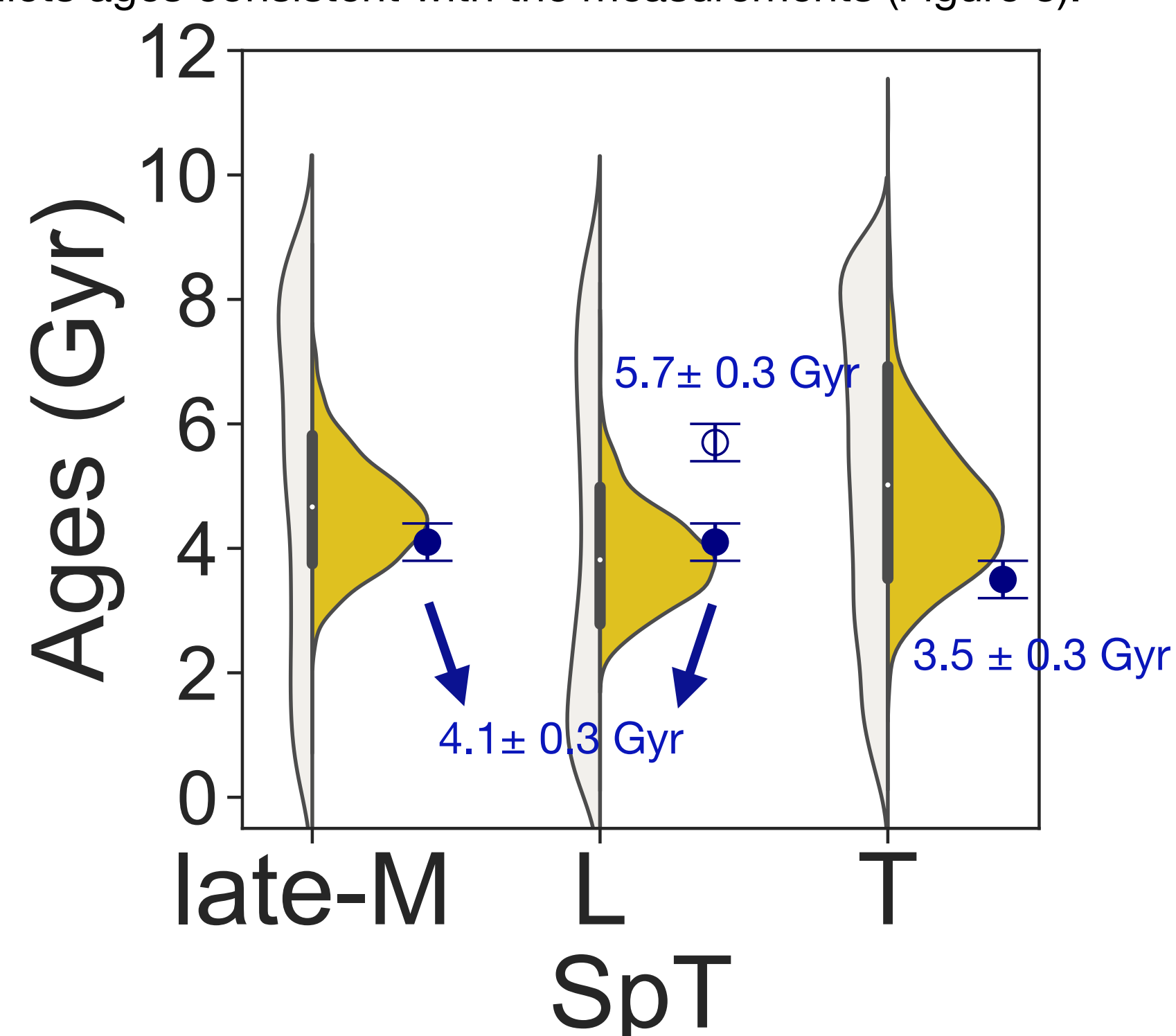


Figure 3. Simulated age distributions (white/yellow violin plots for individual/inferred ages, respectively) and measured kinematic ages for late-M, L, T dwarfs. The L dwarf age with thick disk sources included is indicated by the open circle.

Kinematic H Burning Limit

A kinematic dispersion break is found around the L4–L6 subtypes, which likely reflects the terminus of the stellar Main Sequence (Figure 4), consistent with dynamical mass determinations⁹ but later than radius measurements⁸.

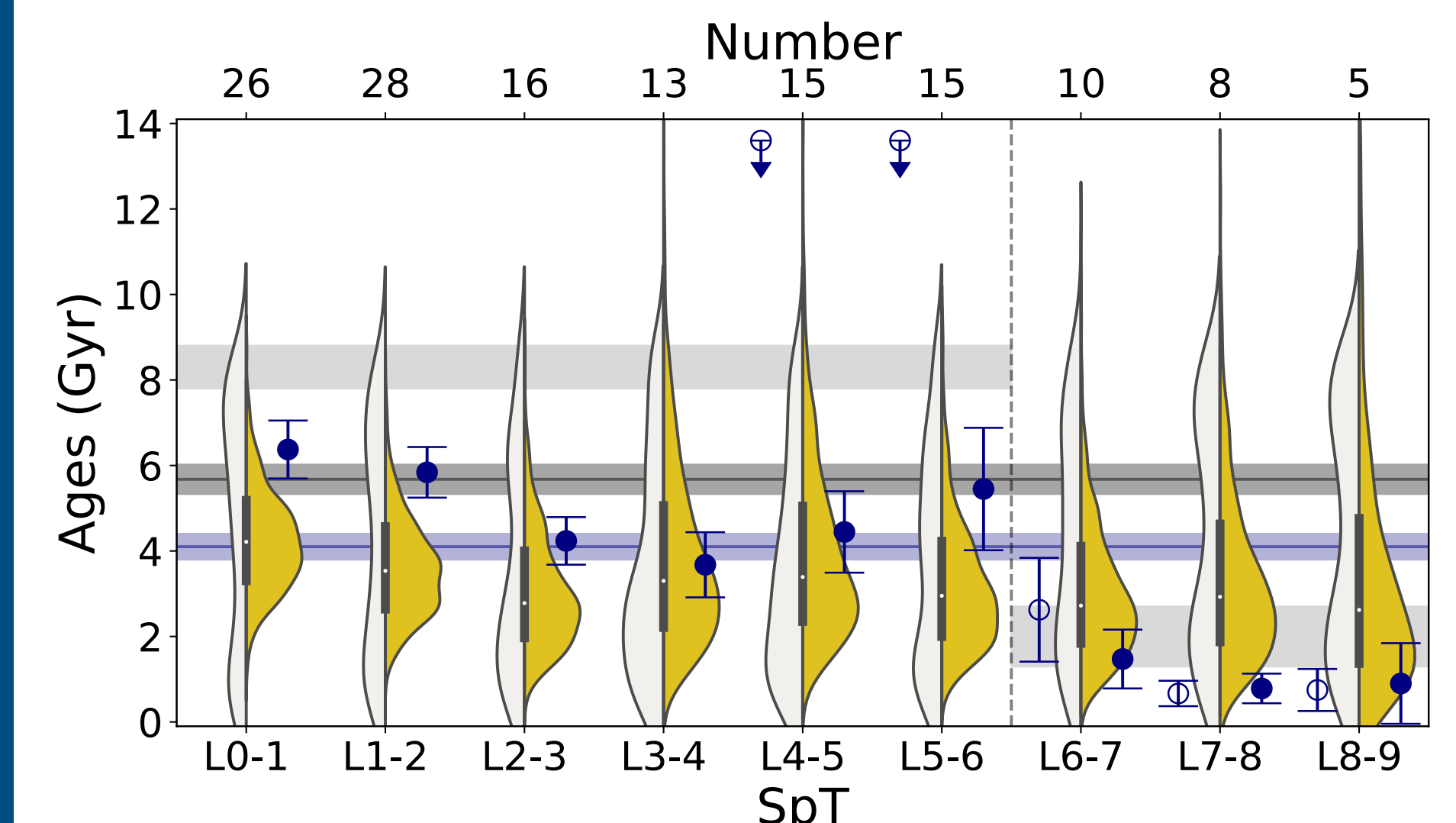


Figure 4. Same as Figure 3 for L subtype distributions.

UCD $v_{\text{sin}i}$ Distribution

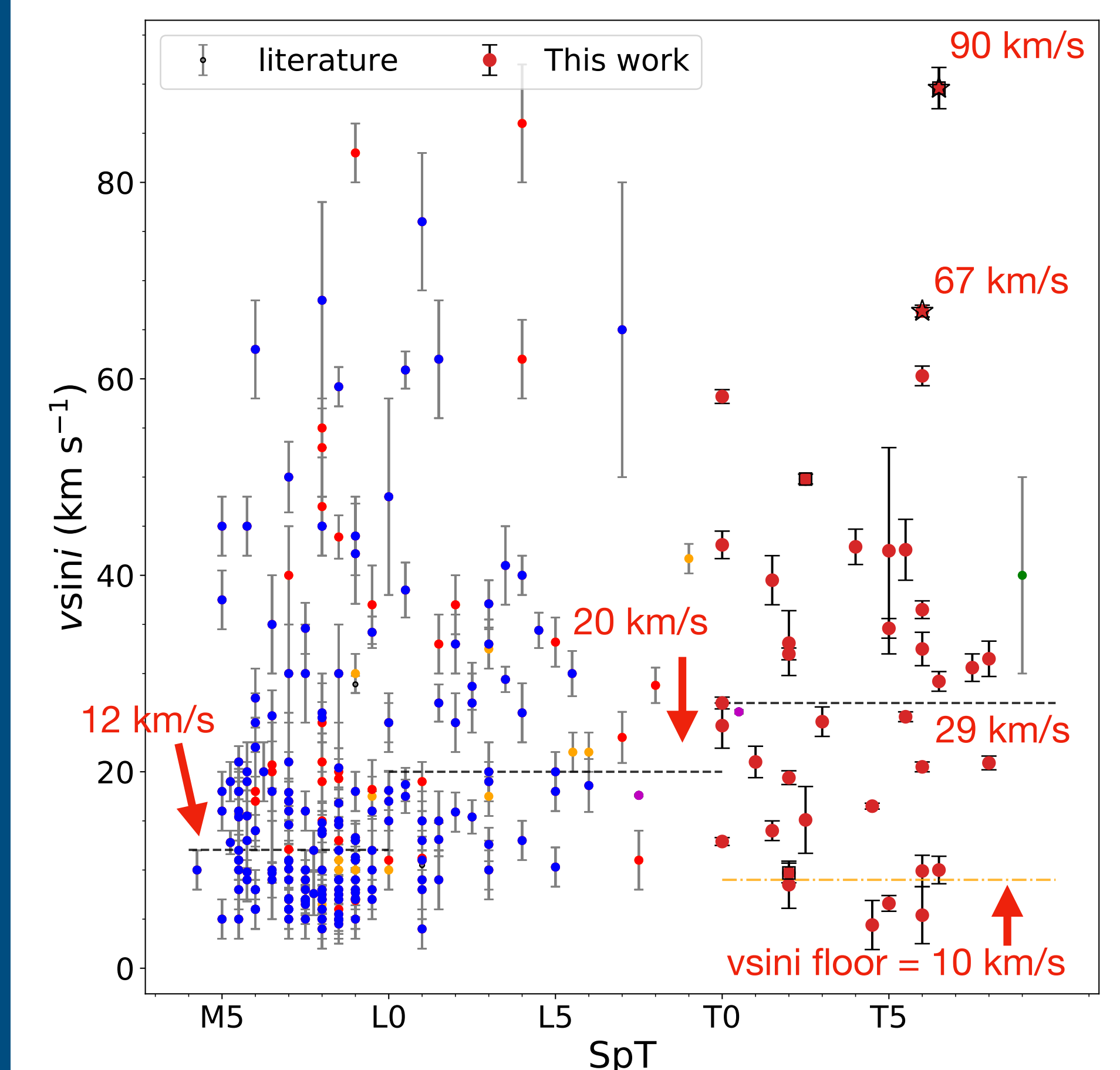


Figure 5. $v_{\text{sin}i}$ measurements as a function of spectral type for a compilation of M4–T9 dwarfs from this work (large symbols) and the literature (small symbols).

We compare $v_{\text{sin}i}$ measurements for M4–T9 dwarfs and our 37 T dwarfs (red circles) in Figure 5. The median $v_{\text{sin}i}$ values increase with later spectral types. The young and fastest T dwarfs are labeled in squares and stars, respectively. T dwarfs are generally fast rotators, indicating little angular momentum loss compared to earlier spectral types.

Acknowledgements

References

- Allard et al. 2012, 2012, RSPTA, 370, 2765
- Aumer & Binney 2009, MNRAS, 397, 1286
- Baraffe et al. 2003, A&A, 402, 701
- Bensby et al. 2003, A&A, 410, 52
- Blake et al. 2010, ApJ, 723, 684
- Burgasser et al. 2015a, ApJS, 220, 18
- Burgasser et al. 2016, ApJ, 827, 25
- Dieterich et al. 2014, AJ, 147, 94
- Dupuy & Liu 2017, ApJS, 231, 15
- Faherty et al. 2009, AJ, 137, 1
- Filippazzo et al. 2015, ApJ, 810, 158
- Johnson & Soderblom 1987, AJ, 93, 864
- Moehler et al. 2014, A&A, 568, A9
- Schonrich et al. 2010, MNRAS, 403, 1829
- Zapatero Osorio et al. 2007, ApJ, 666, 1205
- Zuckerman & Song 2004, ARA&A, 42, 685

