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X-ray study of Abell 3365 with *XMM-Newton*

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ABSTRACT

We present an X-ray spectral analysis using *XMM-Newton*/EPIC observations (~ 100 ks) of the merging galaxy cluster Abell 3365 ($z = 0.093$). Previous radio observations suggested the presence of a peripheral elongated radio relic to the east and a smaller radio relic candidate to the west of the cluster center. We find evidence of temperature discontinuities at the location of both radio relics, indicating the presence of a shock with a Mach number of $\mathcal{M} = 3.5 \pm 0.6$ towards the east and a second shock with $\mathcal{M} = 3.9 \pm 0.8$ towards the west. We also identify a cold front at $r \sim 1.6'$ from the X-ray emission peak. Based on the shock velocities, we estimate that the dynamical age of the main merger along the east-west direction is ~ 0.6 Gyr. We find that the diffusive shock acceleration scenario from the thermal pool is consistent with the electron acceleration mechanism for both radio relics. In addition, we studied the distribution of the temperature, iron (Fe) abundance, and pseudo-entropy along the merging axis. Our results show that remnants of a metal-rich cool-core can partially or entirely survive after the merging activity. Finally, we find that the merger can displace the metal-rich and low entropy gas from the potential well towards the cold front, as has been suggested via numerical simulations.

Key words. X-rays: galaxies: clusters – galaxies: clusters: intracluster medium

1. Introduction

Galaxy clusters are the largest bound structures in the Universe, growing hierarchically through the accretion and merging of less massive subclusters. During this merging process, turbulence, shocks, and cold fronts (CF) arise in the hot intracluster medium (ICM) as a result of the strong merging activity (Markevitch & Vikhlinin 2007). Cold fronts delimit the boundary between the infalling subcluster's cool dense gas and the hot cluster atmosphere. Merger shocks are large-scale structures that propagate along the hot ICM out to the cluster outskirts, where they are usually found. Both can be detected by temperature and density discontinuities in the ICM via spectral analysis and X-ray surface brightness (SB) edges, respectively. Moreover, shocks present a discontinuity also in the pressure distribution, while cold fronts maintain it uniformly across the edge (Vikhlinin et al. 2001). Shocks in galaxy clusters (typically characterized by a Mach number, $\mathcal{M} \leq 3-5$) are thought to (re)accelerate electrons in the ICM via first-order Fermi diffusive shock acceleration (hereafter DSA, Bell 1987; Blandford & Eichler 1987). These relativistic particles may produce non-thermal synchrotron emission in the form of radio relics in the presence of an amplified magnetic field. These radio features are elongated, polarized, and steep-spectrum ($\alpha < -1$, being $S_\nu \propto \nu^\alpha$) structures that are typically located at the cluster periphery (for a theoretical and observational review, respectively, see Brunetti & Jones 2014; van Weeren et al. 2019). Turbulence generated in the merger may also (re)accelerate

particles, generating unpolarized cluster-wide radio phenomena known as radio halos (Brunetti et al. 2001; Feretti et al. 2012; van Weeren et al. 2019).

Today, the number of shock detections in merging galaxy clusters is increasing (see, e.g., Akamatsu et al. 2017; Thölken et al. 2018; Urdampilleta et al. 2018; Di Gennaro et al. 2019; Botteon et al. 2019). In some cases, these shocks appear at the same or a similar location as radio relics (see Table 1 of van Weeren et al. 2019). The study of the shock-relic connection and the relative radial distribution throughout the ICM serves to improve our understanding of the dynamical stage of the cluster, as well as the physics of cosmic ray acceleration. Moreover, a detailed spectral analysis along the merging axis, including a study of their metallicity distribution, can provide valuable information on the chemical evolution and dynamical history of the merged ICM (Urdampilleta et al. 2019).

In this paper, we analyze the temperature and surface brightness discontinuities in the outskirts and central ICM of Abell 3365 (Abell et al. 1989, hereafter A3365), together with the temperature, Fe abundance, and pseudo-entropy distribution along the main merging axis. To this aim, we use observations taken by the *XMM-Newton* satellite. A3365 is a nearby ($z = 0.093$, Struble & Rood 1999) multi-merging galaxy cluster detected in X-rays by *ROSAT* and named RXC J0548.8-2154 (Böhringer et al. 2007). A3365 has $M_{500} = 1.7 \times 10^{14} M_\odot$ obtained from X-ray observations (Lovisari & Reiprich 2019). Based on this value and using the self-similar quantities described in Appendix A of Arnaud et al. (2010), we estimate

$r_{500} = 0.81 \text{ Mpc} = 7.82'$ (at this redshift) and $T_{500} = 3.14 \text{ keV}$. A3365 hosts one radio relic at the east (hereafter relic E) and one radio relic candidate at the west (hereafter, relic CW) discovered in the NVSS survey and observed by van Weeren et al. (2011) with the VLA and WSRT at 1.4 GHz. As described by van Weeren et al. (2011) relic E and relic CW have an angular extension of $5.5'$ or 560 kpc and $2.3'$ or 235 kpc, respectively, at this redshift (see Fig. 1).

A3365 consists of three subclusters (Golovich et al. 2019), see Fig. 1. They all have similar redshifts and the main components, 1 and 3, are separated by a distance of $9'$. They are considered part of the same merging system (Golovich et al. 2019), referred to in this paper as A3365. The mass ratio between these two main subclusters is 5:1, (Bonafede et al. 2017). The most massive component 1 has itself two sub-components. The X-ray emission peak (hereafter 1b) is displaced by $2.5'$ towards the west from sub-component 1a. The recent spectroscopic survey by Golovich et al. (2019) shows that A3365_3 contains 150 cluster members, is located at a redshift of 0.09273 ± 0.00028 and has a velocity dispersion of $981 \pm 58 \text{ km s}^{-1}$. The red sequence distribution suggests that the merging system is formed by three subclusters (see their Fig. 25 for more details): one is coincident with A3365_1a towards the east and the other is in a similar position to A3365_3, namely, in the middle, and the last one is towards the west (hereafter, A3365_2). For more details, see Fig. 1. Golovich et al. (2019), which suggests that the main merging activity occurred along the east-west direction, where A3365 crosses the A3365_1a subcluster, stripping the gas of both subclusters in the east-west direction, intensely disturbing the ICM and forming the cometary tail towards the east. The spatial distribution appears to suggest that the merger between A3365_1a and A3365_3 is associated with the relic E. However, there is no clear connection between this merger and the relic CW, suggesting that it might have a different origin. Golovich et al. (2019) also explain that A3365_1a shows two bright galaxies, which may indicate an ongoing merging activity. Finally, A3365_2 seems to be advancing towards A3365_3, which might be an indication that it has still not merged. Here, we use *XMM-Newton* to investigate the merging scenario in more detail.

For this analysis, we use the values of protosolar abundances (Z_{\odot}) reported by Lodders et al. (2009). We assume the cosmological parameters $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.27$ and $\Omega_{\Lambda} = 0.73$, respectively, which give 104 kpc per 1 arcmin at $z = 0.093$. All errors are given at 1σ (68 %) confidence level unless otherwise stated and all the spectral analysis make use of the modified Cash statistics (Cash 1979; Kaastra 2017).

2. Observations and data reduction

The two available *XMM-Newton* observations of A3365 are listed in Table 1. Both data sets are reduced using the *XMM-Newton* Science Analysis System (SAS) v17.0.0, with the calibration files from June 2018. For this analysis we use only the observations of the EPIC instrument, containing the MOS and pn detectors. We first apply the standard pipeline commands `emproc` and `epproc`. Next, we filter the soft-proton (SP) flares by building Good Time Interval (GTI) files. For this process we use the method detailed in Urdampilleta et al. (2019) and extensively explained in Appendix A.1 of Mernier et al. (2015). We identify and exclude the point sources in the complete FOV with a circular region of $10''$ radius (see Fig. 1), except where higher radii are needed to cover larger sources (see Appendix A.2 of Mernier et al. 2015). We use the SAS task `edetect_chain` for

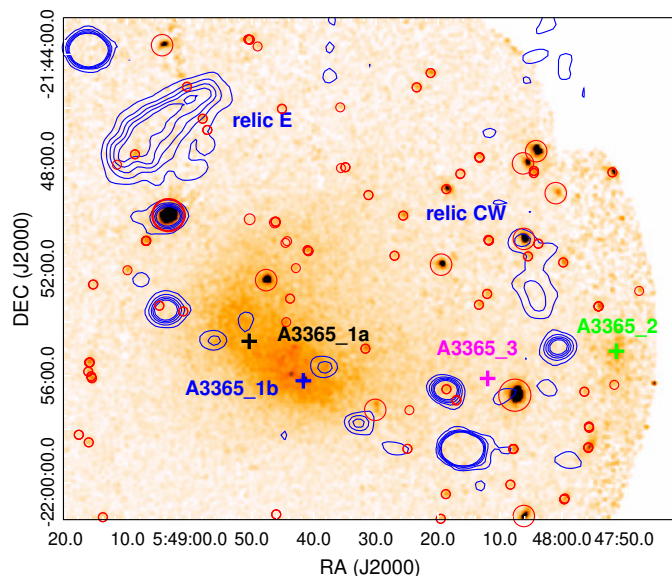


Fig. 1. *XMM-Newton* smoothed image of A3365 in the 0.5–2 keV band. Red circles are the point sources that have been removed. VLA 1.4 GHz radio contours are shown in blue. The center of the three subclusters: A3365_1a, A3365_2 and A3365_3, identified by Golovich et al. (2019) are marked with black, magenta and green crosses, respectively. The X-ray emission peak is shown with a blue cross, A3365_1b.

this purpose. We keep the single, double, triple, and quadruple events in MOS (`pattern ≤ 12`) and only the single pixel events in pn (`pattern == 0`). We also correct the out-of-time events from the pn detector.

3. Spectral analysis

In our spectral analysis of A3365, we assume that the observed spectra include the following components: optically thin thermal plasma emission from the ICM in collisional ionization equilibrium and the background consisting of the local hot bubble (LHB), the Milky Way halo (MWH), the cosmic X-ray background (CXB), the hard particle (HP) background, and residual soft-protons (SP). The models used for each of the background components are detailed in Urdampilleta et al. (2019) and Mernier et al. (2015). The estimation of the sky background components (LHB, MWH, and CXB) is obtained from a *ROSAT* observation of an offset annulus surrounding A3365, using the X-ray Background Tool¹. The parameters of the SP components are obtained fitting the total FOV ($r = 15'$) EPIC spectra, where the ICM *cie* (SPEX) model parameters are also left free. The best-fit parameters of the background components for each observation are listed in Table 2. The emission models are corrected for the cosmological redshift and absorbed by the galactic interstellar medium. We adopt the weighted total hydrogen column density $N_{\text{H,tot}} = 2.94 \times 10^{20} \text{ cm}^{-2}$ (Willingale et al. 2013)².

The spectral analysis software used in this work is SPEX³ (Kaastra et al. 1996, 2017) version 3.05.00 with SPEXACT (SPEX Atomic Code and Tables) version 3.05.00. We fit simultaneously the spectra of the MOS (energy range from 0.5 to 10 keV) and pn (0.6–10 keV) detectors. We apply the method

¹ <https://heasarc.gsfc.nasa.gov/cgi-bin/Tools/xraybg/xraybg.pl>

² <http://www.swift.ac.uk/analysis/nhtot/>

³ <https://www.sron.nl/astrophysics-spx>

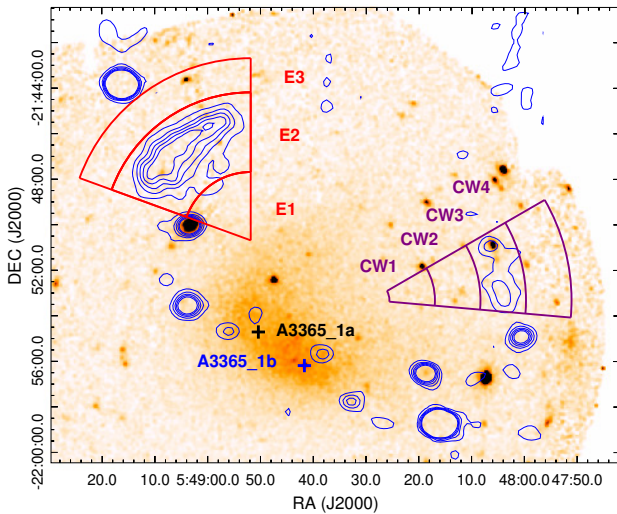
Table 1. Observations and net exposure times.

Sequence ID	Name	Position (J2000) (RA, Dec)	Observation starting date	MOS1 Net (ks)	MOS2 Net (ks)	pn Net (ks)
0760970101	A3365-p1	(05:48:47.00, -21:54:34.0)	2016-03-29	71.1	74.5	64.5
0677181301	A3365-p2	(05:48:50.38, -21:54:43.2)	2011-09-02	26.7	26.5	26.1

Table 2. Best-fit parameter values of the total background estimated for the A3365-p1 and A3365-p2 observations.

	$Norm^{(a)}$	kT (keV)	Γ
MWH	0.10 ± 0.10	0.17 ± 0.02	–
LHB	0.46 ± 0.07	0.08 (fixed)	–
CXB	1.81 ± 0.22	–	1.41 (fixed)
A3365-p1			
SP MOS1	1.32 ± 0.13	–	1.40 ± 0.02
SP MOS2	2.01 ± 0.17	–	1.40 ± 0.02
SP pn	4.88 ± 0.67	–	0.88 ± 0.03
A3365-p2			
SP MOS1	1.65 ± 0.37	–	1.40 ± 0.02
SP MOS2	2.19 ± 0.22	–	1.40 ± 0.02
SP pn	5.05 ± 0.25	–	1.40 ± 0.10

Notes. ^(a)For LHB and MWH norm in units of $10^{70} \text{ m}^{-3} \text{ arcmin}^{-2}$. For CXB norm in units of $10^{50} \text{ ph s}^{-1} \text{ keV}^{-1} \text{ arcmin}^{-2}$. For SP norm in units of $10^{45} \text{ ph s}^{-1} \text{ keV}^{-1} \text{ arcmin}^{-2}$.


Fig. 2. Same as Fig. 1 and indicating spectral extraction regions. The red and purple sectors represent the temperature radial distribution across relic E and CW, respectively. The center of A3365_1a and A3365_1b are marked with black and blue crosses.

of optimal binning (Kaastra & Bleeker 2016). A single temperature plasma is assumed for all the sectors in this study. We leave as free parameters the temperature, kT , the metal abundance, Z , and the normalization, $Norm$. All metal abundances are tied to the Fe abundance. We account for the effect of systematic uncertainties related to the $\pm 10\%$ variation in the normalization of the sky background and non X-ray background components (Mernier et al. 2015; Urdampilleta et al. 2019). The contribution of these systematic uncertainties is further included in our analysis and results.

4. X-ray surface brightness profile

In order to confirm the evidence of shocks and cold fronts characteristic of merging galaxy clusters, the presence of SB profile discontinuities should be found at the same location as temperature discontinuities. We adopt a broken power-law density profile to describe the density. Assuming spherical symmetry, the density distribution is given by:

$$\begin{cases} n_2(r) = n_0 \left(\frac{r}{r_{\text{sh}}}\right)^{-\alpha_2} & r \leq r_{\text{sh}} \\ n_1(r) = \frac{1}{C} n_0 \left(\frac{r}{r_{\text{sh}}}\right)^{-\alpha_1} & r > r_{\text{sh}}, \end{cases} \quad (1)$$

where n_0 is the model density normalization, r is the radius from the center of the sector, r_{sh} is the discontinuity location, and α_1 and α_2 are the power-law indices. At the location of the SB discontinuity the downstream (post-shock or post-CF) density, n_2 , is higher by a factor of $C = n_2/n_1$ compared with n_1 , the upstream (pre-shock or pre-CF) density. This factor, C , is known as the compression factor. We assume that the X-ray emissivity is proportional to the density squared ($S_X \propto n^2$) and integrate the cluster emission along the line of sight in the 0.5–2 keV energy band. We subtract the soft proton and the instrumental contribution, and all the point sources are removed. Afterwards, for the sky background, we fit the X-ray SB distribution with a constant model at large radii and we subtract it. Finally, we fit the X-ray SB, leaving all model parameters listed above as free to vary.

5. Results

5.1. Thermodynamical properties at the location of the radio relics

The diffuse radio emission in the A3365 outskirts (see Fig. 1) was classified as a radio relic in the east and a candidate radio relic in the west by van Weeren et al. (2011), as described in Sect. 1. In this work, we investigate the possible presence of shocks associated with these radio relics by analyzing the thermodynamical properties of the hot ICM. We first obtain a temperature radial distribution along both radio relics as shown in Fig. 2. The best-fit parameters are listed in Table 3. We compared their radial temperature profile with the “universal” distribution of Burns et al. (2010) for relaxed clusters. For that purpose, we assume the location of A3365_1a as the centroid, the average temperature as $T_{500} = 3.14 \text{ keV}$, and $r_{200} \sim 0.65r_{500} = 12'$ (Reiprich et al. 2013). In Fig. 3, both distributions present a clear deviation from the “universal” distribution at the location of the radio relics, which is probably caused by the presence of a shock.

Furthermore, we can calculate the average electron density as $n_e \sim 1.2n_H$ in the sectors shown in Fig. 2, given that $Norm = n_e n_H V$. The $Norm$ values are taken from the best-fit parameters listed in Table 3. V is the emitting volume projected along the line-of-sight (LOS). We assume that only the sphere between the maximum and minimum radius corresponding to each circular sector contributes to the emission (Henry et al. 2004;

Table 3. Best-fit parameters for the temperature radial distributions at the relic E and relic CW shown in Fig. 5.

	kT (keV)	$Norm$ ($10^{70} \text{ m}^{-3} \text{ arcmin}^{-2}$)	C-stat/d.o.f.
Relic E			
E1	3.54 ± 0.73	1.77 ± 0.11	1039/926
E2	4.15 ± 1.35	0.87 ± 0.06	947/1023
E3	0.97 ± 0.16	0.09 ± 0.04	1097/1024
Relic CW			
CW1	2.59 ± 0.46	2.91 ± 0.20	917/835
CW2	3.66 ± 0.80	2.41 ± 0.13	972/885
CW3	7.33 ± 2.25	2.26 ± 0.13	984/912
CW4	0.86 ± 0.10	0.33 ± 0.07	960/923

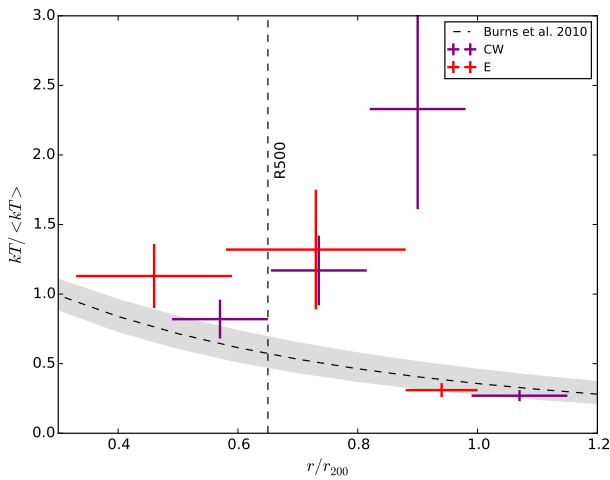


Fig. 3. Temperature distribution of the red and purple sectors of Fig. 2 compared with the Burns et al. (2010) “universal” profile for relaxed clusters. The gray shaded area shows the 1σ scatter.

Mahdavi et al. 2005). We estimate V as $V = 2SL/3$, with $L = 2\sqrt{(R_{\max}^2 - R_{\min}^2)}$, and S is the area of the sectors in the plane of the sky. The density radial distributions along both radio relics are plotted in Fig. 4. The distribution to the west shows a clear density discontinuity at the candidate radio relic edge, consistent with the temperature discontinuity, while a smoother gradient is present at the east.

Next we analyze more specific regions centered on the post and pre-shocks of each radio relic. In order to search for a possible temperature discontinuity, we performed a spectral analysis of the post-shock and pre-shock E and CW regions (see Fig. 5), assuming that the candidate shocks are located at the external edge of the respective radio relics. We maintain a separation between upstream and downstream regions of $\approx 1'$ to avoid any photon leakage from the brighter region. The Fe abundance is fixed to $0.3 Z_{\odot}$ in these regions. The best-fit parameters for the post and pre-shock regions at the relics E and CW are shown in Table 4. They both show a significant drop of the temperature. Shocks are characterized by higher temperature and pressure in the post-shock (downstream) than in the pre-shock (upstream) region. The temperature jumps ($kT_{\text{post}}/kT_{\text{pre}}$) are 4.6 ± 1.4 and 5.5 ± 1.2 for E and CW, respectively.

We also analyzed three regions (see red circular sectors in Fig. 6 left) orthogonal to the merging axis. We have compared their radial temperature profile with the “universal”

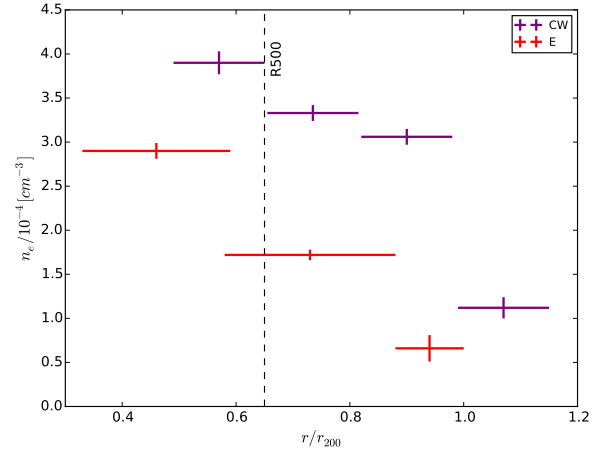


Fig. 4. Radial electron density distribution of the red and purple sectors of Fig. 2.

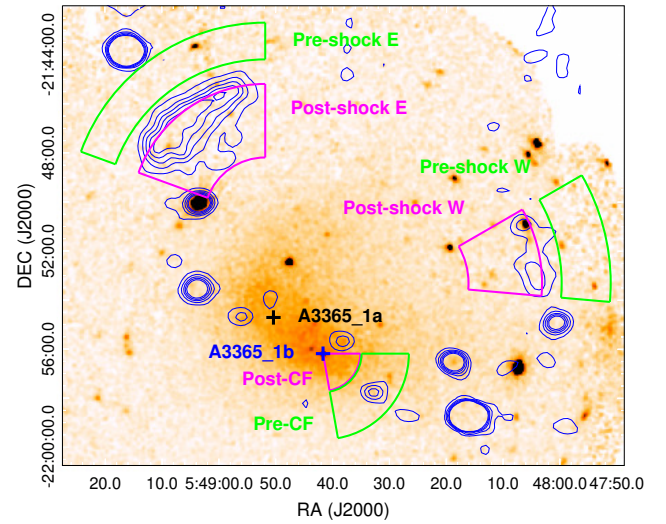


Fig. 5. Same as Fig. 1 and indicating spectral extraction regions. The green sectors represent the pre-shock and pre-cold front regions of relic E, relic CW, and CF. The magenta sectors are the post-shock and post-cold front regions. The center of A3365_1a and A3365_1b are marked with black and blue crosses.

Table 4. Best-fit parameters for the post and pre-shock regions at the relic E and relic CW shown in Fig. 5.

	kT (keV)	$Norm$ ($10^{70} \text{ m}^{-3} \text{ arcmin}^{-2}$)	C-stat/d.o.f.
Relic E			
Post	4.45 ± 1.08	1.10 ± 0.07	883/1009
Pre	0.97 ± 0.16	0.09 ± 0.04	1097/1024
Relic CW			
Post	4.72 ± 0.90	2.12 ± 0.10	1056/933
Pre	0.86 ± 0.10	0.33 ± 0.07	960/923

distribution of Burns et al. (2010) for relaxed clusters. As shown in Fig. 6 (right), the temperature profile in this direction presents a smooth decrease and is consistent with the expected distribution for relaxed clusters.

In these faint peripheral regions, the surface brightness profiles for the sectors corresponding to the E and W temperature

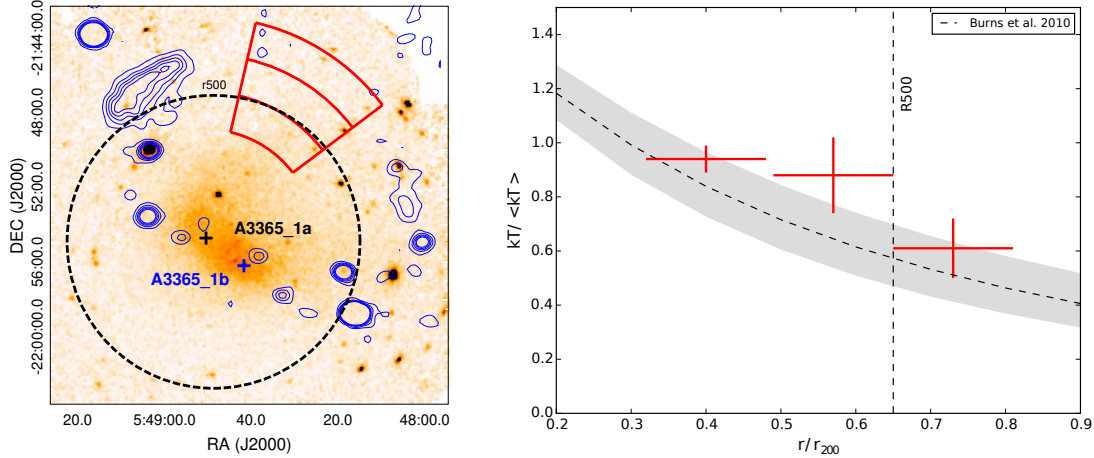


Fig. 6. *Left panel:* same as Fig. 1, showing extraction regions for the “relaxed” direction of the cluster (orthogonal to the merging axis) in red. The black dashed circle represents r_{500} . The center of A3365_1a and A3365_1b are marked with black and blue crosses. *Right panel:* temperature distribution of the red sectors compared with the Burns et al. (2010) “universal” profile for relaxed clusters. The gray shaded area shows the 1 σ scatter.

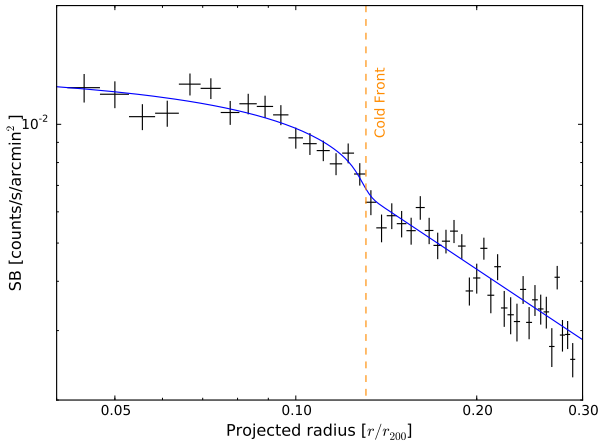


Fig. 7. Radial X-ray SB profile across the cold front in the 0.5–2.0 keV band using the *XMM-Newton* observations. The profile is corrected for exposure, background and soft protons, and point sources are removed. The best-fit model is shown in blue and the vertical dashed orange line represents the estimated position of the cold front.

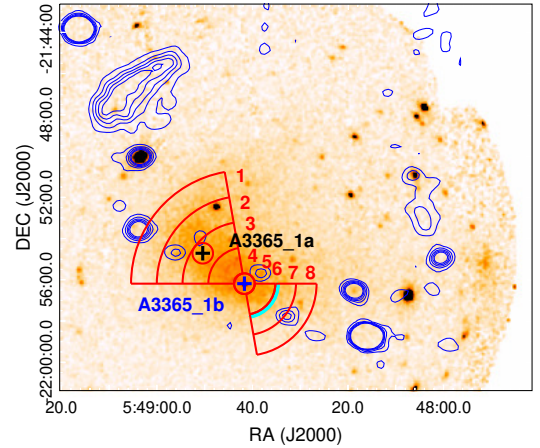


Fig. 8. Same as Fig. 1, where the red sectors represent the regions used for the temperature and abundance distributions. The cyan line shows the location of the cold front. The center of A3365_1a and the A3365_1b are marked with black and blue crosses.

Table 5. Best-fit parameters for the cold front regions shown in Fig. 5.

	kT (keV)	Z (Z_{\odot})	C-stat/d.o.f.
Post-CF	2.62 ± 0.10	0.59 ± 0.09	962/913
Pre-CF	3.30 ± 0.17	0.25 ± 0.08	1039/1003

jumps would need to be very heavily binned in order to detect a cluster signal. Therefore, we do not have sufficient measurements to fit a broken-power law model to these profiles that would allow us to constrain the underlying density jump. Nevertheless, we can estimate the density in the coarse regions used to determine the temperature jumps. These the density jumps should be considered as the upper limits.

For the relic E, we estimate $n_1 = (0.52 \pm 0.11) \times 10^{-4} \text{ cm}^{-3}$ and $n_2 = (1.64 \pm 0.04) \times 10^{-4} \text{ cm}^{-3}$. The compression factor derived from these densities is $C = 3.2 \pm 0.7$, which is in very good agreement with the value obtained from the temperature jump listed in Table 7. For the relic CW, the den-

sity values are $n_1 = (0.97 \pm 0.10) \times 10^{-4} \text{ cm}^{-3}$ and $n_2 = (2.41 \pm 0.06) \times 10^{-4} \text{ cm}^{-3}$. In this case, the compression factor is $C = 2.5 \pm 0.3$. This value is lower but close to the C derived from the temperature discontinuity described in Table 7. For the E and CW relics, therefore, both the density and temperature profiles exhibit evidence that would support the presence of a shock.

5.2. Substructure in the central ICM

We searched for the X-ray SB discontinuity in different edges of the bright and disturbed ICM. We find a clear discontinuity in the western side, close to the X-ray peak. For the X-ray SB profile analysis, we use a circular sector centered on the X-ray peak, with exactly the same shape as the post-CF (magenta) and pre-CF (green) regions together; see Fig. 5. The radial X-ray SB profile obtained is shown in Fig. 7. The best-fit parameters of the model show a discontinuity at $r_{\text{sh}} = 1.57 \pm 0.05$ arcmin with a compression factor value of $C = 1.33 \pm 0.08$. At this same location, we measure the temperature in the post-CF ($r = 0' - 1.57'$) and pre-CF ($r = 1.57' - 3.5'$) regions (see Fig. 5 and Table 5) and we obtain a temperature jump of $T_1/T_2 = 0.79 \pm 0.05$. Using

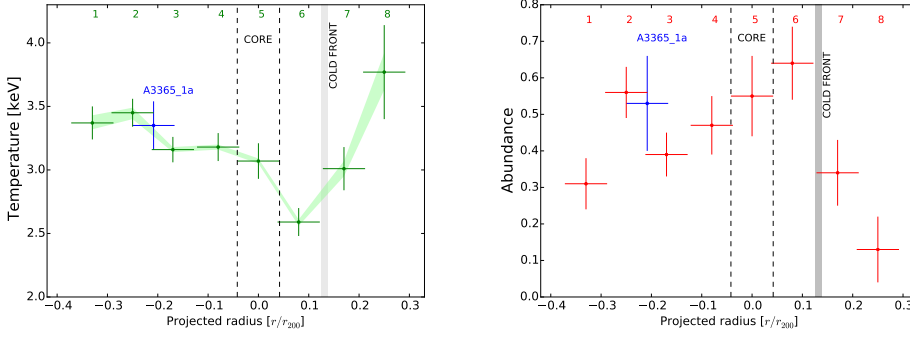


Fig. 9. *Left panel:* temperature distribution of A3365 along the merging direction. *Right panel:* abundance distribution of A3365. The shaded area in the left panel represent the systematic uncertainties. Systematic errors on abundance are much smaller than statistical and thus not shown in the right panel.

the temperature jump and the compression factor, we obtain a pressure ratio of $P_1/P_2 = 1.05 \pm 0.09$, consistent with a constant pressure value across the edge. This proves that the X-ray discontinuity found in the disturbed ICM edge is a cold front.

5.3. Temperature and Fe abundance distributions along the merging axis

As shown in Fig. 8, we cover the disturbed ICM elongation with circular sectors following the X-ray emission in the east-west direction. We dedicate two circles to the subcore of the main subcluster, A3365_1a, and X-ray peak (A3365_1b). We fit the cluster emission and the background spectra, excluding the point sources as explained in Sect. 3. We use the X-ray emission peak as the centroid for our temperature and abundance distributions along the merging axis (see Fig. 9).

The best-fit parameters of the temperature and abundance distributions are presented in Table 6. Sectors one to five, which are contained in the cometary tail, have a uniform temperature distribution between 3–3.5 keV. However, the temperature clearly decreases downstream of the cold front, increasing with a significant gradient in the cold front upstream region. The abundance distribution shows some signatures of two enhanced peaks, one associated with the main subcluster core, A3365_1a, and the other one with the cold front downstream region, both with values close to $\sim 0.6 Z_{\odot}$. The gas located just behind the cold front is displaced from the central potential well by the internal dynamical forces of the cold front (Heinz et al. 2003). These same effects have been already observed in Abell 3667 and Abell 665 (Lovisari et al. 2009; Urdampilleta et al. 2019).

6. Discussion

6.1. Shock jump conditions and properties

Based on the temperature jump found in the A3365 merging galaxy cluster outskirts, we report the presence of two X-ray shock candidates associated with the relic E and relic CW. Here, we present the shock properties (the Mach number, \mathcal{M} , the compression factor, C , and the shock velocity, v_{shock} , among others); see Table 7. The \mathcal{M} and C are calculated from the Rankine-Hugoniot jump condition (Landau & Lifshitz 1959) and assuming that all of the dissipated shock energy is thermalized and the ratio of specific heats (the adiabatic index) is $\gamma = 5/3$:

$$\frac{T_2}{T_1} = \frac{5\mathcal{M}^4 + 14\mathcal{M}^2 - 3}{16\mathcal{M}^2}, \quad (2)$$

$$C = \frac{n_2}{n_1} = \frac{4\mathcal{M}^2}{\mathcal{M}^2 + 3}, \quad (3)$$

where n is the density, and the indices 2 and 1 correspond to post-shock and pre-shock regions, respectively.

The Mach numbers estimated from the temperature discontinuities described in the previous section are $\mathcal{M}_E = 3.5 \pm 0.6$ and $\mathcal{M}_{CW} = 3.9 \pm 0.8$. The presence of shocks with $\mathcal{M} \geq 3.0$ is uncommon in galaxy clusters, with only few examples known until now (CIZA J224.8+5301, Akamatsu et al. 2015; “Bullet”, Shimwell et al. 2015; “El Gordo”, Botteon et al. 2016; A665, Dasadia et al. 2016; A3376, Urdampilleta et al. 2018).

The sound speed at the pre-shock regions is $c_{s,E} \sim 508 \text{ km s}^{-1}$ and $c_{s,W} \sim 479 \text{ km s}^{-1}$, assuming $c_s = \sqrt{\gamma k T_1 / \mu m_p}$ where $\mu = 0.6$. The shock propagation speed $v_{\text{shock}} = \mathcal{M} \cdot c_s$ for the east and west is $v_{\text{shock},E} = 1800 \pm 300 \text{ km s}^{-1}$ and $v_{\text{shock},W} = 1900 \pm 300 \text{ km s}^{-1}$, respectively. If we assume that the eastern shock is generated by the merger of A3365_1a and A3365_3 and it is moving with a constant velocity, we can calculate the dynamical age of the merger. The projected distance between A3365_1a and the relic E is $\sim 10'$, which means $\sim 1.04 \text{ Mpc}$ at $z = 0.093$. Thus, the time required to reach the actual position is $\sim 0.6 \text{ Gyr}$. We carry out the same exercise with the shock in the west, assuming that is generated by the same merger. In this case, the distance between A3365_1a and relic CW is $\sim 11' = 1.17 \text{ Mpc}$, resulting in a time of $\sim 0.6 \text{ Gyr}$. The upper limit of the merging time using the downstream gas velocity is, in both cases, $\sim 2.0 \text{ Gyr}$. Therefore, both shocks required the same time to propagate from the core passage up to their current location. We have no evidence of additional physical processes or mechanisms for the generation of the western shock. This might suggest that the possible origin of both shocks is the merger between A3365_1a and A3365_3.

As we explain in Sect. 1, the DSA is based on first-order Fermi acceleration and considers that there is a stationary and continuous injection. It accelerates relativistic electrons following a power-law spectrum $n(E)dE \sim E^{-p}dE$ with $p = (C + 2)/(C - 1)$, producing a radio synchrotron emission whose emissivity depends on the frequency as $S_{\nu} \propto \nu^{\alpha}$. The two power law indexes p and the radio injection spectral index α are related $\alpha = -(p - 1)/2$. From our analysis, we estimate the injection spectral index values of $\alpha_E = -0.68 \pm 0.06$ and $\alpha_W = -0.64 \pm 0.03$. Unfortunately, there are no available spectral indices derived from radio observations for their comparison to date. Future radio observations are needed to confirm these results.

6.2. Shock acceleration efficiency

Recent studies have demonstrated that the low efficiency of the DSA mechanism for low- \mathcal{M} shocks ($\mathcal{M} \leq 2-2.5$) is not enough for the electron acceleration to proceed from the thermal pool (for observational evidence on radio relics, see Botteon et al. 2020 and theoretical arguments, Vink & Yamazaki 2014, Ha et al. 2018 and Ryu et al. 2019). For this reason,

Table 6. Best-fit parameters for A3365 regions shown in Fig. 8.

Region	Radius ($'$)	kT (keV)	Z (Z_{\odot})	$Norm$ ($10^{70} \text{ m}^{-3} \text{ arcmin}^{-2}$)	C-stat/d.o.f.
1	-4.0 ± 0.5	3.37 ± 0.13	0.31 ± 0.07	10.12 ± 0.27	1077/1006
2	-3.0 ± 0.5	3.45 ± 0.11	0.56 ± 0.07	13.46 ± 0.33	1022/999
3	-2.0 ± 0.5	3.16 ± 0.10	0.39 ± 0.06	18.16 ± 0.43	1092/996
4	-1.0 ± 0.5	3.18 ± 0.11	0.47 ± 0.08	24.92 ± 0.79	967/922
5 (A3365_1b)	0.0 ± 0.5	3.07 ± 0.14	0.55 ± 0.11	25.97 ± 0.10	927/880
6	1.0 ± 0.5	2.59 ± 0.11	0.64 ± 0.10	14.90 ± 0.65	941/893
7	2.0 ± 0.5	3.01 ± 0.17	0.34 ± 0.09	8.63 ± 0.36	1014/927
8	3.0 ± 0.5	3.77 ± 0.37	0.13 ± 0.09	4.25 ± 0.19	1046/935
A3365_1a	-2.5 ± 0.5	3.35 ± 0.19	0.53 ± 0.13	17.44 ± 0.76	929/884

Table 7. Shock properties at relic E and relic CW.

	Mach No. $\mathcal{M}^{(a)}$	v_{shock} (km s^{-1}) ^(b)	Compression $C^{(c)}$	Power-law slope $p^{(d)}$	Spectrum index $\alpha^{(e)}$
relic E	3.5 ± 0.6	1800 ± 300	3.2 ± 0.2	2.36 ± 0.11	-0.68 ± 0.06
relic CW	3.9 ± 0.8	1900 ± 300	3.3 ± 0.2	2.28 ± 0.05	-0.64 ± 0.03

Notes. ^(a) \mathcal{M} is obtained from Eq. (2). ^(b) $v_{\text{shock}} = \mathcal{M} \cdot c_s$, $c_s = \sqrt{\gamma k T_1 / \mu m_p}$. ^(c) C from Eq. (3). ^(d) $p = (C + 2)/(C - 1)$. ^(e) $\alpha = -(p - 1)/2$ predicted.

alternative mechanisms have been proposed lately, as for example the re-acceleration of pre-existing cosmic ray electrons (Markevitch et al. 2005; Kang et al. 2012; Fujita et al. 2015).

In this case, if we assume DSA as the shock acceleration mechanism of the thermal electrons in radio relics, the electron acceleration efficiency, η_e , can be defined as the amount of kinetic energy flux at the shock moving at v_{shock} , which is converted into relativistic electrons and produces the synchrotron luminosity L_{sync} of the radio relic. Both parameters, η_e and L_{sync} , are related according to (Brunetti & Jones 2014):

$$\eta_e = \left[\frac{1}{2} \rho_1 v_{\text{shock}}^3 \left(1 - \frac{1}{C^2} \right) \frac{B^2}{B^2 + B_{\text{CMB}}^2} S \right]^{-1} \Psi(\mathcal{M}) L_{\text{sync}}, \quad (4)$$

where ρ_1 is the total upstream density, C the compression factor, S the shock surface, $B_{\text{CMB}} = 3.25(1+z)^2$ and B are the magnetic field equivalent for the cosmic microwave background radiation (CMBR) and the magnetic field for the radio relic in μG , respectively. The $\Psi(\mathcal{M})$ represents the ratio between the energy flux injected into “all” particles and those that are only visible in radio (for a more detailed description of the computation and equations, see Botteon et al. 2020).

Figure 10 shows the electron acceleration efficiency, η_e , as a function of the magnetic field, B , at both relics. For the relic E we assume $S_{1.4\text{GHz}} = 38.9 \text{ mJy}$ and $S = \pi \times ((LLS/2)^2) = \pi \times ((644/2)^2) \text{ kpc}^2$ with LLS being the radio relic largest linear size (see Nuza et al. 2017). The temperature and electron density in the downstream region are $kT_d = 4.45 \text{ keV}$ and $n_d = 1.64 \times 10^{-4} \text{ cm}^{-3}$, respectively. For the relic CW we assume $S_{1.4\text{GHz}} = 5.4 \text{ mJy}$, $S = \pi \times ((LLS/2)^2) = \pi \times ((235/2)^2) \text{ kpc}^2$ (see van Weeren et al. 2011), $kT_d = 4.72 \text{ keV}$ and $n_d = 2.41 \times 10^{-4} \text{ cm}^{-3}$. The density in the upstream region is obtained from the Eq. (3). The minimum momentum of electrons in both cases has been selected as $p_{\text{min}} = 0.1 m_e c$. We calculated η_e values for a lower value of $p_{\text{min}} = 0.08 m_e c$, based on the Figs. 1 and 2 of Botteon et al. (2020), and the difference can be neglected. We obtain η_e values for three different Mach numbers of $\mathcal{M} = 3.0, 3.5$ and 4.0 . In the case of $\mathcal{M} = 3.5$, for $B > 0.4 \mu\text{G}$, the electron

acceleration efficiency due to the shock is $\eta_e \lesssim 10^{-3}$ for both relics. Therefore, in this case the standard DSA scenario cannot be excluded, where efficiencies of this order are expected for weak shocks (Brunetti & Jones 2014; Caprioli & Spitkovsky 2014; Hong et al. 2014; Ha et al. 2018). These results also agree with the recent findings of Botteon et al. (2020) and Botteon et al. (2016) for “El Gordo” cluster, which suggest that DSA of thermal electrons is a valid mechanism in the case of shocks with high Mach number ($\mathcal{M} \geq 3$).

6.3. Cold front properties

We find evidence of X-ray SB and temperature discontinuities at the western edge of the highly disturbed ICM associated with a cold front. This front is perpendicular to the merging axis and moves presumably in the east-west direction, as shown in Fig. 5. It confines a cool and dense gas cloud moving through a hotter ambient gas (Markevitch & Vikhlinin 2007; Vikhlinin et al. 2001). The velocity of this cool gas cloud can be estimated using the density and temperature values derived from the X-ray SB and temperature discontinuities (Landau & Lifshitz 1959; Vikhlinin et al. 2001; Ichinohe et al. 2017; Urdampilleta et al. 2018). For that purpose, we assume that the cool gas is a spherical body moving through the ambient gas.

The ratio of pressures between the stagnation point (index 0, where the fluid $v = 0$, in front of the blob) and the free stream (index 1) can be expressed as a function of the gas cloud speed v (or equivalently the Mach number assuming $\mathcal{M}_1 = v/c_s$, where c_s is sound velocity of the free stream) (Landau & Lifshitz 1959):

$$\frac{p_0}{p_1} = \begin{cases} \left(1 + \frac{\gamma-1}{2} \mathcal{M}_1^2 \right)^{\gamma/(\gamma-1)} & (\mathcal{M}_1 \leq 1), \\ \left(\frac{\gamma+1}{2} \right)^{(\gamma+1)/(\gamma-1)} \mathcal{M}_1^2 \left(\gamma - \frac{\gamma-1}{2} \mathcal{M}_1^2 \right)^{-1/(\gamma-1)} & (\mathcal{M}_1 > 1), \end{cases} \quad (5)$$

where $\gamma = 5/3$ is the adiabatic index.

The gas dynamic parameters at the stagnation point cannot be measured directly. However, the pressure in the downstream

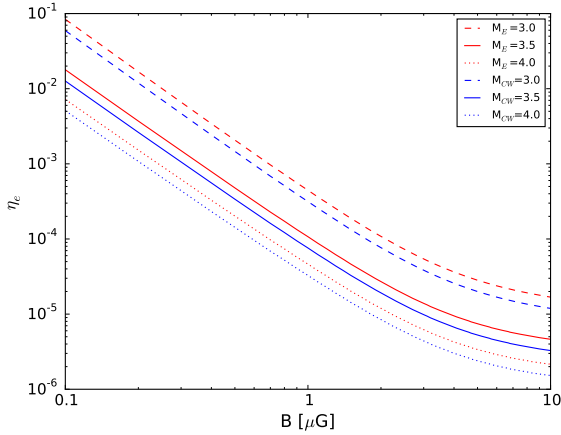


Fig. 10. Electron acceleration efficiency, η_e as a function of the magnetic field, B , for $\mathcal{M}=3.0, 3.5$ and 4.0 . Calculations assume a minimum momentum of electrons $p_{\min}=0.1m_e c$.

region of the cold front has a similar value as the stagnation pressure as explained by Vikhlinin et al. (2001). For this reason, we assume $T_0=2.62 \pm 0.10$ keV as the temperature downstream the cold front and $T_1=3.30 \pm 0.17$ keV as the free stream region, upstream (Vikhlinin et al. 2001; Sarazin et al. 2016). The density of these regions is calculated from the best-fit parameter of the broken power-law model described in Sect. 4 and shown in Fig. 7. The pressure inside the cold front is $p_0=T_0 \times n_0=7.7 \times 10^{-3} \pm 1.0 \times 10^{-3}$ keV cm $^{-3}$ and outside $p_1=T_1 \times n_1=4.2 \times 10^{-3} \pm 0.6 \times 10^{-3}$ keV cm $^{-3}$ (both densities are calculated at the average point of the downstream and upstream regions). Thus, the pressure ratio is $p_0/p_1=1.9 \pm 0.4$, which corresponds to a subsonic Mach number of $\mathcal{M}_1=0.92 \pm 0.07$ in the free stream. This value is consistent with the cold front in A3376 ($\mathcal{M}_1=1.2 \pm 0.2$, Urdampilleta et al. 2018) and slightly higher than A3667 ($\mathcal{M}_1=0.70 \pm 0.06$, Ichinohe et al. 2017). The sound speed in this region is $c_{cf}=\sqrt{\gamma k T_1/\mu m_p}=937 \pm 24$ km s $^{-1}$ with $k T_1=3.30 \pm 0.17$ keV, giving the velocity of the cool gas cloud as $v_{cf}=862 \pm 69$ km s $^{-1}$.

6.4. Pseudo-entropy distribution

We aim to better understand the thermal history of the disturbed ICM of A3365. For that reason, we calculate the pseudo-entropy, $K \equiv kT \times n^{-2/3}$, for each of the sectors described in Sect. 5.3. We use the data obtained from the spectral analysis, namely the temperature, kT , and the emission measure, $Norm$, to derive the density, n .

Figure 11 shows the pseudo-entropy distribution along the merging axis together with the Fe abundance profile. The distribution shows a clear low entropy minimum at the location of the X-ray peak of the merging system. Urdampilleta et al. (2019) observed the same behavior in two of the merging clusters Abell 665 and 1RXS J0603.3+4214 (known as the Toothbrush cluster). This suggests that the cool-core of the progenitor subcluster has totally or partially remained after the merging activity. As a consequence, we may be in the presence of a cool-core remnant (Rossetti & Molendi 2010). Moreover, the presence of the Fe abundance enhancement and the low temperature gas towards the contact discontinuity or cold front, supports the scenario described by the simulations of Heinz et al. (2003). As an effect of the merger, the cold dense core of the cluster is displaced from the center of the potential well. As the velocity of

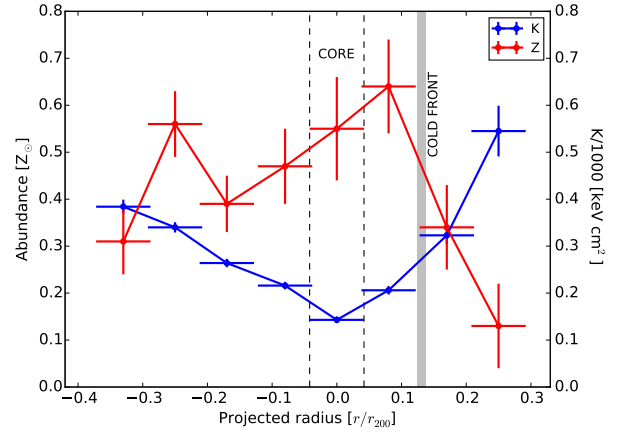


Fig. 11. Scaled pseudo-entropy ($K/1000$) and abundance distributions along the merging axis for A3365.

the core reverses in an attempt to return to equilibrium, a shear layer is created driving the gas along the cold front backwards in the direction of the background flow. This brings the material with the lowest temperature and highest abundance from the center of the cool core towards the cold front, as shown by our results in Fig. 11.

7. Summary

In this work, we present an imaging and spectral analysis of A3365 ($z=0.093$) using *XMM-Newton* observations (~ 100 ks). The recent spectroscopic survey by Golovich et al. (2019) demonstrates that A3365 is a complex merging system formed by three subclusters. Moreover, previous radio observations (van Weeren et al. 2011) suggest this merging galaxy cluster hosts a radio relic to the east and a radio relic candidate to the west of the merging cluster center.

Here, we detect abrupt temperature jumps across both radio relic edges, suggesting the presence of two shocks associated with them in the cluster outskirts. For these shocks, we estimate the Mach number and the shock velocity based on the temperature jump, namely, $\mathcal{M}=3.5 \pm 0.6$ and $v_{\text{shock}} \sim 1800 \pm 300$ for the shock in the east, and $\mathcal{M}=3.9 \pm 0.8$ and $v_{\text{shock}} \sim 1900 \pm 300$ for the shock in the west. Assuming that they are moving with constant velocity, the time since core passage is ~ 0.6 Gyr for both shocks. This might suggest that they originate in the same merger between A3365_1a and A3365_3. We compute the shock acceleration efficiency as a function of the magnetic field in the both relics and we find that DSA scenario is a possible acceleration mechanism. Similar results have been found for shocks with $\mathcal{M} \geq 3$ as “El Gordo” (Botteon et al. 2016) and A3376 (Botteon et al. 2020).

We also confirm, based on temperature and surface brightness discontinuities, the presence of a cold front at $r \sim 1.6'$ from the X-ray emission peak in the merging direction. It delimits a cool gas cloud moving with $\mathcal{M} \sim 0.9$ and a speed of $v \sim 860$ km s $^{-1}$.

In addition, we obtain the temperature, Fe abundance, and pseudo-entropy distribution in the highly disturbed central parts of the ICM along the merging axis. We find signatures of two abundance peaks, one that is coincident with A3365_1a ($\sim 0.6 Z_{\odot}$) and another that is located towards the cold front ($\sim 0.6 Z_{\odot}$) and displaced from the potential well. This is consistent with the simulations of Heinz et al. (2003), which suggest that the shock passage can affect the internal dynamics of the

cool gas and move the low entropy and metal rich gas towards the cold front. Finally, in line with previous results shown in, for example, Urdampilleta et al. (2019), the pseudo-entropy distribution shows a relative low-entropy minimum, which might suggest that remnants of the metal-rich cool-core can survive partially or fully after the merging activity.

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References

- Abell, G. O., Corwin, H. G., Jr, & Olowin, R. P. 1989, *ApJS*, 70, 1
- Akamatsu, H., van Weeren, R. J., O’Greehan, G. A., et al. 2015, *A&A*, 582, A20
- Akamatsu, H., Mizuno, M., Ota, N., et al. 2017, *A&A*, 600, A100
- Arnaud, M., Pratt, G. W., Piffaretti, R., et al. 2010, *A&A*, 517, A92
- Bell, A. R. 1987, *MNRAS*, 225, 615
- Blandford, R., & Eichler, D. 1987, *Phys. Rep.*, 154, 1
- Böhringer, H., Schuecker, P., Pratt, G. W., et al. 2007, *A&A*, 469, 363
- Bonafede, A., Cassano, R., Brügggen, M., et al. 2017, *MNRAS*, 470, 3465
- Botteon, A., Gastaldello, F., Brunetti, G., & Kale, R. 2016, *MNRAS*, 463, 1534
- Botteon, A., Shimwell, T. W., Bonafede, A., et al. 2019, *A&A*, 622, A19
- Botteon, A., Brunetti, G., Ryu, D., & Roh, S. 2020, *A&A*, 634, A64
- Brunetti, G., & Jones, T. W. 2014, *IJMPD*, 23, 1430007
- Brunetti, G., Setti, G., Feretti, L., & Giovannini, G. 2001, *MNRAS*, 320, 365
- Burns, J. O., Skillman, S. W., & O’Shea, B. W. 2010, *ApJ*, 721, 1105
- Caprioli, D., & Spitkovsky, A. 2014, *ApJ*, 783, 91
- Cash, W. 1979, *ApJ*, 228, 939
- Dasadia, S., Sun, M., Sarazin, C., et al. 2016, *ApJ*, 820, L20
- Di Gennaro, G., van Weeren, R. J., Andrade-Santos, F., et al. 2019, *ApJ*, 873, 64
- Feretti, L., Giovannini, G., Govoni, F., & Murgia, M. 2012, *A&ARv*, 20, 54
- Fujita, Y., Takizawa, M., Yamazaki, R., Akamatsu, H., & Ohno, H. 2015, *ApJ*, 815, 116
- Golovich, N., Dawson, W. A., Wittman, D. M., et al. 2019, *ApJ*, 882, 69
- Ha, J.-H., Ryu, D., & Kang, H. 2018, *ApJ*, 857, 26
- Heinz, S., Churazov, E., Forman, W., Jones, C., & Briel, U. G. 2003, *MNRAS*, 346, 13
- Henry, J. P., Finoguenov, A., & Briel, U. G. 2004, *ApJ*, 615, 181
- Hong, S. E., Ryu, D., Kang, H., & Cen, R. 2014, *ApJ*, 785, 133
- Ichinohe, Y., Simionescu, A., Werner, N., & Takahashi, T. 2017, *MNRAS*, 467, 3662
- Kaastra, J. S. 2017, *A&A*, 605, A51
- Kaastra, J. S., & Bleeker, J. A. M. 2016, *A&A*, 587, A151
- Kaastra, J. S., Mewe, R., & Nieuwenhuijzen, H. 1996, *11th Colloquium on UV and X-ray Spectroscopy of Astrophysical and Laboratory Plasmas*, 411
- Kaastra, J. S., Raassen, A. J. J., de Plaa, J., & Gu, L. 2017, <https://doi.org/10.5281/zenodo.2272992>
- Kang, H., Ryu, D., & Jones, T. W. 2012, *ApJ*, 756, 97
- Landau, L. D., & Lifshitz, E. M. 1959, *Fluid Mechanics, Course of Theoretical Physics* (Oxford: Pergamon Press)
- Lodders, K., Palme, H., & Gail, H. P. 2009, *Landolt Börnstein, New Series, Astronomy and Astrophysics* (Springer Verlag), 560 VI/4B
- Lovisari, L., & Reiprich, T. H. 2019, *MNRAS*, 483, 540
- Lovisari, L., Kapferer, W., Schindler, S., & Ferrari, C. 2009, *A&A*, 508, 191
- Mahdavi, A., Finoguenov, A., Böhringer, H., Geller, M. J., & Henry, J. P. 2005, *ApJ*, 622, 187
- Markevitch, M., & Vikhlinin, A. 2007, *Phys. Rep.*, 443, 1
- Markevitch, M., Govoni, F., Brunetti, G., & Jerius, D. 2005, *ApJ*, 627, 733
- Mernier, F., de Plaa, J., Lovisari, L., et al. 2015, *A&A*, 575, A37
- Nuza, S. E., Gelszinnis, J., Hoefl, M., & Yepes, G. 2017, *MNRAS*, 470, 240
- Reiprich, T. H., Basu, K., Ettori, S., et al. 2013, *Space Sci. Rev.*, 177, 195
- Rossetti, M., & Molendi, S. 2010, *A&A*, 510, A83
- Ryu, D., Kang, H., & Ha, J.-H. 2019, *ApJ*, 883, 60
- Sarazin, C. L., Finoguenov, A., Wik, D. R., & Clarke, T. E. 2016, ArXiv e-prints [arXiv:1606.07433]
- Shimwell, T. W., Markevitch, M., Brown, S., et al. 2015, *MNRAS*, 449, 1486
- Struble, M. F., & Rood, H. J. 1999, *ApJS*, 125, 35
- Thölken, S., Reiprich, T. H., Sommer, M. W., & Ota, N. 2018, *A&A*, 619, A68
- Urdampilleta, I., Akamatsu, H., Mernier, F., et al. 2018, *A&A*, 618, A74
- Urdampilleta, I., Mernier, F., Kaastra, J. S., et al. 2019, *A&A*, 629, A31
- van Weeren, R. J., Brügggen, M., Röttgering, H. J. A., et al. 2011, *A&A*, 533, A35
- van Weeren, R. J., de Gasperin, F., Akamatsu, H., et al. 2019, *Space Sci. Rev.*, 215, 16
- Vikhlinin, A., Markevitch, M., & Murray, S. S. 2001, *ApJ*, 551, 160
- Vink, J., & Yamazaki, R. 2014, *ApJ*, 780, 125
- Willingale, R., Starling, R. L. C., Beardmore, A. P., Tanvir, N. R., & O’Brien, P. T. 2013, *MNRAS*, 431, 394