



Universiteit
Leiden
The Netherlands

Newly Discovered Bright $z < 9-10$ Galaxies and Improved Constraints on Their Prevalence Using the Full CANDELS Area

Bouwens, R.J.; Stefanon, M.; Oesch, P.A.; Illingworth, G.D.; Nanayakkara, T.; Roberts-Borsani, G.; ... ; Smit, R.

Citation

Bouwens, R. J., Stefanon, M., Oesch, P. A., Illingworth, G. D., Nanayakkara, T., Roberts-Borsani, G., ... Smit, R. (2019). Newly Discovered Bright $z < 9-10$ Galaxies and Improved Constraints on Their Prevalence Using the Full CANDELS Area. *Astrophysical Journal* (0004-637X), 880(1), 25. doi:10.3847/1538-4357/ab24c5

Version: Not Applicable (or Unknown)
License: [Leiden University Non-exclusive license](#)
Downloaded from: <https://hdl.handle.net/1887/83130>

Note: To cite this publication please use the final published version (if applicable).

NEWLY DISCOVERED BRIGHT $Z \sim 9$ -10 GALAXIES AND IMPROVED CONSTRAINTS ON THEIR PREVALENCE USING THE FULL CANDELS AREA

BOUWENS, R.J.¹, STEFANON, M.¹, OESCH, P.A.^{2,3}, ILLINGWORTH, G.D.⁴, NANAYAKKARA, T.¹, ROBERTS-BORSANI, G.⁵, LABBÉ, I.⁶, SMIT, R.⁷

Draft version May 15, 2019

ABSTRACT

We report the results of an expanded search for $z \sim 9$ -10 candidates over the ~ 883 arcmin² CANDELS+ERS fields. This study adds 147 arcmin² to the search area we consider over the CANDELS COSMOS, UDS, and EGS fields, while expanding our selection to include sources with bluer $J_{125} - H_{160}$ colors than our previous $J_{125} - H_{160} > 0.5$ mag selection. In searching for new $z \sim 9$ -10 candidates, we make full use of all available HST, Spitzer/IRAC, and ground-based imaging data. As a result of our expanded search and use of broader color criteria, 3 new candidate $z \sim 9$ -10 galaxies are identified. We also find again the $z = 8.683$ source previously confirmed by Zitrin et al. (2015). This brings our sample of probable $z \sim 9$ -11 galaxy candidates over the CANDELS+ERS fields to 19 sources in total, equivalent to 1 candidate per 47 arcmin² (1 per 10 WFC3/IR fields). To be comprehensive, we also discuss 28 mostly lower likelihood $z \sim 9$ -10 candidates, including some sources that seem to be reliably at $z > 8$ using the HST+IRAC data alone, but which the ground-based data show are much more likely at $z < 4$. One case example is a bright $z \sim 9.4$ candidate COS910-8 which seems instead to be at $z \sim 2$. Based on this expanded sample, we obtain a more robust LF at $z \sim 9$ and improved constraints on the volume density of bright $z \sim 9$ and $z \sim 10$ galaxies. Our improved $z \sim 9$ -10 results again reinforce previous findings for strong evolution in the UV LF at $z > 8$, with a factor of ~ 10 evolution seen in the luminosity density from $z \sim 10$ to $z \sim 8$.

1. INTRODUCTION

Over the last few years, the search for galaxies in the early universe has revealed sources out to redshifts as high as $z \sim 11$ (Coe et al. 2013; Oesch et al. 2016), corresponding to 400 million years after the Big Bang. Simultaneous with these activities, tens of galaxies have been identified some 50-150 Myr later than this, at $z \sim 9$ -10 (Bouwens et al. 2011; Zheng et al. 2012; Coe et al. 2013; Ellis et al. 2013; McLure et al. 2014; Oesch et al. 2013, 2014; Zitrin et al. 2014; Bouwens et al. 2015, 2016; McLeod et al. 2015, 2016; Ishigaki et al. 2018; Oesch et al. 2018).

In the search for distant galaxies, one surprise was the discovery of very bright ($M_{UV,AB} \lesssim -22$) galaxies at $z \sim 9$ -11 (Oesch et al. 2014). Subsequent work within the HST Cosmic Assembly Near-Infrared Deep Extragalactic Legacy Survey (CANDELS) and Early Release Science (ERS) programs (Grogin et al. 2011; Koekemoer et al. 2011; Windhorst et al. 2011) and also pure parallel HST programs like BoRG/HIPPIES (Yan et al. 2011; Trenti et al. 2011) have added to the number of bright ($M_{UV,AB} \sim$

-21) $z \sim 9$ -10 candidates known (Bouwens et al. 2016; Calvi et al. 2016; Bernard et al. 2016; Livermore et al. 2018; Morishita et al. 2018). Of the known bright $z > 9$ galaxies, the most extreme example has been the $z = 11.1 \pm 0.1$ galaxy GN-z11 (Oesch et al. 2016), which owing to its exceptional brightness ($M_{UV,AB} \lesssim -22$) and high redshift must have required an especially rare, overdense region of the universe to form (Mutch et al. 2016; Waters et al. 2016).

Identifications of such bright galaxies have been useful not only because of the amenability of the sources for spectroscopic follow-up work and redshift determinations (Zitrin et al. 2015; Oesch et al. 2016; Hashimoto et al. 2018), but also because of the utility of such sources to further characterization, i.e., allowing for properties like their dust content (Watson et al. 2015), dynamical properties (Smit et al. 2018), UV-continuum slopes (Wilkins et al. 2016), stellar masses (Lam et al. 2019), and physical sizes (Holwerda et al. 2015) to be examined in detail.

Despite significant work done to the present in searching for $z \sim 9$ -10 galaxies, we can still make progress in expanding current $z \sim 9$ -10 samples using existing data sets. For example, in the Bouwens et al. (2016) search for $z \sim 9$ -10 candidate galaxies over the CANDELS fields, consideration was only given to those WFC3/IR regions with deep optical observations from the CANDELS program, i.e., roughly $\sim 85\%$ of the CANDELS area. In addition, Bouwens et al. (2016) focused on galaxies with the reddest $J_{125} - H_{160}$ colors in obtaining follow-up observations with HST (Bouwens 2014: GO 13752) both to search for the highest redshift sources ($z \gtrsim 8.9$) and to maximize the efficiency of the follow-up observations. By adopting this conservative approach, however, Bouwens et al. (2016) potentially missed some $z \lesssim 8.9$ galaxies

¹ Leiden Observatory, Leiden University, NL-2300 RA Leiden, Netherlands

² Observatoire de Genève, 1290 Versoix, Switzerland

³ International Associate, Cosmic Dawn Center (DAWN) at the Niels Bohr Institute, University of Copenhagen and DTU-Space, Technical University of Denmark

⁴ UCO/Lick Observatory, University of California, Santa Cruz, CA 95064

⁵ Astrophysics Group, Department of Physics and Astronomy, University College London, Gower Street, London, WC1E 6BT

⁶ Centre for Astrophysics & Supercomputing, Swinburne University of Technology, PO Box 218, Hawthorn, VIC 3112, Australia

⁷ Cavendish Laboratory, University of Cambridge, 19 JJ Thomson Avenue, Cambridge CB3 0HE, UK

with somewhat bluer $J_{125} - H_{160}$ colors.

Here, we expand the search for $z \sim 9-10$ galaxies to include the full ~ 883 arcmin² area within CANDELS+ERS. Our new search includes a 147 arcmin² area with deep WFC3/IR observations not utilized in previous work. We have expanded the area we consider within CANDELS, mostly by leveraging ground-based observations where deep ACS/optical data are not available. Our new search results also benefit from our considering sources with a broader set of $J_{125} - H_{160}$ colors than we had previously considered and inclusion of some additional HST follow-up observations taken in cycle 23 (GO 14459: Bouwens 2015). As part of our expanded search, we also pursue the selection of $z \sim 9-10$ sources using the HST and Spitzer/IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ data alone, in case confusion in the ground-based data resulted in our missing some sources in our earlier study (Bouwens et al. 2016).

The plan for this paper is as follows. In §2, we provide a brief description of the observational data we utilize and our procedures for performing photometry. In §3, we describe our selection procedure and results, while taking advantage of some cycle 23 observations to refine our constraints on the redshift of two $z \sim 9$ candidates we had identified. In §4, we take advantage of the new results to obtain improved estimates of the bright ends of the $z \sim 9$ and $z \sim 10$ LFs. We also attempt to quantify variations in the volume densities of bright $z \sim 9$ galaxies across the CANDELS fields. Finally, §5 summarizes the results of this paper. With the two appendices, we consider results from our HST follow-up program in cycle 23, while relaxing further our selection criteria for identifying $z \sim 9-10$ galaxies in the interests of constructing a more complete sample of such sources.

For convenience, we frequently write the HST F606W, F814W, F098M, F125W, F140W, and F160W filters as V_{606} , I_{814} , Y_{098} , J_{125} , JH_{140} , and H_{160} , respectively, throughout this work. Motivated by recent Planck results and for consistency with previous observational work (Planck Collaboration 2018), all results here are presented in terms of the standard concordance cosmology, with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70$ km/s/Mpc. We adopt the AB magnitude system (Oke & Gunn 1983) throughout.

2. OBSERVATIONAL DATA AND PHOTOMETRY

Here we make use of a ~ 883 arcmin² area with the five CANDELS+ERS fields to search for bright candidate $z \sim 9-10$ galaxies. Our total search area includes both previously searched regions of CANDELS (736 arcmin²: §2.2) and new search area (~ 147 arcmin²: §2.1).

2.1. New Search Area Within CANDELS

Here we make use of an additional ~ 147 arcmin² search area within CANDELS not considered in our earlier studies (Bouwens et al. 2015, 2016). Like the previously searched regions, the new region also features deep (~ 26.5 mag, 5σ) $J_{125} + H_{160}$ observations. Observations of this depth are of significant utility for finding $z \sim 9-10$ galaxy candidates, given the increasing prevalence of such candidates at $\gtrsim 25$ mag and especially $\gtrsim 26$ mag.

The primary reason we did not consider this area in Bouwens et al. (2016) was because of the lack of especially deep ACS/optical observations over much of the

area. The extreme outermost regions of the EGS mosaic were also not considered in our earlier search due to our lacking reductions of the HST data in those areas when devising our HST program to follow up specific $z \sim 9-10$ candidates (GO 13792: Bouwens 2014).

As with the case of the new EGS area, the regions considered here are located towards the edges of the CANDELS UDS and COSMOS fields (due to the roll angle constraints in scheduling HST observations and arranging for the ACS/optical parallels to land on the WFC3/IR observations). The new regions of the CANDELS UDS, COSMOS, and EGS mosaics are indicated using the light red shaded regions in Figure 1. The new areas probed in each field subtend 45.3, 48.7, 53.4 arcmin², respectively, or 147.4 arcmin² in total.

In constructing catalogs over these new WFC3/IR areas, we make use of the reductions from the 3D-HST team (Skelton et al. 2014) which are drizzled onto a $0.06''$ grid. Despite the lack of ACS coverage over the northern CANDELS COSMOS region from CANDELS, such coverage is available in the I_{814} band thanks to the original COSMOS program (Scoville et al. 2007). Deep ACS optical V_{606} and I_{814} -band coverage is available from the original CANDELS program over most of the new CANDELS EGS regions we search. The v1.2 reductions of the COSMOS ACS data (Koekemoer et al. 2007) and CANDELS EGS ACS data (Koekemoer et al. 2011) were retrieved from MAST and registered against the WFC3/IR observations.

In addition, we use ground-based observations over the new CANDELS regions to improve our constraints on the photometric redshifts of sources. Over the UDS, COSMOS, and EGS fields, use was made of the Cirasuolo et al. (2010) reductions of deep optical Subaru Suprime-Cam UDS/SXDS observations (Furusawa et al. 2008), the version 7 reductions of the CFHTLS survey observations⁸, and the very deep reductions (Capak et al. 2007) of the Subaru observations in the B , g , V , r , i , and z bands. At near-IR wavelengths, use of the especially sensitive DR3 observations over COSMOS with UltraVISTA in the $YJHK_s$ bands (McCracken et al. 2012), the sensitive version 7 UKIRT/WFCAM JHK_s observations over UDS, and CFHT/WIRCam K_s observations over EGS field (McCracken et al. 2010; Bielby et al. 2012). The depths of the optical data reach to $\sim 27-28$ mag (5σ) and in the near-IR, these data reach to $\sim 25-26$ mag (5σ).

Finally, for Spitzer/IRAC, use is made of the Spitzer/IRAC S-CANDELS and SEDS observations (Ashby et al. 2013, 2015), as well as any other Spitzer/IRAC observations available over the CANDELS regions. Reductions of these data were performed as detailed in Labbé et al. (2015).

2.2. Previously Searched CANDELS Areas

During the process of considering a larger search area within CANDELS, we also took the opportunity to conduct a second search for $z \sim 9-10$ candidates in regions already considered in Bouwens et al. (2016). We utilize the same reductions of the WFC3/IR and ACS observations as were presented in Bouwens et al. (2016). Those data sets reach to ~ 26.5 mag in the WFC3/IR bands (5σ) and ~ 27.0 mag in the optical bands. As in our previous

⁸ <http://www.cfht.hawaii.edu/Science/CFHTLS>

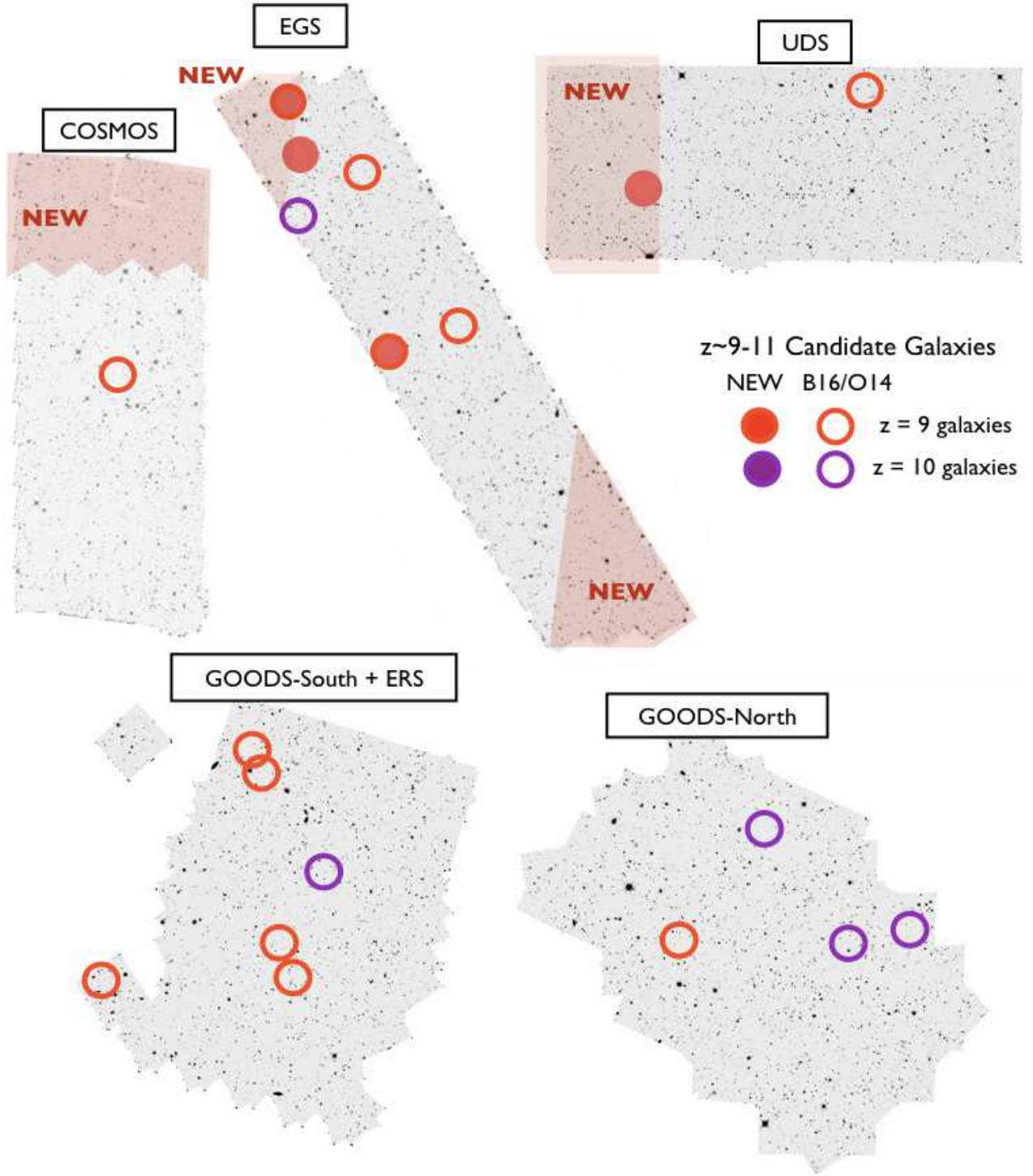


FIG. 1.— Observational footprints showing the layout of the sensitive J_{125} and H_{160} WFC3/IR observations over the five CANDELS fields. The regions enclosed by the red lines indicate the new WFC3/IR areas within CANDELS where searches for $z \sim 9-10$ galaxies are performed in this study. In our earlier study (Bouwens et al. 2016), optical ACS observations were not available to us when we performed our earlier study, and so these regions were not considered. Here ground-based data are utilized when deep ACS V_{606} and I_{814} data from CANDELS are not available. These regions correspond to 53.4 arcmin^2 , 48.7 arcmin^2 , and 45.3 arcmin^2 over the CANDELS EGS, COSMOS, and UDS fields, respectively, for a total area of 147 arcmin^2 . The solid red circles show the position of new ($P(z > 8) > 0.5$) $z \sim 9-10$ candidate galaxies identified by our selection criteria, while the open red and violet circles show the position of similar sources identified in our earlier studies (Oesch et al. 2014; Bouwens et al. 2015, 2016).

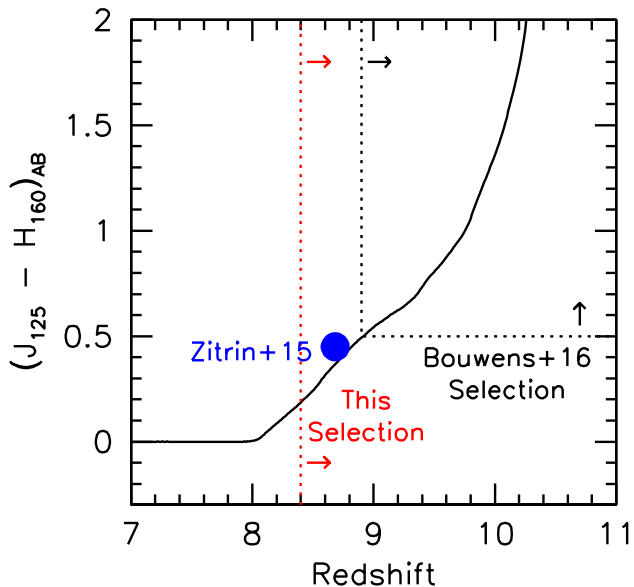


FIG. 2.— Expected $J_{125} - H_{160}$ colors for a star-forming galaxy with UV -continuum slope -2 vs. redshift. The red dotted lines and arrows indicate our inclusion of all sources with a photometric redshift in excess of 8.4. The new selection criteria are shown relative to the $J_{125} - H_{160} > 0.5$ color criterion used in Bouwens et al. (2016; *black dotted lines and arrow*) to identify source for follow-up observations. The redshift and measured $J_{125} - H_{160}$ color of the Zitrin et al. (2015) source is shown with the solid blue circle for context. Bouwens et al. (2016) focused on sources with $J_{125} - H_{160}$ colors > 0.5 to maximize the efficiency of their follow-up observations, but this resulted in their being more incomplete regarding their identification of sources in the redshift range $z \sim 8.4$ -8.9. Here we make use of more inclusive selection criteria to identify a larger number of star-forming galaxies at $z \gtrsim 8.4$.

work, use of the available ground-based + Spitzer/IRAC observations is made to obtain the best constraints on any candidate $z \sim 9$ -10 galaxies that are identified. The properties of these data sets are as described in the previous subsection (see also Bouwens et al. 2016).

2.3. Targeted Follow-up Observations

Also included in the present analysis are targeted follow-up observations of two candidate $z \sim 8.5$ galaxies. The coordinates of those candidates, COS910-5 and COS910-6, are 10:00:31.39, 02:26:39.8 and 10:00:20.12, 02:14:13.0, respectively. 1 orbit of Y_{098} observations were obtained on each (GO 14459: Bouwens 2015) as part of a cycle 23 program.

In executing the follow-up program, we adhered to a similar strategy as Bouwens et al. (2016) utilized in obtaining Y_{105} -band follow-up imaging observations of their candidate $z \sim 9$ -10 galaxies. The goal of the follow-up observations was to test if the sources showed essentially no flux at $\sim 1\mu\text{m}$ and bluer wavelengths. This is what one would expect if they were genuine $z \sim 9$ galaxies.

Given that the sources were identified after our cycle-22 program z9-CANDELS was complete, follow-up was requested in a subsequent HST program (GO 14459). Relative to the sources followed up as part of our cycle-22 program, these sources had slightly bluer $J_{125} - H_{160}$ colors, i.e., < 0.5 mag, and did not satisfy the selection criteria of Bouwens et al. (2016) whose $J_{125} - H_{160} > 0.5$ mag selection criteria preferentially identified sources with redshifts $z \gtrsim 8.9$ (see Figure 2). Both candidates also had H_{160} -band magnitudes brighter than 25.5 mag

and therefore could have an impact on the bright-end shape of the UV LF at $z > 6$. Given that this had been the subject of debate (e.g., Bowler et al. 2014 vs. Bouwens et al. 2015), follow-up of these candidates was considered to be important.

HST follow-up observations of COS910-5 and COS910-6 were obtained on February 27, 2016 and March 1, 2016, respectively. These observations were reduced using our WFC3RED pipeline (Magee et al. 2011) and drizzled onto the same astrometric frame as the ACS + WFC3/IR CANDELS data described in the previous subsection.

2.4. Photometry

Source catalogs were constructed for the new fields using the SExtractor software (Bertin & Arnouts 1996) and essentially an identical procedure to that utilized in previous papers by our team (e.g., Bouwens et al. 2015, 2017). Given that our search is for $z \sim 9$ -10 galaxies, source detection is performed using the H_{160} -band image alone. Our HST color measurements are made in smaller-scalable apertures based on a Kron (1980) factor of 1.2 and make use of the images after PSF correction to the H_{160} -band PSF. These smaller-scalable aperture flux measurements were then scaled up to total magnitudes by first accounting for the additional flux measured in larger scalable apertures (Kron factor of 2.5) relative to smaller scalable apertures and second accounting for the flux outside the larger scalable apertures and on the wings of the PSF. The former correction is made using the detection image, while the latter correction is made using the tabulated encircled energies in the WFC3/IR handbook (Dressel et al. 2012).

Photometry of candidate $z \sim 9$ -10 galaxies using the ground-based data and Spitzer/IRAC observations can also help us constrain their nature. To obtain these flux measurements, we need to cope with the very broad PSF in the ground-based and especially Spitzer/IRAC data which results in substantial overlap between sources. We use the MOPHONGO software (Labbé et al. 2010a, 2010b, 2013, 2015) to obtain flux measurements in the presence of source confusion. MOPHONGO uses the higher resolution HST data to create spatial templates for each source which is then used for modeling the ground-based or Spitzer/IRAC imaging data. The amplitudes of the templates are varied until a good fit to the imaging data is obtained, and then flux from the neighboring sources is subtracted. Photometry is performed on sources in $1.2''$ -diameter apertures for the ground-based data and $1.8''$ -diameter apertures for the Spitzer/IRAC data. Finally, the results are corrected to total based on the PSFs derived from the imaging data.

3. $Z \sim 9$ -10 SAMPLES

3.1. Selection Criteria

In this section, we describe the selection criteria we apply to the CANDELS UDS, COSMOS, and EGS data sets described in §2.1 and §2.2. Collectively, those data sets cover an area of 601 arcmin². As we have emphasized earlier, 147 arcmin² of this area was left unexplored in Bouwens et al. (2016). Meanwhile, the balance of the area, i.e., a 454 arcmin² region with the UDS, COSMOS, and EGS regions, was re-examined using a more inclusive selection criterion than we previously utilized. This will

TABLE 1
NEW CANDIDATE $z \sim 9$ -10 GALAXIES IDENTIFIED OVER THE CANDELS
UDS, COSMOS, AND EGS PROGRAMS (SEE §3.2)

ID	R.A.	Dec	$H_{160,AB}$	z_{phot}^a	$P(z > 8)^a$
UDS910-5	02:18:03.23	-05:13:21.7	25.8 ± 0.1	9.1	0.58
EGS910-8	14:20:52.51	53:04:11.7	25.7 ± 0.1	8.7	0.76
EGS910-9	14:20:45.23	53:02:01.3	26.1 ± 0.1	9.1	0.67
EGS910-10 ^b	14:20:08.50	52:53:26.6	25.3 ± 0.1	8.6 ^c	0.73

^a Best-fit $z > 4$ redshift and integrated $z > 8$ likelihood for source derived from our *HST*+*Spitzer*/IRAC+ground-based photometry (§2.1).

^b Also known as EGSY8p7. This source was previously identified by Roberts-Borsani et al. (2016) as a $z \sim 8.5$ candidate and spectroscopically confirmed by Zitrin et al. (2015).

^c This source has a measured spectroscopic redshift $z = 8.683$ (Zitrin et al. 2015).

TABLE 2
 $z \sim 9$ -11 GALAXY CANDIDATES IDENTIFIED OVER THE CANDELS FIELDS

ID	R.A.	Dec	$H_{160,AB}$	z_{phot}^b	$P(z > 8)$	Ref ^a
$z \sim 9$ Sample						
New Candidates from This Work:						
UDS910-5	02:18:03.23	-05:13:21.7	25.8 ± 0.1	9.1	0.58	
EGS910-8	14:20:52.51	53:04:11.7	25.7 ± 0.1	8.7	0.76	
EGS910-9	14:20:45.23	53:02:01.3	26.1 ± 0.1	9.1	0.67	
EGS910-10	14:20:08.50	52:53:26.6	25.3 ± 0.1	8.683	1.0 ^c	[6,7]
From Oesch et al. (2014) and Bouwens et al. (2016):						
COS910-1	10:00:30.34	02:23:01.6	26.4 ± 0.2	$9.0^{+0.4}_{-0.5}$	0.99	[8]
EGS910-0	14:20:23.47	53:01:30.5	26.2 ± 0.1	$9.1^{+0.3}_{-0.4}$	0.92	[8]
EGS910-3	14:19:45.28	52:54:42.5	26.4 ± 0.2	$9.0^{+0.5}_{-0.7}$	0.97	[8]
UDS910-1 ^b	02:17:21.96	-05:08:14.7	26.6 ± 0.2	$8.6^{+0.6}_{-0.5}$	0.74	[8]
GS-z9-1	03:32:32.05	-27:50:41.7	26.6 ± 0.2	9.3 ± 0.5	0.9992	[1], [8]
GS-z9-2	03:32:37.79	-27:42:34.4	26.9 ± 0.2	$8.9^{+0.3}_{-0.3}$	0.83	[8]
GS-z9-3	03:32:34.99	-27:49:21.6	26.9 ± 0.2	$8.8^{+0.3}_{-0.3}$	0.95	[3],[8]
GS-z9-4	03:33:07.58	-27:50:55.0	26.8 ± 0.1	$8.4^{+0.3}_{-0.3}$	0.97	[3],[8]
GS-z9-5	03:32:39.96	-27:42:01.9	26.4 ± 0.1	$8.7^{+0.8}_{-0.7}$	0.55	[8]
GN-z9-1	12:36:52.25	62:18:42.4	26.6 ± 0.1	9.2 ± 0.3	>0.9999	[1], [8]
$z \sim 10$ Sample						
From Oesch et al. (2014) and Bouwens et al. (2016):						
EGS910-2 ^d	14:20:44.31	52:58:54.4	26.7 ± 0.2	$9.6^{+0.5}_{-0.5}$	0.71	[8]
GN-z10-2	12:37:22.74	62:14:22.4	26.8 ± 0.1	9.9 ± 0.3	0.9994	[1], [2]
GN-z10-3	12:36:04.09	62:14:29.6	26.8 ± 0.2	9.5 ± 0.4	0.9981	[1], [2]
GS-z10-1	03:32:26.97	-27:46:28.3	26.9 ± 0.2	9.9 ± 0.5	0.9988	[1], [2]
$z \sim 11$ Sample						
From Oesch et al. (2016):						
GN-z11	12:36:25.46	62:14:31.4	26.0 ± 0.1	11.1 ± 0.1	1.0 ^e	[1], [2], [4], [5]

^a References: [1] Oesch et al. 2014, [2] Bouwens et al. 2015, [3] McLure et al. 2013, [4] Oesch et al. 2016, [5] Bouwens et al. 2010, [6] Zitrin et al. 2015, [7] Roberts-Borsani et al. 2016, [8] Bouwens et al. 2016

^b The likelihood of the new candidate $z \sim 9$ galaxies being secure $z > 8$ sources is lower than in earlier compilations by Oesch et al. (2014) and Bouwens et al. (2016). This is because the new candidates do not yet have deep HST coverage at $1\mu\text{m}$ from Y_{105} -band observations as the Oesch et al. (2014) and Bouwens et al. (2016) candidates possess.

^c This source has a measured spectroscopic redshift $z = 8.683$ (Zitrin et al. 2015).

^d The follow-up data obtained by the z9-CANDELS program did not significantly clarify the nature of this source (GO 13792: Bouwens 2014). Nevertheless, the available observations still support this source's being a credible $z \sim 9$ candidate.

^e This source has a measured spectroscopic redshift $z = 11.1 \pm 0.1$ (Oesch et al. 2016).

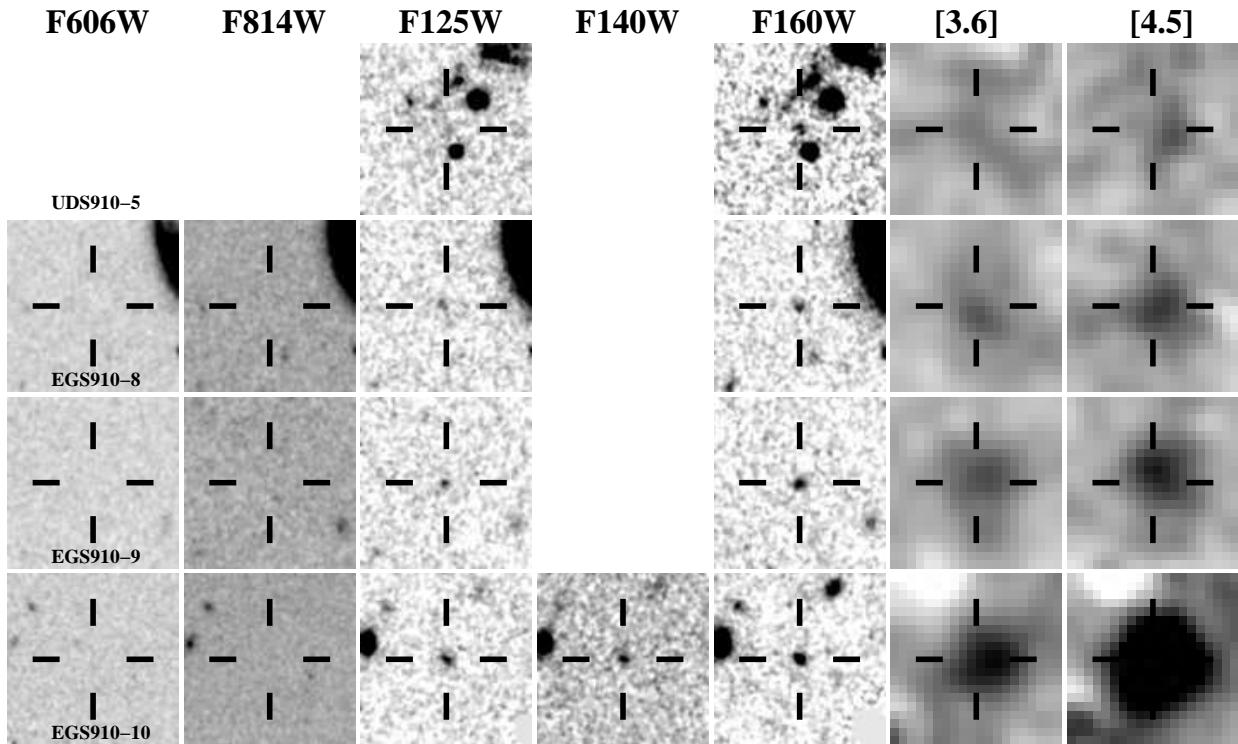


FIG. 3.— Postage stamp images of the $P(z > 8) > 0.5$ $z \sim 9-10$ galaxy candidates we have identified. HST and Spitzer/IRAC images of the candidates are presented (where available) from left to right in the V_{606} , I_{814} , J_{125} , JH_{140} , H_{160} , [3.6], and [4.5] bands. The presented postage stamps for the two Spitzer/IRAC bands is after subtraction of flux from the neighbors. Each of the presented sources is securely detected in the H_{160} band (and also the JH_{140} band when available), but not in the optical ACS V_{606} or I_{814} bands.

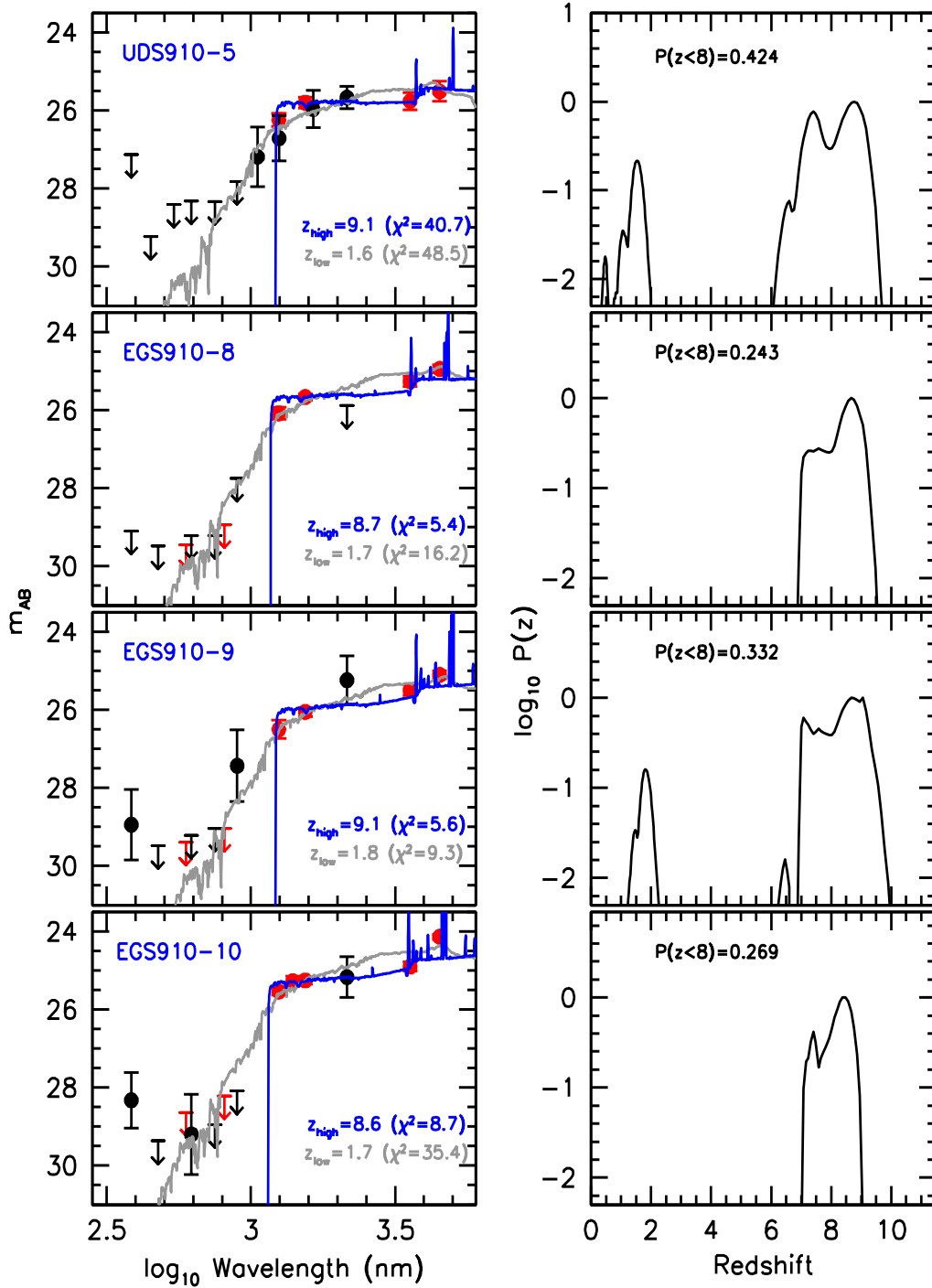


FIG. 4.— Photometric constraints available from HST and Spitzer/IRAC (red circles and upper limits) and also from various ground-based telescopes (black circles and upper limits) on several new $z \sim 9$ candidate galaxies identified within the new 147 arcmin^2 search area we considered. Upper limits are 1σ . The blue line shows the best-fit $z > 8$ model SED we find, while the gray line shows the best-fit lower-redshift SED we derive. The best-fit χ^2 and redshifts we find for the candidates are also indicated in the left-most panels. The right panels show the probability we compute for our $z \sim 9$ candidates to lie at specific redshifts. The total fractional probability that our $z \sim 9$ candidates lie at $z < 8$ is also indicated. Our photometry suggests all four candidates may have redshifts in excess of 8. While EGS910-10 has already been spectroscopically confirmed to have a redshift of 8.683 by Zitrin et al. (2015), the redshifts of the other three candidates are still quite uncertain.

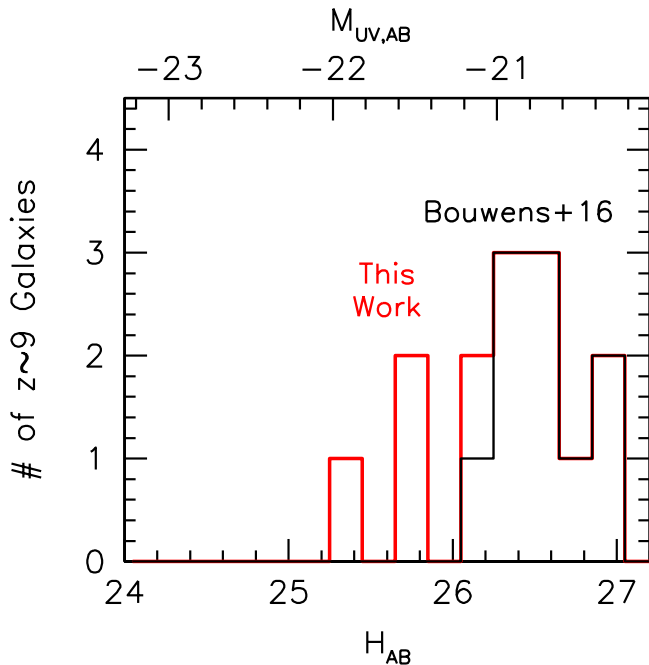


FIG. 5.— Number of $z \sim 9$ Candidates Identified over the CANDELS program vs. the apparent $H_{160,AB}$ -band Magnitude. The red histogram shows the current sample of $z \sim 9$ galaxy candidates, while the black histogram shows sample of $z \sim 9$ candidates identified in Bouwens et al. (2016). The upper axis shows the corresponding absolute magnitude of galaxies at $z \sim 9$ that corresponds to a given H_{AB} -band magnitude. In the present selection of $z \sim 9$ candidates, we find a larger fraction of bright ($H \leq 26.1$ mag) galaxy candidates than we identified in our previous study. The 25.3-mag candidate shown here was previously identified by Roberts-Borsani et al. (2016) and spectroscopically confirmed by Zitrin et al. (2015).

allow for the identification of more star-forming sources between $z \sim 8.4$ and $z \sim 9.0$ (see Figure 2).

In identifying probable candidate $z \sim 9$ -10 galaxies, we are guided by Lyman-break-like selection criteria. Significant spectroscopic work has shown that the Lyman-break selection technique is very effective in identifying galaxies at high redshifts (Steidel et al. 1996; Steidel et al. 2003; Vanzella et al. 2009; Stark et al. 2010; Smit et al. 2018).

To create a sample of potential $z \sim 9$ -10 candidates, we required that sources in our selection be detected at least at 6σ in the H_{160} band in a $0.35''$ -diameter aperture to ensure that sources in our selection are real.

Sources in our selection were also required to show a $\chi^2_{F606W,F814W}$ parameter less than 4. Following Bouwens et al. (2011), we defined the χ^2 parameter as $\chi^2_{F606W,F814W} = \sum_{i=[F606W,F814W]} \text{SGN}(f_i)(f_i/\sigma_i)^2$ where f_i is the flux in band i in a consistent aperture, σ_i is the uncertainty in this flux, and $\text{SGN}(f_i)$ is equal to 1 if $f_i > 0$ and -1 if $f_i < 0$.

For each source that satisfied our H_{160} -band detection criteria and optical non-detection criteria, we computed redshift likelihood distributions based on our HST, Spitzer, and ground-based photometry using the EAZY photometric redshift code (Brammer et al. 2008). In deriving the redshift likelihood distribution for each source, we make use of EAZY_v1.0 template set supplemented by SED templates from the Galaxy Evolutionary Synthesis Models (GALEV; Kotulla et al. 2009). Nebular contin-

uum and emission lines were added to the later templates using the Anders & Fritze-v. Alvensleben (2003) prescription, a $0.2Z_{\odot}$ metallicity, and a rest-frame EW for $H\alpha$ of 1300\AA .

Selected sources were required to have a maximum likelihood redshift $z \geq 8.4$ following the treatment in Bouwens et al. (2016). In Bouwens et al. (2016), we adopted this redshift limit to provide a midway point between our $z \sim 7$ z_{850} and $z \sim 8$ Y_{105} -dropout selections where the median redshift was 6.8 and 7.9. Use of $z = 8.4$ as the lower redshift limit for $z \sim 9$ samples also allows us to slightly increase the number of sources in our $z \sim 9$ samples, providing us with more leverage on the shape of the LF at early times.

In addition to requiring that sources being at $z \geq 8.4$, we also require that $\geq 50\%$ of their total integrated redshift likelihood be at $z > 8$. For two of the $z \sim 9$ -10 candidates over these fields, we could include the constraints obtained by targeted Y_{098} -band observations on the sources from a cycle-23 program (§2.3).

Finally, the $H_{160}-[3.6]$ colors of selected sources were required to be bluer than 1.4 mag to avoid selecting intrinsically red sources at lower redshifts. With such a $H-[3.6]$ limit, we would identify every extreme $z \sim 7$ -11 source in the Bowler et al. (2017), Stefanon et al. (2019), and Oesch et al. (2014) selections over $>1.5 \text{ deg}^2$, as we clearly show in Figure 12 from Appendix B. Conversely, sources with $H-[3.6]$ colors redder than 1.4 mag almost always appear to be at $z < 4$ (for the cases we have examined in Appendix B).

The above criteria differ from those utilized in Bouwens et al. (2015, 2016) in that they allow for the selection of sources with $J_{125}-H_{160}$ colors bluer than 0.5 mag. Note that a $J_{125}-H_{160}$ color of 0.5 mag corresponds to a redshift z of ~ 8.9 for sources with UV -continuum slopes of -2 (see Figure 2). Bouwens et al. (2016) explicitly did not consider such sources to focus on the highest redshift sources over CANDELS with their follow-up observations.

3.2. $z \sim 9$ -10 Sample

Application of the above selection criteria to our photometric catalogs allowed us to identify three new $z \sim 9$ candidates over the previously unexplored area that we searched from CANDELS UDS, COSMOS, and EGS (a $\sim 147 \text{ arcmin}^2$ area). One of these candidates was identified over the CANDELS UDS area, while the other two candidates were from the new CANDELS EGS area we examined. No new $z \sim 9$ candidates were identified over the area we considered from CANDELS COSMOS.

A fourth $z \sim 9$ galaxy was identified over the original 454 arcmin^2 area previously considered by us in Bouwens et al. (2016) over CANDELS UDS, COSMOS, and EGS. This source is the now well-known $z = 8.683 \pm 0.003$ galaxy initially identified by Roberts-Borsani et al. (2016) and spectroscopically confirmed by Zitrin et al. (2015).

Only one of the two $z \sim 9$ candidates targeted with HST Y_{098} -band observations from our cycle 23 program was determined to have a probable redshift in excess of 8. However, even in the case of that source, our best estimate for its redshift is 8.3, which would put it below our selection limit of 8.4.

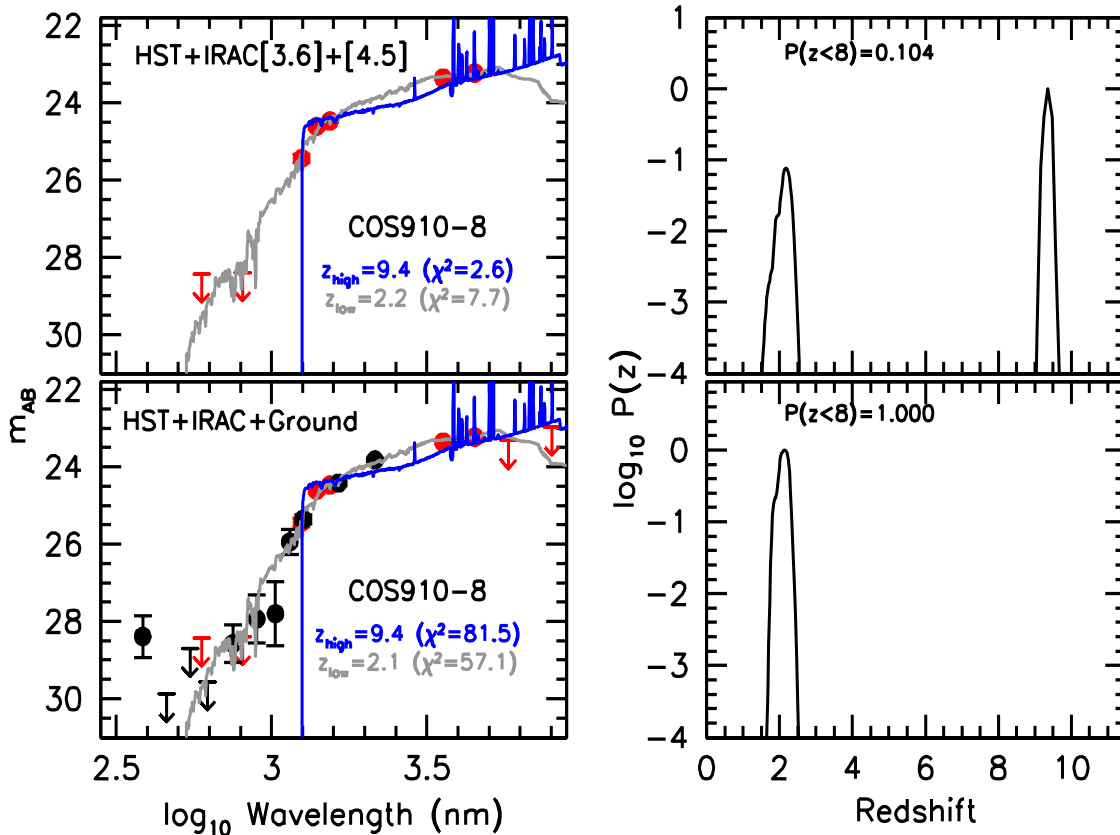


FIG. 6.— The SED fit and redshift likelihood distribution derived for a source COS910-8 (10:00:34.99, 02:14:01.1) from our search fields that prefers a $z > 9$ solution when using the HST+Spitzer/IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ photometry alone (*upper panel*), but which appears to be much more likely a $z \sim 2$ galaxy when incorporating the constraints from our ground-based photometry (*lower panel*: see also Appendix B for other examples). The lines and symbols shown in this figure are similar to Figure 10 from Appendix A. In the case that we rely on HST+IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ photometry alone, we estimate the redshift to be 9.4, with 90% of the probability lying at $z > 8$, but if we fold in the ground-based constraints, the redshift of this source seems much more likely to be ~ 2.1 , largely due to this source showing a 3σ detection in the FourStar J2 band (see Figure 11 from Appendix B) and the shape of continuum SED redward of the H band, including constraints in the four IRAC bands (see Table 6). Making use of our ground-based photometry (where care is exercised in subtracting the bright neighbor) to constrain the redshift of this source, we find that only 2×10^{-6} of the integrated probability is at $z > 8$. If we renormalize the flux uncertainties to obtain a more realistic reduced χ^2 , the probability that COS910-8 is at $z > 8$ is 2×10^{-4} .

Sources in the selected sample had $H_{160,AB}$ band magnitudes between 25.3 and 26.1 mag, similar to those identified by Bouwens et al. (2016) but roughly ~ 0.5 mag brighter in terms of their overall magnitude. See Figure 5. The brightest source in our selection was previously identified by Roberts-Borsani et al. (2016) by focusing on the brightest ($H_{160,AB} < 25.5$) sources and requiring them to be undetected in the optical data while showing very red [3.6]–[4.5] colors, as is expected at $z > 7$ due to the contribution of [OIII]+ $H\beta$ line emission to the $4.5\mu\text{m}$ fluxes.

In addition to the three new $z \sim 9$ candidate galaxies identified here, there are also 28 other $z \geq 7$ candidate galaxies identified over the CANDELS UDS, COSMOS, and EGS fields which, while mostly being lower quality candidates in general, could nevertheless be at $z \sim 9-10$ in a few cases. These sources are presented in Appendix A, B, and in Tables 4 and 5. Three of these candidates were identified as part of our earlier $z \sim 9-10$ search (Bouwens et al. 2016: their Table 7), but which we have been unable to thus far confirm through the acquisition of deeper HST data. 6 of the candidates have redshifts $z \leq 8.4$. 16 of the candidates were identified using the HST + Spitzer/IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ data alone,

in case confusion in the ground-based observations affected our selection (e.g., if there are spurious detections in the ground-based optical data).

Of all the candidate $z \sim 9-10$ sources included in Appendix B, a particularly interesting case is COS910-8, given its exceptional brightness $H_{160,AB} \sim 24.5$ mag. Using the HST+Spitzer/IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ photometry, we estimate a redshift of $z \sim 9.4$ for the source, with 90% of the probability lying at $z > 8$. While it would appear to be quite a compelling candidate based on our HST+Spitzer/IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ photometry, our assessment of this candidate depends sensitively on whether we incorporate constraints from our ground-based photometry or not. While we initially considered the possibility that this source might have been erroneously excluded from our earlier $z \sim 9-10$ selections (Bouwens et al. 2016) due to inclusion of optical flux from a bright neighbor, photometry on the source, after careful subtraction of the neighbor, indicates that this source is much more likely at $z \sim 2.1$, particularly owing to the apparent detection of this source in the $\sim 1.15\mu\text{m}$ J2 band at 3σ (see Figure 11 in Appendix B), the SED shape of the source defined by the H , K_s , and Spitzer/IRAC flux measurements, and tentative $1-2\sigma$ de-

tections of the source in the i , z , and Y bands. On the basis of the computed χ^2 , the probability that COS910-8 is at $z > 8$ is 2×10^{-6} . Given the high value of χ_{min}^2 relative to the number of constraints minus fit degrees of freedom, it is possible that our estimated flux uncertainties are too small. If we renormalize these uncertainties so that the reduced χ^2 is 1 for the best-fit solution, the probability that this candidate is at $z > 8$ is 2×10^{-4} .

We now return our discussion to the new $z \sim 9$ $P(z > 8) > 0.5$ candidates identified as part of this study (Table 1). Each of these candidates show a $>50\%$ probability of having a redshift $z > 8$. Nevertheless, each of them could have a redshift $z \lesssim 8$. Previously, Bouwens et al. (2016) had identified 9 $z \sim 9$ -10 candidates that showed a $\gtrsim 58\%$ probability of lying at $z > 8$ and followed them up with Y_{105} -band observations to test their robustness. Based on the follow-up observations, four sources were confirmed as robust $z > 8$ candidates, four sources were found to prefer a redshift $z \lesssim 8$, and for one source, the follow-up observations were not helpful in clarifying the nature of the source.

We will assume that follow-up of the new $z \sim 9$ candidates with Y_{105} -band observations would yield a similar $\sim 50\%$ contamination fraction to that obtained in our earlier follow-up efforts. Given that one of the four sources in our expanded selection was already spectroscopically confirmed, i.e., the Zitrin et al. (2015) $z = 8.683$ source, we assume that 2.5 of the $z \sim 9$ candidates are bona-fide $z \sim 9$ galaxies and 1.5 of our $z \sim 9$ candidate have redshifts $z < 8$.

We will combine the new identifications of $z \sim 9$ candidate galaxies with the Oesch et al. (2014) and Bouwens et al. (2016) identifications of $z \sim 9$ -10 candidate galaxies. Table 2 provides a comprehensive list of all the $z \sim 9$ -10 candidates we have identified.

In total, our selection includes 9 high-confidence ($P(z > 8) > 0.8$) $z \sim 9$ candidate galaxies, 3 high-confidence $z \sim 10$ candidate galaxies, and 1 spectroscopically confirmed $z \sim 11$ galaxy. Our $z \sim 9$ and $z \sim 10$ selections include 5 sources and 1 source, respectively, with $>55\%$ of the probability lying at $z > 8$. If all of these sources lie at $z > 8$, we find a total of 19 $z \sim 9$ -11 sources in the ~ 883 arcmin² area that make up the 5 CANDELS fields. This translates to a surface density of 1 bright $z \sim 9$ -11 candidate per 47 arcmin² (i.e., ≈ 10 WFC3/IR pointings).

4. LUMINOSITY FUNCTION RESULTS

In the present study, we have expanded our search for $z \sim 9$ -10 galaxies to cover a ~ 883 arcmin² area within CANDELS, which is an improvement on the 736 arcmin² area we previously considered in Bouwens et al. (2016). Thanks to our expanded search area, we were able to expand our $z \sim 9$ -10 selection from 15 candidate $z \sim 9$ -10 galaxies over the CANDELS fields to 19 such candidates.

In this section, we make use of our expanded $z \sim 9$ -10 sample and search area to improve our constraints on the prevalence of bright galaxies at $z \sim 9$ -10.

4.1. Luminosity Function Results

As in other studies of the LF (e.g., Steidel et al. 1999; Bouwens et al. 2007, 2008), we achieve constraints on our model LFs by comparing in detail with the surface density of $z \sim 9$ -10 candidates found in the data sets.

TABLE 3
NEW STEPWISE DETERMINATIONS OF THE
 UV LFs AT $z \sim 9$ AND $z \sim 10$ USING A
 ~ 883 ARCMIN² SEARCH AREA OVER ALL 5
CANDELS FIELDS

$M_{UV,AB}$ ^a	ϕ_k (10^{-3} Mpc ⁻³ mag ⁻¹)
$z \sim 9$ galaxies	
-22.72	$< 0.0014^b$
-21.92	$0.0008^{+0.0018}_{-0.0007}$
-21.12	$0.0074^{+0.0053}_{-0.0034}$
-20.32	$0.0246^{+0.0166}_{-0.0106}$
$z \sim 10$ galaxies	
-22.84	$< 0.0014^b$
-22.05	$< 0.0010^b$
-21.25	$0.0011^{+0.0025}_{-0.0009}$
-20.45	$0.0115^{+0.0111}_{-0.0062}$

^a Derived at a rest-frame wavelength of 1600Å.

^b 1σ upper limit.

Given the small number of $z \sim 9$ -10 candidates in our bright samples and especially per search field, we use luminosity bins of width 0.8 mag in constructing a UV LF and model the counts in each observational bin in the data sets as Poissonian. This results in the following estimated likelihood \mathcal{L} for a model LF given a set of observations:

$$\mathcal{L} = \prod_{i,j} e^{-N_{exp,i,j}} \frac{(N_{exp,i,j})^{N_{obs,i,j}}}{(N_{obs,i,j})!} \quad (1)$$

where $\prod_{i,j}$ is the product symbol and which runs over all search fields each denoted by index i and over all magnitude intervals denoted by index j , where $N_{obs,i,j}$ is the observed number of sources in search field i and magnitude interval j , and where $N_{exp,i,j}$ is the expected number of sources in search field i and magnitude interval j .

We compute the expected number of sources per bin, i.e., N_m , from a model LF as follows:

$$N_m = \sum_k \phi_k V_{m,k} \quad (2)$$

where N_m is the surface density of galaxies in some search field with magnitude m , ϕ_k is the volume density of galaxies with absolute magnitude k , and $V_{m,k}$ is the effective selection volume for which galaxies with absolute magnitude k will both satisfy our dropout selection criteria and be observed to have an apparent magnitude m . The binning we adopt for N_m and $V_{m,k}$ is the same 0.8-mag binning as we adopt for the stepwise LF ϕ_k .

Our procedure for estimating the selection volume $V_{m,k}$ is almost identical to what we do in many other of our recent papers on the high-redshift luminosity function (e.g., Bouwens et al. 2010, 2011, 2015). We begin by first constructing a mock catalog of sources with redshifts lying between $z \sim 7.5$ to $z \sim 11.5$, UV luminosities lying between -23 and -19.5 mag, and with random spatial positions within the CANDELS UDS, COSMOS, and EGS fields. Realistic images for sources from our mock catalogs are first simulated by artificially redshifting similar luminosity $z \sim 4$ sources from the Bouwens et al. (2015) HUDF samples to $z \sim 9$ -11 using the Bouwens et al. (1998, 2003) cloning procedure. Source sizes are scaled as $(1+z)^{-1.1}$ to match the galaxy size vs. red-

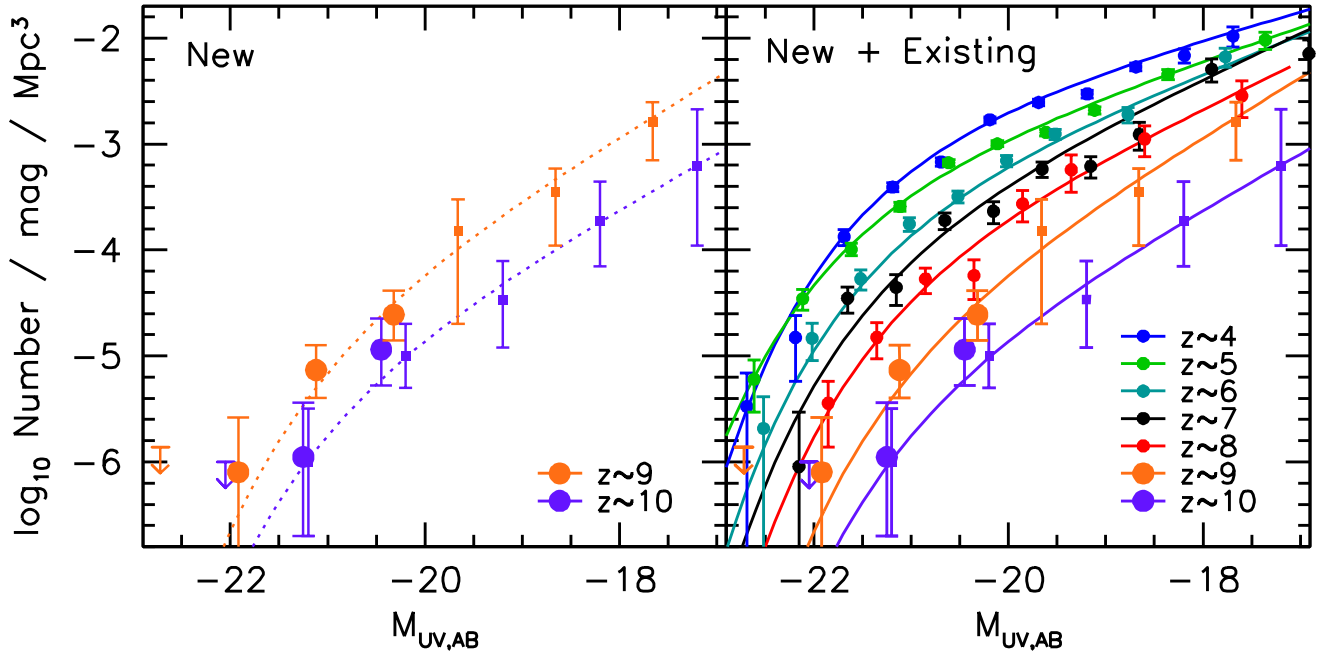


FIG. 7.— (left) Current determinations of the stepwise UV LF at $z \sim 9$ and $z \sim 10$ (large solid orange and purple circles, respectively; §4.1). The plotted error bars are 1σ . In deriving the $z \sim 10$ LF, sources out to $z \sim 11$ (i.e., Oesch et al. 2016) have been included. Table 3 presents our new results in tabular form. The dotted orange and purple lines show the $z \sim 9$ and $z \sim 10$ LF determinations from Bouwens et al. (2019, in prep) and Oesch et al. (2018), respectively. For context, the $z \sim 9$ LF from Oesch et al. (2013; orange solid squares) and $z \sim 10$ LF from Oesch et al. (2018; purple solid squares) at lower luminosities are also shown. (right) Our new LF determinations at $z \sim 9$ and $z \sim 10$ are shown alongside of those determinations by Bouwens et al. (2015) at $z \sim 4$ (blue solid circles), $z \sim 5$ (green solid circles), $z \sim 6$ (cyan solid circles), $z \sim 7$ (black solid circles), and $z \sim 8$ (red solid circles). The $z \sim 9$ and $z \sim 10$ results from Bouwens et al. (2019, in prep), Oesch et al. (2013), and Oesch et al. (2018) from the left panel are also shown in the right panel. The Schechter function fits derived at $z \sim 4$, $z \sim 5$, $z \sim 6$, $z \sim 7$, and $z \sim 8$ by Bouwens et al. (2015) are shown with the blue, green, cyan, black, red, orange, and purple lines, respectively.

shift trend observed to $z \sim 9-10$ (e.g., Ono et al. 2013; Holwerda et al. 2015; Shibuya et al. 2015). The UV -continuum slopes are set equal to -1.8 , consistent with that measured at high luminosities at $z \sim 5-8$ (Bouwens et al. 2012, 2014; Finkelstein et al. 2012; Willott et al. 2013; Rogers et al. 2014), with the dispersion set to 0.3 (Bouwens et al. 2012; Castellano et al. 2012).

In addition to producing HST images for sources in our mock catalogs, ground-based and Spitzer/IRAC images are also generated for all sources in our catalogs. These images are created by convolving the HST images by the appropriate PSF-matching kernel, i.e., HST-to-Spitzer/IRAC or HST-to-ground-based-image. The simulated HST, ground-based, and Spitzer/IRAC images of our mock sources are added to the real data and sources are detected, selected, and characterized in the same way as sources in the real observations. In this way, we compute the selection volume $V_{m,k}$, where sources in the absolute magnitude interval k are selected and found to have an apparent magnitude in the interval m .

We combine our new $z \sim 9$ candidate galaxies with those previously identified over the CANDELS fields. Motivated by the results of Bouwens et al. (2016; see §3.2) who only are able to confirm 50% of the $z > 8.4$ candidates with the Y_{105} -band follow-up observations, we assume the same for our new candidates lacking follow-up Y_{105} -band observations. For two $z \sim 9$ and $z \sim 10$ candidates for which our cycle-22 follow-up observations were indeterminate, i.e., UDS910-1 and EGS910-2, we treat these sources as 0.5 $z \sim 9$ and $z \sim 10$ candidates,

consistent with the 50% confirmation rate achieved with follow-up observations and consistent with the procedure applied in Bouwens et al. (2016). We treat all of the other previously presented candidates from Bouwens et al. (2016) as full candidates, with the exception of GS-z9-5 which we treat as half a candidate.

As discussed in Roberts-Borsani et al. (2016), there is reason to believe that the bright $z = 8.683$ source EGS910-10 (or EGSY8p7 as Zitrin et al. 2015) may benefit from lensing magnification from two massive intermediate redshift galaxies that lie within $3''$ of it. While the degree of magnification is uncertain, we assume it is the same factor of ≈ 2 that Roberts-Borsani et al. (2016) estimate, and therefore shift the source in magnitude by 0.75 mag. Since $z \sim 11$ galaxies also satisfy our $z \sim 10$ selection criteria (but likely constitute a very small fraction of that sample), we include the Oesch et al. (2016) GN-z11 $z = 11.1 \pm 0.1$ source in our $z \sim 10$ sample.

Using our expanded $z \sim 9$ and $z \sim 10$ samples and computed volumes, we recomputed the stepwise rest- UV LFs at $z \sim 9$ and $z \sim 10$. Our results are presented in Table 3 and the left panel of Figure 7. Our results are shown in the context of our previous results at $z \sim 4$, $z \sim 5$, $z \sim 6$, $z \sim 7$, and $z \sim 8$ in the right panel to Figure 7, and there is reasonably smooth evolution with redshift.

We can quantify the redshift evolution better by using our new stepwise UV LF determination to compute the total luminosity density brightward of -20 mag. The integrated luminosity density brightward of -20 is

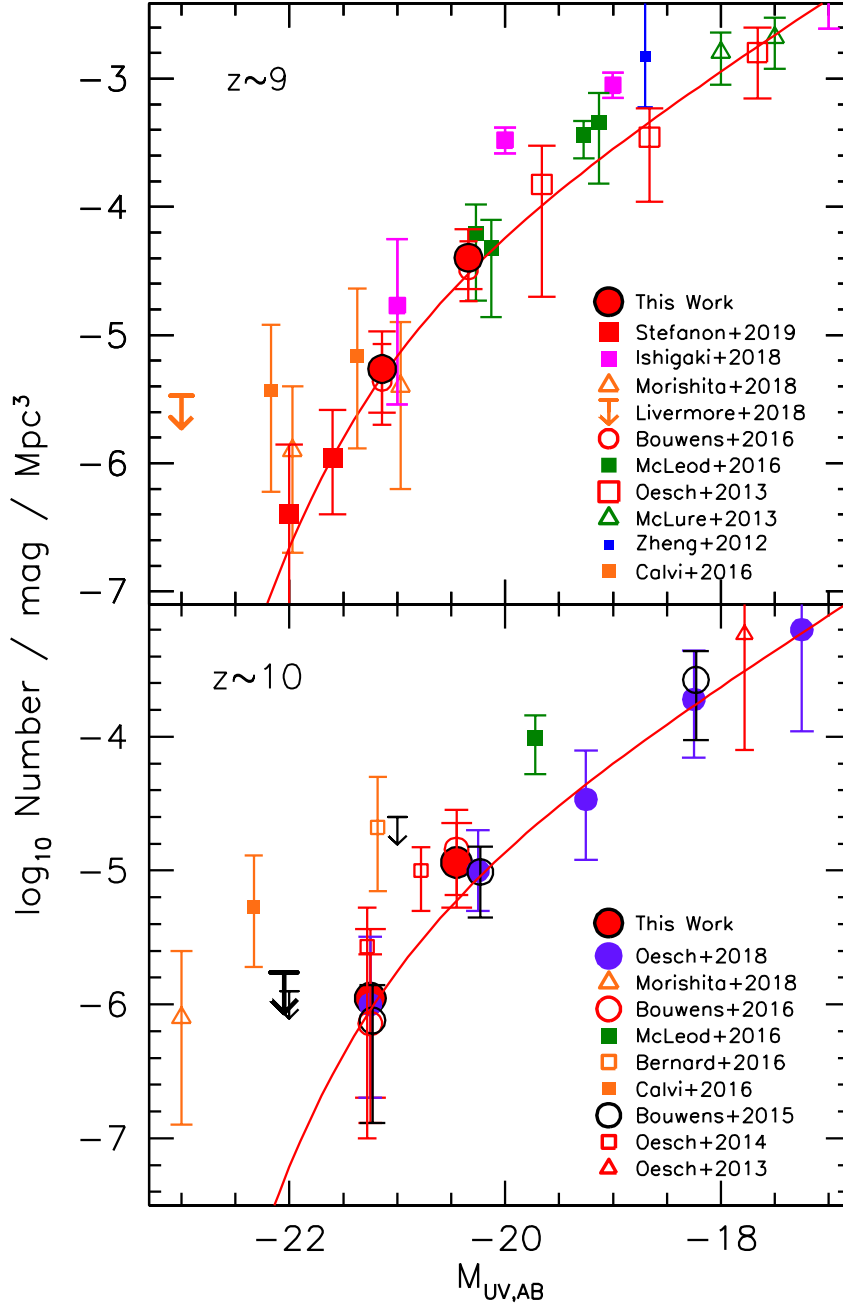


FIG. 8.— Comparison of our $z \sim 9$ and $z \sim 10$ determinations incorporating the present larger search area over CANDELS (*upper and lower panels*), with noteworthy previous determinations of the LF in these same redshift intervals (Zheng et al. 2012; McLure et al. 2013; Oesch et al. 2013, 2014; Bouwens et al. 2015, 2016; McLeod et al. 2016; Oesch et al. 2018; Morishita et al. 2018; Livermore et al. 2018; Ishigaki et al. 2018; Stefanon et al. 2019; §4.2). Two $z \sim 9$ determinations from McLeod et al. (2016) are shown, indicating both their blank-field and lensing field estimates. The red lines show the $z \sim 9$ and $z \sim 10$ LF determinations from Bouwens et al. (2019, in prep) and Oesch et al. (2018), respectively.

$10^{24.23^{+0.11}_{-0.30}}$ ergs/s/Hz/Mpc³ at $z \sim 9$ and $10^{23.89^{+0.16}_{-0.97}}$ erg/s/Hz/Mpc³ at $z \sim 10$. We can compare these luminosity densities with that seen at $z \sim 8$ to the same limiting magnitude, we find $10^{24.91^{+0.06}_{-0.06}}$, using the $z \sim 8$ LF from Bouwens et al. (2015) LF.

As in previous work (e.g., Oesch et al. 2014, 2018), the luminosity density we find at $z \sim 9$, and especially at $z \sim 10$, is much smaller than found at $z \sim 8$, just 100-200 Myr later in cosmic time. The evolution from $z \sim 10$ to $z \sim 8$ is a factor of 10, consistent with estimates from much previous work (e.g., Oesch et al. 2012;

Ellis et al. 2013; Bouwens et al. 2015; Oesch et al. 2018; Ishigaki et al. 2018). The present results are therefore broadly consistent with a rapid evolution in the UV LF results at $z > 8$, as earlier suggested by the “accelerated” evolution scenario of Oesch et al. (2012). Given the small amount of dust extinction that appear to be present in most galaxies at $z \geq 8$ (e.g., Bouwens et al. 2014, 2016), a factor of ten evolution in the luminosity density is equivalent to the star formation rate density evolving by a factor of 10.

4.2. Comparison to Previous $z \sim 9$ -10 Search Results

It is useful for us to compare our new $z \sim 9$ and $z \sim 10$ LF constraints from CANDELS with the substantial number of previous determinations of these LFs (Zheng et al. 2012; McLure et al. 2013; Oesch et al. 2013, 2014; Bouwens et al. 2014b, 2015, 2016; Calvi et al. 2016; McLeod et al. 2016; Bernard et al. 2016; Morishita et al. 2018; Livermore et al. 2018; Stefanon et al. 2019). A comparison of our new results with earlier results is presented in Figure 8. Both the blank-field and the lensing field $z \sim 9$ LF determinations from McLeod et al. (2016) are shown with the green squares.⁹ The red lines in these figures show the $z \sim 9$ and $z \sim 10$ LF determinations from Bouwens et al. (2019, in prep) and Oesch et al. (2018), respectively. The Bouwens et al. (2019, in prep) $z \sim 9$ LF determination leverages both the present $z \sim 9$ search results and those from the Hubble Frontier Field parallels, the HUDF, HUDF09-1, and HUDF09-2.

As the figure illustrates, our results are in reasonable agreement with most previous studies at $z \sim 9$. More scatter is seen in results at $z \sim 10$. Relative to the Bouwens et al. (2016) $z \sim 9$ LF determinations, we infer a $\approx 2\times$ higher volume density of bright $M_{UV,AB} \leq -21.1$ galaxies. This is directly a consequence of the larger number of bright $z \sim 9$ galaxies identified here (Figure 5). At $z \sim 9$, only the Ishigaki et al. (2018) -20 -mag and -19 -mag points lie significantly in excess of the median volume density trend defined by the many different determinations of the LF, i.e., by factors of ~ 6 and ~ 3 .

At $z \sim 10$, there are significant differences between the volume density probes from pure parallel BoRG/HIPPIES observations, e.g., Calvi et al. (2015), Bernard et al. (2016), and Morishita et al. (2018), and those obtained from legacy fields like CANDELS, i.e., Bouwens et al. (2015, 2016) and Oesch et al. (2018). One concern with the $z \sim 10$ candidates identified from the pure-parallel programs is contamination from lower redshift candidates. Not surprisingly, pure-parallel studies like Morishita et al. (2018) – who make use of Spitzer/IRAC data to eliminate lower-redshift contaminants from high-redshift samples – are more in line with results from fields like CANDELS where such multi-wavelength data are available to discriminate against such contaminants.

4.3. Field-to-Field Variations

Finally, with our new expanded sample of bright $z \sim 9$ candidates over the CANDELS fields, we investigate possible field-to-field variations in the volume density of bright sources. We do so, by treating each CANDELS field as independent and using the observed sample from each field to derive the normalization of the UV LF.

For our LF fits, we use the same formalism as described in §4.1, but instead using a Schechter function to create an equivalent binned LF in 0.1 mag intervals. For simplicity, we fix M^* and α to -21.05 and -2.34 , respectively, the values we derive in Bouwens et al. (2019, in prep) from a comprehensive analysis of the LF from $z \sim 2$ to $z \sim 9$, including the HUDF, HUDF parallel

⁹ For simplicity, only the NFW lensing results from McLeod et al. (2016) are presented from the CLASH (Postman et al. 2012) and HFF programs.

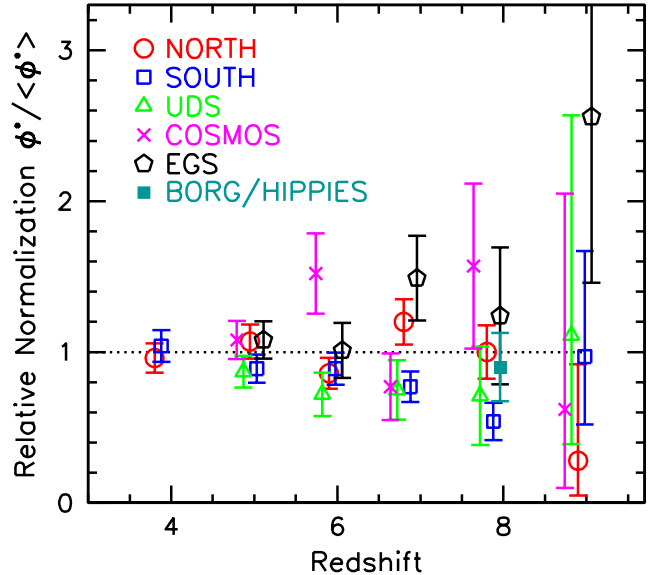


FIG. 9.— Relative normalization inferred for the $z \sim 9$ LF in each of the five CANDELS fields, including GOODS-North (red circles), GOODS-South (blue squares), UDS (green triangles), COSMOS (magenta crosses), and EGS (black pentagons; see §4.3). Both the maximum likelihood determination (solid points) and the 1σ uncertainty (error bars) are shown. Also shown are estimates of the relative normalization of $z \sim 4, 5, 6, 7$, and 8 LFs in each of the same five CANDELS fields, as well as the BoRG/HIPPIES pure-parallel fields at $z \sim 8$ (solid cyan square), as derived in Bouwens et al. (2015).

fields, HFF parallel fields, the ERS fields, all five CANDELS fields, and 50 pure-parallel fields.

We determine the best-fit ϕ^* for each CANDELS field and then compute a relative normalization $\phi^* / \langle \phi^* \rangle$ by comparing the normalization derived for a single CANDELS field with that derived in Bouwens et al. (2019, in prep) considering all CANDELS fields together. Each of the new $z \sim 9$ candidates identified here are treated as $0.5 z \sim 9$ sources to account for the possibility that each may lie at $z < 8$ (which is identical to our treatment of these candidates for our primary LF determinations [§4.1]). We also note that our calculation of the normalization factors for each field requires that we account for both the differing depths and areas of each CANDELS field (as increases in either quantify would increase the total expected number of sources).

The results are presented in Figure 9. Uncertainties are computed based on the Poissonian uncertainties given the small number of $z \sim 9$ galaxies per CANDELS field. The relative normalizations found by Bouwens et al. (2015) for star-forming galaxies at $z \sim 4$ -8 in each of the five CANDELS fields are also shown as a function of redshift.

It is interesting to compute the scatter in the relative normalization and compare it to what is expected from simple Poissonian variations. The RMS scatter in the relative normalizations is 0.67. If we create 10^5 realizations of the CANDELS fields according to the expected number of $z \sim 9$ galaxies per CANDELS field (given their differing depths and areas), the median RMS scatter we compute is 0.56.

While the RMS scatter in the observations is larger than that seen for our median Monte-Carlo simulation,

30% of these simulations give an RMS scatter of similar size or larger than the observed value 0.67. As such, even models with no field-to-field variations are consistent with our observational results.

In summary, we have tried to quantify how the normalization of the $z \sim 9$ UV LF varies from one CANDELS field to another. Unfortunately, the number of $z \sim 9$ candidates per CANDELS field is not sufficiently large to determine this accurately with present data sets.

5. SUMMARY

In this paper, we present new constraints on the bright end of the $z \sim 9$ and $z \sim 10$ LFs based on a search $z \sim 9$ -10 candidate galaxies within a ~ 883 arcmin² area over the five CANDELS fields. The present search includes a 601 arcmin² area over the CANDELS-WIDE UDS, COSMOS, and EGS fields.

The present selection expands on our previous selection of $z \sim 9$ -10 galaxies over these same CANDELS fields (Bouwens et al. 2016) to include an additional ~ 147 arcmin² in search area. We were able to add to our overall selection area within CANDELS by considering those regions which, while having deep WFC3/IR data, did not have deep ACS optical data available from the CANDELS program.

The present selection also considered sources with a broader range of $J_{125} - H_{160}$ colors in our identification of $z \sim 9$ -10 candidate galaxies than in our previous study. Full utilization of the Spitzer/IRAC observations from S-CANDELS (Ashby et al. 2016) and the ground-based optical and near-IR observations is made to refine our selection.

In total, we used the present larger search area to identify three new $z \sim 9$ -10 candidate galaxies. None of these sources were present in our earlier Bouwens et al. (2015, 2016) catalogs or any other published catalogs in the literature. We also identified a fourth candidate with our inclusive selection criteria, which while not identified specifically in our $z \sim 9$ -10 searches, was identified by Roberts-Borsani et al. (2016) using an IRAC [3.6]–[4.5]>0.5 selection designed to pick out bright galaxies at $z > 7$, and which has already been spectroscopically confirmed to lie at $z = 8.68$ (Zitrin et al. 2015).

In creating our expanded $z \sim 9$ samples, we also make use of additional follow-up observations obtained with HST in the Y_{098} -band (GO 14459; Bouwens 2015) of two bright ($H < 25.5$), candidate $z \sim 9$ galaxies identified over the CANDELS fields (Appendix A). These candidates had bluer $J_{125} - H_{160}$ colors than the > 0.5 mag limit we had previously considered. While one candidate is not confirmed to have a redshift of $z > 8$, being well detected in the HST Y_{098} -band data, the other candidate is confirmed to lie at $z > 8$, but with a redshift of 8.3 – too low for inclusion in our $z \sim 9$ selection.

Adding our newly identified $z \sim 9$ candidates to our previous samples (from Bouwens et al. 2016), we identify a total sample of 14 bright $z \sim 9$ galaxy candidates over a ~ 883 arcmin² area in CANDELS. 5 candidate $z \sim 10$ -11 galaxies are found in the same area in CANDELS. This is equivalent to identifying 1 $z \sim 9$ -11 candidate per 47 arcmin² (≈ 10 WFC3/IR fields). Interestingly, our expanded selection of $z \sim 9$ galaxies has UV luminosities which are generally brighter (by 0.1 to 0.4 mag) than in

our previous selection of $z \sim 9$ galaxies (compiled in Bouwens et al. 2016).

In addition to the 19 candidate $z \sim 9$ -10 galaxies we identify over CANDELS that make up our main selection, we also identify 28 mostly lower likelihood candidates (Appendix B). During this process, we consider sources selected from the HST + Spitzer/IRAC $3.6\mu\text{m} + 4.5\mu\text{m}$ data alone in case confusion in the ground-based results in some incompleteness. While a few of those candidates appear to be reliable based on those data, addition of the ground-based constraints show that many are much more likely to be at $z < 4$ (see Figure 6 and 11 from Appendix B). The entire discussion provided in Appendix B and with Figure 6 is useful in illustrating the challenges present in selecting a high-quality sample of $z > 8$ galaxies based on current data sets.

We use this expanded selection of $z \sim 9$ -10 candidate galaxies to refine our determination of the high-luminosity end of the UV LF at $z \sim 9$ and $z \sim 10$. Our revised determinations show a $\approx 2\times$ higher volume density of bright ($M_{UV,AB} \leq -21.1$) $z \sim 9$ galaxies than found by Bouwens et al. (2016). This owes to the increased fraction of bright ($m_{AB} \leq 26.1$) $z \sim 9$ galaxies identified in the new area we probe (Figure 5).

By comparing the number of bright $z \sim 9$ galaxies identified with the number expected, we attempted to estimate the relative volume density of $z \sim 9$ galaxies per CANDELS field. The RMS variation we found was in excess of that expected from Poissonian statistics. Nevertheless, we found that the observed scatter was not especially significant and we could reproduce it, adopting simple Poissonian statistics, in as many as $\sim 30\%$ of our Monte-Carlo trials. To quantify this better, clearly deeper observations are required over all 5 CANDELS fields to identify a larger number of $z \sim 9$ galaxies.

With our new results, we confirm the strong evolution seen in the UV LF at $z > 8$ in previous work (Oesch et al. 2012, 2014; Bouwens et al. 2015, 2016; Oesch et al. 2018), with a factor of 10 evolution from $z \sim 10$ to $z \sim 8$. The present results are broadly consistent with the “accelerated” evolution scenario suggested by Oesch et al. (2012).

Better constraints on the volume density of bright $z \sim 9$ galaxies could be obtained by continuing our exploitation of wide-area VISTA surveys as have been conducted by Bowler et al. (2014), Stefanon et al. (2017), and Stefanon et al. (2019), by surveying much wider area fields to faint magnitudes with HST, by improving our exploitation of the archival and pure parallel HST + Spitzer/IRAC data (e.g., Morishita et al. 2018), and in the future with JWST, Euclid, and WFIRST.

This work is based on observations taken by the CANDELS Multi-Cycle Treasury Program, the 3D-HST Treasury Program (GO 12177 and 12328), and the z9-CANDELS program (GO 13792) with the NASA/ESA HST, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS5-26555. This work is based on observations made with the Spitzer Space Telescope, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. Support for this work was provided by NASA through an

award issued by JPL/Caltech. This work is also based on data products from observations made with ESO Telescopes at the La Silla Paranal Observatory under ESO programme ID 179.A-2005 and on data products produced by TERAPIX and the Cambridge Astronomy Survey Unit on behalf of the UltraVISTA consortium.

We appreciate the feedback on our analysis from an anonymous expert referee whose feedback significantly improved our manuscript. We acknowledge the support of NASA grants HST-AR-13252, HST-GO-13872, HST-GO-13792, and NWO grants 600.065.140.11N211 (vrij competitie) and TOP grant TOP1.16.057.

REFERENCES

- Aihara, H., Arimoto, N., Armstrong, R., et al. 2018a, PASJ, 70, S4
- Aihara, H., Armstrong, R., Bickerton, S., et al. 2018b, PASJ, 70, S8
- Anders, P., & Fritze-v. Alvensleben, U. 2003, A&A, 401, 1063
- Ashby, M. L. N., Willner, S. P., Fazio, G. G., et al. 2013, ApJ, 769, 80
- Ashby, M. L. N., Willner, S. P., Fazio, G. G., et al. 2015, ApJS, 218, 33
- Bernard, S. R., Carrasco, D., Trenti, M., et al. 2016, ApJ, 827, 76
- Bertin, E. and Arnouts, S. 1996, A&AS, 117, 39
- Bielby, R., Hudelot, P., McCracken, H. J., et al. 2012, A&A, 545, AA23
- Bouwens, R., Broadhurst, T. and Silk, J. 1998, ApJ, 506, 557
- Bouwens, R., Broadhurst, T., & Illingworth, G. 2003a, ApJ, 593, 640
- Bouwens, R. J., Illingworth, G. D., Franx, M., & Ford, H. 2007, ApJ, 670, 928
- Bouwens, R. J., Illingworth, G. D., Franx, M., & Ford, H. 2008, ApJ, 686, 230
- Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2010, ApJ, 709, L133
- Bouwens, R. J., Illingworth, G. D., Labbe, I., et al. 2011, Nature, 469, 504
- Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2012, ApJ, 754, 83
- Bouwens, R. 2014, HST Proposal, 13792
- Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2014a, ApJ, 793, 115
- Bouwens, R. J., Bradley, L., Zitrin, A., et al. 2014b, ApJ, 795, 126
- Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2015, ApJ, 803, 34
- Bouwens, R. 2015, HST Proposal, 14459
- Bouwens, R. J., Oesch, P. A., Labbé, I., et al. 2016, ApJ, 830, 67
- Bouwens, R. J., Oesch, P. A., Illingworth, G. D., Ellis, R. S., & Stefanon, M. 2017, ApJ, 843, 129
- Bowler, R. A. A., Dunlop, J. S., McLure, R. J., et al. 2014, MNRAS, 440, 2810
- Bowler, R. A. A., Dunlop, J. S., McLure, R. J., & McLeod, D. J. 2017, MNRAS, 466, 3612
- Brammer, G. B., van Dokkum, P. G., & Coppi, P. 2008, ApJ, 686, 1503
- Calvi, V., Trenti, M., Stiavelli, M., et al. 2016, ApJ, 817, 120
- Capak, P., Aussel, H., Ajiki, M., et al. 2007, ApJS, 172, 99
- Castellano, M., Fontana, A., Grazian, A., et al. 2012, A&A, 540, A39
- Coe, D., Zitrin, A., Carrasco, M., et al. 2013, ApJ, 762, 32
- Dressel, L., et al. 2012. Wide Field Camera 3 Instrument Handbook, Version 5.0 (Baltimore: STScI)
- Ellis, R. S., McLure, R. J., Dunlop, J. S., et al. 2013, ApJ, 763, L7
- Finkelstein, S. L., Papovich, C., Salmon, B., et al. 2012, ApJ, 756, 164
- Furusawa, H., Kosugi, G., Akiyama, M., et al. 2008, ApJS, 176, 1
- Grogin, N. A., Kocevski, D. D., Faber, S. M., et al. 2011, ApJS, 197, 35
- Hashimoto, T., Laporte, N., Mawatari, K., et al. 2018, Nature, 557, 392
- Holwerda, B. W., Bouwens, R., Oesch, P., et al. 2015, ApJ, 808, 6
- Ishigaki, M., Kawamata, R., Ouchi, M., et al. 2018, ApJ, 854, 73
- Koekemoer, A. M., Aussel, H., Calzetti, D., et al. 2007, ApJS, 172, 196
- Koekemoer, A. M., Faber, S. M., Ferguson, H. C., et al. 2011, ApJS, 197, 36
- Kotulla, R., Fritze, U., Weilbacher, P., & Anders, P. 2009, MNRAS, 396, 462
- Kron, R. G. 1980, ApJS, 43, 305
- Labbé, I., González, V., Bouwens, R. J., et al. 2010a, ApJ, 708, L26
- Labbé, I., González, V., Bouwens, R. J., et al. 2010b, ApJ, 716, L103
- Labbé, I., Oesch, P. A., Bouwens, R. J., et al. 2013, ApJ, 777, L19
- Labbé, I., Oesch, P. A., Illingworth, G. D., et al. 2015, ApJS, 221, 23
- Lam, D., Bouwens, R. J., Coe, D., et al. 2019, ApJ, submitted, arXiv:1903.08177
- Livermore, R. C., Trenti, M., Bradley, L. D., et al. 2018, ApJ, 861, L17
- Magee, D. K., Bouwens, R. J., & Illingworth, G. D. 2011, Astronomical Data Analysis Software and Systems XX, 442, 395
- McCracken, H. J., Capak, P., Salvato, M., et al. 2010, ApJ, 708, 202
- McCracken, H. J., Milvang-Jensen, B., Dunlop, J., et al. 2012, A&A, 544, A156
- McLeod, D. J., McLure, R. J., Dunlop, J. S., et al. 2015, MNRAS, 450, 3032
- McLeod, D. J., McLure, R. J., & Dunlop, J. S. 2016, MNRAS, 459, 3812
- McLure, R. J., Dunlop, J. S., Bowler, R. A. A., et al. 2013, MNRAS, 432, 2696
- Morishita, T., Trenti, M., Stiavelli, M., et al. 2018, ApJ, 867, 150
- Mutch, S. J., Liu, C., Poole, G. B., et al. 2016, MNRAS, 463, 3556
- Oesch, P. A., Bouwens, R. J., Illingworth, G. D., et al. 2012a, ApJ, 745, 110
- Oesch, P. A., Bouwens, R. J., Illingworth, G. D., et al. 2013, ApJ, 773, 75
- Oesch, P. A., Bouwens, R. J., Illingworth, G. D., et al. 2014, ApJ, 786, 108
- Oesch, P. A., Brammer, G., van Dokkum, P. G., et al. 2016, ApJ, 819, 129
- Oesch, P. A., Bouwens, R. J., Illingworth, G. D., Labbé, I., & Stefanon, M. 2018, ApJ, 855, 105
- Oke, J. B., & Gunn, J. E. 1983, ApJ, 266, 713
- Ono, Y., Ouchi, M., Curtis-Lake, E., et al. 2013, ApJ, 777, 155
- Postman, M., Coe, D., Benítez, N., et al. 2012, ApJS, 199, 25
- Roberts-Borsani, G. W., Bouwens, R. J., Oesch, P. A., et al. 2016, ApJ, 823, 143
- Rogers, A. B., McLure, R. J., Dunlop, J. S., et al. 2014, MNRAS, 440, 3714
- Scoville, N., Abraham, R. G., Aussel, H., et al. 2007, ApJS, 172, 38
- Shibuya, T., Ouchi, M., & Harikane, Y. 2015, ApJS, 219, 15
- Smit, R., Bouwens, R. J., Carniani, S., et al. 2018, Nature, 553, 178
- Stark, D. P., Ellis, R. S., Chiu, K., Ouchi, M., & Bunker, A. 2010, MNRAS, 408, 1628
- Stefanon, M., Labbé, I., Bouwens, R. J., et al. 2017, ApJ, 851, 43
- Stefanon, M., Labbé, I., Bouwens, R. J., et al. 2019, ApJ, submitted, arXiv:1902.10713
- Steidel, C. C., Giavalisco, M., Pettini, M., Dickinson, M., & Adelberger, K. L. 1996, ApJ, 462, L17
- Steidel, C. C., Adelberger, K. L., Shapley, A. E., et al. 2003, ApJ, 592, 728
- Straatman, C. M. S., Spitler, L. R., Quadri, R. F., et al. 2016, ApJ, 830, 51
- Trenti, M., Bradley, L. D., Stiavelli, M., et al. 2011, ApJ, 727, L39
- Vanzella, E., Giavalisco, M., Dickinson, M., et al. 2009, ApJ, 695, 1163
- Waters, D., Di Matteo, T., Feng, Y., Wilkins, S. M., & Croft, R. A. C. 2016, MNRAS, 463, 3520
- Watson, D., Christensen, L., Knudsen, K. K., et al. 2015, Nature, 519, 327
- Willott, C. J., McLure, R. J., Hibon, P., et al. 2013, AJ, 145, 4

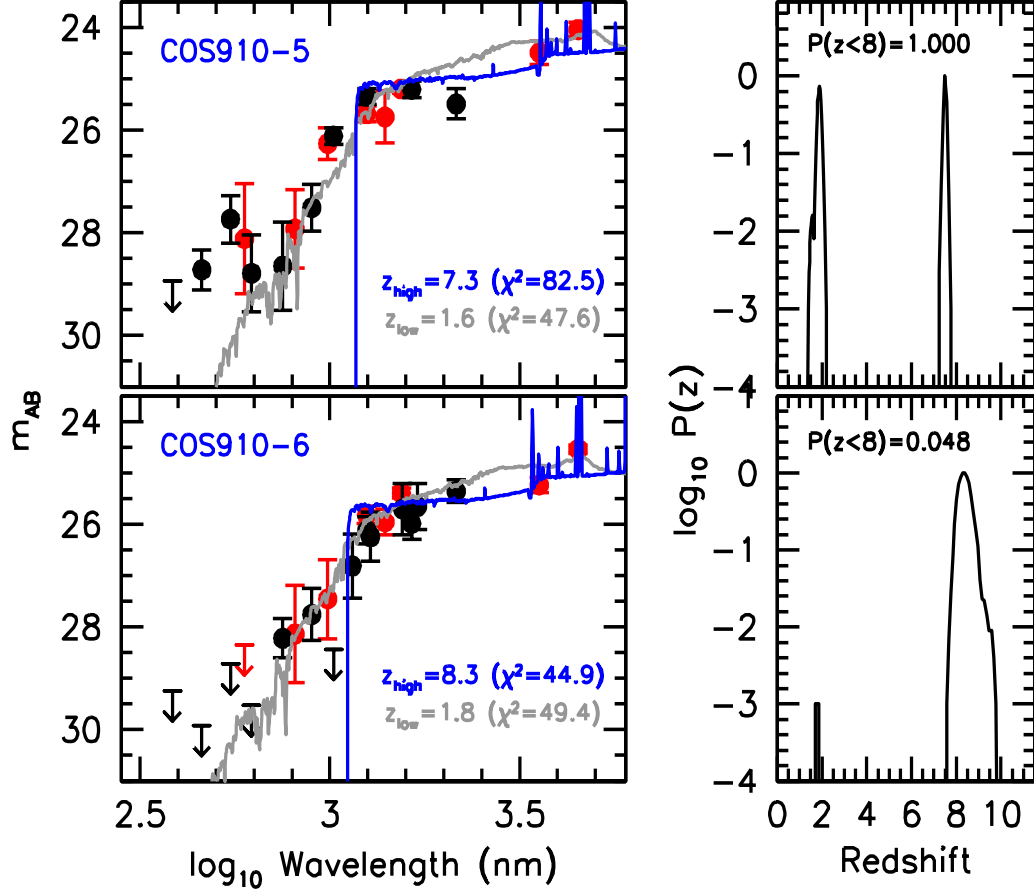


FIG. 10.— (*left panels*) Measured spectral energy distributions for the two candidate $z \sim 9$ galaxies that we identified over the CANDELS fields which had measured $J_{125} - H_{160}$ colors between 0.4 and 0.5 mag and which showed a $>50\%$ probability for having a redshift $z > 8$ using the data that were available in 2015. The black solid circles show ground-based flux constraints (with 1σ error bars plotted), while the red points show the flux constraints from HST and Spitzer/IRAC. Targeted observations of the two candidates have been obtained with HST in the Y_{098} band based on a mid-cycle program (1 orbit each). New observations are also available as a result of the Hyper-Suprime-Cam Ultra-deep observations over the COSMOS field and deeper near-IR observations from UltraVISTA (DR3). The blue and gray lines show the best-fit $z > 6$ and $z < 6$ model fits, respectively, to our measured photometric constraints for the candidates. (*right panels*) The redshift likelihood distributions we derive based on our photometric constraints for the two $z \sim 9$ candidates. While the first candidate COS910-5 shows roughly an equal likelihood of being at $z \sim 1.6$ or $z \sim 7.3$, the redshift of the second candidate COS910-6 appears to be robustly $z \sim 8.3$.

Wilkins, S. M., Bouwens, R. J., Oesch, P. A., et al. 2016, MNRAS, 455, 659
 Windhorst, R. A., Cohen, S. H., Hathi, N. P., et al. 2011, ApJS, 193, 27

Yan, H., Yan, L., Zamojski, M. A., et al. 2011, ApJ, 728, L22
 Zheng, W., Postman, M., Zitrin, A., et al. 2012, Nature, 489, 406
 Zitrin, A., Zheng, W., Broadhurst, T., et al. 2014, ApJ, 793, L12
 Zitrin, A., Labbé, I., Belli, S., et al. 2015, ApJ, 810, L12

APPENDIX

A. NATURE OF TWO CANDIDATE $Z \sim 9$ GALAXIES TARGETED BY HST FOLLOW-UP PROGRAM 14459

Here we present constraints on the spectral energy distributions and redshift likelihood distribution we derived for the two candidate $z \sim 9$ galaxies we targeted with an HST follow-up program in cycle 23 (GO 14459: Bouwens 2015). The sources were selected for follow-up based on their exceptional brightness ($H_{160,AB} < 25.5$) and $>50\%$ probability of lying at $z > 8$ (relying on our earlier photometry). Constraints on the redshifts of the targeted candidates were derived based on the same HST+Spitzer/IRAC+ground-based observations as utilized by Bouwens et al. (2016). We designate them COS910-5 and COS910-6, and they have coordinates of 10:00:31.39, 02:26:39.8 and 10:00:20.12, 02:14:13.0, respectively. These candidates had $H_{160,AB}$ magnitudes of 25.1 mag and 25.3 mag, respectively. The $J_{125} - H_{160}$ colors of both sources were bluer than the >0.5 mag limit we had earlier used in selecting sources for follow-up (Bouwens et al. 2016).

As both of the targetted $z \sim 9$ candidates were found over the CANDELS COSMOS field, we made use of the same imaging data sets and photometry as we used in selecting them. In addition to those photometric constraints, we also made use of the zy imaging observations associated with Hyper Suprime-Cam (HSC) Ultra-Deep survey (DR1: Aihara et al. 2018a,b). The photometric redshift software EAZY (Brammer et al. 2008) is applied to the full set

TABLE 4
A SMALL SAMPLE OF PROBABLE^a $z \sim 8.0$ - 8.4 GALAXIES IDENTIFIED OVER THE
CANDELS UDS, COSMOS, AND EGS PROGRAMS^b

ID	R.A.	Dec	$H_{160,AB}$	z_{phot}	$P(z > 8)$	$P(z > 7)$
COS910-6	10:00:20.12	2:14:13.0	25.3 ± 0.1	8.3	0.87 ^c	0.98
EGS910-11	14:19:59.71	52:51:19.5	26.4 ± 0.2	8.4	0.55	0.93
EGS910-12	14:20:19.08	53:03:14.3	25.9 ± 0.1	8.1	0.50	0.90

^a $P(z > 8) > 0.5$

^b See Appendix A and B

^c See Figure 10 and Appendix A.

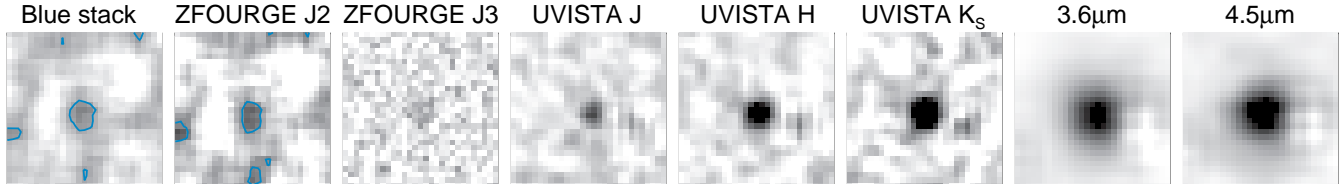


FIG. 11.— Illustration of the neighbor-subtracted postage stamp images ($5'' \times 5''$) of one source COS910-8 (10:00:34.99, 02:14:01.1) in the ZFOURGE $J2$ and $J3$ bands (2nd + 3rd leftmost panels), VISTA J , H , and K_s bands (middle panels), Spitzer/IRAC $3.6\mu\text{m}$ and $4.5\mu\text{m}$ bands (rightmost panels), and in a stack of the imaging data blueward of $1.25\mu\text{m}$ (leftmost panels). The blue stack weights the $<1.25\mu\text{m}$ ground-based images assuming COS910-8 is a $z \sim 2$ galaxy. The blue contours in the leftmost two panels indicate regions detected at $>2\sigma$ significance. COS910-8 seemed likely to be at $z > 9$ based on the HST+IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ photometry, but for which the ground-based photometry indicates is more likely at $z \sim 2$ (see Figure 6). The presented stamps should make it clear that the neighboring source (separated by $1.5''$ [mostly to the west, i.e., on the right-hand side] and with a H -band flux of 22.6 mag) is well subtracted. The apparent detection of this source both in the FourStar $J2$ band from ZFOURGE and also in the other bluer imaging data strongly suggests that COS910-8 is not at $z > 9$.

of flux measurements we have for the candidates to quantify the redshift likelihood distribution. The best-fit high ($z > 6$) and low-redshift ($z < 6$) SEDs for the sources are shown in Figure 10 with the blue and gray lines, respectively. The best-fit redshifts we derive for COS910-5 and COS910-6 are 7.3 and 8.3, respectively, with $P(z < 8) = 1$ and $P(z < 8) = 0.05$. Our first candidate COS910-5 is plausibly a lower redshift ($z < 2$) galaxy based on its photometry, while the photometry of COS910-6 securely places it at $z > 7$.

B. OTHER POTENTIAL $Z \sim 9$ - 10 CANDIDATES

In addition to the high-likelihood candidate $z \sim 9$ - 10 galaxies presented in the body of this manuscript, we also identified other sources which also plausibly correspond to $z \sim 9$ - 10 galaxies. However, because the integrated probability of these sources lying in excess of 8.0 was below 50% or because their maximum likelihood redshift was below 8.4, we excluded them from our primary sample.

In an effort to be comprehensive, we include in our compilation even bright sources which showed at least a 20% probability of lying at $z > 8$. We furthermore relaxed the $H - [3.6] < 1.4$ mag criterion used for our main selection to include sources redder than 1.4 mag, but as we discuss in Appendix B.2 and illustrate in Figure 12, this does not appear to add any probable $z \sim 9$ - 10 candidates.

Table 4 provides a list of sources where the best estimate redshift is between 8.0 and 8.4, with a $>50\%$ probability of lying at $z > 8$. Table 5 features a list of sources where there is a $>20\%$ probability of lying at $z > 8$.

For completeness, we have also included in this table sources which were previously listed in Table 7 and Appendix C of Bouwens et al. (2016).

B.1 Possible $z \sim 9$ - 10 Candidates Identified By Down-Weighting the Ground-Based Data

A fraction of the sources ($\sim 30\%$) found in the HST data are very close by other sources in the ground-based imaging observations. While our photometric software MOPHONGO copes with source overlap (subtracting flux from nearby neighbors before doing photometry), subtractions are not perfect in all cases and this can cause some fraction of bona-fide high-redshift sources to be missed due e.g. to imperfectly subtracted optical flux from neighboring sources. Based on the selection volume simulations run in §4.1, we estimate this incompleteness to be $\sim 30\%$.

Given these incompleteness levels in HST+Spitzer+ground-based $z \sim 9$ - 10 galaxy searches, we repeated our $z \sim 9$ - 10 galaxy searches using the same procedure as in the main text but increasing the uncertainties on the ground-based data by a factor of 10 (and hence significantly downweighting those data). Those candidates are included in Table 5 along with their best-fit redshifts and integrated probabilities of lying at $z > 7$ and $z > 8$. For comparison, we include in the same table the maximum likelihood redshifts and the integrated probabilities of sources lying at $z > 7$ and $z > 8$ when including all the data in the middle columns. This is to show the impact of incorporating the photometric constraints available from the ground-based Y , $J1$, $J2$, $J3$, J , H_s , H , H_l , and K band observations. The optical photometric constraints are also included in the results presented in the middle column, but the photometric uncertainties there are multiplied by a factor of 2 to allow for the possibility of imperfectly subtracted neighbors. Sources with best-fit redshifts $z > 8$ and where the integrated probability at $z > 8$ is $>20\%$ are also included in

TABLE 5
 POSSIBLE^a $z \sim 9-10$ GALAXIES IDENTIFIED OVER THE CANDELS UDS, COSMOS, AND EGS PROGRAMS^b

ID	R.A.	Dec	$H_{160,AB}$	HST+IRAC+Ground ^c			HST+IRAC [3.6]/[4.5]		
				z_{phot}	$P(z > 8)$	$P(z > 7)$	z_{phot}	$P(z > 8)$	$P(z > 7)$
From Appendix C and Table 7 of Bouwens et al. (2016)									
UDS910-2	02:17:13.08	-05:15:55.4	26.6±0.2	10.2	0.68	0.91	—	—	—
UDS910-3	02:17:52.38	-05:15:06.3	26.9±0.2	9.4	0.28	0.60	—	—	—
UDS910-4	02:17:14.61	-05:15:15.7	26.6±0.2	9.1	0.50	0.70	—	—	—
Candidates where $P(z > 8) < 0.5$									
COS910-7	10:00:26.15	02:32:49.8	25.6±0.1	7.4	0.31	1.00	—	—	—
UDS910-5	02:18:20.45	-05:14:07.9	25.9±0.2	8.8	0.46	0.70	—	—	—
EGS910-13	14:20:45.23	53:02:01.4	25.9±0.1	8.5	0.40	0.95	—	—	—
EGS910-14	14:20:05.09	52:58:02.6	25.7±0.1	8.3	0.38	0.77	—	—	—
EGS910-15	14:19:52.22	52:55:58.8	26.1±0.1	8.3	0.38	0.83	—	—	—
EGS910-16 ^d	14:20:47.81	53:02:11.8	25.6±0.1	8.6	0.15	0.23	—	—	—
Candidates Found Downweighting Constraints from the Ground-Based Photometry [†] – Some Foreground Confusion ^e									
COS910-8 [*]	10:00:34.99	02:14:01.1	24.5±0.1	2.1	0.00 [*]	0.00 [*]	9.4	0.90	0.90
COS910-9 ^{**}	10:00:43.84	02:13:50.4	25.6±0.1	1.8	0.03	0.03	9.2	0.63	0.66
COS910-10 ^f	10:00:24.22	02:17:56.5	24.4±0.1	— ^f	—	—	9.2	0.55	0.55
UDS910-6 ^{***}	02:18:00.63	-05:14:43.2	25.7±0.1	3.7	0.00	0.00	8.9	0.76	0.93
UDS910-7 [‡]	02:17:13.13	-05:13:01.6	24.8±0.1	0.6	0.00	0.00	9.2	0.61	0.62
UDS910-8	02:17:25.62	-05:09:38.8	26.1±0.3	9.5	0.31	0.47	9.1	0.56	0.77
UDS910-9	02:17:39.67	-05:14:10.8	25.6±0.1	2.4	0.38	0.38	9.3	0.43	0.46
EGS910-17	14:20:24.88	53:02:35.0	26.5±0.1	9.2	0.55	0.73	9.2	0.60	0.78
EGS910-18	14:20:38.81	53:03:58.1	26.3±0.2	1.6	0.21	0.27	8.6	0.39	0.63
Candidates Found Downweighting Constraints from the Ground-Based Photometry [†] – Minimal Foreground Confusion ^g									
COS910-11 ^{****}	10:00:14.78	02:18:09.6	26.2±0.1	1.6	0.02	0.15	9.0	0.40	0.45
COS910-12	10:00:24.18	02:26:57.2	25.6±0.2	8.6	0.52	0.63	8.6	0.43	0.65
UDS910-10 ^{**}	02:17:38.84	-05:15:29.6	26.1±0.1	1.4	0.00	0.26	9.4	0.57	0.60
UDS910-11	02:17:07.36	-05:11:49.0	26.3±0.1	9.0	0.50	0.54	8.8	0.34	0.45
UDS910-12 ^{***}	02:17:01.03	-05:11:00.4	24.2±0.1	2.3	0.01	0.01	9.2	0.61	0.61
UDS910-13 ^{*****}	02:17:01.36	-05:09:59.6	25.3±0.1	2.3	0.00 ^{*****}	0.00 ^{*****}	9.7	0.68	0.68
UDS910-14	02:17:49.19	-05:15:16.7	26.9±0.2	9.5	0.68	0.73	9.5	0.27	0.29

^a $P(z > 8) > 0.2$

^b See Appendix B

^c In computing the photometric solution for the 12 lowest sources in the table, we downweighted the optical photometric constraints somewhat, by multiplying their uncertainties by a factor of 2. This was done to allow for imperfect subtraction of the optical flux from sources neighboring the indicated sources and also not to overly select against sources that show occasional $\sim 2\sigma$ detections in the optical.

^d Source has a very red $H_{160} - [3.6]$ color of 1.9 mag. Given what has been found for other sources (Figure 12), it seems very likely that this source is at $z < 4$.

^e Sources where subtraction of neighboring sources impacts the ground-based optical flux measurements by more than 1σ in some bands.

^f Ground-based photometry is substantially impacted by a bright nearby neighbor, with subtracted flux from our photometric apertures ($1.2''$ -diameter) 15-20 \times greater than the 1σ noise in many bands. As such, deriving accurate photometry for this source is very challenging. In any case, this source seems consistent with being a $z < 4$ dusty source given that its $H - [3.6]$ color is ~ 1.95 mag (see Figure 12).

^g Sources where subtraction of neighboring sources impacts the ground-based optical flux measurements by less than 1σ in all bands.

[†] Source selected if formal uncertainties on the ground-based flux measurements is increased by a factor of 10. Because of the proximity of some HST-detected sources to nearby neighbors, obtaining accurate flux measurements for some sources is challenging. This can result in some bona-fide $z \sim 9-10$ galaxies being missed, when utilizing the ground-based photometry to determine the likely redshift of a source.

[‡] This source is detected at 1.9σ and 5σ in the I_{814} and Y bands in our photometry.

^{*} This source is detected at $\sim 3\sigma$ in our ground-based FourStar $J2$ photometry of ZFOURGE (Straatman et al. 2016: see Figure 11). As a result of this and the overall shape of the SED as defined by the existing observations in the H , K_s , $[3.6]$, $[4.5]$, $[5.8]$, and $[8.0]$ bands, this source seems very unlikely to be at $z > 9$.

^{**} This source is detected at 3σ in our Y -band ground-based photometry

^{***} This source is detected at 2σ in our Y -band ground-based photometry

^{****} This source shows a $>3\sigma$ detection at $1.1\mu\text{m}$ combining the Y , $J1$ and $J2$ band imaging data.

^{*****} This source is detected at $\sim 4\sigma$ in the ground-based FourStar $J2$ -band data from ZFOURGE. Clearly, this source is not likely to be at $z > 9$.

TABLE 6
PHOTOMETRY OF COS910-8^a

Band	Measured Flux (nJy)
CFHTLS u^*	16 ± 8
SSC B	-11 ± 6
HSC g	12 ± 9
CFHTLS g	-10 ± 7
SSC V	2 ± 12
HST V_{606}	-7 ± 16
HSC r	7 ± 8
CFHTLS r	-8 ± 9
SSC r^+	-1 ± 12
SSC i^+	3 ± 16
CFHTLS y	19 ± 13
CFHTLS i	7 ± 10
HSC i	24 ± 12
HST I_{814}	1 ± 16
CFHTLS z	-17 ± 25
HSC z	44 ± 18
SSC z^+	37 ± 50
HSC y	27 ± 43
UVISTA Y	35 ± 29
ZFOURGE $J1$	12 ± 43
ZFOURGE $J2$	152 ± 46 [†]
HST J_{125}	242 ± 24
UVISTA J	222 ± 33
ZFOURGE $J3$	408 ± 63
HST JH_{140}	520 ± 32 [‡]
HST H_{160}	591 ± 20
ZFOURGE H_s	587 ± 119
UVISTA H	625 ± 42
ZFOURGE H_i	601 ± 115
ZFOURGE K_s	1101 ± 98
UVISTA K_s	1063 ± 55
IRAC $3.6\mu\text{m}$	1645 ± 78
IRAC $4.5\mu\text{m}$	1846 ± 61
IRAC $5.8\mu\text{m}$	-3971 ± 1704 [*]
IRAC $8.0\mu\text{m}$	-2612 ± 2363 [*]

^a Figure 6 shows a fit of the SED of this source to $z \sim 2.1$ galaxy and a $z \sim 9.4$ galaxy, while Figure 11 shows postage stamp images of this source.

[†] No signal is expected in the FourStar $J2$ band from ZFOURGE (which runs from $1.07\mu\text{m}$ to $1.22\mu\text{m}$) if COS910-8 is a $z > 9$ source. Nonetheless, COS910-8 shows a 3.3σ detection in the $J2$ band in our photometry (see second leftmost panel in Figure 11).

[‡] The JH_{140} -band flux we estimate correcting JH_{140} flux measurements made in $0.4''$ -diameter apertures to total is 469 ± 29 nJy.

^{*} A 2σ detection of COS910-8 would be expected at both $5.8\mu\text{m}$ and $8.0\mu\text{m}$, if it was a $z \sim 9.4$ galaxy. By contrast, if the source is instead at $z \sim 2$, no detection is expected in either band due to a turn-over in the SED at $1.6\mu\text{m}$ rest-frame. COS910-8 is not detected in either channel.

Table 5.

B.2 Case of $z \sim 9$ Candidate COS910-8

There are a few sources in Table 5 which seem like reliable $z \sim 9$ -10 candidate galaxies when using only the HST+Spitzer/IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ observations, but which clearly prefer a $z < 4$ solution when including the constraints from the ground-based photometry. In Figure 6 from the main text, we showed one example of this type of source which prefer a $z > 9$ solution using the HST+ $3.6\mu\text{m}+4.5\mu\text{m}$ observations alone. Because of the brightness of that source and to ensure that we were not deriving a low-redshift solution due to imperfect subtraction of a bright nearby source, we took unusual care in subtracting a bright neighboring source before performing our flux measurements ($1.2''$ -diameter apertures). Postage stamp images of this source, after subtraction of a neighboring source, are shown in Figure 11. The leftmost image in this panel show a weighted stack of all the data from the U -band to the FourStar $J2$ band, weighted according to expected contribution in each band if the source were a $z \sim 2$ galaxy.

Particularly concerning for COS910-8 being a $z > 9$ source is the apparent detection of this source at 3σ in the FourStar $J2$ band at the position of the source, tentative detections of the source in the i , z , and Y bands (14 ± 6 nJy, 24 ± 14 nJy, and 27 ± 21 nJy, respectively), and the overall SED shape as defined by the available H , K_s , $[3.6]$, $[4.5]$, $[5.8]$, and $[8.0]$ -band photometry. In particular, the measured K_s -band flux from both UltraVISTA and FourStar/ZFOURGE

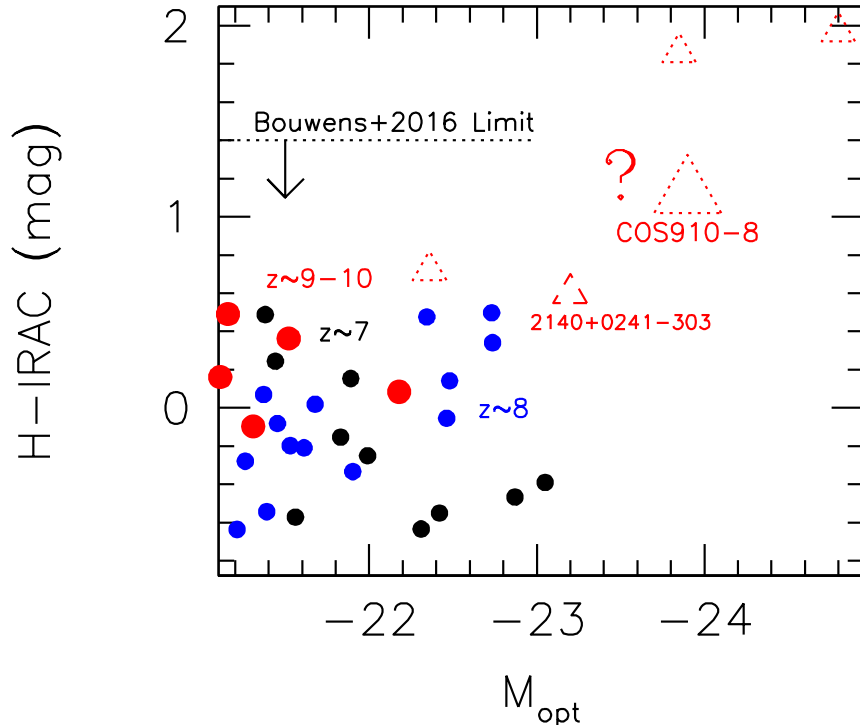


FIG. 12.— $H - IRAC$ colors of candidate $z \sim 7-10$ Galaxies vs. rest-optical luminosities M_{opt} . The brightest $z \sim 7$ and $z \sim 8$ sources in the $>1.5 \text{ deg}^2$ selections of Bowler et al. (2017) and Stefanon et al. (2019) are indicated by the solid black and blue circles, respectively. $z \sim 9-10$ sources (solid red circles) are from Oesch et al. (2014) and Bouwens et al. (2016). The dashed open triangle is a $z \sim 10$ candidate from Morishita et al. (2018). The open dotted red triangles show the colors and nominal luminosities of candidate $z \sim 9-10$ galaxies identified using the HST+Spitzer/IRAC $3.6\mu\text{m}+4.5\mu\text{m}$ data alone and which nominally have fairly extreme properties relative to $z \sim 7-8$ samples. Given that we would tend to expect very wide-area searches such as performed by Bowler et al. (2017) and Stefanon et al. (2019) to yield sources with much more extreme properties than identified over surveys like CANDELS (where a $6\times$ smaller area is available) and given the expectations that galaxies at $z \sim 7-8$ are brighter and redder than at $z \sim 9-10$, it is odd to find candidates with even more extreme properties over CANDELS at $z > 9$. We might thus expect most of the extreme $z \gtrsim 9$ candidates over CANDELS to actually be at $z < 4$; in fact, all of the extreme candidates identified in Appendix B (Table 5) appear to revert to $z < 4$ solutions after incorporating constraints from existing ground-based observations. The limiting $H_{160} - [3.6]$ color for the Bouwens et al. (2016) selection is indicated by the horizontal dotted line. It should be apparent that this limit easily allows for the selection of every bright $z \sim 7-9$ galaxy identified over square degree fields by Bowler et al. (2017) and Stefanon et al. (2019) and also the $z \sim 9-10$ candidates from Oesch et al. (2014).

lie in excess of expectations if the source is at $z > 9$, while our $5.8\mu\text{m}$ and $8.0\mu\text{m}$ flux constraints lie below expectations if the source is at $z > 9$. Including the photometric constraints from all HST, Spitzer, and ground-based data (Table 6), we estimate that the source has a 2×10^{-6} probability of lying at $z > 9$. If slightly larger photometric uncertainties are assumed in all bands (as needed to derive a reduced χ^2 of 1), the integrated likelihood that the source is at $z > 9$ is 2×10^{-4} .

Besides COS910-8, there are 7 other candidates in Table 5, which strongly prefer a low-redshift solution after incorporating the constraints from ground-based observations, mostly as a result of a detection in a band just blueward of $1.2\mu\text{m}$, i.e., Y, J1, or J2. Given the impact of the optical data on the computed $P(z > 8)$'s, clearly it is essential to carefully examine the ground-based optical and near-IR observations when selecting the best $z > 8$ galaxies for further follow-up observations.

Even without a detailed consideration of a vast array of ground-based observations to determine the nature of COS910-8 and a number of similar sources within CANDELS, it is clear that we might have expected most of these sources to correspond to $z < 4$ galaxies given the very high implied luminosities and red $H - IRAC$ colors (treating them as $z > 9$ candidates). To illustrate how extreme e.g. COS910-8 would be as a $z > 9$ candidate, we present its optical luminosity and $H - [3.6]$ color (assuming a $z \sim 9.4$ redshift: red open dotted triangle) relative to the $z \sim 7$ and $z \sim 8-9$ samples identified by Bowler et al. (2017) and Stefanon et al. (2018) over $>1.5 \text{ deg}^2$ in Figure 12. The optical luminosities and $H - IRAC$ colors from the Bowler et al. (2017) and Stefanon et al. (2018) selections only rely on IRAC bands not impacted by the [OIII] or $H\alpha$ nebular emission lines and hence only rely on sources at $z \geq 6.65$. As such, the luminosities are computed from the IRAC $4.5\mu\text{m}$ band for photometric candidates identified in the redshift range $z = 6.65-6.95$ and from the IRAC $3.6\mu\text{m}$ band for sources at redshifts $z > 7$.

Given the $6\times$ larger area of the Bowler et al. (2017) and Stefanon et al. (2019) searches relative to the current CANDELS probes and given that bright, red sources are presumably more common at $z \sim 7-8$ than at $z \sim 9-10$, it would seem odd to find more extreme sources in a $z \sim 9-10$ search over CANDELS than in the aforementioned lower-redshift, wider-area probes, and it therefore seems probable that most of the extreme candidates over CANDELS are actually at $z < 4$.

Also shown in Figure 12 are the luminosities and colors of the $z \sim 9-10$ galaxies in the Oesch et al. (2014), Bouwens et al. (2016), and Morishita et al. (2018) selections and the implied luminosities and colors for the other $z > 8$ candidates from Table 5 (*red open dotted triangles*) that our ground-based photometry largely rule out. The red $H-IRAC$ color selection limit used by Bouwens et al. (2016) for preselecting their $z \sim 9-10$ samples is shown in Figure 12, and it is clear that such a selection would easily include all extreme $z \sim 7-10$ sources identified by Bowler et al. (2017), Stefanon et al. (2019), and Oesch et al. (2014).

The purpose of the entire discussion here is to illustrate the challenges present in selecting high-quality $z > 8$ samples based on current data sets and how important accurate photometry is.

B.3 Possible Impact of these $z \sim 9-10$ Candidates on Our LF Estimates?

If we add up the fractional probability that the 28 sources presented in Table 5 and 4 are at $z > 8$, i.e., $P(z > 8)$, at face value this suggests that 8.67 of the presented candidates are at $z > 8$. As this constitutes a non-negligible fraction of the 19 candidates presented in Table 2, one might wonder whether the presented candidates in Appendix A and B would have an impact on the LF if they were indeed at $z \sim 9-10$.

As we argue below, the answer is that we would not expect the presented candidates to impact our derived LFs significantly, and the reason is that we expect to miss 30% of the bona-fide $z \sim 9-10$ candidates using the selection criteria utilized in this paper. Such a completeness fraction is built into the selection volume estimates we estimate for the CANDELS UDS, COSMOS, and EGS fields. The volumes we estimate for those fields are 30% lower than for fields like CANDELS GOODS-North where deeper ACS/optical and Y_{105} -band coverage is available. In most cases, this incompleteness occurs due to a blending of some $z \sim 9-10$ sources with neighboring sources and some limitations in the sensitivity and wavelength coverage available over the CANDELS WIDE fields.

Let now estimate how many bona-fide $z \sim 9$ candidates we are likely to find if we follow up all the candidates from Tables 5 and 4 with deeper Y_{105} -band observations. Including only those candidates from Tables 5 and 4 which have a redshift $z > 8.4$, the total fractional probability drops from 8.67 to 5.68.

In addition, we should calibrate our estimated $P(z > 8)$'s based on the follow-up results from Bouwens et al. (2015) which confirmed only 50% of 9 $P(z > 8) > 0.5$ candidates as bona-fide $z > 8$ source. As the $P(z > 8)$'s before follow-up sum to 6.5 in Bouwens et al. (2016) but were found to be 4.5 based on the follow-up data, we should renormalize the computed $P(z > 8)$'s down by 31%. Incorporating both effects, the total fractional probability is 3.91. In other words, we would expect to find 4 more bona-fide $z \sim 9$ candidates by following up all the sources in Tables 5 and 4 with Y_{105} -band observations.

There are 9 candidate $z \sim 9-10$ galaxies from the CANDELS UDS, COSMOS, and EGS fields given in Table 2 that we use to derive our $z \sim 9-10$ LF results. After we account for the insecure nature of 5 of them (counting each as 0.5), the total is 6.5. If we compare the effective number of sources used for the present LF determinations, i.e., 6.5 with the total sample including these additional sources, i.e., $6.5+3.91 = 10.46$, the fraction of sources in our main samples is 62%, which is not especially different from the $\sim 70\%$ fraction expected from our selection volume simulations.