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Citation

Zavala, J. A., Aretxaga, I., Dunlop, J. S., Michałowski, M. J., Hughes, D. H., Bourne, N., ...
Werf, P. P. van der. (2018). The SCUBA-2 Cosmology Legacy Survey: The EGS deep field -
II. Morphological transformation and multi-wavelength properties of faint submillimetre
galaxies. *Monthly Notices Of The Royal Astronomical Society*, 475(4), 5585-5602.
doi:10.1093/mnras/sty217

Version: Not Applicable (or Unknown)

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Note: To cite this publication please use the final published version (if applicable).

The SCUBA-2 Cosmology Legacy Survey: The EGS deep field – II. Morphological transformation and multi-wavelength properties of faint submillimetre galaxies

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Accepted 2018 January 23. Received 2018 January 16; in original form 2017 April 25.

ABSTRACT

We present a multi-wavelength analysis of galaxies selected at 450 and 850 μm from the deepest SCUBA-2 observations in the Extended Groth Strip (EGS) field, which have an average depth of $\sigma_{450} = 1.9$ and $\sigma_{850} = 0.46 \text{ mJy beam}^{-1}$ over $\sim 70 \text{ arcmin}^2$. The final sample comprises 95 sources: 56 (59%) are detected at both wavelengths, 31 (33%) are detected only at 850 μm , and 8 (8%) are detected only at 450 μm . We identify counterparts for 75% of the whole sample. The redshift distributions of the 450 and 850 μm samples peak at different redshifts with median values of $\bar{z} = 1.66 \pm 0.18$ and $\bar{z} = 2.30 \pm 0.20$, respectively. However, the two populations have similar IR luminosities, SFRs, and stellar masses, with mean values of $1.5 \pm 0.2 \times 10^{12} L_{\odot}$, $150 \pm 20 M_{\odot}/\text{yr}$, and $9.0 \pm 0.6 \times 10^{10} M_{\odot}$, respectively. This places most of our sources ($\gtrsim 85\%$) on the high-mass end of the ‘main-sequence’ of star-forming galaxies. Exploring the IR excess vs UV-slope (IRX- β) relation we find that the most luminous galaxies are consistent with the Meurer law, while the less luminous galaxies lie below this relation. Using the results of a two-dimensional modelling of the HST H_{160} -band imaging, we derive a median Sérsic index of $n = 1.4^{+0.3}_{-0.1}$ and a median half-light radius of $r_{1/2} = 4.8 \pm 0.4 \text{ kpc}$. Based on a visual-like classification in the same band, we find that the dominant component for most of the galaxies at all redshifts is a disk-like structure, although there is a transition from irregular disks to disks with a spheroidal component at $z \sim 1.4$, which morphologically supports the scenario of SMGs as progenitors of massive elliptical galaxies.

Key words: submillimetre: galaxies – galaxies: high redshift – galaxies: evolution – galaxies: star formation

1 INTRODUCTION

Since their discovery, submillimeter-selected galaxies (hereafter SMGs) have revolutionized the field of galaxy formation and evolution. These sources were detected, for the first

time, by the first submillimeter (850 μm) surveys taken with the James Clerk Maxwell Telescope (JCMT, [Smail et al. 1997](#); [Barger et al. 1998](#); [Hughes et al. 1998](#)), revealing that at least a fraction of the previously detected cosmic infrared background (CIB) by the space-based Cosmic Background Explorer (COBE, [Puget et al. 1996](#); [Fixsen et al. 1998](#)) came from dust-enshrouded galaxies. These results

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changed immediately our understanding of the cosmic star-formation history, and implied that surveys at both ultra-violet (UV)/optical and infrared (IR)/mm are necessary to completely understand it.

Thanks to the extensive follow-up studies carried out during the last two decades, we know that these are typically high-redshift galaxies ($\langle z \rangle \sim 2 - 3$, e.g. Aretxaga et al. 2003, 2007; Chapman et al. 2005; Michałowski et al. 2012a; Yun et al. 2012), with high star formation rates (SFRs, $\gtrsim 300 M_{\odot} \text{ yr}^{-1}$), large far-infrared (FIR) luminosities ($\gtrsim 10^{12} L_{\odot}$), large gas reservoirs ($\gtrsim 10^{10} M_{\odot}$), and orders of magnitude higher number density than local ultra-luminous infrared galaxies (see reviews by Blain et al. 2002 and Casey et al. 2014a). Furthermore, these galaxies are considered to be the progenitors of massive elliptical galaxies (e.g. Lilly et al. 1999; Smail et al. 2002; Simpson et al. 2014; Toft et al. 2014). For these reasons, this population is very important in our general comprehension of the stellar mass assembly, and therefore, in our understanding of the formation and evolution of galaxies over cosmic time. However, despite these significant efforts, our knowledge of this population of galaxies is still not complete, since most of the samples come from single-dish telescope observations with large beam-sizes ($\gtrsim 15$ arcsec) at just one wavelength. This introduces several biases: (1) due to selection effects, observations at a single wavelength are not representative of the whole population of galaxies (Zavala et al. 2014; Casey et al. 2013); (2) the poor angular resolution results in large position uncertainties leading to some misidentifications, as revealed by interferometric observations (Hodge et al. 2013); (3) the high confusion noise caused by the large beam-sizes prevents from detecting galaxies with $L_{\text{FIR}} \lesssim 10^{12} L_{\odot}$ at high redshifts.

Follow-up interferometric observations at submm bands exist now for more than a hundred SMGs (e.g. Iono et al. 2006; Younger et al. 2007, 2009; Wang et al. 2011; Smolčić et al. 2012a; Hodge et al. 2013; Ikarashi et al. 2015; Miettinen et al. 2015; Simpson et al. 2015; Brisbin et al. 2017), which alleviate the problem of positional uncertainty and source blending. Recent deep blank-field observations taken with the Atacama Large Millimeter/submillimeter Array (ALMA) have allowed the detection of galaxies with SFRs $< 100 M_{\odot} \text{ yr}^{-1}$ (e.g. Dunlop et al. 2017; Hatsukade et al. 2016; Umehata et al. 2017). These sources are also detected in small amplified samples towards clusters of galaxies thanks to gravitational amplification (e.g. Pope et al. 2017). However, due to the small surveyed areas only a handful of galaxies are typically detected. Additionally, these observations are taken at just one wavelength and may be not representative of the whole population. The achievement of wide-area surveys at different wavelengths which are necessary to solve these problems, as those planned with TolTEC¹ on the LMT or in a small scale with ALMA, will take several years. In the meanwhile, new deep single-dish telescope observations are coming along to minimize the biases in order to increase our knowledge of SMGs.

The SCUBA-2 Cosmology Legacy Survey (S2CLS; Geach et al. 2017) exploits the capabilities of the SCUBA-

2 camera (Holland et al. 2013) on the JCMT, efficiently achieving large and deep (confusion limited) maps at both 450 and 850 μm simultaneously, allowing us the detection of galaxies with $L_{\text{FIR}} \lesssim 10^{12} L_{\odot}$ up to $z \sim 3$ (see §4.2.1). The higher angular resolution at 450 μm results in a positional uncertainty of $\sim 1 - 2$ arcsec, and therefore, in high accuracy associations than previous studies based on sample of galaxies selected from single-dish telescope observations. Previous SCUBA-2 studies (including those from the S2CLS) have taken these advantages to characterize the physical properties of faint SMGs (Chen et al. 2013; Geach et al. 2013; Roseboom et al. 2013; Hsu et al. 2016; Koprowski et al. 2016; Cowie et al. 2017; Michałowski et al. 2017), although some of these studies have been focused on samples of either 450 or 850 μm -selected galaxies. Here we present a multi-wavelength counterpart analysis of a sample built from both 450 and 850 μm -selected galaxies, in order to minimize selection effects and to investigate any differences in the physical properties of these galaxies.

This paper is organised as follows: sample selection and multi-wavelength data are described in §2. The counterpart matching and identification process are reported in §3. In §4 we derive the physical properties, such as redshifts, luminosities, SFRs, stellar masses, dust properties, and discuss the location of our galaxies in the star-forming main-sequence. The morphology classification as well as a possible morphological evolution is also discussed in §4. Finally, our results are summarized in §5.

All calculations assume a standard Λ cold dark matter cosmology with $\Omega_{\Lambda} = 0.68$, $\Omega_{\text{m}} = 0.32$, and $H_0 = 67 \text{ kms}^{-1} \text{ Mpc}^{-1}$ (Planck Collaboration et al. 2014).

2 MULTI-WAVELENGTH DATA

2.1 S2CLS

The main data for this study comes from the deep 450 and 850 μm observations taken with the SCUBA-2 camera in the EGS field as part of the S2CLS. The characteristics of observations and the data reduction process are described in the first paper of this series (Zavala et al. 2017), where the source catalogues and number counts are also reported. These observations have a mean depth of $\sigma_{450} = 1.9$ and $\sigma_{850} = 0.46 \text{ mJy beam}^{-1}$ (including instrumental and confusion noise) at 450 and 850 μm , respectively, within an area of $\sim 70 \text{ arcmin}^2$. Along with other SCUBA-2 surveys (see §1), these are some of the deepest observations taken at these wavelengths with a single-dish telescope, enabling the detection of relatively faint galaxies (i.e. $0.6 \lesssim S_{850\mu\text{m}} \lesssim 6 \text{ mJy}$) that were unreachable by previous blank-field surveys. For comparison, the sources detected in the LABOCA ECDFS submillimetre Survey (LESS; Weiß et al. 2009), one of the largest and deepest contiguous map at submillimetre wavelengths before the achievement of the S2CLS, were limited to $S_{870\mu\text{m}} \gtrsim 4 \text{ mJy}$. Finally, the angular resolution at 450 μm is $\theta_{\text{FWHM}} \approx 8$ arcsec and ≈ 14.5 arcsec at 850 μm .

¹ <http://toltec.astro.umass.edu>

2.2 Ancillary data

Thanks to the All-wavelength Extended Groth Strip International Survey (AEGIS²), this field has a panchromatic dataset from X-rays to radio wavelengths. We use these observations, among other catalogues described below, to identify the counterpart galaxies and to study their multi-wavelength properties.

The radio and IR images are used as an intermediate step to associate each galaxy in our catalogue with an optical counterpart (see §3.1). For this purpose we use the VLA/EGS 20 cm (1.4 GHz) survey described by Ivison et al. (2007b). These observations have an angular resolution of FWHM ≈ 3.8 arcsec with a 5σ detection limit of $50 \mu\text{Jy}$ in the deepest region. As IR constraints, we use the catalogue derived from *Spitzer*/MIPS observations at $24 \mu\text{m}$ which are part of the Far-Infrared Deep Extragalactic Legacy Survey (FIDEL³). The 3σ detection limit is $30 \mu\text{Jy}$ and it has an angular resolution of 5.9 arcsec. In addition, the $8 \mu\text{m}$ catalogue derived using *Spitzer*/IRAC observations (Barro et al. 2011a) is also exploited for this purpose, with a 5σ limit of 22.3 mag and an angular resolution of 2.2 arcsec.

We use the multi-wavelength catalogues compiled by the 3D-HST team (Brammer et al. 2012; Skelton et al. 2014; Momcheva et al. 2016) to finally associate our radio or IR counterpart with an optical galaxy. These catalogues are built mainly using *H*-band selected galaxies from *HST*(F160W) imaging with a median 5σ depth of 26.4 mag (AB) and include photometry from the *u*-band to $8 \mu\text{m}$. We also use an *Spitzer*/IRAC 3.6+4.5 μm selected catalogue (Barro et al. 2011a; Barro et al. 2011b) which contains photometry from the ultraviolet to $70 \mu\text{m}$. The catalogue includes sources up to 5σ of 23.9 mag at $3.6 \mu\text{m}$. As discussed in §4.1 and §4.3 these catalogues include redshift information and stellar-population parameters, which are used to understand the physical properties of our sample.

To investigate the AGN contamination in our sample, the catalogue of the AEGIS-X Deep (AEGIS-XD) survey⁴ (Nandra et al. 2015) is used. This program combines deep *Chandra* observations in the central region of EGS with previous *Chandra* observations of a wider area (AEGIS-X Wide, Laird et al. 2009), achieving a total nominal exposure depth of 800 ks in the central region. This makes this program one of the deepest X-ray survey in existence.

To extract far-infrared photometry we use *Herschel* observations obtained with the Photodetector Array Camera and Spectrometer (PACS, Poglitsch et al. 2010) and the Spectral and Photometric Imaging Receiver (SPIRE, Griffin et al. 2010) instruments, which are part of the PACS Evolutionary Probe (PEP, Lutz et al. 2011) and the *Herschel* Multi-tiered Extragalactic Survey (HerMES, Oliver et al. 2012) programs. We obtained the *Herschel* fluxes of each SCUBA-2 source as in Michałowski et al. (2017) in the following way. We extracted 120 arcsec wide stamps from all five *Herschel* maps around the position of each SCUBA-2 source. Then we processed the PACS (100 and $160 \mu\text{m}$) maps by simultaneously fitting Gaussian functions with the FWHM of the respective resolution of the

maps, centred at the positions of all $24 \mu\text{m}$ sources located within these cut-outs, and at the positions of the SCUBA-2 IDs. Then, to deconvolve the SPIRE (250, 350 and $500 \mu\text{m}$) maps in a similar way, we used the positions of the $24 \mu\text{m}$ sources detected with PACS ($\geq 3\sigma$), the positions of all SCUBA-2 ID positions (or the submm positions if no radio or mid-IR ID had been secured). The errors were computed from the covariance matrix of the fit, in which the free parameters are simply the heights of the Gaussian beams fitted at each input position. Then the confusion noise of 5.8, 6.3 and $6.8 \text{ mJy beam}^{-1}$ at 250, 350 and $500 \mu\text{m}$, respectively (Nguyen et al. 2010) was added in quadrature. The fitting was performed using the IDL MPFIT⁵ package (Markwardt 2009).

3 SOURCE IDENTIFICATION

For this study, we limit our sample to those sources detected at $> 3.75\sigma$ at 450 or $850 \mu\text{m}$. At this threshold, the contamination due to false-detections is expected to be $\lesssim 5$ per cent (Zavala et al. 2017), which is acceptable for the goals of this paper. On the other hand, sources detected at two (or more) bands have a higher effective S/N, for example, a source detected at 3σ at two different wavelengths has an effective S/N of $\sim 4\sigma$. For this reason, we also include all the sources detected at both wavelengths with S/N > 3.0 . A source is considered detected at both wavelengths if the separation between the 450 and $850 \mu\text{m}$ position is $\lesssim 7$ arcsec, which corresponds to 2.5σ the joined expected positional uncertainty added in quadrature according to the beam-size and S/N (Ivison et al. 2007a).

Following these criteria, our final sample consist of 95 sources: 56 (corresponding to 59% of the total sample) are detected at both wavelengths, 31 (33%) are detected only at $850 \mu\text{m}$, and 8 (8%) are detected only at $450 \mu\text{m}$.

3.1 Counterpart matching

Traditionally, the first step to find a SMG counterpart is to associate it with a radio or IR source, since the source density at optical wavelengths is much higher and therefore many potential counterparts lie within a typical search radius.

The radio band (1.4GHz) is chosen because it also traces recent star formation via synchrotron radiation and there is a well-known FIR-radio correlation (Yun & Carilli 2002; Barger et al. 2014). On the other hand, $24 \mu\text{m}$ observations are sensitive to the warm dust emission and SMGs have been found to be bright at this wavelength (e.g. Pope et al. 2006). Additionally, the surface density of sources is relatively low at these bands and consequently the probability for a misidentification is expected to be low, although some ALMA studies reported a misidentification fraction of up to ~ 30 per cent (Hodge et al. 2013). Over the last decade, observations at $8 \mu\text{m}$ have also been used for counterpart identifications (e.g. Ashby et al. 2006; Pope et al. 2006; Biggs et al. 2011; Michałowski et al. 2012a; Yun et al.

² <http://aegis.ucolick.org/>

³ <http://irsa.ipac.caltech.edu/data/SPITZER/FIDEL/>

⁴ <http://www.mpe.mpg.de/XraySurveys/AEGIS-X/>

⁵ purl.com/net/mpfit

2012) since these trace the emission from the older and mass-dominant stellar populations in this kind of galaxies and, although the surface density is higher, these observations are usually deeper than those at $24\ \mu\text{m}$, increasing the fraction of identifications.

Using these three wavebands independently, we search for counterparts with a variable search radius equal to 2.5σ (corresponding to ~ 96 per cent probability), where σ is the positional uncertainty of each galaxy based on its S/N and the FWHM of the beam, as described by Ivison et al. (2007a). If available, the $450\ \mu\text{m}$ position is favored over the $850\ \mu\text{m}$. Additionally, we imposed a minimum search radius of 4 arcsec to account for systematic astrometry differences between catalogues and for the positional uncertainties of the NIR/radio sources. Furthermore, this minimum search radius help us to increase the fraction of identifications, since it has been found that some SMGs detected with relatively high S/N lie outside of their nominal 2.5σ positional uncertainties (Hodge et al. 2013). Then, if a potential counterpart is found, we estimate the statistical significance of the association or in other words, the corrected Poisson probability, p , that the counterpart candidate has been selected by chance, following the method described by Downes et al. (1986) (see also Dunlop et al. 1989 and Ivison et al. 2007a).

An empirical value of $p < 0.05 - 0.1$ to define robust counterparts for SMGs is commonly adopted in the literature (e.g. Pope et al. 2006; Chapin et al. 2009b; Yun et al. 2012), however, using ALMA observations Chen et al. (2016) showed that there is no statistical difference between the accuracy if we adopt $p < 0.05$ or $p < 0.1$ (i.e. consistent within the error bars) but a better completeness if we select the later. Therefore, we adopt as probable counterparts those galaxies with $p < 0.1$. For the case of galaxies with multiple candidates (20 out of 71), we select the one with the lowest p -value, following Michałowski et al. (2017), who have shown that this procedure recovers the galaxy with the dominant contribution to the submm flux density in most cases ($\sim 86\%$).

Using these criteria, and after rejecting six candidates based on their discrepant FIR photometric redshifts (see §4.1), we achieve a successful identification rate of $\sim 75\%$ (71 out of 95) for the whole sample. In terms of the SCUBA-2-band detections, the successful identification rate is 82% for those galaxies detected at both wavelengths, 88% for those detected only at $450\ \mu\text{m}$, and 62% for those detected only at $850\ \mu\text{m}$. The lower detection rate for the $850\ \mu\text{m}$ -only detected galaxies may reflect the higher redshift nature of these galaxies (see §4.1), albeit the larger beam-size and the source blending translate into a larger positional uncertainty, which can also explain the lower fraction of counterparts.

Once we have the radio/IR counterparts, we match these sources with the 3D-HST catalogue (Skelton et al. 2014; Momcheva et al. 2016) using a search radius of 1.5 arcsec (although all the matches lie within 1 arcsec). For those galaxies which lie outside of the 3D-HST coverage (13 out of 71), we use instead the IRAC catalogue (Barro et al. 2011a,b). All the 71 galaxies with radio/IR counterparts also have an optical counterpart in these catalogues. All properties are summarized in Table A.

3.2 Reliability of galaxy identifications

Before deriving the physical properties of the sample, it is important to test the reliability of our identification methodology and to quantify the misidentification fraction. It is clear that submm/mm interferometry represents the best way of identifying counterparts correctly (e.g. Younger et al. 2007, 2009; Wang et al. 2011; Barger et al. 2012; Smolčić et al. 2012b; Hodge et al. 2013). However, the archival submm/mm interferometry data is scarce in this field (which is, furthermore, inaccessible to ALMA). From the PHIBSS survey (Tacconi et al. 2013), we found 6 fields targetted with IRAM Plateau de Bure millimeter interferometer (PdBI) whose detections lie close (< 10 arcsec) to sources in our catalogue. These observations were designed to map the $^{12}\text{CO}(3-2)$ transition in massive ($M_* > 2.5 \times 10^{10} M_\odot$), main-sequence star-forming galaxies at $z \sim 1.2$ and 2.2 . Since it has been shown that SMGs are bright in CO emission (e.g. Greve et al. 2005) and giving the properties of the PHIBSS targetted galaxies, we can assume that the CO emission is associated with an SMG. Moreover, the probability of finding by chance a CO line emission close to our SMGs is low, based on the results from blind CO surveys (Walter et al. 2016). Assuming then that these six $^{12}\text{CO}(3-2)$ emission lines come from the SMGs, we can test our galaxy identifications. We find that for 5 galaxies we have identified the correct counterpart with our method. The remaining source is one of those with only $850\ \mu\text{m}$ detection, which has a larger positional uncertainty. Alternatively, the CO line might have been too faint to be detected by PHIBSS. Based on these results, we estimate a correct counterpart accuracy of $\gtrsim 83\%$.

On the other hand, Casey et al. (2013) estimated the counterpart contamination when using the SCUBA-2 $450\ \mu\text{m}$ position based on the reduced beamsize with respect to $850\ \mu\text{m}$, assuming a reduction in counterpart contamination proportional to the reduction in sky area searched for potential counterparts, and adopting a $\sim 30\%$ of incorrect associations for the $850\ \mu\text{m}$ sources (Hodge et al. 2013). Based on this, they inferred a $\sim 5\%$ of misidentifications for $450\ \mu\text{m}$ sources. Using these values, and considering that only 33% of our sample lack $450\ \mu\text{m}$ positions, we estimate that $\sim 13\%$ of the sample could have incorrect identifications.

Finally, we compare the IRAC colours of our identified galaxies with the IRAC colour-colour space diagram proposed for SMG counterpart identification by Yun et al. (2008). Securely identified SMGs (with radio or submm interferometric observations) are known to lie within this proposed diagram (Alberts et al. 2013), and may help us to inquire into the reliability of our identifications. As shown in Fig. 1, most of our identified counterparts ($\sim 95\%$) lie on this space (and the remaining sources lie very close to the edges), which further supports the associations.

In summary, based on these three independent tests, we infer that our identification method recovers the correct counterpart with an accuracy of $\gtrsim 85\%$, which is acceptable for an statistical characterization of the population.

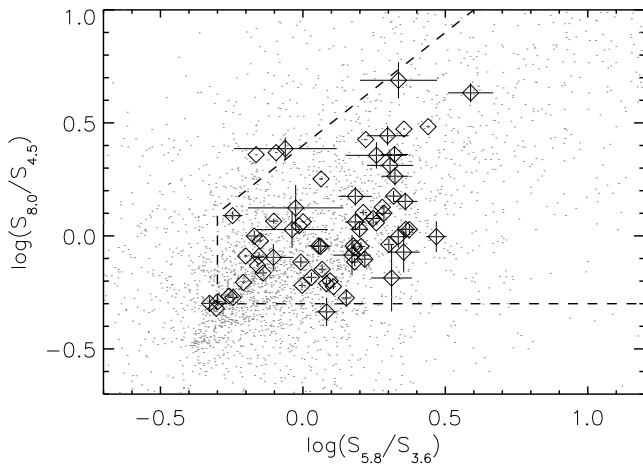


Figure 1. *Spitzer*/IRAC $S_{5.8\mu\text{m}}/S_{3.6\mu\text{m}}$ vs. $S_{8\mu\text{m}}/S_{4.5\mu\text{m}}$ colour-colour diagram for the associated counterparts to our SMGs (black diamonds). The dashed lines enclose the IRAC colour cuts suggested for SMG counterpart identification by Yun et al. (2008), where most of our identified SMGs lies. The grey dots represent other galaxies in the 3D-HST catalogue in the same field.

4 DERIVED PROPERTIES

4.1 Redshift distribution

The optical catalogues described above include redshift estimations for most of the sources. For those galaxies with identified counterparts in the 3D-HST catalogue, we use the combined ‘Z-BEST’ redshift information (hereafter z_{opt}). The best redshift is: (1) spectroscopic long-slit redshift, if available in the compilation by Skelton et al. (2014); (2) spectroscopic *HST* grism redshift (Momcheva et al. 2016), if there is no archive spectroscopic redshift and if the grism spectrum is not flagged as faulty; or (3) optical photometric redshift derived from the EAZY code (Brammer et al. 2008). For the small fraction of sources with no 3D-HST counterpart, we use the photometric redshifts reported in the IRAC catalogue (Barro et al. 2011b). Although these redshifts have been estimated using the *Rainbow* code (Barro et al. 2011a), the authors have shown that there is a good agreement with the values estimated using EAZY. The typical uncertainty for the grism-based redshifts is $\Delta z/(1+z) \approx 0.003$ (Momcheva et al. 2016), and $\Delta z/(1+z) \lesssim 0.04$ for the photometric redshifts derived from EAZY or *Rainbow* (Barro et al. 2011b; Skelton et al. 2014) with a catastrophic failure rate of less than 10 per cent, when considering only $z > 0.5$ sources (although, most of them are optically-selected galaxies). However, when considering only the galaxies in our catalogue, we find a median uncertainty of $\Delta z/(1+z) = 0.05$ for the photometric redshifts.

On the other hand, we estimate photometric redshifts for the whole sample using the rest-frame FIR photometric data provided by *Herschel* 100, 160, 250, 350, 500 μm , and our SCUBA-2 observations at 450 and 850 μm (hereafter z_{FIR}). Though this method is not as accurate as the optical photometric redshifts, it allows us to estimate the redshift for sources with no counterparts, and therefore, to derive

a complete redshift distribution. We fit an average SMG SED template (Michałowski et al. 2010) to the flux density measured at these wavelengths. This is the average SED of 70 SMGs with spectroscopic redshifts, where GRASIL models (Silva et al. 1998) were fitted to the photometry and then averaged. During this procedure the adopted flux calibration errors are 2 & 4% at 100 and 160 μm ⁶, 4% for the SPIRE bands (Bendo et al. 2013), and 10 & 5% for the SCUBA-2 450 and 850 μm (Dempsey et al. 2013). In case of non-detections, we incorporate the upper limits into the fitting through a survival analysis (Isobe et al. 1986) following the same procedure as in Aretxaga et al. (2007). For some sources only lower limits on redshifts could be derived (identified with upward arrows in Fig. 2) due to the high number of non-detections. These z_{FIR} values are compiled in Table A.

In order to investigate the reliability of these photometric redshifts, we compare the z_{FIR} with the optical redshifts described above for those galaxies with optical counterparts. As can be seen in Fig. 2, there is generally a good agreement between both estimations within the large scatter. The relative difference between these methods ($\Delta z = (z_{\text{FIR}} - z_{\text{opt}})/(1 + z_{\text{opt}})$) is well fitted with a Gaussian function (see inset plot in Fig. 2) with a standard deviation of 0.25. However, there are six sources (marked with squares in the figure), for which the two redshift estimations differ dramatically ($z_{\text{FIR}} = 4 - 5$ vs. $z_{\text{opt}} \sim 1$). This may reflect an optical misidentification due to a nearby galaxy in the line of sight (e.g. Bourne et al. 2014) or because the identified optical galaxy is in fact lensing a more distant submm source (e.g. Geach et al. 2015). All of these galaxies have $S/N > 5$ and a low probability of being a false detection. For these reasons, we reject these optical associations (hence also their z_{opt}) for the rest of the analysis, in order to avoid introducing any bias (following Koprowski et al. 2016 and Michałowski et al. 2017).

A trend can be seen in Fig. 2, where our z_{FIR} is systematically lower than the optical redshift at $z_{\text{opt}} > 2$. This may reflect the fact that our SED template is represented by a single temperature and reveals a relation between redshift and dust temperature, where warmer temperatures are necessary at higher redshifts to correct for this effect. The same trend was actually found in the photometric redshift of red *Herschel* sources (Ivison et al. 2016). However, due to the small number of photometric measurements for some sources in our catalogue and to the large scatter of this trend, we decide to keep our single dust temperature template for this purpose. As described below, these results are only used to complete the redshift distribution of all the sources in our catalogue. The rest of the analyses in the paper are focused only on those galaxies with optical counterparts. A detailed discussion of the dust temperature evolution is presented in §4.2.2.

The final redshift distributions for all the 450 and 850 μm -selected galaxies are shown in Fig. 3. These include optical spectroscopic redshifts (36 sources, including *HST* grism-derived redshifts), optical photometric redshifts (35 galaxies), and FIR redshifts for those sources with no counterparts

⁶ http://herschel.esac.esa.int/twiki/pub/Public/PacsCalibrationWeb/pacs_bolo_fluxcal_report_v1.pdf

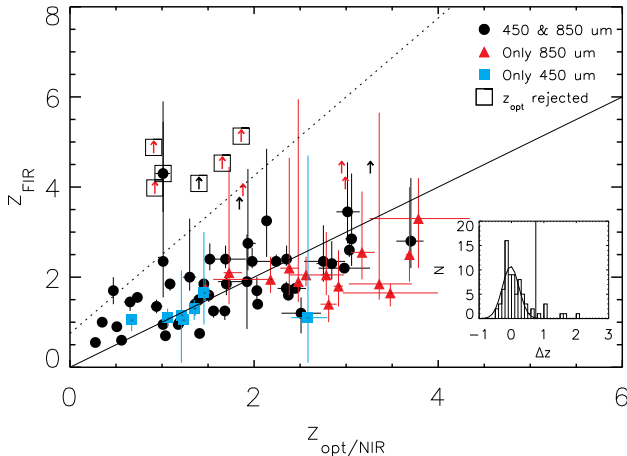


Figure 2. Comparison between FIR photometric redshifts, derived from *Herschel* and SCUBA-2 observations (z_{FIR}), and optical-infrared spectroscopic or photometric redshifts (z_{opt}) of the counterparts (see §4.1). Those sources with only lower limits are marked with upward arrows. The inset plot shows a histogram of the relative difference ($\Delta z = (z_{\text{FIR}} - z_{\text{opt}})/(1 + z_{\text{opt}})$) between the two estimations, which is well fitted with a Gaussian distribution. The sources with a discrepant redshift above 3σ of the Gaussian distribution (represented by the dotted line) are rejected given the potential of misidentification (see §4.1). The colours indicate if the source is detected at both wavelengths (black symbols) or only in a single band (blue and red symbols for 450 or 850 μm , respectively).

(24 sources). Individual values are reported in Table A. In order to take into account the uncertainties in our values, the redshift distributions have been estimated by stacking the redshift likelihood distribution of each source. The stacked redshift distributions expand between $0 < z < 6$, although, as expected due to the k -correction at NIR and radio wavelengths, all the galaxies with identified counterparts lie at $z < 4$. The median redshifts of the distributions are $z_{\text{med}} = 1.66 \pm 0.18$ and 2.30 ± 0.20 (with the errors derived from a bootstrap method) for the 450 and 850 μm -selected galaxies, respectively. This supports the statement that galaxies selected at different wavelengths have different redshift distributions as previously reported (e.g. Zavala et al. 2014; Béthermin et al. 2015b). Actually, all (but one) of the galaxies detected only at 450 μm (i.e. with no 850 μm detection) lie at $z \lesssim 1.5$, while all the galaxies detected only at 850 μm lie at $z \gtrsim 1.5$. As described in the next sections, this is the main difference between the galaxies detected at only one of the two wavebands.

4.1.1 Comparison with previous surveys

We compare our redshift distributions with those derived from previous similar studies in Figure 3.

Roseboom et al. (2013) studied a sample of 450 μm -selected galaxies from the S2CLS observations in the UDS and COSMOS fields with a depth similar to this work, although with a S/N threshold of 4. Their median redshift of 1.4 ± 0.2 is close to our value of 1.66 ± 0.18 , however, as it can be seen in the figure, they have no sources with $z > 3$.

This is expected since their redshift distribution is limited to those galaxies with optical counterparts, and therefore, the highest redshift galaxies are missing. On the other hand, Casey et al. (2013) presented an analysis for 450 and 850 μm galaxies on the COSMOS field covered with SCUBA-2 observations. For the 450 μm sample they reported a median redshift of $z = 1.95 \pm 0.19$. This value is slightly higher than the one found in this work, however, their sample comes from larger but shallower observations selecting, therefore, brighter galaxies that are not present in our smaller field (i.e. sources with $S_{450\mu\text{m}} > 20$ mJy). This is consistent with the picture that galaxies with higher fluxes are located preferentially at higher redshifts as suggested by previous studies (e.g. Pope et al. 2005; Koprowski et al. 2014). Finally, we also compare our 450 μm redshift distribution with theoretical predictions of the GALFORM semi-analytical model (Lacey et al. 2016) for galaxies with $S_{450} > 7$ mJy. This flux density limit corresponds to the average limit of our sample (S/N > 3.75). The model predicts a median redshift of 1.97, which is consistent with our value within $\sim 1.7\sigma$. This model also includes the rare bright galaxies that our relatively small map cannot constrain, which could account for the mild difference between the redshift distributions. Actually, the predictions of the model are in better agreement with the results of Casey et al. (2013), where, as mentioned before, the brightest galaxies are better sampled due to the larger area mapped.

At 850 μm , Casey et al. (2013) reported a median redshift of $z = 2.16 \pm 0.11$ which is just below but consistent with our median value of 2.3 ± 0.2 . Nevertheless, our distribution shows a flat high-redshift tail between $4 < z < 6$ that is not present in the Casey et al. result. This is not surprising since their redshift distribution is limited to those galaxies with optical counterparts. On the other hand, da Cunha et al. (2015) presented the redshift distribution of sources with 870 μm fluxes above 4 mJy (a factor of ~ 2 larger than our threshold) derived from MAGPHYS SED modeling, which has a median redshift of 2.7 ± 0.1 . In contrast to our distribution, they have no sources at $z < 1$ (Fig. 3). This may reflect the fact that we are using the 450 μm information to select faint sources below our formal detection limit, increasing the completeness of our sample. As in this work, Koprowski et al. (2016) used deep S2CLS observations to study the properties of 850 μm galaxies in the COSMOS field. Their redshift distribution with a median of 2.38 is in very good agreement with our result, as can be seen from Figure 3. On the theoretical side, the GALFORM model predicts a median of 2.88 for galaxies with $S_{850} > 1.5$ mJy. This value is higher than ours but consistent within 3σ .

In general, our results are in broad agreement with previous studies and show a trend where longer wavelength observations select, on average, higher redshift galaxies.

4.2 IR luminosities, SFRs, & dust properties

The rest-frame FIR photometry allows us to derive infrared luminosities (which can be converted to star formation rates) and dust temperatures once a SED is fitted to the data. We use a modified blackbody function which is described by

$$S_\nu \propto \{1 - \exp[-(\nu/\nu_0)^\beta]\} B(\nu, T_d), \quad (1)$$

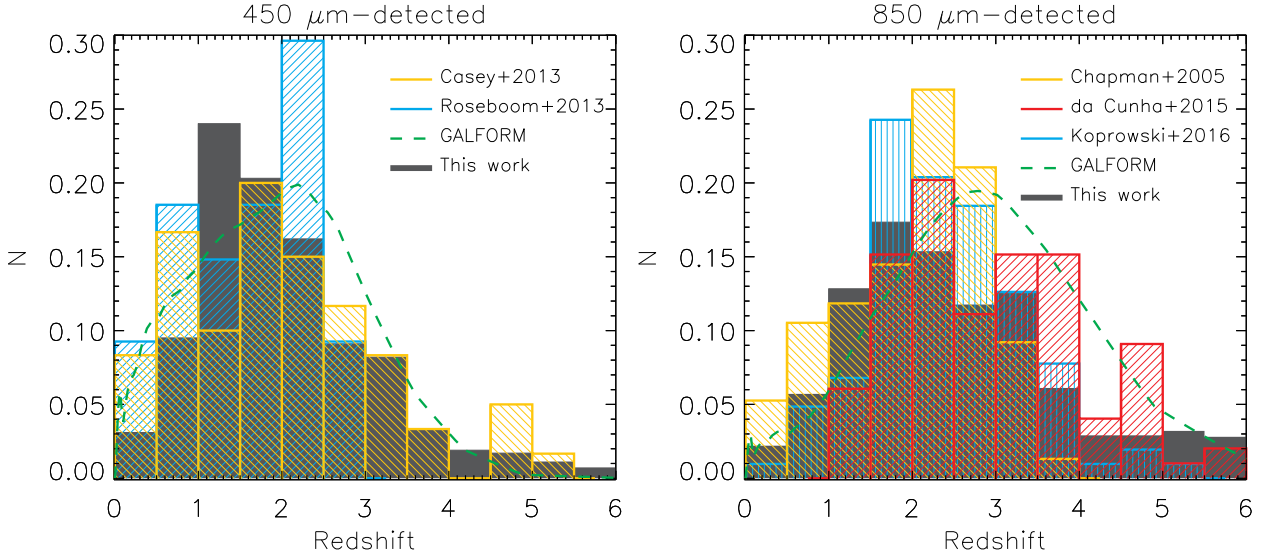


Figure 3. The normalized redshift distributions of our 450 and 850 μm -detected sources (*left and right*, respectively) are represented by the grey histograms. The 450 μm distribution has a median redshift of $z_{\text{med}} = 1.66 \pm 0.18$ and the 850 μm one a median of $z_{\text{med}} = 2.30 \pm 0.20$. For comparison, the results from previous observational studies are plotted in each panel (Roseboom et al. 2013; Chapman et al. 2005; Casey et al. 2013; da Cunha et al. 2015; Koprowski et al. 2016), as well as the theoretical predictions of the GALFORM model (Lacey et al. 2016) matched to the mean S2CLS/EGS depth.

where S_ν is the flux density at frequency ν , ν_0 is the rest-frame frequency at which the emission becomes optically thick, T_d is the dust temperature, β is the emissivity index, and $B(\nu, T_d)$ is the Planck function at temperature T_d . To minimize the number of free parameters, the emissivity index, $\beta = 1.6$, is fixed (previous observational works suggest $\beta = 1.5 - 2$; e.g. Dunne & Eales 2001; Farrah et al. 2003; Chapin et al. 2009b; Magnelli et al. 2012), as well as $\nu_0 = c/100 \mu\text{m}$ (Riechers et al. 2013; Simpson et al. 2016), where c is the speed of light. Furthermore, in order to break the temperature-redshift degeneracy (e.g. Blain et al. 2002), we fix the modified blackbody at the redshift provided by the optical catalogues (see §4.1). Therefore, galaxies with no counterparts (24/95) are not included in this analysis. All the derived properties, named luminosities, SFRs, and dust temperatures, are reported in Table A, and are discussed below. If we choose instead $\beta = 2.0$, the derived IR luminosity does not change (within 1%) but the dust temperature decreases on average by $\sim 10\%$.

4.2.1 Luminosity

In Fig. 4 we plot the IR luminosity (8 – 1000 μm) as a function of redshift for our sources with optical counterparts, and the detection limit (taking the best case between the 450 and 850 μm limit at each redshift) assuming a modified blackbody with $T_d = 47 \text{ K}$ (the average temperature of the sample) and T_d increasing with redshift (as measured in our data, see §4.2.2). As it can be seen, the detection limit predicted by the fixed dust temperature SED does not reproduce our estimations, over-predicting the L_{IR} limit at low redshift and underpredicting it at high redshift. This effect was also noticed by Ivison et al. (2016). On the other hand, the expected detection limit for a modified blackbody with

dust temperature increasing with redshift is in much better agreement with our measurements. This provides additional evidence of a relation between $T_d - z$ (see §4.1 and §4.2.2).

Our survey is deep enough to detect galaxies down to an IR luminosity of $L_{\text{IR}} \sim 1.5 \times 10^{12} L_\odot$ at any redshift below $z \sim 4$ (see Fig. 4), which corresponds to a SFR of $150 M_\odot \text{ yr}^{-1}$, assuming the Kennicutt 1998 relation for a Chabrier 2003 IMF. This highlights the depth of our survey, which actually allows us to detect several galaxies with $20 < \text{SFR} < 100$ at $1 < z < 2$ (as other recent deep S2CLS studies, e.g. Koprowski et al. 2016; Bourne et al. 2017). This population of galaxies was unreachable by previous submillimeter surveys with single-dish telescopes in blank fields. Also shown in Fig. 4 is the evolution of the knee of the IR luminosity function (L^* , Gruppioni et al. 2013), which is slightly fainter (less than a factor of 1.5) than our sensitivity limit.

The mean IR luminosity of our sample is $1.5 \pm 0.2 \times 10^{12} L_\odot$ which corresponds to a mean SFR of $150 M_\odot \text{ yr}^{-1}$ (where the errors have been estimated by a bootstrapping method), covering a range from $L_{\text{IR}} = 0.3 \times 10^{11} - 7.4 \times 10^{12} L_\odot$ (SFR = 3 – 740 $M_\odot \text{ yr}^{-1}$). We notice that at a fixed redshift, the galaxies detected at both wavelengths (black circles in Fig. 4) are, on average, more luminous than those detected at only one wavelength (blue and red circles for 450 and 850 μm -only detected sources, respectively), as expected.

4.2.2 Dust temperature

The mean dust temperature for the whole sample is $\langle T_d \rangle = 47 \pm 15 \text{ K}$, which is in agreement with previous studies. For example, da Cunha et al. (2015) found, using the MAGPHYS code (da Cunha et al. 2008), an average dust temperature of $T = 43 \pm 2 \text{ K}$ for the 870 μm ALMA-detected sources from

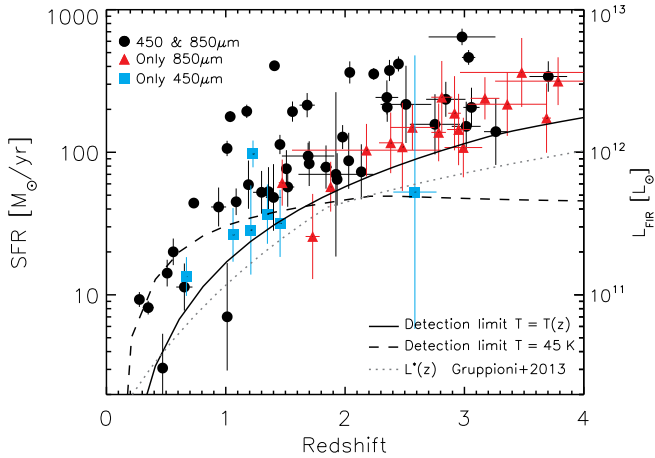


Figure 4. SFR ($\propto L_{\text{IR}}$) as a function of redshift for galaxies with optical counterparts. The solid line represents the sensitivity limit of our survey (considering both wavelengths) assuming a modified blackbody with a dust temperature increasing with redshift (see §4.2.2), and the dashed line a model with fixed dust temperature at $T_d = 47$ K (the average temperature of our sample). As can be seen, the expected detection limit predicted by the SED with variable temperature is in better agreement with the measurements. The depth of our survey allows us to detect galaxies with SFRs down to $\sim 50 M_{\odot} \text{ yr}^{-1}$ at $z \approx 1-2$ and we are sensitive to SFR $> 150 M_{\odot} \text{ yr}^{-1}$ at any redshifts. The dotted line is $L^*(z)$ as reported by Gruppioni et al. (2013), which is slightly below (less than a factor of 1.5) our detection limit.

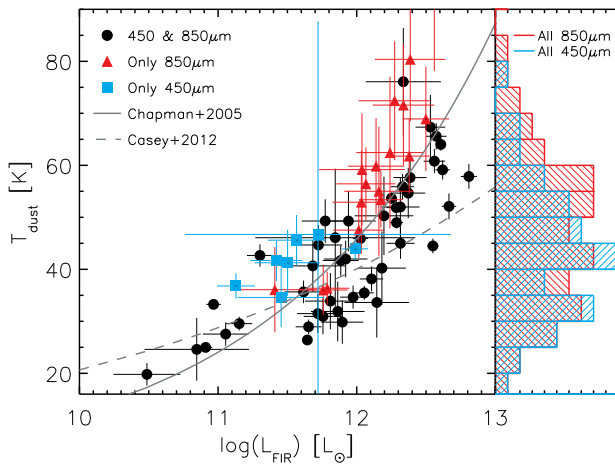


Figure 5. *Left:* Relation between dust temperature (T_d) and IR luminosity (L_{IR}), derived from a modified blackbody SED fitting at fixed redshift. The solid line is the relation reported by Chapman et al. (2005), which qualitatively reproduces our measurements, and the dashed line represents the relation found by Casey et al. (2012). *Right:* Dust temperature distributions for all the 450 and 850 μm -detected sources represented by the blue and red histograms, respectively. The mean dust temperature for each sub-sample is $\langle T_d \rangle = 43 \pm 12$ K and 49 ± 15 K for the 450 and 850 μm galaxies, respectively.

the ALESS survey (Hodge et al. 2013; Karim et al. 2013), which cover a similar range of redshifts and flux densities. On the other hand, Roseboom et al. (2013) derived a mean temperature of $T_d = 42 \pm 11$ K for 450 μm -selected galaxies from the S2CLS in COSMOS and UDS, which have roughly the same depths as our maps. These temperatures are, however, hotter than the ones derived by Swinbank et al. (2014) for the ALESS sample when using modified blackbody fits, which have a median value of ≈ 32 K. These differences may be partially due to differences in the modified blackbody distributions, i.e. optically thin vs optically thick, dissimilar emissivity indices, etc., as discussed in §4.2.

As shown in Fig. 5, the dust temperature follows the well-known temperature-luminosity relation found in previous studies (e.g. Chapman et al. 2005; Kovács et al. 2006; Chapin et al. 2009a; Magnelli et al. 2012; Symeonidis et al. 2013; Swinbank et al. 2014; Hodge et al. 2016). Our results are consistent with the relation found by Chapman et al. (2005) ($T_d \propto L_{\text{FIR}}^{0.28}$; solid line in the figure). This slope is steeper than the one reported by Casey et al. (2012) ($T_d \propto L_{\text{FIR}}^{0.14}$; dashed line in Fig 5), however, their sample peaks at lower redshift (~ 0.85 vs ~ 2.3) and it is selected at shorter wavelengths, so the differences can be explained by selection effects (see discussions by Chapin et al. 2011, Marsden et al. 2011, and Casey et al. 2012).

Interestingly, if we inquire into the dust temperature of the 450 and 850 μm -detected galaxies independently, we find that both sub-samples show similar values. The dust temperature distributions for the 450 and 850 μm -detected galaxies are represented in the right side of Fig. 5. These distributions are not statistically different from each other, with mean dust temperatures of $\langle T_d \rangle = 43 \pm 12$ K and 49 ± 15 K for the 450 and 850 μm galaxies, respectively. This supports the idea that the redshift is the main difference between both populations.

We also find an evolution of the dust temperature, with higher temperatures at higher redshifts, as is clearly seen in Fig. 6. This evolution has been reported in previous observational works (Magdis et al. 2012; Magnelli et al. 2014; Genzel et al. 2015; Kirkpatrick et al. 2015; Hodge et al. 2016), as well as in theoretical predictions (Lagos et al. 2012; Cowley et al. 2017). The best-fitted linear model to our data is described by $T_d = 12(1+z) + 11$ K (see Fig. 6), although with large scatter. As a comparison, we also plot in the figure the relations found by other authors (Casey et al. 2013; Béthermin et al. 2015a). With this evolution we can reproduce the luminosity detection limit seen in our data (see Fig. 4), and this also explains the systematic bias between our FIR photometric redshifts and the optical redshifts z_{opt} (see §4.1).

However, our luminosity detection limit is not flat at all redshifts. This selection effect along with the temperature-luminosity relation may produce an apparent redshift-temperature evolution in our sample, with a similar trend as the one in Fig. 6, since the less luminous galaxies (which have colder dust temperatures) can only be detected at low redshift. To discard that this evolution is not produced only by selection effects, we perform the following simulation. First, we assume the dust temperature-luminosity relation of Chapman et al. (2005), which reproduces well our measurements (Fig. 5). Second, we estimate the flux density at 450 and 850 μm for each simulated source based on the dust

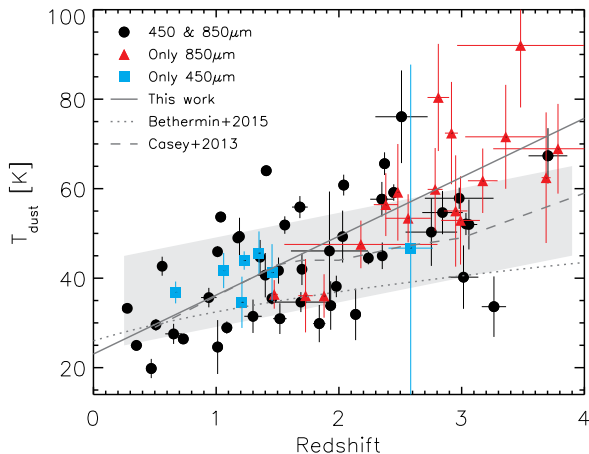


Figure 6. Dust temperature against redshift for sources with optical counterparts. The dust temperature has been estimated through a greybody SED fitting at fixed redshift with fixed $\beta = 1.6$. The dashed line represents the best linear fit for our sample, with dust temperature increasing with redshift (see discussion in §4.2.2). As it can be seen, the galaxies detected at just one single band (blue and red circles for 450 and 850 μm -only detected sources, respectively) follow the same trend that those detected at both wavelengths (black circles). For comparison, we also plot the relation found by Casey et al. (2013) and Béthermin et al. (2015a). The grey shaded region represents the apparent evolution of dust temperature with redshift originated by selection effects, specifically, due to the combination of our luminosity detection limit and the dust temperature-luminosity relation (see main text for details of the simulation). However, this effect cannot totally explain the correlation found in the data, which implies a real evolution of dust temperature with redshift.

temperature, luminosity, and redshift, using equation 1. The redshift of each source is chosen randomly (i.e. without any dependence on dust-temperature or luminosity) from a distribution that resembles those described in §4.1. Then, we impose our flux detection limit at each wavelength to consider only those sources that would be detected in our observations and, finally, we look for any relation between the dust temperature and redshift caused by selection effects. The results of these simulations are represented by the grey region in Fig. 6. Indeed, the combination of our detection limit and the dust temperature-luminosity relation introduces a trend between dust temperature and redshift, however, the relation found in the real data is much steeper (Fig. 6) and cannot be explained only by selection effects.

This implies that there is a more fundamental relation between L_{IR} , z , and T_{d} (and probably other parameters such as M_{dust} or sizes). This should be targeted in future studies of complete samples of galaxies with spectroscopic redshifts in order to understand the physical processes behind this evolution. Some possible explanations include different metallicities at different redshifts (e.g. Magnelli et al. 2014), evolution of the size of the star-forming region where galaxies at high redshifts are more compact and then hotter (e.g. Hodge et al. 2016; Zavala et al. 2018), or a combination of different effects (e.g. Cowley et al. 2017).

4.3 Stellar masses & the main-sequence

In this section we explore the connection between star formation rate and stellar mass for our galaxies. We adopt the stellar masses from the 3D-HST catalogue (Skelton et al. 2014; Momcheva et al. 2016), which are derived from SED fitting with FAST (Kriek et al. 2009) using the ‘Z-BEST’ redshifts (see §4.1). The SEDs are based on Bruzual & Charlot (2003) stellar population synthesis (SPS) models with a Chabrier (2003) IMF. For the small fraction of sources which are outside of the 3D-HST coverage (13 out of 71), we use the stellar mass reported by Barro et al. (2011b) derived using the same SPS models and the same IMF. A comparison between the stellar mass derived for the SCUBA-2 common sources in these catalogues shows that both estimations are in reasonable agreement with an average $\log(M_{*3\text{D}H\text{ST}}/M_{*\text{Barro}2011}) \approx 0.2$ (without any systematic offset).

The mean stellar mass for the galaxies with optical counterpart is $9 \pm 0.6 \times 10^{10} M_{\odot}$, and > 90 per cent of these galaxies have $M_{*} \gtrsim 1 \times 10^{10} M_{\odot}$. This is in good agreement with the estimations of the ALESS sources by da Cunha et al. (2015), with a median stellar mass of $8.9 \pm 0.1 \times 10^{10} M_{\odot}$, as well as with other studies of galaxies with similar flux densities (e.g. Aguirre et al. 2013; Simpson et al. 2014; Koprowski et al. 2016; Michałowski et al. 2017). This is also in agreement with the values found by Dunlop et al. (2017) for the sources detected in the ALMA image of the Hubble Ultra Deep Field, where ~ 80 per cent of the sources have $M_{*} \gtrsim 1 \times 10^{10} M_{\odot}$, even though the ALMA map is deeper and was done at 1.3 mm.

In Fig. 7 we plot the SFR as a function of stellar mass for sources in two redshift bins ($1.0 \leq z < 2.5$ and $2.5 \leq z < 4.0$). This shows the place occupied by our galaxies relative to the ‘main-sequence’ of star-forming galaxies, for which we adopt the parameterization reported by Speagle et al. (2014) based on a compilation of several studies from the literature covering a wide range of stellar masses up to $z \sim 6$. In the low-redshift bin ($1.0 \leq z < 2.5$), we are sensitive down to $\text{SFR} \sim 50 M_{\odot} \text{ yr}^{-1}$ and $M_{*} \sim 3 \times 10^{10} M_{\odot}$, and in the high-redshift bin down to $\text{SFR} \sim 150 M_{\odot} \text{ yr}^{-1}$. This allows us to sample the typical parameter space of the ‘main-sequence’ of star-forming galaxies (for these redshifts and stellar masses) and hence to study the nature of our galaxies in this context. It can be seen that at these redshifts, most of our galaxies (~ 85 per cent) lie on the high-mass end of the ‘main-sequence’ within a factor of $\pm 3\sigma$, where we assume as 1σ the 0.2 dex intrinsic scatter reported by Speagle et al. (2014). This result is consistent with previous studies (e.g. Michałowski et al. 2012b; Koprowski et al. 2016; Michałowski et al. 2017; Dunlop et al. 2017; Schreiber et al. 2017), and it has been predicted by some cosmological simulations (e.g. Davé et al. 2010; Narayanan et al. 2015). On the other hand, we can also represent the ‘main-sequence’ of star-forming galaxies in terms of the specific SFR ($\text{sSFR} = \text{SFR}/M_{*}$), and take into account its evolution with redshift. This is shown in Fig. 8, from which it is clearly seen that our galaxies (even those at $z < 1$) lie on this relation and follow the same redshift evolution.

Only a small fraction ($\lesssim 15$ per cent) of our faint SMGs are above the main-sequence. Actually, from the 8 sources above it three have at least another optical galaxy within our search radius, and therefore, their exceeded SFR may

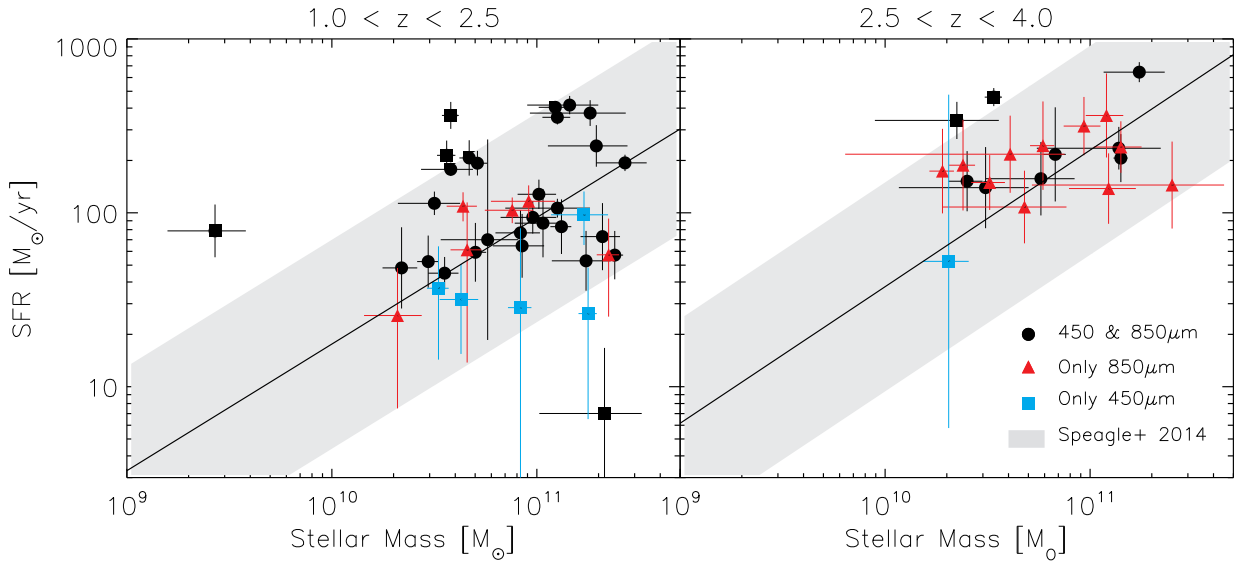


Figure 7. The SFR as a function of stellar mass for those galaxies with optical counterparts in two different redshift bins indicated at the top of the of each panel. The solid line represents the main-sequence of star-forming galaxies at $z \sim 1.5$ (left) and $z \sim 3$ (right) derived from a compilation of different studies by Speagle et al. (2014) and the shaded area shows the $\pm 3\sigma$ scatter over the defined average. As it can be seen, most SMGs lie within a factor of $\pm 3\sigma$ of the main-sequence, represented by the grey shaded region.

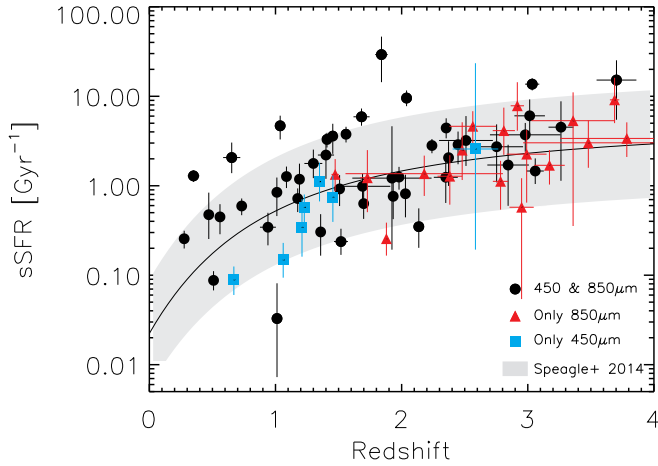


Figure 8. sSFR as a function of redshifts for all the galaxies with optical counterparts. The solid line represents the main-sequence of star-forming galaxies for $M_* = 9 \times 10^{10} M_\odot$ (Speagle et al. 2014). Most of our galaxies (~ 85 per cent) lie within a factor of $\pm 3\sigma$ of the main-sequence (indicated by the grey shaded region).

be explained by the blending of multiple sources. This fraction would be even lower if we adopt the 0.3 dex observed scatter of the ‘main-sequence’ instead of the intrinsic one (Speagle et al. 2014). This suggests that most SMGs can be fully explained as the most massive star-forming main-sequence galaxies, as also discussed by Michałowski et al. 2017.

4.4 The IRX- β relation

The Infrared excess ($\text{IRX} = \log[L_{\text{IR}}/L_{\text{UV}}]$) to UV-slope (β , $f_\lambda \propto \lambda^\beta$) relation, links the amount of dust absorption measured in the UV to the amount of infrared re-emission. It has been shown that local starburst galaxies follow a tight IRX- β relation (e.g. Meurer et al. 1999; Calzetti et al. 2000), which may be used to infer dust obscuration, and hence total SFRs, when IR data is not available. However, this relation is uncertain at high redshifts, with some studies reporting galaxies consistent with the local relation (e.g. Reddy et al. 2012; To et al. 2014; Bourne et al. 2017) while others showing sources above (e.g. Penner et al. 2012; Oteo et al. 2013) and below (e.g. Capak et al. 2015; Álvarez-Márquez et al. 2016; Bouwens et al. 2016; Pope et al. 2017) of this relationship. These high-redshift studies have been done with samples selected with different criteria (e.g. UV-selected vs IR-selected galaxies), which could introduce some extra effects (Casey et al. 2014b). Here we use our deep survey to constrain the dust absorption properties of faint SMGs in the context of the IRX- β relation. For this analysis we limit the sample to those sources detected within the 3D-HST catalogue, which provides estimations for the UV spectral slope. This quantity is determined from a power-law fit of the form $f_\lambda \propto \lambda^\beta$ (Skelton et al. 2014). The IRX is determined as the ratio between the SFR determined from the IR data (see §4.2.1) and the SFR derived from the rest-frame 1600 luminosity (not corrected for dust attenuation, Skelton et al. 2014).

Fig. 9 shows the locus occupied by our galaxies in the IRX- β plane, colour-coded according to their IR luminosities. For comparison, we also plot the Meurer et al. (1999) relation usually adopted for starburst galaxies (see also Calzetti et al. 2000) and its revised version corrected for aperture effects (Takeuchi et al. 2012), along with the relation found in the Small Magallanic Cloud (SMC;

Pettini et al. 1998). This figure suggests that there is a variation in the IRX- β relation as a function of IR luminosity (\propto SFR), where galaxies with $L_{\text{IR}} \gtrsim 1 \times 10^{12} L_{\odot}$ (the most luminous objects) lie above of the corrected starburst relation, while source with $L_{\text{IR}} \lesssim 1 \times 10^{12} L_{\odot}$ preferentially lie below it, in better agreement with the SMC relation (although with a large scatter). In Figure 9, we also plot as comparison other IR-detected galaxies with IR luminosity similar to our less luminous sources (Capak et al. 2015; Pope et al. 2017), which show similar results. The same trend as a function of IR luminosity was reported by Casey et al. (2014a) with a break luminosity of $\approx 10^{11.5}$, and by da Cunha et al. (2015) who show tentative evidence of lower dust attenuation for the lowest luminous galaxies. Exploring other dependencies, we also find a variation as a function of dust temperature and sSFR, which is not surprising given the correlation between these parameters and the IR luminosity.

The explanation for this deviation from the nominal relation is not entirely clear. The interpretations of different observational results and some theoretical models span different scenarios. For example, a more abundant population of young O and B stars for the bluer galaxies (Casey et al. 2014b), dust composition and enrichment (Mancini et al. 2016; Safarzadeh et al. 2016), geometry effects (da Cunha et al. 2015), different star formation histories (Kong et al. 2004), or a combination of these factors (Salmon et al. 2016). Interestingly, some of our galaxies lie below the SMC relation, even when this curve is generally thought to be a limiting case for star-forming galaxies (e.g. Pettini et al. 1998). This could be evidence of older stellar populations in these objects (e.g. Kong et al. 2004; Popping et al. 2017), however, this scenario is not suitable for the $z \sim 5 - 6$ galaxies reported by Capak et al. (2015) which are expected to be dominated by very young systems. In any case, this implies a non-universal IRX- β relation.

4.5 Morphologies

We use the morphological classification derived by van der Wel et al. (2012) and Huertas-Company et al. (2015a) in order to study the morphological properties of our sample. These studies consist of two-dimensional axisymmetric Sérsic models fitted with GALFIT (van der Wel et al. 2012), and visual-like classifications estimated using Convolutional Neural Networks (Huertas-Company et al. 2015a). These methods have been applied to the 5 CANDELS fields (Grogin et al. 2011) using the HST H_{160} -band, which probes the optical rest-frame at the typical redshifts of our sources, and therefore, the mass-dominant component, rather than the high surface-brightness features which commonly dominate in the UV.

The visual-like classification of Huertas-Company et al. (2015a), which is based on neural networks trained to reproduce the visual morphologies published by Kartaltepe et al. (2015), are able to predict the results of experts classifiers with a bias close to zero and a ~ 10 per cent scatter, with a misclassification fraction less than 1 per cent. Each galaxy has five associated real numbers which correspond to the frequencies at which expert classifiers would have flagged the galaxy as having a spheroid, having a disk, presenting an irregularity, being compact or a point source, and being

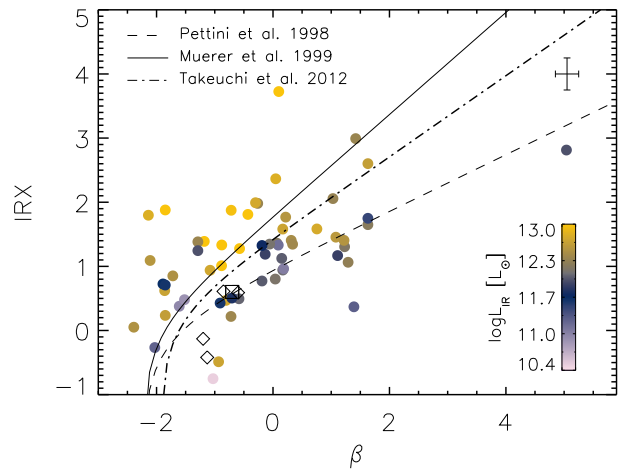


Figure 9. Infrared excess (IRX) versus the UV-continuum slope (β) for galaxies with optical counterpart. The points are colour-coded by their infrared luminosity. The solid line shows the original IRX- β relation for starburst galaxies (Meurer et al. 1999), which was then aperture-corrected by Takeuchi et al. (2012) (dot-dashed line), while the dashed line represents the SMC extinction curve (Pettini et al. 1998). The typical uncertainty of our values is represented on the top right-hand corner. The small diamonds represent ALMA observations of $z \sim 5$ galaxies with similar luminosities ($L_{\text{FIR}} \approx 3 \times 10^{11} L_{\odot}$, Capak et al. 2015), and the square is the measurement of a lensed galaxy at $z \approx 4.1$ with $L_{\text{FIR}} \approx 1 \times 10^{11} L_{\odot}$ detected by the LMT (Pope et al. 2017).

unclassifiable. Based on these numbers, we show in Fig. 10 the probability of our galaxies to be classified as spheroidal, disk, or irregular (the probability of being classified as point source or unclassifiable is negligible for most of our sources). As it can be seen, most of our galaxies have a larger probability of being classified as a disk at all redshifts, but we note an interesting transition at $z \sim 1.4$ with galaxies at high-redshifts showing also a significant probability of being classified as irregulars, or as spheroidals at lower redshifts.

Following the classification example of Huertas-Company et al. (2015b) we defined the following morphological classes based on the aforementioned probabilities:

- Pure disks: $f_{\text{disk}} > 2/3$ AND $f_{\text{sph}} < 2/3$ AND $f_{\text{irr}} < 2/3$
- Pure bulges: $f_{\text{disk}} < 2/3$ AND $f_{\text{sph}} > 2/3$ AND $f_{\text{irr}} < 2/3$
- Irregular disks: $f_{\text{disk}} > 2/3$ AND $f_{\text{sph}} < 2/3$ AND $f_{\text{irr}} > 2/3$
- Disk+bulges: $f_{\text{disk}} > 2/3$ AND $f_{\text{sph}} > 2/3$ AND $f_{\text{irr}} < 2/3$
- Irregulars/mergers: $f_{\text{disk}} < 2/3$ AND $f_{\text{sph}} < 2/3$ AND $f_{\text{irr}} > 2/3$

where f_{disk} , f_{sph} , f_{irr} , are the probabilities of being classified as a disk, spheroidal, or irregular, respectively. Based on this scheme, we can understand the transition in Fig. 10 as a high fraction of irregular disks at high redshifts and a high fraction of disks+bulges at low redshift. This is clearly seen in Fig. 11, where we plot the fraction of pure disks, irregular disks, and disk+bulges, as a function of redshift (we have not found pure bulges or irregular/mergers in our sample). From $z \approx 3$ to 0.5 the galaxies move from being dominated by irregular disks to being dominated by disks+bulges.

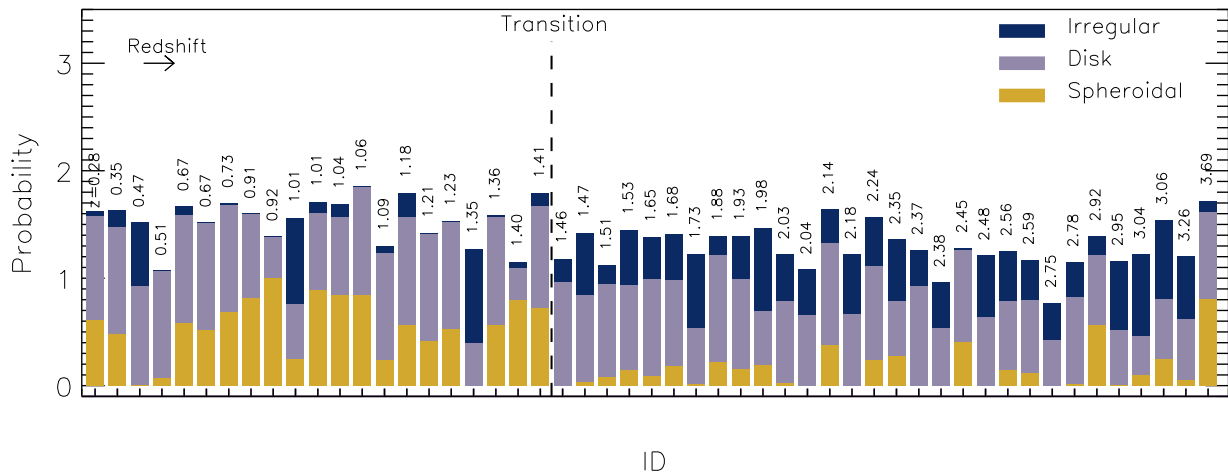


Figure 10. Probability of flagging a galaxy as spheroidal (gold), disk (purple), or irregular (blue), based on a visual-like classification in the H_{160} -band estimated through a neural network algorithm (Huertas-Company et al. 2015a). The probabilities may add up to more than 1 because the categories are not mutually exclusive. The redshift of each galaxy is indicated in the top of each bar and galaxies are sorted from low to high redshift. Most of the sources have a higher probability of being classified as disks. However, the higher redshift sources ($z \gtrsim 1.4$) show also a significant probability of having some irregularities, while the lowest redshift galaxies ($z \lesssim 1.4$) have also a significant probability of having a spheroidal component. This may be interpreted as evidence of structural evolution of SMGs and supports the idea that these objects are the progenitors of massive elliptical galaxies.

If these galaxies follow the same evolutionary path, this implies a morphological transformation from high to low redshift which results in the growth of a bulge component. This results supports, morphologically, the scenario of SMGs as progenitors of present-day massive elliptical galaxies. This has been inferred before by previous studies based on indirect evidence as the number space density, the star formation history and stellar masses, clustering, etc. (e.g. Lilly et al. 1999; Smail et al. 2002; Takagi et al. 2004; Swinbank et al. 2006; Hickox et al. 2012; Simpson et al. 2014). Similar morphological evolution has also been reported in the studies of optically detected massive galaxies ($M_* \gtrsim 1 \times 10^{11} M_\odot$, e.g. Mortlock et al. 2013; Bruce et al. 2014; Huertas-Company et al. 2015b), which supports the scenario that SMGs can be explained as the most massive star-forming main-sequence galaxies (e.g. Michałowski et al. 2017).

One important caveat is that surface brightness and other redshift issues, such as the sampling of different rest-frame wavelengths or the evolution of the galaxies' apparent sizes, may affect the morphological classification, particularly the detection of bulges at high redshifts (which could lead us to a similar picture to the one presented in Fig. 11). However, Huertas-Company et al. (2015a) have shown that their method is able to reproduce the visual classification with a bias close to zero up to (at least) $z \sim 2.7$. Furthermore, the transition occurs abruptly between $z \approx 1-2$, where the H_{160} -band is still sensitive to the optical rest-frame from older stars.

Similar results can be obtained using the parameters of the best-fitting Sérsic models. Structural properties (i.e. effective radii and Sérsic indices; Sérsic 1963) are taken from the public catalogue released by van der Wel et al. (2012),

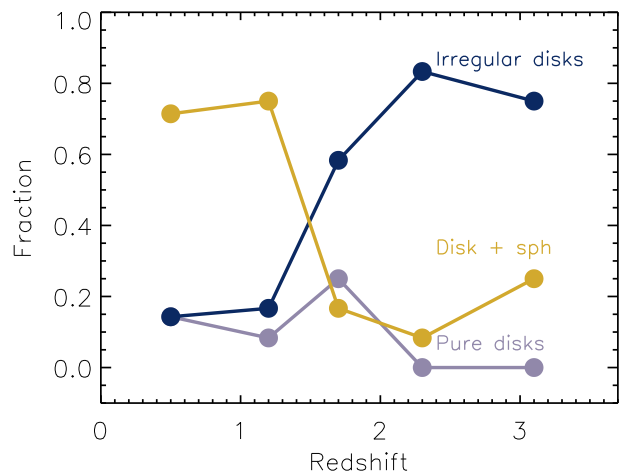


Figure 11. Fraction of galaxies classified as ‘pure disks’, ‘irregular disks’ and ‘disks + spheroids’ (see definitions in the text) based on the visual-like H_{160} -band classification of Huertas-Company et al. (2015a). Although most of the galaxies show a higher probability of being classified as disks, there is a transition from ‘irregular disks’ to disks with a spheroidal component at $z \sim 1.4$.

which were obtained with GALFIT and H_{160} -band images. The median Sérsic index of our sample is $n = 1.4^{+0.3}_{-0.1}$ (see Sérsic index distribution in Fig. 12) with a median half-light radius of $r_{1/2} = 4.8 \pm 0.4$ kpc, where the errors have been estimated via bootstrap resampling. This is in good agreement with the results found by Targett et al. (2013)

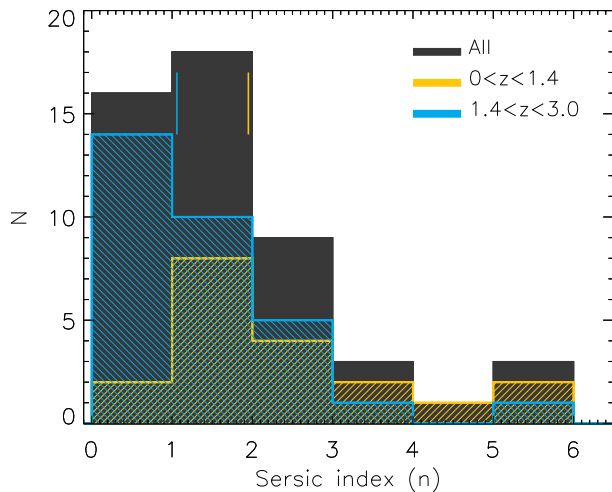


Figure 12. The Sérsic index distribution for our sources with optical counterparts derived from a two-dimensional modelling of the HST H_{160} -band images (van der Wel et al. 2012). The sources have been divided into two sub-samples: galaxies with $z \leq 1.4$ are represented by the blue histogram and the sources with $1.4 < z < 3$ by the yellow one. The vertical lines represent the median value for each sub-sample, where the lowest redshift galaxies show larger sersic index.

and Chen et al. (2015), who also studied the morphology of SMGs in NIR bands. They reported median values of $n = 1.0 \pm 0.1$ and $r_{1/2} = 4.1 \pm 0.5$ kpc, and $n = 1.2 \pm 0.3$ and $r_{1/2} = 4.4^{+1.1}_{-0.5}$ kpc, respectively. These values support the conclusion derived from the visual-like classification, that the dominant component in our galaxies is disc-like, as discussed by Targett et al. (2013).

In order to study the morphological transition described above (Fig. 11), we estimate the median Sérsic index in two redshift bins. The galaxies at higher redshifts ($z > 1.4$) have a median value of $n \approx 1$, while the low-redshift galaxies ($z \leq 1.4$) show a larger value of $n \approx 2$. This increases in the median Sérsic index from high to low redshift also suggests a morphological transformation, which implies the formation of a bulge component. We emphasize that this conclusion is supported by two completely independent morphological classification methods and represents morphological evidence for an evolutionary link in the direction of SMGs evolving towards local elliptical galaxies. The mechanism of bulge formation, however, cannot be deduced from our data, since we cannot discriminate between gas rich major mergers and violent disk instabilities, or other scenarios like minor mergers.

5 CONCLUSIONS

We have used the rich multi-wavelength dataset available in the EGS field to study the physical properties of faint submillimeter galaxies detected in deep SCUBA-2 observations at 450 and 850 μm (Zavala et al. 2017). This survey, together with other similar deep SCUBA-2 maps, represents one of the deepest blank-field observations achieved with a

single-dish telescope, which allows us to detect galaxies with $\text{SFR} \geq 150 M_{\odot} \text{ yr}^{-1}$ at all redshifts. Our sample consists of both, 450 and 850 μm -selected galaxies, which increases the completeness and minimizes the selection effects associated with single-band observations (as discussed by Casey et al. 2013). Additionally, this allows us to explore the physical differences between these populations. The final sample includes 95 sources: 56 (59 %) are detected at both wavelengths, 31 (33 %) are detected only at 850 μm , and 8 (8 %) are detected only at 450 μm .

Using radio and IR observations as an intermediate step to associate our SMGs with optical counterparts, we achieved a robust association rate of 75 %. Combining optical photometric and spectroscopic redshifts (Barro et al. 2011b; Skelton et al. 2014; Momcheva et al. 2016) with FIR photometric redshifts for those galaxies with no counterpart (see §4.1) we derived the redshift distributions for both 450 and 850 μm -selected galaxies, which have a median redshift of 1.66 ± 0.18 and 2.30 ± 0.20 , respectively. This supports the scenario in which galaxies selected at different wavelengths have different redshift distributions (see also Zavala et al. 2014; Béthermin et al. 2015b). Actually, all (except one) of the galaxies detected only at 450 μm (i.e. with no 850 μm detection) lie at $z \lesssim 1.5$, while all the galaxies detected only at 850 μm lie at $z \gtrsim 1.5$.

For those galaxies with optical counterparts we derived infrared luminosities, star formation rates, and dust temperatures. These, combined with their stellar masses, allow us to study the locus occupied by our galaxies relative to the ‘main-sequence’ of star forming galaxies. Additionally, we also studied the dust absorption properties of these galaxies (in the context of the IRX- β relation) as well as the optical rest-frame morphology. The main conclusions from these analyses are:

- The mean IR luminosity of our sample is $1.5 \pm 0.2 \times 10^{12} L_{\odot}$, which corresponds to a mean SFR of $150 \pm 20 M_{\odot} \text{ yr}^{-1}$, covering a range from $L_{\text{IR}} = 0.3 \times 10^{11} - 7.4 \times 10^{12} L_{\odot}$.
- The mean dust temperature for the whole sample is $\langle T_d \rangle = 47 \pm 15$ K, and temperature monotonically increases towards high redshift (see §4.2.2). This $T_d - z$ relation is supported by: (1) higher dust temperatures at higher redshift (Fig. 6); (2) the luminosity detection limit seen in our data (see Fig. 4) which is well fitted when we consider this evolution, and; (3) a trend where the FIR photometric redshifts are systematically lower than the optical ones at $z_{\text{opt}} > 2$ if a single dust temperature SED is assumed, which can be translated into a difference between the real dust temperature of the source and that of the SED (as reported before by Ivison et al. 2016).
- We found that most of our galaxies (~ 85 per cent) lie on the high-mass end of the ‘main-sequence’ within a factor of $\pm 3\sigma$, with a mean stellar mass of $9 \pm 0.6 \times 10^{10} M_{\odot}$.
- In the context of dust absorption, we found a variation in the IRX- β relation as a function of IR luminosity, where galaxies with $L_{\text{IR}} \gtrsim 1 \times 10^{12} L_{\odot}$ lie on (or slightly above for the most luminous objects) the Meurer et al. (1999) relation, while sources with $L_{\text{IR}} \lesssim 1 \times 10^{12} L_{\odot}$ are in better agreement with the SMC relation. A similar behaviour can be found as a function of dust temperature and sSFR.
- Based on visual-like classifications of the HST H_{160} -band (Huertas-Company et al. 2015a) we concluded that the

dominant component for most of our galaxies is a disk-like structure, although we note an interesting transition at $z \sim 1.4$. Galaxies move from being irregular disks-dominated at high redshifts to disks+bulges at low redshift, which morphologically supports the scenario of SMGs being the progenitors of present-day massive elliptical galaxies. Similar conclusion were obtained using two-dimensional axisymmetric Sérsic models (van der Wel et al. 2012), where the highest redshift galaxies have median Sérsic index of $n \approx 1$ while the lowest redshift galaxies show a larger value of $n \approx 2$.

ACKNOWLEDGEMENTS

We would like to thank the anonymous referee for a detailed report that has increased the clarity of the paper. In addition, we thank to James Simpson for his suggestions along this work. This research has been supported by Mexican CONACyT research grant CB-2011-01-167291. JAZ acknowledges support from a CONACyT studentship and from the University of Texas at Austin College of Natural Sciences, and the hospitality at the University of Edinburgh. The James Clerk Maxwell Telescope has historically been operated by the Joint Astronomy Centre on behalf of the Science and Technology Facilities Council of the United Kingdom, the National Research Council of Canada, and the Netherlands Organisation for Science Research. Additional funds for the construction of SCUBA-2 were provided by the Canada Foundation for Innovation. This work is based in part on observations made with the *Spitzer Space Telescope*, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. This research has made use of data from HerMES project (<http://hermes.sussex.ac.uk/>). HerMES is a Herschel Key Programme utilising Guaranteed Time from the SPIRE instrument team, ESAC scientists and a mission scientist.

REFERENCES

- Aguirre P., Baker A. J., Menanteau F., Lutz D., Tacconi L. J., 2013, *ApJ*, **768**, 164
- Alberts S., et al., 2013, *MNRAS*, **431**, 194
- Álvarez-Márquez J., et al., 2016, *A&A*, **587**, A122
- Aretxaga I., Hughes D. H., Chapin E. L., Gaztañaga E., Dunlop J. S., Ivison R. J., 2003, *MNRAS*, **342**, 759
- Aretxaga I., et al., 2007, *MNRAS*, **379**, 1571
- Ashby M. L. N., et al., 2006, *ApJ*, **644**, 778
- Barger A. J., Cowie L. L., Sanders D. B., Fulton E., Taniguchi Y., Sato Y., Kawara K., Okuda H., 1998, *Nature*, **394**, 248
- Barger A. J., Wang W.-H., Cowie L. L., Owen F. N., Chen C.-C., Williams J. P., 2012, *ApJ*, **761**, 89
- Barger A. J., et al., 2014, *ApJ*, **784**, 9
- Barro G., et al., 2011a, *ApJS*, **193**, 13
- Barro G., et al., 2011b, *ApJS*, **193**, 30
- Bendo G. J., et al., 2013, *MNRAS*, **433**, 3062
- Béthermin M., et al., 2015a, *A&A*, **573**, A113
- Béthermin M., De Breuck C., Sargent M., Daddi E., 2015b, *A&A*, **576**, L9
- Biggs A. D., et al., 2011, *MNRAS*, **413**, 2314
- Blain A. W., Smail I., Ivison R. J., Kneib J.-P., Frayer D. T., 2002, *Phys. Rep.*, **369**, 111
- Bourne N., et al., 2014, *MNRAS*, **444**, 1884
- Bourne N., et al., 2017, *MNRAS*, **467**, 1360
- Bouwens R. J., et al., 2016, *ApJ*, **833**, 72
- Brammer G. B., van Dokkum P. G., Coppi P., 2008, *ApJ*, **686**, 1503
- Brammer G. B., et al., 2012, *ApJS*, **200**, 13
- Brisbin D., et al., 2017, *A&A*, **608**, A15
- Bruce V. A., et al., 2014, *MNRAS*, **444**, 1001
- Bruzual G., Charlot S., 2003, *MNRAS*, **344**, 1000
- Calzetti D., Armus L., Bohlin R. C., Kinney A. L., Koornneef J., Storchi-Bergmann T., 2000, *ApJ*, **533**, 682
- Capak P. L., et al., 2015, *Nature*, **522**, 455
- Casey C. M., et al., 2012, *ApJ*, **761**, 139
- Casey C. M., et al., 2013, *MNRAS*, **436**, 1919
- Casey C. M., Narayanan D., Cooray A., 2014a, *Phys. Rep.*, **541**, 45
- Casey C. M., et al., 2014b, *ApJ*, **796**, 95
- Chabrier G., 2003, *PASP*, **115**, 763
- Chapin E. L., Hughes D. H., Aretxaga I., 2009a, *MNRAS*, **393**, 653
- Chapin E. L., et al., 2009b, *MNRAS*, **398**, 1793
- Chapin E. L., et al., 2011, *MNRAS*, **411**, 505
- Chapman S. C., Blain A. W., Smail I., Ivison R. J., 2005, *ApJ*, **622**, 772
- Chen C.-C., Cowie L. L., Barger A. J., Casey C. M., Lee N., Sanders D. B., Wang W.-H., Williams J. P., 2013, *ApJ*, **776**, 131
- Chen C.-C., et al., 2015, *ApJ*, **799**, 194
- Chen C.-C., et al., 2016, *ApJ*, **820**, 82
- Cowie L. L., Barger A. J., Hsu L.-Y., Chen C.-C., Owen F. N., Wang W.-H., 2017, *ApJ*, **837**, 139
- Cowley W. I., Béthermin M., del P. Lagos C., Lacey C. G., Baugh C. M., Cole S., 2017, *MNRAS*,
- Davé R., Finlator K., Oppenheimer B. D., Fardal M., Katz N., Kereš D., Weinberg D. H., 2010, *MNRAS*, **404**, 1355
- Dempsey J. T., et al., 2013, *MNRAS*, **430**, 2534
- Downes A. J. B., Peacock J. A., Savage A., Carrie D. R., 1986, *MNRAS*, **218**, 31
- Dunlop J. S., Peacock J. A., Savage A., Lilly S. J., Heasley J. N., Simon A. J. B., 1989, *MNRAS*, **238**, 1171
- Dunlop J. S., et al., 2017, *MNRAS*, **466**, 861
- Dunne L., Eales S. A., 2001, *MNRAS*, **327**, 697
- Farrah D., Afonso J., Efstathiou A., Rowan-Robinson M., Fox M., Clements D., 2003, *MNRAS*, **343**, 585
- Fixsen D. J., Dwek E., Mather J. C., Bennett C. L., Shafer R. A., 1998, *ApJ*, **508**, 123
- Geach J. E., et al., 2013, *MNRAS*, **432**, 53
- Geach J. E., et al., 2015, *MNRAS*, **452**, 502
- Geach J. E., et al., 2017, *MNRAS*, **465**, 1789
- Genzel R., et al., 2015, *ApJ*, **800**, 20
- Greve T. R., et al., 2005, *MNRAS*, **359**, 1165
- Griffin M. J., et al., 2010, *A&A*, **518**, L3
- Grogin N. A., et al., 2011, *ApJS*, **197**, 35
- Gruppioni C., et al., 2013, *MNRAS*, **432**, 23
- Hatsukade B., et al., 2016, *PASJ*, **68**, 36
- Hickox R. C., et al., 2012, *MNRAS*, **421**, 284
- Hodge J. A., et al., 2013, *ApJ*, **768**, 91
- Hodge J. A., et al., 2016, *ApJ*, **833**, 103
- Holland W. S., et al., 2013, *MNRAS*, **430**, 2513
- Hsu L.-Y., Cowie L., Chen C.-C., Barger A., Wang W.-H., 2016, preprint, ([arXiv:1605.00046](https://arxiv.org/abs/1605.00046))
- Huertas-Company M., et al., 2015a, *ApJS*, **221**, 8
- Huertas-Company M., et al., 2015b, *ApJ*, **809**, 95
- Hughes D. H., et al., 1998, *Nature*, **394**, 241
- Ikarashi S., et al., 2015, *ApJ*, **810**, 133
- Iono D., et al., 2006, *ApJ*, **640**, L1
- Isobe T., Feigelson E. D., Nelson P. I., 1986, *ApJ*, **306**, 490
- Ivison R. J., et al., 2007a, *MNRAS*, **380**, 199
- Ivison R. J., et al., 2007b, *ApJ*, **660**, L77
- Ivison R. J., et al., 2016, *ApJ*, **832**, 78
- Karim A., et al., 2013, *MNRAS*, **432**, 2

- Kartalpepe J. S., et al., 2015, *ApJS*, **221**, 11
- Kennicutt Jr. R. C., 1998, *ARA&A*, **36**, 189
- Kirkpatrick A., Pope A., Sajina A., Roebuck E., Yan L., Armus L., Díaz-Santos T., Stierwalt S., 2015, *ApJ*, **814**, 9
- Kong X., Charlot S., Brinchmann J., Fall S. M., 2004, *MNRAS*, **349**, 769
- Koprowski M. P., Dunlop J. S., Michałowski M. J., Cirasuolo M., Bowler R. A. A., 2014, *MNRAS*, **444**, 117
- Koprowski M. P., et al., 2016, *MNRAS*, **458**, 4321
- Kovács A., Chapman S. C., Dowell C. D., Blain A. W., Ivison R. J., Smail I., Phillips T. G., 2006, *ApJ*, **650**, 592
- Kriek M., van Dokkum P. G., Labbé I., Franx M., Illingworth G. D., Marchesini D., Quadri R. F., 2009, *ApJ*, **700**, 221
- Lacey C. G., et al., 2016, *MNRAS*, **462**, 3854
- Lagos C. d. P., Bayet E., Baugh C. M., Lacey C. G., Bell T. A., Fanidakis N., Geach J. E., 2012, *MNRAS*, **426**, 2142
- Laird E. S., et al., 2009, *ApJS*, **180**, 102
- Lilly S. J., Eales S. A., Gear W. K. P., Hammer F., Le Fèvre O., Crampton D., Bond J. R., Dunne L., 1999, *ApJ*, **518**, 641
- Lutz D., et al., 2011, *A&A*, **532**, A90
- Magdis G. E., et al., 2012, *ApJ*, **760**, 6
- Magnelli B., et al., 2012, *A&A*, **539**, A155
- Magnelli B., et al., 2014, *A&A*, **561**, A86
- Mancini M., Schneider R., Graziani L., Valiante R., Dayal P., Maio U., Ciardi B., 2016, *MNRAS*, **462**, 3130
- Markwardt C. B., 2009, *ASPC*, **411**, 251
- Marsden G., et al., 2011, *MNRAS*, **417**, 1192
- Meurer G. R., Heckman T. M., Calzetti D., 1999, *ApJ*, **521**, 64
- Michałowski M., Hjorth J., Watson D., 2010, *A&A*, **514**, A67
- Michałowski M. J., et al., 2012a, *MNRAS*, **426**, 1845
- Michałowski M. J., Dunlop J. S., Cirasuolo M., Hjorth J., Hayward C. C., Watson D., 2012b, *A&A*, **541**, A85
- Michałowski M. J., et al., 2017, *MNRAS*, **469**, 492
- Miettinen O., et al., 2015, *A&A*, **577**, A29
- Momcheva I. G., et al., 2016, *ApJS*, **225**, 27
- Mortlock A., et al., 2013, *MNRAS*, **433**, 1185
- Nandra K., et al., 2015, *ApJS*, **220**, 10
- Narayanan D., et al., 2015, *Nature*, **525**, 496
- Nguyen H. T., et al., 2010, *A&A*, **518**, L5
- Oliver S. J., et al., 2012, *MNRAS*, **424**, 1614
- Oteo I., et al., 2013, *A&A*, **554**, L3
- Penner K., et al., 2012, *ApJ*, **759**, 28
- Pettini M., Kellogg M., Steidel C. C., Dickinson M., Adelberger K. L., Giavalisco M., 1998, *ApJ*, **508**, 539
- Planck Collaboration et al., 2014, *A&A*, **571**, A16
- Poglitsch A., et al., 2010, *A&A*, **518**, L2
- Pope A., Borys C., Scott D., Conselice C., Dickinson M., Mobasher B., 2005, *MNRAS*, **358**, 149
- Pope A., et al., 2006, *MNRAS*, **370**, 1185
- Pope A., et al., 2017, preprint, ([arXiv:1703.04535](https://arxiv.org/abs/1703.04535))
- Popping G., Puglisi A., Norman C. A., 2017, *MNRAS*, **472**, 2315
- Puget J.-L., Abergel A., Bernard J.-P., Boulanger F., Burton W. B., Desert F.-X., Hartmann D., 1996, *A&A*, **308**, L5
- Reddy N., et al., 2012, *ApJ*, **744**, 154
- Riechers D. A., et al., 2013, *Nature*, **496**, 329
- Roseboom I. G., et al., 2013, *MNRAS*, **436**, 430
- Safarzadeh M., Hayward C. C., Ferguson H. C., 2016, preprint, ([arXiv:1604.07402](https://arxiv.org/abs/1604.07402))
- Salmon B., et al., 2016, *ApJ*, **827**, 20
- Schreiber C., Pannella M., Leiton R., Elbaz D., Wang T., Okumura K., Labbé I., 2017, *A&A*, **599**, A134
- Sérsic J. L., 1963, *Boletín de la Asociación Argentina de Astronomía La Plata Argentina*, **6**, 41
- Silva L., Granato G. L., Bressan A., Danese L., 1998, *ApJ*, **509**, 103
- Simpson J. M., et al., 2014, *ApJ*, **788**, 125
- Simpson J. M., et al., 2015, *ApJ*, **807**, 128
- Simpson J. M., et al., 2016, preprint, ([arXiv:1611.03084](https://arxiv.org/abs/1611.03084))
- Skelton R. E., et al., 2014, *ApJS*, **214**, 24
- Smail I., Ivison R. J., Blain A. W., 1997, *ApJ*, **490**, L5
- Smail I., Ivison R. J., Blain A. W., Kneib J.-P., 2002, *MNRAS*, **331**, 495
- Smolčić V., et al., 2012a, *ApJS*, **200**, 10
- Smolčić V., et al., 2012b, *A&A*, **548**, A4
- Speagle J. S., Steinhardt C. L., Capak P. L., Silverman J. D., 2014, *ApJS*, **214**, 15
- Swinbank A. M., Chapman S. C., Smail I., Lindner C., Borys C., Blain A. W., Ivison R. J., Lewis G. F., 2006, *MNRAS*, **371**, 465
- Swinbank A. M., et al., 2014, *MNRAS*, **438**, 1267
- Symeonidis M., et al., 2013, *MNRAS*, **431**, 2317
- Tacconi L. J., et al., 2013, *ApJ*, **768**, 74
- Takagi T., Hanami H., Arimoto N., 2004, *MNRAS*, **355**, 424
- Takeuchi T. T., Yuan F.-T., Ikeyama A., Murata K. L., Inoue A. K., 2012, *ApJ*, **755**, 144
- Targett T. A., et al., 2013, *MNRAS*, **432**, 2012
- To C.-H., Wang W.-H., Owen F. N., 2014, *ApJ*, **792**, 139
- Toft S., et al., 2014, *ApJ*, **782**, 68
- Umehata H., et al., 2017, *ApJ*, **835**, 98
- Walter F., et al., 2016, *ApJ*, **833**, 67
- Wang W.-H., Cowie L. L., Barger A. J., Williams J. P., 2011, *ApJ*, **726**, L18
- Weiß A., et al., 2009, *ApJ*, **707**, 1201
- Younger J. D., et al., 2007, *ApJ*, **671**, 1531
- Younger J. D., et al., 2009, *ApJ*, **704**, 803
- Yun M. S., Carilli C. L., 2002, *ApJ*, **568**, 88
- Yun M. S., et al., 2008, *MNRAS*, **389**, 333
- Yun M. S., et al., 2012, *MNRAS*, **420**, 957
- Zavala J. A., Aretxaga I., Hughes D. H., 2014, *MNRAS*, **443**, 2384
- Zavala J. A., et al., 2017, *MNRAS*, **464**, 3369
- Zavala J. A., et al., 2018, *Nature Astronomy*, **2**, 56
- da Cunha E., Charlot S., Elbaz D., 2008, *MNRAS*, **388**, 1595
- da Cunha E., et al., 2015, *ApJ*, **806**, 110
- van der Wel A., et al., 2012, *ApJS*, **203**, 24

APPENDIX A: TABLES

Table A1. Counterpart identifications and physical properties of galaxies in our sample. The columns give: (1) & (2) SCUBA-2 850 and 450 μm source names, respectively; (3) & (4) coordinates of the associated optical counterpart; (5) & (6) redshift and the source where it comes from[†]; (7) stellar mass; (8) IR luminosity (9) dust temperature, (10) UV star formation rate, and; (11) UV spectral slope (from Skelton et al. (2014)). The coordinates and flux densities at 450 and 850 μm can be found in Zavala et al. (2017).

850 μm -ID	450 μm -ID	RA _{opt} [hh:mm:ss.s]	DEC _{opt} [°:':"]	z	z_s^\dagger	$\log(M_*)$ [M_\odot]	$\log(L_{\text{FIR}})$ [L_\odot]	T_d [K]	SFR 1600Å [M_\odot/yr]	β
850.001	450.02	214.9109	52.9010	$2.240^{+0.001}_{-0.003}$	2	11.10 ± 0.07	12.53 ± 0.04	40.5 ± 0.9	5.5	-0.43
850.002	450.03	214.9144	52.8760	$0.73^{+0.00}_{-0.00}$	1	10.87 ± 0.10	11.62 ± 0.03	27.2 ± 0.3	30	1.11
850.003	450.05	214.9163	52.8913	$3.04^{+0.06}_{-0.06}$	2	10.53 ± 0.04	12.76 ± 0.04	52.7 ± 1.5	21.5	-0.88
850.004	450.25	214.9467	52.9101	$1.01^{+0.08}_{-0.09}$	3	8.68 ± 0.08	-	-	1.3	-1.49
850.005	450.08	214.9781	52.8115	$2.98^{+0.28}_{-0.28}$	3	11.24 ± 0.17	12.87 ± 0.05	57.4 ± 1.5	-	-
850.006	450.13	214.9741	52.9060	$2.45^{+0.12}_{-0.12}$	2	11.16 ± 0.21	12.67 ± 0.04	57.1 ± 1.2	17	-1.18
850.007	450.20	214.9624	52.8703	$3.02^{+0.01}_{-0.12}$	3	10.40 ± 0.09	12.20 ± 0.11	39.1 ± 4.2	12.3	-2.11
850.008	450.15	215.0144	52.9013	$1.176^{+0.004}_{-0.006}$	2	11.43 ± 0.14	12.21 ± 0.04	46.3 ± 0.8	595	-0.93
850.009	450.12	214.8437	52.9110	$1.45^{+0.03}_{-0.03}$	3	10.50 ± 0.18	12.07 ± 0.05	36.2 ± 0.8	68.5	-0.72
850.010	450.14	214.9197	52.8397	$1.98^{+0.29}_{-0.091}$	2	11.01 ± 0.10	12.14 ± 0.07	37.6 ± 1.5	5.0	0.31
850.011	450.39	214.9224	52.8220	$3.26^{+0.11}_{-0.10}$	3	10.49 ± 0.42	12.16 ± 0.14	34.4 ± 3.8	124	-2.39
850.012	450.06	214.9462	52.8761	$0.35^{+0.05}_{-0.05}$	3	9.80 ± 0.04	10.80 ± 0.04	24.7 ± 0.4	3.4	-1.61
850.013	450.07	-	-	$0.95^{+0.10}_{-0.00}$	4	-	-	-	-	-
850.014	450.46	214.8562	52.9287	$1.69^{+0.00}_{-0.22}$	3	10.98 ± 0.16	11.95 ± 0.08	32.3 ± 1.4	-	-
850.015	450.09	214.9381	52.8743	$1.56^{+0.00}_{-0.01}$	2	10.71 ± 0.05	12.21 ± 0.06	49.2 ± 1.3	5.1	0.17
850.016	450.56	215.0077	52.8271	$1.84^{+0.05}_{-0.05}$	3	9.43 ± 0.23	12.23 ± 0.08	38.5 ± 2.3	-	-
850.017	450.11	214.8988	52.8526	$2.37^{+0.00}_{-0.02}$	2	11.26 ± 0.29	12.63 ± 0.07	63.0 ± 1.7	19.8	-0.57
850.018	450.84	214.9748	52.8606	$3.70^{+0.14}_{-0.16}$	3	10.35 ± 0.40	12.62 ± 0.09	65.3 ± 3.9	0.1	0.10
850.019	450.41	214.9710	52.9574	$2.353^{+0.003}_{-0.003}$	2	10.67 ± 0.05	12.32 ± 0.09	42.0 ± 1.9	23.8	-1.08
850.020	450.23	215.0310	52.9165	$1.087^{+0.001}_{-0.001}$	2	10.55 ± 0.08	11.65 ± 0.08	29.9 ± 0.9	2.9	-0.13
850.021	450.17	215.0542	52.9260	$1.037^{+0.001}_{-0.002}$	2	10.58 ± 0.15	12.13 ± 0.03	51.4 ± 0.8	0.4	1.63
850.022	450.01	214.9276	52.8480	$1.408^{+0.046}_{-0.000}$	2	11.09 ± 0.08	12.50 ± 0.03	61.9 ± 0.9	5.3	-1.85
850.023	450.65	-	-	> 3.95	4	-	-	-	-	-
850.024	450.04	214.9232	52.8822	$1.01^{+0.000}_{-0.000}$	1	11.10 ± 0.25	11.95 ± 0.05	42.7 ± 0.8	1.1	-0.26
850.025	450.29	214.8506	52.8664	$2.14^{+0.03}_{-0.01}$	2	11.32 ± 0.11	11.88 ± 0.13	31.7 ± 3.4	11.5	0.039
850.026	450.31	-	-	$2.25^{+0.35}_{-0.30}$	4	-	-	-	-	-
850.028	450.69	214.8763	52.8521	$2.84^{+0.17}_{-0.17}$	2	11.14 ± 0.40	12.37 ± 0.11	49.3 ± 3.0	3.8	-2.14
850.029	450.21	214.8354	52.8438	$2.039^{+0.001}_{-0.003}$	2	10.58 ± 0.04	12.50 ± 0.06	60.9 ± 1.6	35.5	-0.88
850.030	450.10	214.8780	52.8768	$1.52^{+0.06}_{-0.06}$	3	11.38 ± 0.05	11.87 ± 0.09	33.9 ± 2.0	9.7	-0.19
850.031	450.49	-	-	> 3.65	4	-	-	-	-	-
850.032	450.63	214.9328	52.8330	$3.06^{+0.02}_{-0.07}$	2	11.15 ± 0.04	12.40 ± 0.11	51.5 ± 3.4	70	-0.81
850.033	450.59	-	-	$3.10^{+2.10}_{-0.75}$	4	-	-	-	-	-
850.034	450.66	-	-	> 3.6	4	-	-	-	-	-
850.037	450.64	215.0502	52.8539	$1.30^{+0.07}_{-0.07}$	3	10.47 ± 0.05	11.85 ± 0.11	33.8 ± 2.0	-	-
850.038	450.27	214.8647	52.8991	$1.93^{+0.09}_{-0.01}$	2	10.92 ± 0.14	11.93 ± 0.12	36.4 ± 3.2	7.4	0.17
850.039	450.40	215.0389	52.8544	$2.35^{+0.10}_{-0.10}$	3	11.29 ± 0.24	12.43 ± 0.10	56.5 ± 2.5	-	-
850.040	450.22	-	-	$3.55^{+2.00}_{-0.90}$	4	-	-	-	-	-
850.041	450.16	214.9928	52.8508	$1.68^{+0.06}_{-0.06}$	3	10.56 ± 0.05	12.25 ± 0.07	54.4 ± 1.7	2.2	-0.3
850.042	450.26	214.9800	52.9025	$2.03^{+0.05}_{-0.02}$	2	11.03 ± 0.14	11.91 ± 0.16	46.1 ± 3.6	2.0	1.63
850.043	450.73	214.9501	52.9383	$1.51^{+0.01}_{-0.01}$	2	10.92 ± 0.12	11.88 ± 0.11	40.1 ± 1.8	3.8	1.23
850.046	450.38	-	-	$1.60^{+0.15}_{-0.15}$	4	-	-	-	-	-
850.052	450.76	214.8278	52.9049	$1.70^{+0.04}_{-0.03}$	3	11.12 ± 0.05	11.98 ± 0.13	43.4 ± 2.2	3.6	0.07
850.059	450.24	214.8555	52.8489	$2.75^{+0.24}_{-0.21}$	3	10.76 ± 0.27	12.30 ± 0.16	53.2 ± 4.6	5.5	1.08
850.060	450.75	214.8495	52.9381	$0.94^{+0.06}_{-0.06}$	3	11.08 ± 0.13	11.56 ± 0.11	35.5 ± 1.4	-	-
850.065	450.18	214.8755	52.8665	$0.28^{+0.02}_{-0.02}$	3	10.56 ± 0.09	10.84 ± 0.05	30.7 ± 0.6	1.0	0.18
850.069	450.44	214.9548	52.8766	$0.51^{+0.00}_{-0.00}$	1	11.21 ± 0.06	11.03 ± 0.07	29.6 ± 0.8	6.1	1.39
850.070	450.34	214.9684	52.9251	$0.65^{+0.01}_{-0.01}$	2	9.740 ± 0.04	10.99 ± 0.14	26.8 ± 1.4	0.5	0.09
850.073	450.52	214.9480	52.8407	$1.358^{+0.001}_{-0.001}$	2	11.24 ± 0.17	11.67 ± 0.14	43.3 ± 2.5	0.1	5.04
850.078	450.82	215.0377	52.8708	$0.56^{+0.03}_{-0.04}$	3	10.65 ± 0.16	11.17 ± 0.08	39.4 ± 1.4	0.4	1.63
850.079	450.54	215.0296	52.9364	$0.47^{+0.03}_{-0.03}$	3	9.810 ± 0.09	10.41 ± 0.20	20.2 ± 1.5	17.3	-1.03
850.085	450.33	215.0158	52.8569	$1.40^{+0.07}_{-0.06}$	3	10.34 ± 0.09	11.64 ± 0.19	39.3 ± 3.1	-	-
850.092	450.60	214.8852	52.8158	$1.92^{+0.06}_{-0.31}$	2	10.76 ± 0.23	11.77 ± 0.47	43.3 ± 8.2	3.2	-0.06
850.095	450.37	214.9965	52.8141	$1.19^{+0.03}_{-0.03}$	3	10.70 ± 0.06	11.66 ± 0.14	49.2 ± 3.0	-	-
850.097	450.71	214.9206	52.8659	$1.01^{+0.000}_{-0.000}$	1	11.33 ± 0.32	10.95 ± 0.26	27.7 ± 3.3	2.3	-1.52
850.102	450.83	-	-	$2.35^{+3.10}_{-0.95}$	4	-	-	-	-	-
850.104	450.45	214.9866	52.8021	$2.5^{+0.22}_{-0.21}$	3	10.83 ± 0.06	12.30 ± 0.24	76.1 ± 7.7	-	-

[†](1) spectroscopic redshift; (2) HST grism redshift; (3) optical photometric redshift; (4) FIR photometric redshift (see §4.1).

Table A1. (continuation)

850 μ m-ID	450 μ m-ID	RA _{opt} [hh:mm:ss.s]	DEC _{opt} [°:′:″]	z	z_s	M_* [M_\odot]	L_{FIR} [L_\odot]	T_d [K]	SFR1600 [M_\odot/yr]	β
850.027	-	-	-	> 5.3	4	-	-	-	-	-
850.035	-	-	-	2.05 ^{+0.90} _{-0.30}	4	-	-	-	-	-
850.036	-	-	-	> 3.85	4	-	-	-	-	-
850.044	-	215.0349	52.8914	3.17 ^{+0.14} _{-0.12}	3	11.15 ± 0.13	12.40 ± 0.12	57.0 ± 4.5	1.0	0.05
850.047	-	214.8598	52.8607	3.79 ^{+0.56} _{-0.53}	3	10.97 ± 0.10	12.68 ± 0.12	75.9 ± 6.1	4.2	-0.72
850.048	-	215.0454	52.8946	1.89 ^{+0.01} _{-0.01}	2	11.35 ± 0.06	11.74 ± 0.15	33.7 ± 3.1	4.3	0.15
850.049	-	-	-	> 4.75	4	-	-	-	-	-
850.050	-	214.9968	52.9408	2.56 ^{+0.41} _{-0.19}	3	10.51 ± 0.09	12.23 ± 0.13	52.6 ± 3.4	2.6	0.22
850.051	-	214.9951	52.9062	2.78 ^{+0.06} _{-0.06}	3	11.09 ± 0.19	12.25 ± 0.16	60.9 ± 5.8	19.4	-1.72
850.053	-	214.9935	52.8644	2.99 ^{+0.15} _{-0.16}	3	10.68 ± 0.40	12.12 ± 0.16	52.7 ± 6.5	0.1	1.42
850.054	-	214.8747	52.8438	1.47 ^{+0.03} _{-0.03}	3	10.66 ± 0.08	11.73 ± 0.14	33.2 ± 2.0	19.6	-0.58
850.055	-	-	-	> 4.5	4	-	-	-	-	-
850.056	-	215.0273	52.8946	2.48 ^{+0.000} _{-0.01}	2	10.64 ± 0.08	12.04 ± 0.27	53.0 ± 6.3	4.52	-1.29
850.057	-	214.9969	52.8892	3.69 ^{+0.03} _{-0.04}	2	10.28 ± 0.06	12.22 ± 0.19	56.4 ± 8.8	101	-1.85
850.058	-	-	-	> 5.0	4	-	-	-	-	-
850.061	-	-	-	2.70 ^{+1.40} _{-0.65}	4	-	-	-	-	-
850.062	-	215.0618	52.9010	2.18 ^{+0.01} _{-0.62}	2	10.88 ± 0.13	12.03 ± 0.16	45.6 ± 3.3	0.9	1.03
850.063	-	-	-	1.90 ^{+0.65} _{-0.35}	4	-	-	-	-	-
850.064	-	-	-	> 4.4	4	-	-	-	-	-
850.066	-	-	-	> 5.0	4	-	-	-	-	-
850.067	-	214.8293	52.8940	2.95 ^{+0.10} _{-0.10}	3	11.40 ± 0.68	12.41 ± 0.18	61.4 ± 7.0	5.7	1.22
850.068	-	-	-	> 4.15	4	-	-	-	-	-
850.071	-	-	-	2.70 ^{+2.00} _{-0.75}	4	-	-	-	-	-
850.072	-	215.0159	52.9395	3.36 ^{+0.36} _{-0.33}	3	10.61 ± 0.81	12.46 ± 0.19	71.4 ± 7.5	5.7	0.75
850.074	-	-	-	3.70 ^{+2.03} _{-1.95}	4	-	-	-	-	-
850.075	-	214.8340	52.8701	2.38 ^{+0.10} _{-0.10}	3	10.96 ± 0.18	12.17 ± 0.18	57.6 ± 4.5	5.3	0.34
850.077	-	215.0297	52.8679	3.48 ^{+0.52} _{-0.52}	3	11.08 ± 0.10	12.52 ± 0.21	80.2 ± 8.4	-	-
850.080	-	214.9231	52.9347	2.92 ^{+0.04} _{-0.05}	3	10.38 ± 0.07	12.23 ± 0.23	62.7 ± 6.7	44.9	-1.86
850.081	-	214.9546	52.8991	1.73 ^{+0.05} _{-0.14}	2	10.32 ± 0.16	11.36 ± 0.25	32.9 ± 5.0	9.6	-0.91
850.082	-	215.0557	52.8820	2.81 ^{+0.09} _{-0.09}	3	10.77 ± 0.06	12.34 ± 0.22	73.6 ± 7.7	-	-
850.083	-	-	-	2.55 ^{+2.75} _{-0.85}	4	-	-	-	-	-
-	450.19	214.9169	52.8274	1.228 ^{+0.000} _{-0.002}	2	11.23 ± 0.16	11.89 ± 0.07	43.8 ± 1.4	8.4	1.29
-	450.28	214.9459	52.8942	1.210 ^{+0.005} _{-0.003}	2	10.92 ± 0.06	11.51 ± 0.24	35.9 ± 3.2	5.4	-1.89
-	450.35	214.9024	52.8637	1.061 ^{+0.000} _{-0.000}	1	11.25 ± 0.05	11.35 ± 0.16	39.7 ± 2.6	5.2	-1.85
-	450.36	214.9361	52.8559	1.455 ^{+0.008} _{-0.005}	2	10.63 ± 0.10	11.56 ± 0.19	40.8 ± 3.8	1.5	-0.19
-	450.42	214.9053	52.8510	0.670 ^{+0.005} _{-0.000}	1	11.18 ± 0.09	11.04 ± 0.12	34.6 ± 1.6	24.8	-2.03
-	450.43	215.0204	52.9298	2.59 ^{+0.21} _{-0.18}	3	10.31 ± 0.13	11.52 ± 0.55	35.7 ± 3.4	3.0	-1.29
-	450.47	-	-	1.45 ^{+2.60} _{-1.35}	4	-	-	-	-	-
-	450.48	214.9087	52.9279	1.349 ^{+0.002} _{-0.002}	2	10.52 ± 0.06	11.53 ± 0.18	43.5 ± 3.1	11.5	-0.71

APPENDIX B: POSTAGE STAMPS

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.

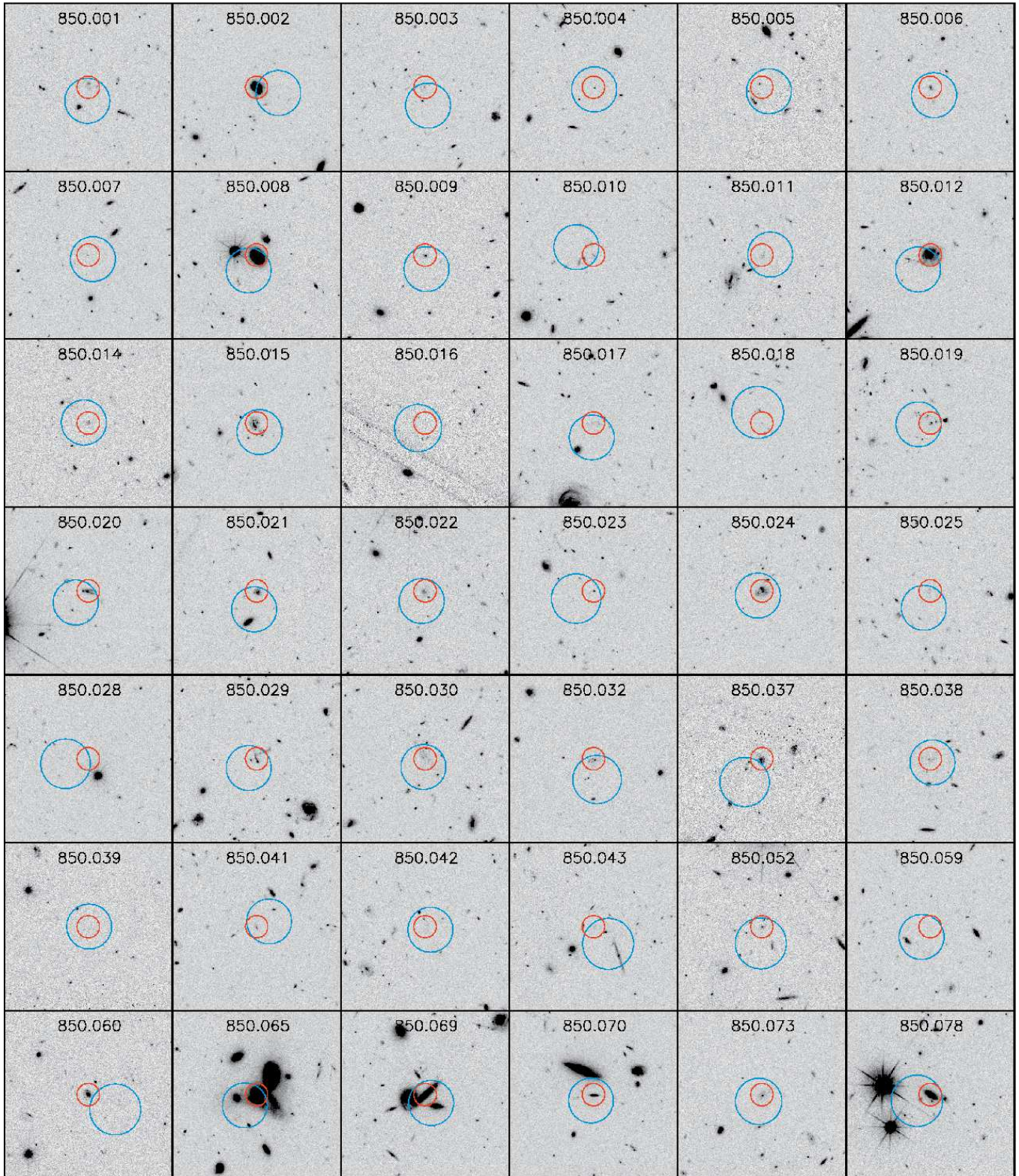


Figure B1. $30'' \times 30''$ postage stamps in the HST/f814w band for our SMGs. The blue circle represents the variable search radius for counterparts in the radio and IR wavelengths (centered at the SCUBA-2 position), and the red circle represents the final optical associated galaxy. The ID of each source is in the top of each panel.

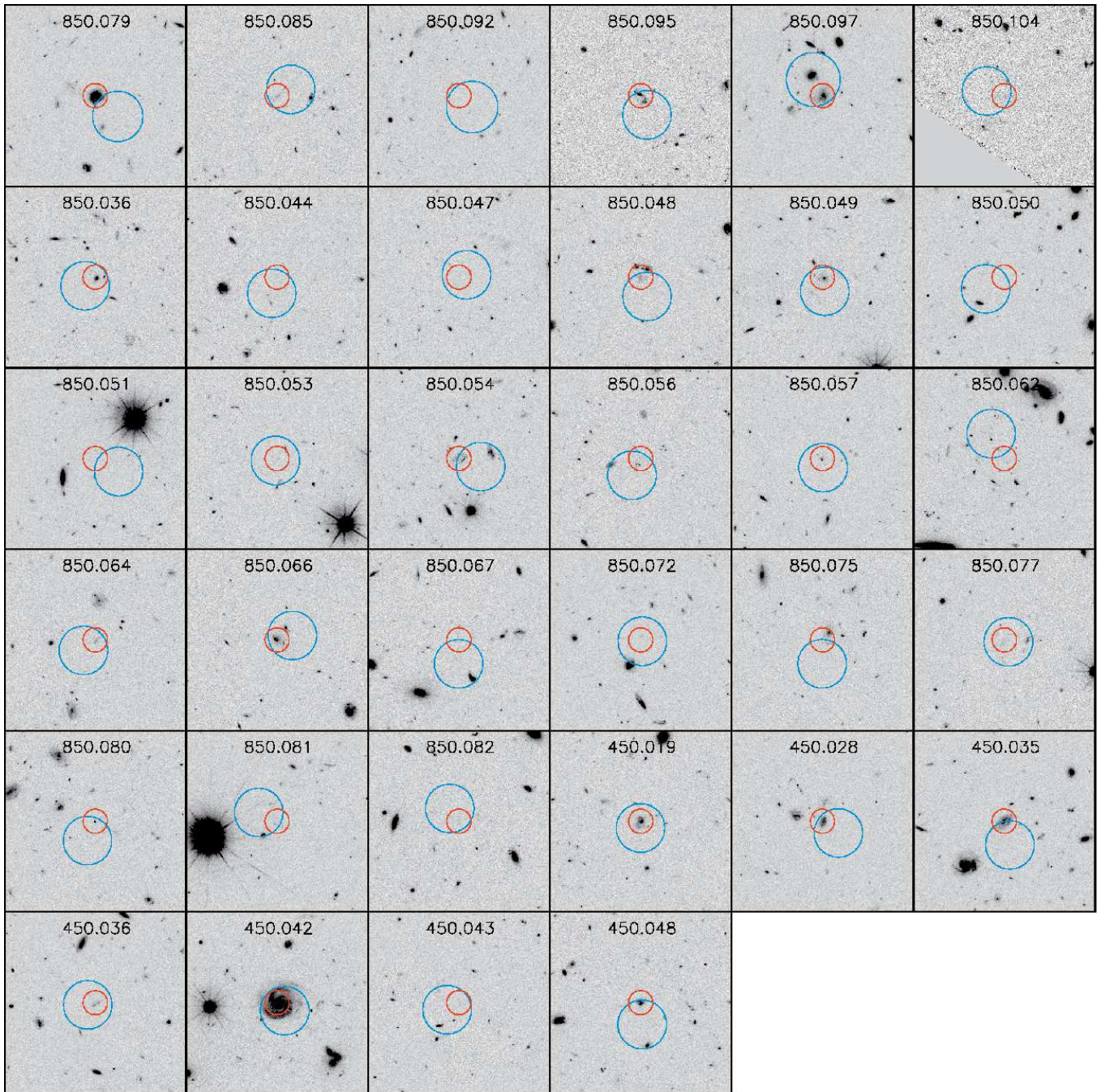


Figure B1. (continuation)