The Atacama Cosmology Telescope: Relation between Galaxy Cluster Optical Richness and Sunyaev-Zel'dovich Effect

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THE ATACAMA COSMOLOGY TELESCOPE: RELATION BETWEEN GALAXY CLUSTER OPTICAL RICHNESS AND SUNYAEV–ZEL'DOVICH EFFECT

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ABSTRACT

We present the measured Sunyaev–Zel'dovich (SZ) flux from 474 optically selected MaxBCG clusters that fall within the Atacama Cosmology Telescope (ACT) Equatorial survey region. The ACT Equatorial region used in this analysis covers 510 deg^2 and overlaps Stripe 82 of the Sloan Digital Sky Survey. We also present the measured SZ flux stacked on 52 X-ray-selected MCXC clusters that fall within the ACT Equatorial region and an ACT Southern survey region covering 455 deg^2. We find that the measured SZ flux from the X-ray-selected clusters is consistent with expectations. However, we find that the measured SZ flux from the optically selected clusters is both significantly lower than expectations and lower than the recovered SZ flux measured by the *Planck* satellite. Since we find a lower recovered SZ signal than *Planck*, we investigate the possibility that there is a significant offset between the optically selected brightest cluster galaxies (BCGs) and the SZ centers, to which ACT is more sensitive due to its finer resolution. Such offsets can arise due to either an intrinsic physical separation between the BCG and the center of the gas concentration or from misidentification of the cluster BCG. We find that the entire discrepancy for both ACT and *Planck* can be explained by assuming that the BCGs are offset from the SZ maxima with a uniform random distribution between 0 and 1.5 Mpc. Such large offsets between gas peaks and BCGs for optically selected cluster samples seem unlikely given that we find the physical separation between BCGs and X-ray peaks for an X-ray-selected subsample of MaxBCG clusters to have a much narrower distribution that peaks within 0.2 Mpc. It is possible that other effects are lowering the ACT and *Planck* signals by the same amount, with offsets between BCGs and SZ peaks explaining the remaining difference between ACT and *Planck* measurements. Several effects that can lower the SZ signal equally for both ACT and *Planck*, but not explain the difference in measured signals, include a larger percentage of false detections in the MaxBCG sample, a lower normalization of the mass–richness relation, radio or infrared galaxy contamination of the SZ flux, and a low intrinsic SZ signal. In the latter two cases, the effects would need to be preferentially more significant in the optically selected MaxBCG sample than in the MCXC X-ray sample.

Key words: cosmic background radiation – galaxies: clusters: general – galaxies: clusters: intracluster medium

Online-only material: color figures

1. INTRODUCTION

Galaxy cluster properties may follow simple scaling laws reflecting their self-similarity. This possibility has given cre-dence to their use as cosmological probes. Cosmological param-eters have been obtained by X-ray, optical, and most recently Sunyaev–Zel'dovich (SZ) surveys of clusters (e.g., Vikhlinin et al. 2009; Mantz et al. 2010; Rozo et al. 2010; Sehgal et al.

2011; Benson et al. 2013). At the same time the scaling laws that feed into these parameter constraints continue to undergo scrutiny.

Recent millimeter-wavelength data have opened a new window whereby these scaling relations can be robustly checked against SZ flux measurements. The SZ cluster signal has been predicted to have a low-scatter correlation with cluster mass (e.g., Motl et al. 2005; Nagai 2006). If true, this would make SZ-detected clusters an excellent tracer of structure growth in the universe (e.g., Wang & Steinhardt 1998; Haiman et al. 2001; Holder et al. 2001; Carlstrom et al. 2002). As steps toward understanding the SZ–mass relation, several studies have shown that the SZ cluster signal correlates well with X-ray signals (e.g., Bonamente et al. 2008, 2012; Andersson et al. 2011), dynamically determined masses (Sifón et al. 2012), and weak-lensing determined masses (e.g., Marrone et al. 2012).

In particular, the *Planck* satellite recently reported a good agreement between the measured and expected SZ–mass relation for a sample of X-ray-selected clusters (Planck Collaboration et al. 2011a, 2011b). However, a similar comparison for optically selected clusters yielded an amplitude of SZ flux lower than expected by about a factor of two, with an even larger discrepancy for lower-mass clusters (Planck Collaboration et al. 2011c). An analysis by Draper et al. (2012) using data from the *WMAP* satellite found a similar result, however with larger uncertainty. Hand et al. (2011), using data from the Atacama Cosmology Telescope (ACT; Swetz et al. 2011) and stacking luminous red galaxies, also suggested a low SZ flux for optically selected halos. Among the possible explanations could be that either the SZ signal is not a robust tracer of galaxy clusters and groups or that optical selection techniques are somehow biased.

Here, we investigate this discrepancy by stacking optically selected clusters in millimeter-wavelength data from ACT that overlaps Stripe 82 of the Sloan Digital Sky Survey (SDSS; York et al. 2000). We also measure the SZ flux for X-ray-selected clusters as a consistency check. Understanding these scaling relations will have important implications for cluster astrophysics as well as for their use in cosmological studies.

This paper is organized as follows. Section 2 discusses the data sets used in this analysis. Section 3 describes the method used to measure cluster SZ flux. Results are presented in Sections 4–6 and discussed in Section 7.

2. DATA SETS

Below we describe the catalogs of optically selected and X-ray-selected clusters and the millimeter-wavelength data used to measure the cluster SZ fluxes. We note that throughout this work M_{500c} refers to the mass within R_{500c}, which is the radius within which the average density equals 500 times the critical density of the universe at the cluster redshift. Similarly, M_{200m} gives the mass within R_{200m}, the radius within which the average density equals 200 times the mean matter density of the universe at the cluster redshift. A fiducial cosmology of $\Omega_m = 0.27$, $\Omega_\Lambda = 0.73$, and $h = 0.71$ is also adopted (Komatsu et al. 2011), where $H(z) = H_0 E(z) = (h \times 100 \text{ km s}^{-1} \text{ Mpc}^{-1}) E(z)$ and $E(z) = [\Omega_m (1 + z)^3 + \Omega_\Lambda]^{1/2}$.

2.1. The MaxBCG Optical Cluster Catalog

The MaxBCG Optical Cluster Catalog consists of 13,823 clusters selected from Data Release 5 (DR5) of the SDSS (Koester et al. 2007a, 2007b). The clusters were selected from a 7500 deg^2 area of sky using the observation that cluster galaxies

tend to be the brightest galaxies at a given redshift, share a similar red color, and are spatially clustered. The catalog consists of clusters that fall in the redshift range of $0.1 < z < 0.3$ and have a richness measure, N_{200m}, within $10 < N_{200m} < 190$. The richness is defined as the number of red-sequence galaxies with $L > 0.4L_*$ (in the i band) within a projected radius of R_{200m}. The catalog provides the brightest cluster galaxy (BCG) position (R.A. and decl.), photometric redshift, richness, BCG luminosity, and total luminosity of each cluster. Applying the cluster detection method to mock catalogs suggests that the catalog should be 90% pure and 85% complete.

Mass estimates of the clusters in the MaxBCG sample were derived by Sheldon et al. (2009) and Mandelbaum et al. (2008a) using weak gravitational lensing. Johnston et al. (2007a) and Rozo et al. (2009) used those mass determinations to construct richness–mass (N_{200m}–M_{500c}) relations. Rozo et al. (2009), in particular, used the masses derived from Mandelbaum et al. (2008a) due in part to the authors' careful treatment of photometric redshift uncertainties (Mandelbaum et al. 2008b). Rozo et al. (2009) also stacked the MaxBCG cluster catalog on X-ray maps from the *ROSAT* All-Sky Survey (Voges et al. 1999) and used the L_X–M relation from Vikhlinin et al. (2009) as a prior to inform their richness–mass relation. Thus the richness–mass relation of Rozo et al. (2009) is expected to be consistent with an L_X–M relation from X-ray clusters. Rozo et al. (2010) applied this richness–mass relation to the MaxBCG sample of optically selected clusters and found a cosmological constraint on σ_8 of $\sigma_8 (\Omega_m/0.25)^{0.41} = 0.832 \pm 0.033$ assuming a flat ΛCDM cosmology. Throughout this work we define the N_{200m}–M_{500c} relation as given by Equations (4), (A20), and (A21) of Rozo et al. (2009).

2.2. The MCXC X-Ray Cluster Catalog

The Meta-Catalog of X-ray detected Clusters of galaxies (MCXC) is presented in Piffaretti et al. (2011). The MCXC cluster catalog is based on publicly available data from a number of different X-ray catalogs including the *ROSAT* All-Sky Survey and comprises 1743 clusters. The catalog provides the position (R.A. and decl.), redshift, X-ray 0.1–2.4 keV band luminosity (L_{500c}), mass (M_{500c}), and radius (R_{500c}) of each system. The redshift distribution of this catalog goes from about 0.05 to 1. The L_X–M relation derived from the MCXC clusters in Piffaretti et al. (2011) is consistent with that of Pratt et al. (2009) and Vikhlinin et al. (2009). The Vikhlinin et al. (2009) L_X–M relation was used to derive a cosmology constraint on σ_8 from a sample of X-ray clusters that is a subsample of the MCXC catalog. From this analysis they found $\sigma_8 (\Omega_m/0.25)^{0.47} = 0.813 \pm 0.013 \text{ (stat)} \pm 0.024 \text{ (sys)}$ (Vikhlinin et al. 2009). Other authors have found similar constraints on σ_8 from *ROSAT* and other X-ray cluster samples (Henry et al. 2009; Mantz et al. 2010).

2.3. Millimeter-wave Data from the Planck Satellite

Given that the optically selected MaxBCG cluster catalog and the X-ray-selected MCXC cluster catalog yield consistent L_X–M relations and constraints on σ_8, one would expect both cluster samples to yield consistent Y_{500c}–M_{500c} relations.[22] However, in a set of papers presented by the Planck collaboration (Planck Collaboration et al. 2011b, 2011c) it was found that the Y_{500c}–M_{500c} relation for X-ray-selected clusters from the MCXC

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[22] Y_{500c} is the SZ flux within R_{500c}, dividing out the frequency dependence of the SZ signal.

sample agrees with expectations, whereas the normalization of the $Y_{500c}-M_{500c}$ relation for the optically selected MaxBCG sample was lower than expectations by about a factor of two. For both cases, expectations were based on X-ray derived cluster profiles from Arnaud et al. (2010). For the latter case, the expectation also folded in the $N_{200m}-M_{500c}$ relation of Rozo et al. (2009). The *Planck* data used in the above analysis consists of six HFI channel millimeter-wave temperature maps as described in Planck HFI Core Team et al. (2011). This data set comprises the first ten months of the survey and covers the full sky.

2.4. Millimeter-wave Data from the Atacama Cosmology Telescope

The ACT is a 6 m telescope operating at an altitude of 5200 meters in the Atacama Desert of Chile. The telescope site allows ACT to observe in both the northern and southern hemispheres. In this work, we use millimeter-wave maps covering two regions of sky: one spanning 510 deg^2 over the celestial equator and one spanning 455 deg^2 in the southern hemisphere. The Equatorial region consists of a 4°.5 wide strip centered at a declination of 0° and running from 20^h20^m through 0^h to 03^h50^m. The Southern region consists of a 7° wide strip centered on $-53°$ and extending from 00^h12^m to 7^h10^m. Both sky regions were observed over the 2008, 2009, and 2010 observing seasons at 148 and 218 GHz. The Equatorial region overlaps the SDSS Stripe 82 and thus overlaps 492 clusters in the MaxBCG catalog. The Equatorial plus Southern regions combined overlap 74 clusters in the MCXC catalog.[23] For a more detailed description of the ACT instrument, observations, and data reduction see Fowler et al. (2007), Swetz et al. (2011), Marriage et al. (2011), Das et al. (2011), Hajian et al. (2011), and Dunner et al. (2013).

3. MEASUREMENTS OF SZ FLUX

3.1. Multi-frequency Matched Filter

We use a multi-frequency matched filter to extract the thermal SZ signal from clusters as described in Haehnelt & Tegmark (1996) and Melin et al. (2006). The filter in Fourier space is given by

$$\Psi(\mathbf{k}) = \sigma_\theta^2 [\mathbf{P}(\mathbf{k})]^{-1} \cdot \tau(\mathbf{k}), \qquad (1)$$

where $\tau(\mathbf{k})$ has the components

$$\tau_\nu(\mathbf{k}) = \tau'(\mathbf{k}) j_\nu B_\nu(\mathbf{k}). \qquad (2)$$

Here j_ν is the frequency dependence of the thermal SZ signal for frequency ν, $\tau'(\mathbf{k})$ is the profile of the cluster in Fourier space, and $B_\nu(\mathbf{k})$ is the profile of the instrument beam in Fourier space. $\mathbf{P}(\mathbf{k})$ is the power spectrum of the noise, both astrophysical and instrumental. The astrophysical noise sources for cluster detection include the primary lensed microwave background, radio galaxies, infrared galaxies, Galactic emission, and the SZ background from unresolved clusters, groups, and the intergalactic medium. Since the power from the cluster thermal SZ signal is subdominant to these astrophysical sources (as evidenced by Lueker et al. 2010; Hall et al. 2010; Fowler et al. 2010; Das et al. 2011), we approximate the power spectrum of the total noise as the power spectrum of the data itself. Here

$$\sigma_\theta^2 = \left[\frac{1}{(2\pi)^2} \int d^2k [\tau(\mathbf{k})]^t \cdot [\mathbf{P}(\mathbf{k})]^{-1} \cdot [\tau(\mathbf{k})] \right]^{-1} \qquad (3)$$

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[23] Only 6 clusters are in common between the 492 cluster MaxBCG sample and the 74 cluster MCXC sample.

is the normalization of the filter that ensures an unbiased estimate of the cluster signal.

3.2. SZ Model Template

We use for the filter's spatial template the empirical universal pressure profile of Arnaud et al. (2010) derived from X-ray observations of the REXCESS cluster sample (Böhringer et al. 2007). The three-dimensional pressure profile is given by

$$P^{3D}(r) \propto \frac{1}{x^\gamma (1 + x^\alpha)^{(\beta-\gamma)/\alpha}} , \qquad (4)$$

where $x = r/r_s$, $r_s = R_{500c}/c_{500}$, $c_{500} = 1.156$, $\alpha = 1.0620$, $\beta = 5.4807$, and $\gamma = 0.3292$. The normalization of this profile is arbitrary for the purposes of the matched filter. We essentially measure this normalization for each cluster when we apply this filter to our maps. The SZ signal is given by the projected gas pressure, so we describe the filter template by integrating the three-dimensional profile above along the line of sight. Thus

$$P^{2D}(\theta) = \int_0^{l_{max}} 2P^{3D}(\sqrt{l^2 + \theta^2 D_A(z)^2}) dl, \qquad (5)$$

where $l_{max} = 5R_{500c}$ and $D_A(z)$ is the angular diameter distance. The filter is truncated at $5R_{500c}/D_A(z) = 5\theta_{500c}$, which contains over 95% of the signal, as was done in Melin et al. (2011) and by the Planck Collaboration et al. (2011b, 2011c).

3.3. Application of Filter

Before filtering the maps, we establish uniform noise properties by creating an effective weight map that has pixel-wise effective weights given by $w_{eff} = ((1/w_1)+(1/w_2))^{-1}$ where w_i is the pixel weight for the ith frequency. Here weight is defined as the number of observations per pixel normalized by the observations per pixel in the deepest part of the map. We multiply the 148 and 218 GHz ACT data maps pixel-wise by the square root of the effective weight map. After we apply the filter to create a filtered map, we divide the filtered map pixel-wise by the square root of the effective weight map.

We apply the matched filter above following the procedure given in Planck Collaboration et al. (2011c). For each cluster in the MaxBCG catalog that falls in the ACT coverage region, we create a unique matched filter using the $N_{200m}-M_{500c}$ relation of Rozo et al. (2009) to determine each cluster's M_{500c} and subsequently R_{500c} from the N_{200m} value given in the catalog. We also derive $D_A(z)$ for each cluster from its photometric redshift. When the filter is applied to the map, the pixel coincident with the location of the cluster center in the filtered map should have a value equal to the normalization of the two-dimensional SZ template given by Equation (5). To simplify extraction of the desired quantity, we normalize $P^{2D}(\theta)$ itself to equal unity when integrated over θ in two dimensions from zero to $5\theta_{500c}$. Thus the pixel value recovered at the cluster center position after applying the filter is $Y_{5\theta_{500c}}^{cyl}$. Here $Y_{5\theta_{500c}}^{cyl} D_A(z)^2 = Y_{5R_{500c}}^{cyl}$, where $Y_{5R_{500c}}^{cyl}$ is the integrated projected SZ signal within a cylinder of radius $5R_{500c}$. We use a geometric factor of $Y_{R_{500c}}^{sph} = (0.986/1.814)Y_{5R_{500c}}^{cyl}$ given in Appendix A of Melin et al. (2011) to convert from $Y_{5R_{500c}}^{cyl}$ to $Y_{R_{500c}}^{sph}$, the integrated SZ flux within a sphere of radius R_{500c}. Throughout this work we plot $\tilde{Y}_{500} \equiv Y_{500} E^{-2/3}(z)(D_A(z)/500 \text{ Mpc})^2$, where $Y_{500} = Y_{\theta_{500c}}^{sph}$, as in Planck Collaboration et al. (2011c).
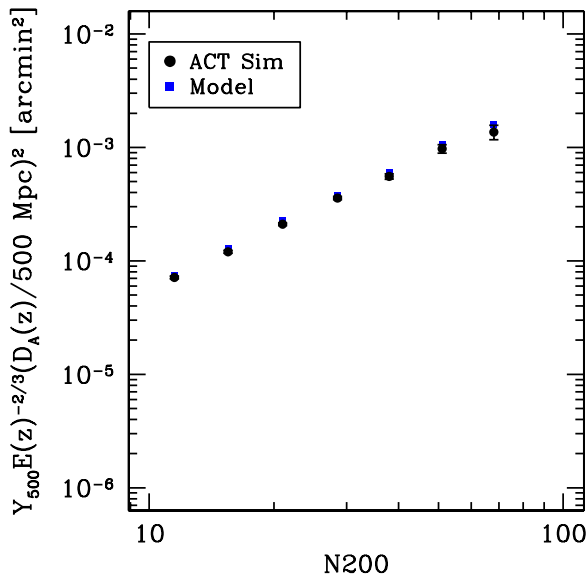
Figure 1. Recovered Y_{500} values from 447 simulated clusters embedded in ACT 148 and 218 GHz maps (black circles). These clusters were simulated to match the properties of the clusters in the MaxBCG catalog (Koester et al. 2007b) that overlap the ACT equatorial region. The simulated clusters were placed at random locations within the ACT maps. The input Y_{500} values are shown as blue squares.

(A color version of this figure is available in the online journal.)

4. SIMULATED ACT SZ SIGNALS

4.1. Embedding Simulated Clusters in ACT Maps

In order to test the analysis pipeline which applies the SZ extraction procedure discussed above, we use simulated clusters embedded within the ACT data maps at random locations. Using the information in the MaxBCG catalog for the 492 clusters that fall within the ACT Equatorial region, we create a unique SZ profile for each cluster using Equation (5) above and the cluster R_{500c} and z given in the catalog. We then add each simulated SZ cluster to the 148 and 218 GHz ACT Equatorial maps, placing it at a random location and scaling the thermal SZ signal to give it the appropriate frequency dependence in each map. The simulated SZ signal is also convolved with the appropriate ACT beam prior to embedding it within each ACT map. We then exclude any simulated clusters from further analysis that happen to be within $5'$ of a point source detected at greater than 5σ in either the 148 or 218 GHz maps. We also exclude any clusters that are within $10'$ from the edge of the map. As an additional cut, we exclude all clusters that are in noisy parts of the map where the local value of the effective weight map is less than 15% of the maximum value. These cuts leave 447 clusters. Similarly, below, when we extract the SZ signal from the real cluster positions using the MaxBCG and MCXC catalogs, we apply the same cuts discussed above.

In Figure 1, we show the results of this SZ extraction procedure. Here the simulated clusters are binned by the optical richness given in the MaxBCG catalog. The blue squares show the input model, and the black circles show the recovered signal of the simulated clusters embedded at random locations within the ACT maps. The error bars are given by $\sigma/\sqrt{N}$ where σ^2 is the variance of $\tilde{Y}_{500}$ in each richness bin, and N is the number of clusters in each richness bin. The variance dominates the uncertainty in each bin, as demonstrated in Figure 4 of Planck

Collaboration et al. (2011c).[24] There is good agreement between input and recovered signals.

4.2. Effect of Cluster Miscentering

The exercise above shows that we should expect excellent recovery of the SZ flux for clusters given in the MaxBCG catalog provided the $N_{200m}-M_{500c}$ and $M_{500c}-Y_{500c}$ relations are correct, and the cluster properties (position, redshift, N_{200m}) listed in the catalog are accurate. However, one source of uncertainty identified in Johnston et al. (2007b, Section 4.3) is the positional accuracy of the cluster center. In the MaxBCG catalog, the cluster position is given by the location of the BCG. Johnston et al. (2007b) suggested two reasons why the BCG found by the MaxBCG cluster identification algorithm may be offset from the true cluster center. One is that the true BCG may be intrinsically offset from the dark matter center or center of the gas concentration presumably due to unrelaxed behavior (e.g., mergers). Another reason may be that the BCG could be misidentified by the cluster finder. Johnston et al. (2007b) explore this latter effect with mock optical cluster catalogs and find that a richness-dependent fraction of the clusters have accurately identified BCGs while the rest are miscentered due to BCG misidentification following a Rayleigh distribution with a scale parameter σ equal to $\sigma_{\mathrm{misc}} = 0.42\ h^{-1}$ Mpc. The distribution of the intrinsic BCG offset from the gas center is unknown.[25]

We explore the potential offset between BCG and gas center by studying the subset of clusters in common to both the MaxBCG and MCXC catalogs. From the full catalogs (which have 13,823 and 1743 clusters, respectively), we identify 208 clusters that are in common in both catalogs. Here we define a cluster as matched in both catalogs when the identified cluster redshifts are within $\Delta z < 0.015$ and the projected physical separation of the identified cluster centers is less than 1.5 Mpc. Using these 208 clusters, we plot the fraction of clusters as a function of separation between BCG and X-ray peak in Figure 2. The solid black line shows the offset distribution for the 208 clusters. Half of the clusters have offsets less than 0.1 Mpc, while the other half have a roughly flat offset distribution between 0.1 and 1.5 Mpc.[26]

The 208 clusters were then divided into rich clusters (with $N_{200m} \geqslant 35$) and poor clusters (with $N_{200m} < 35$). A richness cut of $N_{200m} = 35$ divides the 208 clusters into roughly even subsamples of $\sim$100 clusters each. The dashed red line in Figure 2 shows the offset distribution for the rich clusters, and the dotted blue line shows the distribution for the poor clusters. Similar distributions are found for both subsamples, with the poor clusters showing slightly more offset.

Using the simulations described above, we explore the effect of cluster miscentering on SZ flux recovery. We use the same simulated clusters embedded at random positions in the 148 and 218 GHz ACT maps as before. However, when recovering

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[24] Since no scatter in the $N_{200m}-M_{500c}$ relation has been included here, these error bars reflect only the SZ flux recovery error.

[25] Note that this offset does not arise from pointing uncertainties in optical, X-ray, or millimeter-wave instruments. It is due to either BCG misidentification or cluster astrophysics.

[26] We note that the choice of 1.5 Mpc is somewhat arbitrary. When we allow matches within 1 Mpc, we find 189 clusters. We also find that if we allow matches within 3 Mpc, Figure 2 plateaus instead of dropping to zero at large separation. The plateau is due to poor clusters, as the distribution of rich ones does tend to zero. The size of a cluster (R_{200m}) is usually less than 2 Mpc, so this suggests that some of these poor cluster matches may be spurious.
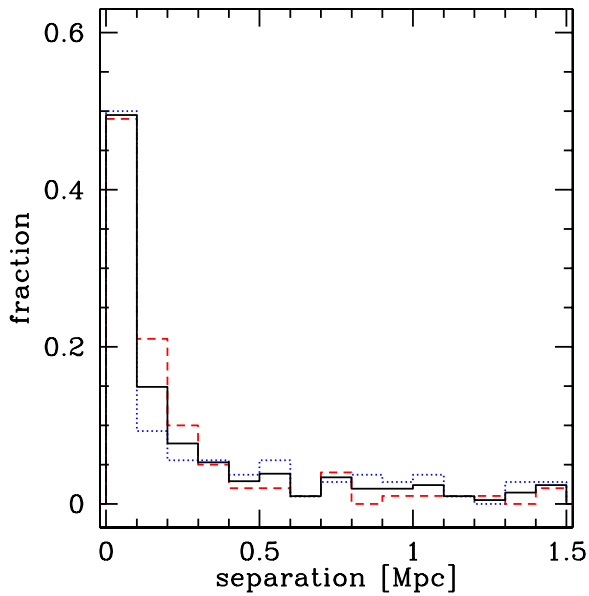
Figure 2. The black solid line shows the distribution of offsets between BCG and X-ray gas peak from 208 clusters found in both the MaxBCG and MCXC cluster catalogs. The red dashed line shows the distribution for the subsample of rich clusters ($N_{200m} \geqslant 35$). The blue dotted line shows the same for the subsample of poor clusters ($N_{200m} < 35$). There are 100 and 108 clusters in each subsample, respectively.
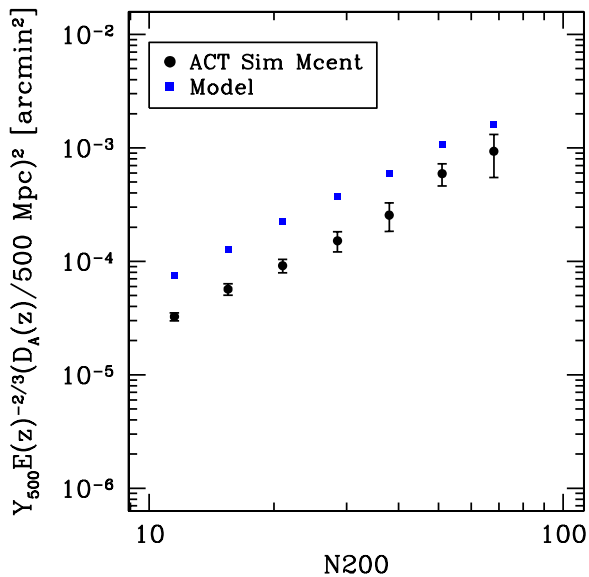
(A color version of this figure is available in the online journal.)

Figure 3. Black circles show recovered Y_{500} values from 446 embedded simulated clusters where the assumed cluster centers used for Y_{500} recovery are offset from the input centers with a random distribution given by the black solid line in Figure 2. "Mcent" stands for miscentered. The input values are shown as blue squares.

(A color version of this figure is available in the online journal.)

the SZ fluxes, we use positions for the clusters that differ from the true cluster gas centers with a random distribution given by the black solid line in Figure 2. We also allow for 10% of the clusters to be false detections to match the purity of the MaxBCG catalog (Koester et al. 2007a). Figure 3 shows the result of the SZ recovery process. The black circles show the recovered Y_{500c} values for simulated clusters embedded in ACT data. To give a sense of the relation between Mpc and arcminutes, for the redshift range of the MaxBCG cluster sample ($z \in (0.1, 0.3)$),
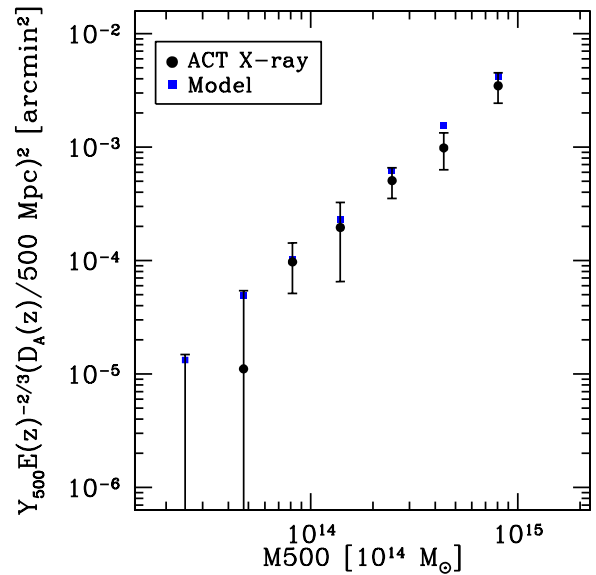

Figure 4. Measured Y_{500} values for 52 MCXC X-ray-selected clusters (Piffaretti et al. 2011) that fall within the ACT equatorial and southern survey regions (black circles). Also shown are expected Y_{500} values based on measured cluster X-ray properties (blue squares). A cluster profile model from Arnaud et al. (2010) was assumed for determining both measured and expected Y_{500} values.

(A color version of this figure is available in the online journal.)

0.5 Mpc corresponds to about $2'$ to $4'.5$. Given the ACT beam size ($1'.4$ at 148 GHz and $1'$ at 218 GHz), we expect a decrease in the recovered Y_{500c} signal due to the amount of miscentering shown in Figure 3. How much the recovered SZ signal decreases depends on the noise properties of the ACT data, which differ significantly from pure white noise. We discuss this further in Section 7.

5. MEASURED ACT SZ SIGNALS

5.1. Stacking X-Ray Selected Clusters

Within the 510 deg² ACT Equatorial region and the 455 deg² ACT Southern region are located 74 clusters found in the MCXC catalog of X-ray-selected clusters (Piffaretti et al. 2011). After making the cuts discussed in Section 4.1 to exclude clusters near bright point sources, near map edges, or in very noisy parts of the map, 52 MCXC clusters remain. Using the R_{500c}, M_{500c}, and redshift of each cluster and the projected SZ profile given in Equation (5), we calculate the expected mean Y_{500c} values in each M_{500c} bin shown as the blue squares in Figure 4. The black circles in Figure 4 show the mean of the recovered Y_{500c} values using the multi-frequency matched filter given in Equation (1) and the projected SZ profile created uniquely for each cluster. The error bars are the error on the mean given by $\sigma/\sqrt{N}$, where σ^2 is the variance and N is the number of clusters in each bin. Figure 4 shows overall agreement between expected and recovered SZ signals for the X-ray-selected clusters. The reduced chi-squared is 0.76 using 7 degrees of freedom. This is consistent with the agreement found by Planck Collaboration et al. (2011b) for a larger sample of X-ray-selected clusters.

5.2. Stacking Optically Selected Clusters

In the ACT Equatorial region there are 492 MaxBCG clusters. This reduces to 474 clusters once the above-mentioned cuts are
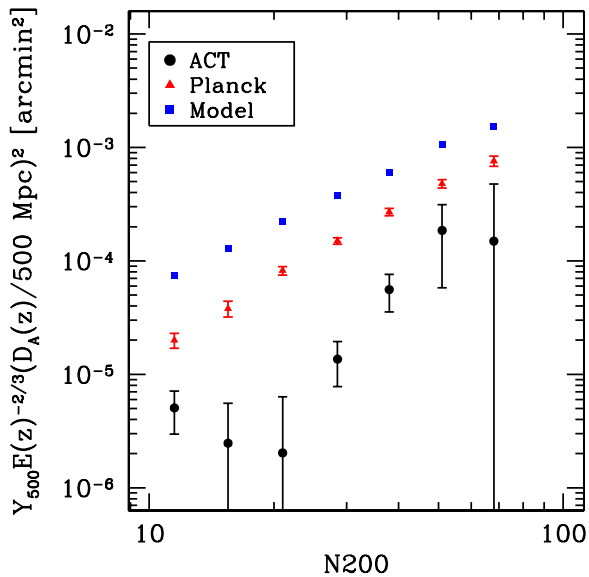
Figure 5. Measured Y_{500} values for 474 MaxBCG optically selected clusters that fall within the ACT equatorial survey region (black circles). Expected Y_{500} values are shown as blue squares. Both measured and expected values assume the $N_{200m}–M_{500c}$ relation from Rozo et al. (2009) and the Arnaud et al. (2010) cluster profile. Red triangles are the measured values from the *Planck* satellite for a sample of 13,104 MaxBCG clusters (Planck Collaboration et al. 2011c).

(A color version of this figure is available in the online journal.)

Figure 6. The black open circles show the measured Y_{500} values from simulations with an offset between identified cluster center and gas concentration center given by a uniform random distribution between 0 and 1.5 Mpc. The solid black circles, blue squares, and red triangles are the same as in Figure 5. The values from simulations have been shifted in the *x*-axis for clarity.

(A color version of this figure is available in the online journal.)

made.[27] Using the $N_{200m}–M_{500c}$ relation of Rozo et al. (2009) and the $M_{500c}–Y_{500c}$ relation of Arnaud et al. (2010), we find the expected and recovered Y_{500c} values of the MaxBCG clusters using the method described in Section 3. Figure 5 shows the expected values as blue squares and the recovered values as black circles. The recovered signal is significantly lower than the expected signal as well as the signal recovered by *Planck* from 13,104 MaxBCG clusters (red triangles; Planck Collaboration et al. 2011c).

From Figure 3, we see that some amount of offset between BCG and gas peak can result in a lower measured signal than expected. However, this figure also shows that the offset distribution given by Figure 2 does not result in a low enough measured signal to explain the measurement shown in Figure 5. To investigate the amount of offset necessary to match the ACT measured Y_{500} values shown in Figure 5, we redo the analysis shown in Figure 3. However, this time we use an offset distribution that is uniformly random between 0 and 1.5 Mpc. We also again allow for 10% false detections. The results are shown as the open black circles in Figure 6. The solid black circles, blue squares, and red triangles in Figure 6 are the same as in Figure 5. We see that this amount of offset between BCG and gas peak roughly matches the measured values. An extensive scan of offset distributions is beyond the scope of this work, but it may be that a more complex or refined distribution could give a better fit to the measurements.

5.3. Stacking Optical Clusters Using BCG Dominant Subsample and New Richness Measure

Using a subsample of clusters with "dominant BCGs" we examine whether the expected SZ signal is closer to the measured
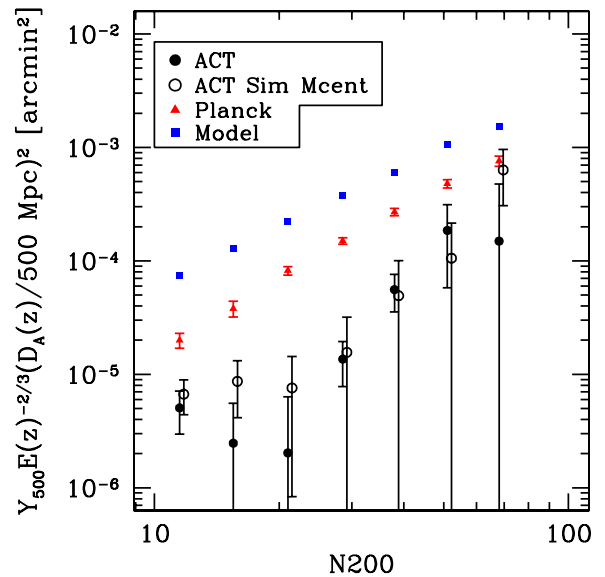
values. Such a subsample may more closely correspond to an X-ray-selected subsample, and with such a subsample Planck Collaboration et al. (2011c) found better agreement between model and measurement. We follow the definition of "BCG dominant" used by Planck Collaboration et al. (2011c) which is defined relative to the quantity $L_{BCG}/(L_{tot} − L_{BCG})$. Here L_{tot} and L_{BCG} are the *R*-band luminosities of the cluster and cluster BCG, respectively. For a "BCG dominant" cluster, this ratio is larger than the average ratio for a given richness bin. From the sample of 474 MaxBCG clusters used above, 126 are "BCG dominant." Figure 7 (left) shows the recovered Y_{500} values versus the model expectation for this subsample. The *Planck* measurements of the sample of 13,104 MaxBCG clusters (not a subsample) is included for reference.

Recently a new measure of cluster richness was developed with less scatter than the measure presented in Koester et al. (2007b; Rykoff et al. 2012). We test whether using this new richness measure will yield differing results with regard to measured versus expected SZ signal. A catalog with a new richness measure assigned to each MaxBCG cluster is available online.[28] While for the previous richness measure the cluster richness–mass relation was calibrated using weak lensing, this new relation was calibrated with an abundance matching technique (Rykoff et al. 2012). Calibration of this new measure via weak lensing is still in progress. We use the richness–mass relation given in Equation (B6) of Rykoff et al. (2012) to determine M_{500c} and subsequently R_{500c}. Figure 7 (right) shows the measured and expected Y_{500} values using this new richness measure for the 474 MaxBCG clusters. The red triangles show the *Planck* measurements using the *old* richness measure for reference. Note that *Planck* would have different results if they use the new richness measure.

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[27] Note that the ACT region is centered on Stripe 82, and extends beyond it. So fewer real clusters are cut than simulated clusters because the real clusters are only in Stripe 82, and thus are located toward the central part of the ACT map which has lower noise.
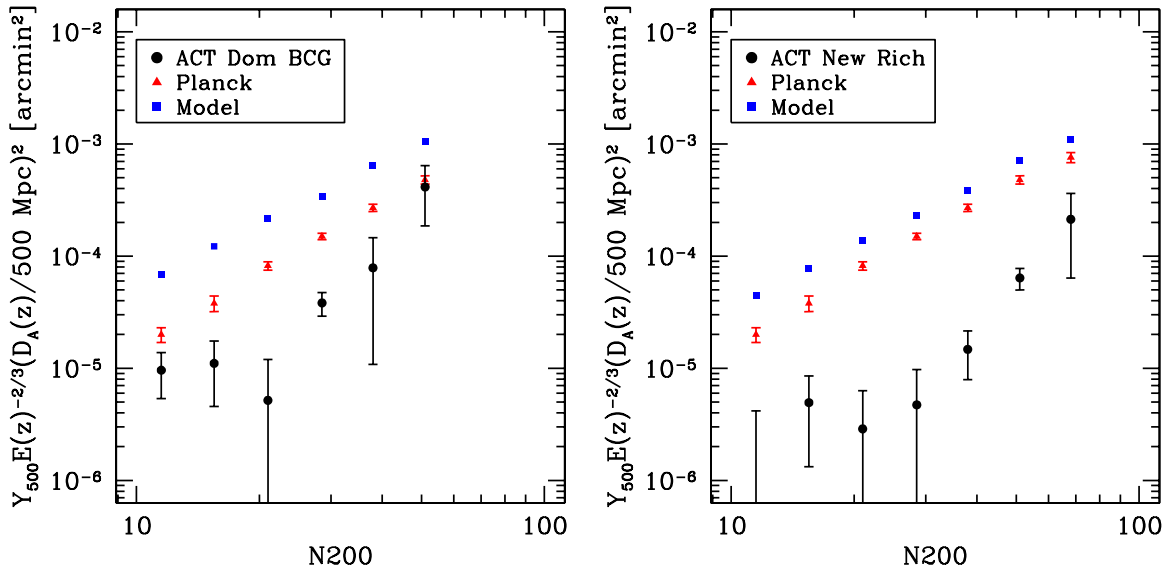
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[28] http://kipac.stanford.edu/maxbcg

Figure 7. Left: shown are the recovered Y_{500} values as in Figure 5, except using a subsample of 126 "BCG dominant" clusters from the 474 MaxBCG clusters used above. Both the measured values (black circles) and expected values (blue squares) change using the new subsample as compared to Figure 5. The red triangles are the measured values from *Planck* for the 13,104 MaxBCG sample. Right: shown are the same points as in Figure 5, however, using the new richness measure for the MaxBCG cluster sample given in Rykoff et al. (2012). Note that both the measured values (black circles) and expected values (blue squares) change using the new richness measure as compared to Figure 5. The red triangles are the *Planck* measured values using the *old* richness measure for reference.

(A color version of this figure is available in the online journal.)

5.4. Contamination from Infrared and Radio Galaxies

To investigate whether infrared galaxies may be reducing the SZ decrement at the MaxBCG cluster positions, we recover the Y_{500} values from the 474 MaxBCG clusters studied above using the single band 218 GHz ACT map alone. We compare that to Y_{500} values extracted at 474 random positions within the 218 GHz map.[29] A positive correlation between MaxBCG clusters and infrared galaxies would result in negative Y_{500} values compared to the random sample. We find that both the MaxBCG cluster sample and the random sample have an SZ flux consistent with zero in the 218 GHz map, and we do not detect any significant excess of infrared signal correlated with the MaxBCG sample. This is shown in Figure 8.

We also cross-correlate the MaxBCG cluster catalog with the Very Large Array FIRST catalog of radio sources to investigate how much radio galaxies may be reducing the measured SZ decrement.[30] The FIRST survey uses the NRAO Very Large Array and covers over 10,000 deg^2 of sky to a sensitivity of about 1 mJy at 1.4 GHz. This survey also overlaps with the SDSS. We cross-correlate to find the fraction of 474 MaxBCG clusters that have a radio source above a given flux threshold, within 5′ of the identified cluster position. Assuming a typical spectral index for radio sources of -0.7 (e.g., Condon 1984; Lin et al. 2009), we choose a flux threshold cut of 50 mJy at 1.4 GHz to yield sources above 2 mJy at 150 GHz. Such a source convolved with the ACT beam would have a temperature increment of about 30 μK at 148 GHz. Since the typical SZ signal from a cluster is about 100 μK to within a factor of a few when smoothed with the ACT beam, a 2 mJy radio source would start to significantly reduce the SZ decrement. We find that about 10% of the 474 MaxBCG clusters investigated above have such a radio source within 5′ of its identified center. This
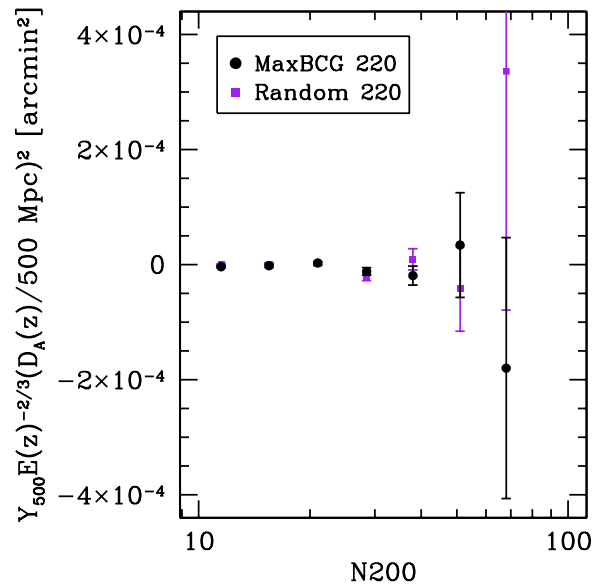

Figure 8. Measured Y_{500} values from 474 MaxBCG clusters using the 218 GHz ACT equatorial map alone (black circles). For comparison is shown the same Y_{500} recovery procedure performed at 474 random locations in the 218 GHz map (purple squares).

(A color version of this figure is available in the online journal.)

small correlation is not enough to explain the large discrepancy between measured and expected SZ signals shown in Figure 5, although it may be a contributing factor.

6. SIMULATED *PLANCK* SZ SIGNALS

To investigate what *Planck*'s measured SZ signal would be if there existed the amount of offset modeled in Figure 6, we again use simulations. We embed 492 simulated MaxBCG clusters in two sets of simulated cosmic microwave background (CMB) maps at 148 and 218 GHz that have *Planck*-like instrument

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[29] For these measurements of Y_{500}, we do not divide out the amplitude of the frequency dependence of the SZ signal, which is close to zero at the null frequency. So these are really measurements of $-\Delta T / T_{cmb}$.
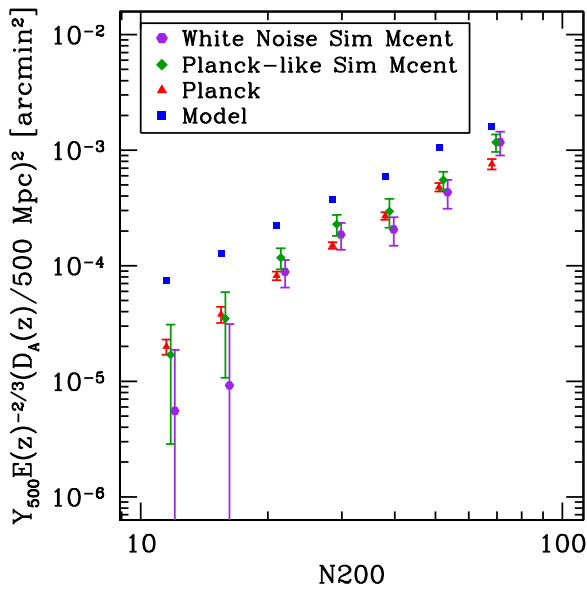[30] http://sundog.stsci.edu/first/catalogs/readme_12feb16.html

Figure 9. Recovered Y_{500} values from 984 simulated clusters embedded in simulated CMB maps at 148 and 218 GHz with *Planck*-like noise (green diamonds) or the corresponding level of white-noise (purple hexagons). The maps were convolved with ACT beams as well as a 5′ Gaussian beam, so that they correspond to *Planck* resolution. The recovered positions are offset from the true positions with a uniform random distribution between 0 and 1.5 Mpc. The input Y_{500} values are shown as blue squares, and red triangles are the measured values from the *Planck* satellite for a sample of 13,104 MaxBCG clusters (Planck Collaboration et al. 2011c).

(A color version of this figure is available in the online journal.)

Figure 10. Shown is the filter response for differing beams and noise models. The blue solid curve is a 5′ Gaussian profile, which roughly approximates the filter response for a cluster profile convolved with the *Planck* beam since *Planck*'s noise is close to white (see Equation (1)). The red dashed curve shows a 1′.4 Gaussian profile, which has been Fourier transformed and has had all Fourier modes set to zero for $l < 2000$, roughly analogous to the effect of the ACT noise. The profile was then Fourier transformed back. The green dotted curve shows the same as the red curve except for a 5′ Gaussian profile.

(A color version of this figure is available in the online journal.)

noise added to them.[31] We model the *Planck* noise using the noise power spectra at 143 and 217 GHz shown in Figure 35 of Planck HFI Core Team et al. (2011). We allow for an offset between the cluster SZ peaks and the identified cluster centers that has a uniform random distribution between 0 and 1.5 Mpc, analogous to Figure 6. We also convolve both maps with ACT beams and a 5′ Gaussian beam to approximate *Planck* resolution. The results of the extracted Y_{500} values are shown as green diamonds in Figure 9. We also show the case where the simulated clusters are embedded in simulated CMB plus white noise maps, using white noise levels that are similar to *Planck* noise levels. These results are shown as purple hexagons in Figure 9.

Planck-like noise is nearly white at these frequencies, so there is not much difference between the *Planck*-like case and the white noise case in Figure 9. The results of the *Planck*-like case in Figure 9 are also very different from those of the simulated ACT case shown by the open black circles in Figure 6. The two differences between the simulations are the beam and the noise. In the ACT data there is $1/f$ noise, atmospheric noise, and noise from the primary microwave background.[32] All of this results in a redder noise spectrum than *Planck*'s. The presence of red noise in the maps causes the matched filter to suppress more power on large scales than would be the case for white noise. This causes the filter to return a lower signal than actually exists if the signal is extracted from a position that is offset from the cluster center. If the signal is extracted at the cluster center, however, the matched filter will return the correct signal
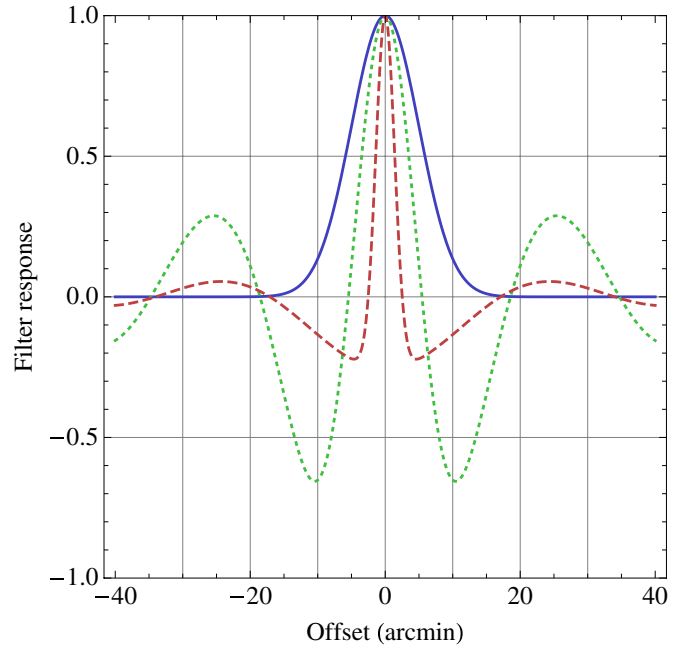
regardless of whether the noise is white or somewhat red (as can be inferred from Figure 1).[33]

We demonstrate this effect in Figure 10, where the blue solid curve is a 5′ Gaussian profile. This is roughly the filter response for a cluster profile convolved with the *Planck* beam as *Planck*'s noise is close to white (see Equation (1)). The red dashed curve shows a 1′.4 Gaussian profile, which had been Fourier transformed and had all Fourier modes set to zero for $l < 2000$, roughly analogous to the effect of the ACT noise. The profile was then Fourier transformed back. One can see how a low flux recovery results from an offset from the center. The green dotted curve shows the same as the red curve except for a 5′ Gaussian profile. The low signal from a miscentered position is still apparent. This is why merely smoothing the ACT maps to match the *Planck* beam would not allow us to reproduce the *Planck* measurement.

It is interesting to see from Figure 9 that the amount of offset modeled in Figure 6 to match the ACT data, can also explain the discrepancy found by *Planck*. If this amount of miscentering actually exists, then it would be the sole explanation of the discrepancies. However, this amount of offset between BCGs and SZ peaks is much larger than the distribution shown in Figure 2. So it may be that the SZ signal is intrinsically low by some amount or that the optical weak-lensing mass calibration is biased high. It is also possible that the fraction of false detections in the optically selected sample is larger than 10%, or some

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[31] We doubled the number of maps and thus the cluster sample to shrink the error bars. This results in 984 embedded clusters in total.
[32] The instrumental noise in the ACT 218 GHz map is not low enough to remove all the primary microwave background signal from the 148 GHz map.

[33] Note that estimating the noise power spectrum as CMB plus white noise when filtering the ACT data results in very large error bars since the matched filter is no longer optimal. High-pass filtering the ACT maps enough to minimize the effect of the red noise causes ringing of the true cluster signal, which also results in a low measurement if the signal is extracted from a position offset from the center.

non-zero amount of radio or infrared galaxy contamination is filling in the SZ decrements (as discussed in Section 5.4). If so, these effects would serve to lower the measured SZ signal by the same amount for both ACT and *Planck*. In that case, a smaller amount of miscentering would be required to match the remaining discrepancy found by these instruments. It is thus likely that a combination of effects may be at work resulting in the measurements found by ACT and *Planck*.

In this vein, we have checked with simulations whether an intrinsic SZ signal as low as measured by *Planck*, combined with the distribution of BCG/SZ peak separations shown in Figure 2, can roughly match the SZ signal measured by ACT. We did this by lowering the SZ signal of the simulated clusters until they matched the *Planck* measurements, and then recovered this signal using wrong gas peak positions given by the distribution in Figure 2, as done in Section 4.2. We find that the recovered simulated SZ signal is at least a factor of two higher than measured by ACT.

7. DISCUSSION

From the analysis of simulated clusters embedded in ACT data presented in Figure 1, we expect robust recovery of the SZ flux from optically selected MaxBCG clusters. This is assuming that the position of the center of the gas concentration is known. From the analysis of 52 MCXC clusters that fall within the ACT Equatorial and Southern survey regions shown in Figure 4, we find agreement between the expected and measured $Y_{500c}-M_{500c}$ relation. Such agreement is consistent with that found in Planck Collaboration et al. (2011b). Both of these figures taken together give confidence in the SZ flux recovery pipeline and the analysis of the ACT maps.

In Figure 5, we find that the recovered ACT SZ flux from 474 optically selected MaxBCG clusters is lower than both the model expectations and the measured *Planck* values given in Planck Collaboration et al. (2011c). Since we expect from Figure 2 some offset between the positions of the BCGs and the centers of the gas concentrations, we explore this possibility in more detail. Given the difference in resolution and noise properties between ACT and *Planck*, such an offset would result in different measured SZ signals between the two instruments (see Figure 10).

Figure 2 gives the offset distribution for clusters in common between the full MaxBCG and MCXC cluster catalogs. Using simulations, Figure 3 shows that this amount of offset is not enough to explain the measured SZ signal shown in Figure 5. Modeling the offset using the Johnston et al. (2007b) distribution, which only includes misidentification of the BCG by the optical cluster finding algorithm, produces a slightly smaller amount of discrepancy between expected and measured SZ signals than what is shown in Figure 3. To match the measured SZ signal found by ACT requires an offset distribution that is significantly larger. A distribution that is uniformly random between 0 and 1.5 Mpc gives a better fit to the ACT data, as shown in Figure 6. This amount of offset also can explain the *Planck* measured discrepancy as shown in Figure 9.

It is possible that the subsample of clusters in common between MCXC and MaxBCG catalogs is not representative of the full MaxBCG cluster sample. This subsample is in fact special in that these clusters are found using both optical and X-ray selection techniques. When *Planck* measured the SZ signal for such a subsample, no significant discrepancy between the measured and expected SZ signals was found

(Planck Collaboration et al. 2011c). This is in contrast to their results for the full optically selected sample. This suggests that the subsample of clusters in common to both catalogs is not representative of the full optically selected sample, at least in some regard.

We also assume throughout this work that the Arnaud et al. (2010) profile accurately describes the gas profiles of clusters in the MCXC and MaxBCG catalogs. The good agreement we find in Figure 4 and also shown in Planck Collaboration et al. (2011b) suggests that this profile works well for clusters in the MCXC sample. However, this profile may not hold for the optically selected MaxBCG sample. In particular, to measure a lower SZ signal with ACT than found by *Planck* would require that the gas distribution is wider (more spread out) than that given by the Arnaud et al. (2010) profile. In such a case, given the finer resolution of ACT, it is conceivable that ACT could measure a lower signal than *Planck* when the SZ signal is extracted using an incorrect narrower profile for both data sets. However, since we expect clusters to be in hydrostatic equilibrium to first order, the consequence of a wider gas distribution is a lower total mass within M_{500c} if the gas mass fraction is fixed. This would imply lower normalizations of the L_X-M_{500c} and $N_{200m}-M_{500c}$ relations. That in turn would result in a lower σ_8 value obtained from the MaxBCG sample than reported in Rozo et al. (2010), the latter of which is currently in agreement with σ_8 values obtained from X-ray-selected cluster samples (Henry et al. 2009; Vikhlinin et al. 2009; Mantz et al. 2010).

Some other recent papers have added more data and analysis shedding light on the measured discrepancy in SZ signal found by *Planck*. Biesiadzinski et al. (2012) argue that there is not really a discrepancy if one considers the 2σ uncertainty on the $N_{200m}-M_{500c}$ relation in Rozo et al. (2009). However, if Rozo et al. (2010) used an $N_{200m}-M_{500c}$ relation that had a normalization that is 2σ away from the true value, then again it is surprising that they find a constraint on σ_8 with an optical cluster sample that is in such good agreement with the σ_8 constraints found using X-ray-selected samples. Such a shift in the $N_{200m}-M_{500c}$ normalization would also put the derived cosmological constraints in some tension with previous results (Rozo et al. 2012a). Angulo et al. (2012), using Millennium-XXL simulations, argue along the same lines as Biesiadzinski et al. (2012) suggesting that the normalization of the richness–mass relation is incorrect (see last paragraph of Section 4 in that work).

Bauer et al. (2012) tested the $N_{200m}-M_{500c}$ relation by verifying the weak-lensing cluster mass estimates upon which this relation was calibrated. They did this by using gravitational lensing magnification of type I quasars in the backgrounds of these clusters. Their results support the mass normalization of the Rozo et al. (2009) $N_{200m}-M_{500c}$ relation. Recent work by Rozo et al. (2012d) suggests that part of the *Planck* measured discrepancy may be due to different X-ray instrument calibrations between *Chandra* and *XMM-Newton* and/or differing analysis procedures when dealing with data from each instrument (see Figure 3 of that work). Rozo et al. (2012d) argue that X-ray scaling relations derived from *Chandra* observations may be systematically high compared to those from *XMM-Newton* observations. Since the $N_{200m}-M_{500c}$ relation was calibrated using an L_X-M_{500c} prior derived from *Chandra* observations (Vikhlinin et al. 2009), and the Arnaud et al. (2010) profile was derived using *XMM-Newton* observations, this may be responsible for part of the discrepancy. Rozo et al. (2012a, 2012b, 2012c) include further discussion in this direction and claim that three

ingredients may be responsible for the tension between measured and expected SZ signals measured by *Planck* for the MaxBCG sample. These ingredients are a mass–richness relation that is normalized high by 10%, *Chandra* derived X-ray masses from Vikhlinin et al. (2009) that are biased low by 20% due to non-thermal pressure, and *XMM-Newton* X-ray derived masses that are ∼10% to 20% lower than those derived from *Chandra*. More data will clarify this issue.

Given the above discussion, the interpretation of both the ACT results presented here and the *Planck* results stacking MaxBCG clusters remains unclear. Since each possibility mentioned above, on its own, would be unlikely to explain the full discrepancy between expectation and measurement of both data sets, it is likely that multiple effects are responsible for the low SZ signals. It is possible that a combination of factors is lowering the SZ signal equally for both ACT and *Planck*, and that offsets between gas peaks and BCGs result in the remaining discrepancy between ACT and *Planck* measurements. We have shown that even in this scenario however, a larger offset distribution than is seen in X-ray-selected cluster samples is still required to explain the discrepancy between ACT and *Planck* measured signals.

Further investigation is required using multi-wavelength techniques to fully contextualize the SZ signal measured by both *Planck* and ACT for this sample of optically selected clusters. High-resolution SZ instruments such as CARMA and MUSTANG2 as well as more sensitive data sets such as those forthcoming from ACTPol and SPTPol will provide more clarity. This will yield a better understanding of the astrophysics of galaxy clusters. Such knowledge will also be vital to assess the use of galaxy clusters as probes of structure growth and the cosmological parameters of the universe.

REFERENCES

Andersson, K., Benson, B. A., Ade, P. A. R., et al. 2011, ApJ, 738, 48
Angulo, R. E., Springel, V., White, S. D. M., et al. 2012, MNRAS, 426, 2046
Arnaud, M., Pratt, G. W., Piffaretti, R., et al. 2010, A&A, 517, A92
Bauer, A. H., Baltay, C., Ellman, N., et al. 2012, ApJ, 749, 56
Benson, B. A., de Haan, T., Dudley, J. P., et al. 2013, ApJ, 763, 147
Biesiadzinski, T., McMahon, J. J., Miller, C. J., Nord, B., & Shaw, L. 2012, ApJ, 757, 1
Böhringer, H., Schuecker, P., Pratt, G. W., et al. 2007, A&A, 469, 363
Bonamente, M., Hasler, N., Bulbul, E., et al. 2012, NJPh, 14, 025010
Bonamente, M., Joy, M., LaRoque, S. J., et al. 2008, ApJ, 675, 106
Carlstrom, J. E., Holder, G. P., & Reese, E. D. 2002, ARA&A, 40, 643
Condon, J. J. 1984, ApJ, 287, 461
Das, S., Marriage, T. A., Ade, P. A. R., et al. 2011, ApJ, 729, 62
Draper, P., Dodelson, S., Hao, J., & Rozo, E. 2012, PhRvD, 85, 023005
Dunner, R., Hasselfield, M., Marriage, T. A., et al. 2013, ApJ, 762, 10
Fowler, J. W., Acquaviva, V., Ade, P. A. R., et al. 2010, ApJ, 722, 1148
Fowler, J. W., Niemack, M. D., Dicker, S. R., et al. 2007, ApOpt, 46, 3444
Haehnelt, M. G., & Tegmark, M. 1996, MNRAS, 279, 545
Haiman, Z., Mohr, J. J., & Holder, G. P. 2001, ApJ, 553, 545
Hajian, A., Acquaviva, V., Ade, P. A. R., et al. 2011, ApJ, 740, 86
Hall, N. R., Keisler, R., Knox, L., et al. 2010, ApJ, 718, 632
Hand, N., Appel, J. W., Battaglia, N., et al. 2011, ApJ, 736, 39
Henry, J. P., Evrard, A. E., Hoekstra, H., Babul, A., & Mahdavi, A. 2009, ApJ, 691, 1307
Holder, G., Haiman, Z., & Mohr, J. J. 2001, ApJL, 560, L111
Johnston, D. E., Sheldon, E. S., Tasitsiomi, A., et al. 2007a, ApJ, 656, 27
Johnston, D. E., Sheldon, E. S., Wechsler, R. H., et al. 2007b, arXiv:0709.1159
Koester, B. P., McKay, T. A., Annis, J., et al. 2007a, ApJ, 660, 221
Koester, B. P., McKay, T. A., Annis, J., et al. 2007b, ApJ, 660, 239
Komatsu, E., Smith, K. M., Dunkley, J., et al. 2011, ApJS, 192, 18
Lin, Y., Partridge, B., Pober, J. C., et al. 2009, ApJ, 694, 992
Lueker, M., Reichardt, C. L., Schaffer, K. K., et al. 2010, ApJ, 719, 1045
Mandelbaum, R., Seljak, U., & Hirata, C. M. 2008a, JCAP, 08, 006
Mandelbaum, R., Seljak, U., Hirata, C. M., et al. 2008b, MNRAS, 386, 781
Mantz, A., Allen, S. W., Rapetti, D., & Ebeling, H. 2010, MNRAS, 406, 1759
Marriage, T. A., Acquaviva, V., Ade, P. A. R., et al. 2011, ApJ, 737, 61
Marrone, D. P., Smith, G. P., Okabe, N., et al. 2012, ApJ, 754, 119
Melin, J.-B., Bartlett, J. G., & Delabrouille, J. 2006, A&A, 459, 341
Melin, J.-B., Bartlett, J. G., Delabrouille, J., et al. 2011, A&A, 525, A139
Motl, P. M., Hallman, E. J., Burns, J. O., & Norman, M. L. 2005, ApJL, 623, L63
Nagai, D. 2006, ApJ, 650, 538
Piffaretti, R., Arnaud, M., Pratt, G. W., Pointecouteau, E., & Melin, J.-B. 2011, A&A, 534, A109
Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2011a, A&A, 536, A11
Planck Collaboration, Aghanim, N., Arnaud, M., et al. 2011b, A&A, 536, A10
Planck Collaboration, Aghanim, N., Arnaud, M., et al. 2011c, A&A, 536, A12
Planck HFI Core Team, Ade, P. A. R., Aghanim, N., et al. 2011, A&A, 536, A6
Pratt, G. W., Croston, J. H., Arnaud, M., & Böhringer, H. 2009, A&A, 498, 361
Rozo, E., Bartlett, J. G., Evrard, A. E., & Rykoff, E. S. 2012a, arXiv:1204.6305
Rozo, E., Evrard, A. E., Rykoff, E. S., & Bartlett, J. G. 2012b, arXiv:1204.6292
Rozo, E., Rykoff, E. S., Bartlett, J. G., & Evrard, A. E. 2012c, arXiv:1204.6301
Rozo, E., Rykoff, E. S., Evrard, A., et al. 2009, ApJ, 699, 768
Rozo, E., Vikhlinin, A., & More, S. 2012d, ApJ, 760, 67
Rozo, E., Wechsler, R. H., Rykoff, E. S., et al. 2010, ApJ, 708, 645
Rykoff, E. S., Koester, B. P., Rozo, E., et al. 2012, ApJ, 746, 178
Sehgal, N., Trac, H., Acquaviva, V., et al. 2011, ApJ, 732, 44
Sheldon, E. S., Johnston, D. E., Scranton, R., et al. 2009, ApJ, 703, 2217
Sifón, C., Menanteau, F., Hasselfield, M., et al. 2012, arXiv:1201.0991
Swetz, D. S., Ade, P. A. R., Amiri, M., et al. 2011, ApJS, 194, 41
Vikhlinin, A., Kravtsov, A. V., Burenin, R. A., et al. 2009, ApJ, 692, 1060
Voges, W., Aschenbach, B., Boller, Th., et al. 1999, A&A, 349, 389
Wang, L., & Steinhardt, P. J. 1998, ApJ, 508, 483
York, D. G., Adelman, J., Anderson, J. E., Jr., et al. 2000, AJ, 120, 1579