

Galaxy And Mass Assembly (GAMA): Galaxy colour gradients versus colour, structure, and luminosity (Research Note)

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ABSTRACT

Using single-component fits to SDSS/UKIDSS images of galaxies in the G09 region of the GAMA survey we study radial colour gradients across the galaxy population. We use the multi-wavelength information provided by MegaMorph analysis of galaxy light profiles to calculate intrinsic colour gradients, and divide into six subsamples split by overall Sérsic index (n) and galaxy colour. We find a bimodality in the colour gradients of high- and low- n galaxies in all wavebands which varies with overall galaxy luminosity. Global trends in colour gradients therefore result from combining the contrasting behaviour of a number of different galaxy populations. The ubiquity of strong negative colour gradients supports the picture of inside-out growth through gas accretion for blue, low- n galaxies, and through dry minor mergers for red, high- n galaxies. An exception is the blue high- n population which has properties indicative of dissipative major mergers.

Key words. galaxies: general – galaxies: structure – galaxies: fundamental parameters – galaxies: formation

1. Introduction

The global colours of galaxies can be driven by a number of factors, including their dust content, the average age of their stellar populations, and metallicity. The colour within a given galaxy also often varies radially, reflecting internal trends in the age, metallicity, and dust extinction of its stellar populations. The existence of negative colour gradients (i.e. redder in the centre and bluer in the outskirts) in elliptical galaxies is well-documented (see Peletier et al. 1990; Gonzalez-Perez et al. 2011, and references therein). La Barbera et al. (2010) find that these colour gradients can be linked to the overall colour and luminosity of an early-type galaxy (ETG); steeper negative colour gradients are more commonly found in bluer or more luminous ETGs. Steep negative colour gradients have also been observed in late-type galaxies (LTGs) and are thought to indicate the presence of young stars in their outer regions (Gonzalez-Perez et al. 2011). There are, however, inherent difficulties in measuring consistently accurate stellar population colours in LTGs due to interstellar extinction caused by dust. This dust extinction is most problematic in optical wavebands (see de Jong & van der Kruit 1994, and references therein). Moving to longer wavelengths combats this problem as observations in the near-infrared (NIR) are less affected by dust. The NIR is also better able to observe the older stellar populations within a galaxy, which contain most of the stellar mass (de Jong & van der Kruit 1994).

Recently, advances have been made in using colour information to constrain age and metallicity gradients in ETGs (e.g. La Barbera et al. 2010; Carter et al. 2011). Fewer studies of a similar nature have been attempted for spiral galaxies, partly due to their increased dust content compared to elliptical galaxies, and partly due to their more complex structure. Bell & de Jong (2000) find that the inner regions of most low-inclination spiral galaxies are older and more metal-rich than their outskirts which supports inside-out formation, although age gradients do not necessarily correlate with metallicity gradients (e.g. Tortora et al. 2010; Sánchez-Blázquez et al. 2014, and references therein). Similarly, $(g-r)$ and $(g-i)$ rest-frame global colours are mainly related to the age of the galaxy, and do not depend strongly on metallicity (Gonzalez-Perez et al. 2011).

In a previous work (Kennedy et al. 2015) we studied the wavelength dependence of galaxy structure to reveal the connections between the stellar populations within a galaxy and their spatial distributions. We found evidence for multi-component structure in many galaxies, even those visually classified as elliptical (e.g. Kennedy et al. 2016). This suggests that considering structural variations with wavelength may provide fundamental insights into galaxy formation. Nevertheless, studying internal colour gradients is a complementary and widely-used approach, and thus it is important to understand how our results on the wavelength-dependence of structure relate to more traditional colour gradients. In this Research Note we study how

colour gradient varies with wavelength (from ∇_{g-r} to ∇_{g-H}) for six subsamples split by overall Sérsic index and galaxy colour, and then consider the luminosity dependence of these colour gradients.

2. Data

Our sample comes from the G09 region of GAMA II (Driver et al. 2009, 2011; Liske et al. 2015), and is volume-limited to $M_r < -19.48$ mag and $z < 0.15$ when studying trends with luminosity (which gives a larger luminosity range but smaller sample of 5317 galaxies, with a stellar mass limit of $\sim 10^9 M_\odot$), and $M_r < -21.2$ mag and $z < 0.3$ elsewhere (which gives a larger sample of 13 825 galaxies at the expense of a smaller luminosity range with a stellar mass limit of $\sim 10^{10} M_\odot$). GALFITM (Häußler et al. 2013; Vika et al. 2013) fits a single-Sérsic wavelength-dependent model to all bands simultaneously, and returns (amongst other parameters) magnitude (m), Sérsic index (n), and effective radius (R_e). For more information on the data used here, see Häußler et al. (2013); Vulcani et al. (2014, hereafter V14); and Kennedy et al. (2015).

As in V14 and Kennedy et al. (2015) we divide our sample into high- and low-Sérsic index at $n_r = 2.5$, and separate red and blue galaxies using a $u-r = 2.1$ colour cut. We then separate the blue galaxies into green and blue at $u-r = 1.6$ in an attempt to separate out the bluest, potentially starburst, population. We note that our green population corresponds to the main population of star-forming galaxies, not specifically the green valley.

In our previous work we measured the wavelength dependence of Sérsic index, $N = n(H)/n(g)$, and effective radius, $R = R_e(H)/R_e(g)$ between the H - and g -bands. We saw that there is a striking difference in the behavior of high- n ($n_r > 2.5$) and low- n ($n_r < 2.5$) galaxies; high- n galaxies show a large change in R_e ($R < 1$) but little change in n ($N \sim 1$), whereas in low- n galaxies show little change in R_e with wavelength but a large change in n with wavelength. Although N and R are useful for considering the dependence of structure on wavelength, they are hard to compare to the literature. Here we therefore convert these measurements into more traditional colour gradients which describe a change in stellar population with radius, rather than a change in radial profile with wavelength.

We derive our colour gradients using the logarithmic slope of a galaxy's radial colour profile between $0.1R_e$ and $1R_e$ (see e.g. Peletier et al. 1990; La Barbera et al. 2010), taking R_e as the effective radius in the r -band. The colours as a function of radius are obtained from our multi-band Sérsic model fit to each galaxy. Both the Sérsic index and effective radius of a galaxy are allowed to vary quadratically with wavelength during the fit. As we measure the colour gradient over a radial range defined in a single band (the r -band), both the variations of n and R_e with wavelength contribute to our measured colour gradients. We then use orthogonal regression to find ∇_{g-x} , the gradient of the best-fitting profile where $x = rizYJH$. Our large sample size allows us to study both early- and LTGs in 7 wavebands. We can therefore provide a more complete picture of how colour gradients – and hence stellar mass growth channels – vary across the galaxy population. We are also able to extend the work of Kennedy et al. 2015 to study in detail the luminosity dependence colour gradients for our subsamples of galaxies.

3. Results

The variation in colour gradient with wavelength is summarised in Fig. 1; we see that the majority of colour gradients are

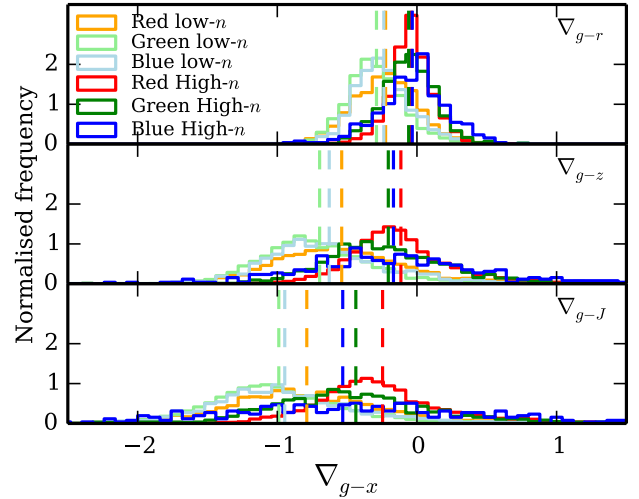


Fig. 1. Distributions of ∇_{g-x} for galaxies in each of our colour/Sérsic index subsamples (see Sect. 2 for more details), normalised to unit integral. Each panel shows the distribution for a different pair of bands. The median of each distribution is indicated by a vertical dashed line. These median values, and their uncertainties, can be found in Table 1. A bimodality in colour gradients for high- and low- n galaxies can be seen in all wavebands.

negative, i.e. appear bluer at larger radii. We have also examined the redshift dependence of these colour gradients in a similar manner to Kennedy et al. (2015), and find that they do not change significantly. As we would expect, colour gradients are stronger for more widely spaced waveband pairs. We see a distinct difference in the colour gradients of low- and high- n galaxies, regardless of overall galaxy colour. Our low- n samples consistently have the strongest colour gradients. In Kennedy et al. (2016) we found these low- n samples to contain galaxies with more significant disc components (i.e. a lower bulge/total flux ratio) than their high- n counterparts, which is in agreement with Gonzalez-Perez et al. (2011) who find that steep negative gradients are more likely to be found in late-types. We can also see that the red subsamples consistently have shallower colour gradients than the blue and green subsamples; once again we are in agreement with the finding of Gonzalez-Perez et al. (2011) that the redder a galaxy, the shallower its colour gradient. Our red, high- n sample is a reasonable proxy for ETGs, and the mean colour gradients we find (given in Table 1 for all our subsamples) agree well with the colour gradients found in La Barbera et al. (2010) for their sample of ETGs over the same range of wavebands. These values are later shown in Fig. 2. Within the uncertainties, the mean colour gradients of our low- n samples also agree with those of Taylor et al. (2005) for “mid-type” (Sc) spiral galaxies, and Gonzalez-Perez et al. (2011) for LTGs in a similar magnitude range.

As we saw in Fig. 1, different galaxy types (in this context, different n and colours) display contrasting colour gradients. We therefore show in Fig. 2 how colour gradients, ∇_{g-x} where $x = rizYJH$, change with luminosity depending on a galaxy's overall $u-r$ colour and Sérsic index, over two radial ranges.

The variety of trends seen in each panel of Fig. 2 shows that the overall trends seen in the literature are actually the result of a number of different populations. Red, high- n galaxies dominate at bright magnitudes (i.e. there are only red, high- n galaxies brighter than $M_r \sim -22.5$) whilst low- n galaxies dominate at faint magnitudes. We note that colour gradients in high- n galaxies do not change when measuring over a wider radius range (as previously found by La Barbera & de Carvalho 2009), for ETGs,

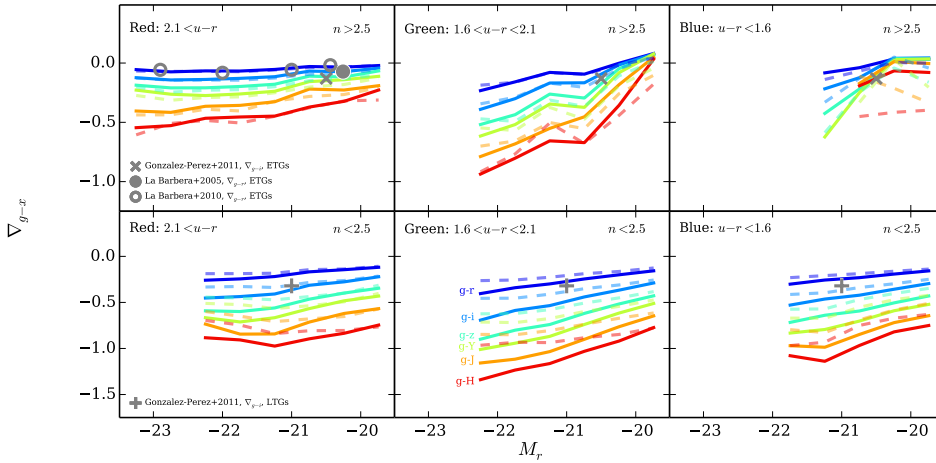


Fig. 2. Median colour gradient, ∇_{g-x} , where $x = rizYJH$, as a function of r -band absolute magnitude, M_r , in each of our colour/Sérsic index subsamples. Solid and dashed lines show colour gradients over the radius ranges $0.1R_e-1R_e$, and $0.1R_e-2R_e$, respectively. Lines are limited to magnitude bins containing at least 10 objects. They are therefore only shown where each subsample makes up a significant fraction ($\geq 1\%$) of the galaxy population. Different colours correspond to different wavebands, with $g-r$ at the top to $g-H$ at the bottom, as indicated. Results of previous studies are overlaid in the panels which most closely match the subsample of that study (see legends). We see that only the red, high- n population has galaxies brighter than $M_r \sim -22.5$, so these galaxies dominate at the bright end.

Table 1. Mean colour gradients, μ , and standard deviations, σ , for each of our colour/Sérsic index subsamples.

	Red, low- n		Green, low- n		Blue, low- n		Red, high- n		Green, high- n		Blue, high- n		Max. err.
	μ	σ	μ	σ	μ	σ	μ	σ	μ	σ	μ	σ	on μ
∇_{g-r}	-0.22	0.24	-0.29	0.21	-0.24	0.23	-0.04	0.16	-0.06	0.19	-0.03	0.25	0.01
∇_{g-i}	-0.39	0.40	-0.51	0.35	-0.44	0.38	-0.07	0.28	-0.12	0.35	-0.07	0.45	0.03
∇_{g-z}	-0.54	0.50	-0.70	0.45	-0.63	0.51	-0.12	0.40	-0.21	0.51	-0.17	0.72	0.04
∇_{g-Y}	-0.64	0.54	-0.82	0.50	-0.76	0.57	-0.16	0.46	-0.28	0.60	-0.29	0.92	0.05
∇_{g-J}	-0.79	0.58	-0.99	0.54	-0.95	0.62	-0.25	0.52	-0.44	0.71	-0.53	1.20	0.07
∇_{g-H}	-0.95	0.61	-1.12	0.55	-1.11	0.68	-0.40	0.52	-0.67	0.80	-0.79	1.44	0.08

Notes. The final column gives the maximum uncertainty on the mean values ($\sigma/\sqrt{N}$), corresponding to the smallest subsample, blue high- n . For the larger subsamples, this uncertainty is up to an order of magnitude smaller.

and shown as dashed lines in Fig. 2). Low- n galaxies do show slightly flatter colour gradients between $0.1R_e-2R_e$ which suggests that bright, low- n galaxies have stronger colour gradients towards their centres. This is consistent with the presence of a bulge within a relatively homogeneous outer disc, with the bulge being more significant in brighter galaxies.

La Barbera et al. (2010) find a “double-valued” behavior in the ∇_{g-r} colour gradient with luminosity for their ETGs. We can see in Fig. 2 that this double-valued behavior can actually be considered a combination of the trends seen in the low- and high- n red populations. The colour gradients seen in a population are sensitive to the selection criteria used to define that population.

Our red, high- n sample shows little change in colour gradient with luminosity in optical wavebands, but at longer wavelengths the faintest galaxies have the shallowest colour gradients. den Brok et al. (2011) study a sample of 142 ETGs and find that the colour gradients of elliptical galaxies (shown in Kennedy et al. 2016 to be analogous to our red, high- n population) become shallower at fainter magnitudes, which is consistent with our results. We note that only the red, high- n population has a significant number of galaxies brighter than $M_r \sim -22.5$ mag, so these galaxies dominate the brightest end of the sample. Tortora et al. (2010) use ∇_{g-i} as a measure of colour gradient, and find that colour gradients decrease with mass; i.e. more massive, and therefore more luminous, systems have the weakest colour gradients. This is consistent with our assumption that our high- n populations contain elliptical galaxies, which tend to be larger and more massive than spirals and S0s (Kelvin et al. 2012; Moffett et al. 2016).

On the whole, the low- n subsamples show little change in colour gradient vs. luminosity with wavelength (i.e. lines are approximately parallel to one another). However, in our high- n subsamples we see that the faintest galaxies show very similar colour gradients, whilst the brightest objects show a larger change in colour gradient with wavelength. La Barbera et al. (2010), who study a sample of ETGs which are most likely to correspond to our high- n subsamples, also find that the behavior of colour gradients with luminosity depends on both the wavebands in which the colour gradient was calculated, and the waveband in which the luminosity was determined.

Overplotted in grey in Fig. 2 we show the colour gradients for samples from some previous studies, in the panel to which their sample is most closely related. In the red high- n panel we show ∇_{g-r} for the La Barbera et al. (2005) sample of luminous ETGs in clusters, which lies within the standard deviation (not shown here, but can be inferred from Table 1) of the faint end of our ∇_{g-r} line. In this panel we also show with open circles ∇_{g-r} for a larger sample of ETGs with $ugrizYJHK$ photometry from La Barbera et al. (2010). The ∇_{g-r} luminosity dependence agrees well with our red, high- n sample, but sits a little lower at intermediate magnitudes. This is likely due to the ETG sample in La Barbera et al. (2010) also containing galaxies which lie in our red, low- n sample, which have stronger colour gradients than their high- n counterparts at $M_r \sim -22$. The ∇_{g-i} colour gradient for the Gonzalez-Perez et al. (2011) ETG sample lies directly on our ∇_{g-i} line at that r -band magnitude, whilst in the red, low- n panel we plot the Gonzalez-Perez et al. (2011) ∇_{g-i} for their brightest late-type sample which again agrees with our ∇_{g-i} for more discy objects.

4. Conclusions

We have measured radial gradients in six colours using multi-wavelength single-Sérsic model fits to galaxies in the GAMA survey. These complement our previous studies (Vulcani et al. 2014; Kennedy et al. 2015) on the variation of structural parameters with wavelength (N and R). Our measurements correspond well to those in the literature, supporting the reliability of our methods and our earlier conclusions regarding the wavelength-dependence of galaxy structure.

We find that galaxies with differing Sérsic index and total colour display contrasting behaviour in terms of both the distributions and luminosity dependence of their colour gradients. This means that care must be taken when comparing results for different sample selections.

The ubiquitous negative gradients we find are indicative of stellar populations becoming younger and/or lower metallicity as one progresses from the centres to the outskirts of galaxies. While this is consistent with simple models of galaxy formation via early dissipative collapse (e.g. Pipino et al. 2010), in our Λ CDM Universe we expect the colour gradients of most galaxies to reflect their long-term hierarchical growth.

Strong negative gradients in low- n systems are consistent with inside-out disc formation via gradual accretion of gas to their outskirts, via smooth flows or (gas-rich) minor mergers (e.g. Lemonias et al. 2011; Wang et al. 2011). Such gradients are present in blue, low-luminosity galaxies, with little bulge component; they must therefore arise from stellar population and/or dust gradients within the discs themselves. Nevertheless, the presence of a substantially redder bulge in more luminous disc galaxies (Kennedy et al. 2016), appears to further enhance colour gradients, particularly within $1R_e$.

The presence of gradients, even if weak, in bright, high- n systems argues against these galaxies being the direct result of major mergers, since such violent interactions should flatten colour gradients (see Kim & Im 2013, and references therein). Instead, their gradients may either result from the fading of bluer and/or lower- n galaxies (e.g. Bell et al. 2006), or from a reintroduction of metal-poor and/or younger stellar populations in the outskirts of these galaxies, for example via dry minor mergers (Eliche-Moral et al. 2013). The latter can also account for the rapid size evolution of passive galaxies (Oser et al. 2010).

In contrast, lower-luminosity high- n galaxies with bluer colours display very flat, or even slightly positive, colour gradients, especially in the optical. We infer that this population is experiencing a period of central star-formation. These galaxies account for 2.4% of our sample, which corresponds well with the local major merger fraction: also $\sim 2\%$ across a wide range of galaxy masses (see Khochfar & Burkert 2001; Keenan et al. 2014, and references therein). At the luminosities of interest, major mergers are mostly mixed elliptical-spiral interactions, and hence gas-rich (see e.g. Khochfar & Burkert 2003; Darg et al. 2010). We therefore associate the blue high- n population with dissipative major mergers. However, less violent causes of central star formation, such as bar instabilities and minor mergers (see e.g. Kormendy & Kennicutt 2004, and references therein), may also contribute. The rarity of these examples suggests their descendants will re-acquire negative colour gradients, via fading of the central star-formation and subsequent accretion.

The colour profiles of most galaxies therefore result from continued accretion, with the ratio of gas to stars in the incoming material varying from high: in the case of lower-luminosity, lower- n and bluer galaxies; to low: for higher-luminosity, higher- n and redder galaxies.

Possible applications for our colour gradient measurements include multi-wavelength strong lensing (information about the likely internal colour gradient of a galaxy will allow tighter constraints on mass models), and comparison with galaxy simulations (colour gradients provide more information with which to challenge models). Our future work will use decomposed bulge and disc properties to further study the relationships between stellar populations and galaxy structure. In particular, we will expand upon the present study by using multi-band method to measuring colour gradients for individual galaxy components, e.g. bulges and discs, separately.

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References

- Bell, E. F., & de Jong, R. S. 2000, *MNRAS*, **312**, 497
- Bell, E. F., Naab, T., McIntosh, D. H., et al. 2006, *ApJ*, **640**, 241
- Carter, D., Pass, S., Kennedy, J., Karick, A. M., & Smith, R. J. 2011, *MNRAS*, **414**, 3410
- Darg, D. W., Kaviraj, S., Lintott, C. J., et al. 2010, *MNRAS*, **401**, 1043
- de Jong, R., & van der Kruit, P. 1994, *A&AS*, **106**, 451
- den Brok, M., Peletier, R. F., Valentijn, E. A., et al. 2011, *MNRAS*, **414**, 3052
- Driver, S. P., Norberg, P., Baldry, I. K., et al. 2009, *Astron. Geophys.*, **50**, 5.12
- Driver, S. P., Hill, D. T., Kelvin, L. S., et al. 2011, *MNRAS*, **413**, 971
- Eliche-Moral, M. C., González-García, A. C., Aguerri, J. A. L., et al. 2013, *A&A*, **552**, A67
- Gonzalez-Perez, V., Castander, F. J., & Kauffmann, G. 2011, *MNRAS*, **411**, 1151
- Häußler, B., Bamford, S. P., Vika, M., et al. 2013, *MNRAS*, **430**, 330
- Keenan, R. C., Foucaud, S., De Propriis, R., et al. 2014, *ApJ*, **795**, 157
- Kelvin, L. S., Driver, S. P., Robotham, A. S. G., et al. 2012, *MNRAS*, **421**, 1007
- Kennedy, R., Bamford, S. P., Baldry, I., et al. 2015, *MNRAS*, **454**, 806
- Kennedy, R., Bamford, S. P., Häußler, B., et al. 2016, *MNRAS*, **460**, 3458
- Khochfar, S., & Burkert, A. 2001, *ApJ*, **561**, 517
- Khochfar, S., & Burkert, A. 2003, *ApJ*, **597**, L117
- Kim, D., & Im, M. 2013, *ApJ*, **766**, 109
- Kormendy, J., & Kennicutt, Jr., R. C. 2004, *ARA&A*, **42**, 603
- La Barbera, F., & de Carvalho, R. R. 2009, *ApJ*, **699**, L76
- La Barbera, F., de Carvalho, R. R., Gal, R. R., et al. 2005, *ApJ*, **626**, L19
- La Barbera, F., De Carvalho, R. R., De La Rosa, I. G., et al. 2010, *AJ*, **140**, 1528
- Lemonias, J. J., Schiminovich, D., Thilker, D., et al. 2011, *ApJ*, **733**, 74
- Liske, J., Baldry, I. K., Driver, S. P., et al. 2015, *MNRAS*, **452**, 2087
- Moffett, A. J., Ingarfield, S. A., Driver, S. P., et al. 2016, *MNRAS*, **457**, 1308
- Oser, L., Ostriker, J. P., Naab, T., Johansson, P. H., & Burkert, A. 2010, *ApJ*, **725**, 2312
- Peletier, R., Valentijn, E., & Jameson, R. 1990, *A&A*, **233**, 62
- Pipino, A., D'Ercole, A., Chiappini, C., & Matteucci, F. 2010, *MNRAS*, **407**, 1347
- Sánchez-Blázquez, P., Rosales-Ortega, F. F., Méndez-Abreu, J., et al. 2014, *A&A*, **570**, A6
- Taylor, V. A., Jansen, R. A., Windhorst, R. A., Odewahn, S. C., & Hibbard, J. E. 2005, *ApJ*, **630**, 784
- Tortora, C., Napolitano, N. R., Cardone, V. F., et al. 2010, *MNRAS*, **407**, 144
- Vika, M., Bamford, S. P., Häußler, B., et al. 2013, *MNRAS*, **435**, 623
- Vulcani, B., Bamford, S. P., Häußler, B., et al. 2014, *MNRAS*, **441**, 1340
- Wang, J., Kauffmann, G., Overzier, R., et al. 2011, *MNRAS*, **412**, 1081