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Remnant radio-loud AGN in the Herschel-ATLAS field

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ABSTRACT

Only a small fraction of observed Active Galactic Nuclei display large-scale radio emission associated with jets, yet these radio-loud AGN have become increasingly important in models of galaxy evolution. In determining the dynamics and energetics of the radio sources over cosmic time, a key question concerns what happens when their jets switch off. The resulting ‘remnant’ radio-loud AGN have been surprisingly evasive in past radio surveys, and therefore statistical information on the population of radio-loud AGN in their dying phase is limited. In this paper, with the recent developments of LOFAR and the VLA, we are able to provide a systematically selected sample of remnant radio-loud AGN in the Herschel-ATLAS field. Using a simple core-detection method, we constrain the upper limit on the fraction of remnants in our radio-loud AGN sample to 9 per cent, implying that the extended lobe emission fades rapidly once the core/jets turn off. We also find that our remnant sample has a wide range of spectral indices ($-1.5 \leq \alpha_{150}^{1400} \leq -0.5$), confirming that the lobes of some remnants may possess flat spectra at low frequencies just as active sources do. We suggest that, even with the unprecedented sensitivity of LOFAR, our sample may still only contain the youngest of the remnant population.

Key words: galaxies: active – galaxies: jets – radio continuum: galaxies – methods: statistical

1 INTRODUCTION

1.1 Radio-loud AGN and their evolution

The large number of Active Galactic Nuclei (AGN) in flux-limited radio surveys has led to valuable statistical information on the population of radio-loud AGN (e.g. Best et al. 2007; Best & Heckman 2012). Multi-wavelength observations of radio-loud AGN, capturing the large-scale radio lobes inflated by the jets, have also given important constraints on the dynamics and energetics of their extended emission and the effects on their surrounding environment (Hardcastle et al. 2002; Croston et al. 2004, 2005). These observations provide tests of (semi-)analytical and numerical models describing the time evolution of powerful FR-II radio galaxies (Fa-

naroff & Riley 1974) and their environmental impact (e.g. Kaiser & Alexander 1997; Blundell & Rawlings 1999; Laing & Bridle 2002; Hardcastle & Krause 2013). In particular, it has been suggested that powerful sources in dense, cluster environments re-heat their surrounding medium as a mechanism to offset cluster cooling – solving the well known ‘cooling flow’ problem in clusters of galaxies (Basson & Alexander 2003; Dunn & Fabian 2006; McNamara & Nulsen 2007; Croston et al. 2008), although direct evidence of this process for the population of radio-loud AGN in the most massive clusters remains elusive (Best et al. 2007). Moreover, it has been suggested that radio-loud AGN disrupt star formation in their host galaxies through gas heating (Best et al. 2006; Wylezalek & Zakamska 2016), regulating the growth of galaxies through the well known AGN feedback cycle (see review by Fabian 2012). These advances have effectively driven the development of a clear picture

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describing the evolution of radio galaxies throughout their ‘active’ lifetimes.

However, much is still not understood regarding the nature of *remnant AGN*, i.e the phase of the AGN life cycle that begins when the jets switch off. Over cosmic time, AGN are expected to have episodic nature (Saikia & Jamrozy 2009; Morganti 2017). Clear evidence of episodic activity has come from observations of radio galaxies with two pairs of radio lobes denoting the inner currently active radio lobes, and the faint outer radio lobes from previous activity, termed ‘double-double’ radio galaxies (DDRGs; Roettiger et al. 1994; Subrahmanyan et al. 1996; Schoenmakers et al. 2000; Saripalli et al. 2002, 2003). Only a few such sources have been observed (e.g. Brocksopp et al. 2011; Nandi & Saikia 2012; Orrù et al. 2015), and this scarcity can be explained by a model in which the currently active radio jets drive into the pre-existing plasma on relatively short time-scales, and therefore merge with the outer remnant lobes (Konar & Hardcastle 2013). Remnant radio-loud AGN, with no evidence for currently active jets, are surprisingly even rarer than DDRGs (Parma et al. 2007; Murgia et al. 2011; Saripalli et al. 2012). During the remnant phase, the bright radio core and jets are expected to disappear, while the lobes may radiate for some time (Slee et al. 2001). Roughly half of the energy ever transported through the jets may remain in the radio-emitting lobes when the jets switch off (Hardcastle & Krause 2013). Whether or not this energy remains in the lobes in the post switch-off phase, for how long it remains there and where the plasma ends up are questions of great interest in studies of the AGN duty cycle.

Knowledge of the dynamics of remnants will also aid analytical models describing the dynamical evolution of the radio plasma in the post switch-off phase. From a physical point of view, as an active radio galaxy expands into its environment, the lobe plasma experiences adiabatic and radiative losses. Meanwhile, the jet continuously replenishes the lobes with new and young electrons which will eventually go through the same loss cycles. The fact that remnant sources have a very low detection rate (e.g. Giovannini et al. 1988) can be explained if the radio lobes quickly become undetectable due to rapid energy losses¹ without the replenishment of fresh electrons by an active jet, as well as adiabatic losses as the lobes continue to expand. Recent analytical modelling by Godfrey et al. (2017) and Brienza et al. (2017) have shown that models of radio galaxy evolution that consider only radiative losses over-predict the number of observed ultra-steep spectrum remnants by at least a factor of two. This further supports the idea that expansion losses in radio galaxies are important in the remnant phase, which might explain why they may quickly escape detection in current flux-limited radio surveys.

It is therefore important to constrain the fading time of remnants from observations. The number of remnants in the sky relative to the number of active sources in a single observation will give us some measure of this, leading to a robust, systematic study of remnant radio-loud AGN as a population. Such a sample will provide a unique opportunity to understand their dynamics, and also aid the development of analytical and numerical models describ-

¹ It is interesting to note that inverse-Compton emission due to the up-scattering of CMB photons by the ageing low-frequency electrons in the lobes can last for a significantly longer period than synchrotron losses at low redshift – giving rise to IC ‘ghosts’ seen in the X-ray (Mocz et al. 2011). Deep X-ray surveys may be able to find remnants previously undetected at radio-wavelengths through this method, although a sample is yet to be developed.

ing the duty cycle of radio-loud AGN – an accurate description of which is currently missing in galaxy evolution models.

1.2 Remnant selection methods

The scarcity of continuum observations of remnant AGN has previously restricted the possibility of a detailed study in a statistical sense. Remnants are expected to be steep-spectrum in nature even below 1 GHz (Giacintucci et al. 2007; Parma et al. 2007; Murgia et al. 2011), owing to the radiative cooling of the lobe plasma without the input of high energy electrons by a jet. However, solely selecting low-frequency steep-spectrum sources (e.g. $\alpha_{150}^{1400} \leq -1.2$, defining the spectral index as the slope of the radio spectrum between two given frequencies in the sense $S \propto \nu^\alpha$) may be biased in the sense that some remnant sources may in fact show flatter spectra (Brienza et al. 2016), while it is also possible for active radio galaxies to possess similarly steep low-frequency spectra (Harwood et al. 2013, 2015), which may contaminate the remnant sample. Moreover, Godfrey et al. (2017) have shown that spectral selection methods preferentially select the oldest remnants – a fraction of the remnant population which may form an unrepresentative sample. Morphological criteria have also been suggested to select remnants (Saripalli et al. 2012), since the ageing large scale lobe emission should generally show a relaxed structure (Saripalli et al. 2012), although recently switched-off FR-II sources can contain bright hotspots (for as long as the light travel time of the radio galaxy – $\gtrsim 1$ Myr for the largest sources), and display less relaxed and more energetic lobes that are typical of active sources (e.g. 3C28; Harwood et al. 2015). Brienza et al. (2016) suggest that morphological, spectral index and other selection methods should be used in conjunction to give a systematic and reliable sample of the remnant population, although multi-frequency observations at comparable resolution can be difficult to obtain for a large sample of sources.

The absence of a core associated with nuclear activity in radio images is a clear signature of a remnant radio galaxy (Giovannini et al. 1988). All genuine remnants are expected to lack a visible radio core at all frequencies – the fading time for the jets, which are visible as the core on pc-scales, is as long as the light travel time of material through the jet ($\sim 10^4$ yr), which is a small fraction of the typical fading time of the large scale radio lobes. Recently, Godfrey et al. (2017) have used the absence of a core in conjunction with ultra-steep spectra as a method to select a sample of candidate remnant FR-II radio galaxies with a flux density limit of > 1.5 Jy from the 74 MHz VLA Low-Frequency Sky Survey Redux catalogue (VLSSr). Two per cent of sources were selected using their sampling criteria as candidate remnants, although this is expected to only be an upper limit, since the steep-spectrum criterion is expected to contaminate the remnant sample with high- z active sources. Selecting sources with low surface-brightness lobes in the absence of a core could form an alternative, unbiased selection criterion.

However, it is difficult to observe faint lobe emission in radio surveys with a high flux limit, as they are sensitive to the brightest radio sources (e.g. 3CRR; Laing et al. 1983), meaning that many remnants may be missed by the sensitivity limits of these surveys. Detailed studies of remnants in the past have been limited to individual sources (e.g. BLOB1, Brienza et al. (2016); PKS B1400-33, Subrahmanyan et al. (2003); B2 0258+35, Shulevski et al. (2012)). Sensitive, low-frequency radio observations over large areas of the sky are required to observe a larger fraction of the faint radio galaxy population associated with the remnant phase. A robust detection of a core will differentiate sources between being remnant and ac-

tive radio galaxies, and hence this method is expected to provide an unbiased sample of remnants.

1.3 Selecting remnants with LOFAR

Faint radio lobes from remnant radio-loud AGN are expected to be detected with the Low-frequency Array (LOFAR; van Haarlem et al. 2013) operating at around 150 MHz, where the ageing radio lobes would be brighter than at GHz frequencies due to the preferential cooling of higher-energy electrons (Kardashev 1962; Pacholczyk 1970). Crucially, LOFAR has the resolution (~ 6 arcsec at 150 MHz with the full Dutch array) and *uv*-coverage to make sensitive and deep surveys of the sky. The mixture of long and short baselines in a single observation and the wide field of view enables a large number of potential remnant sources to be detected. LOFAR thus gives us the opportunity to produce a much-needed systematic survey of the remnant radio-loud AGN population.

Brienza et al. (2017) have recently made use of LOFAR observations of the Lockman Hole field to assess the efficiency of various spectral and morphological criteria in selecting remnant radio galaxies. An initial sample of extended radio galaxies from the LOFAR images was obtained, and cross-matching the sources with the Faint Images of the Radio Sky at Twenty-cm (FIRST; Becker et al. 1995) survey was performed. Those sources that did not show significant evidence of a radio core based on the FIRST images were deemed candidate remnants. Interestingly, the remnant fraction obtained in this sample was 30 per cent. This remnant fraction, if robust, would be considerably higher than that found in the 3CRR sample, which would strengthen the original hypothesis that remnants are more detectable with sensitive telescopes such as LOFAR. Moreover, it implies a much longer remnant lifetime than that generally assumed, potentially giving important constraints on the energy loss processes in these sources. Further samples of candidate remnants obtained using this method would make this result more robust.

An exploratory LOFAR High-Band Antenna (HBA, 110-200 MHz) survey of the *Herschel*-ATLAS North Galactic Pole (H-ATLAS NGP) field has also been carried out recently (Hardcastle et al. 2016, hereafter H16). This sky region covers approximately 142 square degrees centred around RA = 13.5 h and DEC = 30°. Around 15,000 discrete radio sources were detected using observations between 126-173 MHz using the full Dutch array, giving a resolution of $\sim 10 \times 6$ arcsec. For identification of the radio sources in this sample, visual cross-matching with the Sloan Digital Sky Survey Data Release 12 (SDSS DR12; Alam et al. 2015) at *r*-band to find the most likely optical host galaxy, assisted by any 1.4 GHz FIRST core detection, was performed. For sources without FIRST core detections, IDs were made based on the morphology of the LOFAR emission alone. Bright radio sources associated with star formation, which are indistinguishable from radio-loud AGN based on the LOFAR data alone, were separated from the sample using the FIR-radio correlation aided by *Herschel* data. See H16 for full details of observations, data calibration and AGN/star formation separation techniques.

As an initial search for potential remnants, a parent sample of radio-loud AGN selected from the LOFAR observations was made satisfying the following criteria: bright (>80 mJy at 150 MHz), ensuring that the sample is flux-complete while also allowing a faint but active core to be detected from follow-up higher-resolution observations; well resolved (>40 arcsec) to show extended emission; and classed as AGN based on the FIR-radio correlation. Of 127 such sources, 38 sources were then selected as *candidate* remnants

on the basis that there was no evidence of a core detected with FIRST – giving a potential remnant fraction of 30 per cent, in agreement with Brienza et al. (2017).

However, the remnant fraction obtained from LOFAR is limited by the fact that FIRST is not particularly sensitive to the detection of compact cores, meaning that the remnant fractions of H16 and Brienza et al. (2017) must be regarded as only an upper limit. At a resolution of ~ 5 arcsec, FIRST has a typical source detection threshold of 1 mJy. However, if we define the core prominence as the ratio of the FIRST core flux density at 1.4 GHz and the total flux density at 150 MHz (the exact frequency for the measurement of a core flux density is unimportant, since the cores of radio-loud AGN are expected to be flat-spectrum out to high frequencies (Hardcastle & Looney 2008)), then the 3σ upper limit for the core prominence for the faintest objects in the sample detected with FIRST is $0.480 \approx 5 \times 10^{-3}$. On the other hand, the median value of core prominence in the brighter 3CRR sample is $\sim 3 \times 10^{-4}$ (Mullin et al. 2008). Although 3CRR selects for the brightest sources, and so would be expected to have systematically low core prominences relative to that of the LOFAR sample, it is still clearly possible that faint radio cores will be missed in the LOFAR sample if only FIRST is used to identify them. To decide whether this is simply a result of the sensitivity limits of FIRST, or of these sources actually having no radio-bright cores denoting nuclear activity, requires more sensitive and higher-resolution observations. Furthermore, without the detection of a radio core, optical ID cross-matching clearly becomes a challenge, and therefore we cannot rely on the optical ID for sources without a FIRST-detected core.

In this paper, we present new 6 GHz Karl G. Jansky Very Large Array (VLA) snapshot observations of the 38 LOFAR candidate remnants obtained by H16 (sources listed in Table 1). With a higher spatial resolution (~ 0.3 arcsec for the longest baselines at 6 GHz) and an order of magnitude improvement in sensitivity than that of FIRST, we can obtain more robust detections/non-detections of a core. A clear detection of a core in the new 6 GHz images will mean that the corresponding source is not a remnant, and conversely a non-detection of a core will retain the candidate remnant status for the source, thereby giving a more accurate constraint on the remnant fraction. We will then be able to provide a systematic sample of the candidate remnant population in the NGP H-ATLAS field, leading to important statistical information regarding the population of remnant radio-loud AGN. Furthermore, the detection of a compact radio core with the VLA at the position of the host galaxy of the source will enable previously ambiguous optical identifications to be confirmed or rejected.

The remainder of this paper is structured as follows. In Section 2.1, we describe the VLA observations and data reduction processes for the 38 remnant candidate sources from the H-ATLAS NGP field. In Section 2.3 we present the images, followed by a qualitative analysis of each candidate source, outlining their remnant or active status, as well as any optical misidentification, throughout Section 3. We discuss the newly constrained remnant fraction in Section 4, and its implications for the dynamical evolution of radio-loud AGN. We then conclude with a brief summary of the results in Section 5.

Throughout this work we use cosmological parameters based on a Λ CDM cosmology with $H_0 = 70$ km s $^{-1}$ Mpc $^{-1}$, $\Omega_m = 0.3$ and $\Omega_\Lambda = 0.7$. Co-ordinate positions are given in the J2000 system. Spectral indices are defined in the sense $S_\nu \propto \nu^\alpha$.

2 OBSERVATIONS

2.1 VLA Data Reduction

The 38 remnant candidates were observed with the VLA on the 30th September 2016 in the A-configuration and on the 8th September 2016 in the B-configuration (Table 2). Observations were made in the broad-band C-band system in 3-bit mode (4 GHz bandwidth from 4–8 GHz). Since we only required snapshot observations, exposure times were ~ 5 minutes per source, reaching background rms levels of $\sim 10 - 15 \mu\text{Jy}/\text{beam}$ in the final combined images – an improvement in sensitivity to a radio core by an order of magnitude over FIRST. The observations were targeted at the SDSS optical ID for each source (given by H16). The sources lie in a roughly 12×12 -degree region of the sky, including the quasar 3C286 which was used as the flux calibrator. The blazar J1310+3233 was used as the complex gain calibrator, observed 13 times between different scans of target sources. These details are summarised in Table 2.

Prior to calibration the AOFFLAGGER software (Offringa et al. 2012) was used on the data to automatically flag radio-frequency interference (RFI). The data were then reduced using the CASA VLA pipeline version 1.3.5 for reduction using CASA version 4.5.0 (McMullin et al. 2007). A selection of gain calibration tables were inspected to check the quality of the calibration and/or the occurrence of bad baselines that were not automatically flagged previously, prior to imaging. As a secondary flagging process, the CASA tool ‘rflag’ was also applied to both measurement sets to remove additional RFI present in the uv data. Initial images were then produced by CLEANing (Högbom 1974) both A and B-configuration measurement sets separately for each target source, using the imaging parameters detailed in Table 3. For J131040.25+322044.1 (Figure 1.18), a bright quasar in the field prevented suitable deconvolution and resulted in high RMS noise levels in the vicinity of the optical ID. For this source only, the data were calibrated manually, and imaged using the imager WSCLEAN (Offringa et al. 2014) to reduce errors due to poor deconvolution in the wide-field image. Self-calibration was possible for some of the targets with strong core detections, and in the event of non-detections, bright sources present nearby in the field were used. The individually self-calibrated A and B-configuration data sets were then combined in the uv plane, producing the final radio maps with an average background noise level of $\sim 14.6 \mu\text{Jy}/\text{beam}$, shown in Figure 1. For presentation purposes, we smoothed the VLA maps with a Gaussian function having a FWHM of 0.2 arcsec (~ 5 pixels across).

A software problem with the VLA regarding the delay model used for correlation between antennas potentially introduced a problem into observations made between 9th August and 14th November 2016 in the A-configuration. The main effect of this subtle bug would have been a displacement of a source’s position in the direction of elevation. This might have affected our analysis of the data since our science aims depend on our ability to confirm or reject optical IDs based on the position of a VLA core detection, although a potential source offset was only deemed to be serious for sources below a declination of 20 degrees and all of our sources are in fact at higher declination (Table 1). Nevertheless we estimated the magnitude of the potential offset for each target source, and for each source where the calculated offset was larger than the beam-size (5/38 sources), we visually inspected the A-configuration image overlaid on the corresponding SDSS image. We found that, for those targets with detected cores, there was a positional match with the SDSS optical ID. For those without cores, we ensured that the magnitude of the potential offset would not have allowed the target

source to be associated with a different optical host. We concluded that this software problem was not manifested in our observations at a level that would have required re-observation.

2.2 LOFAR data

The LOFAR data used are those presented by H16, but for this paper we use a new direction-dependent calibration procedure. This processing of the H-ATLAS data will be described in more detail elsewhere but, to summarize briefly, it involves replacing the facet calibration method described by H16 with a direction-dependent calibration using the methods of Tasse (2014a,b), implemented in the software package KILLMS, followed by imaging with a newly developed imager DDFACET (Tasse et al. 2017 in prep.) which is capable of applying these direction-dependent calibrations in the process of imaging. The H-ATLAS data were processed using the December 2016 version of the pipeline, DDF-PIPELINE², that is under development for the processing of the LoTSS survey (Shimwell et al. 2017, and in prep.). The main advantage of this reprocessing is that it gives lower noise and higher image fidelity than the process described by H16, increasing the point-source sensitivity and removing artefacts from the data, but it also allows us to image at a slightly higher resolution – the images used in this paper have a 7-arcsec restoring beam. Note that, due to the increased point-source sensitivity and reduction in noise, the most faint and diffuse sources of emission are less well represented than in the previous images. Therefore, for sources J125422.44+304 (Figure 1.4), J132602.42+314 (Figure 1.27) and J132622.56+320 (Figure 1.28), we use the pre-processed LOFAR data presented by H16, for a better visual representation of the source structure. We ensured that these sources still met our sample criteria detailed in Section 1, based on the new re-processed LOFAR data.

After careful visual inspection of the re-processed LOFAR images, we concluded that sources J130849.75+252 (Figure 1.14), J131537.33+310 (Figure 1.23), J131828.97+291 (Figure 1.24) and J131832.33+291 (Figure 1.25) are either imaging artefacts that were misidentified as real sources based on the lower resolution LOFAR images presented by H16, or real sources that now do not meet our flux and angular extent sample criteria. We therefore removed these sources from our sample, leaving our parent AGN sample at 123 sources, and our candidate remnant sample at 34 sources.

2.3 Radio images

In Figure 1 we present the VLA 6 GHz images overlaid with the LOFAR 150-MHz contours (contour levels at increasing powers of 2σ) in order to view the faint low-frequency plasma of the candidate remnants, or the lobes of potentially active sources. To emphasise the visibility of a potential faint and/or compact radio core, we provide a zoom-in of the VLA image at the location of the current optical ID (marked with a cross-hair) on the top-right hand side of each image. We define all 3σ core detections at the optical ID to be active sources. Note that the non-detection of a core ($< 3\sigma$) at the optical ID as seen by these zoom-ins mean that the source may have been optically misidentified, and that there may be a faint active core elsewhere in the 6 GHz map at a location corresponding to a different optical ID associated with the radio source. In this

² See <http://github.com/mhardcastle/ddf-pipeline> for the code.

Table 1. Candidate remnant sources. SDSS ID gives the name of the SDSS source corresponding to the optical host. RA and DEC specify the co-ordinates of this optical ID (also the location of VLA pointing for that source). LOFAR flux density gives the flux density of the radio source at 150 MHz associated with the optical ID. Redshifts are either photometric (p) or spectroscopic (s), as stated below, and where both were available the spectroscopic redshift is given.

LOFAR name	SDSS ID	RA (h:m:s)	DEC (°:′:″)	LOFAR 150 MHz flux density (Jy)	Redshift (z)	Redshift type (p/s)
J125143.00+332020.2	1237665228382077637	12:51:42.90	+33:20:20.66	0.17843	0.3866	p
J125147.02+314046.0	1237665226234659106	12:51:47.02	+31:40:47.66	0.21616	0.3579	s
J125311.08+304029.2	1237665443125723293	12:53:11.62	+30:40:17.35	0.09572	0.3497	s
J125419.40+304803.0	1237667255629643949	12:54:22.44	+30:47:28.10	0.28814	0.1949	p
J125931.85+333654.2	1237665023834915425	12:59:30.80	+33:36:46.96	0.11063	0.5367	p
J130003.72+263652.1	1237667442438111406	13:00:04.24	+26:36:52.70	0.18476	0.3164	s
J130013.55+273548.7	1237667323797766776	13:00:14.65	+27:36:00.07	0.12813	0.3228	p
J130332.38+312947.1	1237665226235773388	13:03:32.47	+31:29:49.54	0.08529	0.7412	s
J130415.46+225322.6	1237667736123933309	13:04:15.96	+22:53:43.49	0.26408	0.3723	p
J130532.49+315639.0	123766522672775688	13:05:32.01	+31:56:34.86	0.13407	0.3318	p
J130548.65+344052.7	1237665025446052014	13:05:48.81	+34:40:53.84	0.19518	0.2318	p
J130640.99+233824.8	1237667446197649528	13:06:41.12	+23:38:23.49	0.16415	0.1829	s
J130825.47+330508.2	1237665228383650573	13:08:26.27	+33:05:15.05	0.09703	0.4298	p
J130849.22+252841.8	1237667912745550347	13:08:49.74	+25:28:40.23	0.08652	0.4157	p
J130916.85+305118.3	1237665225699492186	13:09:16.02	+30:51:21.94	0.11871	0.3459	p
J130915.61+230309.5	1237667783913439350	13:09:16.66	+23:03:11.37	0.18042	0.2281	s
J130917.63+333028.4	1237665023835833039	13:09:17.73	+33:30:35.68	0.27835	0.4922	p
J131040.25+322044.1	1237665330930778227	13:10:40.03	+32:20:47.66	0.33559	0.5517	s
J131039.92+265111.9	1237667442976030877	13:10:40.51	+26:51:05.49	0.08053	0.1853	s
J131235.35+331348.6	1237665126938117173	13:12:36.14	+33:13:38.99	0.08448	0.7846	p
J131405.16+243234.1	1237667911672333046	13:14:05.90	+24:32:40.38	0.25808	0.3806	p
J131446.57+252819.8	1237667448882725588	13:14:46.82	+25:28:20.51	0.10939	0.5438	s
J131536.30+310615.6	1237665226236887706	13:15:37.33	+31:06:15.61	0.09244	0.2775	p
J131827.83+291658.5	1237665442054471904	13:18:28.96	+29:17:26.48	0.13550	0.3027	p
J131833.81+291904.9	1237665442054471911	13:18:32.32	+29:18:38.94	0.08778	0.2705	p
J132402.51+302830.1	1237665225700868124	13:24:03.44	+30:28:22.86	0.35482	0.0488	p
J132602.06+314645.6	1237665227311546592	13:26:02.42	+31:46:50.42	0.12590	0.2370	s
J132622.12+320512.1	1237665227848482831	13:26:22.56	+32:05:02.36	0.16374	0.3489	p
J132738.77+350644.6	1237664671644517294	13:27:35.32	+35:06:36.73	0.83581	0.5003	p
J132949.25+335136.2	1237665128550236906	13:29:48.87	+33:51:52.22	0.14627	0.5601	s
J133016.12+315923.9	1237665227848811145	13:30:16.32	+31:59:19.77	0.20335	0.3106	p
J133058.91+351658.9	1237664852028948635	13:30:57.33	+35:16:50.29	0.22973	0.3158	s
J133309.94+251045.2	1237667321653625933	13:33:10.56	+25:10:44.12	0.10982	0.2639	p
J133422.15+343640.5	1237665129624371524	13:34:22.21	+34:36:34.79	0.13971	0.5575	p
J133502.36+323312.8	1237665023838257509	13:35:02.38	+32:33:13.94	0.24344	0.4282	p
J133642.53+352009.9	1237664852566278355	13:36:43.09	+35:20:11.72	0.13395	0.1145	s
J134702.03+310913.3	1237665330934186147	13:47:01.71	+31:09:24.21	0.28446	0.1995	p
J134802.83+322938.3	1237665024376307813	13:48:02.70	+32:29:40.10	0.13825	0.2101	s

Table 2. VLA observation details

VLA project code	Array	Bandwidth (GHz)	Obs. date	Duration	Flux calibrator	Average RMS (μ Jy/beam)
16B-245	A	4-8	30/09/16	3.4h	3C286	19.9 ¹
16B-245	B	4-8	08/09/16	3.4h	3C286	19.9 ¹

¹ Background RMS taken from radio maps prior to self-calibration, and therefore some maps with bright background sources contain high noise levels.

case, we confirm the active nature of the source and also propose a new SDSS ID, as displayed in Figure 2 and detailed in Table 4. Otherwise, in the absence of a core, we keep the candidate remnant status for the source as given by H16. Physical descriptions of each source, along with confirmation of remnant statuses and optical IDs, based on these images, are given in Section 3.1. The four sources that we removed from the candidate remnant sample of H16 are also shown in Figure 1, and are labelled in their figure caption. Note that in the case of confirming a new optical ID, and hence a different redshift for the source, the physical scale bar on the bottom right of each image is incorrect. We correct these phys-

ical parameters for these few sources in our analysis (Section 4).

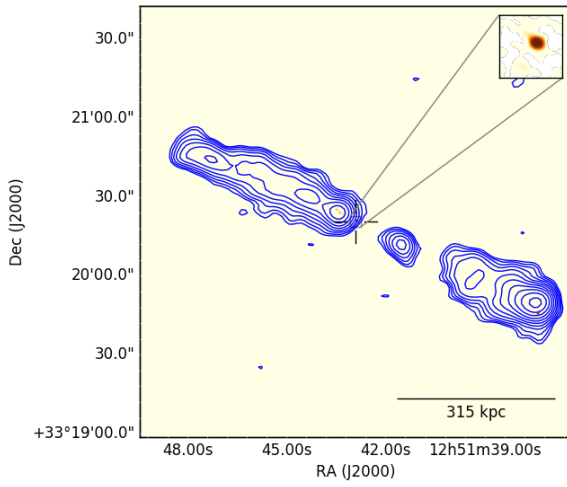
Table 3. Summary of CASA imaging parameters for the 6 GHz VLA observations. Shown are the CLEAN parameters used to image the visibilities from both A- and B-configuration measurement sets.

Parameter	CASA name	Value	Units
Cell size	cellsize	0.04×0.04	arcsec
Image size ¹	imsize	4096×4096	pixels
Noise threshold	noise	0.01	mJy/beam
Weighting ²	robust	0.0	
Average beam major axis ³	BMAJ	0.47	arcsec
Average beam minor axis ³	BMIN	0.28	arcsec

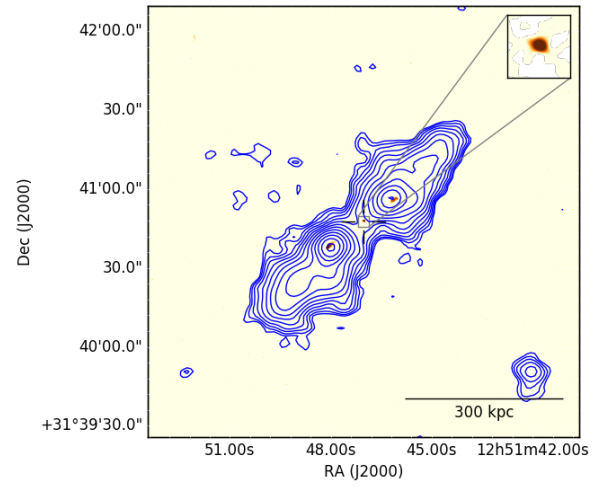
¹ The images shown in Figures 1.4 and 1.29 have a image size of 8192×8192 , due to the large angular extent of the sources.

² Standard ‘Briggs’ weighting characterised by ‘robust’ parameter.

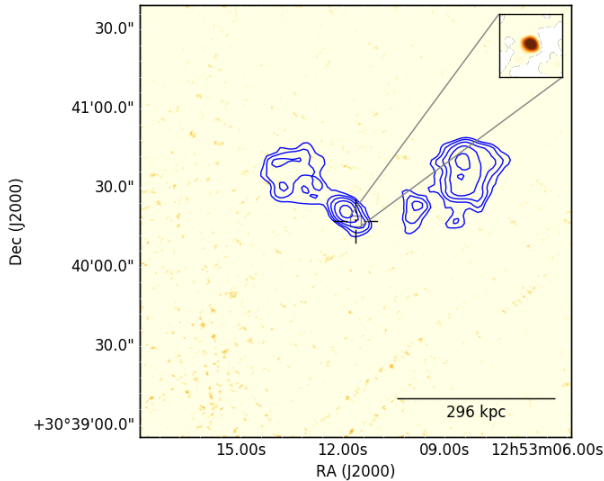
³ Beam size taken as an average from the 38 scans of the target sources.



(1.1) J125143.00+332020.2



(1.2) J125147.02+314046.0



(1.3) J125311.08+304029.2

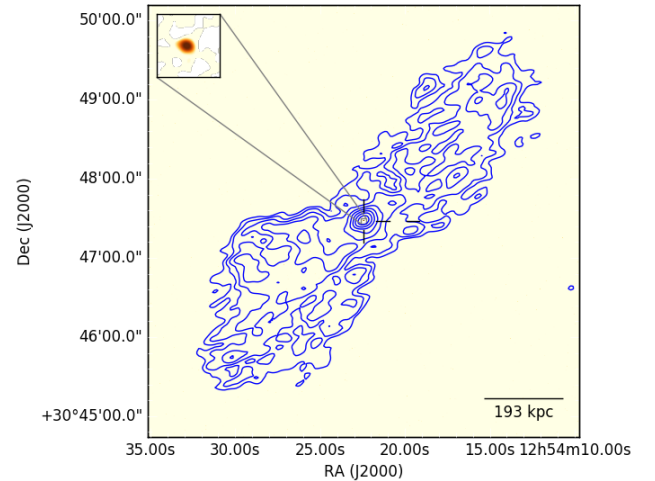
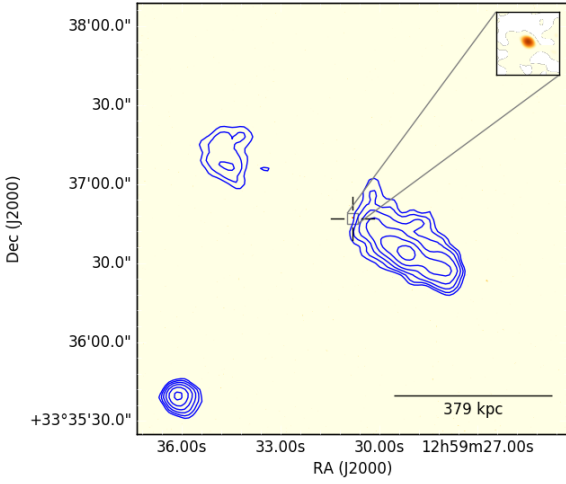
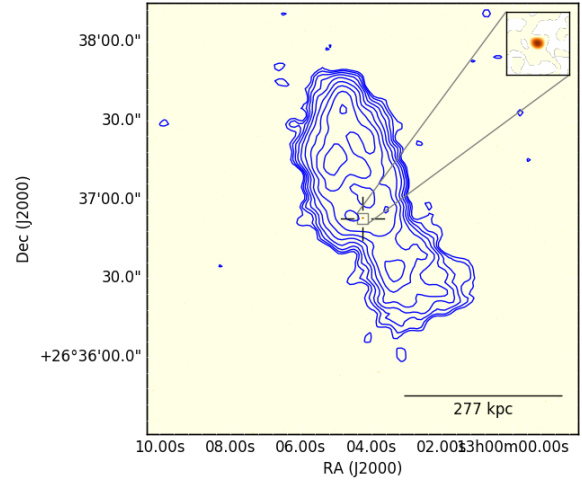
(1.4) J125419.40+304803.0[†]

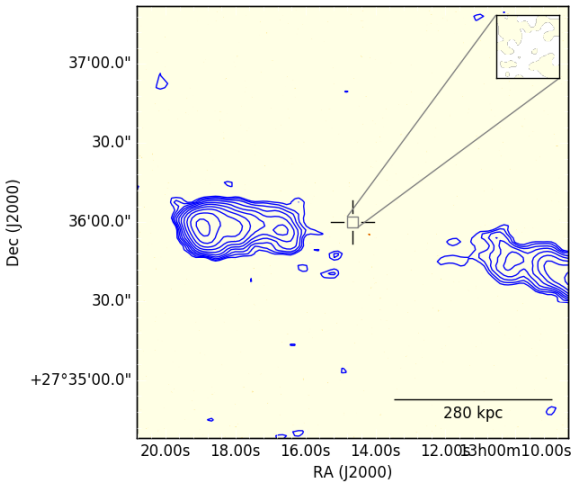
Figure 1. 6-GHz VLA images of the candidate remnant sources (shown as the background), centred on the optical IDs made by H16. A zoom-in of the VLA image at the location of the optical ID is given on the top right hand side of each image, for a clear visual identification of a core. Images are scaled logarithmically, and smoothed with a Gaussian function with FWHM of 0.2 arcsec (5 pixels). 150-MHz LOFAR contours (blue) are set at various powers of 2 multiplied by 2σ (where σ is defined as the local noise level in the LOFAR image), in order to encapsulate as much low surface brightness emission as possible. Sources that we identify as artefacts and not radio-loud AGN are labelled in their figure caption. [†]Contours based on the pre-processed LOFAR data (see Section 2.2).



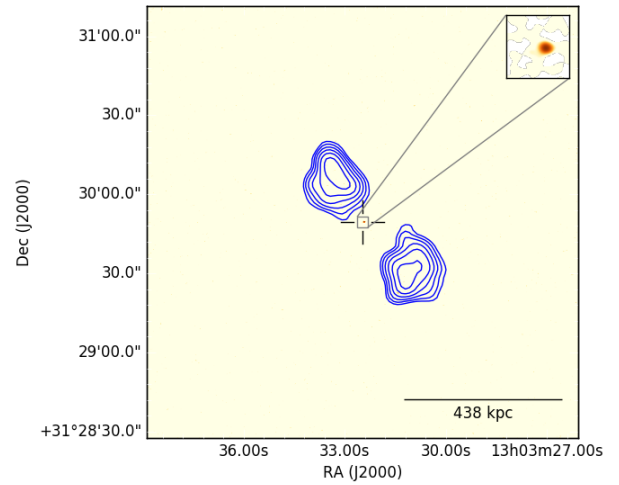
(1.5) J125931.85+333654.2



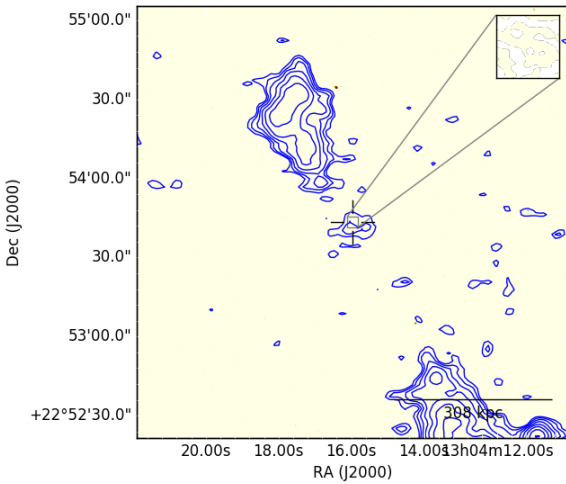
(1.6) J130003.72+263652.1



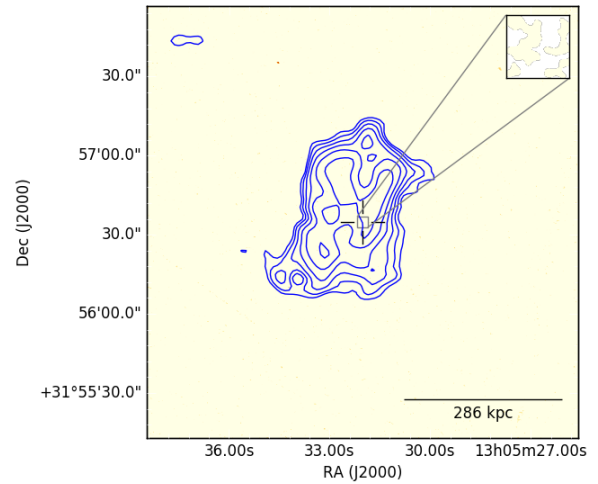
(1.7) J130013.55+273548.7



(1.8) J130332.38+312947.1

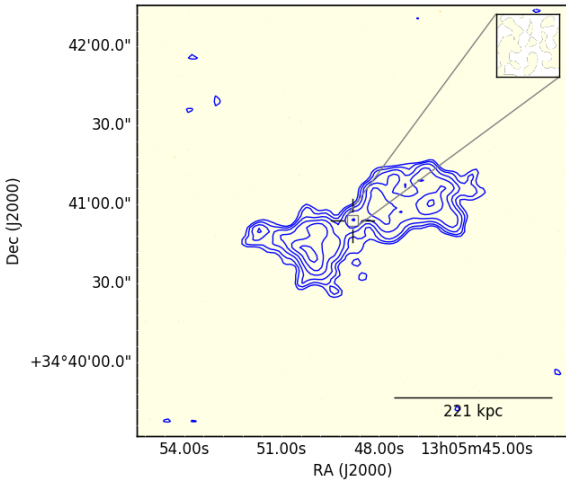


(1.9) J130415.46+225322.6

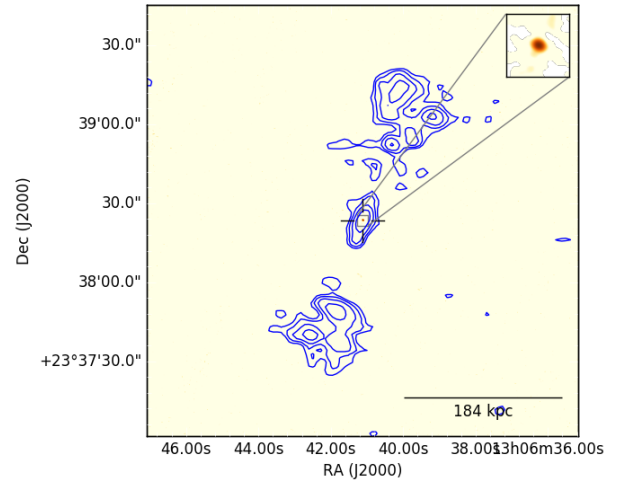


(1.10) J130532.49+315639.0

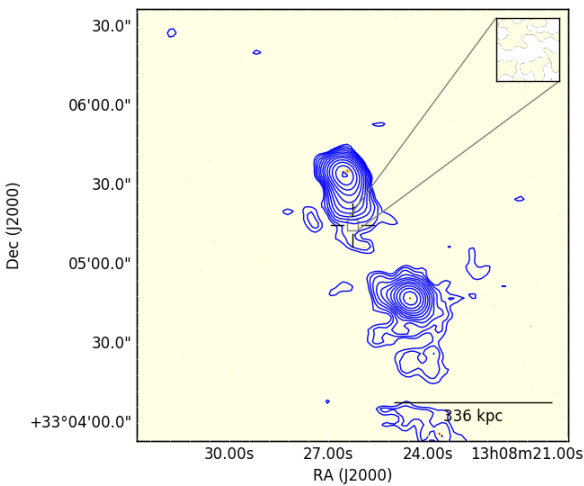
Figure 1. Continued



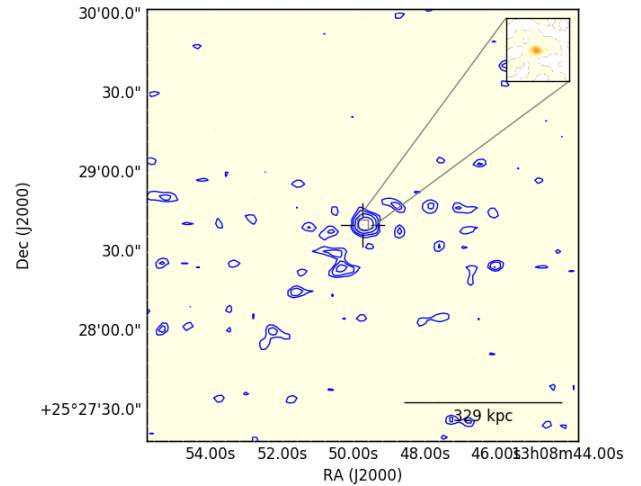
(1.11) J130548.65+344052.7



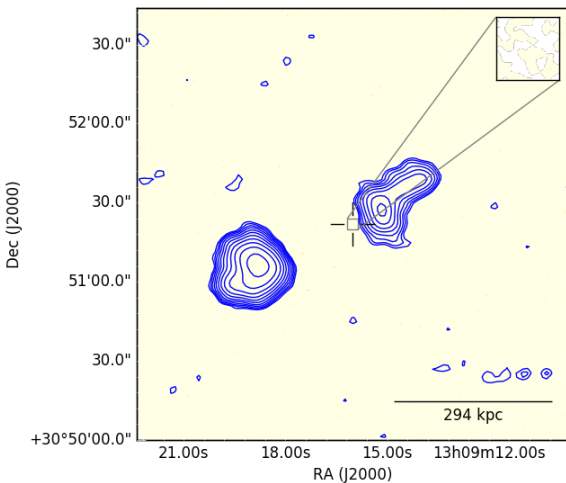
(1.12) J130640.99+233824.8



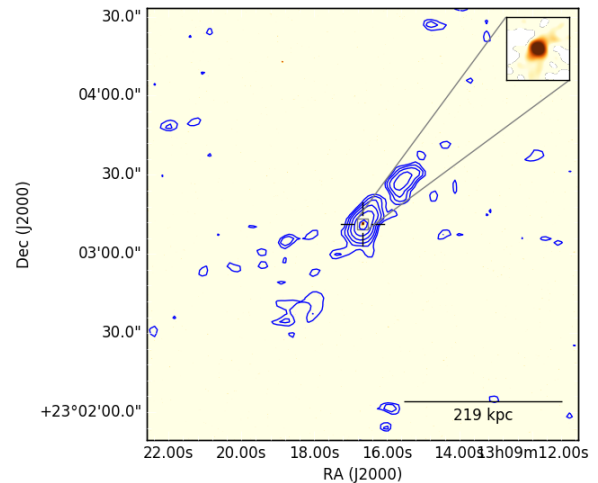
(1.13) J130825.47+330508.2



(1.14) J130849.22+252841.8 (removed from sample)

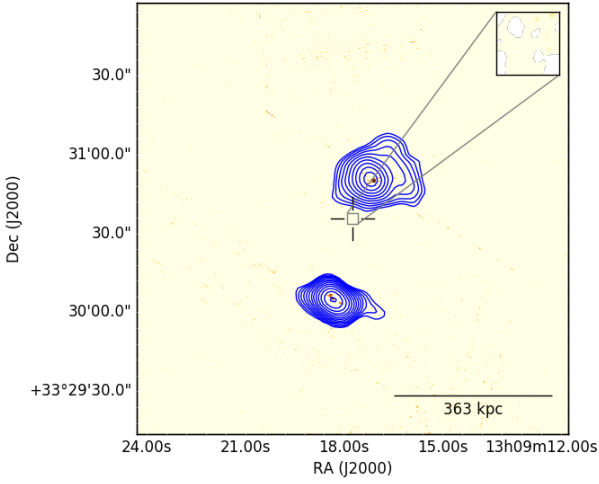


(1.15) J130916.85+305118.3

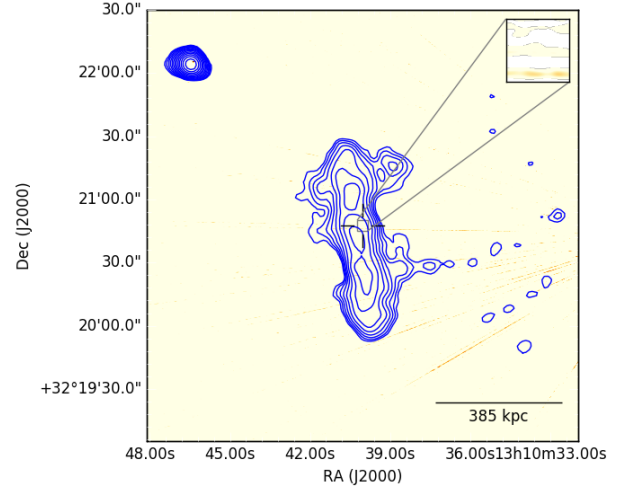


(1.16) J130915.61+230309.5

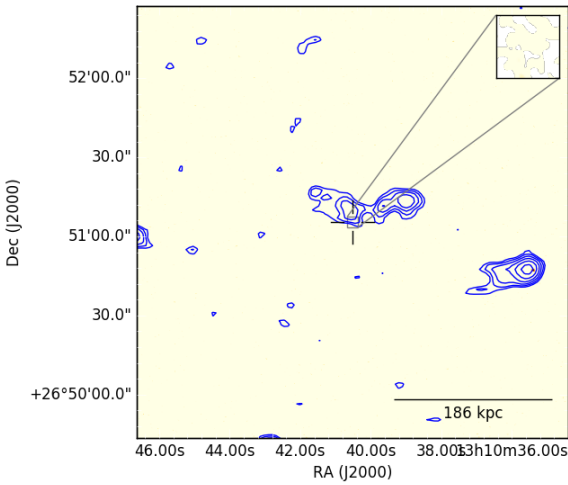
Figure 1. Continued



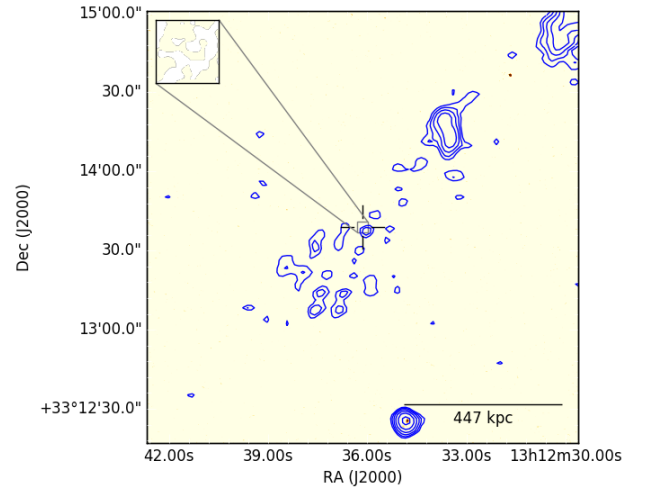
(1.17) J130917.63+333028.4 (removed from sample)



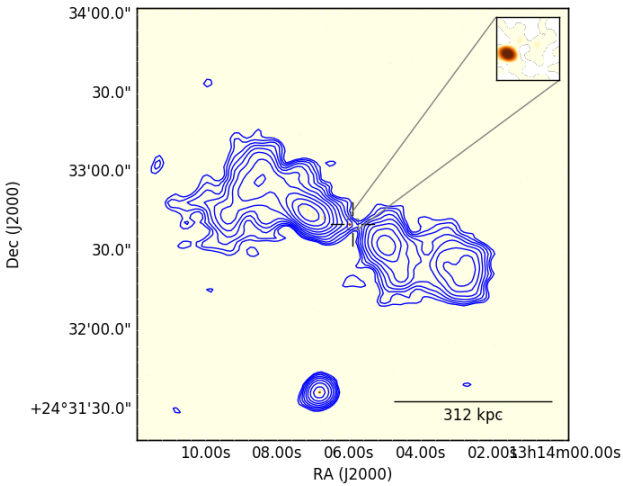
(1.18) J131040.25+322044.1



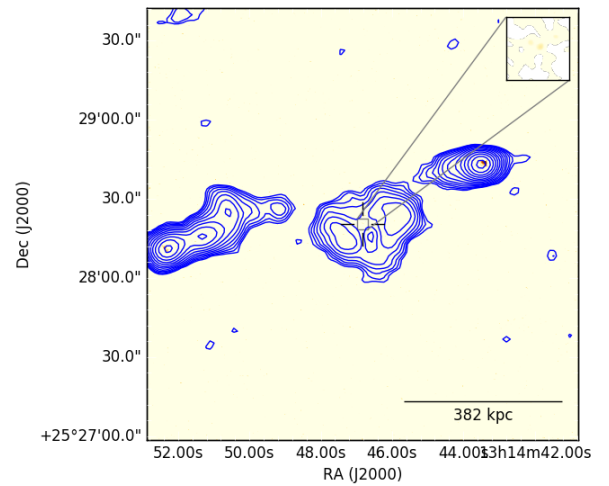
(1.19) J131039.92+265111.9



(1.20) J131235.35+331348.6

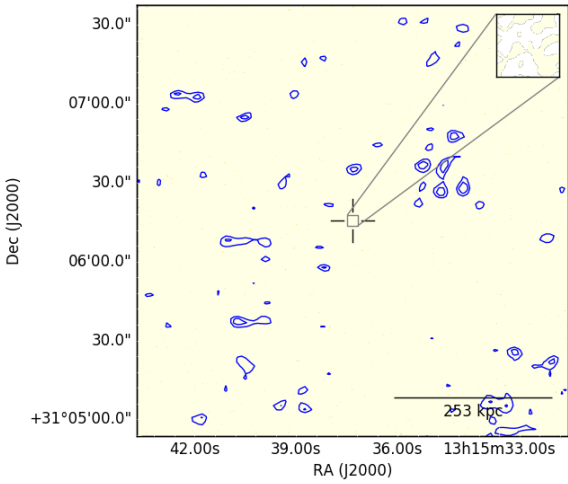


(1.21) J131405.16+243234.1

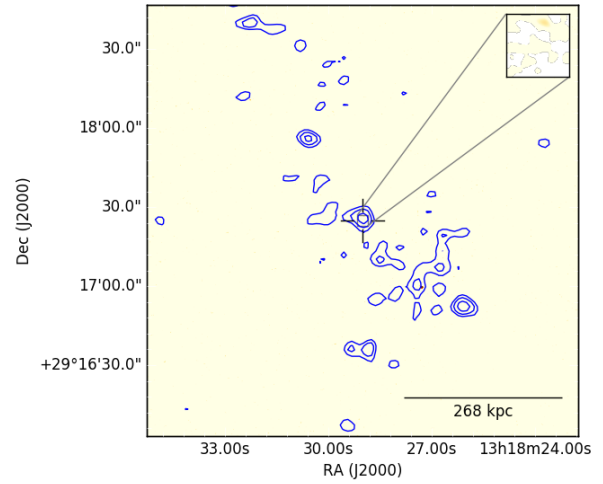


(1.22) J131446.57+252819.8

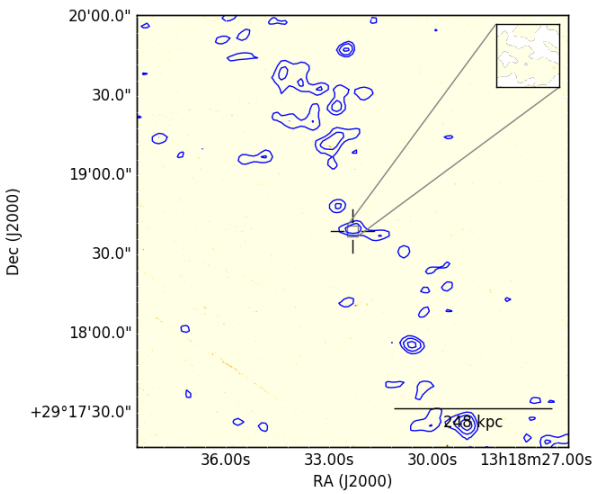
Figure 1. Continued



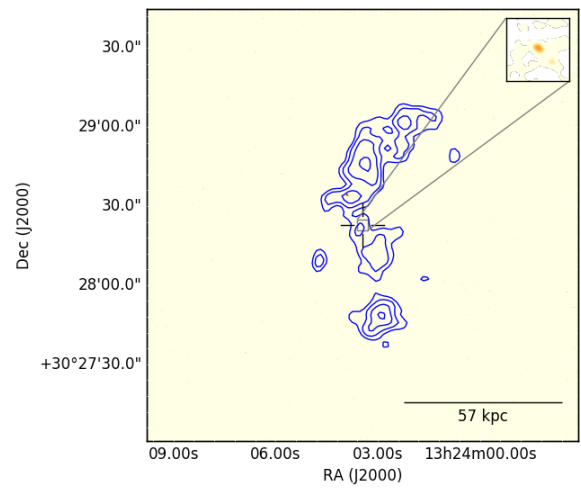
(1.23) J131536.30+310615.5 (removed from sample)



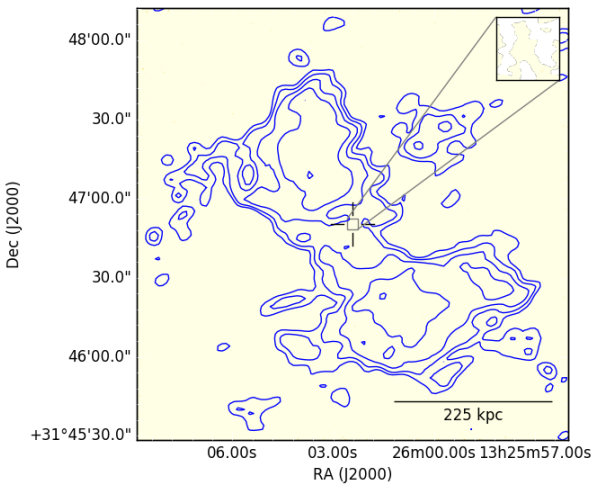
(1.24) J131827.83+291658.5 (removed from sample)



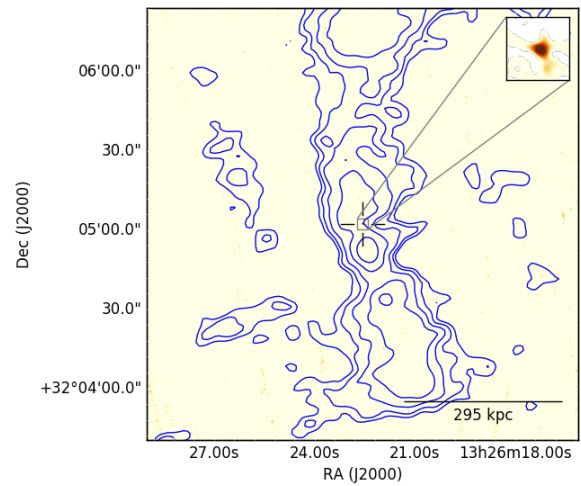
(1.25) J131833.81+291904.9 (removed from sample)



(1.26) J132402.51+302830.1

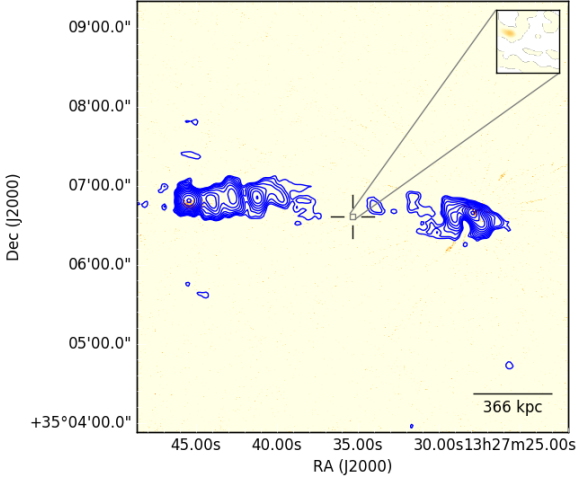


(1.27) J132602.06+314645.6[†]

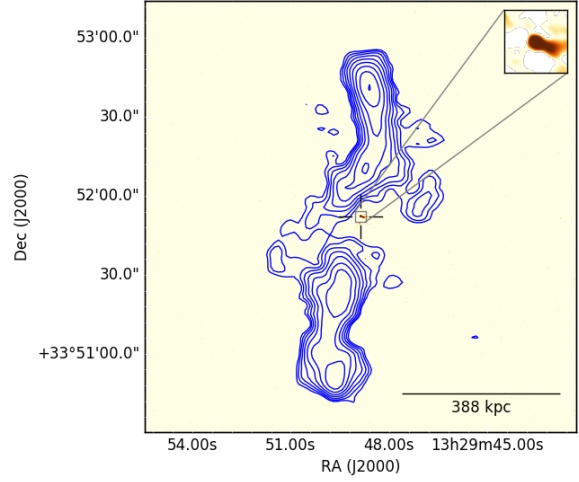


(1.28) J132622.12+320512.1[†]

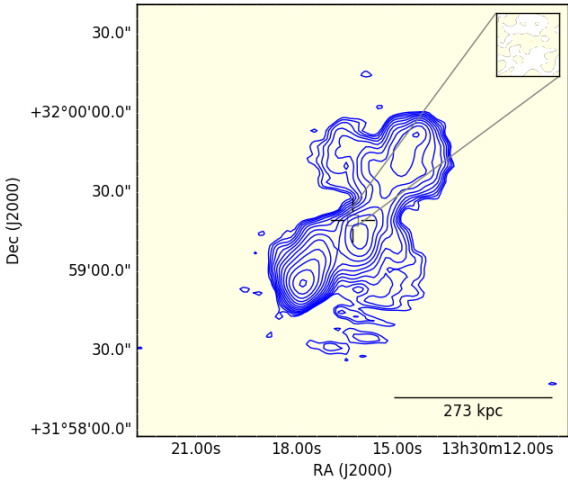
Figure 1. Continued



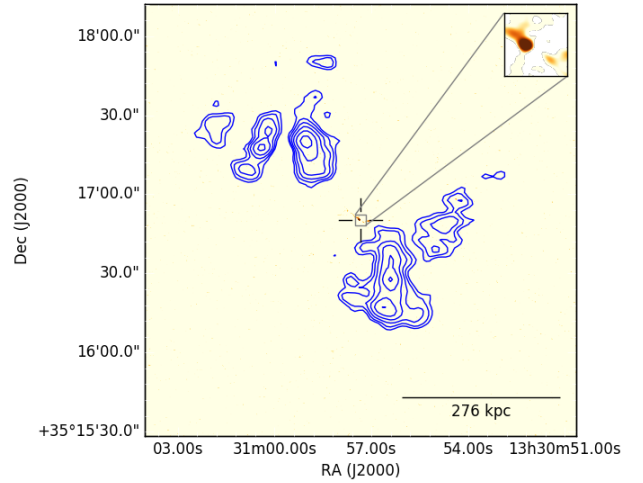
(1.29) J132738.77+350644.6



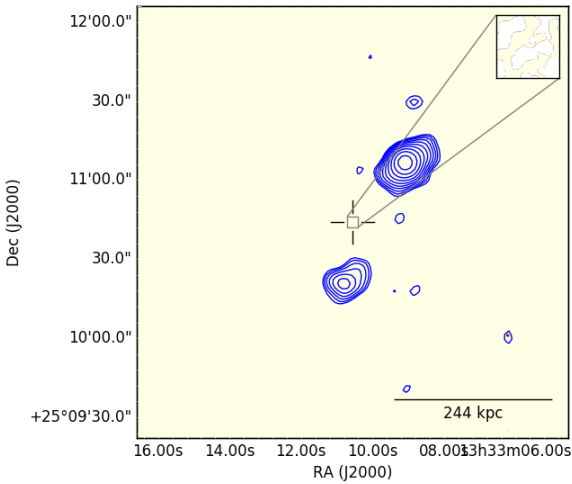
(1.30) J132949.25+335136.2



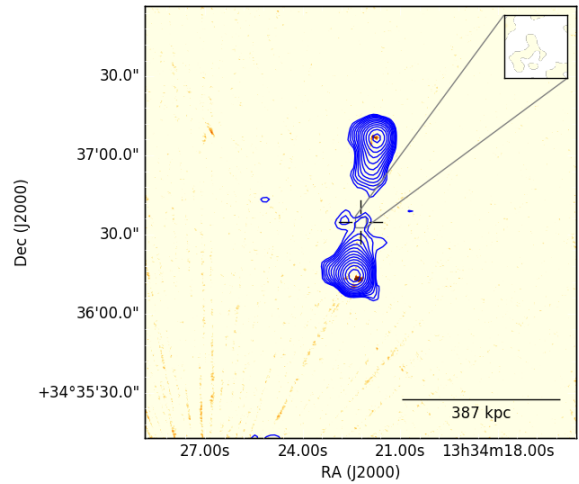
(1.31) J133016.12+315923.9



(1.32) J133058.91+351658.9

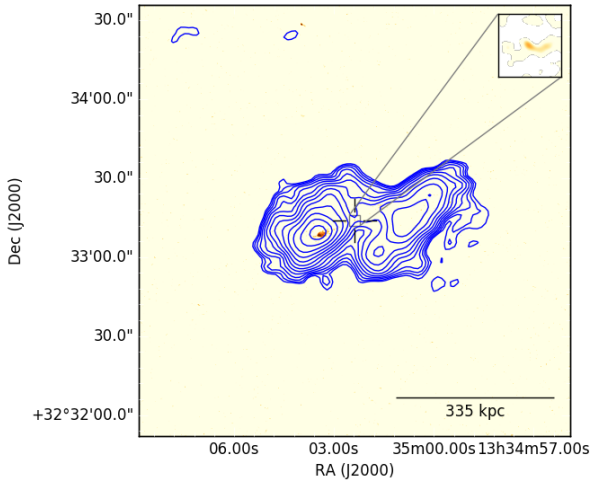


(1.33) J133309.94+251045.2

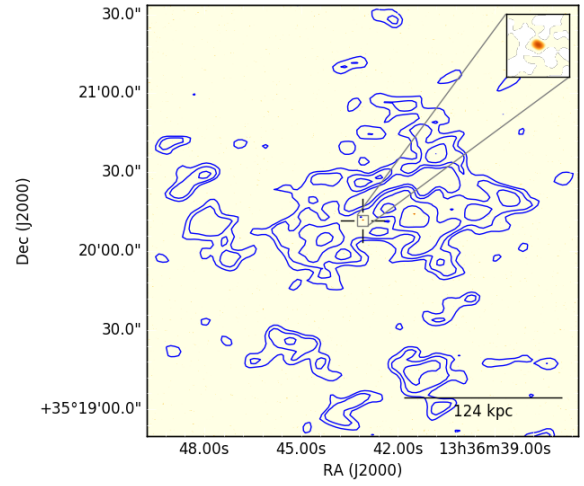


(1.34) J133422.15+343640.5

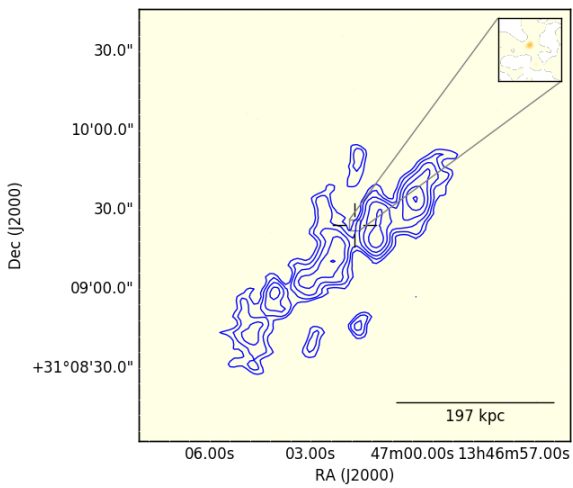
Figure 1. Continued



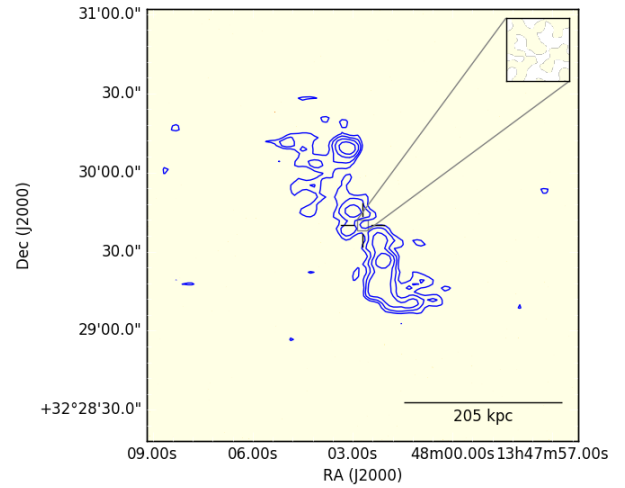
(1.35) J133502.36+323312.8



(1.36) J133642.53+352009.9



(1.37) J134702.03+310913.3



(1.38) J134802.83+322938.3

Figure 1. Continued

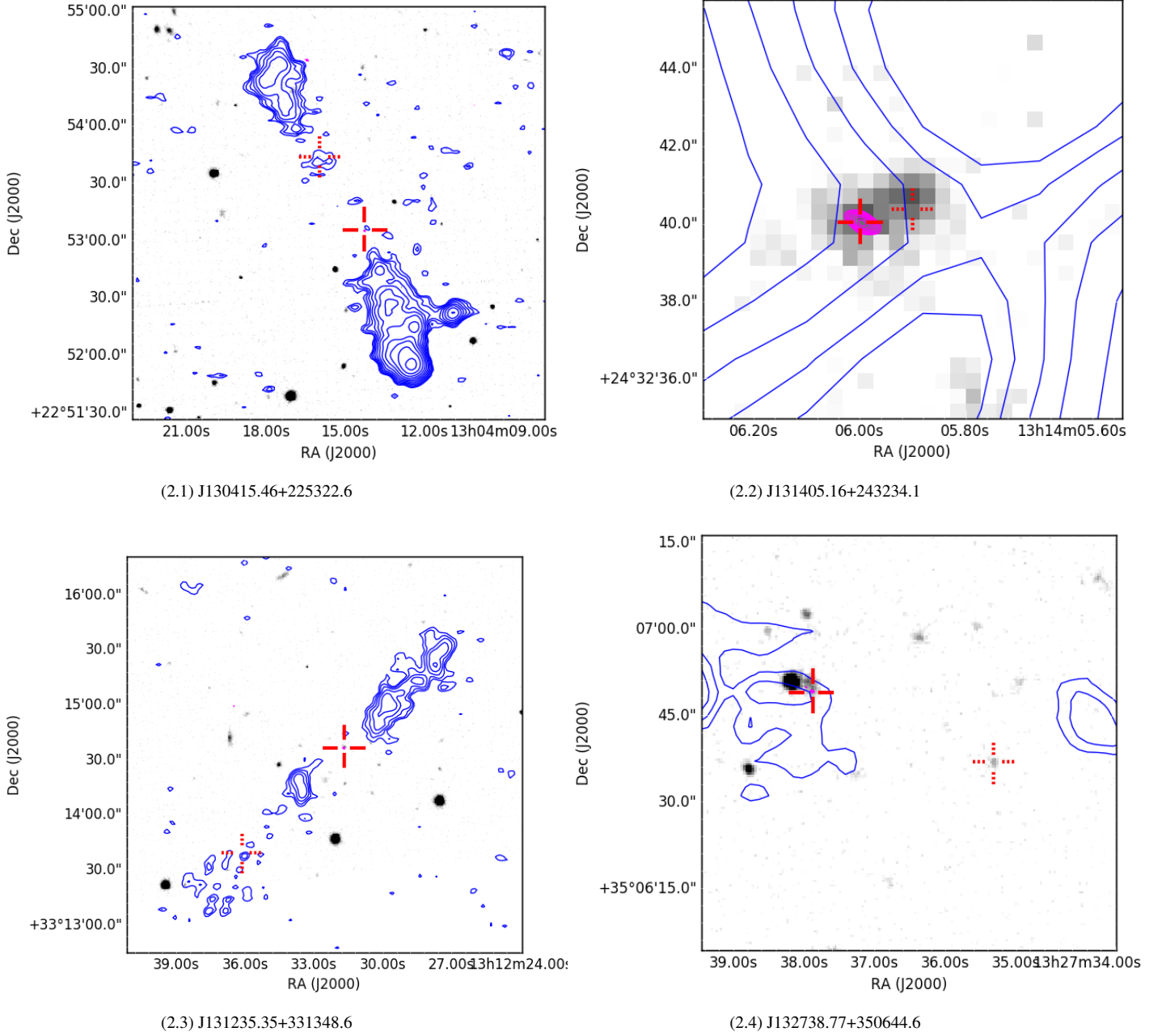


Figure 2. SDSS *r*-band images, shown in grey-scale, of the four new optical identifications made in this sample. Overlaid are LOFAR 150-MHz contours (blue) and VLA 6-GHz contours (magenta). For clarity, the previous optical ID is marked with a red dotted cross-hair, and the new optical ID we make based on the position of the VLA 6 GHz core is marked with a red solid cross-hair.

3 RESULTS

3.1 Remnant statuses

In this section we report on the status of each radio source – either active or remnant. As explained in Section 1, we use the simple definition of a remnant AGN as a radio source, classified as an AGN based on the methods of H16, without a compact radio core and jets, aided by our new 6 GHz VLA observations. Confirmations of optical IDs (previously made by H16) were carried out visually by cross matching our detected cores with catalogues from SDSS Data Release 12 (Alam et al. 2015), in addition to FIRST and NVSS at 1.4 GHz. Where any compact radio core from our new 6 GHz observations coincides with a different optical ID (see Figure 2), we assign the new SDSS optical ID to the radio source and give its IAU name, detailed in the descriptions of each source below and in Table 4.

3.1.1 J125143.00+332020.2: Active

Figure 1.1 shows a clear central core with what seems to be faint jet emission at 6 GHz extending to the north-east, with more extended lower frequency emission at 150 MHz on both sides of the core. A hotspot is also seen near the edge of the western lobe at 6 GHz.

3.1.2 J125147.02+314046.0: Active

This is a clearly active DDRG, as evident in Figure 1.2. The 150 MHz contours show the outer radio lobes indicative of previous AGN activity, while the 6 GHz emission shows the hotspots from the restarted jets as well as a central core.

3.1.3 J125311.08+304029.2: Active

Figure 1.3 shows a clear central core, which coincides with the position of the current optical ID from the SDSS catalogue. Extended structure is also seen with at 150 MHz, although relatively faint, and extending in different axes relative to each other. It is possible that this source is in a group environment, based on the large number of SDSS sources surrounding the radio source, and that the jets are bent due to ram pressure stripping in a rich environment.

3.1.4 J125419.40+304803.0: Active

Figure 1.4 shows a clear central core, with diffuse emission at 150 MHz surrounding the compact source. The extended 150 MHz emission, on a much larger scale, shows an FR-I type morphology, although relatively relaxed in shape. With the number of optical sources surrounding the radio source based on the SDSS image, it is possible therefore that this source is in a dense environment – typically where FR-I sources are found at low redshift (Worrall & Birkinshaw 1994).

3.1.5 J125931.85+333654.2: Active

A faint core detected at 3σ seen in Figure 1.5 indicates that the source is still active. The extended lobe emission as shown by the 150 MHz contours is much fainter for the eastern lobe than the western lobe, possibly indicating a large orientation angle of the radio source with respect to the plane of the sky.

3.1.6 J130003.72+263652.1: Active

A clear central core at 6 GHz is seen in Figure 1.6, with large extended emission at 150 MHz extending to the north and fainter emission in the south. The nature of the source (FR-I or FR-II) remains unclear due to the lack of high resolution observations at 1.4 GHz. The morphology is uncharacteristic of classical FR-I or FR-II sources, and this may suggest that this is a newly active source, with the 150 MHz contours in this image describing the ageing remnant plasma from previous activity.

3.1.7 J130013.55+273548.7: Active (Misidentified)

Figure 1.7 does not show any compact emission at 6 GHz at the position of the optical ID. However, there is a compact source ~ 8 arcsec (37 kpc) to the south-west of the optical ID, which is more in line with the jet axis based on the positions of the radio lobes seen at 150 MHz. Assuming that this is the core, the optical ID for this

source made by H16 is incorrect. However, there is no currently detected optical ID at the position of the core in SDSS, possibly because the host galaxy is at a high redshift and/or dust obscured.

3.1.8 J130332.38+312947.1: Active

Figure 1.8 shows a clear compact core at the position of the optical ID. The extended lobe structure at 150 MHz is that of a classical FR-II radio galaxy, and we therefore confirm this source as being active.

3.1.9 J130415.46+225322.6: Active (Misidentified)

Figure 1.9 does not show any clear compact source at 6 GHz close to the optical ID. There are however, two sources of compact emission in the image – adjacent to the northern lobe and a faint source ~ 15 arcsec above the southern lobe on the jet axis. It is likely that the latter is the radio core of this system and in this scenario therefore, this system is active. The proximity of this core to the southern lobe, relative to the northern lobe, may be explained if the southern lobe is oriented closer towards the line of sight than the plane of the sky, or if the source is intrinsically asymmetrical. The optical ID at the position of this 6 GHz compact source is J130414.25+225305.0, and we assign this as the new optical ID corresponding to this active radio source.

3.1.10 J130532.49+315639.0: Candidate remnant

Figure 1.10 shows no sign of a central core within the extent of the emission at 150 MHz, which seems to be diffuse and showing no signs of typical lobe structure seen in classical FR-I or FR-II. There are no obvious signs, based on the number of catalogued SDSS sources around the radio source, of a dense group or cluster environment. If this is a true remnant radio galaxy, it is likely that it has had a long remnant phase, based on the morphology shown in Figure 1.10 alone.

3.1.11 J130548.65+344052.7: Candidate remnant

Figure 1.11 shows no clear sign of a compact core, indicating that the core and jets have switched off. The 150 MHz contours show double-lobed structure extending to the east and west, although relaxed in shape which is typically expected of remnants.

3.1.12 J130640.99+233824.8: Active

Figure 1.12 shows a clear compact source at 6 GHz at the position of the optical ID, with extended emission surrounding the core at 150 MHz and extending towards the north and south. The lobe emission seems to show a relaxed morphology, in contrast to the emission immediately surrounding the core. This may indicate a restarted source, although we do not detect hotspots.

3.1.13 J130825.47+330508.2: Active (Misidentified)

Figure 1.13 shows no significant compact emission at the location of the optical ID. Compact hotspots at 6 GHz can be seen however, surrounded by bright double radio lobes at 150 MHz. A detailed inspection of the 6 GHz image reveals a very faint, yet significant (3σ), compact object equidistant from the two lobes and along the jet axis. Furthermore, the object lies within the faint extension of

the base of the northern lobe (see Figure 1.13) towards the southern lobe. This object however, is not currently associated with an optical ID from the SDSS DR12 catalogue. A cross-check with the Two-Micron All Sky Survey (2MASS; Skrutskie et al. 2006) and the Wide-field Infrared Survey Explorer (WISE; Wright et al. 2010) catalogues also shows no significantly detected sources. It is likely therefore that the true optical host is a high redshift galaxy. We therefore dismiss the current optical ID for this active radio source.

3.1.14 J130849.22+252841.8: Active

Figure 1.14 shows a faint 6 GHz core, surrounded by diffuse emission at 150 MHz. The morphology of the emission is relaxed in shape, and is typical of what is expected of aged remnant plasma. Given that a core is detected, however, we confirm that this source is still in its active phase.

3.1.15 J130916.85+305118.3: Candidate remnant

Figure 1.15 shows no apparent core near the optical ID, although the 150 MHz contours show characteristics of a bright, double-lobed radio galaxy. However, inspection of the LOFAR 150 MHz image on a larger scale reveals a further pair of lobes to the eastern side of Figure 1.15 with a similar morphology. It is possible that this source is therefore a DDRG, and the non-detection of a compact core within the jet axis of both pairs of lobes makes it possible that this source is a remnant DDRG. Nevertheless, we confirm the remnant candidate status for this source.

3.1.16 J130915.61+230309.5: Active

Figure 1.16 shows a clear, compact and bright core, surrounded by emission at 150 MHz. There is also an extension towards the northwest, which is likely to be a radio lobe.

3.1.17 J130917.63+333028.4: Removed from sample

Figure 1.17 shows compact bright objects at 6 GHz associated with the lobe structures at 150 MHz, particularly in the northern lobe of this source. However, no compact object is visible in this image near the optical ID or along the assumed jet axis. There is however, an SDSS source at the precise location of the compact object seen at 6 GHz in the northern lobe (J130917.11+333049.8). This however, is classified spectroscopically as a quasi-stellar object (QSO), or quasar. Moreover, the 6 GHz emission associated with the southern lobe has two faint components, neither of which have an SDSS optical ID. It is plausible that the northern object is a distinct radio-loud quasar, while the southern source represents a high-redshift radio galaxy. Given that both sources individually do not meet our original AGN sample criteria, we therefore removed this source from our parent radio-loud AGN sample, as well as from our original candidate remnant sample.

3.1.18 J131040.25+322044.1: Candidate remnant

Figure 1.18 shows no visible compact core at 6 GHz, and the 150 MHz contours show a relaxed morphology, with no clear separation of the radio emission from the lobes and the central host. Given the large physical extent of the source (~ 500 kpc) and its morphology, it is most plausible that this is an extremely old source in a long remnant phase.

3.1.19 J131039.92+265111.9: Candidate remnant

Figure 1.19 shows no evidence of a central compact core at 6 GHz, with the extended emission at 150 MHz likely corresponding to lobed remnant emission. A closer inspection of the 6 GHz map reveals some significant compact emission associated with the radio emission at 150 MHz to the south-west of the target source. Furthermore, there are FIRST and NVSS detections associated with this emission. Since, after inspecting the 150 MHz image, there is also an eastern lobe associated with this background source, this is likely a background radio galaxy with a western hotspot. However, we do not detect any significant compact core emission at 6 GHz along the jet axis of this source, nor do we find FIRST or NVSS core detections. Although this source meets the physical selection criteria of our parent AGN sample, we are unable to constrain its optical ID due to the large number of SDSS-detected sources along the jet axis. We therefore interpret this source as a detection of another remnant radio galaxy, although the lack of an optical ID means that we cannot include it in our sample. Nevertheless, we can confirm the candidate remnant status of the original target source.

3.1.20 J131235.35+331348.6: Active (Misidentified)

Figure 1.20 shows no clear 6 GHz emission from the optical ID, with some faint 150 MHz emission extending towards the north-west and south-east. There is however, a compact source at 6 GHz seen towards the north-west (see Figure 1.20), as well as an extension of the 150 MHz emission towards the north-west that extends beyond the image. There is an optical ID corresponding to this compact source (J131231.68+331436.3), implying that this is an active radio galaxy with this new optical ID. Figure 2.3 displays the position of the new optical ID at the position of this 6 GHz core, with respect to the previous optical ID in Figure 1.20.

3.1.21 J131405.16+243234.1: Active (Misidentified)

Figure 1.21 shows a clear core at 6 GHz, surrounded by extended emission at 150 MHz. This source shows the morphology of a classical double radio galaxy, although the 150 MHz contours visually show some sign of outer radio lobes being part of a previous remnant phase. Faint emission associated with the inner lobes can be seen in an un-smoothed 6 GHz map, while 1.4 GHz FIRST detections are also present at these positions, supporting the double-double nature of this source. A faint extension of the outer eastern lobe towards the south can also be seen. We can also dismiss the current optical ID and confirm a new optical ID for this source as J131405.99+243240.1, cross-matching the position of the 6 GHz core and this new optical ID. Due to the proximity with the previous optical ID, the influence of a potential merger event therefore could have disrupted the previous AGN activity, with the subsequent gas infall due to the merger feeding the nucleus to produce the current activity.

3.1.22 J131446.57+252819.8: Candidate remnant

Figure 1.22 shows no apparent 6 GHz core at the optical ID, although two hotspots in the east and west directions are seen, associated with the outer lobes at 150 MHz. Although a double-double nature for this source is plausible, a closer inspection of the 6 GHz un-smoothed map reveals a very faint compact object in line with the jet axis of the outer lobes. We therefore infer that this image

contains two radio galaxies, with the source that has hotspots being active due to a very compact and faint core detected at 3σ lying on the jet axis, and the other source with the inner lobe structure without a core being a remnant. However, there is no currently detected optical ID at the position of the 6 GHz compact object. Given the 3σ detection along the jet axis, and the fact that there is a small extension of the eastern lobe towards the proposed radio core, it is still likely that this faint object corresponds to the core of the active source. Since we cannot optically identify this new source, we do not include this in our sample. Nevertheless, we confirm the remnant status of the original source (J131446.83+252), since it still meets the original AGN sample criteria as detailed in Section 1.

3.1.23 J131536.30+310615.5: Removed from sample

Figure 1.23 shows no compact 6 GHz core or significant emission at 150 MHz resembling a radio galaxy. A lower resolution LOFAR map made by H16 shows some faint diffuse emission that may resemble a radio galaxy, although it is probable that this is an imaging artefact. Given that the 2σ contours of the higher resolution map shown in Figure 1.23 do not bear any resemblance to typical jet or lobe structure, we are confident therefore that this is not a genuine AGN, and remove this source from the original parent radio-loud AGN sample.

3.1.24 J131827.83+291658.5: Removed from sample

Figure 1.24 shows no significant 6 GHz emission corresponding to the optical ID, although there is a faint object towards the south west from the centre of the image (inspecting an un-smoothed 6 GHz image). This compact source is in fact at the position of an SDSS source – namely J131827.29+291659.3. However, this object is currently identified by SDSS as a star and the lack of spectroscopic information available for this object means that it cannot robustly be classified as a radio-loud quasar. Given the brightness of the compact source at 6 GHz (~ 100 mJy) and the relative weakness of 150 MHz emission at this position, it is likely that this radio emission is associated with a flat-spectrum radio-loud quasar. Nevertheless, we find that the 150 MHz emission seen in Figure 1.24 does not satisfy our flux criterion; the source was included in the sample based on a lower resolution LOFAR image. We therefore remove this source from our sample, as explained in Section 2.2.

3.1.25 J131833.81+291904.9: Removed from sample

This source contains no significant compact emission near the optical ID at 6 GHz, as seen in Figure 1.25. There is some faint 6 GHz emission associated with the extended structure at 150 MHz, however this is likely to be due to artefacts in the image from bright sources nearby. A closer inspection reveals that the 150 MHz emission is in fact an extension of J131828.97+291 (Figure 1.24), and it is possible that two separate sources were seen by the lower resolution images made by H16. Nevertheless, similarly to J131828.97+291, this source does not satisfy our flux criterion and we therefore remove it from our sample, as explained in Section 2.2.

3.1.26 J132402.51+302830.1: Active

Figure 1.26 shows a very faint and compact source at 6 GHz at the location of the optical ID, detected at the 3σ level. This source is

surrounded by extended and diffusive emission at 150 MHz, showing a morphology similar to a classical FR-I radio galaxy – consistent with the low-redshift of the object (Table 1) and the large elliptical shape of the host galaxy as seen in the SDSS image. There is another compact region slightly north of the core that also has a significant detection with FIRST, but no corresponding optical ID and is therefore likely to be emission from the same source. Furthermore, the northern jet at 150 MHz seems to change direction and tail towards the west at the position of this compact object, suggesting that it is indeed associated with this radio galaxy.

3.1.27 J132602.06+314645.6: Candidate remnant

Figure 1.27 shows no signs of a compact core in the vicinity of the optical ID, and the 150 MHz emission shows a relaxed, yet double-lobed morphology. We therefore confirm the remnant candidate status of this source.

3.1.28 J132622.12+320512.1: Active

Figure 1.28 shows a clear compact core at 6 GHz centred on the optical ID. The 150 MHz extended emission displays a morphology typical of a classical FR-I radio galaxy. We therefore have robust evidence of the active nature of this source.

3.1.29 J132738.77+350644.6: Active

Figure 1.29 shows this source has a classical FR-II morphology, with a large extension at 150 MHz and hotspot emission at 6 GHz (obscured by the contours in this image). Faint 6 GHz emission can be seen near the optical ID, as seen in the zoom-in – however, a positional cross-check with SDSS shows no sign of an optical host at this position. A closer inspection of the 6 GHz image reveals a brighter compact object along the jet axis towards the east of the current optical ID (unseen in this smoothed map due to obscuration by the 150 MHz contours). Cross-matching its position with SDSS reveals that there is in fact an optical ID associated with this object. Inspecting a colour-scale map of the SDSS image, the morphology of the host suggests that this may be a binary system, although this has not been confirmed by DR12. Nonetheless, we dismiss the current optical ID, and associate the radio source with the SDSS galaxy J132737.92+350650.2.

3.1.30 J132949.25+335136.2: Active

Figure 1.30 shows very clear signs of a prototypical FR-II radio galaxy – a bright central core at 6 GHz and elongated bright lobes at 150 MHz. The radio core is at the location of the optical ID, while there are also 1.4 GHz FIRST and NVSS detections in the lobes. Given its morphology, FIRST and VLA core detections, 150 MHz flux density and its relatively high spectroscopic redshift of 0.5601 compared to the rest of the sources in our sample, this active radio galaxy represents one of the most powerful sources in our core-detected sample.

3.1.31 J133016.12+315923.9: Candidate remnant

Figure 1.31 shows the lack of a significant detection of a compact source at 6 GHz at the optical ID, although the 150 MHz map shows

a classical double lobe structure. The northern lobe shows an extension to the east, and the southern lobe shows an extension to the west.

3.1.32 J133058.91+351658.9: Active

Figure 1.32 shows a clear, bright, compact core with faint collimated extensions at 6 GHz. This is a known FR-II radio galaxy (Kozieł-Wierzbowska & Stasińska 2011). The lobe emission at 150 MHz seems to be dissipative with various extensions and asymmetric with the presumed jet axis – most notably of the southern lobe which extends directly south and also has a separate extension towards the north west. This extension is possibly a separate source – there is a faint 6-GHz detection within the 150 MHz contours at this position, coincident with an SDSS optical ID (J133054.41+351657.8). The redshift of this galaxy ($z = 0.489$) is higher than that of J133057.34+351 ($z = 0.3158$), although the former is based on photometry and the latter on spectroscopy. It is possible therefore that Figure 1.32 shows the image of two interacting radio galaxies. We suggest that this is a separate radio source previously unidentified by H16, but we find that the new radio source does not satisfy our flux density criteria ($> 80\text{mJy}$), and we therefore do not include this as an additional source in our sample.

3.1.33 J133309.94+251045.2: Candidate remnant

Figure 1.33 shows two sources of radio emission at 150 MHz away from the central optical ID, likely corresponding to the two lobes of a radio galaxy. No 6 GHz emission can be seen in Figure 1.33; however, a closer inspection of the 6 GHz map shows some faint emission coinciding with the edge of the northern lobe at 150 MHz (~ 30 arcsec from the central optical ID), and we interpret this as hotspot emission from the northern lobe. However, we do not see any clear evidence of a radio core at low- or high-frequencies in line with the jet axis, and therefore we can confidently confirm the remnant candidate status of this source, albeit with faint hotspot emission at high frequencies.

3.1.34 J133422.15+343640.5: Candidate remnant

Figure 1.34 shows what seem to be two very bright hotspots at 6 GHz, and associated bright lobes based on the 150 MHz contours. However, no core is apparent. It is possible that the presence of a nearby bright object to the south of this source is hindering the detection of a faint and compact core at the optical ID. After visually inspecting the 6 GHz map in detail, and considering the local RMS level at the optical ID is at the sensitivity limit of the sample ($\sim 10\mu\text{Jy}$), the core is clearly not detected. It is likely therefore that this 6 GHz snapshot image has captured the radio galaxy very soon after the core switched off, and that the hotspots are still detected at 6 GHz due to the last injection by the fading jets.

3.1.35 J133502.36+323312.8: Active

Figure 1.35 shows bright, compact emission at 6 GHz at the location of the optical ID. Furthermore, the brighter emission towards the east agrees with the physical inference of the source being active, with an eastern hotspot. There are also 1.4 GHz FIRST detections at the hotspot region as well as at the western lobe, although the western hotspot is not detected in our 6 GHz map.

3.1.36 *J133642.53+352009.9: Active*

Figure 1.36 shows a clear central compact object at 6 GHz at the location of the optical ID. There is also a 6 GHz detection to the west of the optical ID and well within the central area of the 150 MHz emission. A cross-check with SDSS reveals the latter is a radio-bright ‘star’, and therefore we are confident of the current optical ID. Although this source has a core detection, the extended emission shows a relaxed shape, similar to what would be expected from remnants.

3.1.37 *J134702.03+310913.3: Active*

Figure 1.37 shows a very faint, compact object at the centre of the image, coincident with the optical ID. Fitting an elliptical gaussian to this source, we detect the 6 GHz emission at the $\sim 5\sigma$ level, and confirm a core detection. The LOFAR 150 MHz contours show a relaxed morphology of the plasma (~ 400 kpc in extent), indicative of a remnant. However, since we do have a significant core detection at the centre of the radio source, we confirm its nuclear activity.

3.1.38 *J134802.83+322938.3: Candidate remnant*

Figure 1.38 shows no obvious core, or any other 6 GHz emission in the vicinity of the optical ID. The 150 MHz contours show clear extended structure on either side of the host galaxy, resembling a relaxed morphology and signs of the northern lobe diffusing. We therefore confidently confirm the remnant candidate status for this source.

4 DISCUSSION

In this section we interpret our result on the remnant fraction and its implications for the dynamics of radio-loud AGN and their duty cycles.

4.1 Remnant fraction

The results obtained in this paper, which aimed to develop a systematic survey of radio-loud AGN remnants, have been obtained by follow-up observations of low-frequency, wide-area observations of the H-ATLAS field which initially provided a sample of candidate remnant radio-loud AGN. To first order, these potential remnant sources were identified by selecting those sources without detected cores in FIRST images, giving 38 candidate remnants (remnant fraction of 38/127). The resolution and sensitivity of FIRST made this remnant fraction an upper limit, as discussed in Section 1. Our new VLA observations with sub-arcsec resolution at 6 GHz have enabled cores to be detected in this candidate remnant sample (Figure 1), giving revised constraints on the AGN remnant fraction.

From our 34 remnant candidates (after removing four sources that did not meet our sample criteria – see Section 2.2), we also remove the source J130917.74+333, shown in Figure 1.17 (see Section 3.1.17). Our final sample consists of 33 remnant candidates out of 122 radio-loud AGN detected with LOFAR. From these, we see no significant evidence for a radio core in 11. This puts the candidate remnant fraction at $11/122 = 9$ per cent, representing a significant decrease from the upper limit determined by H16 and by Brienza et al. (2017). The implications of this result regarding

the dynamics of remnant radio galaxies and the AGN duty cycle are given in Section 4.6.

4.2 Spectral indices

The radio spectrum of a radio galaxy gives important information on the radiative age of the electrons responsible for the emission. For more accurate constraints on the AGN duty cycle, it is therefore crucial to understand the evolution of the spectral shape of radio galaxies in their remnant phase. In the absence of nuclear activity and/or hotspots, remnant sources are expected to have low-frequency steep spectra, associated with the energy losses (Pacholczyk 1970). However, Murgia et al. (2011); Godfrey et al. (2017); Brienza et al. (2016, 2017) suggest that remnants as a population may show a wide range of spectral characteristics, including flat low-frequency spectral indices consistent with active sources, while also showing steep spectra at higher frequencies. It is important to learn whether these spectral features are characteristic of a significant fraction of remnants, or simply isolated cases.

A spectral analysis of our remnant sample, comparing the 150 MHz-1.4 GHz integrated spectral indices of the remnants (blue) to that of the active sources (red) in our sample, is shown in Figure 3. Spectral indices are based either on catalogued flux density values or on measurements directly from the LOFAR image at 150 MHz, and the corresponding NVSS images at 1.4 GHz. The 150-MHz catalogue used was that from H16 and a 1.4-GHz source catalogue was generated by running PYBDSF³ (Mohan & Rafferty 2015) on the NVSS image covering the same area as the LOFAR image. LOFAR/NVSS cutout images of each remnant candidate were then inspected visually in tandem; regions were defined for each source when the PYBDSF catalogued sources were judged to not accurately capture the source emission in that band. Typically LOFAR regions were identified for large sources where not all the flux was included in the catalogue, and NVSS regions were identified where the source was blended with a nearby source in the lower resolution NVSS image or was split into several sources in the NVSS catalogue. Total flux densities were extracted directly from the relevant images in the aperture defined by the region, and these aperture flux densities were used in preference to the catalogued values in determining the spectral indices. LOFAR source sizes were determined by the largest extent of the 3σ threshold of each source. We also include in Figure 3 the parent AGN sample that have FIRST-detected cores (black), to compare their spectral index distribution with the active sources in our sample that have faint and compact cores. We note that four sources (two core-detected and two candidate remnants) do not have an NVSS detection, and therefore have upper limits to their spectral indices only.

Based on the data presented in Figure 3, the median spectral indices⁴ are: $\alpha_{150}^{1400} \approx -0.97 \pm 0.03$ for the candidate remnants; $\alpha_{150}^{1400} \approx -0.80 \pm 0.01$ for the sources with VLA-detected cores; $\alpha_{150}^{1400} \approx -0.62 \pm 0.00$ for the sources with FIRST-detected cores, indicating a tendency for remnants, on average, to have a slightly steeper low-frequency spectral index than active sources. Parma et al. (2007) select their sample of remnants using $\alpha_{1400}^{327} < -1.3$ as a compromise to include the largest possible number of remnants while minimising the number of steep-spectrum active sources. However, it is clearly possible, based on the median spectral indices

³ <http://www.astron.nl/citt/pybdsf>

⁴ Errors quoted are the standard errors of the median assuming a normal distribution.

of the remnant candidates in our sample, for many remnants to be missed by these spectral selection methods, as suggested by Murgia et al. (2011); Brienza et al. (2016). Moreover, it is possible for remnant samples to be contaminated with active sources that may possess steep spectra, based on this spectral index criterion alone. Figure 3 shows that the remnants and the active sources selected in this study clearly possess a wide range of spectral indices, and this is consistent with the conclusions of Brienza et al. (2017). We computed Wilcoxon-Mann-Whitney rank tests between the three samples to test whether the remnants have significantly steeper spectral indices than the active sources. We found, at the 95 per cent confidence level, significantly steeper spectral indices for the candidate remnants than the sources with FIRST-detected cores. We also found that the spectral indices of the sources with VLA-detected cores are significantly steeper than the sources with FIRST-detected cores. The general trend that the sources with brighter, FIRST-detected cores tend to have flatter spectral indices than the sources with more compact cores may imply a positive correlation between core brightness and spectral index – sources with brighter cores will tend to have extended emission less dominated by energy losses than those with faint and compact cores. However, we found no significant difference in steepness of spectral index between the candidate remnants and the sources with VLA-detected cores. Larger sample sizes will be needed in order to make these statistics robust – our VLA core-detected sample consists of only 22 sources compared to the FIRST core-detected sample of 122.

As a further check on our spectral index statistics, we find that five sources in our sample (remnant or active) have $\alpha_{1400}^{150} \leq -1.2$, giving an ultra-steep-spectrum fraction of 4.1 per cent (5/122). This is comparable with the results of a LOFAR study of the Lockman Hole field (Mahony et al. 2016), where an ultra-steep-spectrum fraction of 4.9 per cent was obtained. We note however, that this percentage includes all radio sources detected with LOFAR including radio-loud quasars and star-forming galaxies, and therefore acts as only an upper limit to the fraction of radio galaxies with ultra-steep-spectra. Nevertheless, these statistics demonstrate the problems with solely applying the classical ultra-steep spectrum criterion to select remnants, as found by Brienza et al. (2016, 2017), since we identify more than twice the number of remnant candidates (11) using our core-selection method than this spectral index selection method (5). A visual inspection of the LOFAR images of these five steep-spectrum sources shows that they have a range of morphologies, from relaxed to powerful FR-II-type, with only one being a candidate remnant *and* having a relaxed morphology.

4.3 Core prominence

The core prominence, as explained in Section 1, describes the brightness of a radio core relative to its extended emission. Since we have sampled active sources with faint and compact cores, we expect to find systematically lower core prominences than those obtained with FIRST by H16. In Figure 4 we present the distribution of core prominences for our entire sample. To determine the flux density of detected cores in our 6 GHz observations, for each source, we fitted an elliptical Gaussian to a small region in the map immediately surrounding the core. For the total flux density of the extended emission of each source we use the LOFAR 150 MHz flux density determined by H16. For completeness, we also present 3σ upper limits on the core prominences of our 11 remnant candidates, defining upper limits on the core flux density of the remnants based on the local RMS level within a square box 20 pixels in size

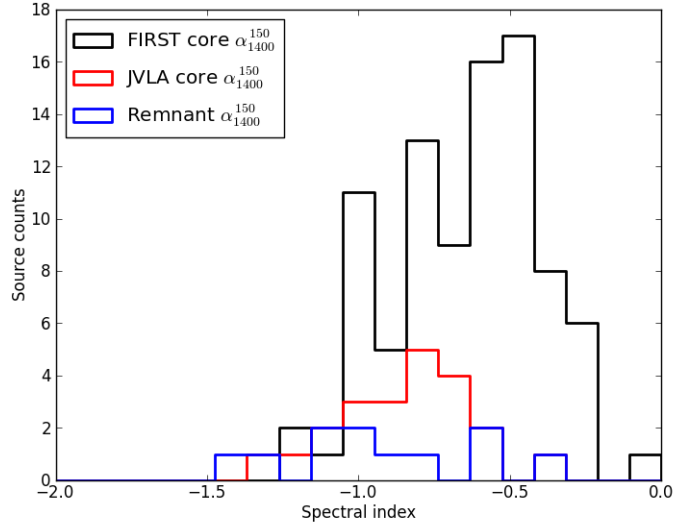


Figure 3. Histogram of 150 MHz-1.4 GHz spectral indices for all sources in our sample colour-coded by core-detections. The candidate remnant with a flat spectral index $\alpha_{1400}^{150} = -0.4$ is only an upper limit due the lack of an NVSS detection at 1.4 GHz.

(dimensions much larger than the VLA beam size) at the location of the optical ID .

Figure 4 shows that, as expected, the core prominences for our compact core-detected sample are significantly lower than the FIRST core-detected sample. The mean core prominence for the faint 6 GHz cores of ~ 0.005 is an order of magnitude lower than the average from the FIRST core detected sample, consistent with the idea that our sensitive, high resolution VLA observations have identified many active sources with faint and compact cores that would be missed by instruments with lower sensitivity and resolution. Since a few of the core prominences of our active sources are comparable to that of the upper limits on the candidate remnants, our results demonstrate that radio-loud AGN with bright, extended emission with faint cores can be misidentified as remnants using observations at lower resolution such as FIRST, if only core prominence is used as a selection criterion.

We also derived core radio luminosities at 6 GHz, listed in Table 5. In comparison to a sample of bright, nearby early-type galaxies formed by Sadler et al. (1989), who infer that the core radio emission for most of their sample is indeed dominated by an active nucleus, the core powers of our sample are higher, including the upper limits on the candidate remnants. This is unsurprising, since our extended, higher-redshift sources are biased towards higher luminosities than that of Sadler et al. (1989), and that some of our candidate remnants may in fact be active. Radio-loud AGN thus have a broad range of possible core luminosities in their active phase – we cannot select a single value below which only remnant AGN may be found using our current sample.

4.4 Power-linear size

Figure 5 shows a power-linear size diagram for the candidate remnant sample (blue), 6 GHz core-detected sample (red) and the sources from the parent AGN sample with FIRST-detected cores (black). To first order, we see that our candidate remnants do not occupy a particularly special position in power-size space with re-

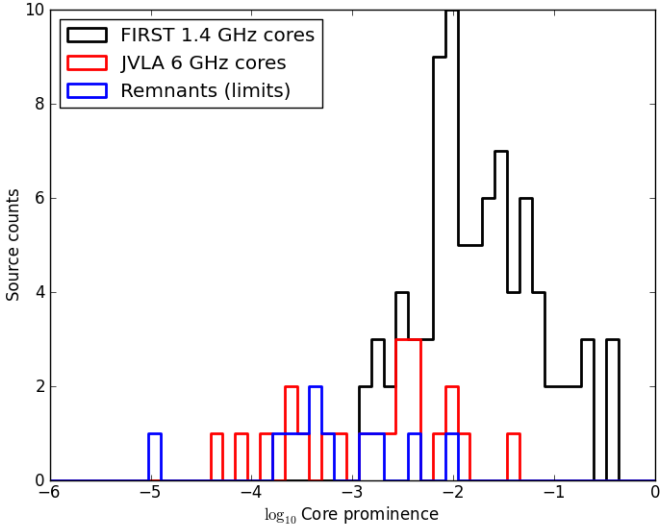


Figure 4. Core prominences for FIRST-detected sources, VLA-detected sources and sources with non-detected cores (remnants), plotted in logarithmic scale. VLA core fluxes were obtained by fitting an elliptical gaussian to the corresponding 6 GHz radio maps. For non-detections (remnants) we use 3σ upper limits on the flux at the positions of their current optical ID.

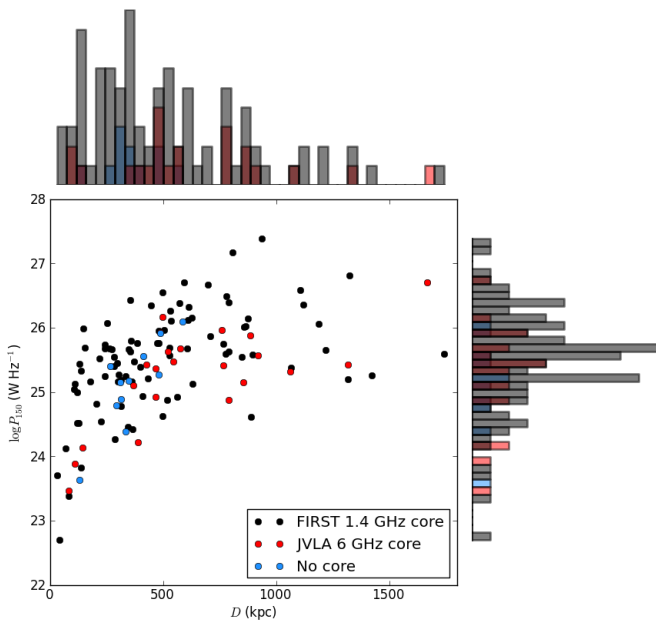


Figure 5. Power-linear size diagram for all sources in this sample. The radio power P_{150} is defined as the 150-MHz luminosity and the linear size D defined as the largest physical extent of the radio source at 150 MHz. We also collapse the distribution of our sources to 1-dimensional histograms, for a clearer view of the distribution in radio power (vertical axis) and linear size (horizontal axis). Note that we have corrected the radio powers and sizes of the four optically misidentified sources (Table 4).

spect to the core-detected sources, as also seen by H16. Performing a two-sample Kolmogorov-Smirnov (K-S) test, we find, at a confidence level of 95 per cent, that there is no statistically significant difference in the distributions of radio luminosities and linear sizes between the candidate remnants and the core-detected sources. We do however, find a significant difference in the physical sizes of the candidate remnants and of the active sources with 6 GHz detected cores – the sources with 6 GHz cores tend to be larger than the candidate remnants for a given radio power. This is surprising, since we might expect remnant sources to be older, and therefore to have expanded to larger distances than active sources, on average. Our analysis, on the other hand, suggests otherwise. This implies that remnants do indeed fade very rapidly, since the intrinsically larger remnants that have experienced more energy losses would have escaped the flux criterion imposed on our sample, and some may have also escaped detection by LOFAR altogether. This would also explain the lack of significantly steeper spectral indices of the remnants compared to the sources with VLA-detected cores (Section 4.2), since the steepest-spectrum sources have escaped detection. Our results can be explained if the remnants in our sample are relatively young compared to the rest of the population (captured soon after switch-off), and therefore generally have physical sizes and luminosities comparable to those of active sources. This is consistent with the conclusions of Godfrey et al. (2017) and of Brienza et al. (2017) – that the core prominence selection method would select young remnants that would be otherwise missed by spectral selection methods. In this scenario, we would expect to observe older and therefore larger remnant sources, that would dominate samples selected by spectral index, in more sensitive surveys.

The radio power and linear size of a particular source are known to be degenerate with other physical parameters, such as radio morphology and environment (Hardcastle & Krause 2013). For example, it can be difficult observationally to distinguish between radio sources of a given power that are young and compact, and aged sources in rich environments whose expansion has been ‘confined’ by a dense external medium (An & Baan 2012; Murgia et al. 2011). This may also explain the lack of a substantial difference in power-size between remnants and active sources in Figure 5. However, as can be shown by the histograms in Figure 5, as well as the results of the K-S test, we see that all classes of sources in our sample are generally clustered around the same values of radio power and physical size. This might suggest that our flux-limited remnant sample may be affected by a degeneracy between radio power/size and environment, although we cannot test the environmental dependency with existing observations. Future observations with increased sensitivity, followed up by optical and/or X-ray surveys to determine the nature of their environment, may be useful in determining a clearer evolutionary path for remnants.

4.5 Optical misidentification

As a sub-product of our VLA 6 GHz observations, we were able to constrain and confirm the SDSS optical identifications of the sources in this sample previously made by H16. By detecting compact cores at high spatial resolution we have been able cross-match their spatial position with the position of the current optical ID, confirming their ID while also making new optical IDs for sources with cores that correspond to different optical sources.

We have confirmed that sources J130415.97+225, J131405.90+243, J130917.74+333 and J132735.32+350 were misidentified by H16, and confirm their new optical IDs based on SDSS DR12, detailed in Table 4. This misidentification fraction

of ~ 10 per cent is high, but it should be borne in mind that these large sources without any core detection in FIRST will be the most challenging of all to optically identify in LOFAR surveys; most sources are unresolved or of small angular size and for these the optical identifications are robust. Nevertheless our work gives an idea of the magnitude of the problem for the largest sources, in the absence of high resolution observations that are sensitive to compact cores. Future wide-area surveys such as LoTSS will need to consider this limit when optically identifying radio-loud AGN that have low core prominences.

4.6 Dynamics of remnant sources

The relatively small upper limit to the remnant fraction (9 per cent) imposed by our observations has strong implications for the understanding of the AGN duty cycle and the dynamics of radio-loud AGN. Previous work on understanding the radio-loud AGN duty cycle based on radiative modelling has suggested, in some cases, that the remnant phase may be comparable or even longer than the active phase (e.g. Parma et al. 2007; Brienza et al. 2016), meaning that remnants should be at least as abundant as active sources using sensitive instruments such as LOFAR. On the other hand, our remnant fraction requires a very rapid fading process. A robust calculation of the duty cycle, of course, would involve determining the spectral ages of all the remnant sources in our sample which are expected to vary between $5 \sim 100$ Myrs, an accurate analysis of which would require using modern, spatially resolved techniques that require deeper observations at a broad range of frequencies with comparable resolution (Harwood et al. 2013) – Harwood (2017) find that spectral ages derived from traditional integrated model fitting techniques give poor descriptions of a source’s spectrum. Recent radiative age mapping of the remnant B2 0924+30 using spatially resolved techniques have shown that its active phase is around a factor of two longer than its remnant phase (Shulevski et al. 2017), agreeing with the general results of this paper. However, Brienza et al. (2017) suggest that radiative ages of remnants would still only yield upper limits, given that the dynamical evolution of radio galaxies becomes an important factor in the remnant phase (also implied by the results of this paper) which is currently missing from models of spectral ageing. We suggest that adiabatic expansion losses must be incorporated into radiative models of radio source evolution, which may lead to more accurate constraints of the AGN duty cycle. Nevertheless, our results imply that the remnant phase of radio-loud AGN is short, relative to the active phase. Therefore, the general picture we can describe for radio-loud AGN is that, once the jets switch off, the radio lobes are visible for a relatively short amount of time – where *both* radiative and adiabatic losses contribute to the shortening of the radiative lifetime. As the lobe plasma expands, its morphology may become relaxed, as seen by the LOFAR contours for most of the candidate remnants in Figure 1. Physical processes within the host galaxy cause the jets to restart in a new active phase, quickly merging with the pre-existing remnant plasma.

Remnant radio-loud AGN hosted by field galaxies would be expected to fade very quickly, whereas in richer and denser environments the lobes are likely to retain the pressure balance with their environment over a longer period, preventing rapid adiabatic losses (Murgia et al. 2011). Given that the physical sizes of the remnants in our sample are similar to those of active sources (Figure 5), it is plausible that we are only able to detect, even with the sensitivity of LOFAR, the youngest and brightest remnants which may tend to reside in dense environments. The effect of an envi-

ronmental dependency on the analysis of our candidate remnants remains unclear without further observations. Nevertheless our results give a clear diagnosis for the lack of a systematic sample of remnant sources until now, in that the sensitivity of instruments and dynamic range are crucial factors in the observations of remnant radio-loud AGN.

In our sample, we have clearly identified a candidate remnant without a core possessing hotspots at 6 GHz – Figure 1.34 (J133422.21+343), similar to the candidate remnant 3C28 (Harwood 2017). This is unsurprising as hotspots are expected to be bright for as long as the light travel time of the jet once the source switches off, and supports the idea, presented in Section 4.4, that the sensitivity limit imposed by current instruments means that we are only able to capture remnants soon after switch-off. We suggest that bright hotspots can be observed for the youngest remnants, although we reiterate that J133422.21+343 is still only a candidate remnant and may still have a compact core that is undetected by the VLA.

Table 5 lists the source properties of the 33 active and candidate remnant sources in our sample, including their core flux density, core luminosity and spectral index of the low-frequency extended emission.

5 CONCLUSIONS AND FUTURE WORK

In this paper we have provided a systematic survey of candidate remnant radio-loud AGN using the unprecedented capabilities of LOFAR and the VLA. Using the detection of compact cores, we are able to constrain the remnant fraction in the H-ATLAS NGP field obtained by H16. In summary we conclude that:

- The small candidate remnant fraction of radio-loud AGN (9 per cent) implies a relatively short fading time of the plasma left behind in the post-switch off phase, before the jets switch on again.
- Our implied short fading time is in contrast to spectral modelling using continuous injection models, which derive longer remnant lifetimes than active lifetimes for a selection of dying sources. This may suggest that these ageing models must incorporate adiabatic losses (as suggested by Harwood 2017) for a more realistic description of remnants, confirming the results of the simulations made by Brienza et al. (2017).
- Unbiased remnant selection strategies in future surveys (e.g. the LOFAR 150-MHz Tier 1 Survey) should bear in mind the uncertainties surrounding the spectral indices of the remnant population (i.e. recently switched off sources may still show flat spectral indices), while allowing the lack of a core detection as a strong proxy for remnant radio-loud AGN.
- We suggest that our remnant sample may only contain the youngest of the population, as also implied by the results of Godfrey et al. (2017) and Brienza et al. (2017). We speculate that some sources may reside in dense environments that prevent strong expansion losses, causing them to remain detectable with LOFAR.

Future work may involve follow up observations of individual remnants from our sample to derive their spectral ages, in order to constrain the duty cycle for AGN. Environmental information is also desirable, with follow-up deep X-ray and/or optical observations potentially providing key constraints on the effect of environment on our sample.

Reference	SDSS Obj ID	SDSS name	RA (deg)	dec (deg)	r-band mag	Redshift <i>z</i>
Figure 1.9	1237667736123933636	SDSSJ130414.25+225305.0	196.059	22.884	21.94	0.795 (p)
Figure 1.20	1237665023836160370	SDSSJ131231.68+331436.3	198.132	33.243	20.52	0.485 (s)
Figure 1.21	1237667911672333044	SDSSJ131405.99+243240.1	198.524	24.544	20.67	0.516 (s)
Figure 1.29	1237664671644516542	SDSSJ132737.92+350650.2	201.908	35.114	20.88	0.500 (p)

Table 4. Table showing the properties and names of the new optical identifications made of four core-detected sources in our sample. We list only the properties of the new optical IDs, referencing the figures from Figure 1 in column 1, and for details of the previous optical ID we refer the reader to Table 1 and Section 3.1. Redshifts given are either photometric (p) or spectroscopic (s).

LOFAR name	SDSS name	6 GHz core flux density (μ Jy)	Flux density error (μ Jy)	log 6 GHz core power (WHz^{-1})	α_{150}^{1400}	log 150 MHz total power (W Hz^{-1})	Linear size (kpc)
J125143.00+332020.2	SDSSJ125142.91+332020.6	442.82	21.083	23.292	-0.9994	25.571	921.01
J125147.02+314046.0	SDSSJ125147.03+314047.6	462.31	39.227	23.237	-0.7919	25.633	522.53
J125311.08+304029.2	SDSSJ125311.62+304017.3	19.958	1.2680	21.850	-0.6406	24.925	467.17
J125419.40+304803.0	SDSSJ125422.44+304728.1	59.711	37.156	21.773	-0.7765	25.322	1063.1
J125931.85+333654.2	SDSSJ125930.81+333646.9	209.23	18.536	23.284	-1.0137	25.414	765.71
J130003.72+263652.1	SDSSJ130004.25+263652.7	201.85	18.201	22.760	-0.9195	25.362	469.07
J130013.55+273548.7	SDSSJ130014.65+273600.0	184.82	22.325	22.740	-0.9937	25.153	855.75
J130332.38+312947.1	SDSSJ130332.47+312949.5	148.87	8.6801	23.450	-0.6124	25.673	575.96
J130415.46+225322.6	SDSSJ130414.25+225305.0	6.4460	3.800	22.155	-0.7534	25.583	342.52
J130532.49+315639.0	SDSSJ130532.02+315634.8	<26.245		<21.918	-0.9945	25.179	349.11
J130548.65+344052.7	SDSSJ130548.81+344053.8	<28.477		<21.614	-0.5835	24.893	314.38
J130640.99+233824.8	SDSSJ130641.12+233823.5	300.95	9.3260	22.416	-0.5683	24.214	387.76
J130825.47+330508.2	SDSSJ130826.28+330515.0	6.8182	6.7591	21.582	-0.9181	25.422	427.12
J130916.85+305118.3	SDSSJ130916.02+305121.9	<20.957		<21.860	-1.3005	25.269	480.09
J130915.61+230309.5	SDSSJ130916.66+230311.3	525.18	2.2780	22.864	-0.3507	24.136	144.66
J131040.25+322044.1	SDSSJ131040.03+322047.6	<61.650		<22.780	-1.4250	26.092	584.94
J131039.92+265111.9	SDSSJ131040.52+265105.5	<25.706		<21.360	<-0.3644	23.637	659.07
J131235.35+331348.6	SDSSJ131231.68+331436.3	671.47	25.829	23.692	-1.0759	25.762	1634.8
J131405.16+243234.1	SDSSJ131405.99+243240.1	645.52	24.911	23.735	-0.7281	25.715	637.17
J131446.57+252819.8	SDSSJ131446.83+252820.5	<34.176		<22.510	-1.1086	25.403	267.44
J132402.51+302830.1	SDSSJ132403.45+302822.8	74.599	11.863	20.612	-1.2779	23.468	83.355
J132602.06+314645.6	SDSSJ132602.42+314650.4	<186.50		<22.451	<-0.9743	24.387	333.47
J132622.12+320512.1	SDSSJ132622.56+320502.3	335.72	5.9261	23.073	<-1.1796	24.871	792.12
J132738.77+350644.6	SDSSJ132737.92+350650.2	154.72	76.843	23.085	-0.8113	26.709	1668.8
J132949.25+335136.2	SDSSJ132948.87+335152.2	978.84	37.237	23.995	-0.8557	25.883	884.27
J133016.12+315923.9	SDSSJ133016.33+315919.7	<317.95		<22.939	-1.0848	25.552	412.15
J133058.91+351658.9	SDSSJ133057.34+351650.2	124.15	20.506	22.546	-0.6594	25.478	544.62
J133309.94+251045.2	SDSSJ133310.56+251044.1	<1.0794		<20.315	-0.9218	25.148	311.21
J133422.15+343640.5	SDSSJ133422.21+343634.8	<25.727		<22.410	-0.8020	25.910	489.14
J133502.36+323312.8	SDSSJ133502.38+323313.9	70.684	26.627	22.594	-0.7828	25.923	489.38
J133642.53+352009.9	SDSSJ133643.09+352011.7	127.35	17.930	21.612	<-1.0860	23.883	111.54
J134702.03+310913.3	SDSSJ134701.71+310924.2	41.891	12.639	21.641	-0.6901	25.105	370.29
J134802.83+322938.3	SDSSJ134802.70+322940.1	<30.713		<21.554	-0.6274	24.787	293.52

Table 5. Table of source properties used for the analysis in Section 4. We also include the SDSS names of the optical host, updating the names for sources with new optical IDs. Candidate remnants only have upper limits to the 6 GHz core flux density, and do not have flux density errors. Core luminosities at 6 GHz are also given, based on the flux densities in column 3, redshifts in Table 1 and Table 4, and assuming a power-law spectral index of -0.5 . α_{150}^{1400} gives the spectral index of the extended emission seen by LOFAR and NVSS. Sources that have no corresponding detection at 1.4 GHz with NVSS have spectral indices with upper limits only. log 150 MHz total power denotes the radio luminosity of the source based on the LOFAR images.

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