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Molecular gas masses of gamma-ray burst host galaxies

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ABSTRACT

Aims. The objectives of this paper are to analyse molecular gas properties of the first substantial sample of GRB hosts and test whether they are deficient in molecular gas.

Methods. We obtained CO(2-1) observations of seven GRB hosts with the APEX and IRAM30m telescopes. We analyse these data together with all other hosts with previous CO observations. From these observations we calculated the molecular gas masses of these galaxies and compared them with the expected values based on their SFRs and metallicities.

Results. We obtained detections for three GRB hosts (980425, 080207 and 111005A) and upper limits for the remaining four (031203, 060505, 060814, 100316D). In our entire sample of twelve CO-observed GRB hosts, three are clearly deficient in molecular gas, even taking into account their metallicity (980425, 060814, and 080517). Four others are close to the best fit-line for other star-forming galaxies on the SFR- M_{H_2} plot (051022, 060505, 080207, and 100316D). One host is clearly molecule-rich (111005A). Finally, for four GRB hosts the data is not deep enough to judge whether they are molecule-deficient (000418, 030329, 031203, 090423). The median value of the molecular gas depletion time, $M_{\text{H}_2}/\text{SFR}$, of GRB hosts is ~ 0.3 dex below that of other star-forming galaxies, but this result has low statistical significance. A Kolmogorov-Smirnov test performed on $M_{\text{H}_2}/\text{SFR}$ shows only $\sim 2\sigma$ difference between GRB hosts and other galaxies. This difference can partially be explained by metallicity effects, since the significance decreases to $\sim 1\sigma$ for $M_{\text{H}_2}/\text{SFR}$ vs. metallicity.

Conclusions. We found that any molecular gas deficiency of GRB hosts has low statistical significance and that it can be attributed to their lower metallicities; and thus the sample of GRB hosts has consistent molecular properties to other galaxies, and can be treated as representative star-forming galaxies. Given the concentration of atomic gas recently found close to GRB and supernova sites, indicating recent gas inflow, our results imply that such inflow does not enhance the SFRs significantly, or that atomic gas converts efficiently into the molecular phase, which fuels star formation. Only if the analysis of a larger GRB host sample reveals molecular deficiency (especially close to the GRB position), then this could support the hypothesis of star formation fuelled directly by atomic gas.

Key words. gamma ray bursts: general – ISM: lines and bands – ISM: molecules – galaxies: ISM – galaxies: star formation – radio lines: galaxies

1. Introduction

Long gamma-ray bursts (GRBs) have long been confirmed to be the endpoints of lives of very massive stars (e.g. Hjorth et al. 2003; Stanek et al. 2003; Hjorth & Bloom 2012). Most of the tracers of the star formation rate (SFR) of galaxies are connected with emission of massive stars (e.g. Kennicutt 1998), so GRBs were also used to measure the star formation history of the Universe (Yüksel et al. 2008; Kistler et al. 2009; Butler et al. 2010; Elliott et al. 2012; Robertson & Ellis 2012; Perley et al. 2016a,b). This approach is valid if GRB hosts are representative star-forming galaxies at a given redshift (Michałowski et al. 2012; Hunt et al. 2014a; Schady et al. 2014; Greiner et al. 2015; Kohn et al. 2015), or if biases are known and can be corrected for (Perley et al. 2013, 2015, 2016a,b; Boissier et al. 2013; Vergani et al. 2015; Schulze et al. 2015; Greiner et al. 2016). Gas is the fuel of star formation, so one of the important aspects of this issue is whether GRB hosts exhibit normal gas properties with respect to other star-forming galaxies.

The information about gas properties of GRB hosts is scarce. Michałowski et al. (2015) and Arabsalmani et al.

(2015) provided the only measurements so far of the atomic gas properties of five such galaxies. This led to a suggestion that GRB hosts have experienced recent inflows of atomic gas. A resulting possibility of using GRBs to select galaxies for the study of gas accretion is important, because the rate of the gas accretion onto galaxies is surprisingly constant since $z \sim 5$, at odds with the significantly changing SFR volume density of the Universe (Spring & Michałowski 2017). Moreover, a fraction of star formation in GRB hosts may be fuelled directly by atomic gas (Michałowski et al. 2015, 2016). The existence of this process is controversial, but it has been predicted theoretically (Glover & Clark 2012; Krumholz 2012; Hu et al. 2016; Elmegreen 2018), and is supported by some observations (Bigiel et al. 2010; Fumagalli & Gavazzi 2008; Elmegreen et al. 2016).

Clearly, most of the star formation in the Universe is fuelled by molecular gas (Fumagalli et al. 2009; Carilli & Walter 2013; Rafelski et al. 2016). There were several unsuccessful searches of CO lines for GRB hosts (Kohn et al. 2005; Endo et al. 2007; Hatsukade et al. 2007, 2011; Stanway et al. 2011) and only four detections so far, for the hosts of GRB 980425 (Michałowski et al. 2016),

051022 (Hatsukade et al. 2014), 080517 (Stanway et al. 2015b), and 080207 (Arabsalmani et al. 2018). These studies resulted in mixed conclusions on whether GRB hosts are deficient in molecular gas with respect to the SFR- M_{H_2} correlation of other star-forming galaxies.

Hence, the objectives of this paper are: *i*) to analyse molecular gas properties of the first substantial sample of GRB hosts; and *ii*) to test whether they are deficient in molecular gas. For this, we combined existing literature data with new observations using the APEX and IRAM30m telescopes.

We use a cosmological model with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_\Lambda = 0.7$, and $\Omega_m = 0.3$. We also assume the Chabrier (2003) initial mass function (IMF), to which all star formation rate (SFR) and stellar masses were converted (by dividing by 1.8) if given originally assuming the Salpeter (1955) IMF.

2. Target selection and data

2.1. APEX

We selected the host galaxies of all known GRBs at $z < 0.12$ in the southern hemisphere (i.e. the sample with HI observations from Michałowski et al. 2015). These criteria were fulfilled by GRB 980425 (the central pointing was published separately in Michałowski et al. 2016), 031203, 060505, 100316D, and 111005A. We performed CO(2-1) observations using the Swedish Heterodyne Facility Instrument (SHeFI; Vassilev et al. 2008; Belitsky et al. 2006) and the Swedish-ESO PI Instrument for APEX (SEPIA; Belitsky et al. 2017; only for the GRB 031203 host) mounted at the Atacama Pathfinder Experiment (APEX; Güsten et al. 2006) (project no. 096.D-0280, 096.F-9302 and 097.F-9308, PI: M. Michałowski). Table 1 shows the observation log with total on-source integration times. Two and three positions were observed for the host of GRB 980425 and 111005A, respectively. The remaining galaxies are smaller than the beam ($\sim 27''$). All observations were carried out in the on-off pattern and the position-switching mode. The fluxes were corrected using the main beam efficiency of 0.75. We reduced and analysed the data using the Continuum and Line Analysis Single Dish Software (CLASS) package within the Grenoble Image and Line Data Analysis Software¹ (GILDAS; Pety 2005).

2.2. IRAM30m

We selected all GRB hosts in the northern hemisphere with infrared or radio detections (Hunt et al. 2014a; Perley et al. 2015; Michałowski et al. 2015) and $z > 1.5$, so that the line is located at lower frequencies and easier to observe. This was fulfilled by GRB 060814 and 080207. We performed observations with the IRAM 30-m telescope (project no. 172-16, PI: M. Michałowski) using the Eight Mixer Receiver² (EMIR; Carter et al. 2012). We implemented wobbler switching mode (with the offset to the reference positions of $60''$), which provides stable and flat baselines and optimises the total observing time. An intermediate frequency (IF) covered the frequency of the CO(2-1) line

Table 1. Log of APEX observations.

| GRB | Obs. Date | time/hr | pwv/mm |
|----------------|-------------|---------|-----------|
| 980425 Center | Total | 4.04 | |
| | 2015 Aug 29 | 0.70 | 1.64–1.70 |
| | 2015 Sep 12 | 0.30 | 0.75–0.85 |
| | 2015 Sep 16 | 0.70 | 1.43–1.57 |
| | 2015 Oct 31 | 1.17 | 1.22–1.96 |
| 980425 WR | Total | 6.57 | |
| | 2015 Nov 02 | 2.17 | 0.75–3.48 |
| | 2016 Apr 03 | 0.10 | 2.02–2.15 |
| | 2016 Apr 04 | 4.30 | 3.33–5.23 |
| 031203 | 2015 Sep 10 | 0.80 | 0.83–0.91 |
| 060505 | Total | 7.00 | |
| | 2015 Aug 28 | 1.20 | 1.50–1.67 |
| | 2015 Aug 29 | 1.40 | 1.38–1.62 |
| | 2015 Sep 02 | 1.40 | 1.55–1.86 |
| | 2015 Sep 03 | 1.00 | 3.36–3.61 |
| | 2015 Sep 04 | 1.00 | 2.50–2.73 |
| | 2015 Sep 06 | 1.00 | 2.45–3.40 |
| 100316D | Total | 6.58 | |
| | 2015 Aug 28 | 2.11 | 1.50–1.62 |
| | 2015 Sep 02 | 1.67 | 1.32–1.93 |
| | 2015 Sep 06 | 2.80 | 2.45–4.80 |
| 111005A Center | Total | 1.65 | |
| | 2015 Sep 01 | 0.75 | 1.00–1.21 |
| | 2015 Sep 12 | 0.20 | 0.72–0.84 |
| | 2015 Sep 15 | 0.70 | 0.64–0.82 |
| 111005A NW | Total | 3.20 | |
| | 2015 Sep 17 | 0.50 | 1.52–1.61 |
| | 2016 Apr 02 | 1.00 | 2.15–2.47 |
| | 2016 Apr 03 | 0.60 | 1.96–2.31 |
| | 2016 Jun 10 | 1.60 | 2.98–3.34 |
| 111005A SE | Total | 2.20 | |
| | 2015 Sep 17 | 0.50 | 1.55–1.65 |
| | 2016 Jun 10 | 0.60 | 3.12–3.32 |
| | 2016 Jun 11 | 1.60 | 2.49–2.83 |

Table 2. Log of IRAM30m observations.

| GRB | Obs. Date | time/hr | $\tau_{225 \text{ GHz}}$ |
|--------|-------------|---------|--------------------------|
| 060814 | Total | 13.10 | |
| | 2017 Feb 01 | 0.40 | 0.29 |
| | 2017 Feb 03 | 1.60 | 0.08–0.23 |
| | 2017 Feb 04 | 3.20 | 0.23–0.51 |
| | 2017 Feb 07 | 1.40 | 0.20–0.39 |
| | 2017 Apr 06 | 0.70 | 0.13–0.17 |
| | 2017 Apr 07 | 2.00 | 0.12–0.20 |
| | 2017 Apr 08 | 2.20 | 0.15–0.19 |
| | 2017 Apr 09 | 1.60 | 0.10–1.60 |
| 080207 | Total | 17.80 | |
| | 2017 Feb 01 | 1.60 | 0.28–0.37 |
| | 2017 Apr 11 | 1.90 | 0.27–0.36 |
| | 2017 Apr 12 | 3.70 | 0.23–0.48 |
| | 2017 Apr 13 | 4.30 | 0.20–0.44 |
| | 2017 Apr 14 | 3.30 | 0.22–0.41 |
| | 2017 May 22 | 3.00 | 0.24–0.36 |

We used the Fourier Transform Spectrometers 200 (FTS-200) providing 195 kHz spectral resolution (corresponding to $\sim 0.8 \text{ km s}^{-1}$ at the frequency of CO(2-1) of our targets) and 16 GHz bandwidth in each linear polarisation. The observations were divided into 6-min scans, each consisting of 12 scans 30 s long. The pointing was verified every 1–2 hr. The observing log is presented in Table 2 with total on-source integration times. The observations were carried

¹ <http://www.iram.fr/IRAMFR/GILDAS>

² www.iram.es/IRAMES/mainWiki/EmirforAstronomers

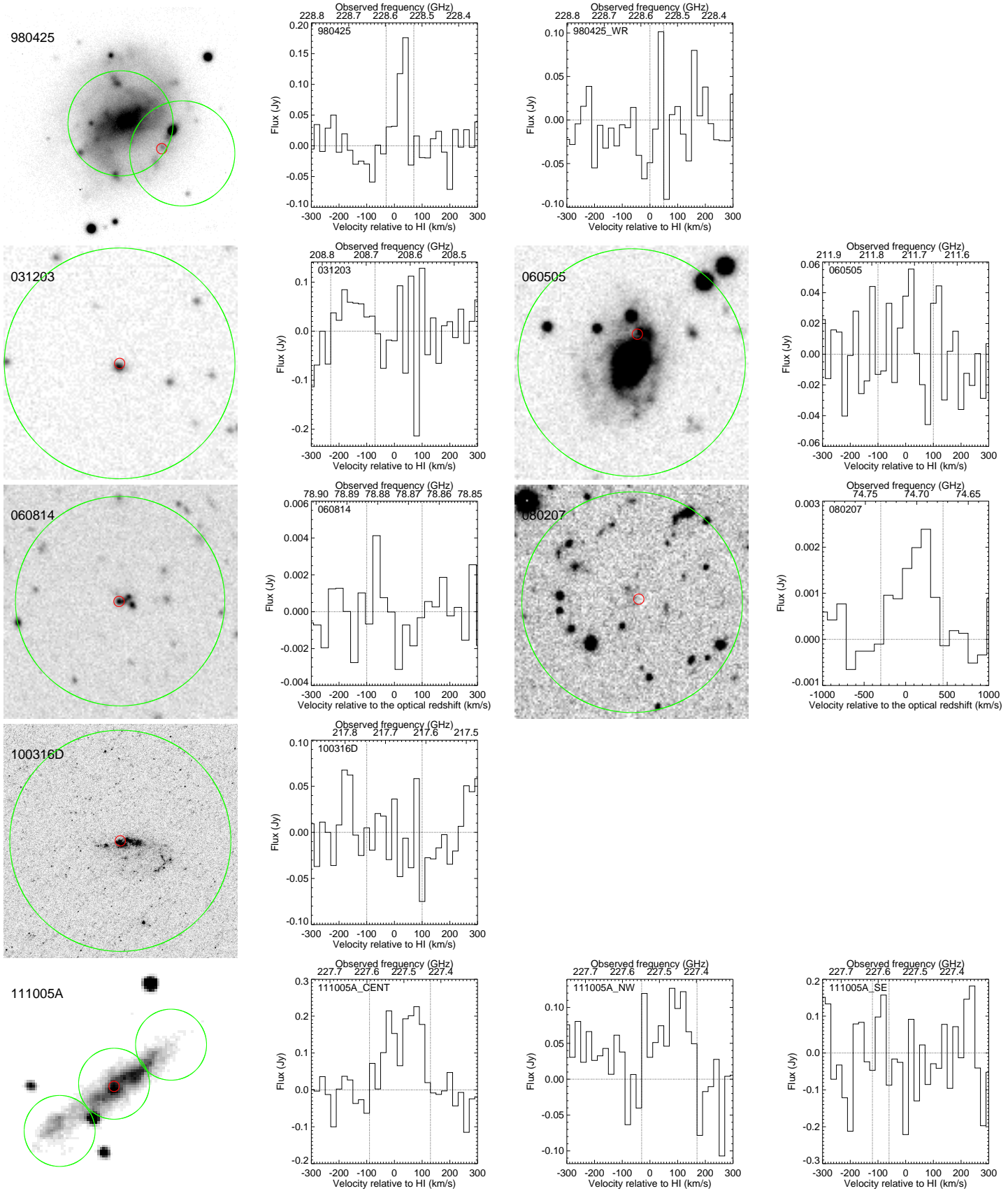


Fig. 1. For each GRB host (labelled in the top-left corner of each panel) the first panel shows the optical image (Sollerman et al. 2005; Mazzali et al. 2006; Thöne et al. 2008; Hjorth et al. 2012; Starling et al. 2011; Michałowski et al. 2018b) together with the *green circles* marking the positions of the pointings and the beam sizes of our CO(2-1) observations. GRB positions are marked as *red circles*. North is up and East is to the left. Further panels show the corresponding CO(2-1) spectra. *Vertical dotted lines* show the velocity intervals within which the line fluxes were measured.

out during good atmospheric conditions and the opacity ($\tau_{225\text{GHz}}$) was uniform across different runs. We reduced the data using the CLASS package within GILDAS (Pety 2005). Each spectrum was calibrated, and corrected for baseline shape. The spectra were aligned in frequency and noise-weight averaged. Some well known platforming, due to the fact that the instantaneous bandwidth of 4 GHz is sampled by three different FTS units, was corrected off-line by a dedicated procedure within CLASS. In all cases, the CO line is far away from the step of the platforming.

2.3. Literature data for additional GRB hosts

In addition to the CO(2-1) measurements obtained here, we included all other CO measurements for GRB hosts from the literature. All molecular masses were converted to $\alpha_{\text{CO}} = 5 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ and to the line luminosity ratios in temperature units $L'_{2-1}/L'_{1-0} = 0.5$, $L'_{3-2}/L'_{1-0} = 0.27$, or $L'_{4-3}/L'_{1-0} = 0.17$ (the Milky Way values, see table 2 of Carilli & Walter 2013) if these masses were based on CO(2-1), CO(3-2), or CO(4-3) observations, respectively. These assumptions lead to conservatively high M_{H_2} , so we are able to robustly test for any molecular deficiency of GRB hosts.

We included the hosts of GRB 000418 (Hatsukade et al. 2011) for which we converted the M_{H_2} upper limit from $L'_{2-1}/L'_{1-0} = 1$ to 0.5 and from $\alpha_{\text{CO}} = 0.8 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ to 5; 030329 (Kohno et al. 2005; Endo et al. 2007) for which we converted the M_{H_2} upper limit from $\alpha_{\text{CO}} = 40 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ to 5; 051022 (Hatsukade et al. 2014) for which we converted the M_{H_2} detection from $L'_{4-3}/L'_{1-0} = 0.85$ to 0.17 and from $\alpha_{\text{CO}} = 4.3 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ to 5; 080517 (Stanway et al. 2015b) for which we converted the M_{H_2} detection from $\alpha_{\text{CO}} = 4.3 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ to 5; 090423 (Stanway et al. 2011) for which we converted the M_{H_2} detection from $L'_{3-2}/L'_{1-0} = 1$ to 0.27 and from $\alpha_{\text{CO}} = 0.8 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ to 5.

We do not use the CO(3-2) observations of GRB 980425 of Hatsukade et al. (2007) because our deeper data resulted in a detection. Moreover we excluded GRB 020819B because the low-redshift galaxy with the existing CO measurement (Hatsukade et al. 2014) has been shown not to be related to the GRB (Perley et al. 2017b). For the GRB 080207 host, the CO(3-2) line observations were recently reported by Arabsalmani et al. (2018). We do not use these values in subsequent analysis, because our lower transition likely traces a larger fraction of the total molecular gas content. We note, however, that the obtained gas masses are consistent (see Sect. 3).

For all GRB hosts in our CO sample we used the literature values for their redshifts (Tinney et al. 1998; Bloom et al. 2003; Greiner et al. 2003; Hjorth et al. 2003; Prochaska et al. 2004; Hjorth et al. 2012; Castro-Tirado et al. 2007; Ofek et al. 2006; Stanway et al. 2015a; Tanvir et al. 2009; Salvaterra et al. 2009; Vergani et al. 2010; Starling et al. 2011; Levan et al. 2011; Michałowski et al. 2018b), SFRs (Michałowski et al. 2009, 2012, 2014, 2015, 2018b; Castro Cerón et al. 2010; Starling et al. 2011; Watson et al. 2011; Tanvir et al. 2012; Perley et al. 2015, 2017a; Hunt et al. 2014a; Stanway et al. 2015a; Walter et al. 2012; Tanga et al. 2017) and metallicities (Sollerman et al. 2005; Christensen et al.

2008; Svensson et al. 2010; Levesque et al. 2010, 2011; Thöne et al. 2008; Krühler et al. 2015; Stanway et al. 2015a; Michałowski et al. 2018b; Tanga et al. 2017).

For the host of GRB 060814 we calculated the metallicity based on the R_{23} method of Kobulnicky & Kewley (2004) based on the [O II], [O III], and H β emission lines, using the fluxes reported in Krühler et al. (2015). We obtained $12 + \log(\text{O}/\text{H}) \sim 8.38 \pm 0.35$.

Additionally we included values measured for the host of SN 2009bb, the relativistic supernova (SN) type Ic (Michałowski et al. 2018a). SNe of this type may have similar engines as GRBs, but we do not use it for statistical analysis, as no γ -rays were detected from it.

2.4. Other galaxy samples

In order to place the GRB hosts in the context of general galaxy populations, we compared their properties with those of the following galaxy samples, chosen based on the availability of the gas mass estimates: the optical-flux-limited spirals and irregulars with IRAS data (Young et al. 1989), local Luminous Infrared Galaxies (LIRGs; Sanders et al. 1991), local Ultra Luminous Infrared Galaxies (ULIRGs; Solomon et al. 1997), the *Herschel* Reference Survey (HRS; Boselli et al. 2010; Cortese et al. 2012, 2014; Boselli et al. 2014; Ciesla et al. 2014), H I-dominated, low-mass galaxies and large spiral galaxies (Leroy et al. 2008), $0.01 < z < 0.03$ mass-selected galaxies with $8.5 < \log(M_*/M_{\odot}) < 10$ (Bothwell et al. 2014), $0.025 < z < 0.2$ mass-selected galaxies with $\log(M_*/M_{\odot}) > 10$ and infrared detections (Bertemes et al. 2018), metal-poor dwarfs (Hunt et al. 2014b, 2015, 2017; Leroy et al. 2007), metal-poor dwarfs from the *Herschel* Dwarf Galaxy Survey (Madden et al. 2013; Cormier et al. 2014), Virgo-cluster dwarfs (Grossi et al. 2016), $z \sim 1.5$ BzK galaxies (Daddi et al. 2010; Magdis et al. 2011; Magnelli et al. 2012), and $1.2 < z < 4.1$ submm galaxies (Bothwell et al. 2013; Michałowski et al. 2010).

All SFRs were converted to the Chabrier (2003) IMF. The molecular masses were converted to $\alpha_{\text{CO}} = 5 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ and to the Milky Way line ratios if they were based on higher CO transitions. Namely, Bothwell et al. (2014), Daddi et al. (2010), and Leroy et al. (2008) assumed $L'_{2-1}/L'_{1-0} = 1$, 0.16, and 0.8 respectively, and Hunt et al. (2014b) assumed $L'_{3-2}/L'_{1-0} = 0.6$. The Galactic value of α_{CO} is appropriate for 0.4–1 solar metallicity galaxies discussed here (Bolatto et al. 2013; Hunt et al. 2014b). Following Hunt et al. (2015), metallicities from Bothwell et al. (2014) were converted from the calibration of Kewley & Dopita (2002, KD02) to that of Pettini & Pagel (2004, PP04 N2) using the equation derived by Kewley & Ellison (2008, their table 3).

Even though SFR estimates of other galaxies are often derived from various diagnostics (UV, H α , IR, radio), they are broadly consistent (Salim et al. 2007; Wijesinghe et al. 2011; Davies et al. 2016; Wang et al. 2016), even in dwarf galaxies, except at very low SFR $< 0.001 M_{\odot} \text{yr}^{-1}$ (Huang et al. 2012; Lee et al. 2009), not discussed here.

3. Results

The positions of our APEX and IRAM30m pointings and the obtained CO(2-1) spectra are shown in Fig. 1. The

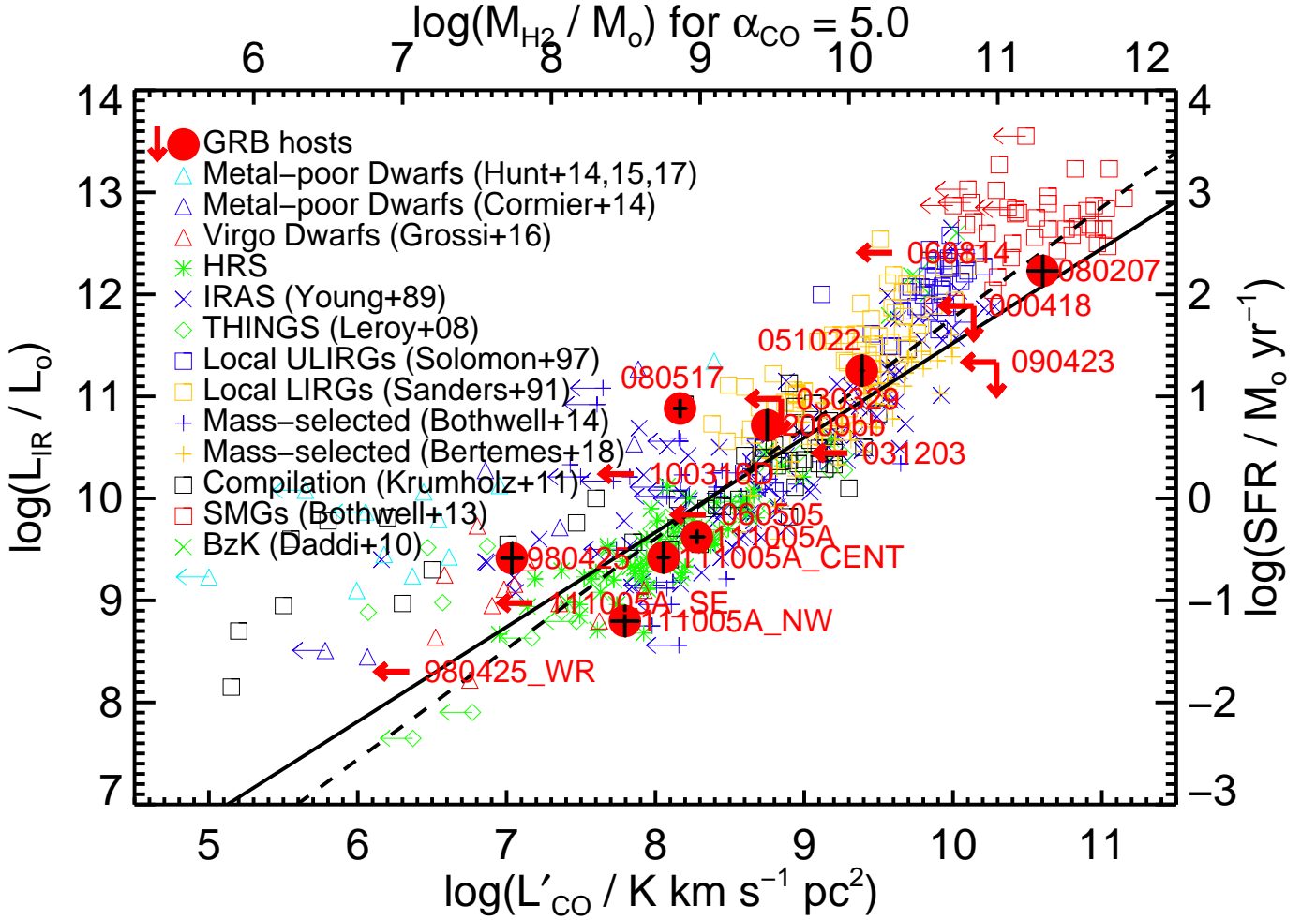


Fig. 2. Infrared luminosity or the corresponding star formation rate (SFR) as a function of CO luminosity, or the corresponding molecular gas mass with the CO-to-H₂ conversion factor $\alpha_{\text{CO}} = 5 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$. GRB hosts are marked with *full red circles or red arrows* with crosses showing the errors. The symbols of other galaxies are indicated in the legend and described in Sect. 2.4. The *solid black line* is a linear fit to the non-GRB galaxies excluding ULIRGs (eq. 1), whereas the *dashed black line* represents the fit including ULIRGs (eq. 2). The ~ 0.3 dex shift for GRB hosts towards lower M_{H_2} is not statistically significant (see Sect. 3.1).

Table 3. APEX and IRAM30m CO(2-1) line fluxes and luminosities.

| GRB | F_{int} (Jy km s ⁻¹) | S/N | F_{int} (10 ⁻²⁰ W m ⁻²) | log L (L_{\odot}) | log L' (K km s ⁻¹ pc ²) | log $M_{\text{H}_2, \text{CO}}$ (M_{\odot}) |
|--------------|--|------|--|----------------------------|---|--|
| (1) | (2) | (3) | (4) | (5) | (6) | (7) |
| 980425 | 6.50 ± 1.25 | 5.2 | 5.00 ± 0.96 | 3.33 ± 0.08 | 6.73 ± 0.08 | 7.73 ± 0.08 |
| 980425_WR | 1.33 ± 1.28 | 1.0 | 1.02 ± 0.98 | 2.64 ± 0.29 | 6.04 ± 0.29 | 7.04 ± 0.29 |
| 031203 | 7.51 ± 3.35 | 2.2 | 5.77 ± 2.58 | 5.58 ± 0.16 | 8.99 ± 0.16 | 9.99 ± 0.16 |
| 060505 | 1.18 ± 1.64 | 0.7 | 0.91 ± 1.26 | 4.63 ± 0.38 | 8.04 ± 0.38 | 9.04 ± 0.38 |
| 060814 | -0.04 ± 0.11 | -0.4 | -0.03 ± 0.09 | <6.51 | <9.92 | <10.92 |
| 080207 | 0.38 ± 0.11 | 3.5 | 0.29 ± 0.08 | 6.90 ± 0.11 | 10.30 ± 0.11 | 11.30 ± 0.11 |
| 100316D | -0.88 ± 2.25 | -0.4 | -0.68 ± 1.73 | <4.76 | <8.16 | <9.16 |
| 111005A_CENT | 28.49 ± 2.94 | 9.7 | 21.91 ± 2.26 | 4.35 ± 0.04 | 7.75 ± 0.04 | 8.75 ± 0.04 |
| 111005A_NW | 15.69 ± 3.93 | 4.0 | 12.07 ± 3.02 | 4.09 ± 0.10 | 7.49 ± 0.10 | 8.49 ± 0.10 |
| 111005A_SE | 3.75 ± 3.08 | 1.2 | 2.89 ± 2.37 | 3.47 ± 0.26 | 6.87 ± 0.26 | 7.87 ± 0.26 |

Notes. (1) GRB (2) Integrated flux within the velocity interval shown by the dotted lines on Fig. 1. (3) Signal-to-noise ratio of the line within this velocity interval. (4) Corresponding integrated flux in W m⁻². (5) Line luminosity. (6) Line luminosity in temperature units based on equation 3 in Solomon et al. (1997). (7) Molecular gas mass estimated assuming $L'_{\text{CO}(1-0)} = 2 \times L'_{\text{CO}(2-1)}$ (see Sect. 2.3 and 3) and the Galactic CO-to-H₂ conversion factor $\alpha_{\text{CO}} = 5 M_{\odot} / \text{K km s}^{-1} \text{pc}^2$.

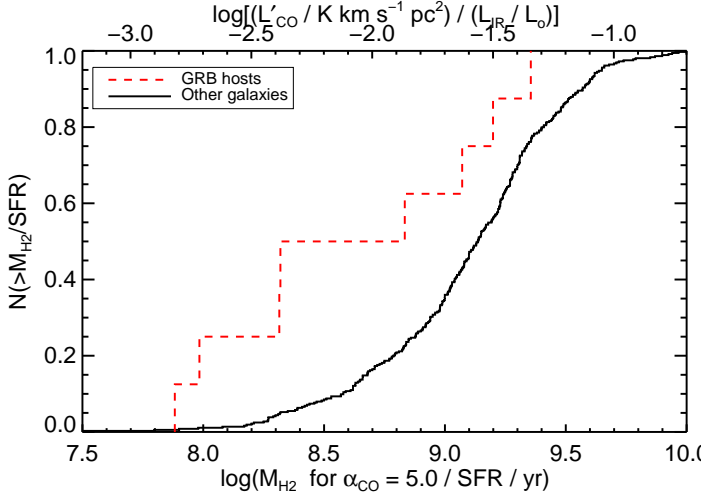


Fig. 3. Cumulative distribution of molecular gas depletion time (or the inverse of the star formation efficiency), i.e. the ratio of the CO luminosity to the infrared luminosity or the corresponding molecular gas mass with the CO-to-H₂ conversion factor $\alpha_{\text{CO}} = 5 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ to the star formation rate (SFR). The distribution of GRB hosts is shown as the *dashed red line*, whereas that of other galaxies is shown as the *solid black line*. We treated the upper limits as actual values, so the histogram for GRB hosts is an upper limit. GRB hosts are systematically shifted to the left on this diagram (lower M_{H_2} given their SFRs), but this is not statistically significant (see Sect. 3.1).

spectra were binned to a velocity resolution of 20 km s^{-1} , except for the GRB 080207 host for which 50 km s^{-1} channels were adopted. The derived parameters are shown in Table 3. The fluxes were integrated within the velocity ranges shown in Fig. 1 as vertical dotted lines. They were chosen to encompass the full extent of the lines for the detected targets, and the most significant positive feature for the non-detected targets in order to obtain the most conservative upper limits. The CO(2-1) line luminosities were calculated using equation 3 in Solomon et al. (1997) and converted to the CO(1-0) luminosities assuming $L'_{1-0} = 2 \times L'_{2-1}$. The Galactic CO-to-H₂ conversion factor $\alpha_{\text{CO}} = 5 M_{\odot} / \text{K km s}^{-1} \text{pc}^2$ was used to calculate molecular gas masses ($M_{\text{H}_2} = \alpha_{\text{CO}} L'_{1-0}$).

3.1. SFR vs. M_{H_2}

The IR luminosity (or SFR) as a function of CO line luminosity (or M_{H_2}) for GRB hosts and other galaxies is shown in Fig. 2. The best linear fit in log-log space to all non-GRB galaxies with SFRs lower than those of ULIRGs (SFR $< 100 M_{\odot} \text{ yr}^{-1}$) is (the solid line in Fig. 2):

$$\log(\text{SFR}/M_{\odot} \text{ yr}^{-1}) = 0.95 \times \log(M_{\text{H}_2}/M_{\odot}) - 8.57 \quad (1)$$

The scatter around this relation is ~ 0.42 dex. If including ULIRGs this equation changes to (the dashed line in Fig. 2):

$$\log(\text{SFR}/M_{\odot} \text{ yr}^{-1}) = 1.10 \times \log(M_{\text{H}_2}/M_{\odot}) - 9.96 \quad (2)$$

As reported in Michałowski et al. (2016), we found low molecular gas content in the GRB 980425 host given its SFR. Similarly, the hosts of GRB 100316D and 060814

are deficient in M_{H_2} given their SFRs. Our M_{H_2} upper limit for the GRB 031203 host is ~ 0.5 dex higher than the value suggested by the best-fit relation of eq. (1) so we cannot conclude much about its molecular gas content. Our M_{H_2} upper limit for the GRB 060505 host is not sufficiently strong to test for any molecular gas deficiency, but it is close to the best-fit line for other star-forming galaxies, so this galaxy is not richer in molecular gas than the average of other galaxies. We found that the GRB 080207 host is very close to the best-fit line for other galaxies on the SFR- M_{H_2} diagram, consistently with the results of Arabalmani et al. (2018) based on the CO(3-2) line. The host of GRB 111005A is molecule-rich with $\log(M_{\text{H}_2}/\text{SFR}/\text{yr}) \sim 9.34$, i.e. ~ 0.24 dex above the best-fit relation for other galaxies at the relevant SFR. Consistently with Michałowski et al. (2018a), we show that the host of SN 2009bb has a few times lower molecular gas mass than its SFR suggests.

The second pointing for the GRB 980425 host, towards the Wolf-Rayet (WR) region (for its properties see Hammer et al. 2006; Le Floch et al. 2006, 2012; Christensen et al. 2008; Michałowski et al. 2009, 2014, 2016; Krühler et al. 2017) resulted in an upper limit close to the best-fit line. Hence, while we cannot establish any molecular deficiency for this region, it is definitely not molecule-rich, in contrast with its high abundance of atomic gas (Arabalmani et al. 2015).

Both the central and NW regions of the GRB 111005A host are molecule-rich, but the SE region is at least slightly molecule-deficient, given its CO upper limit.

Because of our choice to adopt the Milky Way CO line ratios instead of those of M82 (see Sect. 2.3), we obtained approximately five times higher molecular gas mass for the GRB 051022 host, and hence its molecular gas deficiency is not as dramatic as presented originally in Hatsukade et al. (2014), but still apparent (Fig. 2). Our correction for the GRB 080517 is small with respect to the values used in Stanway et al. (2015b), so we recover its reported molecular gas deficiency.

The revised, lower value of the infrared luminosity of the host of GRB 000418 (compare Michałowski et al. 2008 and Perley et al. 2017b) means that the CO observations (Hatsukade et al. 2011) do not provide useful constraints on its location on the SFR- M_{H_2} diagram (see Fig. 2). Similarly, the upper limits on L_{IR} available for GRB 030329 (Endo et al. 2007) and 090423 (Stanway et al. 2011) do not constrain the positions of these galaxies relative to the best-fit SFR- M_{H_2} relation. Hence we do not use these three hosts with upper limits for both SFRs and M_{H_2} in the statistical analysis.

The median value of the molecular gas depletion time for non-GRB galaxies is $\log(M_{\text{H}_2}/\text{SFR}/\text{yr}) = 9.099^{+0.031}_{-0.020}$, whereas for GRB hosts it is $8.83^{+0.24}_{-0.52}$, where we treated the upper limits as actual values, so the value for GRB hosts is an upper limit. Hence, GRB hosts have molecular gas masses ~ 0.3 dex below the expectations from their SFR, but this result has low significance.

The cumulative distributions of the $M_{\text{H}_2}/\text{SFR}$ ratio (molecular gas depletion time) is shown in Fig. 3. For these statistics we excluded hosts with weak upper limits (031203) and those with upper limits for both M_{H_2} and SFRs (000418, 030329, and 090423). Using the Kolmogorov-Smirnov (K-S) test, we found that we can

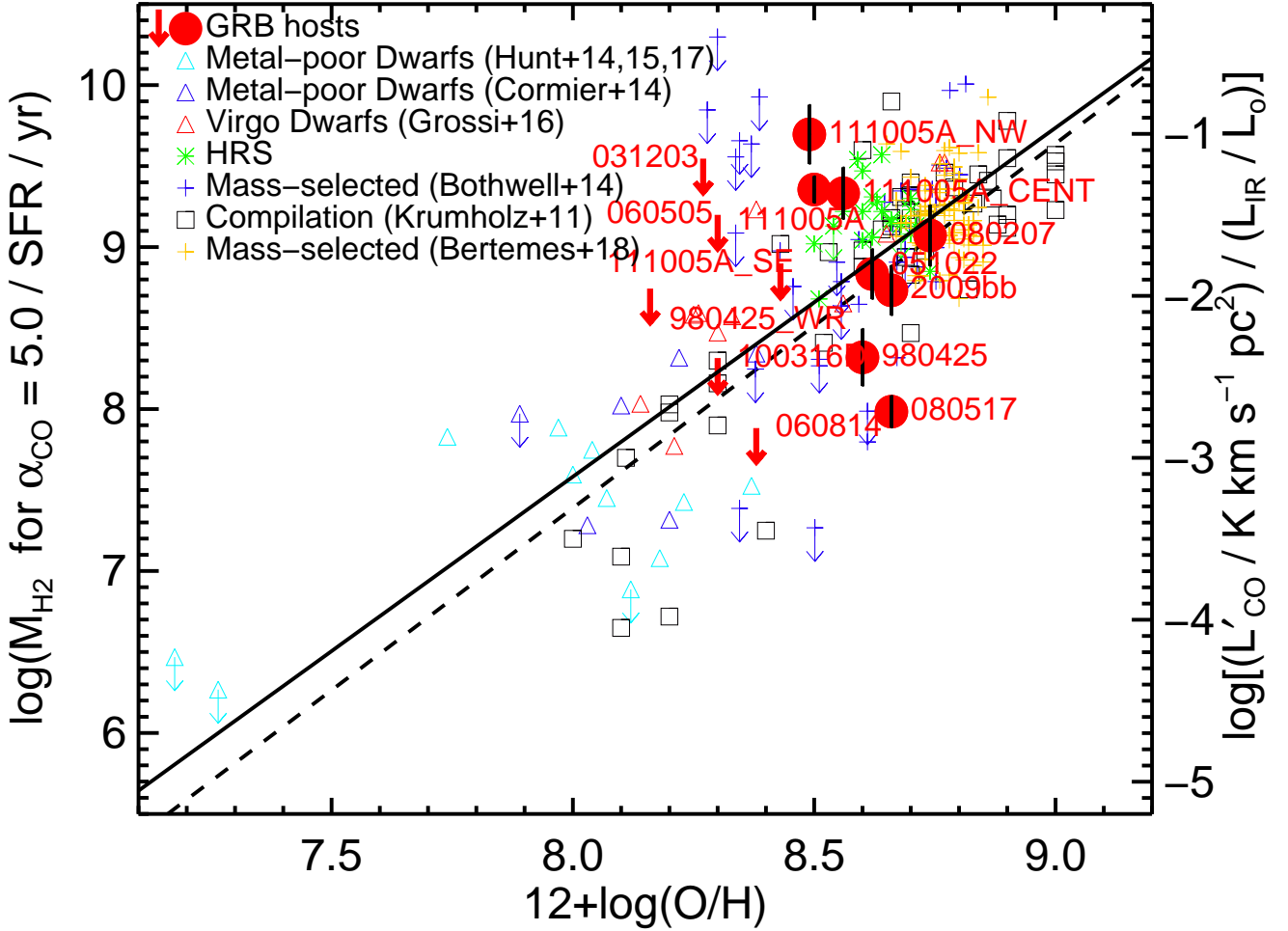


Fig. 4. Molecular gas depletion time (or the inverse of the star formation efficiency), i.e. the ratio of the CO luminosity to the infrared luminosity or the corresponding molecular gas mass with the CO-to-H₂ conversion factor $\alpha_{\text{CO}} = 5 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ to the star formation rate (SFR) as a function of metallicity. GRB hosts are marked with *full red circles or red arrows* with vertical bars showing the errors. The symbols of other galaxies are indicated in the legend and described in Sect. 2.4. The *solid black line* is our fit to the non-GRB galaxies (eq. 3), whereas the *dashed black line* is the relation found by Hunt et al. (2015). GRB hosts are consistent with other galaxies (see Sect. 3.2).

rule out the null hypothesis that the $M_{\text{H}_2}/\text{SFR}$ values of the GRB hosts were drawn from the same distribution as those of other star-forming galaxies at a significance level $p = 0.07$, corresponding to a difference with a low statistical significance of $\sim 1.8\sigma$.

3.2. $M_{\text{H}_2}/\text{SFR}$ vs. metallicity

The CO-to-H₂ conversion factor is metallicity dependent (e.g. Bolatto et al. 2013), so we explored the $M_{\text{H}_2}/\text{SFR}$ ratio as a function of metallicity (Fig. 4). Using the galaxies with metallicity measurement, the linear fit to all non-GRB galaxies is (the solid line in Fig. 4):

$$\log(M_{\text{H}_2}/\text{SFR}/\text{yr}) = 2.33 \times [12 + \log(\text{O}/\text{H})] - 11.1 \quad (3)$$

The scatter around this relation is ~ 0.35 dex.

The molecular deficiency of the GRB 980425 is confirmed, even taking into account its sub-Solar metallicity, i.e. it has a lower molecular gas depletion time than expected for its SFR and metallicity. This is at odds with the

discussion in Arabalsamani et al. (2018) that this galaxy has normal molecular gas properties. However, they compared M_{H_2} with stellar mass, not SFR, as we do here, and also used the dwarf sample of Grossi et al. (2016) as a comparison, but these galaxies exhibit much lower metallicities than the GRB 980425 host (see Fig. 4). Similarly, the molecular gas deficiency of the hosts of GRB 080517 and 060814 are confirmed after taking into account their metallicities.

The hosts of GRB 051022, 080207, and 100316D have depletion times consistent with the expected values given their metallicities (the GRB 100316D host represents an upper limit, so we do not know whether it is actually close to the best-fit relation). Only the GRB 111005A host is clearly molecule-rich for its metallicity. The limits for the hosts of GRB 031203 and 060505 are not constraining, because they are significantly above the best fit line.

Our upper limit for the WR region of the GRB 980425 host is ~ 0.4 dex above the best fit line in Fig. 4, however, the beam size of our observations is much larger than this

region (Fig. 1), so in reality our observations probe also the higher-metallicity regions.

Similarly to the results presented in Sect. 3.1, the central and NW regions of the GRB 111005A host are rich in molecular gas given their SFR and metallicity. On the other hand, the SE region has much lower molecular gas content, close to the best-fit line.

For GRB hosts the median value of the residual from this best fit is $-0.21 \pm 0.07 \text{ yr}^{-1}$, where we treated the upper limits as actual values, so this value is an upper limit.

The cumulative distributions of residuals around the best-fit line (eq. 3) is shown in Fig. 5. For these statistics we excluded hosts with weak upper limits (031203, 060505) and those with upper limits for both M_{H_2} and SFRs (000418, 030329, and 090423). Using the K-S test, we found that we can reject the null hypothesis that the residuals around the best-fit line for GRB hosts were drawn from the same distribution as those for other star-forming galaxies only at a significance level $p = 0.33$, corresponding to a $\sim 1\sigma$ difference.

3.3. Molecular gas fraction

Using the HI data from Michałowski et al. (2015) we can constrain the molecular gas fraction ($M_{\text{H}_2}/(M_{\text{H}_2} + M_{\text{HI}})$) to be $\sim 7\%$ for the GRB 980425 host, $< 15\%$ for the GRB 060505 host, and $\sim 13\%$ for the GRB 111005A host. This is within scatter of, but on the lower side compared to, other star-forming galaxies (a few to a few tens %; Young et al. 1989; Devereux & Young 1990; Leroy et al. 2008; Saintonge et al. 2011; Cortese et al. 2014; Boselli et al. 2014), and SN hosts (Galbany et al. 2017; Michałowski et al. 2018a).

4. Discussion

We obtained mixed results analysing CO data for twelve GRB hosts from our survey and from the literature. Three GRB hosts are clearly deficient in molecular gas, even taking into account their metallicity (980425, 060814, and 080517). Four others are close to the best fit-line for other star-forming galaxies on the SFR- M_{H_2} plot (051022, 060505, 080207, and 100316D). One host is clearly molecule-rich (111005A). Finally, for four GRB hosts the data is not deep enough to judge whether they are molecule-deficient (000418, 030329, 031203, 090423).

These results suggest that GRB hosts may be preferentially found in galaxies with lower molecular gas content than other star-forming galaxies, as there are more examples of GRB hosts in the M_{H_2} -poor part of the M_{H_2} -SFR diagram, and the median molecular depletion timescale ($M_{\text{H}_2}/\text{SFR}$) of GRB hosts is ~ 0.3 dex below that of other galaxies. However, the difference between GRB hosts and other star-forming galaxies is significant only at the $\sim 2\sigma$ level when analysing $M_{\text{H}_2}/\text{SFR}$ (Figs. 2 and 3). Moreover, the statistical significance of this tentative difference decreases further to the $\sim 1\sigma$ level when taking the metallicity into account (Figs. 4 and 5). Hence, our sample is statistically consistent with other star-forming galaxies.

Recent high-resolution observations of GRB and SN hosts showed concentrations of atomic gas close to the GRB and SN positions (Michałowski et al. 2015, 2018a; Arabsalmani et al. 2015), strongly supporting the hypothesis of recent inflow of gas at these sites. The sample of

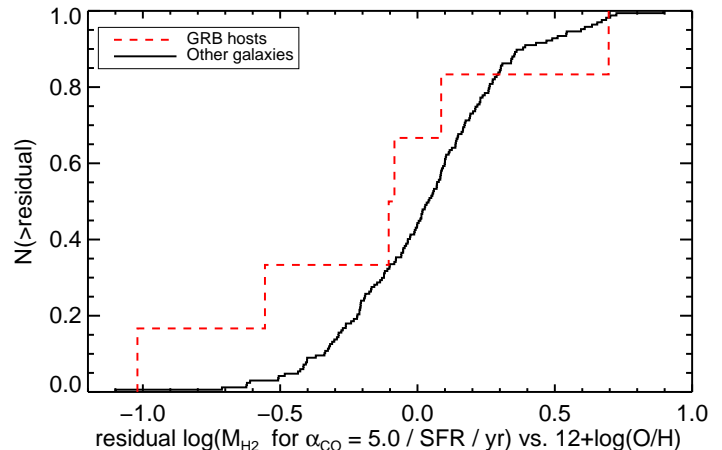


Fig. 5. Cumulative distribution of the residuals with respect to the solid line in Fig. 4 (eq. 3), showing the relation between metallicity and molecular gas depletion time (or the inverse of the star formation efficiency), i.e. the ratio of the CO luminosity to the infrared luminosity or the corresponding molecular gas mass with the CO-to- H_2 conversion factor $\alpha_{\text{CO}} = 5 M_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$ to the star formation rate (SFR). The distribution of GRB hosts is shown as the dashed red line, whereas that of other galaxies is shown as the solid black line. We treated the upper limits as actual values, so the histogram for GRB hosts is an upper limit. GRB hosts are systematically shifted to the left on this diagram (lower M_{H_2} given their SFRs and metallicity), but this is not statistically significant (see Sect. 3.2).

GRB/SN hosts can then be used to study recent gas inflow. Our result of very weak molecular deficiency (if any) implies that either the SFRs of GRB/SN hosts are not significantly enhanced by such inflow, or that atomic gas is efficiently converted into the molecular phase, so SFR and M_{H_2} increase hand in hand.

However, if molecular deficiency is confirmed with a larger sample of GRB hosts, then this will be consistent with a scenario in which their SFRs are enhanced by a recent inflow of atomic gas, which did not have time to convert to the molecular phase. Moreover, small molecular gas content would be consistent with star formation fuelled directly by atomic gas (Michałowski et al. 2015).

Two other issues need to be pointed out. First, most of our M_{H_2} estimates are based on the CO(2-1) line or higher transitions. In order to calculate molecular gas masses we converted these line luminosities to those of the CO(1-0) line assuming a conservatively low Milky Way 2-1/1-0 ratio, giving high M_{H_2} . If however, the gas in GRB hosts is even less excited than the Milky Way, then the real 2-1/1-0 ratio is even lower, and our assumption would result in too low M_{H_2} . This is however unlikely, as GRB hosts are usually found to have large SFR given their stellar masses (Castro Cerón et al. 2006, 2010; Savaglio et al. 2009; Thöne et al. 2009), likely leading to high excitations (see Michałowski et al. 2016). If this is the case generally, then our M_{H_2} are overestimated, and the difference between GRB hosts and other galaxies is stronger than suggested by

our analysis. This can be tested with sensitive observations of other CO transitions (especially 1-0).

Second, our tentative molecular deficiency could result from the assumption of too low α_{CO} . We did take into account the variation of α_{CO} with metallicity (Fig. 4), but it is possible that other properties (e.g. gas density, turbulence) lead to high α_{CO} and result in weak CO emission. This aspect is much more difficult to investigate (also for non-GRB galaxies), because there is no robust way of measuring α_{CO} , especially in non-standard environments.

We also stress that it is important to investigate the molecular gas properties with high-resolution observations. If molecular deficiency is found locally close to the GRB positions, then this will be consistent with star formation fuelled directly by atomic gas. In such a scenario, we are not able to capture this effect using the existing CO data with low spatial resolution, as the hosts on average are not significantly molecule-poor.

This analysis can be improved via the investigation of a larger sample of GRB hosts, and possibly with deeper observations, allowing to probe well below the average molecular gas depletion time of other star-forming galaxies. Moreover, the caveat of our sample is that it is heterogeneous, including low- z hosts and highly star-forming hosts at higher redshifts (Hunt et al. 2011, 2014a; Svensson et al. 2012; Perley et al. 2015). This demonstrates the need of obtaining CO data for a larger sample of homogeneously-selected GRB hosts. This is likely possible only with ALMA, because we have targeted nearby and bright hosts with CO emission potentially easier to detect. ALMA will be able to detect fainter targets, and so will enable studies of a larger and unbiased sample.

5. Conclusions

We observed the CO(2-1) line for seven GRB hosts, obtaining detections for three GRB hosts (980425, 080207 and 111005A) and upper limits for the remaining four (031203, 060505, 060814, 100316D). In our entire sample of twelve CO-observed GRB hosts, including objects from the literature, three are clearly deficient in molecular gas, even taking into account their metallicity (980425, 060814, and 080517). Four others are close to the best fit-line for other star-forming galaxies on the SFR- M_{H_2} plot (051022, 060505, 080207, and 100316D). One host is clearly molecule-rich (111005A). Finally, for four GRB hosts the data is not deep enough to judge whether they are molecule-deficient (000418, 030329, 031203, 090423). The median value of the molecular gas depletion time, $M_{\text{H}_2}/\text{SFR}$, of GRB hosts is ~ 0.3 dex below that of other star-forming galaxies, but this result has low statistical significance. A Kolmogorov-Smirnov test performed on $M_{\text{H}_2}/\text{SFR}$ shows only $\sim 2\sigma$ difference between GRB hosts and other galaxies. This difference can partially be explained by metallicity effects, since the significance decreases to $\sim 1\sigma$ for $M_{\text{H}_2}/\text{SFR}$ vs. metallicity.

We found that any molecular gas deficiency of GRB hosts has low statistical significance and that it can be attributed to their lower metallicities; and thus the sample of GRB hosts has consistent molecular properties to other galaxies, and can be treated as representative star-forming galaxies. Given the concentration of atomic gas recently found close to GRB and SN sites, indicating recent gas inflow, our results imply that such inflow does not

enhance the SFRs significantly, or that atomic gas converts efficiently into the molecular phase, which fuels star formation. Only if the analysis of a larger GRB host sample reveals molecular deficiency (especially close to the GRB position), then this could support the hypothesis of star formation fuelled directly by atomic gas.

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