

ALMA observations of AGN fuelling: the case of PKS B1718–649

F. M. Maccagni^{1,2}, R. Morganti^{1,2}, T. A. Oosterloo^{1,2}, J. B. R. Oonk^{2,3}, and B. H. C. Emonts⁴

¹ Kapteyn Astronomical Institute, University of Groningen, Postbus 800, 9700 AV Groningen, The Netherlands

² ASTRON, Netherlands Institute for Radio Astronomy, Postbus 2, 7990 AA, Dwingeloo, The Netherlands

³ Leiden Observatory, Leiden University, Postbus 9513, 2300 RA Leiden, the Netherlands

⁴ National Radio Astronomy Observatory, 520 Edgemont Road, Charlottesville, VA 22903, USA

January 12, 2018

ABSTRACT

We present ALMA observations of the ^{12}CO (2–1) line of the newly born ($t_{\text{radio}} \sim 10^2$ years) active galactic nucleus (AGN), PKS B1718–649. These observations reveal that the carbon monoxide in the innermost 15 kpc of the galaxy is distributed in a complex warped disk. In the outer parts of this disk, the CO gas follows the rotation of the dust lane and of the stellar body of the galaxy hosting the radio source. In the innermost kiloparsec, the gas abruptly changes orientation and forms a circumnuclear disk ($r \lesssim 700$ pc) with its major axis perpendicular to that of the outer disk. Against the compact radio emission of PKS B1718–649 ($r \sim 2$ pc), we detect an absorption line at red-shifted velocities with respect to the systemic velocity ($\Delta v = +365 \pm 22 \text{ km s}^{-1}$). This absorbing CO gas could trace molecular clouds falling onto the central super-massive black hole. A comparison with the near-infrared H_2 1–0 S(1) observations shows that the clouds must be close to the black hole ($r \lesssim 75$ pc). The physical conditions of these clouds are different from the gas at larger radii, and are in good agreement with the predictions for the conditions of the gas when cold chaotic accretion triggers an active galactic nucleus. These observations on the centre of PKS B1718–649 provide one of the best indications that a population of cold clouds is falling towards a radio AGN, likely fuelling its activity.

Key words. galaxies: individual: PKS B1718–649 – galaxies: active – galaxies: ISM – galaxies: kinematics and dynamics – ISM: clouds - submillimeter: ISM

1. Introduction

The presence of an active galactic nucleus (AGN) in the centre of a galaxy is associated with the accretion of material onto the central super-massive black hole (SMBH), and the consequent release of energy into the host galaxy, either through radiation or relativistic jets of radio plasma, or both. The effect of the nuclear activity onto the interstellar medium (ISM) likely plays a fundamental role in regulating the star formation of the host galaxy as well as the observed relation between the masses of the bulge and the SMBH (Bower et al. 2006; Croton et al. 2006; Booth & Schaye 2009; Ciotti et al. 2010; Faucher-Giguère & Quataert 2012; Debuhr et al. 2012; King & Pounds 2015; King & Nixon 2015). The energy released by the AGN in the ISM depends on the efficiency of the accretion onto the SMBH. However, identifying the gas that is falling onto the SMBH and characterising its kinematics and physical properties has proven to be a particularly difficult task.

AGN are broadly classified as radio-loud when the nucleus expels a considerable fraction of the accreted energy through jets of radio plasma, otherwise they are called radio-quiet. Radio-loud AGN are divided into two categories depending on the efficiency of their accretion rate (e.g. Best et al. 2005; Best & Heckman 2012; Alexander & Hickox 2012; Heckman & Best 2014). On the one hand, radiatively efficient AGN have high accretion rates (between one and ten per cent of their Eddington

rate, λ_{Edd}^1), and the SMBH expels most of the energy through radiation which ionises the surrounding ISM. On the other hand, jet-mode radio AGN have lower accretion rates ($\lambda_{\text{Edd}} \lesssim 1\%$), and most of the energy of the AGN is deposited mechanically in the ISM by jets of radio plasma. The radio jets produce some of the most directly observable phenomena of feedback from AGN, for example inflated bubbles, cavities in the hot gas, or fast outflows of gas driven by the expansion of the radio jets (see e.g. Fabian 2012 for a review).

While radiative mode AGN are considered to be triggered mostly via mergers or secular events (e.g. Hernquist & Spergel 1992; Kormendy & Kennicutt 2004; Comerford et al. 2008; Hopkins & Elvis 2010; Alexander & Hickox 2012), it is uncertain which processes trigger jet-mode AGN. The duration of the nuclear activity ($\sim 10^8$ years) is much shorter than the lifetime of the host galaxy, hence the accretion event is likely short and could occur multiple times throughout the evolution of the galaxy (e.g. Clarke & Burns 1991; Jones & Preston 2001; Shabala et al. 2008; Saikia & Jamroz 2009; Brienza et al. 2015; Morganti 2017). The mechanism by which the fuelling gas is transported to the vicinity of the black hole, and the nature of the accretion flow onto the black hole must play a role determining the efficiency of accretion onto the SMBH.

¹ $\lambda_{\text{Edd}} = \frac{L_{\text{bol}}}{L_{\text{Edd}}}$, L_{bol} is the bolometric luminosity of the source and the Eddington luminosity is $L_{\text{Edd}} = \frac{4\pi G m_p c}{\sigma_T} M_{\text{SMBH}}$; where m_p is the mass of the proton, σ_T the cross section for Thomson scattering, G is the gravitational constant, c the speed of light and M_{SMBH} the mass of the SMBH.

Different accretion mechanisms have been proposed to explain the triggering of jet-mode AGN. One possibility (e.g. Bondi 1952; Narayan & Yi 1995; Ho 2009) is that the accreting material has an inflow time much shorter than its radiative cooling time and hot gas is possibly the main fuel for the nuclear activity (e.g. Hardcastle et al. 2007). The simplest model for accretion of hot gas (Bondi 1952) assumes that the accreting material is initially distributed in a spherically symmetric geometry with negligible angular momentum. Bondi accretion does not realistically describe the circumnuclear environments of a radio AGN but it provides an estimate of the accretion rate which correlates with the power released by the radio AGN (i.e. Allen et al. 2006; Russell et al. 2013).

A more plausible model for the accretion mechanism of jet-mode radio AGN has been proposed based on different numerical simulations which consider that the gas cools down to temperatures lower than 10^3 K prior to the accretion onto the SMBH. These simulations predict that the circumnuclear regions of radio AGN are rich in molecular gas distributed in a clumpy disk, with kinematics characterised both by rotation and turbulence (e.g. Wada & Tomisaka 2005; Wada et al. 2009). Accretion of small ($30 \leq r \leq 150$ pc) clouds of gas is a process that could explain both the low accretion rates typical of these active nuclei as well as the short time-scale of the fuelling of the radio AGN (Nayakshin & Zubovas 2012; Gaspari et al. 2013a; Hillel & Soker 2013; King & Nixon 2015; Gaspari et al. 2013b; Gaspari & Sadowski 2017; Gaspari et al. 2017a). In this process, dense clouds of cold gas form within a hot medium because of turbulence, and only within a few Bondi radii the chaotic inelastic collisions between the clouds are strong enough to cancel their angular momentum and the clouds accrete onto the SMBH.

Over the past few years, different observations have allowed us to characterise the kinematics and the physical conditions of the cold atomic and molecular gas in the circumnuclear regions of a handful of radio AGN (e.g. Davies et al. 2009, 2012; Müller-Sánchez et al. 2013; Combes et al. 2013, 2014; García-Burillo et al. 2014; Viti et al. 2014; Davies et al. 2014; Morganti et al. 2015; Martín et al. 2015; Dasyra et al. 2016; Garcia-Burillo et al. 2016; Tremblay et al. 2016; Russell et al. 2016; Espada et al. 2017). This enabled us to study the effect of the active nucleus on the surrounding molecular medium and, to some degree, to trace gas that may contribute to fuel the central SMBH. Hints that chaotic cold accretion can fuel active galactic nuclei have been found in different objects. For example, PKS 2322-12 (Tremblay et al. 2016), NGC 5044 (David et al. 2014), PKS 0745-191 (Russell et al. 2016) show molecular clouds falling towards the active nucleus, that, if located close to the SMBH, could possibly accrete onto it. Clouds of H I gas detected through absorption could contribute to sustain the nuclear activity of, for example, NGC 315 (Morganti et al. 2009). Observations of the molecular gas in a few young radio sources (e.g. PKS B1718-649, Maccagni et al. 2016; 4C +31.04, Labiano et al. 2013a; 3C 293, Labiano et al. 2013b; Centaurus A, Espada et al. 2017) infer a clumpy circumnuclear disk and find indications that the cold gas may contribute to fuel the central radio source.

In most AGN where indications of accretion of cold gas have been found, it has been very difficult to determine the distance of the in-falling material from the SMBH, and to understand under which physical conditions of the ISM a radio AGN can be triggered. Often, the identification of the in-falling material is subject to many uncertainties, and the radio jets have already expanded for hundreds of parsecs into the ISM, possibly changing the physical conditions of the gas compared to before the nuclear activity was triggered. To study the fuelling of the central

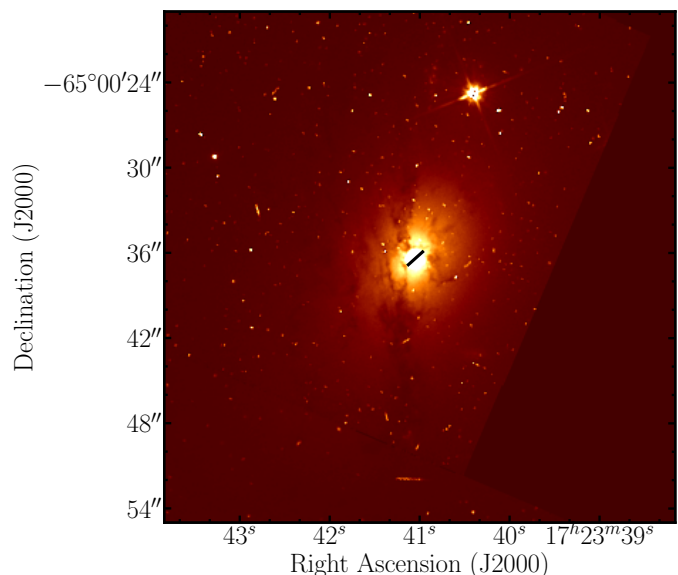


Fig. 1: Hubble Space Telescope WFC2 image of PKS B1718-649 within the field of view of the ALMA telescope. The black line shows the position angle of the radio AGN ($PA_{\text{radio}} = 135^\circ$).

nucleus, it is important to select a radio AGN where the activity did not have time to perturb the surrounding ISM. The ideal candidate would be a very young radio source, where the radio jets extend for at most a few parsecs, and their impact on the ISM is limited.

Young AGN with compact radio jets (i.e. the jets are embedded within the host galaxy) are rich in atomic and molecular gas compared to other more evolved AGN (e.g. Emonts et al. 2010; Guillard et al. 2012; Geréb et al. 2015; Maccagni et al. 2017). The kinematics of the gas in the proximity of the AGN seems to be more unsettled than in older AGN and likely fuelling is still on-going (e.g. Allison et al. 2012; Geréb et al. 2015; Maccagni et al. 2017). These observations lack the spatial resolution to resolve individual cloud-like structures, but their results strongly suggest that compact and young radio AGN are the best candidates to study the triggering and fuelling of low-efficiency active nuclei.

In the following section, we introduce the main characteristics of PKS B1718-649, which is a nearby ($z=0.01442$)² young radio AGN, thus, an ideal candidate to study the triggering of young active galactic nuclei.

1.1. A baby radio source: PKS B1718-649

PKS B1718-649 is one of the youngest sources ($t_{\text{radio}} \sim 10^2$ years, Giroletti & Polatidis 2009) among nearby radio AGN. The radio source is compact, extending only 2 pc and it is classified as a Giga-Hertz Peaked Spectrum source (GPS, Tingay et al. 1997, 2002). Previously, we observed the source in atomic hydrogen (H I) with the Australia Telescope Compact Array (ATCA) and in warm molecular hydrogen (H₂, $2.12\mu\text{m}$) with SINFONI. The results are shown in Maccagni et al. (2014, 2016).

² Throughout this paper we use a Λ CDM cosmology, with Hubble constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_\Lambda = 0.7$ and $\Omega_M = 0.3$. At the distance of PKS B1718-649 (62.4 Mpc), $1'' = 294 \text{ pc}$.

PKS B1718–649 is hosted by an early-type galaxy and is embedded in a large-scale H I disk with regular kinematics (Veron-Cetty et al. 1995; Maccagni et al. 2014). The regularity of the H I disk rules out interaction events as responsible for the fuelling of the radio source (Maccagni et al. 2014).

Figure 1 shows the optical emission of the central regions of PKS B1718–649 observed by the Hubble Space Telescope WFPC2. A dust-lane oriented north-south crosses the central bulge. Clouds of dust are detected also in other regions of the field. The black line in the figure shows the orientation of the compact radio emission of the AGN ($PA_{\text{radio}} = 135^\circ$, $r_{\text{radio}} \sim 2$ pc, Tingay et al. 1997).

The observations of the H₂ 1-0 S(1) line emission at $2.12\,\mu\text{m}$ (Maccagni et al. 2016) show that in the innermost kilo-parsec warm H₂ is rotating in a disk that abruptly changes its orientation at $r \sim 650$ pc. At outer radii, the disk is oriented north-south and it follows the rotation of the other components of the galaxy (e.g. stars, H I disk), while at radii $\lesssim 650$ pc the disk is oriented along the east-west direction. The inner circumnuclear disk is clumpy and is characterised by rotation around the centre.

However, in the innermost regions of this disk, both the H I and the H₂ gas show a component with complex kinematics deviating from the regular rotation of the disk. Against the compact radio continuum emission, we detect two H I absorption lines with opposite velocities with respect to the systemic velocity of the galaxy. These lines trace two clouds of cold gas which are not rotating within the galactic disk. The clouds are small ($M_{\text{H I}} < 10^7 M_\odot$), otherwise they would have been detected in emission. As discussed in Maccagni et al. (2014), these clouds may belong to a larger population of clouds surrounding the radio source and may contribute to its fuelling. In the innermost hundreds of parsecs, a component of H₂ is also found to deviate from the regular rotation of the disk with red-shifted velocities with respect to the systemic ($\gtrsim +150\text{ km s}^{-1}$, Maccagni et al. 2016). Given the clumpiness of the ISM surrounding the radio source, it is likely that the redshifted H₂ and the absorbed H I gas belong to the same nuclear components, in front of the radio source within the innermost few hundred parsec. If this were the case, since the velocities of the H₂ gas are red-shifted, we would observe molecular hydrogen directly falling towards the SMBH.

The spatial resolution of the $2.12\,\mu\text{m}$ observations does not allow us to further constrain the kinematics of the molecular hydrogen close to the SMBH, neither to understand if and how molecular clouds could fuel the newly born radio AGN. Nevertheless, complementary observations also suggest that gas in the form of dense clouds may actively contribute to the accretion onto the SMBH. In the optical band, PKS B1718–649 is a weak AGN, classified as a LINER. From the study of the optical emission lines, Filippenko (1985) suggested that the circumnuclear environment of PKS is populated by dense clouds ($\sim 1.5 \times 10^{21}\text{ cm}^{-2}$). A clumpy circumnuclear ISM could also explain the variability of the peak of the radio spectrum due to free-free absorption (Tingay et al. 2015). While relatively weak in the optical, PKS B1718–649 has an excess of X-ray emission. This could be due to an extended thermal component typical of complex gaseous environments (Siemiginowska et al. 2016). Moreover, PKS B1718–649 was the first compact young radio AGN where γ -ray emission has been detected (Migliori et al. 2016).

In PKS B1718–649 we may be witnessing the signatures of clouds falling towards the AGN, actively contributing to its feeding, through the H I lines and the broad red-shifted molecular component (Maccagni et al. 2014, 2016). If this was the case, higher resolution and sensitivity observations of the molecular gas in the central regions of the galaxy could provide us a de-

tailed view of the interplay between the gas and the nuclear activity, and to understand how this radio AGN was triggered.

In this paper, we present results from the ALMA observations of the ¹²CO (2–1) line of the central 15 kpc of PKS B1718–649. In Sect. 2, we describe the observations and the data reduction. In Sect. 3, we reveal the complexity of the distribution and kinematics of the molecular gas. In Sect. 3.2, we focus on the study of kinematics of the gas in the circumnuclear ($r < 700$ pc) regions. In Sect. 3.3, we analyze the kinematics of the ¹²CO (2–1) gas detected in absorption against the compact radio AGN. In Sect. 3.4, we determine the masses of the molecular gas in different regions of PKS B1718–649. In Sect. 4.1, we compare the distribution and kinematics of the ¹²CO (2–1) with the warm H₂ (Maccagni et al. 2016). In Sect. 4.2, we discuss evidence for infalling clouds towards the radio source, and in Sect. 4.3 we analyze how accretion of these clouds onto the SMBH could sustain the radio nuclear activity of PKS B1718–649. Section 5 summarizes our results and conclusions.

2. Observations

The ¹²CO (2–1) line of PKS B1718–649 was observed by ALMA during Cycle 3 in September 2016 (ID: 2015.1.01359.S, PI: F. M. Maccagni). The observing properties are summarized in Table 1. The observations cover the innermost 15 kpc of PKS B1718–649 with a single-pointing field of view (FoV) of $\sim 51.2''$. The ALMA extended antenna configuration (C36 – 6, maximum baseline of 16 km) allows us to reach maximum spatial resolution of $0.22''$, and is sensitive to a maximum recoverable scale of $2''$. At the redshift of PKS B1718–649, the ¹²CO (2–1) line is found at a frequency of $\nu_{\text{obs}} = 227.591\text{ GHz}$. The Band 6 receivers were used for the observations. The spectral window of the observation was 1.875 GHz , divided in 240 channels of 7.81 MHz ($\sim 10.3\text{ km s}^{-1}$). The duration of the observation, including the time for calibration, was 3.4 hours. The initial calibration was done in CASA (McMullin et al. 2007) using the ALMA data reduction scripts. These calibrated uv -data were subsequently exported to MIRIAD (Sault et al. 2006) which was used to perform additional self-calibration of the continuum point source. This improved the quality of the images and data cubes. All further reduction steps (continuum subtraction, mapping/cleaning) were also done in MIRIAD.

Table 1: Specifications of the ALMA final data cube

Parameter	Value
Field of view	$51.2'' \times 51.2''$ ($15 \times 15\text{ kpc}$)
Synthesized beam	$0.28'' \times 0.19''$ ($82 \times 82\text{ pc}$)
Velocity Resolution	48 km s^{-1}
r.m.s noise per channel	$0.09\text{ mJy beam}^{-1}$

The data cubes were made using various Briggs weightings (Briggs 1995) to explore which one would allow us to better image the molecular gas. In this paper, we show the data cube obtained using natural weighting and smoothed to a velocity resolution of 48 km s^{-1} , except when indicated otherwise. The r.m.s. noise per channel of the data cube is $0.09\text{ mJy beam}^{-1}$. The size of the synthesized beam of the final data cube is $0.28'' \times 0.19''$ ($PA = 41^\circ$), which corresponds to a spatial resolution of $\sim 82\text{ pc}$.

A continuum image was also produced from the uv -data using uniform weighting. The r.m.s noise of the continuum image is 0.1 mJy beam^{-1} and the restoring beam is $0.24'' \times 0.15''$

($PA = 37^\circ$). Since the emission of the AGN is strong at 230 GHz ($S_{\text{cont}} = 303$ mJy), the quality of the continuum image is limited by the dynamic range. The continuum source is unresolved, as expected from its measured linear size at cm wavelengths ($r_{\text{radio}} \sim 2$ pc). The 230 GHz continuum flux found in the ALMA observations is about a factor two lower than what expected from the extrapolation of the data at frequencies of a few GHz (e.g. Tingay et al. 1997, 2002; Giroletti & Polatidis 2009). Likely, this occurs because the radio source is known to be variable (Tingay et al. 2015) and monitoring observations (Moss et al. in prep) show that on the time scale of a year, the flux of the source varies by more than 50%.

3. Results

In this section, we present the analysis of the data cube of the observations of the ^{12}CO (2–1) line of PKS B1718–649. The main results are that the gas detected in emission reveals a complex clumpy distribution overall dominated by rotation (Sect. 3.1) and that against the central compact radio continuum, ^{12}CO (2–1) gas is detected in absorption at red-shifted velocities ($\sim +365$ km s $^{-1}$) with respect to the systemic velocity (Sect. 3.3).

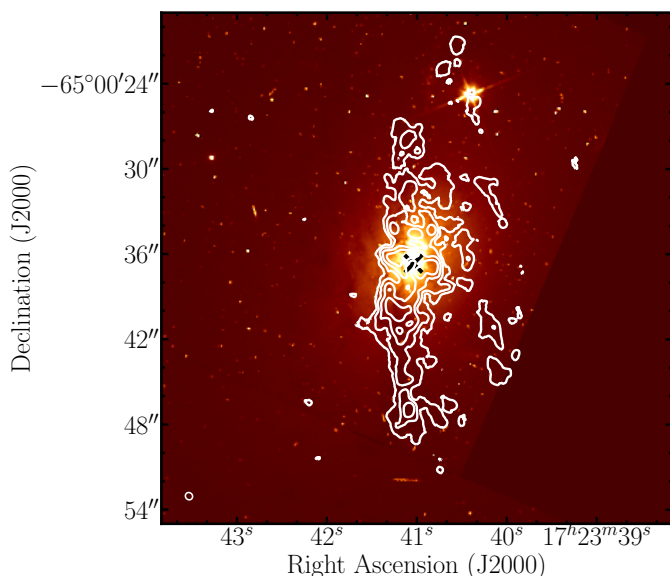


Fig. 2: Hubble Space Telescope WFPC2 image of PKS B1718–649 with overlaid the contours of the total intensity map of the ^{12}CO (2–1) line detected in the innermost 15 kpc. Contour levels are 3, 6, 12 and 18σ . This and all intensity maps shown in this paper are not corrected for the primary beam of ALMA. The 230 GHz continuum emission of the radio AGN is indicated with a black cross. The resolution of the total intensity map is $0.5''$. The beam is shown in white in the bottom left corner.

3.1. Molecular clouds rotating in a warped disk

The ALMA observations of PKS B1718–649 show that the innermost 15 kpc are rich in molecular gas distributed in clouds and filamentary structures. Figure 2 displays the total intensity map of the ^{12}CO (2–1) line emission (extracted from a data cube with $0.5''$ of spatial resolution) overlaid on the optical image taken with the Hubble Space Telescope WFPC2 camera. Overall, the ^{12}CO (2–1) emission is located along the dust lane

visible in the optical band and oriented north-south, and in the circumnuclear regions, close to the radio AGN. The distribution of the carbon monoxide is clumpy, with many molecular clouds spatially resolved in the field of view. From the total intensity map, corrected for the primary beam of ALMA, we measure the total flux of the ^{12}CO (2–1) line, $S_{\text{CO}}\Delta\nu = 174 \pm 17.4$ Jy km s $^{-1}$. We estimate that the error in the flux is $\sim 10\%$, to allow for uncertainties in the flux of the gain calibrator.

Figure 3 (left panel) shows the velocity field of the ^{12}CO (2–1) line emission extracted from the same data cube as the total intensity map. Overall, the kinematics of the gas is dominated by rotation. The observed velocity ranges between ~ 3950 km s $^{-1}$ and ~ 4600 km s $^{-1}$ along the north-south axis. We measure the systemic velocity of the ^{12}CO (2–1) as the mean point of this range of velocities, 4250 ± 20 km s $^{-1}$. This velocity is comparable to the systemic velocity measured from the kinematics of the H I gas ($v_{\text{sys}} = 4274 \pm 7$ km s $^{-1}$). The position angle of the disk and the direction of rotation is the same as the H I at outer radii (between 3 and 8 kpc) suggesting that overall the ^{12}CO (2–1) is following the rotation of the other components of the galaxy.

The observations of the atomic and molecular hydrogen suggested that the cold gas of PKS B1718–649 is distributed a warped structure. Here, we analyse a number of different signatures indicating that this is the case. The left panel of Fig 3 shows the velocity field of the ^{12}CO (2–1) gas seen by ALMA. Moving towards the centre, the kinematic major axis changes orientation from north-south in the outer regions to almost east-west in the inner part of the gas distribution. The figure suggests that the carbon monoxide in the circumnuclear regions and in the dust lane are connected in velocity and are distributed in a warped disk. The right panel of Fig. 3 shows in red contours the ^{12}CO (2–1) emission detected at 4440 km s $^{-1}$ and in blue contours the emission detected at 4100 km s $^{-1}$, extracted from a data-cube at the spatial resolution of $1''$. These two channels have opposite velocities with respect to the systemic velocity of the galaxy and show a good symmetry in the distribution of the ^{12}CO (2–1) emission. This panel shows that the velocity gradient in the central regions is perpendicular to that of the larger disk. This feature is very similar to that found in the prototypical warped galaxy NGC 3718 (Sparke et al. 2009), even though over larger spatial scales. In this galaxy, Sparke et al. (2009) showed that the H I gas is distributed in a single extremely warped disk, which varies its position angle while its inclination remains fairly edge-on. Likely, also the molecular gas of PKS B1718–649 is distributed in a similar warped disk.

Even though the kinematics of the molecular gas of PKS B1718–649 are consistent with a warped disk, bars in the stellar distribution can induce non-circular motions which can mimic the kinematics of a warped disk. We investigate whether a bar could be present, studying the light distribution seen in our SINFONI data (Maccagni et al. 2016). At radii larger than a few kilo-parsecs, the light distribution is well modelled using elliptical isophotes (ellipticity $\epsilon = 1 - b/a = 0.18$ and position angle, $PA = 150^\circ$) and a deVaucouleur's law (Veron-Cetty et al. 1995). We subtract this model from the light distribution seen in the innermost regions of PKS B1718–649. The residuals (Fig. 4) show that no signature of a bar is found. The figure also shows that the only deviations from the model are found in the centre ($r \lesssim 300$ pc), and suggests that a small spherical component may be present there.

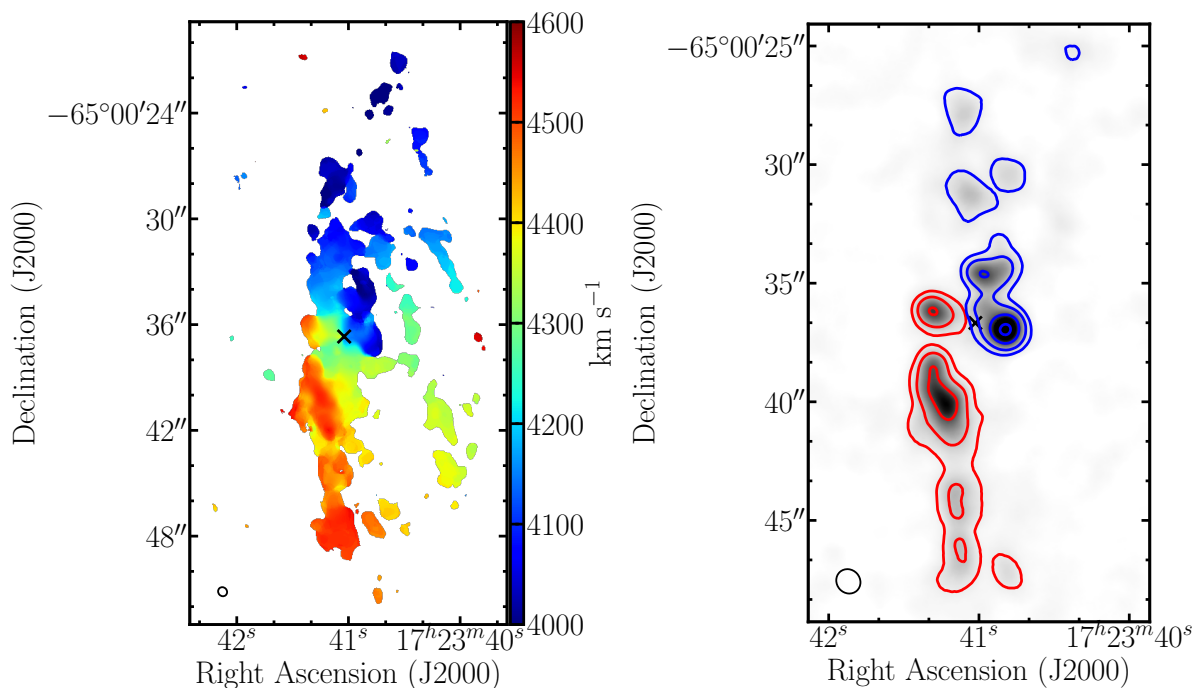


Fig. 3: *Left Panel*: Velocity field of the ^{12}CO (2–1) line detected in the innermost 15 kpc of PKS B1718–649. The systemic velocity ($\sim 4250 \pm 20 \text{ km s}^{-1}$) has green colours. The radio continuum emission is indicated by a black cross. The beam is shown in black in the bottom left corner. *Right Panel*: Superposed channel maps for the ^{12}CO (2–1) gas at velocities displaced roughly 190 km s^{-1} on either side of the systemic velocity. Upper contours, in blue, show gas centred at 4086 km s^{-1} while lower contours, in red, show gas centred at 4444 km s^{-1} .

3.2. A clumpy circumnuclear disk

Figure 5 shows the total intensity map of the ^{12}CO (2–1) emission line in the circumnuclear regions of PKS B1718–649. The rotating structure is oriented along the east-west direction ($\text{PA} = 72^\circ$), it has a radius of 700 pc. The radio source is almost orthogonal to the position angle of the major axis of the disk, $\text{PA} = 135^\circ$ (Tingay et al. 1997, 2002). The circumnuclear disk shows a clumpy structure embedded in a more diffuse medium. Within the disk, we resolve spatially and in velocity a few clouds that have approximate sizes of $r \lesssim 120 \text{ pc}$. In the very centre, where the radio AGN is located (black cross in the figure), we do not detect ^{12}CO (2–1) in emission at the sensitivity of our observations. This implies a 3σ upper limit to the column density of the molecular hydrogen of $N_{\text{H}_2} \sim 4.8 \times 10^{21} \text{ cm}^{-2}$, for a Galactic conversion factor (see Sect. 3.4).

The total flux density (corrected for the primary beam of ALMA) of the circumnuclear disk is $S_{\text{CO}}\Delta\nu = 29.5 \pm 2.95 \text{ Jy km s}^{-1}$ (see Table 3).

Figure 6 shows that the circumnuclear disk is dominated by rotation. The velocity field, as well as the single channel maps, suggest the disk is approximately edge-on ($i \sim 60^\circ - 80^\circ$).

Figure 7 shows a position-velocity diagram taken along the major axis of the circumnuclear disk. The smooth velocity gradient along the x-axis clearly shows the disk is overall regularly rotating and has a radius of 700 pc. In the innermost $\sim 300 \text{ pc}$, the gas in the disk spans a broader range of velocities than at larger radii. This may suggest an increase in the velocity dispersion of the gas in the central regions, even if part of this broadening of the velocity gradient can be explained by the geometry of the system, because the disk is observed edge-on.

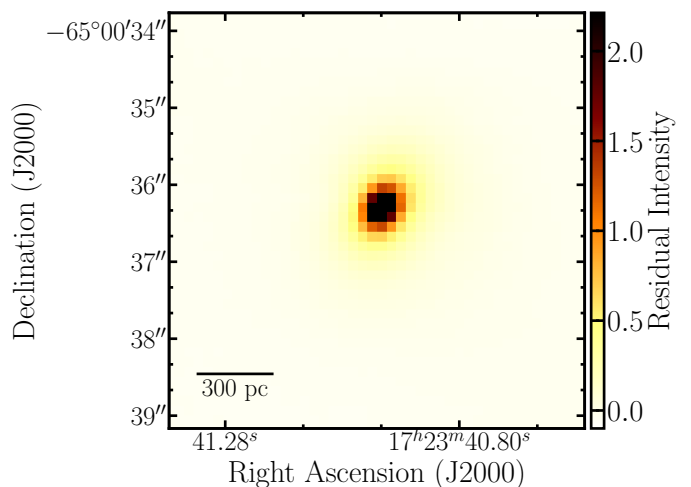


Fig. 4: Map of the residuals between the light distribution of the stellar body, seen by SINFONI (Maccagni et al. 2016) and a model with isophotes of constant ellipticity ($\epsilon = 1 - b/a = 0.18$) and position angle ($\text{PA} = 150^\circ$). No indications of the presence of a bar are found, the only significant deviations from an elliptical light distribution are found in the innermost 300 pc, where a spherical component may be present.

3.3. The ^{12}CO (2–1) detected in absorption

Figure 7 shows that the ^{12}CO (2–1) is detected in absorption against the radio AGN, at $v = +4615 \pm 20 \text{ km s}^{-1}$ (in black dashed contours). To investigate the line profile in more detail,

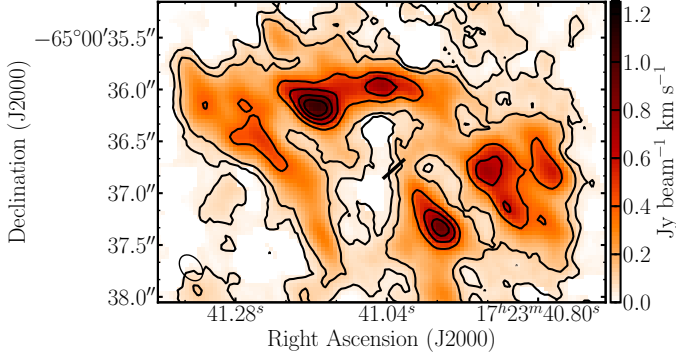


Fig. 5: Total intensity map of the ^{12}CO (2–1) line detected in the innermost 1.5 kpc of PKS B1718–649. Contour levels are 3, 5, 8, 12, 15 and 18σ ($1\sigma = 0.14 \text{ Jy beam}^{-1} \text{ km s}^{-1}$). The black line shows the position angle of the radio AGN. The beam ($0.2''$) is shown in black in the bottom left corner.

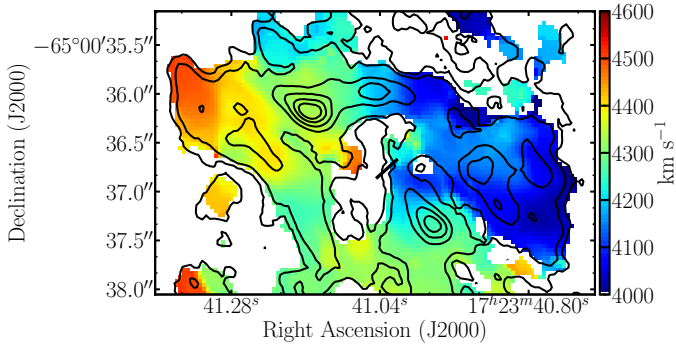


Fig. 6: Velocity field of the ^{12}CO (2–1) line detected in the circumnuclear regions of PKS B1718–649. Green velocities mark approximately the systemic velocity of the galaxy ($v_{\text{sys}} = 4250 \pm 20 \text{ km s}^{-1}$). Contour levels are as in Fig. 5. The black line shows the position angle of the radio AGN. The beam ($0.2''$) is shown in black in the bottom left corner.

in Fig. 8 we show the spectrum extracted against the radio continuum emission from a data cube at a higher spectral resolution (20 km s^{-1}). The noise of the spectrum is $1.12 \text{ mJy beam}^{-1}$. The peak of the absorption line ($S_{\text{peak}} = -0.813 \text{ mJy}$) is detected with $\sim 7\sigma$ significance. The total absorbed flux is $2.59 \text{ mJy km s}^{-1}$. Given that the flux density of the continuum at 230 GHz is 304 mJy , the line has peak optical depth $\tau_{\text{peak}} = 0.003$ and integrated optical depth $\int \tau dv = 0.221 \text{ km s}^{-1}$. Our observations spectrally resolve the absorption line and we measure a Full Width at Half Maximum, $\text{FWHM} = 57.4 \pm 20 \text{ km s}^{-1}$.

In the centre, if the ^{12}CO (2–1) was regularly rotating within the disk, we would have detected it at the systemic velocity, given the small size of the radio source. Instead, the absorption line peaks at $v_{\text{peak}} = 4615 \pm 20 \text{ km s}^{-1}$ and has an offset with respect to the systemic velocity of $\Delta v = v_{\text{peak}} - v_{\text{sys}} = +365 \pm 22 \text{ km s}^{-1}$ (see Table 2). Since the gas is detected in absorption it is located in front of the radio source, it is falling towards it and the in-fall velocity is given by the shift of the line $v_{\text{infall}} \lesssim \Delta v = +365 \pm 22 \text{ km s}^{-1}$.

If we assume that the gas detected in absorption is homogeneously distributed in front of the background continuum source,

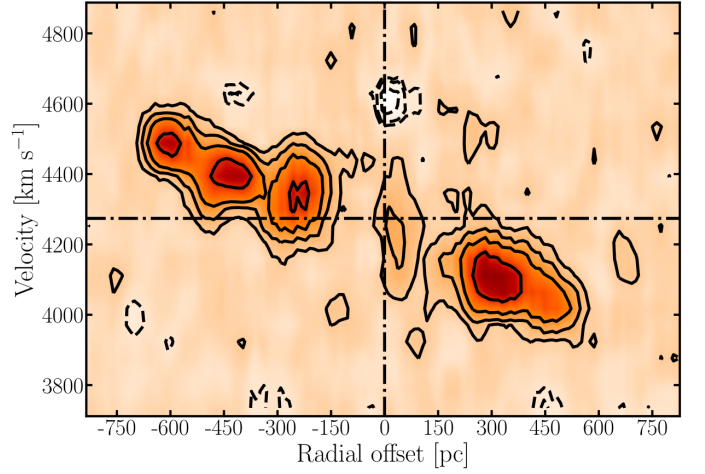


Fig. 7: Position-velocity diagram taken along the major axis of the circumnuclear disk of PKS B1718–649 ($\text{PA} = 72^\circ$), with a slit of $0.2''$. The dashed horizontal line shows the systemic velocity of the source. The vertical dashed line shows the position of the radio continuum sources. Contour levels are $-5, -3, -2\sigma$, in white, and $3, 6, 12, 18, 24\sigma$, in black. Absorption is detected with $\sim 7\sigma$ significance at $v \sim 4615 \text{ km s}^{-1}$ and zero offset.

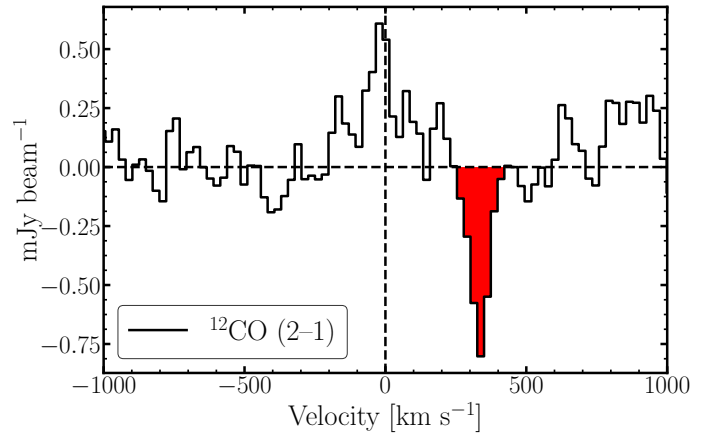


Fig. 8: Spectrum extracted against the radio continuum emission of PKS B1718–649. ^{12}CO (2–1) is detected in absorption at red-shifted velocities with respect to the systemic ($v = +365 \pm 20 \text{ km s}^{-1}$). The peak of the line is detected with $\sim 7\sigma$ significance. The red shaded regions show the integrated flux of the line.

and that it is in local thermal equilibrium (LTE), we can determine the column density of the carbon monoxide in the level $J = 2$ ($N_{\text{CO}(J)}$) from the integrated optical depth of the absorption line ($\int \tau dv$) using equation (e.g. Wiklind & Combes 1995; Bolatto et al. 2013):

$$N_{\text{CO}(J)} = \frac{3h}{8\pi^3 \mu^2} \frac{g_J}{J \cdot \left(e^{\frac{h\nu_{\text{rest}}}{k_B T_{\text{ex}}}} - 1 \right)} \cdot \int \tau_{\text{CO}(J)} dv \quad (1)$$

where $\mu = 0.112 \text{ cm}^{5/2} \text{ g}^{1/2} \text{ s}^{-1}$ is the dipole moment of the CO molecule, $J = 2$ is the upper level of the $J \rightarrow J - 1$ transition, $g_J = 2J + 1$ is the statistical weight of the transition,

Table 2: Properties of the ^{12}CO (2–1) detected in absorption

Parameter	Value
$S_{230\text{ GHz}}$	$304 \pm 30.4\text{ mJy}$
S_{abs}	$2.6 \pm 0.3\text{ mJy}$
S_{peak}	$-0.8 \pm 0.1\text{ mJy}$
FWHM	$57 \pm 20\text{ km s}^{-1}$
$\Delta v = v_{\text{peak}} - v_{\text{sys}}$	$345 \pm 20\text{ km s}^{-1}$
Peak optical depth	0.003
Integrated optical depth	0.221 km s^{-1}
N_{CO}	$\gtrsim 1.0 \times 10^{17}\text{ cm}^{-2}$
N_{H_2}	$\gtrsim 1.6 \times 10^{21}\text{ cm}^{-2}$
$M(\text{H}_2) (r = 2 - 75\text{ pc})$	$3 \times 10^2 M_{\odot} - 5 \times 10^5 M_{\odot}$

T_{ex} its excitation temperature and v_{rest} its rest frequency, h and k_B are the Planck's and Boltzmann constants, respectively. To determine the column density of the carbon monoxide we need to assume its excitation temperature.

In the Milky Way, in conditions of thermal equilibrium the typical temperature of the CO gas is $\sim 16\text{ K}$ (e.g. Heyer et al. 2009). Nevertheless, close to the nuclear activity, the gas could be warmer, which would increase the column density. For example, in the circumnuclear regions of 4C +12.50 the gas has likely temperature $T_{\text{ex}} \sim 60\text{ K}$ (Dasyra & Combes 2012; Dasyra et al. 2014) which would increase the column density of the ^{12}CO (2–1). Here, for simplicity, we assume the ^{12}CO (2–1) is in thermal equilibrium at $T_{\text{ex}} \sim 16$ and the column densities we derive are lower limits, if the temperature of the gas was higher.

The column density of the ^{12}CO (2–1) gas is $N_{\text{CO}(2)} = 9 \times 10^{14}\text{ cm}^{-2}$, respectively. To derive the column density of the total CO gas, we multiply $N_{\text{CO}(2)}$ by the partition function of the 2–1 rotational transition of the CO gas, $\sum_J (2J+1)e^{-E_J/k_B T} = \frac{k_B T_{\text{ex}}}{B}$ (where B is the rotational constant of the molecule expressed in cm^{-1}). We divide the result by the relative population of the upper level of the transition, $g_J \exp(-E_u/k_B T_{\text{ex}})$ (where $E_u = 11.5\text{ cm}^{-1}$), obtaining $N_{\text{CO}} = 1.0 \times 10^{17}\text{ cm}^{-2}$. To convert it to the column density of the molecular hydrogen we assume a CO-to- H_2 abundance ratio, $A_C = 1.6 \times 10^{-4}$ (Sofia et al. 2004), and we find $N_{\text{H}_2} = 1.6 \times 10^{21}\text{ cm}^{-2}$. This result is compatible with the 3σ detection limit of the ^{12}CO (2–1) line emission on the same line of sight ($N_{\text{H}_2} \sim 4.8 \times 10^{21}$, see Sect. 3.2).

The width of the absorption line also suggests that the absorbing ^{12}CO (2–1) is part of a large and complex ensemble of molecular clouds. In the spectrum at the original channel width of the observations (10.3 km s^{-1}), we resolve the absorption line over five channels. The width of the absorption line is much higher than the typical velocity dispersion of the molecular gas in a cloud ($\lesssim 4\text{ km s}^{-1}$, e.g. Heyer et al. 2009). According to the scaling relations for molecular clouds in the normal ISM (Larson 1981; Solomon et al. 1987; Heyer et al. 2009), the size of a cloud corresponding to the width of the absorption line would be $\sim 1\text{ kpc}$, which is the extent of the entire circumnuclear disk. Hence, the absorption line is likely tracing an ensemble of transient clouds, that are not in equilibrium with the surrounding medium, have non-circular kinematics, and are being shredded and torn apart while moving towards the SMBH (see Sect. 4.2 for further details).

We estimate the range of possible masses of the absorbed in-falling molecular gas assuming that the gas could extend only exactly over the background radio continuum emission ($r \sim 2\text{ pc}$) or that it could fill the entire region where ^{12}CO (2–1) emission is not detected ($r \sim 75\text{ pc}$). Given these two extreme pos-

sible cases, for a column density of $\sim 1.6 \times 10^{21}\text{ cm}^{-2}$ (see Table 2), the mass of the in-falling molecular gas ranges between $3 \times 10^2 M_{\odot}$ and $5 \times 10^5 M_{\odot}$.

3.4. Luminosities and masses of the molecular gas

The emission of the ^{12}CO (2–1) line allows us to determine the molecular masses of the different structures we identified in the previous sections, i.e. the mass derived from the total flux density of the ^{12}CO (2–1) line in PKS B1718–649 and from the flux of the circumnuclear disk (see Table 3).

To convert the flux of the ^{12}CO (2–1) line to mass we need to assume the ratio between the luminosity of the observed line and the ^{12}CO (1–0) line. Optically thick gas in thermal equilibrium has brightness temperature and line luminosity independent of the rotational energy level J of the line, $L'_{\text{CO}(2-1)} = L'_{\text{CO}(1-0)}$ (in units of $\text{K km s}^{-1}\text{ pc}^2$). If the gas was optically thin, the ratio could be higher than one. Here, we are not able to set constraints on the optical thickness of the ^{12}CO (2–1) (observations of different excitation levels, or of different isotopologues of the CO, would be needed) and, for simplicity, we assume $L'_{\text{CO}(2-1)} = L'_{\text{CO}(1-0)}$. If instead the molecular gas was sub-thermally excited, the values we derive for the luminosity of the total ^{12}CO would be lower limits.

According to Solomon & Bout 2005; Bolatto et al. 2013 (and references therein), the luminosity of the CO line is given by:

$$L'_{\text{CO}} [\text{K km s}^{-1}\text{ pc}^2] = 3.25 \times 10^7 S_{\text{CO}} \Delta v [\text{Jy km s}^{-1}] \times \\ \times \nu_{\text{obs}}^{-2} [\text{GHz}] \times \frac{D_L^2 [\text{Mpc}]}{(1+z)^3}$$

where ν_{obs} is the frequency at which the line is observed, and D_L is the luminosity distance of the source in Mpc. The total luminosity of the ^{12}CO (2–1) emission in PKS B1718–649 is $L'_{\text{CO}} = 7.4 \times 10^7\text{ K km s}^{-1}\text{ pc}^2$ and the circumnuclear disk has luminosity of $L'_{\text{CO}} = 1.2 \times 10^7\text{ K km s}^{-1}\text{ pc}^2$.

From these we determine the total mass of the molecular hydrogen traced by the ^{12}CO (2–1) line and the molecular hydrogen mass of the circumnuclear disk. The H_2 mass-to-CO luminosity relation can be expressed as $M_{\text{H}_2} = \alpha_{\text{CO}} L'_{\text{CO}}$ (Solomon et al. 1997; Downes & Solomon 1998; Solomon & Bout 2005; Dale et al. 2005; Bolatto et al. 2013). In the Milky Way, most of the CO is detected in optically thick molecular clouds and $\alpha_{\text{CO}} = 4.6 M_{\odot} \cdot (\text{K km s}^{-1}\text{ pc}^2)^{-1}$. If the temperature of the gas was higher, e.g. in Ultra Luminous InfraRed Galaxies (ULIRGs) and AGN, the conversion factor could be as low as $\alpha_{\text{CO}} = 0.8$ (Downes & Solomon 1998; Bolatto et al. 2013; Geach et al. 2014). In Table 3, we show that the total mass of the H_2 varies between $3.26 \times 10^8 M_{\odot}$ for ($\alpha_{\text{CO}} = 0.8 M_{\odot} \cdot (\text{K km s}^{-1}\text{ pc}^2)^{-1}$), to $1.88 \times 10^9 M_{\odot}$ for a Galactic conversion factor ($\alpha_{\text{CO}} = 4.6 M_{\odot} \cdot (\text{K km s}^{-1}\text{ pc}^2)^{-1}$). The mass of the circumnuclear disk ranges between $5.35 \times 10^7 M_{\odot}$ and $3.08 \times 10^8 M_{\odot}$.

4. Discussion

4.1. The ^{12}CO (2–1) and the H_2 in PKS B1718–649

In this section, we compare the SINFONI observations of the warm H_2 (Maccagni et al. 2016), with the ALMA observations of the ^{12}CO (2–1). Figure 9 (left panel) shows the ^{12}CO (2–1) emission of the circumnuclear disk overlaid with the H_2 1-0 S(1) emission detected in the innermost 8 kpc of PKS B1718–649.

Table 3: Properties of the ^{12}CO (2–1) detected in emission

Parameter		Value
$S_{\text{CO}(2-1)}$	total	$174 \pm 17 \text{ Jy km s}^{-1}$
$L'_{\text{CO}(2-1)}$		$7.3 \times 10^7 \text{ K km s}^{-1} \text{ pc}^2$
$M(\text{H}_2)$	$\alpha_{\text{CO}} = 0.8 M_{\odot} \cdot (\text{K km s}^{-1} \text{ pc}^{-2})^{-1}$	$3.2 \times 10^8 M_{\odot}$
	$\alpha_{\text{CO}} = 4.6 M_{\odot} \cdot (\text{K km s}^{-1} \text{ pc}^{-2})^{-1}$	$1.9 \times 10^9 M_{\odot}$
$S_{\text{CO}(2-1)}$	circumnuclear disk	$29 \pm 3 \text{ Jy km s}^{-1}$
$L'_{\text{CO}(2-1)}$		$1.2 \times 10^7 \text{ K km s}^{-1} \text{ pc}^2$
$M(\text{H}_2)$	$\alpha_{\text{CO}} = 0.8 M_{\odot} \cdot (\text{K km s}^{-1} \text{ pc}^{-2})^{-1}$	$5.3 \times 10^7 M_{\odot}$
	$\alpha_{\text{CO}} = 4.6 M_{\odot} \cdot (\text{K km s}^{-1} \text{ pc}^{-2})^{-1}$	$3.1 \times 10^8 M_{\odot}$

Notes. We estimate the total H_2 mass assuming two different conversion factors between the luminosity of the ^{12}CO (2–1) line and the mass of the molecular gas. The lower values of the mass are obtained adopting the typical value for the optically thick gas in ULIRG and AGN ($\alpha_{\text{CO}} = 0.8 M_{\odot} \cdot (\text{K km s}^{-1} \text{ pc}^{-2})^{-1}$) and the upper limit are given for a Galactic conversion factor ($\alpha_{\text{CO}} = 4.6 M_{\odot} \cdot (\text{K km s}^{-1} \text{ pc}^{-2})^{-1}$).

The two lines trace the H_2 in its cold ($T \sim 10 - 100 \text{ K}$) and warm ($T \sim 10^3 \text{ K}$) phase, respectively. The spatial resolution of the SINFONI observations ($0.55''$) is lower than the ALMA observations. The figure shows that, overall, the warm H_2 and the ^{12}CO (2–1) gas have a similar spatial distribution within the central region, a circumnuclear disk with radius of 700 pc oriented east-west and an external disk oriented north-south. The SINFONI observations detected H_2 gas distributed in two orthogonal disks, with possibly gas between the two structures. The ALMA observations trace ^{12}CO (2–1) gas connecting the two disks, which appear to be part of the same structure.

In the outer disk ($r \gtrsim 1 \text{ kpc}$), the H_2 is detected where the ^{12}CO (2–1) is brightest. Fig. 9 (right panel) shows that this occurs also along the circumnuclear disk except for the innermost 75 pc. There, the H_2 emission is brightest while the ^{12}CO (2–1) emission is dimmest. In particular, within 75 pc from the SMBH, the ratio between the column density of the H_2 traced by the H_2 1–0 S(1) line and by the ^{12}CO (2–1) line is ten times higher than in the circumnuclear disk ($N_{\text{H}_2}(\text{H}_2 1-0 \text{ S}(1))/N_{\text{H}_2}(\text{CO}(2-1)) = 110$ and 16, respectively). Since the ^{12}CO (2–1) and the H_2 1–0 S(1) are both tracers of the molecular hydrogen, this suggests that in the centre the conditions of the gas are different than at larger radii. In the centre of PKS B1718–649, the ^{12}CO (2–1) line may not be tracing all molecular hydrogen present, but only its coldest component. This can happen when, in the presence of a strong UV-radiative field, but also of X-rays and γ -rays, the CO gas is photo-dissociated while the H_2 gas is heated (e.g. Lamarche et al. 2017). In PKS B1718–649, a radiation field with this properties could have been recently originated by the newly born nuclear activity and indications of its presence could be found in the excess of radiation in the X-rays and γ -rays detected in this galaxy (Siemiginowska et al. 2016; Migliori et al. 2016). If this was the case, we would be observing the radio nuclear activity of a newly born AGN actively changing the physical conditions of the surrounding ISM.

PKS B1718–649 is not the only radio AGN where the nuclear radiation may be changing the conditions of the surrounding ISM, and where cold gas clouds of molecular gas may fall onto the SMBH. For example, Centaurus A is a very well studied low efficiency accretion radio AGN where the last episode of nuclear activity was recently triggered ($t_{\text{radio}} \sim 10^6 \text{ yr}$, see, e.g. Morganti 2010 for a review). Interestingly, recent ALMA observations of Centaurus A (Espada et al. 2017), reveal a clumpy circumnuclear disk of molecular gas in its innermost 500 pc. In the

centre of the disk, there is a lack of CO gas, while the warm H_2 is detected down to small radii ($r \lesssim 10 \text{ pc}$). The results we found for PKS B1718–649 indicate that this may be due to different conditions of the molecular gas rather than a complete depletion of CO gas in the very central regions. Moreover, the circumnuclear disk of Centaurus A may be strongly warped, as the one we detect in PKS B1718–649. In the innermost regions of the disk ($r \sim 10 - 50 \text{ pc}$) different tracers (CO, HCN, HCO^+ (4–3)) reveal cold clouds of molecular gas with non circular orbits.

Figure 10 shows the position-velocity diagram extracted along the major axis of the circumnuclear disk and integrated over a width of $0.5''$ in the direction perpendicular to the major axis. In Maccagni et al. (2016), we estimated a model of the rotation curve of the inner disk (shown by the white dashed line in Fig. 10) considering the contribution of the central SMBH and of the stellar mass distribution ($M_{\star} = 4 \times 10^{11} M_{\odot}$, effective radius $r_e = 9.7 \text{ kpc}$). The rotation curve matches the smooth velocity gradient of the gas, hence, overall, the warm H_2 and the ^{12}CO (2–1) gas regularly rotate in the circumnuclear disk. Nevertheless, in the centre, there is gas that is deviating from the rotation at red-shifted velocities, as also seen in the SINFONI data (Maccagni et al. 2016).

Overall, the warm H_2 has kinematics characterised by regular rotation, except that in the central resolution element where warm H_2 emission is detected at the same red-shifted velocities of the ^{12}CO (2–1) gas detected in absorption ($\Delta v = +365 \pm 20 \text{ km s}^{-1}$, see Sect. 3.3). This is the only component of H_2 gas in the entire field of view with kinematics deviating from rotation, and it appears to be kinematically connected to the circumnuclear disk (see Fig. 10). This likely suggests that the H_2 gas with red-shifted velocities lies within the innermost 75 pc of the galaxy. Given that the ^{12}CO (2–1) absorption is found on the same line of sight of this H_2 gas and that the two components have the same red-shifted velocities, they are likely located in the same region. Hence, the upper limit on the projected-distance from the SMBH of the in-falling ^{12}CO (2–1) gas is $d_{\text{infall}} \lesssim 75 \pm 30 \text{ pc}$.

4.2. Molecular clouds fuelling the newly born radio source

Figure 11 shows the spectrum of the ^{12}CO (2–1) extracted against the central compact radio source ($r_{\text{radio}} \sim 2 \text{ pc}$, Sect. 3.3) overlaid with the spectrum of the H I (Maccagni et al. 2014) and of the warm H_2 gas (Maccagni et al. 2016) extracted along the

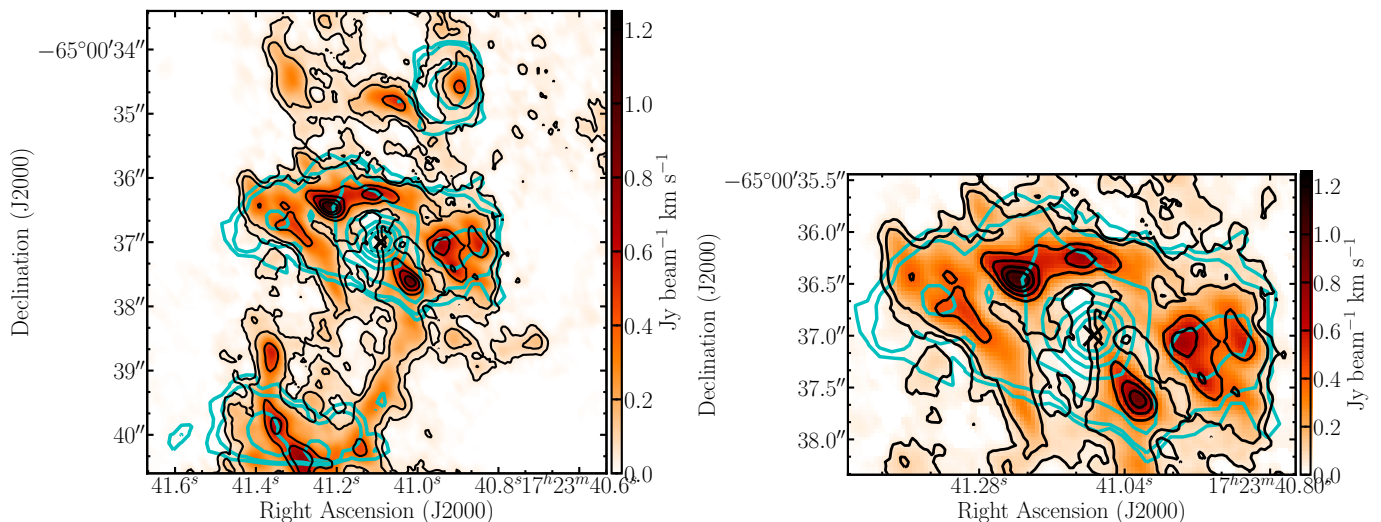


Fig. 9: *Left panel*: Total intensity image of the distribution of the ^{12}CO (2–1) gas in PKS B1718–649 (black contours) overlaid with the distribution of the H_2 detected by SINFONI, in cyan contours (Maccagni et al. 2016). The resolution of the ^{12}CO (2–1) data is $0.28''$, the resolution of the H_2 data is $0.55''$. *Right panel*: Total intensity map of the circumnuclear disk of ^{12}CO (2–1) gas of PKS B1718–649 overlaid with the warm H_2 emission. The radio continuum is shown by a black cross.

same line of sight. The spectrum of the H_2 1–0 S(1) gas is normalised to the peak of the ^{12}CO (2–1) emission line. The shaded regions of the figure highlight the parts of the lines tracing gas that is not regularly rotating within the circumnuclear disk. The two H I absorption lines cannot rotate with the disk because they trace gas that is on a line of sight of only 2 pc and has opposite velocities with respect to the systemic velocity ($v_{\text{peak,blue}} = -74 \text{ km s}^{-1}$ and $v_{\text{peak,red}} = +26 \text{ km s}^{-1}$, respectively). The warm H_2 gas has a component of the line with red-shifted velocities with respect to the systemic velocity ($v_{\text{peak,H}_2} = +4600 \text{ km s}^{-1}$) and it extends at the same velocities of the ^{12}CO (2–1) absorption line outside of the range of rotational velocities of the circumnuclear disk (see Sect. 4.1). The spectra show a first order similarity, despite the differences in velocity and spatial resolution (i.e. 6.7 km s^{-1} and $\sim 11''$ for the H I , 75 km s^{-1} and $\sim 0.55''$ for the H_2 and 24 km s^{-1} and $\sim 0.28''$ for the ^{12}CO (2–1)).

Figure 11 shows that clouds of cold gas are found in various tracers (H I , H_2 and ^{12}CO (2–1)) at different velocities with respect to the systemic velocity, along the same (very narrow) line of sight of the radio continuum. The clouds are embedded in a circumnuclear disk, and have kinematics with clear deviations from the regular rotation of the disk. This suggests that in the nuclear region of PKS B1718–649 there is a reservoir of clouds of cold gas with strong radial motions. Some of these clouds have red-shifted velocities and are located in front of the SMBH, suggesting they are likely falling onto it.

In Sect. 3.4, we provide a range of the mass of the absorbing molecular clouds, $3 \times 10^2 M_\odot - 5 \times 10^5 M_\odot$. Given that the infall velocity is $v_{\text{infall}} = +365 \pm 20 \text{ km s}^{-1}$ and assuming that all the absorbing ^{12}CO (2–1) gas is at the maximum distance of infall, $d_{\text{infall}} \sim 75 \text{ pc}$ (see Sect. 4.1), the time-scale of accretion is $t_{\text{accretion}} \sim 5 \times 10^5 \text{ years}$. The accretion rate onto the SMBH is $1.6 \times 10^{-3} M_\odot \text{ yr}^{-1} \lesssim \dot{M}_{\text{H}_2} \lesssim 2.2 M_\odot \text{ yr}^{-1}$. This range is broad because we consider two extreme approximations for the extent and column density of the in-falling clouds (see Sect. 3.3).

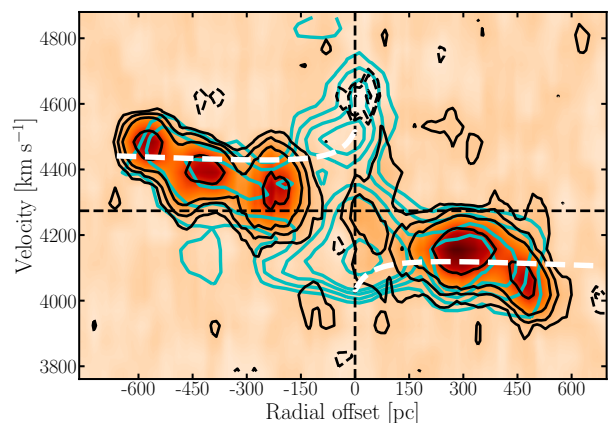


Fig. 10: Integrated position-velocity diagram taken along the major axis of the circumnuclear disk PKS B1718–649 (PA = 72°), integrated over $0.5''$ in the direction perpendicular to the major axis. Black solid and dashed contours show the ^{12}CO (2–1) emission and absorption, respectively. Cyan contours show the H_2 emission. Contour levels are $-5, -3, -2\sigma$ and $3, 6, 12, 18, 24\sigma$, respectively. The white dashed line shows the rotation curve predicted from the stellar distribution and the contribution of the SMBH to the rotation (see the text for further details).

4.3. Accretion of molecular gas in radio AGN

In the previous sections, we defined the main properties of the in-falling clouds of molecular hydrogen that we have detected in the innermost 75 pc of PKS B1718–649. Here, we relate these properties to the accretion mechanism that may have triggered the radio source and may still fuel it. In low-efficiency radio AGN, the energetic output of the nuclear activity provides an estimate of the accretion rate required to sustain the AGN. The

luminosity at 1.4 GHz allows us to measure the mechanical energy released by the radio jets (e.g. Cavagnolo et al. 2010), while the luminosity of the the [O III] line provides information on the radiative energy of the AGN (e.g. Allen et al. 2006; Balmaverde et al. 2008; Best & Heckman 2012).

In PKS B1718–649, the expected accretion rate is $\dot{M} \sim 10^{-2} M_{\odot} \text{ yr}^{-1}$ (see Maccagni et al. 2014). The range of accretion rates we determine $1.6 \times 10^{-3} M_{\odot} \text{ yr}^{-1} \lesssim \dot{M}_{\text{H}_2} \lesssim 2.2 M_{\odot} \text{ yr}^{-1}$ is compatible with the expected accretion rate of PKS B1718–649. Likely, the accretion of all atomic and molecular gas with strong radial motions in the innermost 75 pc of PKS B1718–649 (i.e. H I clouds detected in absorption, and molecular clouds traced by the H₂ 1-0 S(1) and ¹²CO (2–1) with red-shifted velocities) can fuel the nuclear activity. Even though the difference between the upper limit on the accretion of the ¹²CO (2–1) clouds and the expected accretion rate suggest that not all gas within the 75 pc is directly accreting onto the SMBH.

Chaotic cold accretion could be the triggering and fuelling mechanism of PKS B1718–649. Different numerical simulations of chaotic cold accretion (King & Nixon 2015; Gaspari et al. 2013b; Gaspari & Sadowski 2017; Gaspari et al. 2017a) show that small clouds of multi-phase gas condense out of the hot gas turbulent eddies of the halo of the host galaxy. Overall, this gas would not show kinematics with clear deviations from circular orbits at kiloparsec distance (as found for example in PKS 2322–12, Tremblay et al. 2016; NGC 5044, David et al. 2014; PKS 0745–191, Russell et al. 2016 and Centaurus A, Espada et al. 2017), while within a few Bondi radii, i.e. the radius of influence of the SMBH, the inelastic collisions between the clouds are so intense to cancel angular momentum so that accretion onto the SMBH may occur. This breaks the spherical symmetry assumed so far in models of low-efficiency accretion, and can boost the efficiency of accretion in radio AGN. Chaotic cold accretion is predicted to manifest through giant molecular associations of clouds (with size between 50 and 100 pc and mass up to several $10^5 M_{\odot}$), with strong non circular orbits. Gaspari et al. (2017a,b) predict that structures falling into radio AGN should appear as H I or CO absorption features.

The in-falling clouds we detect in the centre of PKS B1718–649, as well as the different physical conditions of the H I, H₂ 1-0 S(1) and ¹²CO (2–1) gas in the innermost 75 pc compared to the circumnuclear disk and the complex warped disk to which it is connected, all well match the properties of the accreting gas predicted by the models of chaotic cold accretion. Moreover, the maximum distance from the AGN where the in-falling gas could be located is 75 pc. This is less than twice the radius within which the SMBH dominates the kinematics of the galaxy ($r \sim 45$ pc, assuming that the velocity dispersion of the stars is $\sim 200 \text{ km s}^{-1}$ and $M_{\text{SMBH}} \sim 4 \times 10^8 M_{\odot}$, Willett et al. 2010). This is also in good agreement with the predictions of chaotic cold accretion models (e.g. Gaspari & Sadowski 2017; Gaspari et al. 2017a). Therefore, it is possible that chaotic cold accretion of molecular clouds triggered the radio nuclear activity of PKS B1718–649, and is currently fuelling it.

PKS B1718–649 could be an optically elusive AGN (Schawinski et al. 2015), i.e. the AGN is weak in the optical band (classified as LINER) because it is living an intermediate phase of nuclear activity, where accretion has already triggered the mechanism of radiation of radio and X-ray emission, but did not have enough time to sufficiently ionize the circumnuclear ISM to be bright in the optical. Given the age of the radio source, the radiation field of the nuclear activity could have influenced the ISM only within the innermost ~ 28 pc. Even though the information on the ionized gas surrounding the AGN is limited (e.g. Filip-

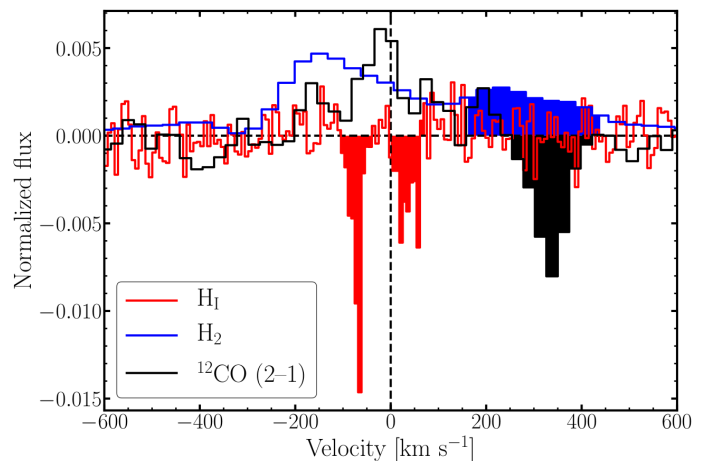


Fig. 11: Spectra of the H I (red), H₂ (blue) and ¹²CO (2–1) (black) gas extracted against the compact radio continuum emission of PKS B1718–649. Two separate lines are detected in the H I gas with opposite velocities with respect to the systemic. The ¹²CO (2–1) detected in absorption peaks at red-shifted velocities ($\Delta v \sim 343 \text{ km s}^{-1}$) with respect to the systemic (black dashed line). The H₂ gas is detected in emission at the systemic velocity and at the velocities of the red-shifted absorption line of ¹²CO (2–1) gas. The spectrum of the H₂ gas has been normalized to the maximum of the spectrum of the ¹²CO (2–1) line (see Sect. 3.3 for further details). The shaded regions of the lines are tracing gas that is not regularly rotating.

penko 1985), PKS B1718–649 is optically a LINER AGN, it has a thermal X-ray excess (Siemiginowska et al. 2016), and has a young powerful radio AGN carving its way through a dense gas, which is mostly molecular.

5. Conclusions

In this paper, we presented new ALMA observations of the ¹²CO (2–1) gas in the central 15 kpc around the young radio source PKS B1718–649. These observations follow the H I (Maccagni et al. 2014) and warm H₂ (Maccagni et al. 2016) observations that revealed cold gas in the centre of the galaxy that may contribute fuelling the radio AGN. The ALMA observations allow us to trace with high spatial resolution ($0.28''$) the distribution of the carbon monoxide, study its kinematics in relation to the recent triggering of the radio nuclear activity, and detect in-falling molecular clouds that may accrete onto the SMBH.

The ALMA observations of the ¹²CO (2–1) line reveal that multiple clouds are distributed in a complex warped disk. Between 7.5 kpc and 2.5 kpc from the central radio AGN, the disk is oriented as the dust lane of the galaxy in the north-south direction ($PA \sim 180^\circ$). At inner radii (2.5 kpc - 700 pc), the disk changes slightly in orientation ($PA \sim 150^\circ$) and molecular clouds are found in proximity of dust (Fig. 2) and of warm H₂ gas (see the left panel of Fig. 9). In the very central regions ($r \sim 700$ pc), the disk changes abruptly its orientation and the molecular clouds form an edge on circumnuclear disk oriented approximately east-west ($PA \sim 72^\circ$).

The carbon monoxide is a tracer of the cold phase of the molecular hydrogen. Depending on the conversion factor assumed, the total mass of cold molecular hydrogen is $M(\text{H}_2) 3.2 \times 10^8 - 1.9 \times 10^9 M_{\odot}$ (see Table 3), while the circumnu-

clear disk has total mass of cold molecular hydrogen between $5.3 \times 10^7 M_{\odot}$ and $3.1 \times 10^8 M_{\odot}$.

The circumnuclear disk of molecular gas of PKS B1718–649 may form the reservoir of fuel for the radio emission. Although the rotation of the gas follows what expected from the mass of the central SMBH and the distribution of the stellar body, the velocity dispersion of the gas is higher in the centre of the disk ($r < 150$ pc, $0.45''$) than at larger radii (Fig. 7). This suggests that close to the radio source the gas may have non circular orbits.

The circumnuclear disk is also traced by the H_2 1-0 S(1) line emission (Maccagni et al. 2016). The ratio between the column density of the H_2 gas traced by the H_2 1-0 S(1) line and by the ^{12}CO (2–1) line increases by a factor ten in the innermost 75 pc of the circumnuclear disk with respect to its outer regions (~ 110 and ~ 16 respectively). This suggests that, in the centre, the physical conditions of the molecular gas are different. It is possible that the radiative field of the AGN is photo-dissociating the cold component of the molecular gas, traced by the ^{12}CO (2–1) line, before the warmer component, traced by the H_2 1-0 S(1) line. This would indicate that the newly born AGN is actively changing the physical conditions of the surrounding ISM.

Against the 230 GHz emission of the radio AGN, we detect ^{12}CO (2–1) gas in absorption (see Sect. 3.3). The absorption line has a width (FWHM = 54 km s^{-1}) much larger than the typical velocity dispersion of molecular clouds. Hence, either the absorption is tracing multiple clouds, or a single cloud with non circular orbits, or a combination of both. The absorption line is detected at red-shifted velocities with respect to the systemic velocity ($\Delta v = +365 \pm 20 \text{ km s}^{-1}$). Its kinematics deviate from the rotation of the circumnuclear disk and the gas may be falling towards the radio source. At the same red-shifted velocities we also detect in emission a component of warm H_2 (see Fig. 10). This allows us to constrain the distance from the SMBH of the red-shifted ^{12}CO (2–1) to 75 pc. Likely, these clouds are accreting onto the SMBH. This is one of the best indications that accretion of cold molecular clouds can trigger and fuel a young radio source. These clouds may have originated from condensation of the hot gas present in the centre of PKS B1718–649, as suggested by the excess in soft X-ray emission (Siemiginowska et al. 2016). In Section 4.2, we suggest that the in-falling molecular clouds may trace chaotic cold accretion (Gaspari & Sadowski 2017; Gaspari et al. 2017a,b) onto the SMBH of PKS B1718–649.

PKS B1718–649 is the youngest radio source showing strong indications that cold gaseous clouds are accreting onto the SMBH. Indications that chaotic cold accretion can fuel an AGN have been found also in a handful of other sources. For example, in Centaurus A (Espada et al. 2017), NGC 5044 (David et al. 2014; Russell et al. 2016; David et al. 2017) and PKS 2322–12 (Tremblay et al. 2016) cold gas clouds with non circular orbits are found in proximity of the AGN. While in NGC 3115 (Wong et al. 2013) and M 87 (Russell et al. 2015) similar clouds are likely condensing from the hot circumnuclear environment.

The ALMA observations of only the ^{12}CO (2–1) transition do not allow us to constrain the physical conditions of the circumnuclear disk, such as temperature, pressure and heating mechanism of the gas. This would be a crucial step forward in the study of accretion of cold gas in low-efficiency radio AGN, and in understanding how the radio nuclear activity can change the conditions of the surrounding molecular gas. For these reasons, we planned to continue the study of the molecular gas in PKS B1718–649 and its contribution to fuel the newly born radio AGN with follow-up ALMA observations of the ^{12}CO (3–2)

and HCO^+ (4–3) at the same spatial and spectral resolution. We will determine where along the disk the molecules are heated by photo-dissociation and where by X-ray dissociation, and we will measure if there are gradients in pressure between the clouds of the circumnuclear disk. This will allow us to relate the properties of the molecular gas to the radio and X-ray properties of the AGN and, possibly, to set further constraints on the processes of accretion onto the SMBH of young radio sources.

Acknowledgements. ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada), NSC and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO and NAOJ. The National Radio Astronomy Observatory (NRAO), which is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. The research leading to these results has received funding from the European Research Council under the European Union's Seventh Framework Programme (FP/2007-2013) / ERC Advanced Grant RADIOLIFE-320745. The authors wish to thank Dr. Massimo Gaspari for the useful discussions and suggestions.

References

- Alexander, D. & Hickox, R. 2012, *New Astron. Rev.*, 56, 93
 Allen, S. W., Dunn, R. J. H., Fabian, A. C., Taylor, G. B., & Reynolds, C. S. 2006, *MNRAS*, 372, 21
 Allison, J. R., Curran, S. J., Emonts, B. H., et al. 2012, *MNRAS*, 423, 2601
 Balmaverde, B., Baldi, R. D., & Capetti, A. 2008, *A&A*, 486, 119
 Best, P. N. & Heckman, T. M. 2012, *MNRAS*, 421, 1569
 Best, P. N., Kauffmann, G., Heckman, T. M., & Ivezić, Ž. 2005, *MNRAS*, 362, 9
 Bolatto, A. D., Wolfire, M., & Leroy, A. K. 2013, *ARA&A*, 51, 207
 Bondi, H. 1952, *MNRAS*, 112, 195
 Booth, C. M. & Schaye, J. 2009, *MNRAS*, 398, 53
 Bower, R. G., Benson, A. J., Malbon, R., et al. 2006, *MNRAS*, 370, 645
 Brienza, M., Godfrey, L., Morganti, R., et al. 2015, *A&A*, 585, A29
 Briggs, D. S. 1995, PhD Thesis, New Mex. Inst. Min. Tech
 Cavagnolo, K. W., McNamara, B. R., Nulsen, P. E. J., et al. 2010, *ApJ*, 720, 1066
 Ciotti, L., Ostriker, J. P., & Proga, D. 2010, *ApJ*, 717, 708
 Clarke, D. A. & Burns, J. O. 1991, *ApJ*, 369, 308
 Combes, F., García-Burillo, S., Casasola, V., et al. 2013, *A&A*, 558, A124
 Combes, F., García-Burillo, S., Casasola, V., et al. 2014, *A&A*, 565, A97
 Comerford, J. M., Gerke, B. F., Newman, J. A., et al. 2008, *ApJ*, 698, 956
 Croton, D. J., Springel, V., White, S. D., et al. 2006, *MNRAS*, 365, 11
 Dale, D. A., Sheth, K., Helou, G., Regan, M. W., & Hüttemeister, S. 2005, *AJ*, 129, 2197
 Dasyra, K. M. & Combes, F. 2012, *A&A*, 541, L7
 Dasyra, K. M., Combes, F., Novak, G. S., et al. 2014, *A&A*, 565, A46
 Dasyra, K. M., Combes, F., Oosterloo, T., et al. 2016, *A&A*, 595, L7
 David, L. P., Lim, J., Forman, W., et al. 2014, *ApJ*, 792
 David, L. P., Vrtillek, J., O'Sullivan, E., et al. 2017, *ApJ*, 842, 84
 Davies, R., Mark, D., & Sternberg, A. 2012, *A&A*, 537, A133
 Davies, R. I., Maciejewski, W., Hicks, E. K. S., et al. 2014, *ApJ*, 792, 101
 Davies, R. I., Maciejewski, W., Hicks, E. K. S., et al. 2009, *ApJ*, 702, 114
 Debuhr, J., Quataert, E., & Ma, C. P. 2012, *MNRAS*, 420, 2221
 Downes, D. & Solomon, P. M. 1998, *ApJ*, 507, 615
 Emonts, B. H., Morganti, R., Struve, C., et al. 2010, *MNRAS*, 406, 987
 Espada, D., Matsushita, S., Miura, R. E., et al. 2017, *ApJ*, 843, 136
 Fabian, A. C. 2012, *ARA&A*, 50, 455
 Faucher-Giguère, C. A. & Quataert, E. 2012, *MNRAS*, 425, 605
 Filippenko, A. V. 1985, *ApJ*, 289, 475
 García-Burillo, S., Combes, F., Almeida, C. R., et al. 2016, *ApJL*, 823, 1
 García-Burillo, S., Combes, F., Usero, A., et al. 2014, *A&A*, 567, A125
 Gaspari, M., McDonald, M., Hamer, S. L., et al. 2017a
 Gaspari, M., Ruszkowski, M., & Oh, S. P. 2013a, *MNRAS*, 432, 3401
 Gaspari, M., Ruszkowski, M., & Oh, S. P. 2013b, *MNRAS*, 432, 3401
 Gaspari, M. & Sadowski, A. 2017, *ApJ*, 837, 149
 Gaspari, M., Temi, P., & Brighenti, F. 2017b, *MNRAS*, 466, 677
 Geach, J. E., Hickox, R. C., Diamond-Stanic, A. M., et al. 2014, *Nature*, 516, 68
 Geréb, K., Maccagni, F. M., Morganti, R., & Oosterloo, T. A. 2015, *A&A*, 575, A44
 Giroletti, M. & Polatidis, A. 2009, *Astron. Nachr.*, 330, 193
 Guillard, P., Ogle, P. M., Emonts, B. H. C., et al. 2012, *ApJ*, 747, 95
 Hardcastle, M. J., Evans, D. A., & Croston, J. H. 2007, *MNRAS*, 376, 1849
 Heckman, T. M. & Best, P. N. 2014, *ARA&A*, 52, 589
 Hernquist, L. & Spitzer, D. N. 1992, *ApJ*, 399, L117
 Heyer, M., Krawczyk, C., Duval, J., & Jackson, J. M. 2009, *ApJ*, 699, 1092
 Hillel, S. & Soker, N. 2013, *MNRAS*, 430, 1970

- Ho, L. C. 2009, *ApJ*, 699, 626
- Hopkins, P. F. & Elvis, M. 2010, *MNRAS*, 401, 7
- Jones, D. L. & Preston, R. A. 2001, *AJ*, 122, 2940
- King, A. & Nixon, C. 2015, *MNRAS*, 453, L46
- King, A. & Pounds, K. 2015, *ARA&A*, 53, 115
- Kormendy, J. & Kennicutt, R. C. 2004, *ARA&A*, 42, 603
- Labiano, A., García-Burillo, S., Combes, F., et al. 2013a, *A&A*, 549, A58
- Labiano, A., García-Burillo, S., Combes, F., et al. 2013b, *A&A*, 549, A58
- Lamarche, C., Stacey, G., Brisbin, D., et al. 2017, *ApJ*, 836, 123
- Larson, R. B. 1981, *MNRAS*, 194, 809
- Maccagni, F. M., Morganti, R., Oosterloo, T. A., Geréb, K., & Maddox, N. 2017, *A&A*, 604, A43
- Maccagni, F. M., Morganti, R., Oosterloo, T. A., & Mahony, E. K. 2014, *A&A*, 571, A67
- Maccagni, F. M., Santoro, F., Morganti, R., et al. 2016, *A&A*, 588, A46
- Martín, S., Kohno, K., Izumi, T., et al. 2015, *A&A*, 573, A116
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, *Astron. Data Anal. Softw. Syst. XVI*, 376, 127
- Migliori, G., Siemiginowska, A., Sobolewska, M., et al. 2016, *ApJ*, 821, L31
- Morganti, R. 2010, *PASA*, 27, 463
- Morganti, R. 2017, *Nat. Astron.*, 1, 596
- Morganti, R., Oosterloo, T., Oonk, J. B. R., Frieswijk, W., & Tadhunter, C. 2015, *A&A*, 580, A1
- Morganti, R., Peck, A. B., Oosterloo, T. A., et al. 2009, *A&A*, 505, 559
- Müller-Sánchez, F., Prieto, M. A., Mezcuca, M., et al. 2013, *ApJ*, 763, L1
- Narayan, R. & Yi, I. 1995, *ApJ*, 452, 710
- Nayakshin, S. & Zubovas, K. 2012, *MNRAS*, 427, 372
- Russell, H. R., Fabian, A. C., McNamara, B. R., & Broderick, A. E. 2015, *MNRAS*, 451, 588
- Russell, H. R., McNamara, B. R., Edge, A. C., et al. 2013, *MNRAS*, 432, 530
- Russell, H. R., McNamara, B. R., Fabian, A. C., et al. 2016, *MNRAS*, 458, 3134
- Saikia, D. J. & Jamrozy, M. 2009, *Bull. Astr. Soc. India*, 37, 27
- Sault, R. J., Teuben, P. J., & Wright, M. C. H. 2006, *ASPC*, 77, 433
- Schawinski, K., Koss, M., Berney, S., & Sartori, L. F. 2015, *MNRAS*, 451, 2517
- Shabala, S. S., Ash, S., Alexander, P., & Riley, J. M. 2008, *MNRAS*, 388, 625
- Siemiginowska, A., Sobolewska, M., Migliori, G., et al. 2016, *ApJ*, 823, 57
- Sofia, U. J., Lauroesch, J. T., Meyer, D. M., & Cartledge, S. I. B. 2004, *ApJ*, 605, 272
- Solomon, P. M. & Bout, P. A. V. 2005, *ARA&A*, 43, 677
- Solomon, P. M., Downes, D., Radford, S. J. E., & Barrett, J. W. 1997, *ApJ*, 478, 144
- Solomon, P. M., Rivolo, a. R., Barrett, J., & Yahil, A. 1987, *ApJ*, 319, 730
- Sparke, L. S., van Moorsel, G., Schwarz, U. J., & Vogelaar, M. 2009, *AJ*, 137, 3976
- Tingay, S. J., Jauncey, D. L., Reynolds, J. E., et al. 1997, *AJ*, 113, 2025
- Tingay, S. J., Macquart, J.-P., Collier, J. D., et al. 2015, *AJ*, 149, 74
- Tingay, S. J., Reynolds, J. E., Tzioumis, A. K., et al. 2002, *ApJS*, 141, 311
- Tremblay, G. R., Oonk, J. B. R., Combes, F., et al. 2016, *Nature*, 533, 1
- Veron-Cetty, M.-P., Woltjer, L., Ekers, R., et al. 1995, *A&A*, 297, L79
- Viti, S., García-Burillo, S., Fuente, A., et al. 2014, *A&A*, 570, A28
- Wada, K., Papadopoulos, P., & Spaans, M. 2009, *ApJ*, 702, 63
- Wada, K. & Tomisaka, K. 2005, *ApJ*, 619, 93
- Wiklind, T. & Combes, F. 1995, *A&A*, 299, 382
- Willett, K. W., Stocke, J. T., Darling, J., & Perlman, E. S. 2010, *ApJ*, 713, 1393
- Wong, K.-W., Irwin, J. A., Shcherbakov, R. V., et al. 2013, *ApJ*, 780, 9