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A STRONGLY LENSED MASSIVE ULTRACOMPACT QUIESCENT GALAXY AT $z \sim 2.4$ IN THE COSMOS/UltraVISTA FIELD*

ADAM MUZZIN¹, IVO LABBÉ¹, MARIJN FRANX¹, PIETER VAN DOKKUM², J. HOLT¹, DANIEL SZOMORU¹, JESSE VAN DE SANDE¹,
 GABRIEL BRAMMER³, DANILO MARCHESINI⁴, MAURO STEFANON⁵, F. BUITRAGO⁶, K. I. CAPUTI⁷, JAMES DUNLOP⁶,
 J. P. U. FYNBO⁸, OLIVIER LE FÉVRE⁹, HENRY J. MCCrackEN¹⁰, AND BO MILVANG-JENSEN⁸

¹ Leiden Observatory, Leiden University, P.O. Box 9513, 2300 RA Leiden, The Netherlands

² Department of Astronomy, Yale University, New Haven, CT 06520-8101, USA

³ European Southern Observatory, Alonso de Córdova 3107, Casilla 19001, Vitacura, Santiago, Chile

⁴ Department of Physics and Astronomy, Tufts University, Medford, MA 06520, USA

⁵ Observatori Astronòmic de la Universitat de València, E-46980 Paterna, València, Spain

⁶ SUPA, Institute for Astronomy, University of Edinburgh, Royal Observatory, Edinburgh EH9 3HJL, UK

⁷ Kapteyn Astronomical Institute, University of Groningen, P.O. Box 800, 9700 AV Groningen, The Netherlands

⁸ Dark Cosmology Centre, Niels Bohr Institute, University of Copenhagen, Juliane Maries Vej 30, DK-2100 Copenhagen, Denmark

⁹ Laboratoire d'Astrophysique de Marseille, CNRS and Aix-Marseille Université, 38 rue Frédéric Joliot-Curie, F-13388 Marseille Cedex 13, France

¹⁰ Institut d'Astrophysique de Paris, UMR7095 CNRS, Université Pierre et Marie Curie, 98 bis Boulevard Arago, F-75014 Paris, France

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ABSTRACT

We report the discovery of a massive ultracompact quiescent galaxy that has been strongly lensed into multiple images by a foreground galaxy at $z = 0.960$. This system was serendipitously discovered as a set of extremely K_s -bright high-redshift galaxies with red $J - K_s$ colors using new data from the UltraVISTA $YJHK_s$ near-infrared survey. The system was also previously identified as an optically faint lens/source system using the COSMOS Advanced Camera for Surveys (ACS) imaging by Faure et al. Photometric redshifts for the three brightest images of the source galaxy determined from 27-band photometry place the source at $z = 2.4 \pm 0.1$. We provide an updated lens model for the system that is a good fit to the positions and morphologies of the galaxies in the ACS image. The lens model implies that the magnification of the three brightest images is a factor of 4–5. We use the lens model, combined with the K_s -band image, to constrain the size and Sérsic profile of the galaxy. The best-fit model is an ultracompact galaxy ($R_e = 0.64^{+0.08}_{-0.18}$ kpc, lensing-corrected), with a Sérsic profile that is intermediate between a disk and a bulge profile ($n = 2.2^{+2.3}_{-0.9}$), albeit with considerable uncertainties on the Sérsic profile. We present aperture photometry for the source galaxy images that have been corrected for flux contamination from the central lens. The best-fit stellar population model is a massive galaxy ($\log(M_{\text{star}}/M_{\odot}) = 10.8^{+0.1}_{-0.1}$, lensing-corrected) with an age of $1.0^{+1.0}_{-0.4}$ Gyr, moderate dust extinction ($A_v = 0.8^{+0.5}_{-0.6}$), and a low specific star formation rate ($\log(\text{SSFR}) < -11.0 \text{ yr}^{-1}$). This is typical of massive “red-and-dead” galaxies at this redshift and confirms that this source is the first bona fide strongly lensed massive ultracompact quiescent galaxy to be discovered. We conclude with a discussion of the prospects of finding a larger sample of these galaxies.

Key words: galaxies: evolution – galaxies: fundamental parameters – galaxies: high-redshift – galaxies: structure – infrared: galaxies

Online-only material: color figures

1. INTRODUCTION

The brightening and magnifying effects from strong gravitational lensing are a powerful tool that permit us to study high-redshift galaxies with better signal-to-noise ratio (S/N) and higher spatial resolution than is normally possible with current instrumentation. In the last decade, significant successes in finding samples of lensed galaxies have come from surveys that have searched behind clusters (e.g., Bayliss et al. 2011b) and luminous red galaxies (e.g., Bolton et al. 2008; Faure et al. 2008; Cooray et al. 2011). These samples are now giving us insights into the evolution of typical galaxies at $1 < z < 5$, which are most frequently low-mass blue star-forming galaxies (e.g., Rigby et al. 2011; Wuyts et al. 2012; Sharon et al. 2012; Brammer et al. 2012). With the advent of the *Herschel* and South Pole Telescope submillimeter observatories we have also

gained access to a population of strongly lensed submillimeter galaxies that are readily identifiable in wide-field surveys (e.g., Negrello et al. 2010; Vieira et al. 2010) or behind clusters (e.g., Egami et al. 2010; Gladders et al. 2012). These lensed submillimeter galaxies are providing intriguing results about how dusty and strongly star-forming galaxies evolve at high redshift (e.g., Swinbank et al. 2010, 2011).

What has remained elusive are strongly lensed examples of the proverbial “red-and-dead” massive galaxies at $z \sim 2$. These galaxies have generated significant interest in the past few years after it was discovered that they are smaller in size than their similar-mass counterparts in the local universe by a factor of ~ 5 (e.g., Daddi et al. 2005; Trujillo et al. 2006; van Dokkum et al. 2008, 2010; Buitrago et al. 2008; Szomoru et al. 2010; Bruce et al. 2012, and numerous others). Understanding how they grow so substantially in size without adding significant stellar mass is currently one of the major open questions in galaxy evolution theory.

Finding gravitationally lensed examples of these sources could prove to be valuable for understanding their evolution.

* Based on data products from observations made with ESO Telescopes at the La Silla Paranal Observatory under ESO program ID 179.A-2005 and on data products produced by TERAPIX and the Cambridge Astronomy Survey Unit on behalf of the UltraVISTA consortium.

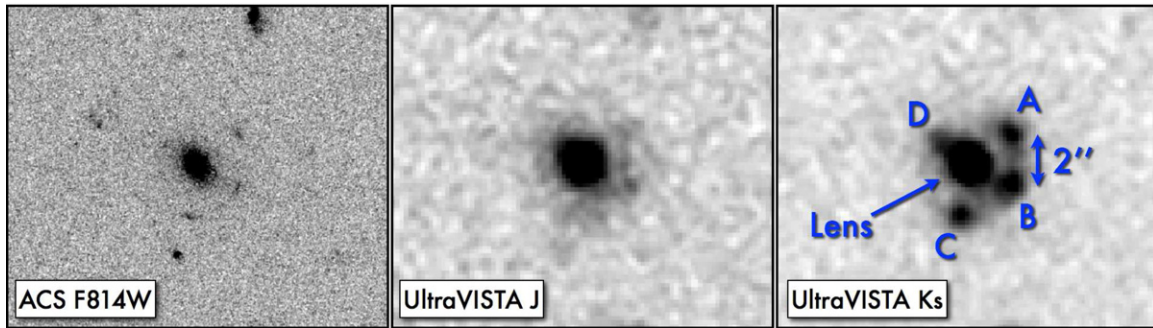


Figure 1. Postage stamp optical/NIR images of the lens/source system with a high-contrast stretch. The field of view in each image is $12'' \times 10''$. The lensed galaxy is faint in the observed optical and has an extremely red $J - K_s$ color. The separation between the multiple images is $\sim 2''$.

(A color version of this figure is available in the online journal.)

The magnifying effects from lensing will allow us to potentially resolve the central regions of these compact sources and understand what their central stellar density profiles are. The brightening effects from lensing will also allow us to obtain higher S/N spectroscopy of these systems. Currently, even the brightest of these systems are extremely faint, and determining quantities such as velocity dispersions and dynamical masses requires substantial integration times even on the world’s largest telescopes (e.g., van Dokkum et al. 2009; Cappellari et al. 2009; van de Sande et al. 2011; Toft et al. 2012).

Unfortunately, these galaxies are much more challenging to detect as lensed sources, and until now none have been securely identified. The challenge is that while they compose $\sim 50\%$ of the population of massive galaxies ($\log M_{\text{star}}/M_{\odot} > 11.0$) at $z \sim 2$ (e.g., Kriek et al. 2008; Brammer et al. 2011), they are still quite rare (~ 1 every 10 arcmin 2), making it unlikely to find one in a favorable alignment with a foreground lensing structure. Furthermore, they are extremely faint in the observed optical bands and also have very red $J - K_s$ near-infrared (NIR) colors (e.g., Franx et al. 2003; Kriek et al. 2008), making both deep and wide-field NIR imaging a requirement for their detection.

In this paper, we report the discovery of the first example of a massive ultracompact quiescent galaxy that has been strongly lensed by a foreground galaxy. The galaxy was serendipitously discovered as a set of bright and extremely red high-redshift galaxies using new data from the 1.8 deg 2 UltraVISTA $YJHK_s$ NIR survey (see McCracken et al. 2012). The lens/source system was previously identified as an optically faint quadruply lensed galaxy in the COSMOS field using Advanced Camera for Surveys (ACS) imaging (Faure et al. 2008, 2011); however, with the addition of the UltraVISTA NIR data it is now clear that the source is a massive quiescent galaxy.

This paper is organized as follows: In Section 2 we discuss how the lens/source system was discovered. In Section 3, we present an updated lens model for the system and use this to constrain the size and Sérsic profile of the source galaxy. In Section 4, we present updated 27-band photometry for the source galaxies that has been deblended from contamination from the central source. In Section 5, we use this photometry to determine an accurate photometric redshift and stellar population parameters for the galaxy. We conclude in Section 6 with a summary and a discussion and prospects of finding more strongly lensed massive ultracompact quiescent galaxies in future deep, wide-field NIR surveys. Throughout this paper we assume a $\Omega_{\Lambda} = 0.7$, $\Omega_m = 0.3$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ cosmology. All magnitudes are in the AB system.

2. IDENTIFICATION OF THE LENS/SOURCE SYSTEM

The lens/source system was serendipitously discovered using photometry from a K_s -selected catalog of the COSMOS/UltraVISTA field. The catalog contains photometry in 27 photometric bands, including the publicly available 7 optical broadband ($u^*g^+r^+i^+z^+B_jV_j$) and 12 optical medium-band (IA427 – IA827) imaging of the COSMOS field from Capak et al. (2007). It also includes the public-release $YJHK_s$ NIR imaging of the COSMOS field from the UltraVISTA survey (McCracken et al. 2012), as well as the public four-channel IRAC imaging from the S-COSMOS survey (Sanders et al. 2007). Object detection for the catalog was performed in the K_s band, which has a 5σ depth in a $2''$ aperture of $K_s = 23.9$ AB (McCracken et al. 2012). Photometry was measured using the SExtractor package (Bertin & Arnouts 1996) in $2''$ apertures using point-spread function (PSF) matched images in all bands. Full details of the COSMOS/UltraVISTA catalog will be presented in a future paper (A. Muzzin et al. 2012, in preparation).

The system (R.A., 10:00:50.55, decl., +02:49:01, J2000) was identified as a strong-lensing system via eye examination of the ACS F814W (Koekemoer et al. 2007) and UltraVISTA K_s -band (McCracken et al. 2012) images. This process was performed as a quality check of the catalog for a subsample of galaxies that were bright ($K_s < 21.0$ AB), red ($J - K_s > 1.5$ AB), and had photometric redshifts $z_{\text{photo}} > 2$. In Figure 1, we plot postage stamp images of the system in the ACS F814W and UltraVISTA J and K_s bands with a high-contrast stretch. Sources A and C (see labels in Figure 1) were the sources in the catalog that matched the selection criteria. Sources B and D lie along the major axis of the lens galaxy and hence are more blended with the lens and were not identified as unique sources by SExtractor. The extraordinary brightness of sources A and C given their z_{photo} , combined with the distinct cross-shape pattern of four sources around the brighter central lens (similar to quasar strong lenses), made it clear that the system was not just a very bright set of high-redshift galaxies, but in fact was a candidate strong-lensing system.

A literature search for known lens/source systems in the COSMOS field showed that this system had already been reported by Faure et al. (2008, 2011). Faure et al. (2008, 2011) designated the system as “COSMOS 0050+4901” and also noted that it is a clear quadruple-lens in the ACS images. Faure et al. (2011) report a spectroscopic redshift for the lens at $z_{\text{spec}} = 0.960$. No spectroscopic redshift was determined for the source galaxies, which are extremely faint in the observed optical bands; however, Faure et al. (2008) estimate the redshift

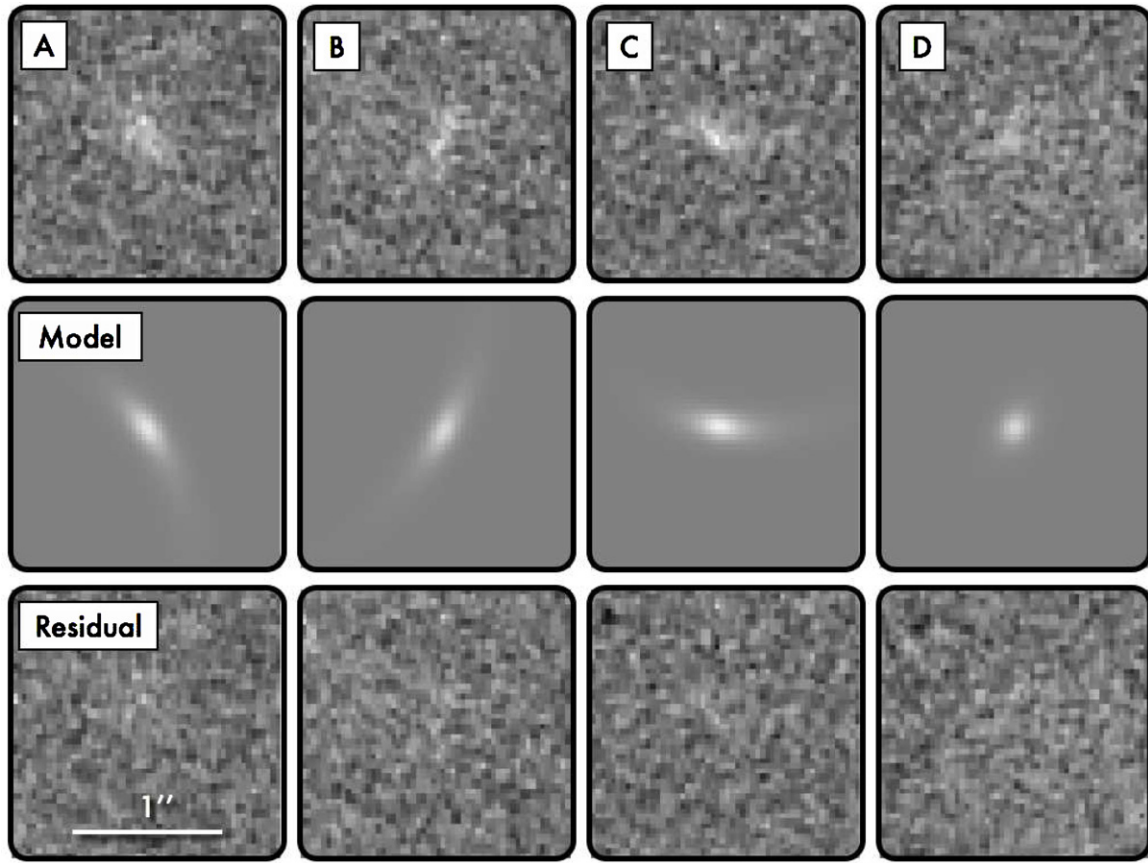


Figure 2. Top panels: ACS F814W postage stamps of the four images of the source galaxy. Middle panels: the best-fit ray-traced models for the source galaxy. Bottom panels: residuals of subtracting the model for the lens. The best-fit model simultaneously reproduces the relative brightness and positions of the source, as well as its morphology and size.

of the source as $z_{\text{photo}} = 3.34$. This does not agree well with our photometric redshift of the source (Section 5.1); however, the source galaxies are extremely faint in the observed optical, and their spectral energy distributions (SEDs) have no notable features there (see Section 5.2), so it is unsurprising that a highly uncertain photometric redshift was estimated without high-quality NIR data such as those from UltraVISTA.

The analysis presented in Faure et al. (2008, 2011) focuses on the properties of the lens galaxy. In this paper, we focus on the properties of the source galaxy, which is currently the only known candidate for a strongly lensed compact quiescent galaxy.

3. LENS MODEL

In order to measure the structural parameters of the source galaxy, we construct a model of the lens using the position and brightness of each of the four images of the source. We assume that the lens galaxy is embedded in a dark matter halo with an elliptical isothermal sphere profile with an additional core component. The lens model has six free parameters: the Einstein radius (R_{Ein}), core radius (R_{core}), ellipticity (e), position angle (P.A.), and X and Y positions of the lensed galaxy in the source plane (X_s, Y_s).

In order to constrain this large parameter space, we assumed a simple Sérsic profile for the source galaxy with an effective radius (R_e) of 1.0 kpc and a Sérsic index (n) of 2.5. We then generated a grid of models spanning the six-dimensional parameter space with fine sampling and used ray tracing to

construct mock images at ACS resolution of the four images of the lensed galaxy. The residuals from a subtraction of the ACS image and the mock images were then evaluated using χ^2 to determine the best-fit lens model parameters. The images, model, and residual for the best-fit lensing model are shown in Figure 2.

With four images of the source, the model is well constrained and the best-fit parameters are $R_{\text{Ein}} = 3''.94^{+0.15}_{-0.11}$, $R_{\text{core}} = 0''.5^{+0.4}_{-0.5}$, $e = 0.075^{+0.015}_{-0.010}$, P.A. = $-31.0^{+0.3}_{-0.5}$ deg, $X_s = 0''.15^{+0.02}_{-0.01}$, and $Y_s = -0''.11^{+0.01}_{-0.01}$. The error bars have been calculated using 100 Monte Carlo simulations where the background noise in the image was resampled and the lensing model was re-fit. The lensing model implies that the magnifications for the source images A, B, C, and D are factors of $4.5^{+0.4}_{-0.6}$, $5.3^{+0.7}_{-0.6}$, $4.6^{+0.3}_{-0.5}$, and $2.1^{+0.4}_{-0.3}$, respectively.

3.1. The Size and Sérsic Index of the Source Galaxy

With the parameters of the lensing model constrained by the positions and brightnesses of the multiple images, we fix those parameters and re-perform the ray tracing and χ^2 -fitting of the ACS images, this time allowing the R_e and n of the source galaxy as free parameters in the fit. This produced a best fit of $R_e = 0.43^{+2.40}_{-0.40}$ kpc and $n = 1.9^{+2.1}_{-0.9}$. Although formally fit as an ultracompact galaxy with a profile intermediate between a disk and a bulge, the uncertainties are extremely large due to the low S/N of the galaxy in the ACS F814W image.

The S/N of the ground-based K_s -band image is substantially higher than the ACS image (see Section 4), so we attempted

to obtain better constraints on the size and profile using those data. With the higher S/N K_s -band data we measure an $R_e = 0.64^{+0.08}_{-0.18}$ kpc and $n = 2.2^{+2.3}_{-0.9}$. This is remarkably consistent with the ACS measurement despite the much larger uncertainties on the ACS-measured R_e .

If we use the best-fit stellar mass of the system (Section 5.2), then the galaxy lies a factor of ~ 4.5 below the R_e versus stellar mass relation for galaxies in the local universe defined in van Dokkum et al. (2008). If we compare to the stellar mass–size relation at $z \sim 2$ defined in Bruce et al. (2012), we find that it is consistent with the most compact galaxies in that sample. Taken together, this shows that the source galaxy appears to be a member of the massive ultracompact galaxy population seen at $z \sim 2$.

Unfortunately, at ground-based resolution (or with the current S/N in the ACS image) we cannot make stronger statements about the Sérsic profile of the galaxy. We note that while uncertain, the best-fit Sérsic index, $n = 2.2^{+2.3}_{-0.9}$, suggests that the galaxy is intermediate between a disk and a bulge. This value is also similar to the median n of galaxies in the van Dokkum et al. (2008) sample ($n = 2.3$); however, we note that that sample has a large range of Sérsic indices ($n = 0.5$ – 4.5). Unfortunately, better constraints on the light profile of the source will not be possible until higher S/N imaging with *Hubble Space Telescope* resolution is available.

4. PHOTOMETRIC DATA

4.1. Deblending Method

The lens galaxy is bright and resolved at ground-based resolution. Examination of the PSF-matched images shows that its extended light distribution causes non-negligible contamination to the photometry of the source galaxies. In order to remove this contamination, we used the GALFIT package (Peng et al. 2010) to model the lens/source system and subtract the flux of the lens. For this procedure we do not use the PSF-matched images, which are convolved to the worst seeing image (the g^+ image, $\sim 1''.2$ seeing). Instead, we use the original images, which have typical seeing of $< 1''.0$ in the optical bands and $\sim 0''.8$ in the NIR band. This subtraction procedure is illustrated in Figure 3 and is as follows.

In the J , H , and K_s bands the source galaxies A, B, C, and D are clearly detected and can be distinguished from the lens in the images (see Figures 1 and 3). For these bands we used GALFIT to perform a simultaneous fit to all five galaxies. Given the complexity of a five-object simultaneous fit, the process was somewhat unstable; however, we continuously refined the initial guess parameters until a good fit with null residuals for all five galaxies was achieved. In these three bands the final fit values for the Sérsic index (n_{lens}), effective radius ($R_{e,\text{lens}}$), axis ratio (q_{lens}), and position angle (P.A._{lens}) of the central lens converged to similar values.

Once a good fit was determined, we subtracted only the model for the lens from the images. The model for the lens galaxy and the residuals after subtraction in the K_s and H bands are shown in Figure 3.

In the remaining 19 optical filters, as well as the UltraVISTA Y band, the source galaxies are barely detected (see Figure 3), and attempts at a simultaneous five-object fit using a wide range of initial guess parameters with GALFIT did not return reasonable values. In these bands we performed only a fit to the lens. This fitting was stable and in all bands converged to similar values of n_{lens} , $R_{e,\text{lens}}$, q_{lens} , and P.A._{lens}. For these bands

we subtracted the GALFIT model for the central lens (the Y -band subtraction is shown in Figure 3).

4.2. Optical and NIR Photometry

With the contribution from the lens subtracted, we performed aperture photometry on the residual images using SExtractor in dual image mode. The source galaxies are brightest in the K_s band, so the K_s -band residual image was used as the detection image. Photometry was performed in $1''$ diameter apertures on the original seeing images in each band. These images have different PSFs; however, using the PSF-matched images for photometry would be sub-optimal because they have had the good seeing data smoothed to the larger seeing disk of the worse seeing image. Photometry on these images requires larger apertures that are more susceptible to contamination from residual flux missed in the fitting of the lens, as well as flux from the other source images that are separated by $\sim 2''$. We correct the photometry in each individual band to a total flux by assuming that the source galaxies are point sources and extrapolating the growth curves of bright stars in the field. Formally, this is not completely correct because the source galaxy is resolved at ground-based resolution; however, as discussed in Section 3, the R_e of the source is a factor of two smaller than the PSF size, so a point source provides a reasonable first-order approximation to the observed profile.

The PSFs in the UltraVISTA $YJHK_s$ are remarkably similar, and the seeing is $0''.75$, $0''.76$, $0''.79$, and $0''.82$ in the K_s , H , J , and Y bands, respectively (McCracken et al. 2012). The correction to total magnitudes in these bands ranges between a factor of 2.04 and 2.44; hence, the NIR colors do not depend strongly on the aperture corrections. The PSF and seeing range in the optical is larger, $0''.5$ – $1''.2$ (Capak et al. 2007). The aperture corrections range between a factor of 1.79 and 3.33. This is larger; however, as we discuss below, the source galaxies are not detected at $> 2\sigma$ in the vast majority of the optical bands (the exception is the deep COSMOS i' filter, where they are detected at $\sim 4\sigma$ – 5σ), so these bands provide only weak constraints on the SED of the galaxies.

Photometry and associated errors for the source galaxies and the lens are listed in Table 1. Calculation of the photometric errors includes three sources of uncertainty. We estimate the uncertainty in the background subtraction by calculating the variance of the background in $1''$ empty apertures in regions around the lens/source system. We also estimate a contribution from photon noise using the gain of the detector and the original exposure times, although we note that this contribution is much smaller than the background variations for sources this faint. Lastly, we perform aperture photometry on the GALFIT model at the location of each source. We use the square root of this flux as an estimate in the additional uncertainty from subtraction of the model. These three sources of uncertainty are added in quadrature to determine the total photometric error. The aperture-corrected photometry, the photometric errors, and the aperture corrections are listed in Table 1.

4.3. IRAC Photometry

The S-COSMOS IRAC imaging of the source is also deblended using GALFIT. The source galaxies are clearly detected in all four IRAC channels; however, the large FWHM of the IRAC PSF creates stronger blends between the lens and source than in the optical and NIR bands, and simultaneous five-object fits were highly unstable. The problem was further complicated

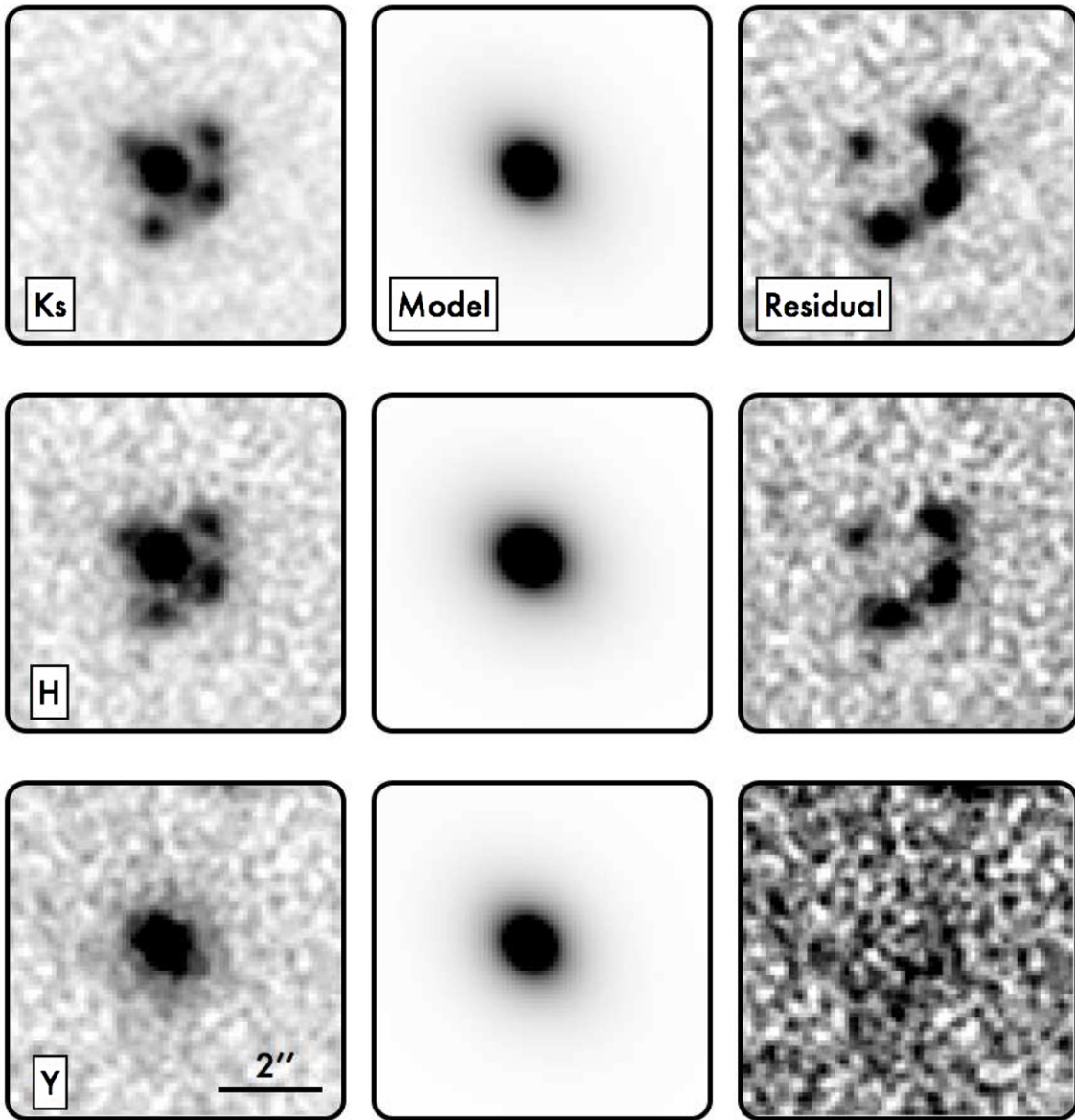


Figure 3. Left panels: images of the lens/source system in the K_s , H , and Y bands. Middle panels: models for the lens determined using GALFIT. Right panels: residuals of subtracting the model for the lens. Aperture photometry is performed on these residual images.

by the fact that the source galaxies are almost as bright as the lens in the $5.8 \mu\text{m}$ and $8.0 \mu\text{m}$ channels. Therefore, instead of a simultaneous five-object fit, we used the same approach as in the optical bands and fit only the central lens. Given the strong blending, we left the only free parameter in the fit as the flux of the lens and constrained the position, n_{lens} , $R_{e,\text{lens}}$, q_{lens} , and $\text{P.A.}_{\text{lens}}$ using the best-fit GALFIT parameters from the NIR fits. This constrained fit was much more stable, and the residuals at the position of the lens after subtraction of the model were low.

With the lens subtracted, aperture photometry was also performed in $1''.0$ diameter apertures and corrected to a total flux using the growth curve of bright stars. The aperture corrections for the IRAC photometry are substantially larger than the optical data, by a factor of ~ 6 (see Table 1). Although this correction is large, the IRAC PSF is well determined and quite stable, so the correction should be well determined. Regardless, we note that the inclusion of the IRAC photometry does not change the best-fit values of the parameters derived from the SED fitting

(Section 5.2). It does considerably reduce the uncertainties in these parameters, particularly on the dust extinction. This is typical for high-redshift massive galaxies (e.g., Muzzin et al. 2009b), and therefore we note that any uncertainties in the IRAC photometry will not change the interpretation of the source galaxy's SEDs.

5. PROPERTIES OF THE LENSED GALAXY

5.1. Photometric Redshifts

Photometric redshifts (z_{photo}) for both the lens and source galaxy images are determined using the EAZY photometric redshift code (Brammer et al. 2008). All 27 photometric bands are used in the determination of the redshift, and the values are listed in Table 2. The photometry for the source galaxies is derived from the deblended images, whereas we have derived the photometry for the lens from the original PSF-matched K_s -selected catalog. The z_{photo} of the lens galaxy is 0.98 ± 0.03 ,

Table 1
Photometric Data

Filter	Lens Flux Density	Source A Flux Density	Source B Flux Density	Source C Flux Density	Source D Flux Density	Ap Corr
(1)	(2)	(3)	(4)	(5)	(6)	(7)
u^*	1.28 ± 0.47	0.06 ± 0.84	0.35 ± 0.84	0.37 ± 0.84	0.29 ± 0.84	2.33
B	2.34 ± 0.25	0.28 ± 0.46	0.28 ± 0.47	0.43 ± 0.46	0.10 ± 0.47	1.79
g'	3.12 ± 0.30	0.49 ± 1.14	0.64 ± 1.15	1.52 ± 1.14	0.01 ± 1.15	3.33
V	5.32 ± 0.26	0.07 ± 0.48	0.35 ± 0.49	0.69 ± 0.48	0.12 ± 0.49	2.00
r'	10.82 ± 0.29	0.57 ± 0.43	0.71 ± 0.44	0.88 ± 0.43	0.36 ± 0.45	2.08
i'	24.66 ± 0.28	0.92 ± 0.25	1.15 ± 0.26	1.04 ± 0.25	0.02 ± 0.28	2.00
z'	41.25 ± 0.39	0.53 ± 0.64	0.91 ± 0.66	1.48 ± 0.64	0.56 ± 0.68	2.33
$IA427$	2.55 ± 0.89	0.63 ± 1.14	0.72 ± 1.14	-0.27 ± 1.14	-0.37 ± 1.14	2.04
$IA464$	1.75 ± 0.77	0.93 ± 1.81	0.97 ± 1.81	0.21 ± 1.81	0.96 ± 1.81	2.94
$IA484$	2.24 ± 0.62	0.07 ± 0.82	0.33 ± 0.83	0.75 ± 0.82	0.01 ± 0.83	1.96
$IA527$	4.50 ± 0.68	0.02 ± 0.67	0.49 ± 0.68	0.59 ± 0.67	0.00 ± 0.68	1.89
$IA505$	2.83 ± 0.73	-0.66 ± 0.99	0.31 ± 0.99	0.23 ± 0.99	0.52 ± 0.99	2.04
$IA574$	6.43 ± 0.79	0.48 ± 1.09	1.25 ± 1.10	0.30 ± 1.09	0.46 ± 1.10	2.63
$IA624$	10.17 ± 0.74	0.49 ± 0.57	0.40 ± 0.57	0.80 ± 0.57	0.06 ± 0.59	1.89
$IA679$	14.30 ± 0.65	1.13 ± 0.76	0.68 ± 0.78	0.75 ± 0.76	0.04 ± 0.79	2.63
$IA709$	16.08 ± 0.76	1.11 ± 0.59	1.04 ± 0.59	1.30 ± 0.58	0.83 ± 0.61	2.38
$IA738$	17.19 ± 0.79	1.13 ± 0.59	0.64 ± 0.60	1.16 ± 0.58	0.02 ± 0.62	2.04
$IA767$	18.22 ± 0.77	0.58 ± 0.86	0.57 ± 0.87	1.58 ± 0.86	0.26 ± 0.89	2.94
$IA827$	31.68 ± 0.97	0.04 ± 0.90	1.32 ± 0.92	1.42 ± 0.90	-0.03 ± 0.96	3.23
Y	44.47 ± 0.48	1.24 ± 1.10	2.44 ± 1.13	2.05 ± 1.10	0.83 ± 1.17	2.44
J	46.44 ± 0.35	2.55 ± 0.78	3.18 ± 0.80	2.98 ± 0.78	0.77 ± 0.85	2.27
H	44.78 ± 0.28	11.06 ± 0.57	12.00 ± 0.58	9.88 ± 0.57	5.65 ± 0.61	2.13
K_s	36.38 ± 0.23	10.68 ± 0.31	10.95 ± 0.32	10.23 ± 0.31	4.72 ± 0.34	2.04
$3.6 \mu\text{m}$	17.58 ± 0.07	7.45 ± 0.27	7.38 ± 0.32	7.07 ± 0.24	2.00 ± 0.32	5.88
$4.5 \mu\text{m}$	9.85 ± 0.06	6.11 ± 0.20	6.52 ± 0.21	5.75 ± 0.18	3.52 ± 0.23	5.88
$5.8 \mu\text{m}$	5.13 ± 0.19	3.28 ± 0.42	4.06 ± 0.42	3.47 ± 0.42	1.27 ± 0.43	6.67
$8.0 \mu\text{m}$	1.94 ± 0.11	1.05 ± 0.21	1.32 ± 0.21	0.90 ± 0.21	0.52 ± 0.21	5.88

Note. All flux densities are listed in units of $10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$.

Table 2
Stellar Population Parameters

Parameter	Lens	Source A	Source B	Source C	Source D
(1)	(2)	(3)	(4)	(5)	(6)
Bruzual & Charlot (2003) Models					
z_{photo}	0.98 ± 0.03	2.41 ± 0.13	2.37 ± 0.18	2.36 ± 0.15	2.22 ± 0.22
$\log(\text{Stellar Mass}) (M/M_{\odot})$	$11.49^{+0.11}_{-0.19}$	$10.82^{+0.05}_{-0.07} (11.48)$	$10.78^{+0.07}_{-0.10} (11.50)$	$10.70^{+0.17}_{-0.01} (11.36)$	$10.68^{+0.16}_{-0.17} (11.00)$
$\log(\tau) \text{ (yr)}$	$8.5^{+0.3}_{-1.5}$	$8.0^{+0.3}_{-1.0}$	$8.1^{+0.5}_{-1.1}$	$8.2^{+0.3}_{-0.3}$	$7.0^{+1.6}_{-0.0}$
$\log(\text{Age}) \text{ (yr)}$	$9.4^{+0.3}_{-0.3}$	$9.0^{+0.2}_{-0.2}$	$9.0^{+0.4}_{-0.2}$	$9.1^{+0.4}_{-0.2}$	$9.0^{+0.5}_{-0.2}$
A_v	$0.6^{+0.3}_{-0.4}$	$0.9^{+0.2}_{-0.6}$	$1.0^{+0.2}_{-0.7}$	$0.6^{+0.5}_{-0.4}$	$0.8^{+0.5}_{-0.8}$
$\log(\text{SFR}) (M_{\odot} \text{ yr}^{-1})$	$-0.21^{+0.07}_{-0.99}$	$-1.26^{+0.81}_{-0.99} (-0.61)$	$-0.54^{+0.13}_{-0.99} (0.18)$	$-0.71^{+0.82}_{-0.07} (-0.05)$	$-99.9^{+99.5}_{-0.0} (-99.9)$
$\log(\text{SSFR}) \text{ (yr}^{-1}\text{)}$	$-11.70^{+0.09}_{-0.99}$	$-12.10^{+0.82}_{-0.99}$	$-11.32^{+0.07}_{-0.99}$	$-11.41^{+0.69}_{-0.20}$	$-99.9^{+87.60}_{-0.0}$
Maraston (2005) Models					
z_{photo}	0.98 ± 0.03	2.41 ± 0.13	2.37 ± 0.18	2.36 ± 0.15	2.22 ± 0.22
$\log(\text{Stellar Mass}) (M/M_{\odot})$	$11.40^{+0.00}_{-0.00}$	$10.76^{+0.17}_{-0.09} (11.41)$	$10.69^{+0.16}_{-0.14} (11.41)$	$10.68^{+0.13}_{-0.08} (11.34)$	$10.72^{+0.16}_{-0.14} (11.04)$
$\log(\tau) \text{ (yr)}$	$8.7^{+0.0}_{-0.0}$	$8.2^{+0.3}_{-1.2}$	$8.3^{+0.2}_{-1.3}$	$8.3^{+0.2}_{-0.2}$	$7.9^{+0.7}_{-0.9}$
$\log(\text{Age}) \text{ (yr)}$	$9.6^{+0.0}_{-0.0}$	$9.2^{+0.2}_{-0.2}$	$9.2^{+0.4}_{-0.2}$	$9.2^{+0.4}_{-0.2}$	$9.3^{+0.2}_{-0.1}$
A_v	$0.0^{+0.0}_{-0.0}$	$0.2^{+0.2}_{-0.2}$	$0.3^{+0.2}_{-0.3}$	$0.2^{+0.2}_{-0.2}$	$0.2^{+0.4}_{-0.2}$
$\log(\text{SFR}) (M_{\odot} \text{ yr}^{-1})$	$-0.53^{+0.00}_{-0.00}$	$-0.97^{+0.04}_{-0.77} (-0.32)$	$-1.64^{+0.81}_{-0.99} (-0.92)$	$-0.79^{+0.07}_{-0.99} (-0.13)$	$-8.22^{+7.1}_{-99.9} (-7.56)$
$\log(\text{SSFR}) \text{ (yr}^{-1}\text{)}$	$-11.94^{+0.00}_{-0.00}$	$-12.32^{+0.77}_{-0.99}$	$-11.55^{+0.10}_{-0.99}$	$-11.55^{+0.10}_{-0.10}$	$-18.59^{+7.10}_{-99.9}$

Notes. The values listed for $\log(\text{Stellar Mass})$ and $\log(\text{SFR})$ are corrected based on the lensing model. The values in parentheses are the direct measurements with no correction for lensing.

which agrees well with the spectroscopic redshift of $z = 0.960$ measured by Faure et al. (2011). The photometric redshifts for the images A, B, C, and D are 2.41 ± 0.13 , 2.37 ± 0.18 , 2.36 ± 0.15 , and 2.22 ± 0.22 , respectively, and are all consistent within the 1σ uncertainties.

5.2. Stellar Mass, Age, and Dust Content

We perform SED fitting on the photometric data to determine the stellar population parameters using the FAST fitting code (Kriek et al. 2009). For the SED fitting we assume the redshift

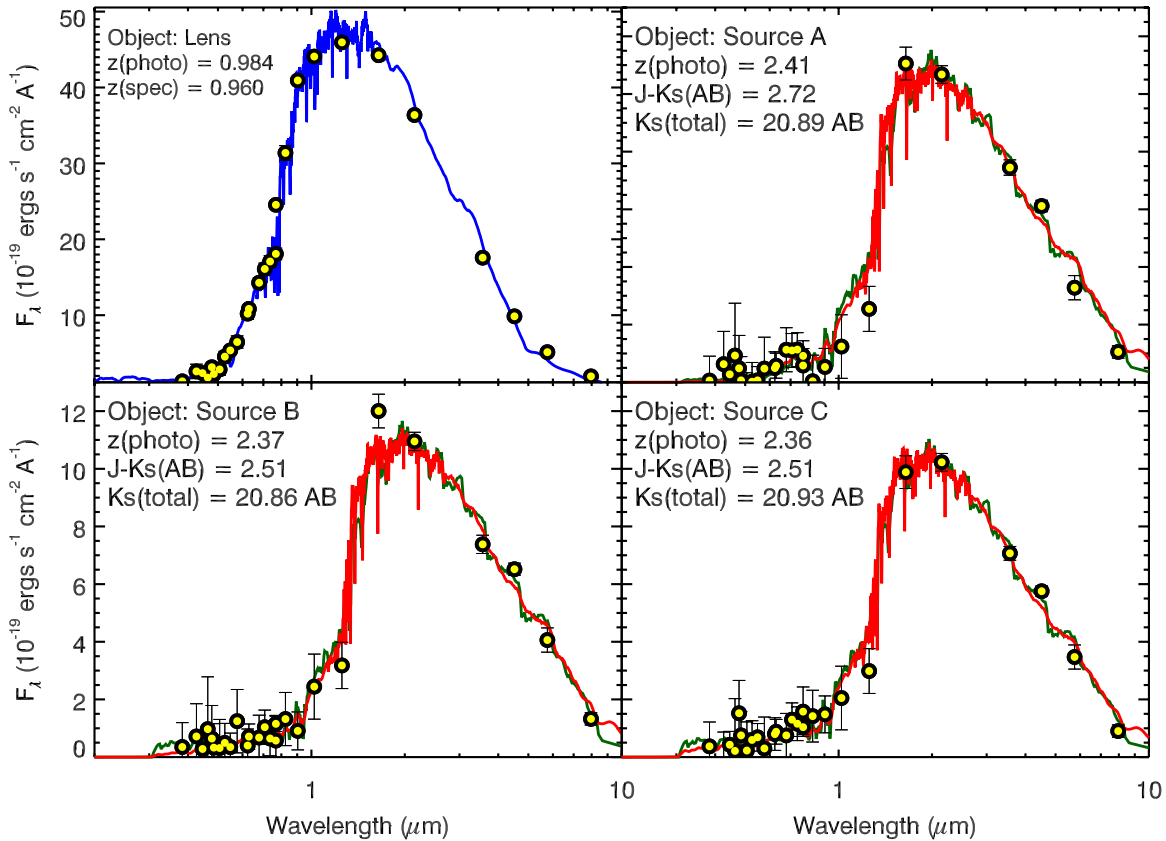


Figure 4. Spectral energy distributions for the lens (upper left) and the three brightest images of the source galaxy (sources A, B, C). The SEDs of the sources have been measured from aperture photometry on the images after subtraction of the central source using GALFIT. The best-fit Bruzual & Charlot model is plotted in blue for the lens and in red for the source galaxies. The best-fit Maraston model for the sources is shown in green. The photometric redshifts of all three images are consistent at $z \sim 2.4$. The source galaxies are extremely red ($J - K_s(\text{AB}) = 2.6$) and are best fit by (relatively) old templates with a moderate amount of dust ($A_v \sim 0.8$).

(A color version of this figure is available in the online journal.)

as the best-fit z_{photo} from EAZY. We use the stellar population models from Bruzual & Charlot (2003) as our default model, but we also fit to the Maraston (2005) models. We assume solar metallicity, a Calzetti et al. (2000) dust law, and a Chabrier (2003) initial mass function. We parameterize the star formation history as a declining exponential with an e -folding timescale of τ . Overall we fit four parameters: the stellar mass (M_{star}), age since the onset of star formation (t), τ , and the V -band dust attenuation (A_v). Integration of the star formation history also returns a star formation rate (SFR) and a star formation rate per unit stellar mass (specific star formation rate, SSFR). The best-fit SEDs for both the lens and the three brightest sources, A, B, and C, are plotted along with the corresponding photometry in Figure 4, and the best-fit stellar population parameters and associated errors are listed in Table 2.

The lens is best fit as an extremely massive ($\log(M_{\text{star}}/M_{\odot}) = 11.49^{+0.11}_{-0.19}$), old ($\log(\text{Age}/\text{yr}^{-1}) = 9.4^{+0.3}_{-0.3}$), and quiescent galaxy ($\log(\text{SSFR}/\text{yr}^{-1}) = -11.70^{+0.09}_{-0.99}$). The M_{star} derived from our photometry agrees reasonably well with the value of $\log(M_{\text{star}}/M_{\odot}) = 11.64 \pm 0.03$ measured by Faure et al. (2011).

As Figure 4 shows, the SEDs of images A, B, and C are all similar within the uncertainties. This is not necessarily expected, as small perturbations in the lens potential and/or a non-symmetric source galaxy can result in different locations of the source being more or less magnified in each image (e.g., Seitz et al. 1998; Sharon et al. 2012). We evaluate the

overall best-fit stellar population parameters for the source by considering those derived from images A, B, C, which are the most strongly magnified images. We take the median value of all parameters but consider the full range of uncertainty allowed by the fits. Source D, which is less magnified, is excluded in this calculation because the uncertainties are much larger than the other three images of the source. We find that the source is massive ($\log(M_{\text{star}}/M_{\odot}) = 10.8^{+0.1}_{-0.1}$, corrected for lensing), old ($\log(\text{Age}/\text{yr}^{-1}) = 9.0^{+1.0}_{-0.4}$), and quiescent ($\log(\text{SSFR}/\text{yr}^{-1}) = -11.41^{+0.69}_{-0.99}$) with moderate dust extinction ($A_v = 0.8^{+0.5}_{-0.6}$).

These stellar population parameters are quite typical of the well-studied population of massive ultracompact quiescent galaxies at $z \sim 2$ (e.g., Kriek et al. 2008; van Dokkum et al. 2008; Muzzin et al. 2009a, 2009b; van de Sande et al. 2011; Cassata et al. 2011). This suggests that the source is likely to be a strongly lensed member of that population. The implied SSFR is extremely low; however, it is derived from the rest-frame UV/optical SED. In order to test for any potential obscured star formation, we examined the S-COSMOS MIPS $24 \mu\text{m}$ image of the lens/source system.

5.3. MIPS $24 \mu\text{m}$ Photometry

The S-COSMOS $24 \mu\text{m}$ imaging is deep enough to detect galaxies with $\sim 100 M_{\odot} \text{ yr}^{-1}$ of star formation at $z \sim 2.4$. Unfortunately, the FWHM resolution of the data is $5''.5$, which is larger than the separation between both the source and lens and the source galaxies themselves. Given that individual sources

cannot be resolved at 24 μm , we performed photometry in a 7'' diameter aperture to get a total flux density for the complete lens/source system. Within this aperture we measure a flux of $26.3 \pm 6.3 \mu\text{Jy}$. Based on the aperture corrections in the MIPS data handbook, we correct this to a total flux of $67.3 \pm 16.1 \mu\text{Jy}$.

The aperture photometry implies that there is a significant 4σ detection at the location of the lens/source system. Examination of the image shows evidence for a faint object within the aperture; however, we note that it is offset by a few arcseconds to the southwest of the lens/source system. This clear offset suggests that the lens/source is not the correct counterpart. There are, however, no other plausible counterparts in the optical or NIR images near the 24 μm detection, so if the lens/source is not the counterpart, then the true counterpart must be highly extincted.

We investigate the implications if the lens/source is the counterpart of the MIPS source. If we assume that 100% of the detected flux comes from the source and none comes from the lens, then the implied total SFR rate derived from the logarithmic average of the Dale & Helou (2002) templates (see Wuyts et al. 2008; Muzzin et al. 2010), uncorrected for lensing, is $111^{+26}_{-27} M_{\odot} \text{ yr}^{-1}$. The total M_{star} in all four images, uncorrected for lensing, is $\log(M_{\text{star}}/M_{\odot}) = 11.97^{+0.17}_{-0.07}$. This implies a hard upper limit for the SSFR of the source as $\log(\text{SSFR}) = -9.93^{+0.20}_{-0.20}$. This is ~ 1.4 dex higher than the SSFR determined from the SED fitting (see Table 2). If 100% of the flux actually comes from the source, then it implies that there may be some optically thick star formation continuing within the galaxy, or that it contains an obscured active galactic nucleus (AGN), or both. We also checked if the lens/source was detected in either the *XMM-Newton* X-ray and Very Large Array radio observations and catalogs of the COSMOS field but found no detection at the location of the lens/source.

Recent papers have shown evidence for a “main sequence” of star-forming galaxies out to $z \sim 3$ (e.g., Noeske et al. 2007; Elbaz et al. 2007; Wuyts et al. 2011; Whitaker et al. 2012). We note that while the MIPS flux of the source galaxies may imply that it is not completely quiescent, the typical SSFR of galaxies with $\log(M_{\text{star}}/M_{\odot}) = 10.8$ at $z \sim 2.5$ is $\log(\text{SSFR}) = -8.6 \text{ yr}^{-1}$ (Whitaker et al. 2012). This is more than an order of magnitude higher than the upper limits on the SSFR, suggesting that if the source is a star-forming galaxy, it lies well off the star-forming main sequence and may be en route to becoming fully quenched.

We also cannot rule out the possibility that some or all of the 24 μm flux comes from the lens. The SED-measured SFR of the lens is only $0.6 M_{\odot} \text{ yr}^{-1}$, which would only account for 4 μJy , or 6% of the observed flux. However, we note that only $10 M_{\odot} \text{ yr}^{-1}$ of obscured star formation in the lens could account for 100% of the observed 24 μm flux. Likewise, an obscured AGN could account for a fraction or all of the total flux.

6. DISCUSSION AND CONCLUSION

Overall, the size constraints from the lens modeling and the SED modeling confirm that the strongly lensed galaxy COSMOS 0050+4901 is in fact the first example of a strongly lensed massive ultracompact quiescent galaxy. The high magnification from the strong lensing implies that with future high-S/N, high-resolution imaging we may have the possibility to peer into the cores of these compact systems. Likewise, the brightening effects should allow us to determine a high-quality velocity

dispersion with a reasonable integration time on an 8 m class telescope.

Discovering a larger representative sample of strongly lensed massive ultracompact quiescent galaxies would be an important step forward to obtaining a more detailed picture of their evolution. It would also allow comparison studies with the samples of strongly lensed blue star-forming and submillimeter galaxies that already exist. One question that needs to be answered in order to find such samples is, how frequently do these lensed sources appear, and therefore how much area with high-quality NIR imaging needs to be searched? This is a difficult question to answer precisely, as it requires knowledge of the underlying set of lenses and sources, as well as an understanding of the observational selection effects. Here we present an order-of-magnitude estimate based on what is currently known about current strong-lensing samples, the stellar mass function, and the fact that we have discovered one such source in the UltraVISTA field.

If we make the optimistic assumption that searches for strong lenses can identify all strongly lensed galaxies down to $\log(M_{\text{star}}/M_{\odot}) = 9.5$ (these are the lowest mass galaxies that have been identified with strong lensing, e.g., Rigby et al. 2011; Wuyts et al. 2012), then the fraction of high-redshift strong lenses that will be massive galaxies at $z \sim 2$ can be estimated using the stellar mass function. If we integrate the stellar mass function at $2 < z < 3$ from Marchesini et al. (2009), we find that the ratio of galaxies with $\log(M_{\text{star}}/M_{\odot}) > 11.0$ compared to those with $9.5 < \log(M_{\text{star}}/M_{\odot}) < 11.0$ is a factor of 20.7. This suggests that roughly 1 in 20 high-redshift strong lenses will be a massive galaxy, although the true number recovered will depend on observational selection effects. Of these massive galaxies, only $\sim 50\%$ will be quiescent (e.g., Kriek et al. 2008; Muzzin et al. 2009b).

In their “high-quality” sample of 16 strong lenses in the COSMOS field, Faure et al. (2008) estimate that 9/16 strong lenses are at $z > 2$, and this is consistent with the median redshift of the brightest strongly lensed galaxies in cluster surveys (e.g., Bayliss et al. 2011a, 2011b).

To estimate the frequency of strongly lensed massive quiescent galaxies at $z > 2$, we use the total sample of 67 strong lenses in Faure et al. (2008) as an estimate of the total number of strongly lensed galaxies per COSMOS field ($\sim 1.5 \text{ deg}^2$). Using this empirically derived number as a starting point does account for some of the observational selection effects. The total number of strongly lensed massive quiescent galaxies that are expected within the COSMOS field should be (total lenses) \times (fraction of lenses at $z > 2$) \times (fraction of lensed sources that are $\log(M_{\text{star}}/M_{\odot}) > 11.0$) \times (fraction of $\log(M_{\text{star}}/M_{\odot}) > 11.0$ galaxies that are quiescent). Filling in the numbers above gives $67 \times 0.56 \times 0.05 \times 0.5 = 0.94$. The number is of order unity, suggesting that there should be approximately one such source per COSMOS field, consistent with the single source we have discovered. This agreement is comforting and suggests that the order-of-magnitude calculation is reasonable; however, all of the numbers used in the estimate are highly uncertain, and most likely there are additional selection effects that are not accounted for.

The true frequency of strongly lensed massive ultracompact quiescent galaxies may be factors of several larger or smaller. Still, the fact that we found only one source in the COSMOS/UltraVISTA survey would seem to rule out the possibility that these sources are significantly more abundant than our estimate. Also, it is worth noting that no other such sources

have yet been reported in similar-sized surveys such as the UDS (Williams et al. 2009) and NMBS (Whitaker et al. 2011), which would again support the idea that the abundance is not much higher than this estimate.

It is possible that the abundance is lower than our estimate, and that it is fortuitous to find such a system in the COSMOS/UltraVISTA field. The lack of such sources in the UDS and in the 50% of the NMBS that is not in the COSMOS field means that the current total area of deep-wide NIR surveys is $\sim 2.5 \text{ deg}^2$, with only one lensed massive, ultracompact, quiescent galaxy discovered so far. It is difficult to say more without better statistics, but if the order-of-magnitude estimate is correct, it suggests that the space density of strongly lensed massive quiescent galaxies at $z > 2$ is roughly one per $1\text{--}2 \text{ deg}^2$. This suggests that in order to obtain a sample of 10 such galaxies, medium-deep optical/NIR photometry with good angular resolution (to avoid lens/source blending) covering $\sim 10\text{--}20 \text{ deg}^2$ would be required. Optical/NIR surveys with these requirements are currently being performed, which suggests that if searches are careful, the prospects are good for obtaining a real sample of strongly lensed massive ultracompact quiescent galaxies within the next few years.

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