

# The rise and fall of star formation in $z \sim 0.2$ merging galaxy clusters

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Accepted 2014 November 19. Received 2014 November 6; in original form 2014 October 10

## ABSTRACT

CIZA J2242.8+5301 (‘Sausage’) and 1RXS J0603.3+4213 (‘Toothbrush’) are two low-redshift ( $z \sim 0.2$ ), massive ( $\sim 2 \times 10^{15} M_{\odot}$ ), post-core passage merging clusters, which host-shock waves traced by diffuse radio emission. To study their star formation properties, we uniformly survey the ‘Sausage’ and ‘Toothbrush’ clusters in broad- and narrow-band filters and select a sample of 201 and 463 line emitters, down to a rest-frame equivalent width (13 Å). We robustly separate between  $H\alpha$  and higher redshift emitters using a combination of optical multiband ( $B, g, V, r, i, z$ ) and spectroscopic data. We build  $H\alpha$  luminosity functions for the entire cluster region, near the shock fronts, and away from the shock fronts and find striking differences between the two clusters. In the dynamically younger, 1 Gyr old ‘Sausage’ cluster we find numerous (59)  $H\alpha$  emitters above a star formation rate (SFR) of  $0.17 M_{\odot} \text{ yr}^{-1}$  surprisingly located in close proximity to the shock fronts, embedded in very hot intracluster medium plasma. The SFR density for the cluster population is at least at the level of typical galaxies at  $z \sim 2$ . Down to the same SFR, the possibly dynamically more evolved ‘Toothbrush’ cluster has only nine  $H\alpha$  galaxies. The cluster  $H\alpha$  galaxies fall on the SFR–stellar mass relation  $z \sim 0.2$  for the field. However, the ‘Sausage’ cluster has an  $H\alpha$  emitter density  $>20$  times that of blank fields. If the shock passes through gas-rich cluster galaxies, the compressed gas could collapse into dense clouds and excite star formation for a few 100 Myr. This process ultimately leads to a rapid consumption of the molecular gas, accelerating the transformation of gas-rich field spirals into cluster S0s or ellipticals.

**Key words:** shock waves – galaxies: clusters: individual: CIZA J2242.8+5301 – galaxies: clusters: individual: 1RXS J0603.3+4213 – galaxies: evolution – cosmology: observations – large-scale structure of Universe.

## 1 INTRODUCTION

Galaxy clusters grow by merging with other clusters and via accretion of galaxies (e.g. Sarazin 2002). Even at low redshifts ( $z < 0.5$ ), a significant population of galaxy clusters are undergoing mergers, with clear evidence from their disturbed intracluster medium (ICM) X-ray emission. Merging clusters are a unique probe of the interaction between dark matter, the ICM and the galaxies. They provide

us with a way to test hierarchical structure formation, cosmic ray physics and galaxy evolution in dense environments. Major mergers have been argued to lead to increased turbulence within the ICM. In a number of cases merging clusters have been observed to produce travelling shock waves with Mach numbers ( $M$ ) below 4 (e.g. Brunetti & Jones 2014). The shock fronts (re-)accelerate ICM electrons, which radiate synchrotron emission, observed in the radio as giant relics at cluster peripheries (Brunetti & Jones 2014).

### 1.1 Star-forming galaxies in clusters

The ICM interacts strongly with the cluster galaxies and is efficient in transforming the star-forming properties of member galaxies

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and/or maintaining them quenched (e.g. Butcher & Oemler 1978a,b; Dressler 1980). The cluster galaxy population is dominated by passive, massive elliptical galaxies. The total galaxy number density in cluster environments is higher than in the field. Nevertheless, owing to a low fraction of blue, late types within clusters, the number density of star-forming cluster galaxies is generally lower than in the field (e.g. Dressler 1980; Goto et al. 2003). For example, by using the H $\alpha$  emission line which traces recent (<10 Myr) star formation (SF), multiple authors have found that the number density of star-forming galaxies is  $\sim 50$  per cent lower than in blank fields (e.g. Couch et al. 2001; Balogh et al. 2002; Kodama et al. 2004). Neutral hydrogen (H I) observations also show that cluster spirals contain significantly less H I gas than their field counterparts (e.g. Cayatte et al. 1990).

Hence, dense cluster environments seem to suppress SF and probably lead to a morphological transformation from gas-rich spirals into gas-poor ellipticals. The deficit of star-forming galaxies in clusters is thought to occur through the process of ram pressure stripping (e.g. Gunn & Gott 1972; Fumagalli et al. 2014). Evidence of ram pressure stripping of the H I and H $\alpha$  gas in infalling cluster galaxies has been observed in the form of tails, knots and filaments (e.g. Gavazzi et al. 2001; Oosterloo & van Gorkom 2005). *N*-body, smooth particle hydrodynamical simulations by Steinhauser et al. (2012), in line with previous work by Bekki & Couch (2003), Kronberger et al. (2008) and Bekki (2009), show that relatively weak ram pressure can compress the interstellar medium of the galaxy and lead to an increase of SF. By contrast, high environmental densities and strong ram pressure can remove most of the gas from the host galaxy.

Other processes, such as galaxy harassment (Moore et al. 1996), where galaxies are distorted by tidal forces, are also important. Tidal forces can be caused by the gravitational potential of the cluster or by encounters with other galaxies. The relative movement of galaxies as they fall into the cluster potential with respect to the ICM leads to a truncation of the outer galactic halo and disc. Simulations by Moore, Lake & Katz (1998) indicate that galaxy harassment of small disc galaxies in clusters produces distorted spirals, often seen in  $z \sim 0.4$  clusters (e.g. Couch et al. 1994), which evolve into spheroidal systems observed in local clusters.

Larson, Tinsley & Caldwell (1980) proposed the process of galaxy strangulation as another means of transforming field spirals into cluster ellipticals and S0s. Gas from infalling galaxies escapes its host because of tidal forces created by the cluster potential well. With a limited supply of its main ingredient, the SF in the galaxy is effectively shut-down after over the course of a few Gyr.

## 1.2 Merging clusters with shocks

Even though the majority of galaxies in relaxed clusters are quenched, recent observations suggest that vigorous SF can be observed in merging clusters. By studying a sample of  $z > 0.3$  clusters, Ebeling, Stephenson & Edge (2014) found that gas within infalling galaxies is first shock compressed and then removed from the host galaxies. Therefore, if observed at the right time, gas-rich galaxies possibly shocked by infalling into the cluster or by the passage of a shock wave can exhibit high star formation rates (SFRs).

There has been recent evidence that merging cluster processes such as increased turbulence and shocks affect the SF properties of associated galaxies. Pranger et al. (2013) find a population of quenched spirals at 3–4 Mpc distances from the core of Abell 3921, which they attribute to shocks and cluster mergers. In the post-merger cluster Abell 2384, Pranger et al. (2014) find a significant

population of disc galaxies in the cluster core, which is expected to be devoid of star-forming galaxies. Owers et al. (2012) find three star-forming tails and filaments in galaxies nearby the X-ray shock front in Abell 2744. Ferrari et al. (2003) and Umeda et al. (2004) find a significant population of luminous, H $\alpha$ -emitting, star-forming galaxies in the merging cluster Abell 521 (Ferrari et al. 2003, 2006), which hosts a radio-detected shock front (Giacintucci et al. 2008).

Galaxy formation simulations can be used to model the impact of mergers on galaxy properties. For example, models of ram pressure stripping indicate that SF in galaxies within merging clusters can be quenched (e.g. Kapferer et al. 2009). More recently, hydrodynamical simulations of shocks passing through galaxies have reproduced star-forming tails trailing behind their parent galaxy (Roediger et al. 2014). The SF persists for a few 100 Myr after the shock passage, in line with observations of star-forming tails in cluster infalling galaxies by Owers et al. (2012).

## 1.3 The ‘Sausage’ and ‘Toothbrush’ clusters

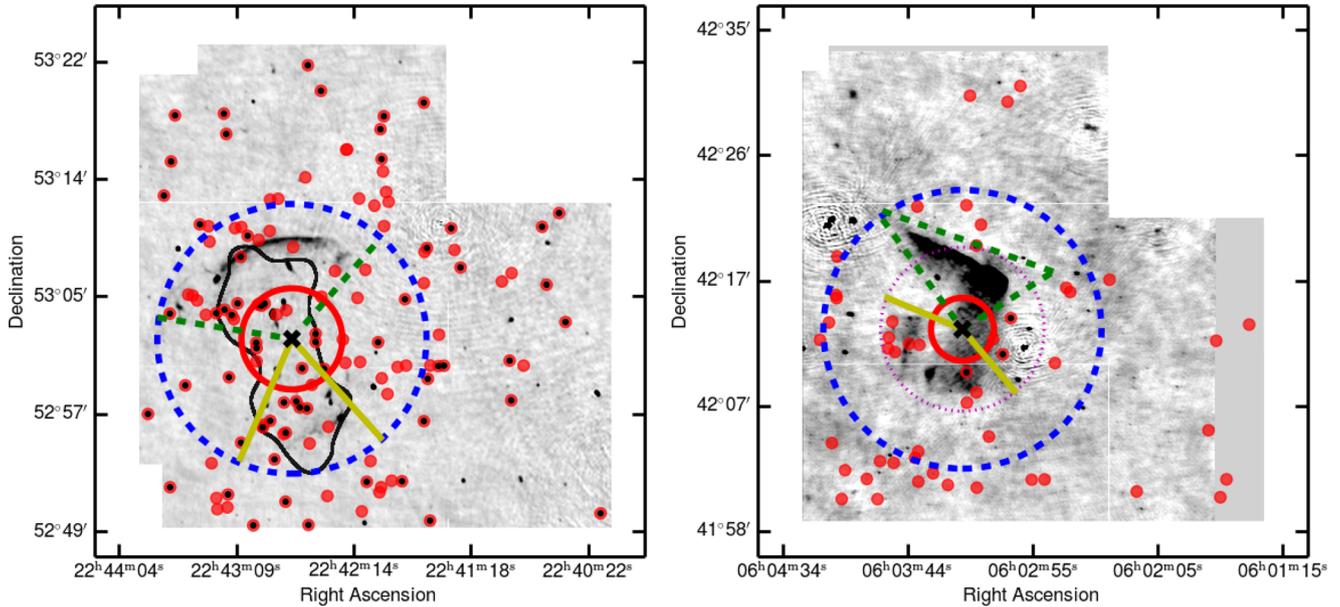
To probe the effects of shocks in transforming cluster galaxies, we started an observing campaign of two major-merging,  $z \sim 0.2$  clusters hosting some of the strongest radio shocks detected to date: CIZA J2242.8+5301 (nicknamed the ‘Sausage’; van Weeren et al. 2010) and 1RXS J0603.3+4213 (nicknamed the ‘Toothbrush’; van Weeren et al. 2012). The peculiar morphology of the relics explains the nickname of each cluster (see Fig. 1). The bright radio relics show clear signs of steepening and curving radio spectrum from the shock front into the downstream area, suggesting a scenario where the synchrotron electrons are shock accelerated and subsequently cool (van Weeren et al. 2010, 2012; Stroe et al. 2013). Nevertheless, the high-frequency 16 GHz observations of the ‘Sausage’ relic suggests a more complicated scenario in which the electrons are injected and accelerated also in the downstream area (Stroe et al. 2014b). Both clusters are massive, X-ray luminous and present elongated X-ray morphologies suggesting a merger in the plane of the sky (Akamatsu & Kawahara 2013; Ogrea et al. 2013a,b, 2014). They host two radio relics, which trace  $M \sim 2$ –4 Mach number shocks. In both clusters, one relic is significantly larger and brighter than its counterpart, suggesting the merging subclusters were at close, but not 1 : 1 mass ratio. A weak-lensing analysis by Jee et al. (2014) indicates that the ‘Sausage’ cluster is among the most massive clusters discovered to date, with a total mass exceeding of  $M_{200} > 2.5 \times 10^{15} M_{\odot}$ . The northern ( $M_N = 11.0^{+3.7}_{-3.2} \times 10^{14} M_{\odot}$ ) and the southern subclusters ( $M_S = 9.8^{+3.8}_{-2.5} \times 10^{14} M_{\odot}$ ; Jee et al. 2014) are very similar in mass. Dawson et al. (2014) derive velocity-dispersion-based mass estimates ( $M_N = 16.1^{+4.6}_{-3.3} \times 10^{14} M_{\odot}$  and  $M_S = 13.0^{+4.0}_{-2.5} \times 10^{14} M_{\odot}$ ), which are in agreement with the weak-lensing results.

On the basis of its X-ray luminosity, van Weeren et al. (2012) and Brüggén, van Weeren & Röttgering (2012) conclude the ‘Toothbrush’ is also a very massive cluster of about  $1$ – $2 \times 10^{15} M_{\odot}$ .

Hydrodynamical simulations, radio spectral modelling and an analytical dynamics analysis suggest that the ‘Sausage’ core-passage has happened  $\sim 1.0$  Gyr ago, at a relative speed of  $\sim 2000$ – $2500$  km s $^{-1}$  (van Weeren et al. 2011; Dawson et al. 2014; Stroe et al. 2014c), making it a younger merger than the possibly  $\sim 2$  Gyr old ‘Toothbrush’ merger (Brüggén et al. 2012).

## 1.4 This paper

In this paper, we aim to study the SF properties of galaxies within the ‘Sausage’ and the ‘Toothbrush’ clusters. We derive SFRs and



**Figure 1.** The emitters for the ‘Sausage’ field (left-hand panel) and the ‘Toothbrush’ field (right-hand panel) in red circles and sources selected as  $H\alpha$  emitters in the filled circles, overlaid on a Giant Meterwave Radio Telescope 323 MHz radio images in grey intensity. The red circles with black dots at the centre represent spectroscopically confirmed  $H\alpha$  emitters. The radio image is cut according to the NB FOV coverage, which was used for the selection of the emitters. The arcsectors define the areas around the radio relics which were considered for producing relic LFs. The cluster centres are defined to be at the location of the black crosses. Left: ‘Sausage’ cluster; the northern relic was captured by a section between the solid red and the dashed blue circles, bound by the dashed green lines. The southern relic area was defined between the solid red and the dashed blue circles and the solid yellow radii. The weak-lensing area, enclosing a mass of about  $1 \times 10^{15} M_{\odot}$ , is marked by the solid black curve. The non-relic areas are defined as the arcsectors between the solid red and dashed blue circles and the dashed green and solid yellow lines. The entire cluster was assumed to occupy the volume defined by the dashed blue circle, which has an  $\sim 1.85$  Mpc radius. Right: ‘Toothbrush’ cluster; the northern relic was captured by a section of the solid red circle, bound by the dashed green lines. The southern relic area was defined between the solid red and dotted purple circles and the solid yellow radii. The non-relic areas are defined as the arcsectors between the solid red and dotted purple circles and the dashed green and solid yellow lines. The cluster was assumed to occupy the volume defined by the dashed blue circle of  $\sim 2.2$  Mpc radius.

masses for the  $H\alpha$  sample and build luminosity functions (LFs) for different environments in and around the cluster volumes.

By using narrow-band (NB) observations, Stroe et al. (2014a) constrained the  $H\alpha$  LF of the cluster galaxies and found striking differences between the ‘Sausage’ and the ‘Toothbrush’ clusters. Stroe et al. (2014a) observed a notable enhancement in the normalization of the LF around the relics within the ‘Sausage’ cluster, where the travelling shock is expected to have passed 10–100 Myr ago. By contrast, the ‘Toothbrush’ cluster is almost devoid of line emitters, similar to a relaxed cluster. The relatively short time span when enhanced SF is seen in the ‘Sausage’ cluster could explain the differences found between the two clusters, in line with simulations from Roediger et al. (2014).

In Stroe et al. (2014a), lack of multiband photometry prevented the galaxy-by-galaxy separation between  $H\alpha$  emitters and higher redshift, lower rest-frame wavelength emission lines. Without such data, disentangling the drivers of the enhanced  $H\alpha$  emission in the ‘Sausage’ clusters is not possible, specifically the role of the post-merger time-scales in shaping the  $H\alpha$  properties of the galaxies. Instead, Stroe et al. (2014a) applied a statistical correction for the fraction of  $H\alpha$  emitters from the total number of emitters, which was based on deep, NB observations on the Cosmic Evolution Survey field (COSMOS; Shioya et al. 2008).

In this paper, we present an optical multiwavelength analysis of the ‘Sausage’ and ‘Toothbrush’ clusters. A combination of photometric and spectroscopic data from the Isaac Newton, William Herschel, Canada–France–Hawaii, Subaru and Keck telescopes spanning the entire optical spectrum through the  $B$ ,  $V$ ,  $g$ ,  $r$ ,  $i$ ,  $z$  bands

enables us to properly separate  $H\alpha$  emitters from higher redshift emitters. Compared to Stroe et al. (2014a), we are going 0.2–0.4 magnitudes (mag) deeper in our detection band ( $i$ ) and in the NB data, resulting in a larger sample of emitters, of which more than 50 per cent have been followed up spectroscopically and confirmed as  $H\alpha$  emitters.

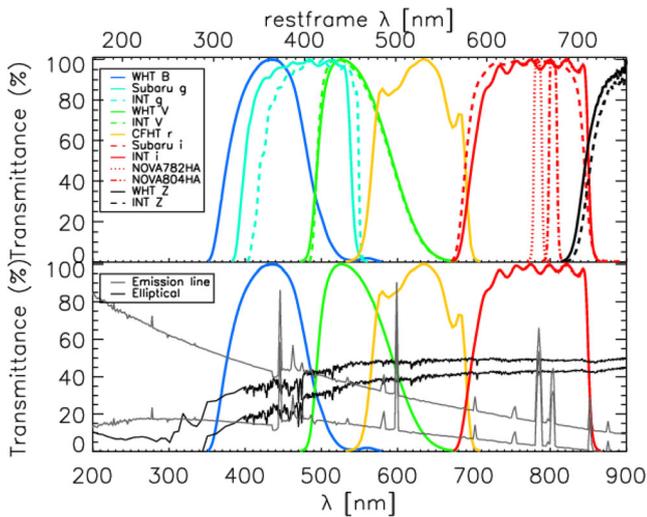
In Section 2, we present the reduction of the photometric and spectroscopic observations and the source extraction. Section 3 describes the emitter selection and separation, the LFs and properties we derive for the galaxies. In Section 4, we compare our results with those for field galaxies and galaxies hosted in other clusters. We assume a flat  $\Lambda$  cold dark matter cosmology with  $H_0 = 70.5 \text{ km s}^{-1} \text{ Mpc}^{-1}$ , matter density  $\Omega_M = 0.27$  and dark energy density  $\Omega_\Lambda = 0.73$  (Dunkley et al. 2009). We make use of Edward Wright’s online cosmological calculator (Wright 2006). 1 arcmin measures 0.191 Mpc at  $z = 0.192$  (‘Sausage’), while at  $z = 0.225$  (‘Toothbrush’) it corresponds to a physical size of 0.216 Mpc. All images are in the J2000 coordinate system. Magnitudes are in the AB system.

## 2 DATA REDUCTION AND ANALYSIS

We use a multitude of photometric and spectroscopic instruments mounted on a range of optical telescopes. We describe the data acquisition, reduction and processing below. Table 1 and Fig. 2 display the filter properties, while the observations and integration times can be found in Table 2.

**Table 1.** Filter properties: type (narrow-band, NB, or broad-band, BB), weighted central wavelength and FWHM. The redshift range  $z_{H\alpha}$  for which the  $H\alpha$  line is detected within the FWHM of the NB filters is also given.

Filter	Type	$\lambda_c$ (Å)	FWHM (Å)
NOVA782HA ( $z_{H\alpha} = 0.1865\text{--}0.2025$ )	NB	7839.0	110
NOVA804HA ( $z_{H\alpha} = 0.2170\text{--}0.2330$ )	NB	8038.5	110
INT <i>i</i>	BB	7746.0	1519
WHT <i>B</i>	BB	4332.7	1065
Subaru <i>g</i>	BB	4705.5	1393
INT <i>g</i>	BB	4857.3	1290
WHT <i>V</i>	BB	5488.1	990
INT <i>V</i>	BB	5483.4	990
CFHT <i>r</i>	BB	6257.9	1200
Subaru <i>i</i>	BB	7676.0	1555
WHT <i>z</i>	BB	8720.9	–
INT <i>z</i>	BB	8749.3	–



**Figure 2.** Top panel: normalized transmittance profiles for the filters used in the analysis. The top  $x$ -axis marks the rest-frame wavelength coverage of the filters, assuming a redshift of 0.2. Bottom panel: emission spectra for passive and emission line galaxies redshifted to  $z = 0.2$ , at arbitrary normalization, taken from the PEGASE 2.0 template set (Grazian et al. 2006). Note how the  $H\alpha$  line falls within the corresponding NB filter NOVA782HA, as expected. Both intrinsic and dust-attenuated spectra are shown. Note that the effect of Galactic dust extinction, especially in the blue side of the spectrum.

## 2.1 Observations and data processing

### 2.1.1 Isaac Newton Telescope observations

The broad-band (BB) *i* and NB NOVA782HA and NOVA804HA imaging data presented in Stroe et al. (2014a) have been supplemented with new *g*, *V*, *z*, *i* and NB data taken in four photometric nights in 2013 October–November with the Wide Field Camera (WFC)<sup>1</sup> mounted on the Isaac Newton Telescope (INT, PI Stroe).<sup>2</sup> The camera, a mosaic of four chips, has a 0.33 arcsec pixel<sup>−1</sup> scale and a square field of view (FOV) of 34.2 arcmin × 34.2 arcmin, with the top north-western corner missing. Individual exposures of

200 s for the BB and 600 s for the NB were taken in five dithered positions to cover the chip gaps, under seeing conditions varying from 0.7 to 2.0 arcsec. A total of ∼90 ks and ∼51 s were observed in the NB and ∼13 ks and ∼12 ks in the BB for the ‘Sausage’ and ‘Toothbrush’, respectively. For details see Table 2.

Note that the NB filters have a full width at half-maximum (FWHM) of 110 Å. They were designed to capture  $H\alpha$  emission ( $\lambda_{\text{restframe}} = 6562.8$  Å) at the redshift of the ‘Sausage’ and ‘Toothbrush’ clusters. See Stroe et al. (2014a) for further details.

### 2.1.2 Canada–France–Hawaii Telescope observations

Under OPTICON programme 13B055 (PI Stroe), service mode *r*-band images were taken using the Megacam imager<sup>3</sup> installed on the 3.6-m Canada–France–Hawaii Telescope (CFHT),<sup>4</sup> under excellent seeing conditions (<0.8 arcsec), between 2013 July and December. The 36-chip camera has an ∼1 deg<sup>2</sup> FOV, with a 0.187 arcsec pixel<sup>−1</sup>. To obtain a contiguous FOV coverage, 600 s exposures were taken in two dither positions spaced at 15 arcmin. 18 ks were observed in the ‘Sausage’ field and 5.4 ks in the ‘Toothbrush’.

### 2.1.3 Subaru observations

Images in the *g* and *i* band (PI Wittman) were taken with Subaru’s<sup>5</sup> Prime Focus Camera (Suprime-Cam),<sup>6</sup> a 10-chip mosaic with a 34 arcmin × 27 arcmin FOV and 0.2 arcsec pixel scale. For the full details of the observations, we refer the reader to Dawson et al. (2014) and Jee et al. (2014).

### 2.1.4 William Herschel Telescope imaging

*B*-, *V*- and *z*-band data were taken on 2013 November 1–3 with the newly-commissioned Prime Focus Imaging Platform (PFIP)<sup>7</sup> on the 4.2-m William Herschel Telescope (WHT, PI Stroe).<sup>8</sup> The single-chip camera has a pixel scale of 0.27 arcsec with a FOV of 18.0 arcmin × 18 arcmin. We took a series of 600-s exposures in three pointings, jittered over five positions, to roughly cover the FOV of INT’s WFC. The seeing varied between 0.7 and 2.0 arcsec. See Table 2 for the integration times.

## 2.2 Spectroscopy

We have also obtained spectroscopic data using two instruments. Spectra of a sample of 27 emission line, star-forming galaxies in the ‘Sausage’ and the ‘Toothbrush’ clusters were taken with the DEep Imaging Multi-Object Spectrograph (DEIMOS; Faber et al. 2003)<sup>9</sup> mounted at the Nasmyth focus of the Keck II telescope at the W. M. Keck Observatory (PI Wittman).<sup>10</sup> These observations are described in detail in Dawson et al. (2014).

Spectra of line emitters within the ‘Sausage’ cluster roughly covering the 5500–8500 Å range were obtained on 2014 July 2 and

<sup>3</sup> <http://www.cfht.hawaii.edu/Instruments/Imaging/Megacam/>

<sup>4</sup> <http://www.cfht.hawaii.edu/>

<sup>5</sup> <http://www.naoj.org>

<sup>6</sup> <http://www.naoj.org/Observing/Instruments/SCam/index.html>

<sup>7</sup> <http://www.ing.iac.es/Astronomy/instruments/pfip/index.html>

<sup>8</sup> <http://www.ing.iac.es/Astronomy/telescopes/wht/>

<sup>9</sup> <http://www2.keck.hawaii.edu/inst/deimos/>

<sup>10</sup> <http://www.keckobservatory.org/>

<sup>1</sup> [http://www.ing.iac.es/engineering/detectors/ultra\\_wfc.htm](http://www.ing.iac.es/engineering/detectors/ultra_wfc.htm)

<sup>2</sup> <http://www.ing.iac.es/Astronomy/telescopes/int/>

**Table 2.** Details of the observations: filters, total integration times, effective integration times after removing bad frames and observing dates.

Field	RA	Dec.	Filter	Int. time (ks)	Eff. time (ks)	Dates
‘Sausage’	22 <sup>h</sup> 42 <sup>m</sup> 50 <sup>s</sup>	53°06′30″	NOVA782HA	78.9	47.4	2012 October 13–15, 20–22; 2013 November 1–6
			INT <i>i</i>	13.2	9.0	2012 October 15, 20–22; 2013 November 1–6
			WHT <i>B</i>	12.0	9.6	2013 November 1 and 3
			Subaru <i>g</i>	3.3	3.3	2013 July 13
			WHT <i>V</i>	9.0	9.0	2013 November 1 and 3
			CFHT <i>r</i>	18.0	18.0	2013 July 3, 5–8, 11 and 12
			Subaru <i>i</i>	3.3	3.3	2013 July 13
			WHT <i>z</i>	10.8	9.0	2013 November 1 and 3
			INT <i>z</i>	1.0	1.0	2013 November 1–6
			‘Toothbrush’	06 <sup>h</sup> 03 <sup>m</sup> 30 <sup>s</sup>	42°17′30″	NOVA804HA
INT <i>i</i>	11.8	11.8				2012 October 15, 16 and 21; 2013 November 1–6
WHT <i>B</i>	12.0	9.6				2013 November 1 and 3
INT <i>g</i>	6.0	6.0				2013 November 1–6
WHT <i>V</i>	9.0	9.0				2013 November 1 and 3
INT <i>V</i>	2.0	2.0				2013 November 1–6
CFHT <i>r</i>	5.4	5.4				2013 December 4 and 5
WHT <i>z</i>	10.8	9.0				2013 November 1 and 3
INT <i>z</i>	5.0	5.0				2013 November 1–6

3 with the multi-object, wide-field AutoFib2 (AF2)<sup>11</sup> fibre spectrograph mounted at the prime focus of WHT (PI Stroe). A series of 30 min exposures in two configurations allowed us to observe ~100 sources within the cluster and its northern periphery. Six line emitters in this sample were also targeted with DEIMOS. The data were analysed similarly to the method described in Dawson et al. (2014). In short, the fibre traces on the CCD was corrected for curvature using the lamp flats. After bias-subtracting, flattening and sky-background subtracting the data, the spectra were extracted and wavelength-calibrated using the lamp flats and sky lines. Full details can be found in Sobral et al. (2015) and Dawson et al. (2014).

Note that the observing strategies for the DEIMOS and AF2 data were fundamentally different. The DEIMOS observations, tailored for a redshift analysis were mainly targeting the red sequence galaxies. The AF2 observations, however, specifically targeted emission line galaxies presented in this paper.

### 2.3 Photometric reduction and source extraction

We reduced the BB and NB optical photometry from the INT, WHT, Subaru and CFHT using the standard approach for reducing imaging data, implemented in our in-house PYTHON-based pipeline. We rejected data affected by cloud extinction, pointing, focusing, read-out issues and very poor seeing (>2 arcsec). The data for each chip and each filter in the WFC (INT), Suprime-Cam (Subaru) and Megacam (CFHT) CCD mosaics were processed independently. Note that the PFIP (WHT) imager contains a single CCD.

The sky flats for each filter on each instrument were median-combined to obtain a ‘master-flat’. A ‘master-bias’ for each night of observing was obtained by median-combining biases taken with each instrument. The individual exposures were bias-subtracted and sky-flattened to remove electronic camera noise, shadowing effect and normalize for the pixel quantum efficiency. Science exposure pixels that deviated by more than  $3\sigma$  from the local median were blanked as non-responsive, hot or as cosmic rays. We additionally normalized the WHT and INT *i* and *z* bands by a ‘super-flat’, obtained by combining science frames with masked sources. This step

is necessary to remove the effects of significant thin-film interference (‘fringing’) for images taken in the red and near-infrared part of the spectrum.

We used recursive rounds of SCAMP (Bertin 2006) to find astrometric solutions for the processed exposures. 0.2–0.3 arcsec root-mean-square (rms) accuracy per object was obtained by comparing source positions with 2MASS astrometry (Skrutskie et al. 2006). The exposures were normalized to the same zero-point (ZP) by referencing to the closest photometric band measured in the fourth United States Naval Observatory (USNO) CCD Astrograph Catalog (UCAC4; Zacharias et al. 2013). The fully-processed science exposures for each filter and each instrument were median-combined to obtain final stacked images using SWARP (Bertin et al. 2002).

The Sloan Digital Sky Survey (SDSS) does not cover our fields. We therefore used the USNO-B1.0 catalogue (Monet et al. 2003) to derive the photometric calibration, as outlined in Stroe et al. (2014a). The USNO-B1.0 magnitudes were converted to Johnson system *B*, *V*, *z* and Sloan system *g*, *r*, *i* based on relations derived from SDSS Data Release 7 (SDSS DR7; Abazajian et al. 2009) on a 9 deg<sup>2</sup> field with overlapping SDSS DR7 and USNO-B1.0 coverage. Bright but not saturated stars in our fields were matched to magnitudes of USNO-B1.0 sources, converted to the equivalent filter, in order to obtain the photometric ZP. Given the high number of sources matched an accuracy of ~0.05 mag was attained in the calculation of the ZP. The calibration was performed independently for the four WFC chips (INT) and the three PFIP (WHT) pointings.

We extracted sources using SExtractor (Bertin & Arnouts 1996), measuring magnitudes in 5 arcsec apertures, corresponding to a physical diameter of ~17 kpc at the redshift of the clusters. This aperture ensures that we encompass the full disc of the galaxies. Subsequently, the magnitudes were corrected for dust absorption by the Milky Way, using the reddening values from Schlafly & Finkbeiner (2011), interpolated to the effective wavelength of each of our filters. The clusters are located at low Galactic latitude and suffer from significant dust extinction ( $A_\lambda$ ) which varies across the relatively large FOV of our observations (see Appendix A, Figs A1 and A2). If uncorrected for, the dust extinction can shift galaxy *B–V* colours by up to 0.5 mag and *B–z* up to 1.5 mag.

We used the rms noise reported by SExtractor to calculate the  $1\sigma$  and  $3\sigma$  limiting magnitudes for our observations. Note that the

<sup>11</sup> <http://www.ing.iac.es/Astronomy/instruments/af2/index.html>

**Table 3.** Observed  $1\sigma$  error and  $3\sigma$  limiting magnitudes (measured in 5 arcsec apertures) for the ‘Sausage’ and ‘Toothbrush’ observations, uncorrected for the effects of Galactic dust attenuation. The extinction  $A_\lambda$  ranges for that filter are also given.

‘Sausage’	INT NB	INT $i$	INT $z$	WHT $B$	WHT $V$	WHT $z$	Subaru $g$	Subaru $i$	CFHT $r$
$1\sigma$	21.7	21.8	19.4	24.0	23.1	21.5	24.2	23.7	23.4
$3\sigma$	20.5	20.7	18.2	22.8	21.9	20.2	23.1	22.5	22.2
$A_\lambda$	0.6–1.0	0.6–1.0	0.5–0.8	1.3–2.1	1.0–1.6	0.5–0.8	1.2–1.9	0.6–1.0	0.8–1.3
‘Toothbrush’	INT NB	INT $i$	INT $g$	INT $V$	INT $z$	WHT $B$	WHT $V$	WHT $z$	CFHT $r$
$1\sigma$	21.9	22.5	23.5	22.8	20.8	24.0	23.6	21.5	23.8
$3\sigma$	20.7	21.3	22.3	22.3	19.6	22.8	22.4	20.3	22.6
$A_\lambda$	0.32–0.43	0.34–0.46	0.27–0.37	0.65–0.88	0.57–0.76	0.74–1.00	0.56–0.76	0.27–0.37	0.47–0.68

Galactic dust extinction is substantial for our field and rises steeply in the blue side of the spectrum (see Table 3).

### 3 METHODS AND RESULTS

#### 3.1 NB emitter selection

In order to select line-emitting candidates, we study the excess of the NB emission as compared to the BB continuum (Fig. 3). If an emission line is present, the source will have a significant BB–NB colour excess. We use the same approach described in detail in Stroe et al. (2014a), based on the methods of Bunker et al. (1995) and Sobral et al. (2009, 2012). We refer the interested readers to those papers for the details of the selection criteria.

The different effective central wavelengths of the NB and BB filters (see Table 1,  $\sim 100$  Å for the ‘Sausage’ and  $\sim 300$  Å for the ‘Toothbrush’) cause systematic BB–NB colour offsets. A constant offset was sufficient to correct the excess in the ‘Sausage’ field, as there was no dependence of NB–BB excess on the NB magnitude. The NOVA804HA filter peaks 300 Å redder than the INT  $i$  BB filter. Therefore, the  $i$  and  $z$  magnitudes were used to correct for the colour offsets that vary with NB magnitude. For sources without a  $z$ -band magnitude, a statistical correction was applied based on the average  $z-i$  colour. Note this is a significant improvement, greatly

reducing scatter compared to Stroe et al. (2014a), where  $z$ -band data were not available.

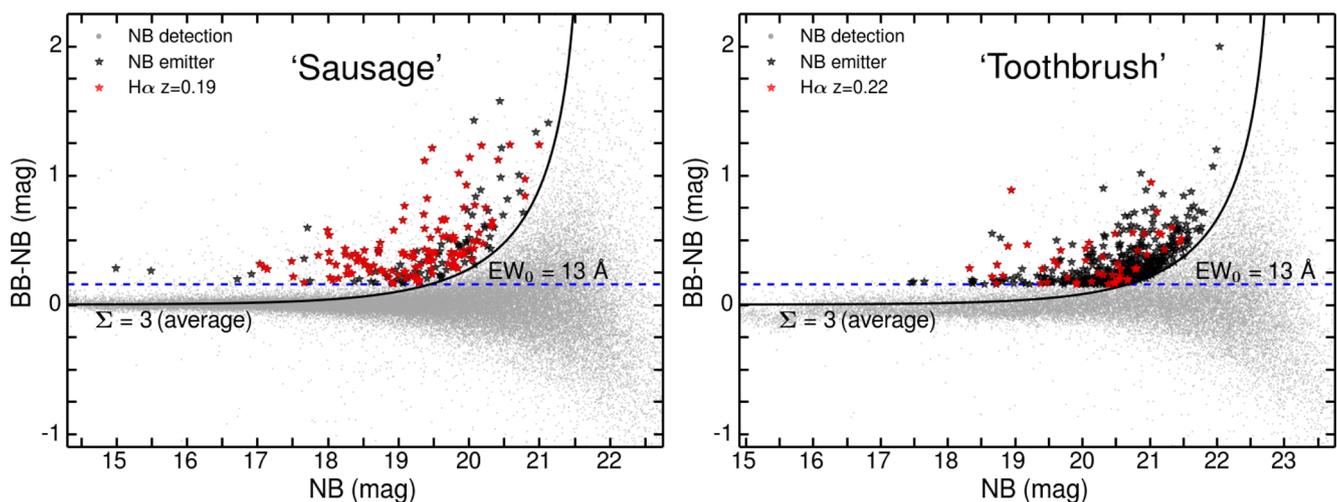
The selection of emitters is performed anew, since our new NB and BB data are deeper by 0.2–0.4 mag than in Stroe et al. (2014a), with a better colour correction, allowing us to probe fainter and lower equivalent width (EW) emitters. In short, to be selected as a line emitter, a source must fulfil three criteria (for details, see Stroe et al. 2014a).

(i) Significant excess ( $\Sigma$ ) NB emission with respect to the broadband (BB,  $\Sigma > 3$ ), based on the scatter of the faint end of the NB magnitudes. This criterion rejects faint, low-signal-to-noise sources from entering the emitter catalogue.

(ii) A NB minus BB colour cut, intended to retain sources with an intrinsic emission line EW higher than 13 Å (assuming the sources are at  $z \sim 0.2$ ). This ensures that we select sources with strong spectral features and remove stars or sources without an emission line, but which have steep continuum. The EW cut value was chosen to reflect the  $3\sigma$  scatter of the BB–NB excess around 0, for bright, but not saturated NB magnitudes.

(iii) Visual inspection to remove saturated stars, double stars and false positives at the edge of chips. These types of sources can mimic line emitters.

Our sample consists of 201 emitters for the ‘Sausage’ field (0.020 sources per kpc<sup>2</sup>, down to a line flux of  $5.2 \times 10^{-16}$  erg s<sup>-1</sup> cm<sup>-2</sup>)



**Figure 3.** BB minus NB as function of NB magnitude for the ‘Sausage’ (left) and ‘Toothbrush’ (right) fields. The INT  $i$  band was used for BB subtraction. The blue dashed, horizontal line represents the limiting rest-frame EW, while the curve marks the  $3\Sigma$  colour significance limit for choosing sources as NB emitters (masked in the black stars). The sources selected as  $H\alpha$  emitters at  $z \approx 0.2$  (according to Fig. 4) are shown in the red stars.

**Table 4.** Statistics for the emitters detected in our ‘Sausage’ and ‘Toothbrush’ surveys. A non-detection in a particular band can be caused by lack of FOV coverage or limited depth. For emitters with detections in all bands or with significant upper limits, a colour–colour selection criterion could be applied to select H $\alpha$  emitters. Additionally, spectroscopic or photometric redshifts were used to select potential H $\alpha$  sources. Note that even though the sample of emitters is larger for the ‘Toothbrush’ FOV, the number of H $\alpha$  emitters is smaller, reinforcing the effect seen in Fig. 5, where the photometric redshift distribution is dominated by high- $z$  emitters.

Field	Emitters	Detected in						Not detected		$z_{\text{spec}}$ selected	H $\alpha$ selected
		$B, g, r, i$	$B, g, i$	$B, r, i$	$g, r, i$	$B$	$g$	$r$	in $B, g, r$		
‘Sausage’	201	79	0	2	95	0	0	25	0	65	124
‘Toothbrush’	463	120	9	51	114	3	6	141	19	4	50

and 463 for the ‘Toothbrush’ field (0.036 sources per kpc<sup>2</sup>, see also Table 4, down to a line flux of  $9.8 \times 10^{-17}$  erg s<sup>-1</sup> cm<sup>-2</sup>). This is a substantial increase compared to Stroe et al. (2014a), where 181 emitters were found for the ‘Sausage’ field and 141 in the ‘Toothbrush’ FOV.

### 3.2 Identifying H $\alpha$ emitters among the line emitters

Our NB emitter population is composed of a sample of H $\alpha$  sources at  $z \sim 0.2$ , together with other strong, higher redshift line-emitters that fall within the passband of our NB filter. Nevertheless, given the moderate depth of our NB survey, we expect our emitter population to have a higher fraction of H $\alpha$  emitters compared to what is measured from deep surveys such as COSMOS, which are saturated at bright luminosities because of long individual integration times. Emission line sources strong enough to potentially be detected are: H $\beta$  ( $\lambda_{\text{rest}} = 4861$  Å), [O III] $\lambda\lambda 4959, 5007$  emitters at  $z \sim 0.61$ – $0.65$  and [O II] ( $\lambda_{\text{rest}} = 3727$  Å) emitters at  $z \sim 1.15$ . We might also be contaminated by  $z \sim 0.8$  4000 Å break galaxies.

In order to differentiate between these emitter populations, we use colour–colour separation (e.g. Shioya et al. 2008; Sobral et al. 2013), in combination with spectroscopic and photometric redshifts. In order to do so, we fully exploit the wealth of multiband photometry we have acquired. We use our INT  $i$  catalogue as our main detection catalogue (used for subtraction of the continuum, with the INT NB data). The Subaru  $i$  band catalogue, albeit deeper than the INT  $i$ , does not have perfectly matching FOV coverage to our NB observations. We note that all our emitters have a detection in the BB and NB. We further use all the data available from the other bands ranging from  $B$  up to  $z$  band. Given the different depths and FOV coverage of these data, not all emitters have detections in all seven ancillary bands. About 40 per cent of sources have detections in all bands and another  $\sim 45$  per cent miss a detection in one single band. The remaining  $\sim 15$  per cent of sources lack two or more bands.

#### 3.2.1 Colour–colour separation

We base our colour–colour selection on the  $z = 0.24$  H $\alpha$  emitters selected in COSMOS (Capak et al. 2007; Ilbert et al. 2009) from Subaru NB NB816 imaging (Shioya et al. 2008). Since COSMOS goes to much fainter magnitudes than our data, we select only emitters with line emission greater than  $2.5 \times 10^{-16}$  erg s<sup>-1</sup> cm<sup>-2</sup> to match the range observed in our survey. We explore possible colour–colour selections which best separate the low-redshift H $\alpha$  emitters from the higher redshift interlopers. We use the photometric and spectroscopic redshifts available for COSMOS to test how many emitters are correctly classified by each colour–colour selection. We adopt the  $B - g$  versus  $r - i$  colour–colour plane as best discriminator between the low- and high-redshift emitters (see Fig. 4).

Sources are selected as potential H $\alpha$  emitters if they simultaneously fulfil the two colour requirements:

$$(B - g) > (0.6(r - i) - 0.3) \quad (1)$$

$$(B - g) > (1.6(r - i) - 1.1) \quad (2)$$

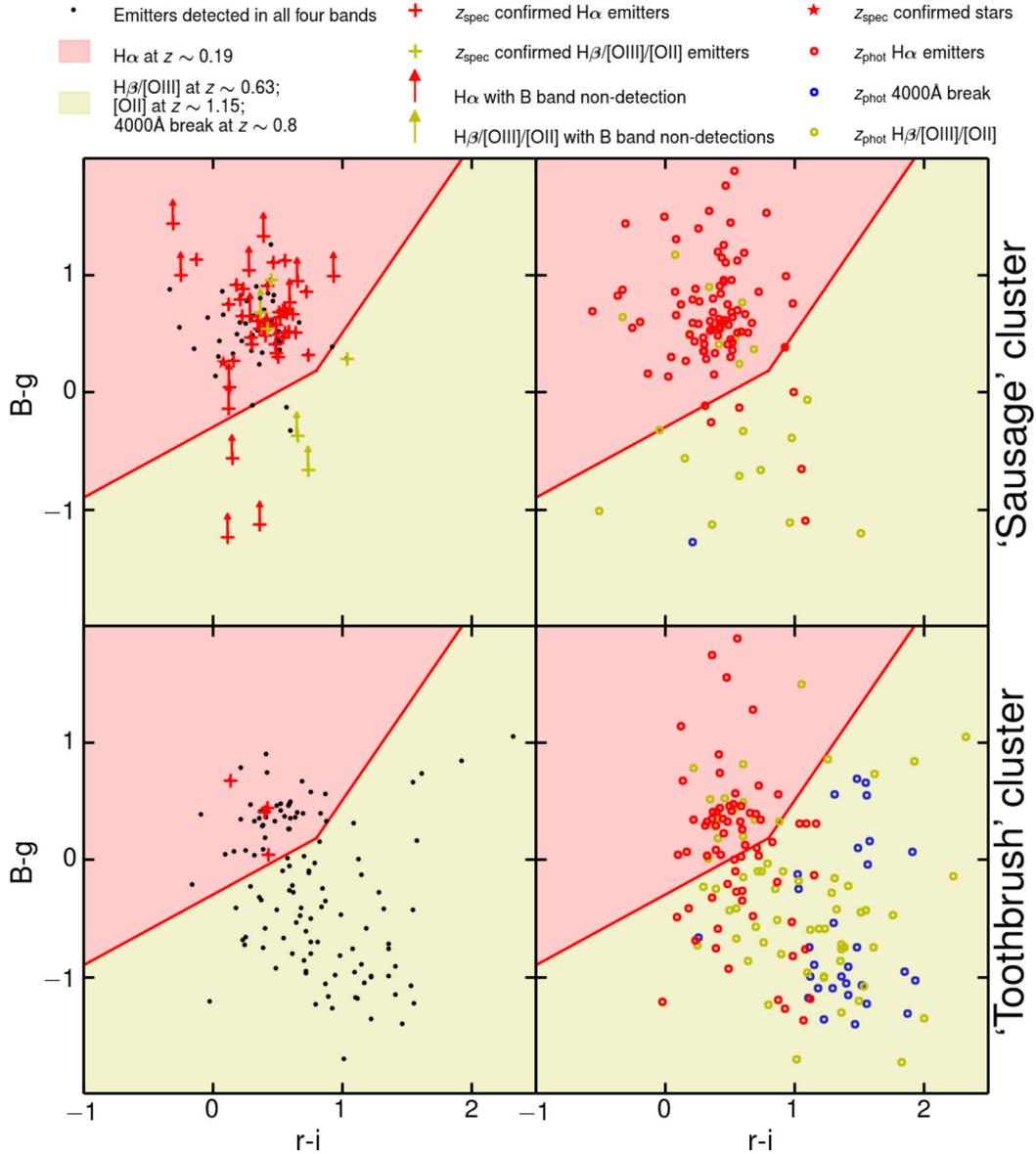
These separation lines are marked in Fig. 4 by thick red lines.

#### 3.2.2 Spectroscopic and photometric redshifts

The redshift of the sources was found by measuring the position of the brightest emission lines (e.g. H $\alpha$ , H $\beta$ , [O III] [N II] and [S III]) in each spectrum. For the ‘Sausage’ cluster, based on Keck spectra, there are 23 emitters which have been spectroscopically confirmed as H $\alpha$  line galaxies at  $z \sim 0.2$  (Sobral et al. 2015; Dawson et al. 2014). 48 H $\alpha$  sources are detected among the WHT spectra (see Sobral et al. 2015, for details). Six emitters were targeted by both surveys and are detected as H $\alpha$  in both data sets, confirming the robustness of our analysis. Therefore, we have 65 spectroscopically confirmed H $\alpha$  sources out of a sample of 201 emitters. We confirm 8 H $\beta$ /[O III] emitters at  $z \sim 0.6$ , 1 [O II] emitter at  $z \sim 1.1$  and 1 passive galaxy at  $z \sim 0.8$ . Hence, out of 75 emitters with spectra, 65 are H $\alpha$  (87 per cent) and the rest are higher redshift emitters. This low number of non-H $\alpha$  sources is partly driven by our selection of bright emitters (with a mean 5 arcsec  $i$ -band magnitude of 19.6 compared to 19.9 for the entire emitter population) for spectroscopic follow up, which have very high chances of being low rather than high redshift. In addition, the Keck spectra confirm only three stars which contaminate our emitter population and we note that we have not done any star rejection due to the lack of near-infrared data. We exclude these sources from our H $\alpha$  catalogue.

For the ‘Toothbrush’ cluster, we have only four spectroscopically confirmed H $\alpha$  emitters from Keck data. Only these four emitters were serendipitously targeted as the survey was targeting passive members.

As we do not have spectroscopy for all sources, we compute photometric redshifts. Using the comprehensive optical photometry available for these two fields ( $B, g, V, r, i, z, \text{NB}$ ), relatively precise photometric redshifts (with errors  $\Delta z_{\text{phot}}/z_{\text{phot}} < 20$  per cent) can be derived. For this purpose, we performed a grid-based redshift search between 0.01 and 1.3 with the EAZY code (Brammer, van Dokkum & Coppi 2008). Full freedom has been otherwise allowed in the fitting process of the full set of PEGASE 2.0 templates (described in Grazian et al. 2006), which includes a range of early to late-type galaxies, with a range of stellar ages. We used magnitudes measured in 5 arcsec apertures. The large apertures bias against the detection of high-redshift emitters, which are more compact and faint and are likely missed by our selection. Therefore, we expect our emitter population to be predominantly H $\alpha$  emitters at  $z \sim 0.2$ . The distribution of



**Figure 4.** Colour–colour plot of ‘Sausage’ (top plot) and ‘Toothbrush’ (bottom plot) emitters, selected according to the criteria described in Fig. 3.  $B - g$  versus  $r - i$  colours are plotted.  $H\alpha$  emitters are expected to lie above and to the left of the thick red lines, in the red shaded areas.  $H\beta$ , [O III], [O II] emitters and 4000 Å break galaxies lie in the yellow shaded area. Left-hand panels: sources with clear detections are plotted with black dots. Spectroscopically confirmed  $H\alpha$  emitters are plotted in red crosses. Right-hand panel: sources with photometric redshifts consistent with being  $H\alpha$  emitters at  $z \sim 0.2$  are plotted with red circles. Possible 4000 Å break galaxies are plotted in blue. Sources with photo- $z$ s around 0.63 (between 0.5 and 0.7) and 1.15 (between 1.0 and 1.2) are selected as  $H\beta$ , [O III] and [O II] emitters (plotted in yellow).

photometric redshifts  $z_{\text{phot}}$  can be found in Fig. 5. A natural spread in the redshifts is expected given the uncertainties in fitting photometric redshifts, especially for star-forming galaxies, which can be relatively featureless in the continuum, e.g. they do not have strong 4000 Å breaks. The majority of the line emitters (>85 per cent) are correctly fitted with templates that include line emission features, while galaxies at  $z \sim 0.8$ – $0.9$  are fitted with passive galaxy templates marked by absorption features. Therefore, not only the redshifts, but also the correct spectral type can be recovered from the template fitting.

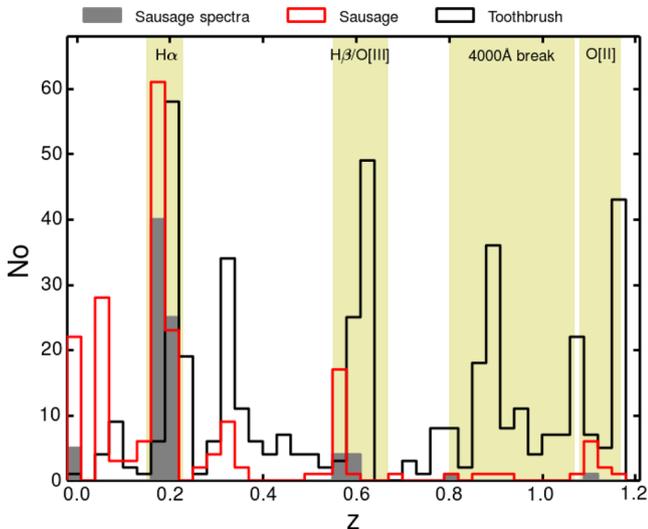
The effect of Galactic extinction (see Figs A1 and A2) is evident in the photometric redshift fitting: if we use magnitudes uncorrected for dust attenuation, the bulk of the line emitters are fitted as higher redshift passive galaxies. Hence, as explained

in Section 2.3, correcting for Milky Way dust is of the utmost importance.

Emitters are selected as  $H\alpha$  emitters if the photometric redshift lies between 0.16 and 0.23.  $H\beta$ /[O III] explains the emission if  $0.5 < z_{\text{phot}} < 0.7$ , [O II] if  $1.0 < z_{\text{phot}} < 1.2$  and 4000 Å break galaxies if  $0.7 < z_{\text{phot}} < 0.9$  (Figs 4 and 5).

### 3.2.3 $H\alpha$ emitter selection

A source is selected as an  $H\alpha$  emitter if it fulfils any of the criteria listed below (see Fig. 4). We give the number of sources selected through each criterion for each field in the parentheses. Note that the spectroscopic confirmation overlaps with the other criteria in most cases, confirming the robustness of our selection:



**Figure 5.** Redshift distribution of emitters in the ‘Sausage’ and ‘Toothbrush’ FOV, as reported photometrically by *EAZY*. Spectroscopic redshifts obtained from Keck and WHT are also overplotted. Note that the clear peaks around  $z \approx 0.2$ . The ‘toothbrush’  $H\alpha$  emitter distribution peaks at a slightly higher redshift than the ‘Sausage’ which is in line with the design of the NB filters. The ranges of higher redshift emitters possibly captured by our NB filters are marked in green. More high-redshift emitters are captured in the ‘Toothbrush’ field, given its slightly deeper data.

(i) passes the colour–colour selection (equation 1, either has clear detections in all bands  $B, g, r, i$  or an upper limit in  $B$ , i.e. the source is not detected in the  $B$  band below the detection limit in the  $B$  band, see Table 3) and a  $z_{\text{phot}}$  placing it at the cluster redshift (63 sources for the ‘Sausage’ and 24 sources for the ‘Toothbrush’ field) or

(ii) passes the colour–colour selection (either has clear detections in all bands  $B, g, r, i$  or significant upper limit in  $B$ ), but not the  $z_{\text{phot}}$  criterion due to insecure  $z_{\text{phot}}$  (the  $\chi^2$  distribution does not have a clear minimum around  $z \sim 0.2$ , with minima of similar significance at other redshifts; we selected sources with a difference greater than 0.15 between the primary redshift solution and the redshift marginalized over  $p(z|C) = \exp(-0.5\chi^2)$ , or  $z_{\text{m1}}$  as denoted in *EAZY*) (42 sources for the ‘Sausage’ and 15 sources for the ‘Toothbrush’ field) or

(iii) does not pass the colour–colour selection, but has a secure  $z_{\text{phot}}$  (the  $\chi^2$  distribution has a clear minimum around  $z \sim 0.2$ , without minima of similar significance at other redshifts, i.e. difference between the primary redshift solution and  $z_{\text{m1}}$  is lower than 0.15) (2 sources for the ‘Sausage’ and 11 sources for the ‘Toothbrush’ field) or

(iv)  $z_{\text{spec}}$  confirms it is at the cluster redshift (65 sources for the ‘Sausage’ and 4 sources for the ‘Toothbrush’ field).

The fraction of spectroscopically confirmed  $H\alpha$  emitters in the ‘Sausage’ is extremely high ( $\sim 52.5$  per cent, see Fig. 5). The location of the spectroscopically confirmed sources fully validates the colour–colour selection (Fig. 4). Based on the colour–colour selection the photometric redshift catalogue suffers from 20 per cent misclassifications. Most of these are sources where the photometric redshift probability distribution was roughly equal for classifying the source as  $H\alpha$ ,  $H\beta/[O\text{III}]$  or  $[O\text{II}]$ . In the case of the ‘Sausage’ field, out of 95 sources that would be classified as  $H\alpha$  by the photo- $z$ , 37 were targeted by spectroscopy. Out of 76 emitters confirmed through spectroscopy, 40 were also correctly classified by the photometric redshift method. Another 11 sources were assigned as

$H\beta/[O\text{III}]$  or  $[O\text{II}]$  emitters, instead of their right type. For the rest of the sources photometric redshifts between 0.05 and 0.35 were assigned. In the case of the ‘Sausage’ cluster, out of 129 potential  $H\alpha$  emitters selected through the criteria above, 5 were removed as confirmed stars of higher redshift emitters (amounting to a contamination of less than  $<4$  per cent). The emitters were also visually inspected to check for possible interlopers and obvious  $H\alpha$  emitters not selected by our method. The visual inspection indicates a rate of  $\sim 10$  per cent possible  $H\alpha$  sources not categorized as  $H\alpha$  by our method (or 90 per cent incompleteness), which is similar to Shioya et al. (2008).

We select a total of 124  $H\alpha$  emitters located at the ‘Sausage’ cluster redshift and 50 for the ‘Toothbrush’ FOV (Table 4 and Figs 1 and 4). For similar  $H\alpha$  luminosities, the typical fraction for  $H\alpha$  emitters at  $z \sim 0.2$  out of a population of emitters selected in blank fields with an NB filter is  $\sim 15$ – $20$  per cent (e.g. Shioya et al. 2008). The fraction of  $H\alpha$  emitters (based on Table 4) in the ‘Toothbrush’ FOV resembles that of blank fields (11 per cent, 50  $H\alpha$  out of 463 emitters), while in the ‘Sausage’ the fraction ( $\sim 62$  per cent, 124 out of 201) is significantly above field levels.

### 3.3 Removing $[N\text{II}]$ contamination

The  $[N\text{II}]$  forbidden line is very close in wavelength to the  $H\alpha$  line ( $\sim 20$  Å away, rest frame). Since our filters are 110 Å wide, we expect to pick up emission from both the  $H\alpha$  and  $[N\text{II}]$ . We remove the contribution to the line flux, using the relation from Sobral et al. (2012). The median contribution of  $[N\text{II}]$  to the  $H\alpha + [N\text{II}]$  flux is 0.32, consistent with solar metallicity.

### 3.4 $H\alpha$ luminosity

The  $H\alpha$  luminosity  $L_{H\alpha}$  can be calculated from the  $H\alpha$  flux  $F(H\alpha)$ , corrected for  $[N\text{II}]$ , as described in Section 3.3:

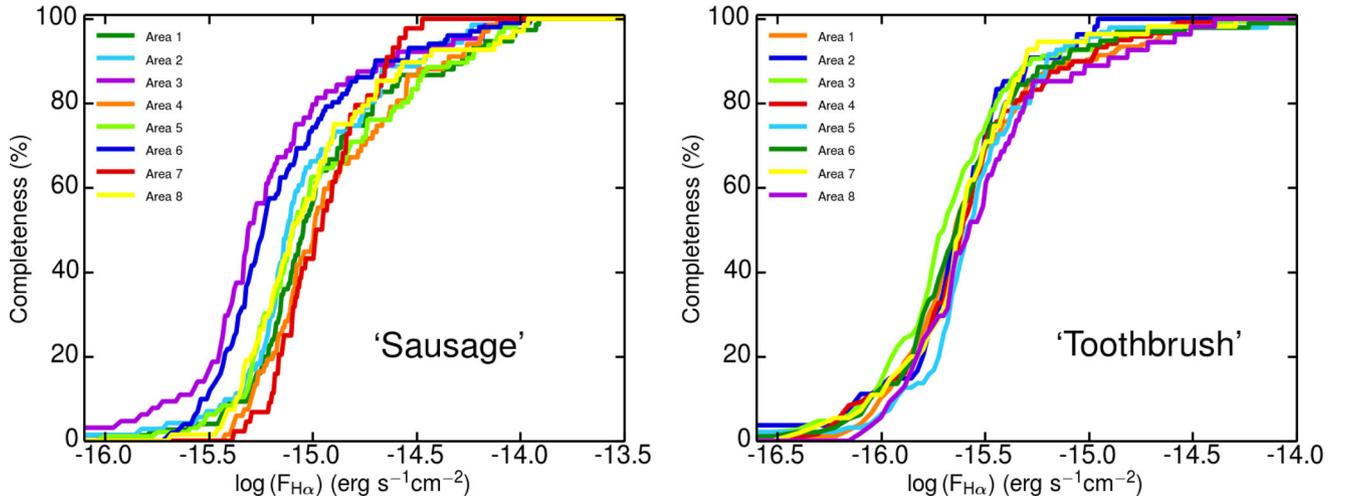
$$L_{H\alpha} = 4\pi d_L^2 F(H\alpha), \quad (3)$$

where  $d_L$  is the luminosity distance (941 Mpc for a redshift of 0.1945 and 1107 Mpc for 0.2250, respectively, for the two clusters). The emitters are binned based on their luminosity and normalized by the survey volume to form an LF.

### 3.5 Completeness correction

Fainter  $H\alpha$  emitters and the emitters with lower line EWs will not enter our  $H\alpha$  emitter catalogue given our emitter selection criteria on limiting  $\Sigma$  and EW. This results in incompleteness. We study the way our completeness rate varies as a function of line luminosity following the method of Sobral et al. (2012). For this, we pass subsamples of our emitters population through our selection criteria for  $H\alpha$  emitters described in Section 3.2.3.

We select a sample of sources, consistent with being non-emitters, but which pass our colour–colour criteria as being located at the cluster redshift. We add fake  $H\alpha$  emission lines to the flux of these galaxies and fold them through our EW and  $\Sigma$  emitter selection criteria (as shown in Section 3.1) and study the recovery fraction. We perform this study independently for eight areas, given the variation of the dust extinction across the FOV (as noted by Stroe et al. 2014a, see also Fig. A1). The variable dust extinction has a non-trivial effect on the recovery rate of  $H\alpha$  emitters, given the different way dust extinction affects the blue side and the red side of the spectrum, affecting the perceived colours of the emitters.



**Figure 6.** Survey completeness for the ‘Sausage’ (left) and ‘Toothbrush’ (right) fields as a function of  $H\alpha$  flux, performed separately for eight subareas within the FOV, to account for the varying dust extinction (see Figs A1 and A2). The areas are colour-coded according the dust extinction, from low dust extinction in indigo, though blue, green yellow, orange and red for the highest extinction. Note the correlation between completeness and amount of extinction: areas 3 and 6 in the ‘Sausage’ field have the highest level of completeness, being the least affected by dust.

The recovery rate of  $H\alpha$  emitters as function of their flux, for the eight areas of the FOV can be found in Fig. 6.

### 3.6 Volume and filter profile corrections

Assuming a top-hat (TH) shape with FWHM of  $110 \text{ \AA}$  for the NB filter transmission profiles and given the FOV coverages (see Fig. 1), we are surveying a total comoving volume of  $3.435 \times 10^3 \text{ Mpc}^3$  for the ‘Sausage’ and  $4.625 \times 10^3 \text{ Mpc}^3$  for the ‘Toothbrush’. Since our filters are not perfect THs (Fig. 2), we correct the volumes for the possible sources which might be missed in the wings of the filter as shown in Sobral et al. (2009, 2012).

### 3.7 Survey limits

We probe luminosities down to a 10 per cent completeness limit (see Fig. 6). This is equivalent to an average limiting  $H\alpha$  luminosity of  $10^{40.64} \text{ erg s}^{-1}$  for the ‘Sausage’ field and  $10^{40.14} \text{ erg s}^{-1}$  for the ‘Toothbrush’ field. The actual limiting magnitude will vary across the FOV much as the completeness varies, due to the varying dust extinction. Assuming that all the  $H\alpha$  luminosity comes from SF, we can use the Kennicutt (1998) relation, with a Chabrier (2003) initial mass function (IMF) to obtain the limiting SFR of our surveys (see equation 5 in Section 3.10). The ‘Sausage’ field average limiting SFR is  $0.17 \text{ M}_{\odot} \text{ yr}^{-1}$ , while for the ‘toothbrush’ we reach down to an average of  $0.06 \text{ M}_{\odot} \text{ yr}^{-1}$ .

### 3.8 $H\alpha$ LF and SFR

We fit LFs to different regions within the two clusters (e.g. relic areas that are aligned with the merger axis, the sides of the cluster perpendicular to the axis where there is no radio emission, weak-lensing area where most of the mass is contained etc.). For this purpose, we bin the data based on luminosity over each area. We bin the data using a range of bins. We define the bins by varying the minimum luminosity and bin width (100 random choices with uniform distribution). The errors on the  $\log \phi$  values are Poissonian. By resampling the LF in different ways, we can obtain a more robust determination of the fit parameters, which are not dominated by a

particular choice of binning. To compare with other studies, we use the popular parametrization of the LF defined by Schechter (1976):

$$\phi(L) dL = \phi^* \left( \frac{L}{L^*} \right)^{\alpha} e^{-(L/L^*)} d \left( \frac{L}{L^*} \right), \quad (4)$$

where  $L^*$  is the characteristic luminosity of the emitters where the power-law cuts off.  $\phi^*$  is the density of  $H\alpha$  emitters and provides the normalization.  $\alpha$  is the faint-end slope of the LF, which we fixed to  $-1.2$  (Stroe et al. 2014a).

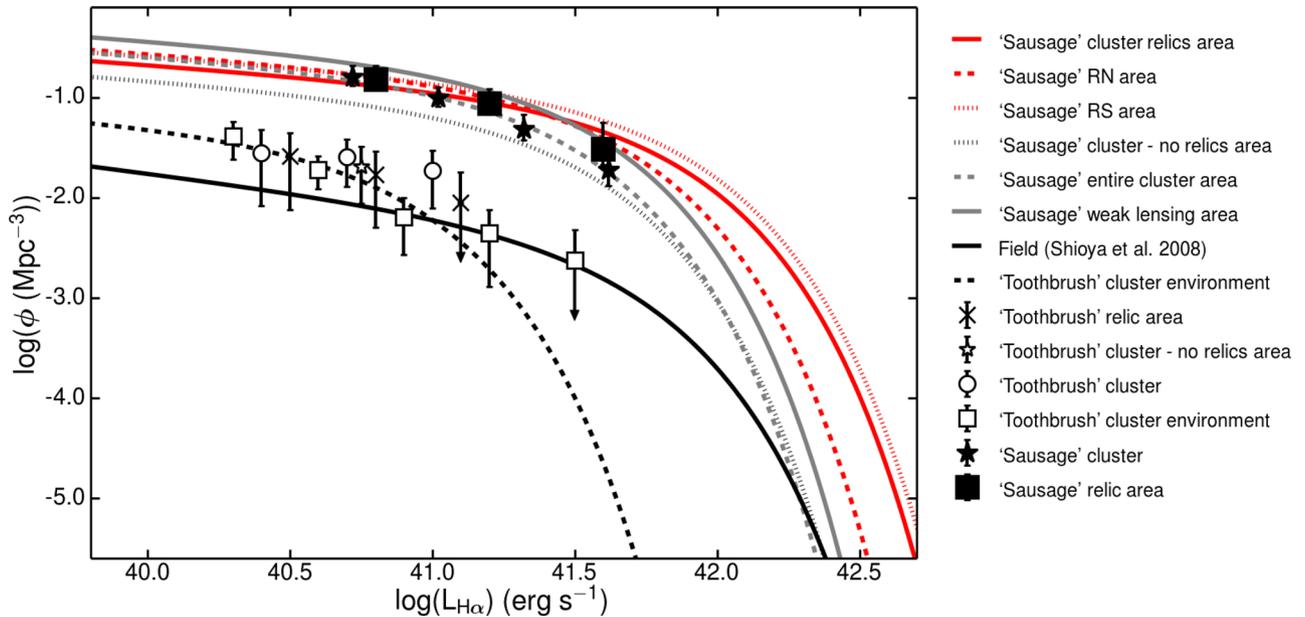
We produce a 2D distribution of number of realizations as function of LF parameters (i.e. how many of the randomly generated binnings were best fitted with a particular combination of  $\phi^*$  and  $L^*$ ). The reported  $\log \phi^*$  and  $\log L^*$  are determined as the mean of this distribution obtained by resampling the LF. The errors are  $1\sigma$  standard deviations away from the mean. The results of the fit for different regions within and around the clusters (as defined in Fig. 1) can be found in Table 5 and Fig. 7. There is a striking difference between the normalization of the two cluster LFs, as discussed in more detail in Section 4. The ‘Toothbrush’ LF is similar to a blank field, while the ‘Sausage’  $\phi^*$  is about a factor of 10 larger than that (Fig. 7).

We can use the Kennicutt (1998) conversion from  $H\alpha$  luminosity to SF activity, using a Chabrier (2003) IMF:

$$\text{SFR}(\text{M}_{\odot} \text{ yr}^{-1}) = 4.4 \times 10^{-42} L_{H\alpha} (\text{erg s}^{-1}). \quad (5)$$

### 3.9 Stellar masses

Galaxy spectral energy distributions were generated with the Bruzual & Charlot (2003) software package. We used stellar synthesis models from Bruzual (2007), a Chabrier IMF with exponentially declining SF histories and a range of metallicities. These models were fitted to the full BB data ( $BgVrIz$ ) to obtain stellar masses for the  $H\alpha$  galaxies, following the method presented in Sobral et al. (2011, 2014, for further details, see Sobral et al. 2015). A histogram of the masses of  $H\alpha$  emitters in the ‘Sausage’ and ‘Toothbrush’ field is shown in Fig. 8. The values are normalized by the volume of each survey. The masses of the ‘Sausage’  $H\alpha$  emitters are on average  $\sim 10^{9.8} \text{ M}_{\odot}$ . On average, the ‘Toothbrush’  $H\alpha$  emitters are  $\sim 3$  times



**Figure 7.** LF for the two clusters, obtained by fixing the faint-end slope  $\alpha$  to  $-1.2$  (as derived in Stroe et al. 2014a) and varying the normalization  $\log \phi^*$  and the characteristic luminosity  $\log L^*$ . The regions are defined according to Fig. 1. Downward pointing arrows represent upper limits equivalent to 1 source per volume element, per bin. The fit for a  $z = 0.24$  blank field from Shioya et al. (2008) is overlotted. Note the difference in normalization between the two clusters: the density of  $H\alpha$  emitters in the ‘Sausage’ is a factor of 10 higher than in the ‘Toothbrush’. The typical luminosity  $L^*$  of ‘Toothbrush’ emitters is lower than the blank field.

**Table 5.** Parameters of the LFs fitted with a Schechter function (see equation 4). The faint-end slope was fixed to a value of  $\alpha = -1.2$ , which was obtained by Stroe et al. (2014a) by combining data for the entire FOV of the two clusters. The fits were possible only for some areas within the cluster, as the number statistics were not in all cases high enough to allow the fitting of the independent Schechter parameters. For each area, the data were binned in range of different ways (by varying the luminosity of the first bin and the bin width) and fit independently to obtain an average, characteristic fit that best describes the LF shape. The average and standard deviation over these independent fits is reported in the table. The blank field fit for the COSMOS field by Shioya et al. (2008) is also reported. We also list  $\rho_{\text{SFR}}$  corresponding to each fit. The number of  $H\alpha$  emitters employed in each fit is given in the last column.

Field	$\alpha$	$\log \phi^*$ ( $\text{Mpc}^{-3}$ )	$\log L^*$ ( $\text{erg s}^{-1}$ )	SFRD ( $M_{\odot} \text{ yr}^{-1} \text{ Mpc}^{-3}$ )	Number of $H\alpha$	$z_{\text{spec}}$ confirmed $H\alpha$
‘Sausage’ field						
Cluster relics area	$-1.2$	$-1.37 \pm 0.04$	$41.69 \pm 0.09$	$0.22 \pm 0.05$	32	15
RN area	$-1.2$	$-1.22 \pm 0.04$	$41.51 \pm 0.06$	$0.21 \pm 0.04$	20	7
RS area	$-1.2$	$-1.29 \pm 0.04$	$41.72 \pm 0.19$	$0.29 \pm 0.147$	12	8
Weak-lensing area	$-1.2$	$-1.07 \pm 0.04$	$41.40 \pm 0.05$	$0.23 \pm 0.03$	26	19
Entire cluster area	$-1.2$	$-1.21 \pm 0.02$	$41.33 \pm 0.02$	$0.14 \pm 0.01$	59	30
Cluster – no relics area	$-1.2$	$-1.46 \pm 0.03$	$41.38 \pm 0.06$	$0.09 \pm 0.01$	14	5
‘Toothbrush’ field						
Cluster environment	$-1.2$	$-1.76 \pm 0.04$	$40.75 \pm 0.01$	$0.01 \pm 0.001$	25	4
Shioya et al. (2008)	$-1.35^{+0.11}_{-0.13}$	$-2.65^{+0.27}_{-0.38}$	$41.57^{+0.38}_{-0.23}$	$0.018^{+0.007}_{-0.004}$	–	–

less massive than those in the ‘Sausage’ ( $\sim 10^{9.3} M_{\odot}$ ). Note that this is driven by the deeper ‘Toothbrush’ data: the faintest  $H\alpha$  flux detected in the ‘Sausage’ survey is  $3.4 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2}$ , while in the ‘Toothbrush’ we probe down to  $4.2 \times 10^{-17} \text{ erg s}^{-1} \text{ cm}^{-2}$ , a factor of  $\sim 8$  deeper. Note that despite the similar volumes probed by the two surveys, there are significantly fewer massive  $H\alpha$  emitters in ‘Toothbrush’ compared to the ‘Sausage’.

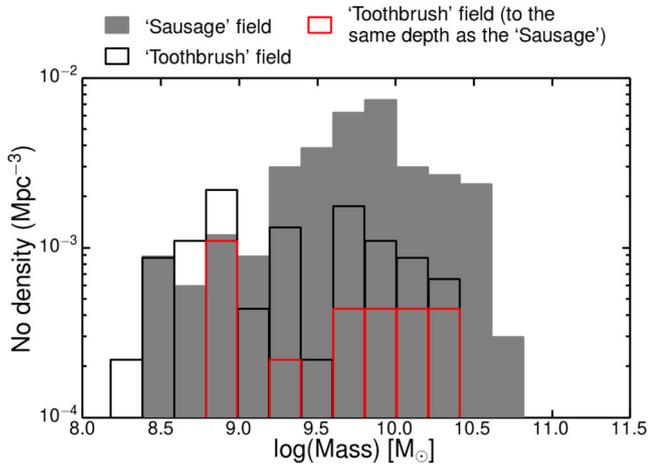
In Fig. 9, we plot the SFR versus the mass of the cluster  $H\alpha$  emitters, on top of the results for the blank field obtained from COSMOS (Shioya et al. 2008). We compute SFR for individual galaxies using the dust correction based on stellar mass (Garn & Best 2010). While our cluster  $H\alpha$  emitters fall on the SFR–mass relationship as defined from blank fields, it is important to note

that the COSMOS data were obtained over a volume  $\sim 10$  times higher than the volume probed by our ‘Sausage’ and ‘Toothbrush’ surveys. We are detecting high numbers of very high-mass, highly star-forming galaxies which are relatively rare in the field.

### 3.10 The SFR density

Given an LF and using the conversion from  $H\alpha$  luminosity and SFR, we can also calculate the SFR density  $\rho_{\text{SFR}}$  within that particular volume. The luminosity density is obtained by integrating the LF:

$$\rho_L = \int_0^{\infty} \phi(L)L dL = \Gamma(\alpha + 2)\phi^*L^*, \quad (6)$$



**Figure 8.** Mass distribution of  $H\alpha$  emitters in the fields of the two clusters, normalized to the volume of each respective survey. Note that even though the ‘Toothbrush’ data probes to  $H\alpha$  fluxes fainter by a factor of 8 compared to the ‘Sausage’, the ‘Toothbrush’ cluster is almost devoid of massive star-forming galaxies.

where  $\Gamma(n) = (n - 1)!$  is the gamma function.  $\rho_{\text{SFR}}$  is then:

$$\rho_{\text{SFR}} = \Gamma(\alpha + 2) 10^{\phi} 10^{L^*} 10^{0.4} (1 - f_{\text{AGN}}), \quad (7)$$

where we are assuming that part of the  $H\alpha$  emitters are powered by active galactic nuclei (AGNs), rather than SF. From the spectroscopic observations of the ‘Sausage’  $H\alpha$  emitters, emission line ratios (comparison of  $[\text{O III}]/\text{H}\beta$  versus  $[\text{N II}]/\text{H}\alpha$  (BPT diagram; Baldwin, Phillips & Terlevich 1981) enable the separation between star-forming galaxies and AGN (Sobral et al. 2015). This analysis indicates a fraction of 10 per cent pure AGN-powered  $H\alpha$  emitters. We thus assume that a fraction  $f_{\text{AGN}}$  of the emitters is powered by AGN activity, chosen to be 10 per cent according to Sobral et al. (2015) and in line with Garn & Best (2010) and Sobral et al. (2013). We also correct the luminosity for the dust extinction intrinsic to the line emitter. A conservative value of 1.0 mag has been shown to be appropriate by various authors (e.g. Sobral et al. 2012; Ibar et al. 2013) and has been widely assumed in the literature (e.g. Fujita et al. 2003; Ly et al. 2007; Geach et al. 2008; Sobral et al. 2009). The error on the  $\rho_{\text{SFR}}$  is  $\Delta\rho_{\text{SFR}}$ :

$$\Delta\rho_{\text{SFR}} = \rho_{\text{SFR}} \ln(10) \sqrt{(\Delta\phi)^2 + (\Delta L^*)^2}. \quad (8)$$

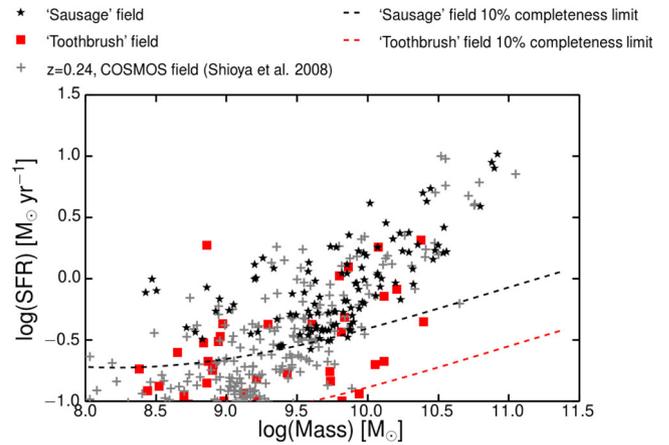
The comparison of  $\rho_{\text{SFR}}$  in and around our clusters and the field is illustrated in Fig. 10. We also overplot the value obtained by Shioya et al. (2008) for the COSMOS field. For comparison, we show the parametrization derived by Sobral et al. (2013) using a method similar to ours from a set of consistent  $H\alpha$  NB observations at four redshift slices (2.23, 1.47, 0.84 and 0.40). The dependence of the  $\rho_{\text{SFR}}$  on redshift for the blank field is

$$\log_{10} \rho_{\text{SFR}} = -2.1(1 + z)^{-1}. \quad (9)$$

We find that  $\rho_{\text{SFR}}$  for the Sausage cluster is significantly enhanced compared to the field. See Section 4 for a more detailed discussion.

## 4 DISCUSSION

The fraction of star-forming galaxies drops steeply from field environments towards the cores of massive, hot, relaxed clusters (e.g. Gómez et al. 2003). To study the effect of cluster mergers and shocks on the star-forming properties of cluster galaxies, we performed observations of two disturbed,  $z \sim 0.2$  clusters hosting radio relics,



**Figure 9.** SFR as function of stellar mass for the two clusters and the blank field (Shioya et al. 2008), using a Chabrier IMF. The Shioya et al. (2008) results for the COSMOS field were corrected for the different aperture sizes. The average 10 per cent completeness line for the ‘Sausage’ and ‘Toothbrush’ surveys is shown in the dashed lines. The SFR–mass relation holds for both the blank field and the two clusters. Note the overdensity of high-mass, high-SFR galaxies within the ‘Sausage’ cluster, which surveys a volume  $\sim 10$  times smaller than the COSMOS field. This point is illustrated further in Figs 12 and 13.

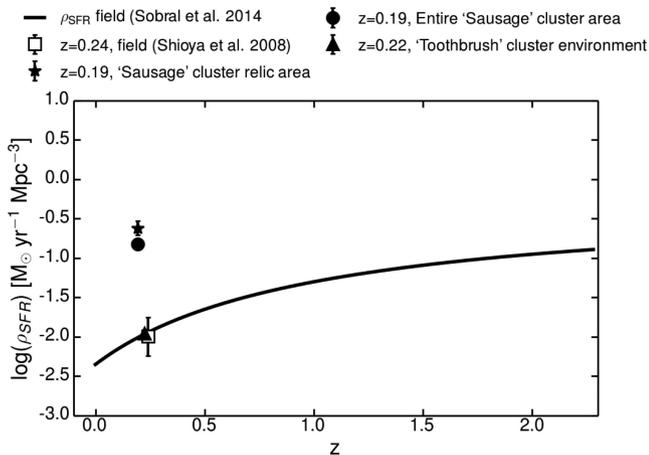
tracers of ICM shock waves of Mach number 3–4. We look at the trends in the normalization and specific luminosity of the  $H\alpha$  LF for the clusters and the masses, SFR and SFR density for the cluster star-forming galaxies. We discuss possible scenarios in which we expect differences between the  $H\alpha$  population of the two clusters.

### 4.1 $H\alpha$ LF for the clusters

#### 4.1.1 ‘Sausage’ $H\alpha$ LF

In the case of the ‘Sausage’, we build independent  $H\alpha$  LFs for the areas around the relics, along the merger axis of the cluster, and also away from the relics, perpendicular to merger axis (see Fig. 1). We study the LF for the entire cluster, by encompassing it with a circular aperture of  $\sim 1.85$  Mpc radius. This radius was chosen as function of the X-ray temperature peaks (Ogrea et al. 2013a) and the positions of the radio relics, which should be located at the cluster outskirts. A weak-lensing analysis of the cluster (Jee et al. 2014), indicates a value of  $r_{200} = 2.63$  Mpc for the radial extent, assuming that the cluster is fitted with a single Navarro–Frenk–White dark matter halo (Navarro, Frenk & White 1997). Therefore, our chosen radius for the cluster is smaller than  $r_{200}$ . We also build LFs for the area where most of the weak-lensing mass is contained, as per Jee et al. (2014). We smooth the luminosity density map with a Gaussian kernel of 125 arcsec. We consider the weak-lensing area to be contained within the contour defined by a 0.4 value of the peak in the smoothed map. This corresponds to a mass of about  $1 \times 10^{15} M_{\odot}$  (half the mass of the entire cluster) that is enclosed within this contour.

The normalization of the  $H\alpha$  LF for all areas within the ‘Sausage’ cluster is higher by a factor  $> 15$  compared to the fit for a blank field at  $z = 0.24$  by Shioya et al. (2008). The field LF parameter fits are expected to be only moderately affected by cosmic variance given the volume probed ( $3.1 \times 10^4 \text{ Mpc}^3$ ) and the small correlation length of  $r_0 = 1.8 \pm 0.2$  Mpc of the observed  $H\alpha$  emitters at  $z = 0.24$ , roughly corresponding to emitters residing in typical dark matter haloes of  $10^{11} M_{\odot}$  masses (Sobral et al. 2010). There is evidence that the

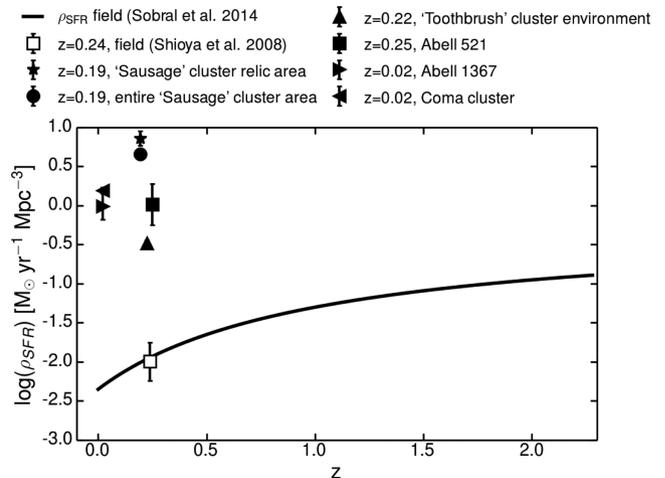


**Figure 10.**  $\rho_{\text{SFR}}$  for the ‘Sausage’ and ‘Toothbrush’ fields plotted over the  $H\alpha$  parametrization of the  $\rho_{\text{SFR}}$  history from Sobral et al. (2013), corrected for a Chabrier IMF. The entire volume spanned by the NB filter is considered. Note that the  $\rho_{\text{SFR}}$  for the ‘Sausage’ cluster relic area is at the level of  $z \sim 2.3$  field galaxies. By contrast the ‘Toothbrush’ cluster environment is at the level of  $z \sim 0.2$  field galaxies. The ‘Toothbrush’ behaves more like one would expect for a cluster, where the fraction of  $H\alpha$  galaxies steeply drops towards the cluster core. Error bars are plotted, but sometimes they are smaller than the symbol.

number of emitters within the weak-lensing area, which traces the direction of the merger, is enhanced compared to other areas within the cluster. The LF normalization  $\log \phi^*$  in the weak-lensing area is  $-1.07 \pm 0.04$  compared to an average of  $-1.37 \pm 0.04$  near the relics (difference of  $5.0\sigma$ ) and  $-1.46 \pm 0.03$  in the region away from the relics ( $7.8\sigma$ ), perpendicular to the merger axis. This fully confirms the results from Stroe et al. (2014a), where, even if a conservative, low fraction of  $H\alpha$  emitters was considered, an enhancement in the  $\log \phi^* = -1.77$  around the relics area was observed by comparison to field galaxies ( $\log \phi^* = -2.65$ ).

These very high numbers of  $H\alpha$  galaxies within the densest parts of the ‘Sausage’ cluster are surprising, given the number of studies showing that the fraction of star-forming galaxies drops steeply towards the cluster core. Even in the case of medium-redshift clusters such as the  $z = 0.81$  RXJ1716.4+6708 cluster, Koyama et al. (2010) have found that the fraction of  $H\alpha$  emitters drops steeply towards the regions of highest density, down to values of just 10 percent of the total population. The authors use a similar method to ours: an NB  $H\alpha$  selection technique to identify star-forming galaxies. A different trend is found at higher redshifts, in dynamically young clusters. For example, Hayashi et al. (2011) have found that at  $z \sim 1.5$ , the centre of a dynamically young cluster is still full of star-forming galaxies and AGN. Therefore, the ‘Sausage’ cluster presents features very similar to a young cluster at the time where the progenitors of the most massive clusters have started forming.

Note that we calculate the volume by taking the entire redshift slice captured by the NB filter. The volume occupied by the cluster which, given its projected dimension, is expected to be  $<3$  Mpc (real-space) across in the redshift direction. If we were to assume the cluster occupies a cylindrical volume, with a circular shape projected on to the sky, the cluster volume would be a maximum of  $\sim 32$  Mpc<sup>3</sup> (real-space volume). We subtract the typical level of field emitters ( $\log \phi^* = -2.65$ , as per Shioya et al. 2008) and consider all additional  $H\alpha$  emitters projected on the cluster area to be cluster members. Note that this method assumes that there are field galaxies mixed within the cluster volume, but this is not expected



**Figure 11.**  $\rho_{\text{SFR}}$  for the ‘Sausage’ and ‘Toothbrush’ clusters, plotted over the parametrization of the  $\rho_{\text{SFR}}$  history from Sobral et al. (2013, corrected for a Chabrier IMF) together with Abell 521, 1367 and the Coma cluster. We correct for the cluster volume.

to make a big difference in the final numbers. This exercise leads to an enhancement by  $\sim 1.5$  dex (a factor of 30) compared to the  $\log \phi^*$  calculated in Table 5 (0.10 versus  $-1.27$ ). If the background level of field  $H\alpha$  galaxies is not subtracted  $\log \phi^*$  is 0.12. The same enhancement would be seen in all areas if the volume was adjusted.

Even more so, there are a high number of spectroscopically confirmed  $H\alpha$  cluster members close to the relics and along the merger axis. Within the entire cluster encompassed by the large circle there are 30 confirmed cluster members, compared to 59 sources in total. This means that even if we build an  $H\alpha$  LF using the spectroscopically confirmed members and a cluster volume of  $\sim 32$  Mpc<sup>3</sup>,  $\log \phi^*$  is  $-0.18$ , significantly above the field where  $\log \phi^* = -2.65$ . The difference lies at  $>8\sigma$  significance.

We also detect a mild boost in the characteristic luminosity  $\log L^*$  of the cluster galaxies located in proximity of the relics compared to the field ( $41.69 \pm 0.09$  versus  $41.57^{+0.38}_{-0.23}$ ). By contrast, the population located away from the shocks seems to have lower luminosities on average, pointing towards a quenching of luminous  $H\alpha$  emitters in that direction. The difference between the  $H\alpha$  emitters located around the relics and away from the relics is at the level of at least  $3\sigma$ .  $H\alpha$  luminosity scales with SFR, which means the ‘Sausage’ relic emitters are slightly more star forming than field galaxies at the same redshift and cluster galaxies located away from the shock fronts. Balogh et al. (1998) found, from comparing a sample of field star-forming galaxies with galaxies in X-ray luminous,  $0.18 < z < 0.55$  clusters, that the dense cluster environment suppresses the SFR of galaxies. This is consistent with the galaxies located away from the shock fronts, which have average SFR  $1 M_{\odot} \text{ yr}^{-1}$ , compared to  $2 M_{\odot} \text{ yr}^{-1}$  for the  $H\alpha$  emitters located close to the radio relics.

Comparison with other clusters is challenging, because there are only a few studies that use a uniform NB,  $H\alpha$  selection method. Umeda et al. (2004) derive  $H\alpha$  LFs for the merging cluster Abell 521, assuming that all the  $H\alpha$  emitters captured by their filter are encompassed in a spherical volume of radius 2 Mpc (with a volume of  $33.5$  Mpc<sup>3</sup>). The volume is therefore comparable to our method. The LF normalization obtained by Umeda et al. (2004) is  $-0.25 \pm 0.20$  compared to our estimation of 0.12. Note that the authors do not subtract a background level of field galaxies, which would make

the difference between the two results more significant. The typical  $\log_{10}$  luminosity of Abell 521 is  $41.73 \pm 0.17$ , which is consistent with the values obtained for the ‘Sausage’ cluster relic areas. Note that the luminosity distance depends on the chosen cosmology and affects the derived  $H\alpha$  luminosity, when converting from  $H\alpha$  flux. To correct for this, we correct for the different cosmology used by the other authors.

Iglesias-Páramo et al. (2002) derived LFs for the local cluster Abell 1367 and Coma, which have also have a lower  $\log \phi^*$  of  $0.06^{+0.14}_{-0.12}$  and  $-0.07^{+0.03}_{-0.02}$ , estimated in the same way as Umeda et al. (2004). The typical luminosities are  $\log L^* = 41.00^{+0.07}_{-0.09}$  and  $\log L^* = 40.97^{+0.01}_{-0.02}$ , which are much lower than the overall cluster and subareas ( $\log L^* = 41.33 - 41.72$ ), at the level of  $5.0\sigma - 8.0\sigma$ .

Abell 1367 is a relatively cold, but dynamically active cluster with temperature averages for its two subclusters of  $4.2 \pm 0.3$  keV and  $3.2 \pm 0.01$  keV (using Advanced Satellite for Cosmology and Astrophysics data; Donnelly et al. 1998). Coma is a dynamically more evolved cluster with a low spiral fraction of  $\sim 13$  per cent (Bahcall 1977) and an average temperature of  $7.9 \pm 0.03$  keV (Mushotzky & Smith 1980). By contrast, in the ‘Sausage’ cluster the numerous emitters along the merger axis are located within areas of extremely hot ICM. As shown in Ogrea et al. (2013a), the cluster temperatures along the merger axis are higher than 9 keV, reaching 13 keV. The enhanced numbers of SF galaxies seems to indicate that there is an inversion of the environmental trends in the ‘Sausage’ cluster. Typically, the interaction with the hot ICM should quench the SF by removing gas from the galaxies, while the ‘Sausage’, a dynamically active cluster, has a very high fraction of star-forming galaxies.

#### 4.1.2 ‘Toothbrush’ $H\alpha$ LF

In the case of the ‘Toothbrush’, because of a lack of emitters within the cluster potential, we could not separate the cluster in multiple areas. There are no emitters in the region close to the northern relic and a few (9) close to the southern relic, with three emitters towards the right-hand side of the cluster, away from the relic areas. Compared to the blank field around the cluster there is an underdensity of emitters within the cluster. The few  $H\alpha$  emitters within the cluster region, between and around the relics are clumped together close to the southern relic. To gain enough statistics, we produce an  $H\alpha$  LF only for the cluster environment encompassed by a large circular aperture of  $\sim 2.2$  Mpc radius, chosen to reflect the locations of the X-ray peak and the radio relics. A very preliminary weak-lensing analysis of the cluster (Jee, private communication), coupled with the X-ray temperature of this cluster (Ogrea et al. 2013b), indicates the ‘Toothbrush’ cluster is similarly massive to the ‘Sausage’ cluster, suggesting a similar or slightly larger radial extent of  $r_{200} > 2.6$  Mpc. The four spectroscopically confirmed  $H\alpha$  emitters are located towards the south-west of the cluster.

The number density of emitters is typical of what is measured for the COSMOS field (Shioya et al. 2008). Over comparable volumes and down to the same  $H\alpha$  flux, the ‘Toothbrush’ FOV has only about  $\sim 12$  per cent of the number of  $H\alpha$  galaxies hosted by the ‘Sausage’. Within the cluster volumes, down to the same  $H\alpha$  flux, the ‘Toothbrush’ cluster contains only 9 emitters compared to 59 in the ‘Sausage’. In the case of the ‘Toothbrush’, there is evidence for a suppression of  $H\alpha$  emission at high luminosities, as indicated by the lower value of the specific luminosity  $\log L^*$  compared to the field (40.75 versus 41.57). Nevertheless, because of low number statistics, the error in the ‘Toothbrush’  $\log L^*$  measurement is high.

The suppression of bright  $H\alpha$  emission is consistent with the dense, hot environment in which galaxies are residing. X-ray observations of the ‘Toothbrush’ cluster from Ogrea et al. (2013b) indicate that the cluster, while highly disturbed in its temperature structure, has temperatures towards the southern part of the cluster ranging between 6 and 8 keV.

Following the procedure described in Section 4.1, we subtract the background level of field emitters and assume all other  $H\alpha$  emitters located in projection in the cluster area are actually cluster members and we adjust the volume as we did in Section 4.1 (the cluster volume is  $\sim 46$  Mpc<sup>3</sup>), pushing the LF normalization to a value of  $\log L^* = -0.41$ .

#### 4.2 SFR density for the two clusters

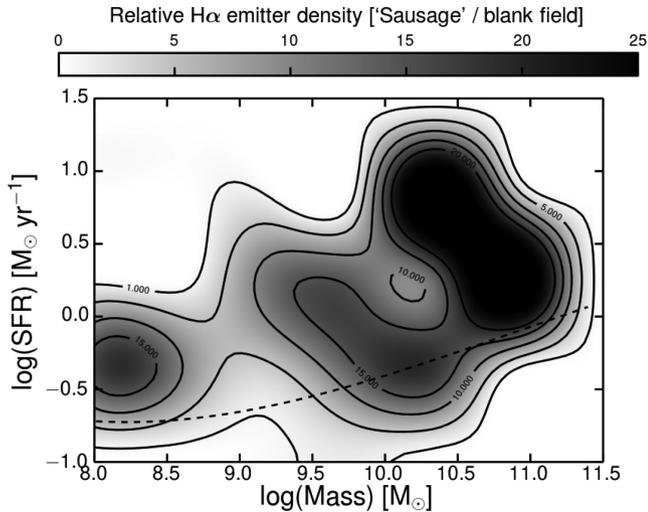
The  $\rho_{\text{SFR}}$  values for the ‘Sausage’ relic area and the entire cluster area are significantly above what is expected for galaxies at redshift 0.2 (Fig. 10). Note that these values are obtained if one uses the entire redshift span of the filter, hence using comoving volumes. The emitters in the cluster and the relic area behave on average like typical blank field galaxies at  $z \sim 2.3$ .

If we correct for the limited real-space volume the cluster is expected to occupy (as shown in Section 4.1),  $\rho_{\text{SFR}}$  for the cluster is  $6.9 M_{\odot} \text{ yr}^{-1} \text{ Mpc}^{-3}$ , more than 15 times the level of typical galaxies located at the peak of the SFR history (Fig. 11). Note that here we calculate the SFR for the cluster volume using real-space volumes, while the volume for the blank fields are calculated using comoving volumes. If we use only the spectroscopically confirmed galaxies,  $\rho_{\text{SFR}}$  is at least at the level seven times the peak of the SFR. Note however that the volume is small and the stellar-mass density added is therefore very little, equivalent to adding until  $z = 0.2$  just one galaxy with a stellar mass of about  $4.5 \times 10^{10} M_{\odot}$  (see also Section 4.4).

By contrast, the lower LF normalization and specific  $H\alpha$  luminosity in the ‘Toothbrush’ cluster drive  $\rho_{\text{SFR}}$  to very low values, consistent with  $z = 0.2$  blank field galaxies.

#### 4.3 Stellar masses of cluster galaxies

The cluster galaxies ( $10^8 - 10^{10.5} M_{\odot}$ ) fall on the mass–SFR relation (see Fig. 9). Therefore, even though an exceptionally high fraction of galaxies are star forming, given they are located in a cluster, they have SFR typical for their mass. This is in agreement with Koyama et al. (2013) that study the star-forming properties of the cluster Cl0939+4713 at  $z = 0.4$ . Nevertheless, the distribution of  $H\alpha$  galaxies within the mass–SFR relation is strikingly different from the field. To quantify the distribution of  $H\alpha$  galaxies in our two clusters in the stellar mass–SFR plane, we compare with the results from Shioya et al. (2008), obtained over volumes  $\sim 10$  times larger than the survey of the ‘Sausage’ and ‘Toothbrush’ clusters. We bin the  $H\alpha$  emitters in mass and SFR and divide the 2D histogram by the different volumes probed by the three surveys. This ensures we obtain comparable results for the clusters and the blank field. We then divide the normalized mass–SFR distribution of emitters for the clusters by the normalized  $H\alpha$  distribution for the COSMOS blank field. The bicubically-interpolated results for the ‘Sausage’ and the ‘Toothbrush’ fields can be visualized in Figs 12 and 13, respectively. The ‘Sausage’ field contains elevated numbers of  $H\alpha$  emitters at all masses and SFRs. The cluster is exceptionally abundant in high-mass ( $> 10^{10} M_{\odot}$ ), highly star-forming ( $1 - 10 M_{\odot} \text{ yr}^{-1}$ ) galaxies. The ‘Sausage’ has a density of  $H\alpha$  emitters 20–25 times higher than a blank field. Note that towards low SFRs, our results suffer



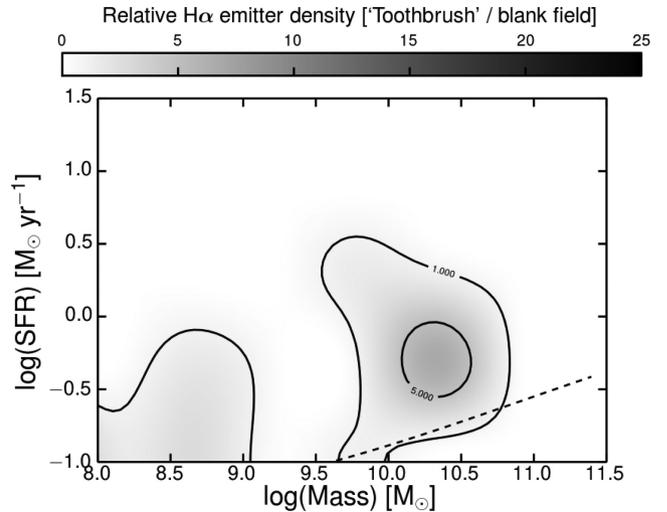
**Figure 12.** Relative density of H $\alpha$  emitters of the ‘Sausage’ field compared to a blank field from Shioya et al. (2008), plotted in the stellar mass–SFR plane. The 10 per cent completeness in the SFR measurements is plotted in the dashed line. The ‘Sausage’ field has a higher density of H $\alpha$  emitters at all SFRs and stellar masses, as compared to a blank field. The ‘Sausage’ cluster contains up to 15–20 times more highly star-forming, high-mass galaxies than a blank field.

from incompleteness so the relative abundance of H $\alpha$  emitters is underestimated. In conclusion, the enhancement in the ‘Sausage’ H $\alpha$  emitter density is seen at all masses, but particularly at the high masses. The ‘Toothbrush’ field H $\alpha$  emitters density is consistent with results obtained from blank fields, with mild overdensities of relatively high-mass, low-SFR H $\alpha$  emitters ( $\sim 10^{10} M_{\odot}$ ,  $\sim 0.3 M_{\odot} \text{ yr}^{-1}$ ). However, this enhancement in one mass–SFR bin could easily be explained by Poissonian errors and/or cosmic variance.

#### 4.4 Cause for enhanced SF in the ‘Sausage’ cluster

Both the ‘Sausage’ and the ‘Toothbrush’ clusters have undergone mergers that induced travelling shock waves to propagate through the ICM and accelerate ICM electrons to relativistic speeds. In the case of the ‘Sausage’ cluster, a binary merger leads to the formation of two unequal relics along the merger axis (van Weeren et al. 2011). The northern relic has a stronger Mach number than the southern one by a factor of  $\sim 2$  (Stroe et al. 2013). In the ‘Toothbrush’ cluster, simulations and X-ray data indicate a more complicated merger scenario with a main binary merger, possibly followed by a smaller mass interloper participating in the merger towards the end (Brüggen et al. 2012; Ogarean et al. 2013b). The northern relic is also larger than the relic located towards the south-east side of the cluster. Simulations indicate that the ‘Toothbrush’ could also be a more advanced merger than the ‘sausage’ with a core passage time of the main subclusters happening  $\sim 2$  Gyr ago, compared to  $\sim 0.5$ –1 Gyr ago for the ‘Sausage’ (van Weeren et al. 2011; Brüggen et al. 2012).

Through our H $\alpha$  observations we discover and confirm numerous star-forming galaxies nearby the relics in the ‘Sausage’ cluster and along the merger axis. Also in the ‘Toothbrush’ cluster there is a clump of star-forming galaxies in the vicinity of the southern shock. The concentration of emitters close to the southern relic in the ‘Toothbrush’ shows that the shock front might have passed



**Figure 13.** Relative density of H $\alpha$  emitters of the ‘Toothbrush’ field compared to a blank field from Shioya et al. (2008), plotted in the stellar mass–SFR plane. The 10 per cent completeness in the SFR measurements is plotted in the dashed line. Note the stark difference with the ‘Sausage’ field. The ‘Toothbrush’ field contains a lower or similar number of emitters as compared to blank fields at all masses and SFRs. The slight enhancement up to a factor of 5 can be explained by Poissonian errors or cosmic variations. The ‘Toothbrush’ not only follows the SFR–mass relation, but also has the same density of H $\alpha$  emitters as blank fields in each part of the relation.

through a group of gas-rich galaxies. There is indication that the cluster is a more complicated merger with a third smaller cluster responsible for the straight northern relics and that could also cause the formation of the southern relic, at a later stage than the northern one. If this smaller subclump was a low-mass group and still contained numerous gas-rich galaxies that could explain why there is a striking difference between the galaxy populations towards the northern and the southern relic in the ‘Toothbrush’ cluster.

Another point to note is that the distribution of H $\alpha$  emitters around the clusters is very different. The environment of the ‘Toothbrush’ cluster is at the level expected for the field, while the ‘Sausage’ clusters seems to be embedded in a more SFR-rich region with filamentary H $\alpha$  structures.

The gravitationally decoupled shock travels ahead of the ICM gas, dark matter and galaxies and therefore could have interacted with these components at some point since the subclusters merged. In terms of its interaction with galaxies, a shock wave will heat any available intragalactic gas and increase turbulence and induce instabilities in the gas. The gas can then collapse into clouds which are dense enough to start forming stars. The scenario was proposed by Rees (1989) for high-redshift galaxies where the powerful radio jet was suggested to induce SF along its propagation axis. Radio jets with high Mach numbers ( $M \sim 10$ –100) at their terminating shocks have too much power and entrain the low-entropy gas away from the galaxy uplifting it in buoyant bubbles (e.g. Werner et al. 2010). Werner et al. (2010) found H $\alpha$  filaments in the immediate downstream area from the jet-termination shock front in M87. Crockett et al. (2012) and Hamer et al. (2014) found evidence for shock-induced shells around star-forming clumps. The locations of these shells along the optical filaments connecting the two radio jets in the Centaurus A galaxy is consistent with material being swept along the back flow of the jet and ionized through shocks. In the case of the high-redshift galaxy 4C+41.17, Steinbring

(2014) confirmed the presence of a bow-shock structure inducing SF near the core of the galaxy. In the context of interacting galaxies, smooth particle hydrodynamical simulations by Barnes (2004) show that a shock-induced SF prescription fits very well observations of NGC 4676. Saitoh et al. (2009) similarly found in their simulations a vigorous, shock-induced starburst event at the interface of two colliding disc galaxies. By contrast to radio-jet terminating shocks, the merger shocks in the ‘Sausage’ and the ‘Toothbrush’ cluster have Mach numbers of at most 4. O’grea et al. (2014) have shown using *Chandra* data that minor shock fronts are ubiquitous within the ICM of the ‘Sausage’ cluster.

We speculate that low Mach number, extended shocks might have a higher chance of not removing the gas from the galaxy, but increasing the turbulence. We suggest that there is a tight correlation between the strength of the shock and the ultimate fate of the gas. If the shock is too strong, then the gas can be stripped from the host galaxy in the same fashion as it happens when a galaxy is infalling into the cluster potential with relative speeds of 1000–2000 km s<sup>-1</sup>. For example, there is evidence that the galaxies in the ‘Sausage’ cluster are moving along the merger axis, following the merger direction (Stroe et al. 2013). Given the speed at which the collisionless shock front travels with respect to the ICM (~2500 km s<sup>-1</sup>; Stroe et al. 2014c) and the relative collision speed (~2250 km s<sup>-1</sup>; Dawson et al., in preparation), the galaxies seem to be trailing behind the shock front, but are ahead of the ICM, indicating they are possibly moving at a speed of 500–2000 km s<sup>-1</sup> with respect to the ICM gas. This is comparable or lower than the speeds reached by galaxies infalling into the cluster via accretion. Note that there is a velocity dispersion of ~1000 km s<sup>-1</sup>, so some galaxies will be going much faster/slower relative to the shock.

The shock-induced SF interpretation is in line with simulations by Roediger et al. (2014). They show that SF lasting for up to a few 100 Myr can happen with a delay of ~10 Myr after a passage of the shock. The newly-born stars build up with a tail trailing the shock direction. Multiple episodes of SF can happen after the passage of the shock and therefore a gradient of ages within these tails or between galaxies located at different distance from the shock is not necessarily expected. Given the delayed start of SF, our H $\alpha$  selection might miss the galaxies very close to the shock fronts, where the SF might not have started yet.

A requirement for the shock and merger to increase the SFR is that the galaxies within the subclusters are still relatively gas rich. The H $\alpha$  emitters in the ‘Sausage’ cluster are relatively massive (significant numbers of star-forming galaxies with stellar masses over 10<sup>10</sup> M $\odot$ , see Fig. 8). The shock is expected to traverse a galaxy within a very short time-scale of about 10–50 Myr. Hence, the shock quickly induces turbulence in the gas, after which the gas cools and collapses. The high rate of SF following the shock passage can quickly deplete the gas reservoir. Part of the gas fuels SF, while the rest, is removed from the galaxy through strong outflows. A fraction of the gas located towards the outer disc of the galaxy could be easily stripped by the shock. We therefore expect the passage of the shock to lead to a steep rise in SF for 10–100 Myr, followed by a quick quenching of the galaxy and a shut-down in the formation of new stars.

As shown in Sobral et al. (2015), there is evidence for strong outflows in some of the ‘Sausage’ cluster galaxies with asymmetric and P-Cygni profiles in the H $\alpha$ , [N II] and [S II] lines and broad components for the [N II] and [S II] forbidden lines. This suggests that even though galaxies are highly star forming as we observe them, they will quickly evolve into gas-poor galaxies. Given the galaxies reside in a very massive cluster, any outflowing material will easily

escape the host galaxy. The high mass-loss rate caused by these outflows leads to a quick depletion of the gas reservoir effectively shutting down SF. Hence, the outflows indicate that H $\alpha$  galaxies will be transformed into passive galaxies within a short time-scale. We can make an estimate of the rate at which this can happen by comparing the two clusters. The ‘Toothbrush’ cluster hosts about 60 per cent (34) fewer H $\alpha$  galaxies than the ‘Sausage’ cluster. If the time since core passage is 2 Gyr for the ‘Toothbrush’ compared to 1 Gyr for the ‘Sausage’ cluster, SF in these galaxies is being shut-down at a rate of 34 galaxies per Gyr, or one galaxy every ~30 Myr. If no additional accretion of fresh matter happens, the ‘Toothbrush’ will have a complete shut-down of SF within 1 Gyr of when we observe it. ‘Right now’, the ‘Toothbrush’ galaxy population is fully passively-evolving. For the ‘Sausage’, the process of SF shut-down will take another 2 Gyr. Given that the stellar mass of the ‘Sausage’ galaxies is 10<sup>8–10.5</sup> M $\odot$ , and assuming a molecular gas  $M_{\text{gas}}$  to gas plus stellar mass  $M_{\text{gas}} + M_*$  ratio of 0.1 (Young & Scoville 1991), our galaxies have about 10<sup>7–9.5</sup> M $\odot$  molecular gas. The molecular gas has two possible fates: it either contributes to SF or it is removed from the host galaxy, either via outflows or ram pressure stripping. The characteristic H $\alpha$  luminosity of the ‘Sausage’ galaxies of 10<sup>41.6</sup> erg s<sup>-1</sup> is equivalent to a rate of ~3 M $\odot$  yr<sup>-1</sup> of molecular gas being converted into new stars. If all the molecular gas is consumed through SF, the low-mass galaxies would use up their molecular gas within 1 Myr, while the galaxies located at the high-mass end would require 1 Gyr. Assuming a maximal mass-loss through outflows equal to the rate of conversion to SF, SF in the ‘Sausage’ galaxies would last up to 0.5 Gyr.

Another calculation we can make using the difference of age between the clusters is the quantity of stellar mass being added to the cluster volume. We assume that the two clusters are relatively similar in their properties apart from the ‘Toothbrush’ being an older merger compare to the ‘Sausage’. In this scenario, the ‘Toothbrush’ represents a ‘look into the future’ of a ‘Sausage’-like cluster in 0.5 Gyr. The ‘Sausage’ cluster  $\rho_{\text{SFR}}$  is ~6.9 M $\odot$  yr<sup>-1</sup> Mpc<sup>-3</sup> (corrected for the volume occupied by the cluster), while the ‘Toothbrush’ is only ~0.3 M $\odot$  yr<sup>-1</sup> Mpc<sup>-3</sup>, a drop which could happen, given the simulations of the two clusters, over 0.5 Gyr. Within every Mpc<sup>3</sup>, a mass of ~1.5 × 10<sup>9</sup> M $\odot$  of stars would be formed. Given the entire cluster volume is expected to be about 30 Mpc<sup>3</sup>, the total stellar mass added to the cluster is ~4.5 × 10<sup>10</sup> M $\odot$ , less than 3 per cent of the mass of the Milky Way. A similar mass is lost through outflows. Therefore, even though many galaxies within merging clusters can go through episodes of vigorous SF, once enough time passes, this will not necessarily reflect in the total mass of passive galaxies.

As an interpretation for our results, we suggest that even in a hot cluster atmosphere, as long as galaxies retain some of their gas content until the passage of the merger shock, it is possible to observe high levels of SF in a large fraction of cluster galaxies. The influence of shocks seems to supersede the passive evolution of galaxies within the cluster environment and prevent the rapid loss of gas through interactions with the ICM. Rather, the shock favours the retainment of gas within its host galaxy. This is consistent with the finding that the number of high-mass star-forming galaxies is especially boosted in the ‘Sausage’ cluster. High stellar-mass galaxies reside in massive dark matter haloes with a strong gravitational pull capable of holding the gas. Shocks probably induce instabilities in the gas which collapses into star-forming clouds and also increase AGN activity that produces large mass outflows. Therefore, while momentarily the galaxies close to the shock front will exhibit high levels of H $\alpha$  emission, the fast consumption of their gas will lead to an accelerated evolution from gas-rich to gas-poor

ellipticals or S0s, compared to other cluster galaxies allowed to passively evolve.

## 5 CONCLUSIONS

We conducted an  $H\alpha$  survey of two post-core passage, merging clusters which host Mpc-wide travelling shock waves (the ‘Sausage’ and the ‘toothbrush clusters’). Using optical BB and NB data and spectroscopy, we are able to draw a number of conclusions.

(i) We robustly select line emitters towards the two clusters using custom-made NB filters. Down to a similar EW, but different luminosity limits, we select 201 and 463 line emitters towards the ‘Sausage’ and the ‘Toothbrush’ FOV, respectively. We separate between  $H\alpha$  emitters at the cluster redshift and other higher redshift emitters using colour–colour diagnostics, photometric and spectroscopic data.

(ii) Based on photometric redshifts, the emitter population for the ‘Sausage’ cluster is clearly dominated by  $H\alpha$  emitters with a fraction of 62 per cent, 52.5 per cent of which are confirmed by spectroscopy. The bulk of the emitters are located in the cluster.

(iii) In the case of the ‘Toothbrush’, 89 per cent of the emitters are not  $H\alpha$ . The bulk of the  $H\alpha$  emitters are located in the field environment around the cluster.

(iv) We find a clear enhancement in the number density of  $H\alpha$  emitters in the ‘Sausage’ compared to a blank field  $\log \phi^*$  ( $-1.37 \pm 0.04$  versus  $-2.65^{+0.27}_{-0.38}$ ), pointing towards a very high fraction of the cluster population being star forming. This is a highly surprising result given the high X-ray temperature of the ICM (9–13 keV). The results also hold if we look only at emitters around the shock fronts or within the cluster volume (as defined from weak-lensing data). The SFR density for the cluster is of the order of 15 times the peak of the SF history of the Universe. The average SFR of galaxies along the merger axis is higher than away from the shock fronts. The cluster has a density of high-mass, highly star-forming galaxies 20–25 higher than blank fields.

(v) The normalization  $\log \phi^*$  of the  $H\alpha$  LF in the ‘Toothbrush’ cluster is  $-1.76 \pm 0.02$ , compared to  $-2.65^{+0.27}_{-0.38}$  for a blank field. Owing to a specific luminosity  $\log L^* = 40.75 \pm 0.01$  in the cluster lower than the field ( $41.57^{+0.38}_{-0.23}$ ), the ‘Toothbrush’ overall SFR density is consistent with blank fields at  $z = 0.2$ .

(vi) We find that the relation between the SFR and stellar mass for the two clusters is very similar to that of blank fields. However, the density of  $H\alpha$  emitters in the ‘Sausage’ cluster is boosted compared to the field, especially in the high-mass, high-SFR regime, where the cluster is 20–25 times denser than the COSMOS field.

(vii) Accounting for the different ages of the two clusters, we measure a rate of one star-forming galaxy being transformed into a non-star-forming galaxy every  $\sim 30$  Myr to transform a SF rich, ‘Sausage’-like cluster into a ‘Toothbrush’-like cluster devoid of star-forming galaxies.

(viii) We interpret our results as shock-induced SF. In line with simulations and other observational results, the merger and shock waves lead to a momentary increase in SF in gas-rich galaxies. This is turn accelerates the turn-off of SF owing to a rapid consumption of the molecular gas supply. This effects seems to be happening at all masses and SFR, but predominantly in the high-mass regime.

## ACKNOWLEDGEMENTS

We thank the anonymous referee for his/her comments that helped improving the clarity of the paper. We thank Florian Pranger,

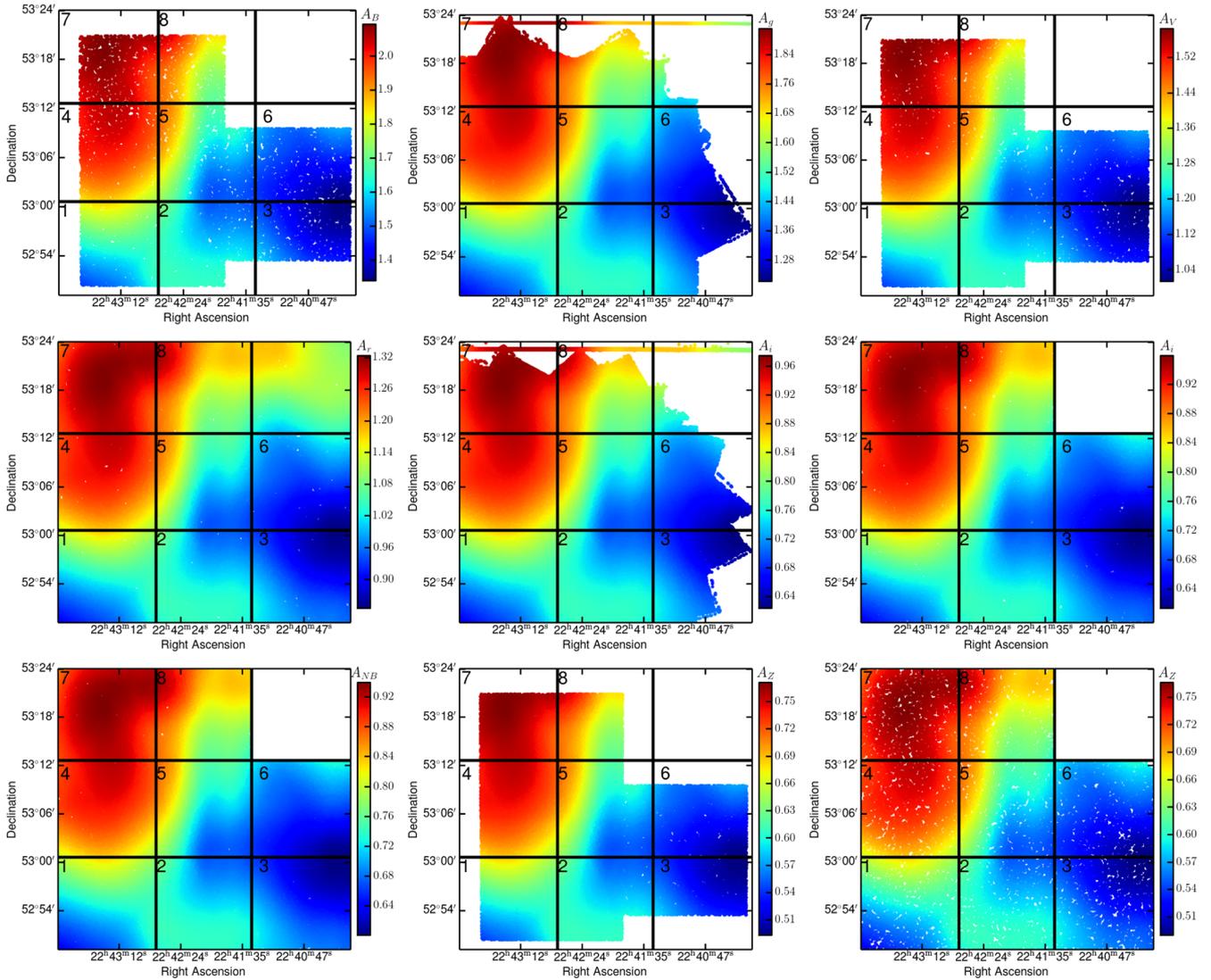
David Carton, Francois Mernier, Monica Turner and Mattia Fumagalli for useful discussions. Based on observations made with the INT (proposals I12BN003 and I13BN006) and the WHT (proposal W13BN006, W14AN012) operated on the island of La Palma by the Isaac Newton Group in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias. Based in part on data collected at Subaru Telescope, which is operated by the National Astronomical Observatory of Japan. Also based on observations obtained through the OPTICON programme 13B055 with MegaPrime/MegaCam through, a joint project of CFHT and CEA/DAPNIA, at the CFHT which is operated by the National Research Council (NRC) of Canada, the Institut National des Sciences de l’Univers of the Centre National de la Recherche Scientifique of France, and the University of Hawaii. The research leading to these results has received funding from the European Community’s Seventh Framework Programme (FP7/2007-2013 and FP7/2013-2016) under grant agreement numbers RG226604 and 312430 (OPTICON). Some of the data presented herein were obtained at the W.M. Keck Observatory, which is operated as a scientific partnership among the California Institute of Technology, the University of California and the National Aeronautics and Space Administration. The Observatory was made possible by the generous financial support of the W.M. Keck Foundation. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. This research has made use of NASA’s Astrophysics Data System. AS acknowledges financial support from the Netherlands Organisation for Scientific Research (NWO). DS acknowledges financial support from NWO through a VENI fellowship, from FCT through an FCT Investigator Starting Grant ad Start-up Grant (IF/01154/2012/CP0189/CT0010) and from FCT grant PEst-OE/FIS/UI2751/2014. RJvW acknowledges support provided by NASA through the Einstein Postdoctoral grant number PF2-130104 awarded by the Chandra X-ray Center, which is operated by the Smithsonian Astrophysical Observatory for NASA under contract NAS8-03060. Part of this work performed under the auspices of the US DOE by LLNL under Contract DE-AC52-07NA27344.

## REFERENCES

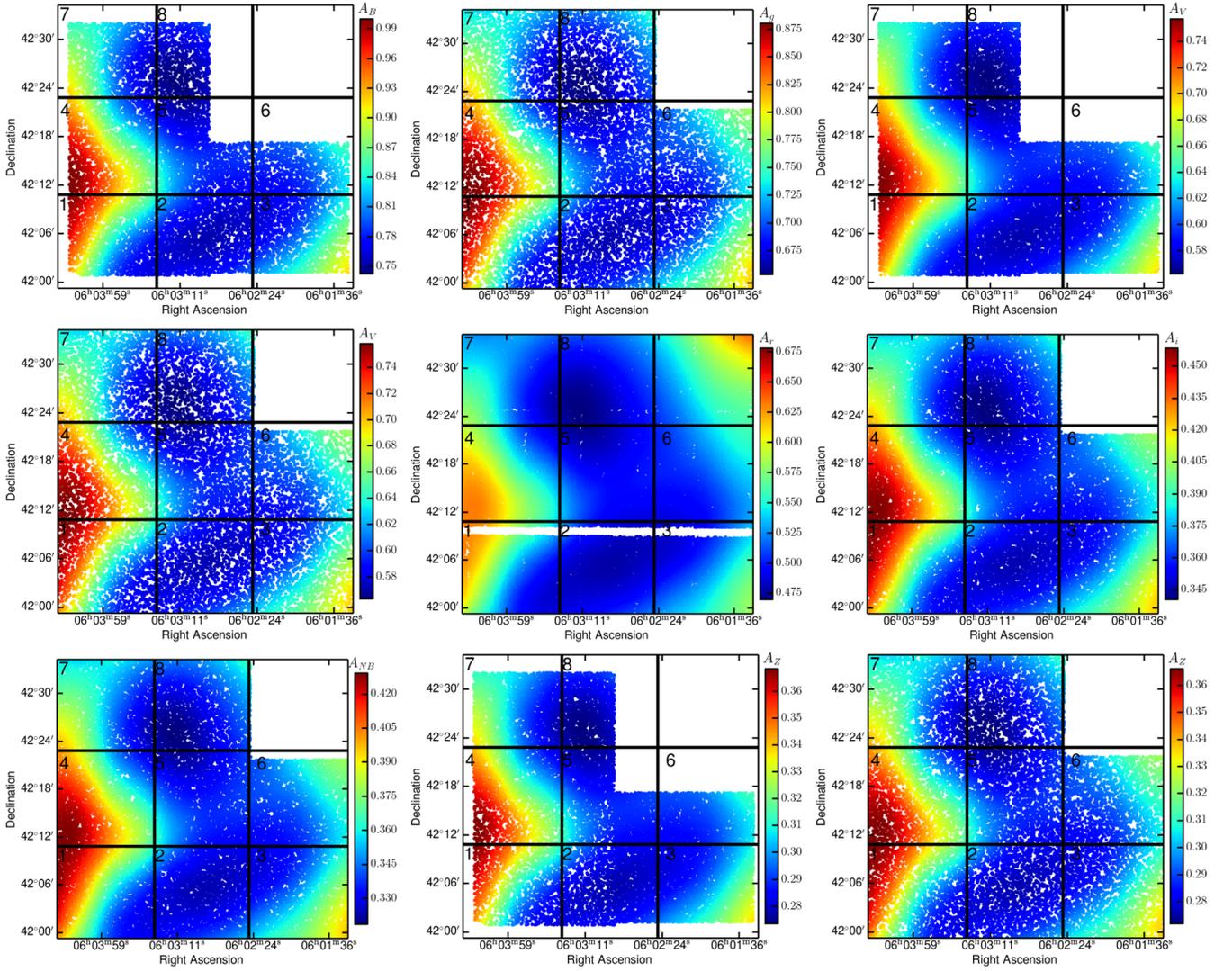
- Abazajian K. N. et al., 2009, *ApJS*, 182, 543  
 Akamatsu H., Kawahara H., 2013, *PASJ*, 65, 16  
 Bahcall N. A., 1977, *ApJ*, 218, L93  
 Baldwin J. A., Phillips M. M., Terlevich R., 1981, *PASP*, 93, 5  
 Balogh M. L., Schade D., Morris S. L., Yee H. K. C., Carlberg R. G., Ellingson E., 1998, *ApJ*, 504, L75  
 Balogh M. L., Couch W. J., Smail I., Bower R. G., Glazebrook K., 2002, *MNRAS*, 335, 10  
 Barnes J. E., 2004, *MNRAS*, 350, 798  
 Bekki K., Couch W. J., 2003, *ApJ*, 596, L13  
 Bekki K., 2009, *MNRAS*, 399, 2221  
 Bertin E., Arnouts S., 1996, *A&AS*, 117, 393  
 Bertin E., Mellier Y., Radovich M., Missonnier G., Didelon P., Morin B., 2002, in Bohlender D. A., Durand D., Handley T. H., eds, *ASP Conf. Ser. Vol. 281, Astronomical Data Analysis Software and Systems XI*. Astron. Soc. Pac., San Francisco, p. 228  
 Bertin E., 2006, in Gabriel C., Arviset C., Ponz D., Solano E., eds, *ASP Conf. Ser. Vol. 351, Astronomical Data Analysis Software and Systems XV*. Astron. Soc. Pac., San Francisco, p. 112  
 Brammer G. B., van Dokkum P. G., Coppi P., 2008, *ApJ*, 686, 1503  
 Brügger M., van Weeren R. J., Röttgering H. J. A., 2012, *MNRAS*, 425, L76

- Brunetti G., Jones T. W., 2014, *Int. J. Mod. Phys. D*, 23, 30007
- Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000
- Bruzual G., 2007, in Vallenari A., Tantaló R., Portinari L., Moretti A., eds, *ASP Conf. Ser. Vol. 374, From Stars to Galaxies: Building the Pieces to Build Up the Universe*. Astron. Soc. Pac., San Francisco, p. 303
- Bunker A. J., Warren S. J., Hewett P. C., Clements D. L., 1995, *MNRAS*, 273, 513
- Butcher H., Oemler A., Jr, 1978a, *ApJ*, 219, 18
- Butcher H., Oemler A., Jr, 1978b, *ApJ*, 226, 559
- Capak P et al., 2007, *ApJS*, 172, 99
- Cayatte V., van Gorkom J. H., Balkowski C., Kotanyi C., 1990, *AJ*, 100, 604
- Chabrier G., 2003, *PASP*, 115, 763
- Couch W. J., Ellis R. S., Sharples R. M., Smail I., 1994, *ApJ*, 430, 121
- Couch W. J., Balogh M. L., Bower R. G., Smail I., Glazebrook K., Taylor M., 2001, *ApJ*, 549, 820
- Crockett R. M. et al., 2012, *MNRAS*, 421, 1603
- Dawson W. A. et al., 2014, preprint ([arXiv:1410.2893](https://arxiv.org/abs/1410.2893))
- Donnelly R. H., Markevitch M., Forman W., Jones C., David L. P., Churazov E., Gilfanov M., 1998, *ApJ*, 500, 138
- Dressler A., 1980, *ApJ*, 236, 351
- Dunkley J. et al., 2009, *ApJS*, 180, 306
- Ebeling H., Stephenson L. N., Edge A. C., 2014, *ApJ*, 781, L40
- Faber S. M. et al., 2003, *Proc. SPIE*, 4841, 1657
- Ferrari C., Maurogordato S., Cappi A., Benoist C., 2003, *A&A*, 399, 813
- Ferrari C., Arnaud M., Ettori S., Maurogordato S., Rho J., 2006, *A&A*, 446, 417
- Fujita S. S. et al., 2003, *ApJ*, 586, L115
- Fumagalli M., Fossati M., Hau G. K. T., Gavazzi G., Bower R., Sun M., Boselli A., 2014, *MNRAS*, 445, 4335
- Garn T., Best P. N., 2010, *MNRAS*, 409, 421
- Giacintucci S. et al., 2008, *A&A*, 486, 347
- Gavazzi G., Boselli A., Mayer L., Iglesias-Paramo J., Vilchez J. M., Carrasco L., 2001, *ApJ*, 563, L23
- Geach J. E., Smail I., Best P. N., Kurk J., Casali M., Ivison R. J., Coppin K., 2008, *MNRAS*, 388, 1473
- Gómez P. L. et al., 2003, *ApJ*, 584, 210
- Goto T., Yamauchi C., Fujita Y., Okamura S., Sekiguchi M., Smail I., Bernardi M., Gomez P. L., 2003, *MNRAS*, 346, 601
- Grazian A. et al., 2006, *A&A*, 449, 951
- Gunn J. E., Gott J. R., III, 1972, *ApJ*, 176, 1
- Hamer S., Salomé P., Combes F., Salomé Q., 2014, *A&A*, doi:10.1051/0004-6361/201424808
- Hayashi M., Kodama T., Koyama Y., Tadaki K.-I., Tanaka I., 2011, *MNRAS*, 415, 2670
- Ibar E. et al., 2013, *MNRAS*, 434, 3218
- Ilbert O. et al., 2009, *ApJ*, 690, 1236
- Iglesias-Paramo J., Boselli A., Cortese L., Vilchez J. M., Gavazzi G., 2002, *A&A*, 384, 383
- Jee M. J. et al., 2014, preprint ([arXiv:1410.2898](https://arxiv.org/abs/1410.2898))
- Kapferer W., Sluka C., Schindler S., Ferrari C., Ziegler B., 2009, *A&A*, 499, 87
- Kennicutt R. C., Jr, 1998, *ARA&A*, 36, 189
- Kodama T., Balogh M. L., Smail I., Bower R. G., Nakata F., 2004, *MNRAS*, 354, 1103
- Koyama Y., Kodama T., Shimasaku K., Hayashi M., Okamura S., Tanaka I., Tokoku C., 2010, *MNRAS*, 403, 1611
- Koyama Y. et al., 2013, *MNRAS*, 434, 423
- Kronberger T., Kapferer W., Ferrari C., Unterguggenberger S., Schindler S., 2008, *A&A*, 481, 337
- Larson R. B., Tinsley B. M., Caldwell C. N., 1980, *ApJ*, 237, 692
- Ly C. et al., 2007, *ApJ*, 657, 738
- Monet D. G. et al., 2003, *AJ*, 125, 984
- Moore B., Katz N., Lake G., Dressler A., Oemler A., 1996, *Nature*, 379, 613
- Moore B., Lake G., Katz N., 1998, *ApJ*, 495, 139
- Mushotzky R. F., Smith B. W., 1980, *Highlights Astron.*, 5, 735
- Navarro J. F., Frenk C. S., White S. D. M., 1997, *ApJ*, 490, 493
- Ogrea G. A., Brüggén M., Röttgering H., Simionescu A., Croston J. H., van Weeren R., Hoeft M., 2013a, *MNRAS*, 429, 2617
- Ogrea G. A., Brüggén M., van Weeren R. J., Röttgering H., Croston J. H., Hoeft M., 2013b, *MNRAS*, 433, 812
- Ogrea G. A., Brüggén M., van Weeren R., Röttgering H., Simionescu A., Hoeft M., Croston J. H., 2014, *MNRAS*, 440, 3416
- Oosterloo T., van Gorkom J., 2005, *A&A*, 437, L19
- Owers M. S., Couch W. J., Nulsen P. E. J., Randall S. W., 2012, *ApJ*, 750, L23
- Pranger F. et al., 2013, *A&A*, 557, A62
- Pranger F., Böhm A., Ferrari C., Maurogordato S., Benoist C., Höller H., Schindler S., 2014, *A&A*, 570, A40
- Rees M. J., 1989, *MNRAS*, 239, 1 p
- Roediger E., Brüggén M., Owers M. S., Ebeling H., Sun M., 2014, *MNRAS*, 443, L114
- Saitoh T. R., Daisaka H., Kokubo E., Makino J., Okamoto T., Tomisaka K., Wada K., Yoshida N., 2009, *PASJ*, 61, 481
- Sarazin C. L., 2002, in Ferretti L., Gioia I. M., Giovannini G., eds, *As-trophysics and Space Science Library, Vol. 272, Merging Processes in Galaxy Clusters*. Kluwer, Dordrecht, p. 1
- Schechter P., 1976, *ApJ*, 203, 297
- Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Shioya Y. et al., 2008, *ApJS*, 175, 128
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Sobral D. et al., 2009, *MNRAS*, 398, 75
- Sobral D., Best P. N., Geach J. E., Smail I., Cirasuolo M., Garn T., Dalton G. B., Kurk J., 2010, *MNRAS*, 404, 1551
- Sobral D., Best P. N., Smail I., Geach J. E., Cirasuolo M., Garn T., Dalton G. B., 2011, *MNRAS*, 411, 675
- Sobral D., Best P. N., Matsuda Y., Smail I., Geach J. E., Cirasuolo M., 2012, *MNRAS*, 420, 1926
- Sobral D., Smail I., Best P. N., Geach J. E., Matsuda Y., Stott J. P., Cirasuolo M., Kurk J., 2013, *MNRAS*, 428, 1128
- Sobral D., Best P. N., Smail I., Mobasher B., Stott J., Nisbet D., 2014, *MNRAS*, 437, 3516
- Sobral D. et al., 2015, *MNRAS*, in press
- Steinbring E., 2014, *ApJ*, in press
- Steinhauser D., Haider M., Kapferer W., Schindler S., 2012, *A&A*, 544, A54
- Stroe A., van Weeren R. J., Intema H. T., Röttgering H. J. A., Brüggén M., Hoeft M., 2013, *A&A*, 555, A110
- Stroe A., Sobral D., Röttgering H. J. A., van Weeren R. J., 2014a, *MNRAS*, 438, 1377
- Stroe A. et al., 2014b, *MNRAS*, 441, L41
- Stroe A., Harwood J. J., Hardcastle M. J., Röttgering H. J. A., 2014c, *MNRAS*, 445, 1213
- Umeda K. et al., 2004, *ApJ*, 601, 805
- van Weeren R. J., Röttgering H. J. A., Brüggén M., Hoeft M., 2010, *Science*, 330, 347
- van Weeren R. J., Brüggén M., Röttgering H. J. A., Hoeft M., 2011, *MNRAS*, 418, 230
- van Weeren R. J., Röttgering H. J. A., Intema H. T., Rudnick L., Brüggén M., Hoeft M., Oonk J. B. R., 2012, *A&A*, 546, A124
- Werner N. et al., 2010, *MNRAS*, 407, 2063
- Wright E. L., 2006, *PASP*, 118, 1711
- Young J. S., Scoville N. Z., 1991, *ARA&A*, 29, 581
- Zacharias N., Finch C. T., Girard T. M., Henden A., Bartlett J. L., Monet D. G., Zacharias M. I., 2013, *AJ*, 145, 44

## APPENDIX A: DUST EXTINCTION SCREENS



**Figure A1.** ‘Dust screens’ for the ‘Sausage’ for the different filters from top-left to bottom-right: WHT  $B$ , WHT  $V$ , Subaru  $g$ , CFHT  $r$ , Subaru  $i$ , INT  $i$ , INT NB, WHT  $z$ , INT  $z$ . The figures are all on the same RA and Dec. scale and serve to display the FOV coverage with each camera. The values of the dust attenuation are calculated for each source in the FOV and all sources are plotted as individual points. All filters cover the full extent of the cluster. Note that the NB extinction is effectively the same as for the  $i$ -band filter, as the NB filter is centred close to the central wavelength of the  $i$ -band filter. The dust extinction, measured in magnitudes, is based on measurements from Schlafly & Finkbeiner (2011). Note the different scales of the figures: the dust extinction and its variation across the FOV increase significantly towards the blue side of the optical spectrum. The dust extinction variations across the FOV lead to differences in depth across the FOV; therefore, the FOV has been divided into eight areas for completeness study purposes.



**Figure A2.** The same as Fig. A1, but for the ‘Toothbrush’ field. The figures are in order from top-left to bottom-right: WHT  $B$ , INT  $g$ , WHT  $V$ , INT  $V$ , CFHT  $r$ , INT  $i$ , INT NB, WHT  $z$ , INT  $z$ . Note that because the dust attenuation is plotted for individual sources, the lower source density than the ‘Sausage’ field is apparent.

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