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A multiplicity survey of the ρ Ophiuchi molecular clouds^{★,★★}

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Abstract. We present a volume-limited multiplicity survey with magnitude cutoff ($m_K \leq 10.5$ mag) of 158 young stellar objects located within or in the vicinity of the ρ Ophiuchi Dark Cloud. With exception of eleven already well observed objects, all sources have been observed by us in the K-band with 3.5 m telescopes by using speckle techniques. The separation range covered by our survey is $0.13'' \leq \theta \leq 6.4''$, where the lower limit is given by the diffraction limit of the telescopes and the upper limit by confusion with background stars. The multiplicity survey is complete for flux ratios ≥ 0.1 ($\Delta m_K \leq 2.5$) at the diffraction limit. After taking the background density into account the degree of multiplicity is $29.1\% \pm 4.3\%$ and thus only marginally higher than the value $23.5\% \pm 4.8\%$ derived for the given separation range for the main-sequence solar-like stars in the solar neighbourhood (Duquennoy & Mayor 1991). We discuss the implications of these findings.

Key words. stars: pre-main-sequence – binaries: visual – infrared: stars – surveys – techniques: interferometric

1. Introduction

The detection of an overabundance by a factor of two of binaries among the young stars in Taurus when compared to the results for the main sequence (Ghez et al. 1993; Leinert et al. 1993; Reipurth & Zinnecker 1993) made it very clear that binarity indeed is the dominant mode of star formation. Consequently, in the years after these studies both theoretical and observational work on binaries among young stars and binary formation was intensified. Observationally, two main routes were followed: studying the fraction of binaries in associations and in young clusters, both with the aim to learn about the conditions which influence the preference of binary over single star formation. The study of associations, all of which were about at the same distance of ≈ 150 pc and are about equally young (several million years) has so far not given a clear picture, see e.g. the summary by Duchêne (1999): the duplicity is high in Taurus (see references above), CrA (Ghez et al. 1997; Reipurth & Zinnecker 1993) and Scorpius (Köhler et al. 2000), while it almost corresponds to main-sequence values or is even lower in Chamaeleon and Lupus (Brandner et al. 1996; Reipurth & Zinnecker 1993; Köhler 2001). To identify the reason for this different behaviour will need further observational studies and continued discussions on the interpretation, although Durisen & Sterzik (1994) proposed a possible

explanation, namely that fragmentation should lead to lower fractions of binaries for higher initial cloud temperature.

The situation is somewhat more settled in the case of clusters. The advantage here is that clusters of different age can be studied in order to get information on the temporal evolution. The result is that even the youngest of them, the Trapezium, does not show an overabundance of binaries (Prosser et al. 1994; Petr et al. 1998; Padgett et al. 1997). It is true that N -body simulations, e.g. Kroupa (1995) indicate that in dense clusters the fraction of binaries could be reduced by gravitational interactions within 1 million of years from “high” to “normal”. But the assumption that the lower fraction of binaries in dense clusters may be intrinsic and determined by the density as a parameter remains an attractive hypothesis (Duchêne et al. 1999).

Although there is evidence for an overabundance of multiple systems in the Ophiuchus star forming region compared to the main-sequence (Ghez et al. 1997; Simon et al. 1995; Duchêne 1999), the statistics are based only on a small number of systems observed with various techniques. To derive a survey for the ρ Ophiuchi cloud complex (Sect. 2) that is comparable to our surveys of the Taurus star forming region (Leinert et al. 1993; Köhler & Leinert 1998) we created a magnitude-limited sample (Sect. 3) based on previous work that determined the cloud membership of our targets. With exception of some well-studied objects we observed our complete survey by using speckle techniques at 3.5 m telescopes (Sect. 4). To reduce the data we used a software package (Sect. 5) developed in our group during the last years. After the correction of the raw data (Sect. 6) we discuss the results in terms of age and

[★] Based on observation with the New Technology Telescope (NTT, proposals 65.I-0067 and 67.C-0354) and the 3.6 m telescope (proposal 65.I-0086) at the European Southern Observatory (ESO), La Silla, Chile and the 3.5 m telescope at Calar Alto, Spain.

^{★★} Table 2 and Appendices A and B are only available in electronic form at <http://www.edpsciences.org>

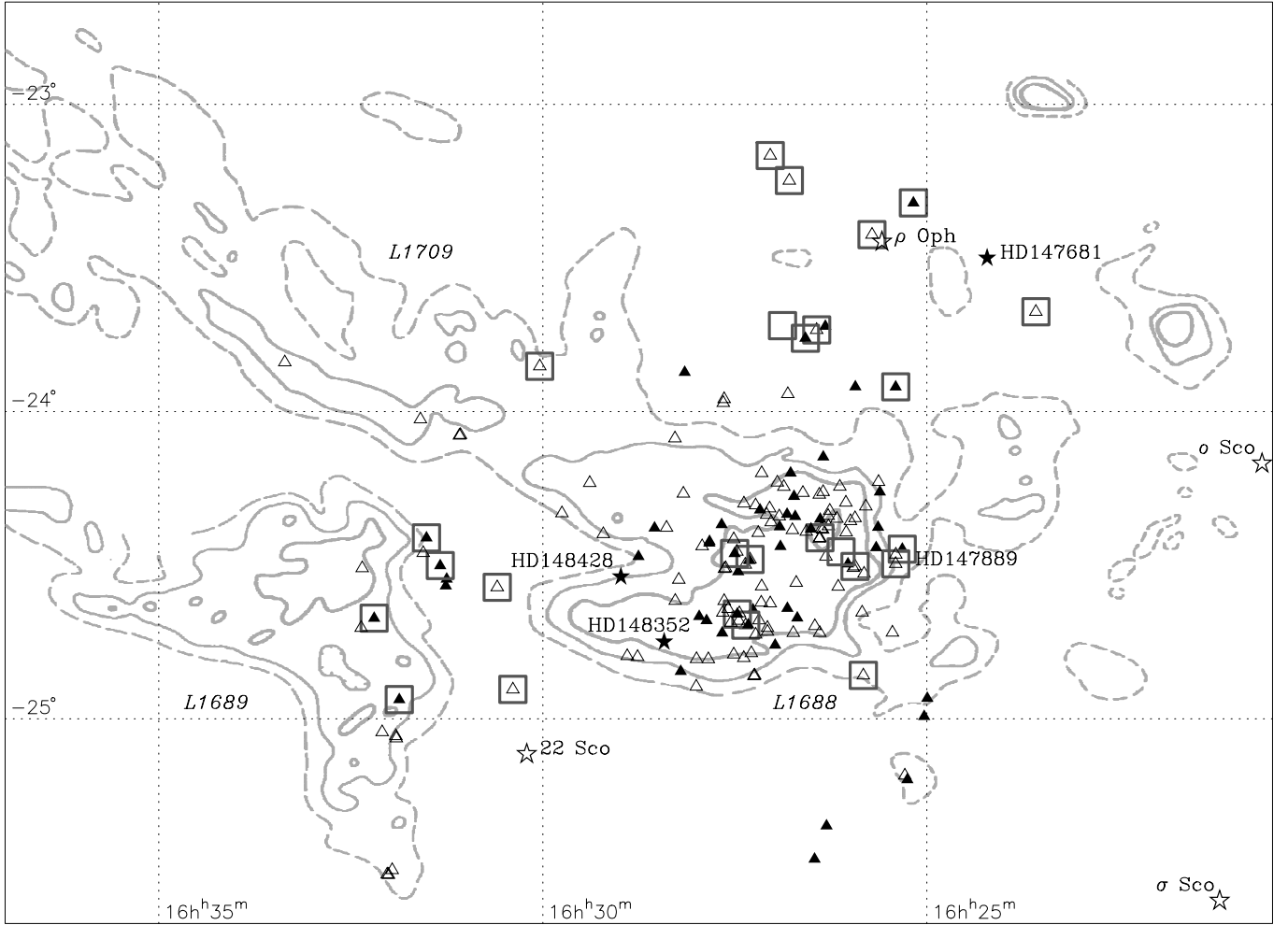


Fig. 1. The ^{13}CO ($J = 0-1$) contours of the ρ Oph molecular clouds for $T_{\text{A}}^*(^{13}\text{CO}) = 2, 6, 10$ and 20 K (Loren 1989). Each triangle marks a star of our sample. Filled triangles indicate double or multiple systems. The squares frame those areas that are used for the determination of the stellar background density. With exception of two cases the squares are centered around or near stars included in our sample. While the empty square close to the center of the core contains ISO-Oph 13 and ISO-Oph 14, the second one southeast of ρ Oph is centered around VSS 28. Bright prominent stars not included in our sample are marked by a white asterisk, while black ones represent stars used as a PSF reference. All coordinates are in equinox J2000.0.

density effects (Sect. 7). A summary of the results is given in Sect. 8.

2. The cloud complex

The ρ Ophiuchi Dark Cloud (L1688, see Fig. 1) is the densest part of a complex of vast dark nebulae and molecular clouds that extends from $l \approx 345^\circ$ to 10° and from $b \approx 0^\circ$ to $+25^\circ$. The eastern part of this complex is dominated by long elongated filaments. A scenario presented by de Geus (1992) assumes that early-type stars located in the Upper-Scorpius OB association ($l \approx 360^\circ \dots 343^\circ$, $b \approx +10^\circ \dots +30^\circ$) produced a shock-wave that encountered the dense precursor of the ρ Oph cloud from behind, swept away material and deposited it in the present day filaments. This encounter may also have triggered the continuing low-mass star formation within this cloud, resulting in an extremely young population of stars with a median age of ≈ 0.3 Myr (Greene & Meyer 1995; Luhman & Rieke 1999).

A recent paper by Sartori et al. (2003) investigates the star-formation process on a larger scale. They found that the pre-main-sequence stars within the Ophiuchus, Lupus and Chamaeleon molecular cloud complexes follow a similar spatial distribution as the early-type stars in the subgroups of the Scorpius-Centaurus OB association and a newly found OB association in Chamaeleon. Furthermore, the young objects form an almost uniform group with respect to their kinematics and ages. The most natural scenario to explain the measurements is a spiral arm passing close to the Sun. The global distribution of HII regions (Lépine et al. 2001) supports this hypothesis.

The distance to the ρ Oph cloud is still the subject of discussions. A lower limit was published by Knude & Hog (1998) who found a steep increase of reddening at 120 pc. An upper limit of 145 ± 2 pc has been derived by de Zeeuw et al. (1999). They determined the distance of the Scorpius-Centaurus OB association in the background of the ρ Oph cloud by analysing positions, proper motions, and parallaxes of its members. In

Table 1. Contributions from different papers for $m_K = m_{2\text{MASS}} < m_{\text{lim}} = 10.5$ mag.

#	Paper	Sources	$m_K < m_{\text{lim}}$			Survey	Region	Criterion	Association
			Total	New	Obs				
1	Casanova et al. (1995) (Table 2)	87	61	61	59	61	core	X-ray + NIR, visual	bona fide
2	Casanova et al. (1995) (Table 1)	19	8	8	7	7	core	X-ray + NIR	probable
3	Casanova et al. (1995) (Table 3)	22	2	2	2	2	core	X-ray + NIR	candidate
4	Wilking et al. (1989) (Table 4)	74	56	5	5	5	L1688	visual to FIR SED	bona fide
5	Greene et al. (1994)	47	37	14	14	14	L1688, L1689, L1709	visual to MIR SED	bona fide
6	Bouvier & Appenzeller (1992)	30	30	13	13	13	whole complex	visual spectra	bona fide
7	Grosso et al. (2000)	54	46	9	7	7	core	X-ray + NIR, MIR	bona fide
8	Bontemps et al. (2001)	212	98	20	15	15	L1688, L1689	MIR excess	bona fide
9	Herbig & Bell (1988)	24	24	6	6	6	whole complex	visual spectra	bona fide
10	Wilking et al. (1987)	57	53	19	15	15	whole complex	H α	probable
11	Elias (1978)	26	26	3	3	3	whole complex	NIR to MIR SED	bona fide
12	Wilking et al. (1989) (Table 6)	38	24	13	10	10	L1688	visual to FIR SED	candidate
			Σ			158			

this paper we assume a value of 140 pc. The same distance as measured for the Taurus-Auriga association and used by Köhler & Leinert (1998). This allows a direct comparison of the results.

3. The sample

Our sample of 158 young stellar objects (YSOs) recruits from surveys at optical, infrared, and X-ray wavelengths (Table 1). From these surveys we selected objects which can be considered as cloud members using criteria commonly applied in distinguishing young stars from background or foreground stars. The most convincing are detailed studies of the optical spectra (Herbig & Bell 1988; Bouvier & Appenzeller 1992), infrared spectral energy distributions (Wilking et al. 1989; Greene et al. 1994; Elias 1978), mid-infrared colour-magnitude relations (Bontemps et al. 2001), and X-ray detections combined with optical/infrared information (Casanova et al. 1995; Grosso et al. 2000). We tried to combine and observe a sample as reliable and complete as possible down to magnitude $m_K \leq 10.5$ mag. Therefore, we preferably included objects fulfilling more than one of the criteria infrared excess, X-ray detection, and H α emission. The criteria met by the individual sources of our sample are indicated in Table 2 together with the number of the catalogue that lead to their selection. If the criteria did not appear strong, we marked the source with an “U”. Our sample may be characterised as volume-limited with magnitude cutoff. It was intended to be larger in size and more statistically complete than earlier surveys.

The coordinates and the magnitudes in the K -band presented in Table 2 are taken from the Two Micron All-Sky-Survey (2MASS) Catalog of Point Sources. At the time of the preparation of our survey, we had the slightly different K -band magnitudes of Greene & Young (1992) and Barsony et al. (1997) available and used them to determine the magnitude cutoff. This means that some sources close to the cutoff and bright enough in the 2MASS survey were not observed, and vice versa.

We started to build the sample with the then new list of 87¹ confirmed cloud members presented in Casanova et al. (1995). These authors analysed a deep ROSAT image of the central region of the ρ Oph star-forming region and compared the sources with a list of confirmed members mainly derived from the infrared surveys of Wilking et al. (1989) and Greene et al. (1994). This list was completed by including ROXs 4, SR 2 and VLA 1623. It includes X-ray sources with an IR-counterpart but not detected in the visible. This strengthens the role of X-ray observations as a criterion of cloud membership. Since in Casanova et al. (1995) 67% of the found X-ray sources and 42% of the candidate X-ray sources are common with the list of confirmed cloud members, hitherto unconfirmed cloud members coinciding with the remaining (candidate) X-ray sources are probable new cloud members and thus also targets of our survey. After removal of the background giant VSSG 6 (Luhman & Rieke 1999) we are left with 61 certain and 10 probable cloud members (see Table 1). For the catalogue differences just mentioned IRS 46 and IRS 54, bright enough in 2MASS, did not make it into our sample, while WL 5 and WL 6 were observed.

From the table of cluster members in Wilking et al. (1989), H α 38, H α 60, SR 20 and H α 63 are missing in the list of Casanova et al. (1995) due to their position outside the investigated core region, as well as the objects IRS 7, IRS 8, IRS 14 where the IRAS association was uncertain. Also the source VSSG 12 was ignored for inconsistencies in the coordinates². With exception of the spurious VSSG 12 and the faint objects IRS 7 and IRS 14 we reinserted these sources.

The multicolour infrared study by Greene et al. (1994) includes also sources in L1689 and L1709. We removed VSSG 13, VSSG 15, and VSSG 16, because they had been identified by Elias (1978) as field sources, and also the background giants GY 45, GY 65, GY 232, GY 411, and VSSG 6

¹ To allow easier comparison of different papers the wide binaries SR 12, SR 24, ROXs 31, ROXs 43 are always counted as one object.

² Although SIMBAD identifies VSSG 12 with GSS 30-IRS 2 Wilking et al. (1989) give other coordinates north of GSS 29.

(Luhman & Rieke 1999). From the remaining 37 sources brighter than $m_K = 10.5$ mag 14 were new and are part of our sample.

Bouvier & Appenzeller (1992) searched for counterpart candidates of X-ray sources detected with the Einstein satellite. Studying 46 optically visible stars lying in the error circles of 29 ROX sources with spectroscopic and photometric methods resulted in the identification of 29 certain and one probable (ROXs 45D = DoAr 48) cloud member. Of these, we added the 13 until now unaccounted cloud members to our list.

Out of 63 sources found with the ROSAT High Resolution Imager Grosso et al. (2000) could identify 54 with optical, infrared and radio sources. This emphasises again the usefulness of X-ray emission as criterion for membership. We observed seven of the nine new targets, omitting two which were close to the brightness limit.

Recently (Bontemps et al. 2001) presented an extensive mid-infrared survey of L1688, L1689N and L1689S performed with the ISOCAM camera on board the ISO-satellite at $6.7 \mu\text{m}$ and $14.3 \mu\text{m}$. A catalogue of 212 sources detected at both wavelengths and classified as cloud members on the basis of colour-magnitude relations is now available. This catalogue includes 98 objects brighter than $m_K = 10.5$ mag, of which twenty were not already included in our list. With exception of five sources close to the brightness limit all could be observed. If a source of our survey is included in this catalogue, the number therein is given in the second column of Table 2.

In the third edition of their catalogue Herbig & Bell (1988) listed 24 sources towards the ρ Oph molecular clouds. In our complete field six sources were new and thus added to our list.

Wilking et al. (1987) used objective-prism plates to survey 40 square degrees toward the Sco-Cen OB association including much of the ρ Oph cloud complex for H α emission. Of the 57 objects not far from the central cloud L1688, nineteen were not yet included in our sample. They are mainly located in the western part of the complex. All sources of this catalogue that have been observed are indicated by their number in the third column of Table 2. From the sources in Elias (1978) 26 fall into our region, of which three add to our catalogue. From the list of unidentified sources given in Wilking et al. (1989) thirteen sources are new and we could observe 10 of them.

The full list of sources brighter than 10.5 mag in the K -band would include 173 objects spread over the molecular clouds with a natural concentration in L1688. Our multiplicity survey covers 156 of these and 2 slightly fainter young stellar objects. Eleven well known sources among them have been already observed in the last decade by Ghez et al. (1993) and Simon et al. (1995) with speckle imaging and during lunar occultations, i.e. with sufficient resolution and sensitivity. So, there was no necessity to observe these sources again. They are marked in Table 2 with an “O”.

4. Observations

The principle part of the speckle observations (Table 3) were carried out with the camera SHARP I (System for High Angular Resolution Pictures) of the Max-Planck-Institute for Extraterrestrial Physics (Hofmann et al. 1992) mounted on the

Table 3. Journal of observations.

Camera	Telescope	Date
SHARP I	NTT, La Silla	2000, June 17–22 2001, June 28–July 4
BlackMAGIC	3.5 m, Calar Alto	2000, June 22
SHARP II+ / ADONIS	3.6 m, La Silla	2000, June 5–6
Ω -Cass (background)	3.5 m, Calar Alto	2001, May 31–June 1

ESO New Technology Telescope (NTT) at La Silla, Chile. Further we obtained observations with BlackMAGIC (Herbst et al. 1993) on the 3.5 m telescope at Calar Alto, Spain and with SHARP II+ with the adaptive optics system ADONIS on the ESO 3.6 m telescope at La Silla. In Table 2 objects observed with BlackMAGIC are marked with a “B” while those observed with ADONIS/SHARP II+ are indicated by an “A”. All observations have been performed in the K -band at $2.2 \mu\text{m}$.

The cameras are equipped with an 256×256 pixel NICMOS3 array. To derive the exact pixel scale and orientation of the chips we took images of the Galactic Center and/or the Orion Trapezium during each observing campaign. We compared the instrumental positions of the stars with the very accurate coordinates given in Genzel et al. (1996), Menten et al. (1997) and McCaughrean & Stauffer (1994) by using the astrometric software *ASTROM*. In the case of the observations with BlackMAGIC no such calibration frames are available. Here we compared the position angle and separation of H α 71 given in Koresko (2002) with our result.

For each of the scientific targets we took between 500 and 1000 frames with an exposure time of 200 to 500 ms each to create the two fitscubes required for the data analysis (see Sect. 5). We centered the primaries in one of the four quadrants of the detector and shifted the target after half of the frames had been taken to another quadrant. If no companion was visible below the primary we used the lower two quadrants. The advantage of this shifting is the exact measurement of the background, both at the same time in different areas of the chip, and in the same area at a different time.

To analyse the data we also need speckle-images of stars that have no companions as PSF-references. In most cases we obtained time series of the star HD 148352 with a spectral type of F2V (Houk & Smith-Moore 1988) and a K -band magnitude of 6.511 ± 0.018 mag (2MASS). Another reference is the high proper-motion star HD 148428 with a spectral type of K0III+G (Houk & Smith-Moore 1988) and a K -band magnitude of 5.925 ± 0.024 mag (2MASS). These two stars are located in the foreground of the southern wing of L1688 and they are comparable to our brightest targets in the K -band. An additional reference used during the observations with BlackMAGIC is the G1 main-sequence star HD 147681. Its brightness is 7.508 ± 0.023 mag in the K -band (2MASS). The references for the objects observed with SHARP II+ are the single targets themselves.

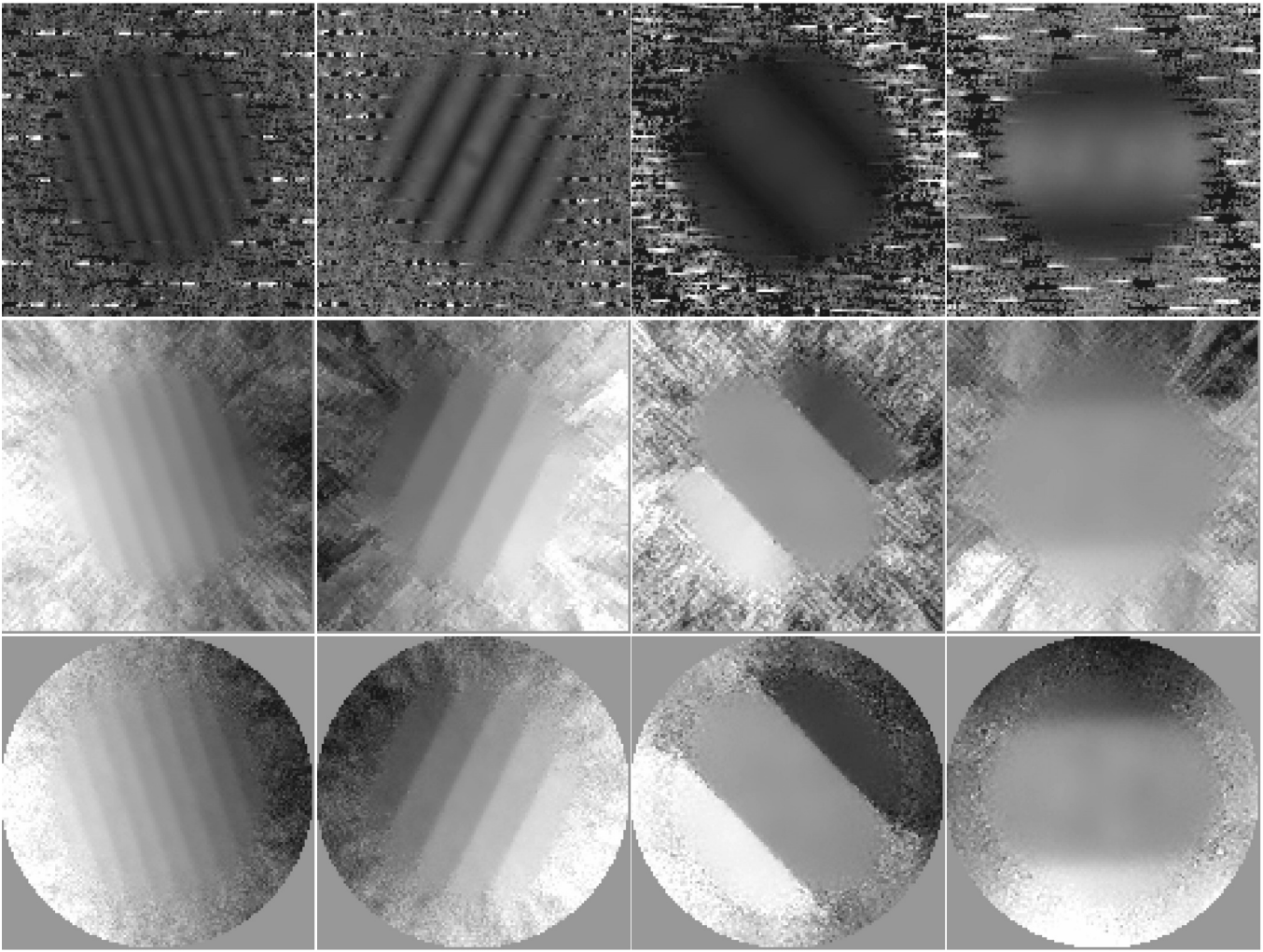


Fig. 2. The visibility, the Knox-Thompson phase, and the bispectrum phase (*from top to bottom*) of the sources IRS 3, ROXs 31, VSSG 5, and VSSG 11 (*from left to right*) as reconstructed by our software. The increasing gap between the maxima in the visibility and the decreasing number of steps in the phases clearly indicate a decrease in separation: 0.663'', 0.396'', 0.148'', and 0.107''. VSSG 11 is an example of an object falling below the diffraction limit of the telescope, i.e. we are not able to decide whether it is a binary or an elongated structure. The spatial vector between the two components of a binary is perpendicular to the stripes in the visibility and the phases. The overall gradient of the phases eliminates the 180° ambiguity. We get position angles of 115.5°, 251.3°, 133.9°, and 180.1°. The flux ratio (0.323, 0.655, 0.873, and 0.584) can be determined by the amplitude of the sinusoidal wave in the visibility and the transition between the steps in the phases. The smaller the flux ratio the smaller the height of the steps. The equally spaced horizontal stripes in the visibilities are artefacts, probably from an interference with the readout electronics.

5. Data analysis

We used our program *speckle* which has already been used for the surveys in Taurus-Auriga (Köhler & Leinert 1998), Scorpius-Centaurus (Köhler et al. 2000), and Chamaeleon (Köhler 2001). In this program, the modulus of the complex visibility (i.e., the Fourier transform of the object brightness distribution) is determined from power spectrum analysis, and the phase is computed using the Knox-Thompson algorithm (Knox & Thompson 1974) and from the bispectrum (Lohmann et al. 1983). For a more detailed description see Köhler et al. (2000). A few examples are presented in Fig. 2.

If the object turns out to be a binary or multiple star, we obtain the position angle, separation and brightness ratio of the components from a multidimensional least-square fit. Our program tries to minimize the difference between modulus and

phase computed from a model binary and the observational data by varying the binary parameters. Fits to different subsets of the data give an estimate for the standard deviation of the binary parameters.

For the search of wide and faint companions we coadded the frames within one fitscube. This increases the sensitivity, because signals are amplified while the noise is reduced. On the other hand this merging process would introduce the atmospheric distortion again that has been “frozen” by the short integration time of speckle images. The technique of shift-and-add partially circumvents this problem. Here the frames are centered before they are summed up. This centering can be done either on the brightest pixel or the center of light. Afterwards, these pictures can be analysed with standard tools, like *daophot*.

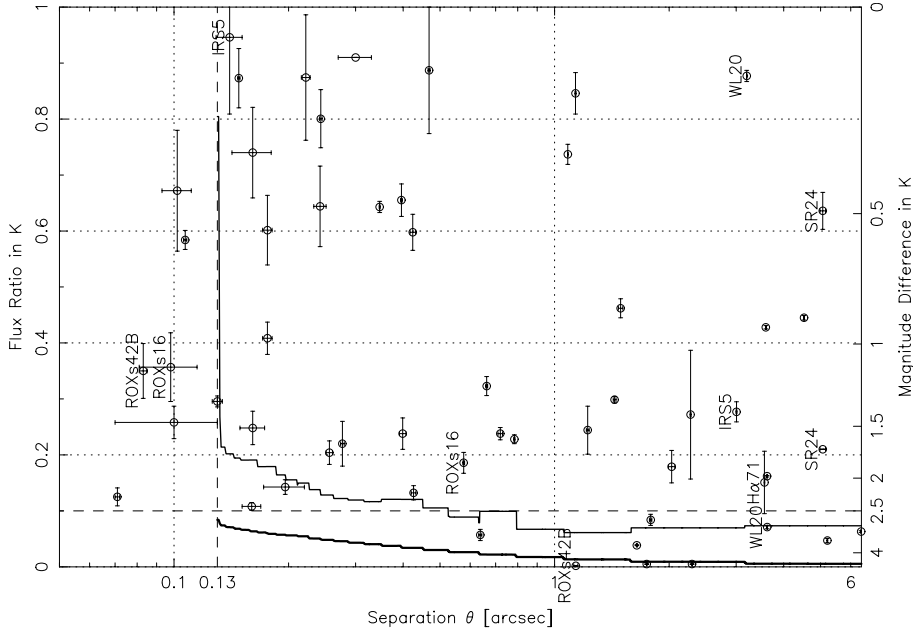


Fig. 3. Results of our multiplicity survey in a plot of flux ratio or magnitude difference vs. binary star separation. The data points mark the detected companion stars. If a companion is a component of a triple star it is labeled with the name of the system. The thick line is the average, and the thin line the worst sensitivity for undetected companions. The dashed vertical line at $0.13''$ shows the diffraction limit for a 3.5 m telescope at K . This is the limit for unambiguous identification of binary stars. The dashed horizontal line shows the completeness limit in flux ratio for the whole survey.

If the object appears unresolved, we compute the maximum brightness ratio of a companion that could be hidden in the noise of the data. The principle is to determine how far the data deviate from the nominal result for a point source (modulus = 1, phase = 0) and to interpret this deviation as caused by a companion. The procedure is repeated for different position angles and the maximum is used as an upper limit for the brightness ratio of an undetected companion (Leinert et al. 1997). In Table 4 we list the values at a distance of 0.15 arcsec and 0.50 arcsec from the primary. After subtraction of the companion(s) the first value is also calculated for double or multiple stars as an indicator for the quality of the fit (Table 5).

6. Results

6.1. Uncorrected data

In Tables 4 and 5 we list our results. Objects also observed in other near-infrared high-resolution studies are identified. In total, among the 158 targets of our sample, we find up to separations of $6.4''$ 45 binaries, 5 triple systems (ROXs 16, WL 20, ROXs 42B, L1689-IRS 5, and SR 24), and no quadruples. The flux ratio or magnitude difference vs. the separation of these systems is plotted in Fig. 3.

6.2. Completeness

The sensitivity of our survey as a function of the separation (see Fig. 3) depends on factors such as atmospheric conditions at the time of the observations or the brightness of the target star. Since we derive for each dataset with our reduction method the maximum brightness ratio of a possible undetected companion (see Table 5), we can continuously monitor the quality of our data. At the diffraction limit we reached in 85% of the observations our quality criterion of a flux ratio ≤ 0.1 (≥ 2.5 mag) in the K -band. Twenty-two observations are not quite sensitive enough to fit this request. The maximum brightness of an undetected companion at the diffraction limit varies here between

0.11 and 0.19. In the case of IRS 44 where the data are very noisy, we provide in Table 5 the flux ratio of the detected companion (~ 0.2) at a separation of $0.26''$ as upper limit for the brightness of an undetected companion.

Based on the surface density of companions found in Fig. 3 at separations larger than $0.13''$ in the range between the requested flux ratio of 0.1 and the detection limits of the twenty-two measurements described above, the probability to have missed one companion is 40%. Since the real sensitivity deficit is only relevant for separations below $1''$, this estimate represents an upper limit. We are thus confident, that we have found all companions with a magnitude difference ≤ 2.5 mag.

6.3. Lower separation limit

The lower limit of $0.13''$ is given by the diffraction limit λ/D of a 3.5 m telescope in K . Nevertheless, it is possible to detect under good circumstances companions down to a separation of $\frac{1}{2}\lambda/D$, where the first minimum of the modulus of the complex visibility can be seen. In these cases it is not longer possible to definitely distinguish between an elongated structure and a binary star. Figure 3 shows that we actually found such candidates: ROXR1-12, VSSG 11, ROXs 16, H α 59 and ROXs 42B. Also the close companion of SR 20 would fall below our diffraction limit. It was detected below the diffraction limit of the Hale 5 m Telescope of Palomar Observatory by Ghez et al. (1993).

6.4. Background

Since it is not possible with our data to distinguish between gravitationally bound companions and mere chance projections of background stars, it is necessary to quantify this bias. Therefore, we analysed 104 fields covering ≈ 6.5 arcmin $^{-2}$ each, that are centered around 24 infrared sources of our sample, ISO-Oph 13 and VSS 28 (see Fig. 1). These fields were

Table 4. Upper limits for the relative brightness of an undetected companion to the unresolved stars in our survey, measured at 0.15'' and 0.50''.

Object	Date	0.15''	0.50''	References*	Object	Date	0.15''	0.50''	References*
H α 16	2001, July 4	0.09	0.03	A1	IRS 32	2000, June 22	0.08	0.04	
H α 22	2001, July 4	0.08	0.03		VSSG 24	2001, June 29	0.07	0.03	
SR 22	2000, June 17	0.07	0.07	B2	IRS 32b	'91, Aug./'92, July	0.5 (0.02'')		S2
SR 1	2000, June 6	0.09	0.03		ROXs 20A	2000, June 20	0.10	0.10	
SR 8	2001, July 4	0.14	0.07		ROXs 20B	2000, June 20	0.09	0.07	
Elias 12	2000, June 17	0.03	0.02		H α 47	2000, June 21	0.09	0.04	
H α 24	2001, July 4	0.12	0.05		WL 5	2000, June 20	0.18	0.13	
IRS 8	2000, June 21	0.05	0.02		IRS 42	2000, June 21	0.06	0.03	C, S2
IRS 9	2000, June 20	0.09	0.06	C	WL 6	2001, July 1	0.11	0.05	
ROXs 3	2000, June 17	0.06	0.04	S2	VSSG 22	2000, June 20	0.05	0.04	C
VSS 23	2000, June 17	0.05	0.05	B2	H α 49	2000, June 20	0.06	0.06	
IRS 11	2001, July 3	0.09	0.04		GY 262	2000, June 21	0.10	0.04	
SR 4	1990, Aug. 7	0.05	0.04	C, G2, S2	IRS 43	2001, June 29	0.07	0.07	C, S2
GSS 20	2000, June 17	0.05	0.04	A1, C	VSSG 18	2000, June 20	0.12	0.06	
Chini 8	2001, June 30	0.16	0.04		GY 284	2001, July 1	0.08	0.04	
DoAr 21	1990, July 9	0.06	0.06	A1, C, G2, S2	J162730-244726	2001, June 29	0.07	0.05	
VSSG 19	2000, June 21	0.09	0.04		GY 292	2000, June 20	0.03	0.03	B2
Chini 11	2001, July 3	0.13	0.06		H α 50	2000, June 21	0.05	0.02	
SR 3	2000, June 6	0.05	0.04	C, S2	IRS 48	2000, June 20	0.07	0.02	C, S2
GSS26	2000, June 20	0.06	0.03	C	IRS 50	2000, June 20	0.10	0.04	C
SKS 1-7	2001, June 30	0.05	0.02		IRS 49	2000, June 21	0.04	0.02	C, S2
GSS29	2000, June 17	0.04	0.04	C, S2	ROXs 30B	2000, June 21	0.07	0.03	A1, B2
DoAr 24	1990, Aug. 7	0.09	0.07	B2, C, G2	ROXs 30C	2000, June 21	0.08	0.02	A1
VSSG1	2000, June 20	0.04	0.03	C	H α 52	2000, June 21	0.04	0.02	B2, S2
J162621-241544	2001, June 29	0.08	0.03		IRS 56	2001, July 3	0.11	0.06	S2
Elias 21	2000, June 20	0.04	0.02	C	SR 10	2000, June 21	0.04	0.02	C, R3, S2
GSS 30 - IRS 2	2000, June 20	0.10	0.06		H α 58	2001, June 29	0.12	0.03	
LFAM 3	2001, June 30	0.10	0.04	C	J162800-245340	2001, June 30	0.05	0.03	
DoAr 25	2000, June 6	0.05	0.03	C	VSS 35	2001, July 4	0.10	0.05	R3**
GSS 32	2000, June 18	0.03	0.03	R3, S2	J162813-243249	2001, July 1	0.09	0.04	
Elias 24	2000, June 17	0.10	0.03	C	H α 60	2000, June 21	0.03	0.02	S2
H α 33	2001, July 4	0.06	0.03		ISO-Oph 195	2001, June 29	0.05	0.04	
GY 33	2001, June 30	0.06	0.02		SR 20 W (GWAYL)	2001, July 3	0.08	0.04	
S1	2000, June 21	0.05	0.02	A1, C, S2**, R3**	VSS 38	2000, June 17	0.03	0.02	
J162636-241554	2001, July 3	0.12	0.05		H α 63	2000, June 17	0.07	0.05	S2
WL 8	2001, June 29	0.09	0.04		VSS 42	2001, July 4	0.07	0.02	R3
GY 112	2001, June 30	0.19	0.05		IRAS 64a	2000, June 21	0.06	0.02	
GSS39	2000, June 20	0.06	0.03	C	VSS 41	2001, July 4	0.04	0.03	
Haro 1-8	2000, June 22	0.04	0.02		Elias 41	2001, July 3	0.19	0.06	
H α 40	2001, July 3	0.06	0.03		H α 67	2001, July 4	0.14	0.04	S2
VSSG 10	2001, July 3	0.10	0.05		ROXs 39	2000, June 22	0.09	0.06	A1
VSSG 7	2001, June 30	0.14	0.03		Haro 1-14/c	2000, June 22	0.06	0.06	B2
J162656-241353	2001, June 30	0.04	0.02		Haro 1-14	2000, June 22	0.05	0.03	B2,G2
VSSG 8	2001, June 29	0.06	0.03			2001, July 4	0.07	0.04	
H α 44	2001, July 4	0.10	0.06		Haro 1-16	1990, Aug. 6	0.05	0.05	B2, G2, R3, S2
WL16	2000, June 18	0.06	0.03	C, S2	IRS 63	2001, July 4	0.08	0.03	
VSSG 9	2001, July 3	0.06	0.03		H α 73	2001, July 2	0.09	0.07	S2
GY 193	2001, June 30	0.09	0.03		H α 74	2001, July 2	0.07	0.04	B2
GY 194	2001, June 30	0.10	0.04		ROXs 45D	2001, July 2	0.07	0.02	
VSSG 21	2001, July 3	0.12	0.05		ROXs 45E	2001, July 2	0.10	0.04	
J162708-241204	2001, June 30	0.05	0.02		ROXs 45F	2001, July 2	0.08	0.04	
WL 10	2000, June 21	0.08	0.03		H α 75	2001, July 1	0.06	0.02	
Elias 29	2000, June 21	0.03	0.02	C, S2	L1689 - IRS 7	2000, June 22	0.05	0.04	
	2000, June 21	0.09	0.04		Haro 1-17	2001, July 2	0.05	0.02	
GY 224	2000, June 22	0.09	0.05		Elias 45	2001, July 3	0.13	0.05	

Names adopted from Barsony et al. (1997) are given without the leading "BKLT" and thus start with "J16".

* References are given in Table A.1 in the appendix.

** Additional lunar occultation companions.

Table 5. The double and multiple stars in our sample. Given are the position angles, the separations, and the flux ratios. The upper limit for the relative brightness of an additional undetected companion at the diffraction limit is provided in the seventh column.

Object	Date	PA [deg]	Separation ["]	Flux ratio	0.15''	Remarks*
H α 18	2001, July 4	82.3 $\pm$ 0.1	1.083 $\pm$ 0.002	0.737 $\pm$ 0.018	0.15	
H α 19	2001, July 4	262.9 $\pm$ 0.1	1.491 $\pm$ 0.020	0.462 $\pm$ 0.017	0.05	
Haro 1-4	1990, July 9	27 $\pm$ 1	0.72 $\pm$ 0.01	0.238 $\pm$ 0.011	0.05	G2
H α 21	2001, July 4	57.6 $\pm$ 1.6	0.161 $\pm$ 0.019	0.740 $\pm$ 0.081	0.07	
SR 2	2000, June 5	122.4 $\pm$ 0.6	0.222 $\pm$ 0.006	0.874 $\pm$ 0.112	0.06	G2
ROXs 2	2000, June 22	345.5 $\pm$ 1.4	0.424 $\pm$ 0.007	0.598 $\pm$ 0.032	0.05	B2, C
IRS 2	2000, June 17	78.6 $\pm$ 0.4	0.426 $\pm$ 0.006	0.132 $\pm$ 0.013	0.10	B2, C
J162538-242238	2001, July 4	170.2 $\pm$ 0.5	1.788 $\pm$ 0.013	0.084 $\pm$ 0.010	0.06	
IRS 3	2001, June 29	115.5 $\pm$ 0.6	0.663 $\pm$ 0.004	0.323 $\pm$ 0.017	0.04	
ROXs 5	2000, June 22	327.3 $\pm$ 1.7	0.176 $\pm$ 0.005	0.408 $\pm$ 0.029	0.03	A1
ROXR1-12	2001, June 30	18.5 $\pm$ 2.9	0.102 $\pm$ 0.009	0.672 $\pm$ 0.108	0.12	
H α 26	2001, July 4	25.8 $\pm$ 0.5	1.135 $\pm$ 0.004	0.846 $\pm$ 0.037	0.11	
DoAr 22	2001, July 2	258.9 $\pm$ 0.2	2.297 $\pm$ 0.004	0.005 $\pm$ 0.000	0.03	
H α 28	2001, June 29	357.8 $\pm$ 0.1	5.209 $\pm$ 0.013	0.047 $\pm$ 0.004	0.09	
DoAr 24E	1990, July 9	150 $\pm$ 1	2.03 $\pm$ 0.04	0.179 $\pm$ 0.029	0.05	A1, C, G2, S2
ROXs 12	2001, July 2	10.3 $\pm$ 0.1	1.747 $\pm$ 0.002	0.005 $\pm$ 0.000	0.06	
VSSG 27	2000, June 20	66.8 $\pm$ 0.5	1.222 $\pm$ 0.010	0.244 $\pm$ 0.043	0.05	C
H α 35	2001, July 4	132.2 $\pm$ 0.1	2.277 $\pm$ 0.007	0.272 $\pm$ 0.115	0.10	
H α 37	2000, June 20	65 $\pm$ 2	0.16 $\pm$ 0.01	0.108 $\pm$ 0.007	0.10	not seen by C, PA mod 180°
GSS 37	2000, June 18	69.5 $\pm$ 0.3	1.438 $\pm$ 0.012	0.299 $\pm$ 0.006	0.05	C
VSSG 11	2001, July 1	180.1 $\pm$ 0.6	0.107 $\pm$ 0.001	0.584 $\pm$ 0.017	0.04	
ROXs 16	Aa-Ab 2000, June 21	24.2 $\pm$ 7.5	0.098 $\pm$ 0.017	0.357 $\pm$ 0.061	0.05	
	Aa-B	105.4 $\pm$ 0.6	0.577 $\pm$ 0.003	0.186 $\pm$ 0.019		A1, C
WL18	2000, June 22	292.4 $\pm$ 0.2	3.617 $\pm$ 0.001	0.162 $\pm$ 0.001	0.04	
VSSG 3	2000, June 21	53.8 $\pm$ 0.5	0.243 $\pm$ 0.002	0.801 $\pm$ 0.052	0.07	C
VSSG 5	2001, June 30	133.9 $\pm$ 1.3	0.148 $\pm$ 0.001	0.873 $\pm$ 0.053	0.04	
GY 156	2000, June 21	201.9 $\pm$ 1.8	0.161 $\pm$ 0.012	0.248 $\pm$ 0.030	0.07	
SR 24	S-N 1999, Apr. 17	349.4 $\pm$ 1.3	5.065 $\pm$ 0.086	0.636 $\pm$ 0.033	0.06	G2, S2 (flux limit at 0.02'')
	Na-Nb 1991, Aug. 19	84	0.197 $\pm$ 0.020	0.21		C, S2
Elias 30	2000, June 21	175.6 $\pm$ 0.2	6.388 $\pm$ 0.013	0.063 $\pm$ 0.002	0.06	S2, R3, not seen by C
WL 20	A-B 2001, July 1	269.9 $\pm$ 0.1	3.198 $\pm$ 0.000	0.877 $\pm$ 0.010	0.06	
	A-C	232.3 $\pm$ 0.1	3.619 $\pm$ 0.001	0.071 $\pm$ 0.003		
WL 4	2000, June 20	284.2 $\pm$ 2.3	0.176 $\pm$ 0.005	0.602 $\pm$ 0.062	0.06	not seen by C
SR 12	'86, Jan. / '91, Aug.	85	0.300 $\pm$ 0.030	0.91	0.33	C, S2 (flux limit at 0.02'')
VSSG 25	2000, June 20	173.3 $\pm$ 0.3	0.468 $\pm$ 0.003	0.887 $\pm$ 0.113	0.10	C
IRS 44	2001, June 30	246.6 $\pm$ 5.1	0.256 $\pm$ 0.005	0.204 $\pm$ 0.021	0.2	not seen by C and S2, bad s/n
VSSG 17	2000, June 21	260.2 $\pm$ 0.8	0.242 $\pm$ 0.009	0.644 $\pm$ 0.072	0.04	C
IRS 51	2000, June 20	9.6 $\pm$ 0.3	1.645 $\pm$ 0.005	0.039 $\pm$ 0.001	0.07	not seen by C and S2
SR 9	2001, July 3	353.3 $\pm$ 0.5	0.638 $\pm$ 0.006	0.057 $\pm$ 0.010	0.17	B2, G2
GY 371	2001, June 30	198.1 $\pm$ 0.3	0.347 $\pm$ 0.001	0.643 $\pm$ 0.010	0.06	
VSSG 14	2000, June 18	83.6 $\pm$ 1.5	0.130 $\pm$ 0.004	0.296 $\pm$ 0.010	0.04	S2, R3, not seen by C
ROXs 31	2001, June 29	251.3 $\pm$ 0.2	0.396 $\pm$ 0.002	0.655 $\pm$ 0.029	0.05	A1, C, S2
GY 410	2000, June 20	277.0 $\pm$ 1.4	0.196 $\pm$ 0.024	0.143 $\pm$ 0.013	0.05	
H α 59	2001, July 4	103.2 $\pm$ 2.4	0.100 $\pm$ 0.030	0.258 $\pm$ 0.029	0.07	
J162812-245043	2001, July 3	101.7 $\pm$ 0.1	3.591 $\pm$ 0.001	0.428 $\pm$ 0.003	0.15	
SR 20	1990, July 9	225 $\pm$ 5	0.071 $\pm$ 0.001	0.125 $\pm$ 0.016	0.03	G2, R3, S2, not seen by C
V 853 Oph	1990, Aug. 7	96 $\pm$ 2	0.399 $\pm$ 0.008	0.238 $\pm$ 0.028	0.14	C, G2, S2 (triple**)
ROXs 42B	Aa-Ab 2001, July 1	157.9 $\pm$ 1.7	0.083 $\pm$ 0.002	0.350 $\pm$ 0.049	0.06	R3, S2, not seen by A1
	A-B	268.0 $\pm$ 0.3	1.137 $\pm$ 0.014	0.002 $\pm$ 0.001		
ROXs 42C	2001, July 1	151.0 $\pm$ 0.7	0.277 $\pm$ 0.003	0.220 $\pm$ 0.040	0.05	B2, G2
ROXs 43A/B	2001, July 4	11.9 $\pm$ 0.1	4.523 $\pm$ 0.004	0.445 $\pm$ 0.004	0.07	A1, G2, S2 (quadruple**)
H α 71	2000, June 22	35.0 $\pm$ 1.4	3.560 $\pm$ 0.006	0.151 $\pm$ 0.056	0.04	S2
L1689 - IRS 5	A-Ba 2001, July 2	241.2 $\pm$ 0.1	3.006 $\pm$ 0.009	0.277 $\pm$ 0.018	0.05	
	Ba-Bb	84.4 $\pm$ 6.1	0.140 $\pm$ 0.011	0.946 $\pm$ 0.137		
DoAr 51	2001, July 2	79.3 $\pm$ 0.2	0.784 $\pm$ 0.003	0.228 $\pm$ 0.008	0.06	B2

Names adopted from Barsony et al. (1997) are given without the leading "BKLT" and thus start with "J16".

* References are given in Table A.1 in the appendix.

** Additional spectroscopic or lunar occultation companions.

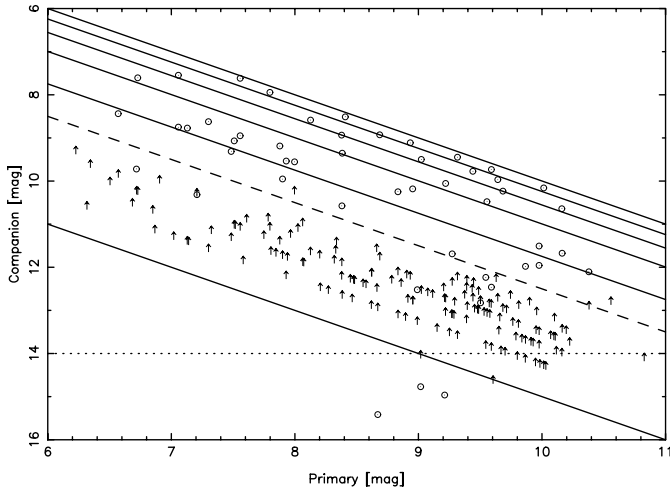


Fig. 4. The brightness of detected ($\odot$) and the upper limit for non-detected companions at a separation of $0.5''$ ($\uparrow$) vs. the brightness of the primaries. The diagonal lines indicate flux ratios of 0.01, 0.1 (=completeness, dashed), 0.2, 0.4, 0.6, 0.8, and 1.0. The dotted horizontal line gives the magnitude used for the background determination.

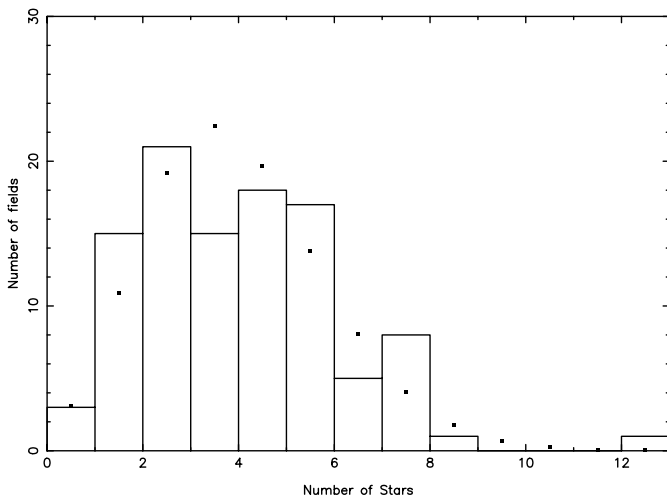


Fig. 5. Background statistics: in the 104 fields all stars down to a brightness limit of 14 mag are included. The dots represent a Poisson distribution with the mean value of ≈ 3.5 .

created by mosaicing images obtained with the infrared camera Ω -Cass in the K - or K_s -band (see Table 3). Ω -Cass was mounted on the 3.5 m telescope at Calar Alto, Spain. After excluding the central region with a radius of 6.4 arcsec corresponding to the largest separation found in our sample, we divided each mosaic into four equal fields.

Although we detected three companions with a K -band magnitude around 15 (see Fig. 4), these were found in the shift-and-add images and are thus not representative for the detection limit of our survey. The upper limits provided by the speckle software for non-detected companions are much better suited for this purpose. As shown in Fig. 4 they correspond to $m_K = 14$ mag for the fainter primaries.

The results of counting the stars down to the 14th magnitude in each field is plotted in Fig. 5. The histogram can be fitted by a Poisson distribution with a mean value of ≈ 3.5 ,

which corresponds to an absolute value of $1.5 \times 10^{-4} \text{ arcsec}^{-2}$. Defining the area within $16^{\text{h}}25^{\text{m}} \dots 16^{\text{h}}30^{\text{m}}$ in right ascension and $-25^\circ \dots -24^\circ$ in declination as center and the remaining area as periphery, we find no significant difference between them. The background density of $1.5 \times 10^{-4} \text{ arcsec}^{-2}$ is in good agreement with the value $1.6 \dots 1.7 \times 10^{-4} \text{ arcsec}^{-2}$ that we derived from the survey of Barsony et al. (1997) by counting the stars brighter than $m_K = 14$ mag.

Could chance coincidences with background galaxies have led to a false classification as young stellar object or to a spurious detection of a companion? The answer is no. The galaxy counts performed with the same instrument as used by Bontemps et al. (2001) give values of about 25000 sr^{-1} (Serjeant et al. 2000) for the mid-infrared and the relevant sensitivity limits of 5 mJy at $6.7 \mu\text{m}$ and 10 mJy at $14 \mu\text{m}$. The probability is only 1% that any one of the objects in our sample could be close enough to such a galaxy, i.e. within $9''$, to have its mid-infrared photometry affected by the presence of this galaxy. Similarly, the K -band galaxy counts (Gardner et al. 1993; Huang et al. 2001) result in about 0.1 galaxies per magnitude interval per square degree at $m_K = 10$ mag. The number increases with magnitude $\propto 10^{0.67m_K}$ down to $m_K = 16.5$ mag. This gives a probability of 1.4% that any one of the objects could have the photometry affected by a close galaxy, i.e. residing within a $4''$ diameter. A probability of only 14% is found that any of the companions within our limits of $6.4''$ radius and $m_K \leq 14$ mag would be a background galaxy.

6.5. Surface density

An interesting property of a star forming region is the surface density $\Sigma(\theta)$ of companions (see Fig. 6). Over the separation range $0.13'' \leq \theta \leq 6.4''$ a linear regression of the surface density leads to

$$\Sigma(\theta) \propto \theta^{-2.13 \pm 0.07}, \quad (1)$$

which means that the number of companions is almost constant per logarithmic separation interval (see Fig. 3). This is nearly the same value as derived for the Taurus star forming region (Köhler & Leinert 1998). Due to the enlarged samples, both results put the conclusion of Simon (1997) on a firmer footing, that the surface density of companions in the binary regime in different star forming regions (Taurus, Ophiuchus, Orion Trapezium) can be approximately described by θ^{-2} .

The surface density of the companions is used to provide an upper limit for the separations in our survey. We choose $6.4''$ (half the field of view of the SHARP cameras), because chance projections of background or foreground objects would become important at larger separations.

6.6. Wide companions

The field of view of the SHARP I and the SHARP II+ cameras is $12.8''$. Usually, we centered the object in one of the lower quadrants, where the number of bad pixels was lowest. We therefore should have seen all companions out to $3.2''$ in these measurements. For the wider companions we searched

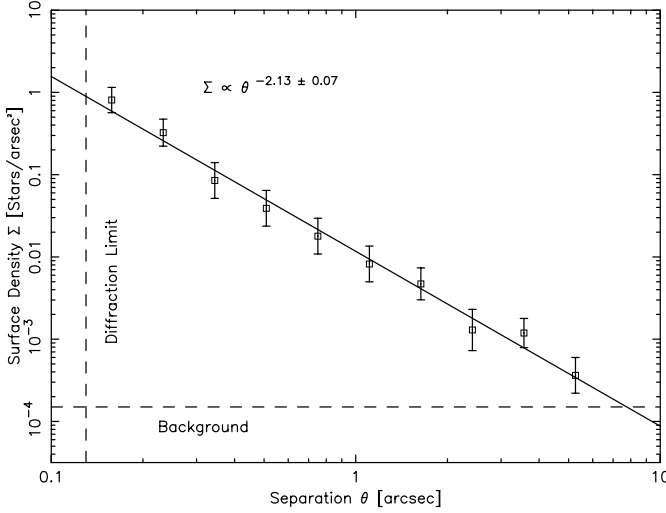


Fig. 6. Surface density of the companions, compared to the surface density of the background stars.

the 2MASS database. We found 11 infrared sources with a second source detected in the K -band within our separation range. These objects are marked with a “C” in Table 2. None of them added to our list of companions (Table 5), because they had been already detected in our speckle data or were doubtful.

The “wide companions” of the sources Elias 21 ($m_K^{\text{comp}} = 11.015$ mag), VSSG 18 (12.284 mag) and VSSG 17 (13.291 mag) are indicated in the 2MASS All Sky Catalog as point sources falling within the elliptical boundary of an extended source. This suggests that the point sources are extractions of pieces of underlying nebulae. A visual inspection of the 2MASS images strengthens this suspicion. A similar case is the spurious source 5.5'' west of GSS 32 with a brightness of $m_K^{\text{comp}} = 13.339$ mag. It is probably an artifact. Although Simon et al. (1995) found GSS 32 single, they did not reach the necessary sensitivity to falsify the wide companion ($m_K \leq 9.2$ mag). Terebey et al. (2001) and Haisch et al. (2002) classified GSS 32 as a single star. In our fitscubes of LFAM 3 GSS 32 appears in the upper, i.e. eastern quadrants of the chip. No companions are visible.

We preferred in the case of the south and the north component of SR 24 the separation and position angle following from the coordinates given in 2MASS over the discrepant relative position reported in Simon et al. (1995) from where we adopted the values for the close pair Na-Nb.

6.7. Number of systems after background subtraction

We find in our sample of 158 targets 49 fully resolved companions in the separation range $0.13'' \leq \theta \leq 6.4''$. It thus contains 112 single stars, 43 binaries, 3 triples, and no quadruples. In addition we have to take into account that the probability p to detect a background star close to a surveyed star is (Sect. 6.4)

$$p = \pi \cdot (6.4 \text{ arcsec})^2 \cdot 1.5 \times 10^{-4} \text{ arcsec}^{-2} \approx 0.019 \quad (2)$$

or ≈ 3 companions in the whole sample. Therefore, three of the companions should be chance projections. This leads to a companion star frequency of 0.29 ± 0.04 . To correct the number of

single, binary, and triple systems, we have to take into account that, e.g. “false” triple systems can be produced with a probability of p by the “true” binaries and with a probability of p^2 by “true” single star. Otherwise, e.g. the number of “true” single systems is increased when compared to the number of “observed” single systems by a factor $1/(1-p+O(p^2))$, because projected companions reduce their number. A brief calculation leads to 114.2 “real” single stars, 41.7 binaries, and 2.2 triple systems.

6.8. The restricted sample

For statistical purposes we also define a restricted sample, excluding all targets with uncertain association (“U” in Table 2) and including only companions with brightness ratios ≥ 0.1 where we are complete and with separations exceeding the diffraction limit. The brightness ratio of 0.1 for these young stars approximately corresponds to the limit in mass ratio of 0.1 used for the work on solar-like main-sequence stars (Duquennoy & Mayor 1991). This restricted sample contains 38 companions around 139 primaries. For the restricted sample we find 103 single stars, 34 binaries, and 2 triple systems. The background density is only $0.6 \times 10^{-4} \text{ arcsec}^{-2}$ for a detection limit of 12 mag on average. With $p = 0.008$ this sample thus contains 103.8 “real” single stars, 33.4 binaries, and 1.7 triple systems. The companion star frequency is 0.27 ± 0.04 .

7. Discussion

7.1. Comparison to main-sequence stars

To compare our results with the solar-type main-sequence sample surveyed by Duquennoy & Mayor (1991) we transform their lognormal period distribution

$$f[\lg(P)] = C \exp \left[-\frac{1}{2\sigma^2} (\lg(P) - \langle \lg(P) \rangle)^2 \right] \quad (3)$$

with $\langle \lg(P) \rangle = 4.8$, $\sigma_P = 2.3$, and P in days into a lognormal distribution of separations. This is not trivial since our observations are snapshots, i.e. we cannot derive periods by fitting the orbits.

For random distribution of orbital planes the relation between semi-major axis and actual observed separation is given by (Leinert et al. 1993)

$$\langle r \rangle = \frac{\pi}{4} a \left(1 + \frac{e^2}{2} \right). \quad (4)$$

The combined reduction of the average separation with respect to the semi-major axis would be by a factor of 0.98 if the eccentricities follow the distribution (Duquennoy & Mayor 1991)

$$f(e) = 2e. \quad (5)$$

This allows to convert the orbital periods to average observed separations using Kepler’s third law. With an assumed system mass of $1 M_\odot$ and r in astronomical units $\langle \lg(P) \rangle$ and σ_P transform into $\langle \lg(r) \rangle = 1.48$ and $\sigma_r = 1.53$. The observed separation r scales with the cubic root of the total system mass.

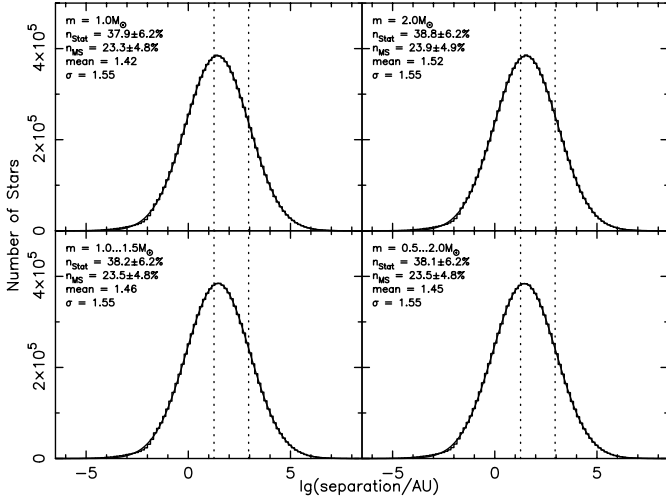


Fig. 7. Simulated distributions of projected separations for four samples of 10 million main-sequence binaries each with different system masses or mass ranges. The histogram shows the simulated data; the line is a lognormal distribution fitted to the histogram. The dotted vertical lines border the separation range we have observed when assuming a distance of 140 pc to the cloud complex.

Alternatively, we use the well-known properties of main-sequence binaries to predict their number within the observed separation range. For Fig. 7 we simulated a sample of 10^7 systems with different masses or mass ranges (values in the plots). These systems have orbital elements according to Duquennoy & Mayor (1991), i.e. the periods have the lognormal distribution (3) and the distribution of eccentricities is (5). The inclinations are distributed isotropically and the other parameters uniformly. After binning the results we fitted the distribution with a Gaussian (solid line). For a total system mass of $1 M_\odot$ we obtain $\langle \lg(r) \rangle = 1.42$ and $\sigma_r = 1.55$. While σ_r is constant for all masses and mass ranges the mean value increases as expected from lower to higher masses. For the plots in the following sections we will use $\langle \lg(r) \rangle = 1.45$ and $\sigma_r = 1.55$, i.e. the average of the values resulting from the system mass distributions $0.5 M_\odot \dots 2.0 M_\odot$ and $1.0 M_\odot \dots 1.5 M_\odot$. They are in good agreement with our results above.

We assume 140 pc for the distance to the Ophiuchus Dark Cloud. The separation range thus covered by our sample is marked by the vertical dotted lines in Fig. 7. n_{Stat} is the percentage of the systems falling within these limits. After multiplying this value with the corrected (Duquennoy & Mayor 1991) multiplicity of the main-sequence sample (101 companions out of 164 systems) we find n_{MS} , the number of companions we would have found if we had observed a sample of main-sequence stars in our survey. Due to the fact that with increasing masses the peak of the Gaussian drifts to larger separations and into our observation range, n_{MS} also increases with the mass of the systems. However, for the masses considered here this effect is negligible. From the two plots with mass ranges we find

$$n_{\text{MS}} = (23.5 \pm 4.8)\%. \quad (6)$$

We will use this value as reference.

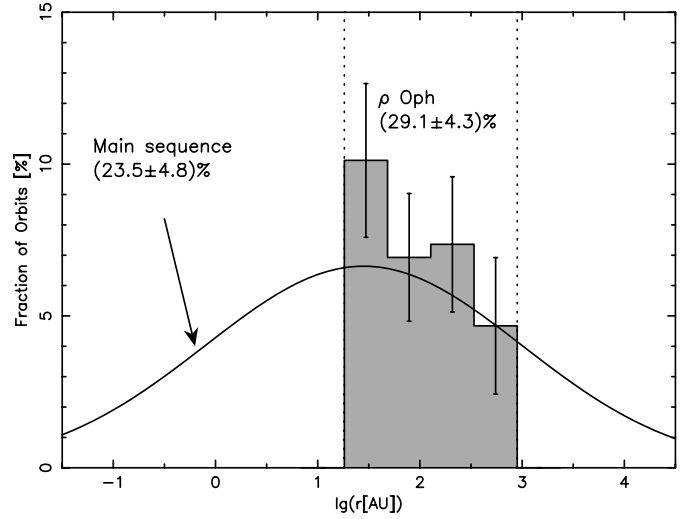


Fig. 8. Binary frequency as a function of separation for the total sample. The curve is the distribution of binaries among solar-type main-sequence stars (Duquennoy & Mayor 1991).

After binning all companions within the separation range $0.13'' \leq \theta \leq 6.4''$ into four bins and subtracting the background we plot the result of our survey in Fig. 8. Four bins are chosen since the original histogram by Duquennoy & Mayor (1991) also contains approximately four bins for the relevant range of separations or periods, respectively. The error is estimated as $\sqrt{N}$. Comparing the slope of the distribution with that of the main-sequence we find good agreement. An exception is the overabundance of close companions (see Sect. 7.3.3). With a value of

$$n_{\text{Oph}} = (29.1 \pm 4.3)\% \quad (7)$$

the multiplicity is only 1.24 ± 0.31 times larger than for the main-sequence stars (6). For the restricted sample (see Fig. 12) we have

$$n_{\text{Oph}}^{\text{res}} = (26.6 \pm 4.4)\% \quad (8)$$

or 1.13 ± 0.30 times the value for a main-sequence sample (6). We find the multiplicity in Ophiuchus marginally larger than for the main-sequence, but the difference is on the level of one σ only.

7.2. Comparison to previous surveys

The appraisal of multiplicity among young stars “in Ophiuchus” keeps changing. From the beginning, it has been centered on a comparison to the multiplicity observed in the Taurus-Auriga star-forming region.

Ghez et al. (1993) observed the 24 known young stars brighter than $m_K = 8.5$ mag in Scorpius and Ophiuchus and found no difference with respect to Taurus in the range of separations between 16 AU and 252 AU ($0.1''$ – $1.8''$), but a value of duplicity by a factor of 4 greater than that of the solar-type main-sequence stars (Duquennoy & Mayor 1991). Simon et al. (1995) had a sample of 35 sources, selected by the randomness of lunar occultation observations (location of observatories, committee approval, weather, instrumental efficiency).

Supplemented by imaging for larger separations, they found that between 3 AU and 1400 AU ($0.05''$ – $10''$) Ophiuchus had a binary frequency of 1.1 ± 0.3 times that of nearby solar-type stars, while for Taurus this number was 1.6 ± 0.3 . These are lower limits because no corrections for incompleteness were applied. Duchêne (1999) added the Gunn z band observations of Reipurth & Zinnecker (1993) to these earlier surveys, corrected for incompleteness and found an enhancement of multiplicity by a factor of 1.5 ± 0.3 (2.0 ± 0.3 , when Simon et al. (1995) is not included) over the main-sequence value, quite the same as for the Taurus-Auriga association.

Barsony et al. (2003) restricted the sample to objects searched by high-resolution near-infrared techniques. Adding new observations of this type for 19 optically selected sources from the environment of the main cloud L1688, they arrived at an overabundance of a factor of 2 ± 1 with respect to the main sequence for their sample of 80 objects, consistent with the values for the Taurus-Auriga star forming region. Duchêne et al. (2004) did a deep ($3 \text{ mag} \leq \Delta m_K^{\text{lim}} \leq 7 \text{ mag}$) near-infrared imaging survey of 63 embedded young stellar objects in Ophiuchus and Taurus, concluding that in the range of 110–1400 AU ($0.8''$ – $10''$) the multiplicity is about twice as large as for nearby solar-type main-sequence stars, and with no difference between Taurus and Ophiuchus. In this study the most embedded sources showed the highest multiplicity, still by a factor of 1.5 larger than the average. The latter result is similar to the findings in Haisch et al. (2002) on a sample of 19 embedded objects in Ophiuchus and Serpens. Our survey, with a duplicity of young stars in Ophiuchus close to that of the main-sequence sample of Duquennoy & Mayor, is similar in result to the study of Simon et al. (1995) again.

While the studies of Haisch et al. (2002) and Duchêne et al. (2004), performed on small samples, delineate interesting and important trends with age of the objects, the difference of our work to the work of Barsony et al. (2003) needs some explanation. As shown in the appendix, the difference will not lie in the different efficiency of the surveys, since binary young stars are consistently found in the overwhelming majority of cases by both surveys with quantitatively good agreement. Differences then should result from the selection of the sample and the angular limits over which duplicity is considered. Barsony et al. (2003) in their survey and compilation of 80 objects, found 0.24 ± 0.11 companions per primary for the range of $0.1''$ – $1.1''$. Choosing from their paper companions in the range of $0.13''$ to $6.4''$, as applied in our study, the resulting number of companions per primary would increase to 0.33 ± 0.07 , or 1.4 ± 0.4 above the expectation for the main-sequence sample of solar-type stars. Otherwise, when restricting our sample to separations between $0.13''$ and $1.1''$ we find a companion star frequency of 0.16 ± 0.03 . The differences are thus within the errors and naturally to be explained by differences in the samples. This just shows again the importance of large samples and to keep the sample by definition as complete as possible. The current survey with the selection criterion to take *all* stars brighter than $m_K = 10.5 \text{ mag}$ that have shown convincing signs of youth compares well with previous work.

7.3. Implications for the formation process

The general frame in which we are looking at the data is the scenario that stars originally form with a high multiplicity, which then is reduced to the main-sequence value in dense environments by dynamical interactions on a short time scale. This does not mean that we want primarily to confirm this image, but that we want to check which comments or corrections with respect to this picture result from our study.

7.3.1. Density

Both the Taurus-Auriga and the ρ Ophiuchi molecular clouds are located at a distance of about 140 pc, contain of the order of $10^4 M_\odot$ of gas and dust and harbour several hundreds of young stars with an age of at most a few million years. What causes the smaller binary frequency found in our survey when compared with the result

$$n_{\text{Tau}} = (48.9 \pm 5.3)\% = (1.93 \pm 0.26)n_{\text{MS}} \quad (9)$$

found by Köhler & Leinert (1998) for Taurus-Auriga? n_{MS} is the main-sequence binary fraction between their diffraction limit of $0.13''$ and their upper limit of $13''$. We find

$$n_{\text{Tau}}^{\text{res}} = (39.7 \pm 4.8)\% = (1.56 \pm 0.31)n_{\text{MS}} \quad (10)$$

after all companions with a flux ratio less than 0.1 have been removed.

Taurus-Auriga is the prototypical site of low-mass star-formation. Various studies of the large-scale structure have revealed a complex, irregular, and filamentary appearance. Embedded along this filamentary structures small ($\approx 0.1 \text{ pc}$) and dense ($\geq 10^4 \text{ cm}^{-3}$) cores have been identified in which the young stars are forming. Their typical mass is $1 M_\odot$ and their kinetic temperature about 10 K. Typical visual extinctions are between 5 and 10 mag. The whole Taurus-Auriga aggregate covers an area of 300 pc^2 and thus the stellar surface density is a few stars pc^{-2} . Only weak clustering is apparent.

Similar conditions are found when studying the outer regions of the ρ Oph complex. Loose filamentary and clumpy structures can be easily identified. A different environment is present in the main cloud L1688. This westernmost cloud contains in an area of only $1 \times 2 \text{ pc}$ a centrally condensed core of $600 M_\odot$ with active star formation. A large fraction of all young stellar objects in the ρ Oph molecular cloud are concentrated in this cluster. Stellar surface densities one or two orders of magnitudes higher than the values found in Taurus-Auriga are the result. Peak values of $5 \times 10^3 \text{ stars pc}^{-2}$ within the densest cores (Allen et al. 2002) are almost comparable to the values that are found in the Orion nebula cluster, although there the high densities extend over larger scales. The high star-formation efficiency in L1688 suggests that the cluster may remain gravitationally bound and thus survive as an open cluster (Lada et al. 1993). The visual extinction can reach values between 50 and 100 mag. While the radiation field caused by massive stars plays a minor or even no role for the Taurus-Auriga complex (Zinnecker et al. 1993), the ρ Ophiuchi molecular clouds are highly influenced by the nearby Upper Scorpius-Centaurus OB-association. This may be reflected by the cometary shape

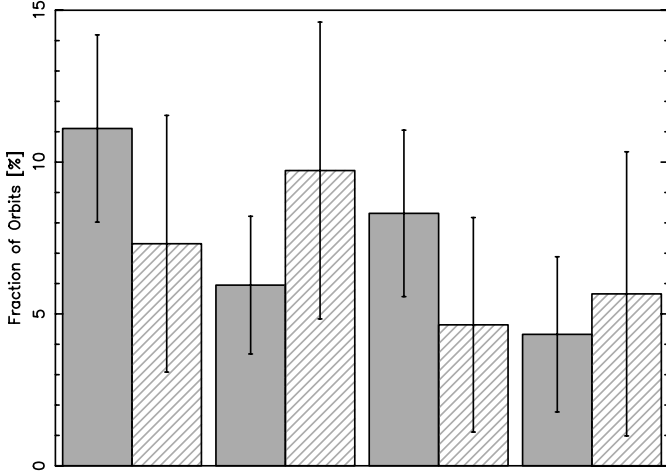


Fig. 9. The four bins plotted in Fig. 8, but plotted for both the 117 primaries in the center (solid) and the 41 primaries in the periphery (hatched).

of the complex and the high density reached within L1688. The influence of nearby massive stars may have also triggered the rapid rise of star-formation about 1 million years ago in the central cloud L1688 (Palla & Stahler 2000). All in all the ρ Oph Dark Cloud with its embedded cluster seems to be an important link between loose T associations and dense clusters.

Duchêne (1999) in his quantitative comparison of various multiplicity surveys found that all dense clusters have binary fractions compatible with the main-sequence, while all the regions with a binary excess are loose associations. This favours a tight correlation between the density or a related parameter and the multiplicity of a star forming region. Our results seem to fit very well in this picture with a duplicity value lying between those of Taurus-Auriga and the main-sequence and hence dense clusters, both for the full and the restricted samples. In the context of dependence of duplicity on density of the star-forming region this would be a plausible result, more plausible than the large overabundance of companions found in some of the earlier surveys with smaller samples.

One consequence of this density hypothesis would be a difference between the multiplicity of the dense central region (L1688) and that of the less dense outer regions. Recalculating the companion frequency of the total sample for both the 117 sources within L1688 ($16^{\text{h}}25^{\text{m}} \dots 16^{\text{h}}30^{\text{m}}$, $-25^{\circ} \dots -24^{\circ}$) and for the 41 sources in the periphery reveals that the multiplicity of both is very similar:

$$n_{\text{Cen}} = (29.7 \pm 5.0)\% = (1.26 \pm 0.34) n_{\text{MS}}, \quad (11)$$

$$n_{\text{Per}} = (27.3 \pm 8.2)\% = (1.16 \pm 0.42) n_{\text{MS}}. \quad (12)$$

Replotting the four bins displayed in Fig. 8 shows that the distribution of the separations of both subsamples differ only slightly from each other (Fig. 9). A Kolmogorov-Smirnov test (Fig. 10) also favours a common distribution. The corresponding values for the restricted sample with 104 targets in the center and 38 targets in the periphery are

$$n_{\text{Cen}}^{\text{res}} = (26.2 \pm 5.0)\% = (1.11 \pm 0.31) n_{\text{MS}}, \quad (13)$$

$$n_{\text{Per}}^{\text{res}} = (27.8 \pm 8.9)\% = (1.18 \pm 0.45) n_{\text{MS}} \quad (14)$$

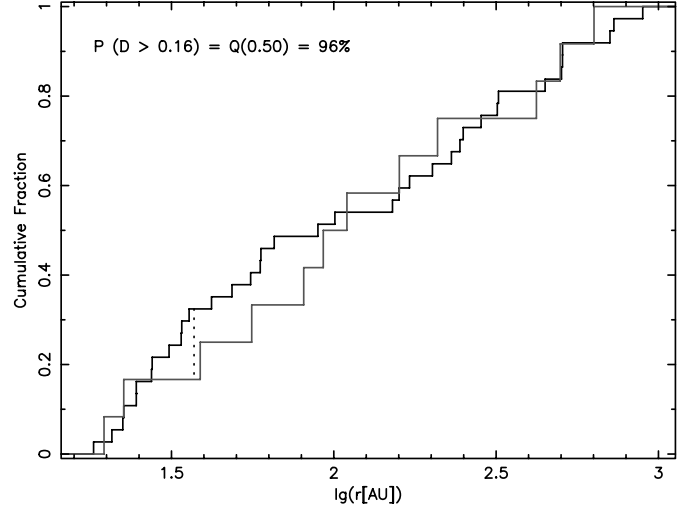


Fig. 10. A Kolmogorov-Smirnov test for the two datasets (center: black, periphery: grey). With a probability of 96% the two data sets are from the same sample.

with a probability of 29% that the two distributions are drawn from the same sample.

This means that locally we cannot see a density effect of duplicity within the errors. The sources in the surroundings appear older on average than those within L1688 and part of them could have formed in a denser environment now dissolved. However, this is nothing more than a somewhat vague possibility. Our conclusion therefore is not as clear as one might want it to be.

Although density, or a related parameter seems to play a crucial role in the formation of binaries on a global scale, there is no statistical significance within the ρ Oph molecular cloud complex that areas with different densities show different multiplicities.

7.3.2. Temporal evolution

Star forming regions with a main-sequence binary fraction are found at all ages, e.g. IC 348 (Duchêne et al. 1999), Orion (Petr et al. 1998) with an age of a few million years, the Pleiades (Bouvier et al. 1997) with 120 Myr, and the Praesepe (Bouvier et al. 2001) with 700 Myr. This suggests that dynamic interactions, if responsible for reducing an originally high duplicity to much lower values, act very quickly in dense clusters, while little future effect has to be expected for low-density regions like Taurus-Auriga. Thus temporal evolution of the binary frequency is not in general responsible for the difference between the overabundance of companions found in Taurus-Auriga when compared to the main-sequence. The fact that in the young but not too dense Ophiuchus star-forming region there remains an overabundance of companions, with the most embedded sources showing the highest degree of multiplicity (Duchêne et al. 2004), would be compatible with the dynamical evolution of binarity in cluster environments (Kroupa 1995).

We tried to see this effect in our sample with respect to the age of the different objects. We thus searched for their infrared classes in the literature. Although this classification scheme is

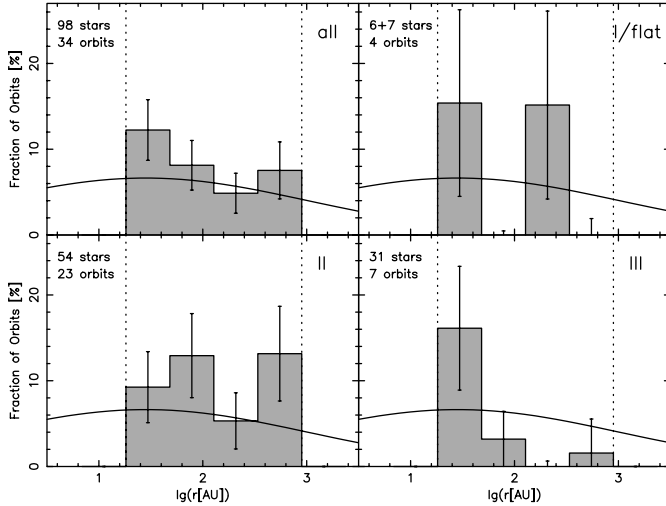


Fig. 11. Binary frequency as a function of separation and class. The upper left panel shows the distribution for all classified sources. The remaining panels display the combined sample of flat spectrum and class I sources, the class II, and the class III samples. The curve is the distribution of binaries among solar-type main-sequence stars (Duquennoy & Mayor 1991).

more a morphological description than a direct indicator of the age, it provides the best approach when no spectroscopic data are available. To avoid systematic errors from different surveys we only used the classification provided by the mid-infrared survey of Bontemps et al. (2001) and the near-infrared study of Greene et al. (1994). To be consistent with the classification in Bontemps et al. (2001) we decided to classify in Greene et al. (1994) objects with a spectral slope $a > 0.55$ as class I and those with $a > -0.05$ as flat spectrum sources. To distinguish between more evolved class II and class III objects in Greene et al. (1994) we used $a = -1.6$ as limit. WL 5 is an exception, since it is classified as an heavily reddened class III source. This conclusion is in agreement with the result in Bontemps et al. (2001). All sources included in both samples are classified consistently with exception of L1689-IRS 5 and LFAM 3 that are according to Greene et al. (1994) flat spectrum sources, but are classified as class II sources in Bontemps et al. (2001). Since LFAM 3 lies only marginally below the limit in Bontemps et al. (2001) we decided to classify it as flat spectrum source. Otherwise, L1689-IRS5 is only slightly above the limit in Greene et al. (1994) and well below in Bontemps et al. (2001). We thus classified it as class II object. Class I and class II sources in Greene et al. (1994) with an upper limit for a are ignored. This leads to a sample of 6 class I, 7 flat spectrum, 54 class II, and 31 class III sources. The multiplicity of this subsample is

$$n_{I-III} = (32.8 \pm 5.8)\% = (1.39 \pm 0.38)n_{MS} \quad (15)$$

and thus compatible with the result found in (7) and (8) within the error bars. After separating the different evolutionary states we are left with subsamples that are no longer free from small number statistics (see Fig. 11):

$$n_{I/flat} = (29 \pm 15)\% = (1.2 \pm 0.7)n_{MS}, \quad (16)$$

$$n_{II} = (41 \pm 9)\% = (1.7 \pm 0.5)n_{MS}, \quad (17)$$

$$n_{III} = (21 \pm 8)\% = (0.9 \pm 0.4)n_{MS}. \quad (18)$$

There appears to be a trend that class III systems (WTTS) have fewer companions and at smaller separations than their class II (CTTS) counterparts. This was not found in Taurus. Ghez et al. (1993) suggested from a similar result on a smaller sample that close companions may help to clear circumstellar disks earlier and therefore appear more frequently in WTTS.

Temporal evolution may be important in dense environments at early stages. In our sample of stars located in a cluster of medium density we are less sensitive to such an effect. However, the difference in the multiplicity and separation distribution between class II and class III sources and with respect to Taurus could nevertheless show real changes, maybe temporal evolution.

Another possibility is a biasing of the sample by a yet not distinguished older population of lower multiplicity. In the last section we excluded a strong influence of such a population in the periphery. Nevertheless, if the stars reside instead in the foreground, they could mimic the here discussed difference between the classes. Precise measurements, e.g. with GAIA of the parallaxes will test this idea.

Although temporal evolution seems to be not responsible for the reduction of the binary frequency in general except for the earliest stages, our survey indicates statistical differences between the infrared classes with respect to their companion frequencies and separation distributions.

7.3.3. Missing companions

Two possible explanations have been discussed by Duchêne (1999) for a low multiplicity of the Ophiuchus star forming region when compared to Taurus-Auriga. a) The distribution of the projected separations can be shifted to lower values, i.e. the “missing” companions are *too close* to be resolved and are hidden from our survey below the diffraction limit. b) The flux ratio of the companions is smaller for Ophiuchus, i.e. the “missing” companions are *too faint* to be detected.

To conclude on the first possibility, very high resolution observations (lunar occultations or interferometry) would have to be available for most of the sources of our survey, which is not yet the case. From Simon et al. (1995) and Barsony et al. (2003) there is at least evidence that no overabundance of stars with very close companions is present. However, both studies suffer from poor statistics. Figure 12 that displays the multiplicity as function of the separation for our restricted sample shows a trend that the sample is dominated by close companions. This overabundance is more apparent in the restricted sample than in the total sample (Figs. 8 and 12).

Concerning the second suggestion, Duchêne (1999) found from Ghez et al. (1993) that 73% of the binaries in Taurus, but only 23% of the binaries in Ophiuchus exhibit a magnitude difference between companion and primary of $\Delta m_K < 1.5$ mag. We want to check whether this also holds for our larger sample. In Fig. 13 the flux ratios of our companions in the total sample are displayed. Indeed, the whole sample is dominated by small flux ratios. About 59% of the systems show flux ratios below 0.4 and 33% below 0.2. This tendency to favour

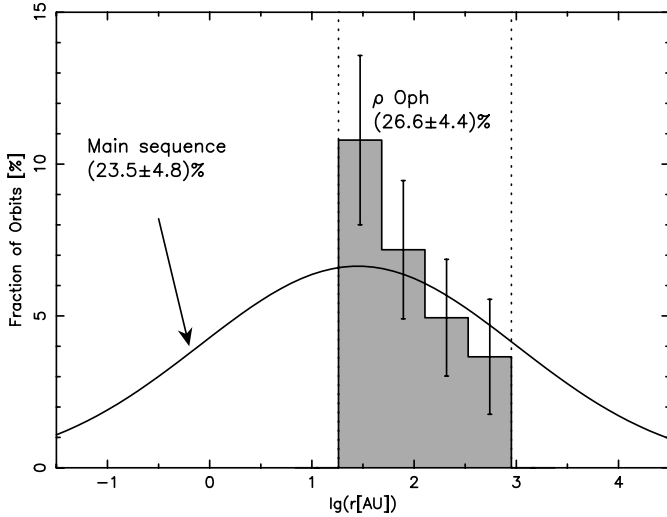


Fig. 12. Binary frequency as a function of separation for the restricted sample. The curve is the distribution of binaries among solar-type main-sequence stars (Duquennoy & Mayor 1991).

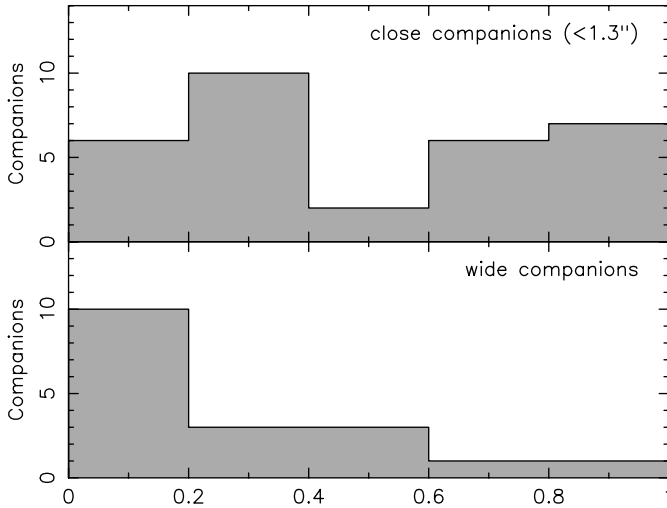


Fig. 13. Flux ratio for close (<1.3'') and wide companions.

small flux ratios is introduced by the wide pairs (>1.3''). The close companions are almost equally distributed. Even when the boundary between close and wide companions is varied, Kolmogorov-Smirnov tests show that the probability that the two distributions have a common origin is below 10%.

In Köhler & Leinert (1998) wide pairs (>1.3'') are also dominated by small flux ratios, similar to the result displayed in Fig. 13. On the other hand there is a clear tendency in Taurus-Auriga for close binaries to exhibit a large fraction of equally bright systems possibly caused by a lack of close binaries with small flux ratios that are present in Ophiuchus. Such a population may be the reason for the finding in Duchêne (1999).

A combination of the two trends, i.e. the high fraction of close binaries and the presence of close companions with low flux ratios, leads to the conclusion that “missing” companions may play a role with the implication that the full binary fraction over all separations would be more clearly enhanced than the binarity in our restricted sample.

8. Summary

- We presented a volume-limited multiplicity survey with magnitude cutoff ($m_K \leq 10.5$ mag) of 158 young stellar objects located within or in the vicinity of the ρ Ophiuchi Dark Cloud (L1688). The survey covers separations between 0.13'' (diffraction limit) and 6.4'' (background contamination) and is complete for flux ratios ≥ 0.1 ($\Delta m_K \leq 2.5$) at the diffraction limit. A restricted sample has been defined that is complete and excludes all uncertain cloud members.
- The detection limit is $m_K \approx 14$ mag, and the stellar background density at this brightness is $\approx 1.5 \times 10^{-4}$ arcsec $^{-2}$.
- Among the 147 targets newly observed with speckle techniques in the K -band we found 48 companions (40 binary and 4 triple systems). Five of these companions are below the diffraction limit of the telescopes and thus only marginally resolved. From the remaining 43 companions (39 binary and 2 triple systems) 14 are new detections including a third component in the previously known binary system ROXs 42B and the resolution of the previously known companion of L1689-IRS 5 into two sources.
- The surface density of the companions Σ as a function of the separation θ can be well fitted by the power law $\Sigma(\theta) \propto \theta^{-2.13 \pm 0.07}$.
- Within the range $0.13'' \leq \theta \leq 6.4''$ our multiplicity is $(29.1 \pm 4.3)\%$ for the total and $(26.6 \pm 4.4)\%$ for the restricted sample.
- This value is 1.24 ± 0.31 , respectively 1.13 ± 0.30 times the main-sequence value. The close similarity between Taurus and Ophiuchus found in most previous surveys is questioned by our result, which is based on a larger and more complete sample.
- The idea that the observed duplicity in star-forming regions is governed by some process related to the density of the stellar environment gets global support from our observations. This process has been suggested earlier to be related either to the formation process or to dynamical interaction afterwards. Observations like those of Duchêne (1999) and Haisch et al. (2002) tend to favour the second scenario. Our data are not sensitive to this alternative.
- There seems to be a relation between spectral classes and binary fraction. Class II objects have a multiplicity twice that of class III objects. This relation has not been found in the Taurus-Auriga survey (Köhler & Leinert 1998).
- Our results find their place in the paradigm of originally very high multiplicity of young stellar objects that then is reduced by dynamical interactions to different degrees in environments of different densities. This may be the global picture, however, locally within our sample we see no significant difference between the ρ Ophiuchi Dark Cloud (L1688) and its less dense environment. Only the differences between class II and class III sources may point to evolution.
- A population of close binaries with low flux ratios not present in Taurus, for which first indications exist, may be a partial answer to the question why the multiplicity in Taurus-Auriga is so clearly higher than in Ophiuchus.

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Online Material

Table 2. The 158 sources of our survey.

Object	ISO No.	H α No.	J2000.0		B1950.0		$K_{2\text{MASS}}$ [mag]	K_{BKLT} [mag]	Selection #	Note crit.	Other Designation	
H α 16	–	16	16 23 34.83	–23 40 30.3	16 20 34.46	–23 33 35.0	8.552	–	10	–v–	U	–
H α 18	–	18	16 24 59.74	–24 56 00.8	16 21 57.68	–24 49 11.1	9.441	–	10	–v–	–	–
H α 19	–	19	16 25 02.09	–24 59 32.3	16 21 59.95	–24 52 42.8	9.220	–	10	–v–	–	–
Haro 1-4	–	20	16 25 10.52	–23 19 14.6	16 22 10.51	–23 12 25.7	7.509	–	9	–v–	O	HBC 257
H α 21	–	21	16 25 15.22	–25 11 54.1	16 22 12.80	–25 05 05.4	9.642	–	10	–v–	U	–
H α 22	–	22	16 25 17.27	–25 11 05.5	16 22 14.87	–25 04 16.9	9.911	–	10	–v–	–	–
SR 2	–	–	16 25 19.24	–24 26 52.6	16 22 17.79	–24 20 04.2	7.798	–	1	x–i	A	Elias 6
SR 22	–	23	16 25 24.34	–24 29 44.3	16 22 22.83	–24 22 56.2	9.446	–	1	x v i	–	V 852 Oph, HBC 258
SR 1	–	–	16 25 24.31	–24 27 56.6	16 22 22.83	–24 21 08.5	4.582	–	1	x–i	AE	Elias 9
ROXs 2	–	–	16 25 24.35	–23 55 10.3	16 22 23.57	–23 48 22.3	8.379	–	6	x v–	B	HBC 635
SR 8	–	–	16 25 26.87	–24 43 09.0	16 22 25.06	–24 36 21.1	8.662	8.72	2	x––	–	–
IRS 2	1	–	16 25 36.74	–24 15 42.4	16 22 35.51	–24 08 55.2	8.380	8.36	1	x–i	–	–
Elias 12	–	–	16 25 37.81	–24 13 43.6	16 22 36.62	–24 06 56.5	6.867	6.95	11	––i	–	–
J162538-242238	2	–	16 25 38.12	–24 22 36.3	16 22 36.75	–24 15 49.1	9.545	9.68	7	x–i	–	–
IRS 3	3	–	16 25 39.58	–24 26 34.9	16 22 38.12	–24 19 47.9	8.954	8.93	2	x–i	–	–
H α 24	–	24	16 25 42.89	–23 25 26.1	16 22 42.73	–23 18 39.3	9.204	–	10	–v–	–	–
IRS 8	–	–	16 25 47.69	–24 18 31.1	16 22 46.39	–24 11 44.6	9.701	9.69	4	––i	–	–
IRS 9	–	–	16 25 49.05	–24 31 39.1	16 22 47.48	–24 24 52.7	9.487	9.41	1	x–i	–	–
ROXs 3	–	–	16 25 49.64	–24 51 31.9	16 22 47.64	–24 44 45.5	8.784	8.78	6	x v–	–	HBC 636
VSS 23	5	–	16 25 50.53	–24 39 14.5	16 22 48.78	–24 32 28.2	8.331	8.24	1	x v i	–	IRS 10, ROXs 4
ROXs 5	–	–	16 25 55.83	–23 55 10.0	16 22 55.02	–23 48 24.0	8.383	–	6	x v–	B	–
IRS 11	–	–	16 25 56.10	–24 30 14.9	16 22 54.54	–24 23 28.9	9.764	9.76	12	––i	U	–
SR 4	6	25	16 25 56.16	–24 20 48.2	16 22 54.81	–24 14 02.3	7.518	7.25	1	x v i	O	ROXs 6, HBC 259, Elias 13
GSS 20	7	–	16 25 57.52	–24 30 31.7	16 22 55.96	–24 23 45.9	8.319	8.37	1	x v i	–	ROXs 7
Chini 8	–	–	16 25 59.65	–24 21 22.3	16 22 58.28	–24 14 36.6	9.531	9.52	2	x––	–	–
ROXR1-12	8	–	16 26 01.61	–24 29 45.0	16 23 00.06	–24 22 59.4	9.629	9.55	7	x–i	–	J162601–242945
DoAr 21	10	–	16 26 03.02	–24 23 36.0	16 23 01.60	–24 16 50.6	6.227	6.16	1	x v i	O	ROXs 8, HBC 637, Elias 14
VSSG 19	11	–	16 26 03.29	–24 17 46.5	16 23 02.00	–24 11 01.0	9.115	9.07	5	x–i	–	–
Chini 11	–	–	16 26 08.01	–24 14 39.2	16 23 06.77	–24 07 54.1	9.812	9.79	12	––i	U	–
SR 3	16	–	16 26 09.31	–24 34 12.1	16 23 07.66	–24 27 27.1	6.504	6.42	1	x–i	A	Elias 16
GSS 26	17	–	16 26 10.33	–24 20 54.8	16 23 08.96	–24 14 09.8	8.475	9.38	1	x–i	–	–
SKS 1-7	18	–	16 26 15.81	–24 19 22.1	16 23 14.47	–24 12 37.5	10.031	9.98	8	––i	–	–
GSS 29	19	–	16 26 16.84	–24 22 23.2	16 23 15.44	–24 15 38.6	8.201	8.19	1	x–i	–	Elias 18
DoAr 24	20	27	16 26 17.06	–24 20 21.6	16 23 15.70	–24 13 37.1	8.063	8.09	1	x v i	O	GSS 28, ROXs 10A, HBC 638, Elias 19
H α 26	–	26	16 26 18.40	–25 20 55.7	16 23 15.73	–25 14 11.2	8.934	–	10	–v–	–	DoAr23
VSSG 1	24	–	16 26 18.87	–24 28 19.7	16 23 17.33	–24 21 35.3	8.072	8.68	1	x–i	–	Elias 20
DoAr 22	–	–	16 26 19.32	–23 43 20.5	16 23 18.75	–23 36 36.2	9.018	–	9	–v–	–	HBC 260
H α 28	27	28	16 26 20.97	–24 08 51.9	16 23 19.85	–24 02 07.6	9.502	9.48	7	x v i	C	–
J162621-241544	28	–	16 26 21.02	–24 15 41.5	16 23 19.75	–24 08 57.2	9.273	9.27	8	––i	–	–
Elias 21	29	–	16 26 21.38	–24 23 04.1	16 23 19.96	–24 16 19.8	8.835	8.32	1	x–i	CE	GSS 30 – IRS 1
GSS 30 - IRS 2	34	–	16 26 22.39	–24 22 52.9	16 23 20.96	–24 16 08.8	9.957	9.60	1	x–i	–	–
DoAr 24E	36	30	16 26 23.36	–24 20 59.8	16 23 21.98	–24 14 15.7	6.571	6.44	1	x v i	O	GSS 31, ROXs 10B, HBC 639, Elias 22
LFAM 3	37	–	16 26 23.58	–24 24 39.5	16 23 22.12	–24 17 55.4	10.224	9.94	1	x–i	–	GY 21
DoAr 25	38	29	16 26 23.68	–24 43 13.9	16 23 21.82	–24 36 29.8	7.847	7.57	1	x v i	A	–
GSS 32	39	–	16 26 24.04	–24 24 48.1	16 23 22.58	–24 18 04.0	7.324	7.20	1	x–i	C	S 2, Elias 23, GY 23
Elias 24	40	31	16 26 24.07	–24 16 13.5	16 23 22.79	–24 09 29.4	6.685	6.77	1	x v i	–	–
H α 33	–	33	16 26 26.06	–23 44 02.6	16 23 25.46	–23 37 18.7	9.928	–	10	–v–	U	–
GY 33	43	–	16 26 27.54	–24 41 53.5	16 23 25.70	–24 35 09.7	9.983	9.83	8	––i	–	–
ROXs 12	–	32	16 26 27.75	–25 27 24.7	16 23 24.93	–25 20 40.8	9.211	–	6	x v–	–	–
VSSG 27	46	–	16 26 30.47	–24 22 57.1	16 23 29.04	–24 16 13.5	9.977	10.72	1	x–i	–	–
S1	48	–	16 26 34.17	–24 23 28.3	16 23 32.72	–24 16 44.9	6.317	6.32	1	x v i	–	ROXs 14, Elias 25
H α 35	–	35	16 26 34.90	–23 45 40.6	16 23 34.26	–23 38 57.3	8.834	–	10	–v–	–	–

Table 2. continued.

Object	ISO No.	H α No.	J2000.0		B1950.0		K_{2MAS} [mag]	K_{BKL} [mag]	Selection #	Note crit.	Other Designation
			α	δ	α	δ					
J162636-241554	51	–	16 26 36.83	–24 15 51.9	16 23 35.54	–24 09 08.7	9.589	9.38	7	x–i	–
H α 37	56	37	16 26 41.26	–24 40 18.0	16 23 39.44	–24 33 35.0	9.273	9.20	1	x v i	–
WL 8	58	–	16 26 42.02	–24 33 26.2	16 23 40.35	–24 26 43.4	9.578	9.44	8	–i	–
GSS 37	62	–	16 26 42.86	–24 20 29.9	16 23 41.47	–24 13 47.1	7.878	8.00	1	x–i	VSSG 2, Elias 26
VSSG 11	64	–	16 26 43.76	–24 16 33.3	16 23 42.45	–24 09 50.6	9.604	9.58	1	x–i	–
GY 112	66	–	16 26 44.30	–24 43 14.1	16 23 42.42	–24 36 31.4	9.573	9.54	2	x–i	–
GSS 39	67	–	16 26 45.03	–24 23 07.7	16 23 43.58	–24 16 25.1	8.955	8.88	1	x–i	Elias 27
ROXs 16	68	38	16 26 46.43	–24 12 00.1	16 23 45.22	–24 05 17.5	7.485	7.51	4	x v i	VSS 27
Haro 1-8	–	39	16 26 47.42	–23 14 52.2	16 23 47.42	–23 08 09.7	8.619	–	9	–v–	B HBC 261
H α 40	–	40	16 26 48.65	–23 56 34.2	16 23 47.76	–23 49 51.8	8.449	–	10	–v–	–
WL 18	72	–	16 26 48.98	–24 38 25.2	16 23 47.20	–24 31 42.8	9.977	9.82	1	x–i	E
VSSG 3	73	–	16 26 49.23	–24 20 02.9	16 23 47.85	–24 13 20.5	8.687	8.62	1	x–i	–
VSSG 10	–	–	16 26 51.69	–24 14 41.6	16 23 50.41	–24 07 59.3	9.713	9.77	12	–i	U
VSSG 5	78	–	16 26 54.44	–24 26 20.7	16 23 52.92	–24 19 38.7	10.014	9.88	1	x–i	–
GY 156	80	–	16 26 54.97	–24 22 29.7	16 23 53.53	–24 15 47.7	10.163	10.19	3	x–i	–
VSSG 7	81	–	16 26 55.31	–24 20 27.8	16 23 53.91	–24 13 45.8	9.789	9.69	8	–i	–
J162656-241353	83	–	16 26 56.77	–24 13 51.6	16 23 55.51	–24 07 09.7	9.251	9.57	8	–i	–
SR 24	88	41/42	16 26 58.51	–24 45 36.9	16 23 56.56	–24 38 55.1	7.057	7.08	1	x v i	CO HBC 262, Elias 28
VSSG 8	91	–	16 27 01.62	–24 21 37.0	16 24 00.19	–24 14 55.5	9.389	9.32	8	–i	–
H α 44	–	44	16 27 02.37	–23 09 59.2	16 24 02.46	–23 03 17.8	9.435	–	10	–v–	U
WL 16	92	–	16 27 02.34	–24 37 27.2	16 24 00.57	–24 30 45.7	8.064	7.92	1	x–i	–
VSSG 9	–	–	16 27 02.85	–24 18 54.7	16 24 01.48	–24 12 13.2	10.116	10.07	12	–i	U
GY 193	96	–	16 27 04.52	–24 42 59.7	16 24 02.62	–24 36 18.3	9.837	9.80	2	x–i	–
GY 194	97	–	16 27 04.56	–24 42 14.0	16 24 02.69	–24 35 32.6	9.809	9.82	2	x–i	–
VSSG 21	–	–	16 27 05.16	–24 20 07.7	16 24 03.76	–24 13 26.4	9.374	9.27	12	–i	U
J162708-241204	106	–	16 27 09.07	–24 12 00.8	16 24 07.84	–24 05 19.7	9.800	9.76	8	–i	–
WL 10	105	–	16 27 09.10	–24 34 08.1	16 24 07.40	–24 27 27.1	8.915	8.85	1	x–i	–
Elias 29	108	–	16 27 09.43	–24 37 18.8	16 24 07.66	–24 30 37.7	7.140	7.54	1	x–i	–
Elias 30	110	–	16 27 10.28	–24 19 12.7	16 24 08.89	–24 12 31.8	6.719	6.30	1	x–i	C SR 21
GY 224	112	–	16 27 11.18	–24 40 46.7	16 24 09.33	–24 34 05.7	10.196	10.79	1	x–i	–
IRS 32	113	–	16 27 11.68	–24 23 42.0	16 24 10.20	–24 17 01.1	10.107	10.06	5	–i	–
VSSG 24	116	–	16 27 13.73	–24 18 16.9	16 24 12.36	–24 11 36.1	9.287	9.32	8	–i	–
IRS 32b	117	–	16 27 13.82	–24 43 31.7	16 24 11.91	–24 36 50.9	9.978	10.13	8	–i	O
ROXs 20A	–	45	16 27 14.49	–24 51 33.5	16 24 12.40	–24 44 52.8	10.381	10.39	5	x v i	HBC 640
ROXs 20B	–	46	16 27 15.13	–24 51 38.8	16 24 13.05	–24 44 58.1	9.392	9.51	5	x v i	HBC 641
WL 20	121	–	16 27 15.88	–24 38 43.4	16 24 14.07	–24 32 02.8	9.590	9.21	1	x–i	CE
H α 47	–	47	16 27 17.08	–24 47 11.2	16 24 15.08	–24 40 30.6	9.487	9.47	10	–v–	–
WL 5	125	–	16 27 18.17	–24 28 52.7	16 24 16.57	–24 22 12.2	10.558	10.28	1	x–i	L
WL 4	128	–	16 27 18.49	–24 29 05.9	16 24 16.88	–24 22 25.5	9.683	9.13	1	x–i	–
SR 12	130	–	16 27 19.51	–24 41 40.4	16 24 17.64	–24 35 00.0	8.408	8.41	1	x v i	O ROXs 21, HBC 263
IRS 42	132	–	16 27 21.47	–24 41 43.1	16 24 19.59	–24 35 02.8	8.483	8.41	1	x–i	–
WL 6	134	–	16 27 21.80	–24 29 53.4	16 24 20.18	–24 23 13.1	10.827	10.04	1	x–i	L
VSSG 22	135	–	16 27 22.91	–24 17 57.4	16 24 21.54	–24 11 17.2	9.454	9.41	1	x–i	–
H α 49	–	49	16 27 22.97	–24 48 07.1	16 24 20.95	–24 41 27.0	9.390	9.30	1	x v i	–
GY 262	140	–	16 27 26.49	–24 39 23.1	16 24 24.66	–24 32 43.2	9.952	9.77	1	x–i	–
IRS 43	141	–	16 27 26.94	–24 40 50.8	16 24 25.07	–24 34 10.9	9.745	9.46	1	x–i	YLOW 15A
VSSG 25	142	–	16 27 27.38	–24 31 16.6	16 24 25.72	–24 24 36.7	9.316	9.30	1	x–i	Elias 31
IRS 44	143	–	16 27 28.03	–24 39 33.5	16 24 26.19	–24 32 53.7	10.379	9.65	1	x–i	YLOW 16A
VSSG 18	144	–	16 27 28.45	–24 27 21.0	16 24 26.87	–24 20 41.3	10.101	9.39	1	x–i	C Elias 32
VSSG 17	147	–	16 27 30.18	–24 27 43.4	16 24 28.59	–24 21 03.7	9.024	8.95	1	x–i	C Elias 33
GY 284	151	–	16 27 30.84	–24 24 56.0	16 24 29.32	–24 18 16.4	10.070	10.04	3	x–i	–
J162730-244726	149	–	16 27 30.84	–24 47 26.8	16 24 28.83	–24 40 47.2	9.502	9.44	7	x–i	–
GY 292	155	–	16 27 33.11	–24 41 15.3	16 24 31.23	–24 34 35.8	7.806	7.92	1	x–i	–
H α 50	156	50	16 27 35.26	–24 38 33.4	16 24 33.44	–24 31 54.1	9.668	9.64	5	–v–	GY 295
IRS 48	159	–	16 27 37.19	–24 30 35.0	16 24 35.54	–24 23 55.8	7.582	7.42	1	x–i	–
IRS 50	–	–	16 27 38.13	–24 30 42.9	16 24 36.47	–24 24 03.8	9.658	9.59	1	x–i	–

Table 2. continued.

Object	ISO No.	H α No.	J2000.0		B1950.0		$K_{2\text{MASS}}$ [mag]	K_{BKLT} [mag]	Selection #	Note crit.	Other Designations
			α	δ	α	δ					
IRS 49	163	–	16 27 38.32	–24 36 58.6	16 24 36.53	–24 30 19.5	8.271	8.31	1	x – i	–
ROXs 30B	–	51	16 27 38.33	–23 57 32.4	16 24 37.38	–23 50 53.3	7.940	–	6	x v –	DoAr 32
ROXs 30C	–	53	16 27 39.01	–23 58 18.7	16 24 38.05	–23 51 39.7	8.206	–	6	x v –	–
H α 52	166	52	16 27 39.43	–24 39 15.5	16 24 37.59	–24 32 36.5	8.464	8.35	1	x v i	GY 314
IRS 51	167	–	16 27 39.83	–24 43 15.1	16 24 37.90	–24 36 36.0	8.991	8.93	1	x – i	–
SR 9	168	54	16 27 40.29	–24 22 04.0	16 24 38.81	–24 15 25.0	7.207	7.20	1	x v i	ROXs 29, HBC 264, Elias 34
GY 371	178	–	16 27 49.78	–24 25 22.0	16 24 48.23	–24 18 43.6	10.161	10.17	8	– – i	–
VSSG 14	180	–	16 27 49.87	–24 25 40.2	16 24 48.31	–24 19 01.9	7.301	7.32	1	x – i	Elias 36
IRS 56	–	–	16 27 50.74	–24 48 21.6	16 24 48.69	–24 41 43.3	8.337	8.23	12	– – i	U
ROXs 31	184	–	16 27 52.09	–24 40 50.4	16 24 50.20	–24 34 12.2	8.126	8.09	1	x v i	IRS 55, HBC 642
SR 10	187	57	16 27 55.58	–24 26 17.9	16 24 54.00	–24 19 39.9	8.896	8.74	1	x v i	HBC 265
GY 410	188	–	16 27 57.83	–24 40 01.8	16 24 55.95	–24 33 24.0	9.866	9.78	1	x – i	–
H α 58	–	58	16 27 59.97	–24 48 19.3	16 24 57.91	–24 41 41.6	9.269	9.26	2	x v –	U
J162800-245340	–	–	16 28 00.11	–24 53 42.7	16 24 57.94	–24 47 05.1	9.651	9.63	7	x – –	U
H α 59	–	59	16 28 09.21	–23 52 20.5	16 25 08.35	–23 45 43.5	9.075	–	10	– v –	U
VSS 35	–	–	16 28 10.22	–24 16 01.0	16 25 08.85	–24 09 24.0	7.965	7.89	12	– – i	U
J162812-245043	–	–	16 28 12.28	–24 50 44.6	16 25 10.16	–24 44 07.7	9.555	9.36	7	x – –	CU
J162813-243249	194	–	16 28 13.79	–24 32 49.4	16 25 12.06	–24 26 12.7	10.096	10.04	8	– – i	–
H α 60	196	60	16 28 16.51	–24 36 58.0	16 25 14.68	–24 30 21.4	9.316	9.43	4	– v i	–
ISO-Oph 195	195	–	16 28 16.73	–24 05 14.3	16 25 15.59	–23 58 37.7	8.860	–	8	– – i	–
SR 20 W (GWAYL)	–	–	16 28 23.33	–24 22 40.6	16 25 21.81	–24 16 04.5	8.623	8.55	5	– – i	–
SR 20	198	61	16 28 32.66	–24 22 44.9	16 25 31.13	–24 16 09.4	6.850	7.16	4	x v i	O ROXs 33, HBC 643
V 853 Oph	199	62	16 28 45.28	–24 28 19.0	16 25 43.61	–24 21 44.4	7.997	7.88	6	x v i	O ROXs 34, HBC 266
VSS 38	–	–	16 28 45.98	–24 47 55.3	16 25 43.89	–24 41 20.7	5.960	6.14	12	– – i	U
H α 63	–	63	16 28 54.07	–24 47 44.2	16 25 51.98	–24 41 10.2	8.905	8.96	4	– v i	–
VSS 42	–	–	16 29 12.73	–24 23 55.3	16 26 11.14	–24 17 22.6	5.825	6.27	12	– – i	U
IRAS 64a	–	–	16 29 23.39	–24 13 56.9	16 26 22.00	–24 07 24.8	7.020	6.45	5	– – i	–
VSS 41	–	–	16 29 45.12	–24 19 50.5	16 26 43.59	–24 13 19.9	8.664	–	12	– – i	U
Elias 41	–	–	16 30 02.41	–23 51 09.1	16 27 01.47	–23 44 39.7	6.905	–	11	– – i	–
H α 67	–	67	16 30 23.40	–24 54 16.2	16 27 21.08	–24 47 48.1	9.293	–	10	– v –	–
ROXs 39	–	–	16 30 35.63	–24 34 18.9	16 27 33.74	–24 27 51.7	8.025	–	5	x v i	B
Haro 1-14/c	–	–	16 31 04.37	–24 04 33.1	16 28 03.09	–23 58 07.8	7.784	–	9	– v –	B HBC 644
Haro 1-14	–	69	16 31 05.17	–24 04 40.1	16 28 03.89	–23 58 14.9	8.576	–	9	– v –	B HBC 267
ROXs 42B	–	–	16 31 15.02	–24 32 43.7	16 28 13.13	–24 26 19.1	8.671	–	6	x v –	B
ROXs 42C	–	–	16 31 15.75	–24 34 02.2	16 28 13.83	–24 27 37.7	7.129	–	6	x v –	–
ROXs 43A/B	–	–	16 31 20.12	–24 30 05.2	16 28 18.29	–24 23 41.0	6.729	–	5	x v i	C GWAYL 1
H α 71	–	71	16 31 30.88	–24 24 40.0	16 28 29.15	–24 18 16.4	7.900	–	5	– v i	BC GWAYL 2
Haro 1-16	–	72	16 31 33.46	–24 27 37.3	16 28 31.67	–24 21 13.9	7.610	–	5	x v i	O GWAYL 3, ROXs 44, HBC 268
IRS 63	–	–	16 31 35.66	–24 01 29.5	16 28 34.42	–23 55 06.3	9.219	–	5	– – i	– GWAYL 4
L1689-IRS 5	204	–	16 31 52.11	–24 56 15.7	16 28 49.68	–24 49 53.6	7.557	–	5	– – i	– GWAYL 5
H α 73	206	73	16 31 54.42	–25 03 49.3	16 28 51.82	–24 57 27.4	9.899	–	8	– v i	–
H α 74	207	74	16 31 54.73	–25 03 23.8	16 28 52.14	–24 57 01.9	7.749	–	8	– v i	–
ROXs 45D	–	–	16 31 57.68	–25 29 33.7	16 28 54.51	–25 23 11.9	9.865	–	6	x v –	–
ROXs 45E	–	–	16 32 00.59	–25 30 28.7	16 28 57.39	–25 24 07.1	9.478	–	6	x v –	–
ROXs 45F	–	–	16 32 01.61	–25 30 25.3	16 28 58.41	–25 24 03.8	9.395	–	6	x v –	–
H α 75	–	75	16 32 05.52	–25 02 36.2	16 29 02.94	–24 56 15.0	9.947	–	10	– v –	U
DoAr 51	–	76	16 32 11.79	–24 40 21.4	16 29 09.69	–24 34 00.6	7.929	–	6	x v –	– ROXs 47A, HBC 647
L1689-IRS 7	212	–	16 32 21.05	–24 30 35.8	16 29 19.15	–24 24 15.7	8.620	–	5	– – i	B GWAYL 7
Haro 1-17	–	77	16 32 21.93	–24 42 14.8	16 29 19.77	–24 35 54.7	9.151	–	9	– v –	– HBC 648
Elias 45	–	–	16 33 21.54	–23 50 21.4	16 30 20.46	–23 44 05.3	6.345	–	11	– – i	–

Names adopted from Barsony et al. (1997) are given without the leading “BKLT” and thus start with “J16”.

A: observed with ADONIS/SHARP II+

B: observed with BlackMAGIC

C: companion in 2MASS

E: 2MASS photometric quality flag E

L: $m_K > 10.5$ (see text)

O: well observed

Selection criteria: X-ray detection (x), visual spectra / H α -emission (v), infrared excess (i) – Selection # refers to Table 1.

Appendix A: Comparison with previously known binary and multiple systems

Barsony et al. (2003) compiled a list of all known binary and multiple systems associated with the ρ Ophiuchi Dark cloud. In Table A.1 we list all sources of this compilation with a separation in the range $0.05'' \leq \theta \leq 6.7''$ that have been newly observed within the scope of our survey. The first column gives the name used in our survey. If Barsony et al. (2003) used another designation it appears in the last column. The second column gives their references. The references printed in bold font are those that we used to derive the values for the position angles, the separations and the flux ratios in the Cols. 4, 5, and 6. The date of observation in this reference is given in Col. 3. For comparison our results are listed in the subsequent columns.

Only one companion was not detected in our survey. This is probably caused by the small separation below the diffraction limit of our telescope. Otherwise, we detected additional companions of ROXs 42B and L1689-IRS 5 transforming these objects into triple systems. The new companion of ROXs 42B (Fig. A.1) is a faint knot west of the primary. In the case of L1689-IRS 5 (Fig. A.2) the previously known companion splits up into two point sources with equal fluxes. They are visible as a elongated structure in the shift-and-add images and are very prominent in the visibility. Unfortunately, a probable third component of ROXs 16 is too close to distinguish between an elongated structure and a point source.

Although the binary parameters derived in our analysis are in general very similar to those provided by the papers used as references in Table A.1, some important differences exist. The position angle of the binaries ROXs 5, VSSG 25 and IRS 44 have changed approximately 180° since the last observations. Due to the fact that the separations did not change either the flux ratio has changed significantly or the measurements suffer from the 180° ambiguity. Furthermore, some flux ratios have changed at least by a factor of two: ROXs 16 (maybe mainly an effect of the probable new companion), VSSG 3, Elias 30, VSSG 17, and of the close companion of ROXs 42B. Finally, the relative positions of the companions orbiting SR 2 and ROXs 42C have changed significantly.

Appendix B: Additional sources

During our speckle observations we recorded 12 sources not part of the final sample (see Table B.1). BKLT J162637-241602 has been observed by chance due to its location close to BKLT J162636-241554. Although VSS 28 is probably not a member of the Ophiuchus molecular clouds, we observed it with speckle techniques, because the area around VSS 28 has been used to determine the background density. While DoAr 58 was listed as HBC 269 in Herbig & Kameswara Rao (1972), it has been removed in the third edition of this catalogue (Herbig & Bell 1988). The remaining objects are either background giants or foreground dwarfs.

VSS 28 is a binary that has been observed with a separation of $0.344'' \pm 0.005''$ at a position angle of $308.3^\circ \pm 0.7^\circ$ in the night following the 4th July 2001. The flux ratio of the two components is 0.203 ± 0.026 and we can exclude at a distance

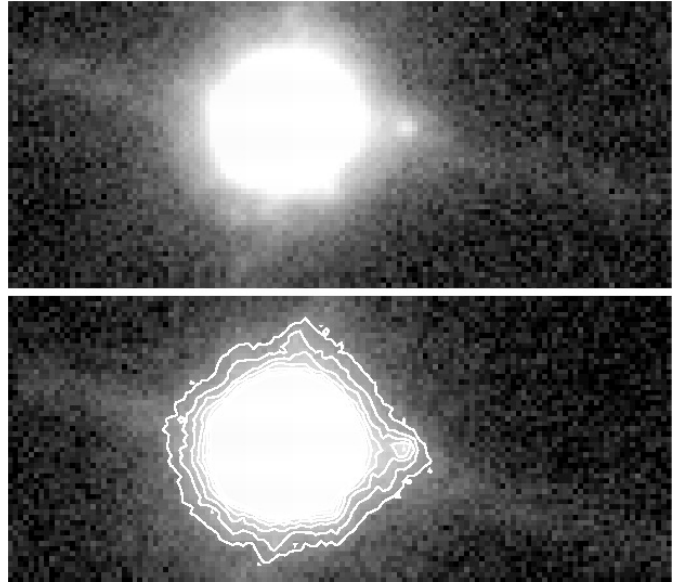


Fig. A.1. The faint companion of ROXs 42B as seen in the shift-and-add images of the two fitscubes. The images are printed in logarithmic scale. Linear contours are overlaid on the second image. The flux ratio is 0.002.

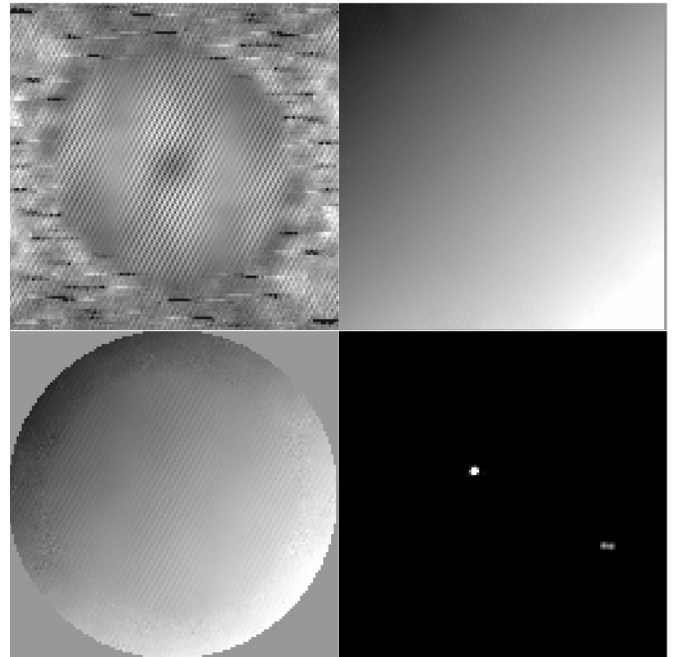


Fig. A.2. The newly detected triple system L1689-IRS 5. Displayed are the visibility, the Knox-Thompson phase, the bispectrum phase, and the reconstructed image. The visibility clearly shows that the wide pair (narrow stripes) is superimposed by a close pair almost at the diffraction limit (the two wide stripes).

of $0.15''$ ($0.50''$) from the main component companions with flux ratios larger than 0.07 (0.04). Another binary is GY 45. The 2MASS images of this source show a very symmetric extension to the south that has been detected in all three wave-length bands. We inspected our shift-and-add images and found a shallow glow at the calculated position only 2σ above the background. The parameters of this binary can be derived from

Table A.1. Comparison to previously known binaries.

Object	Ref	Date	PA _{Lit} [deg]	θ _{Lit} [″]	K ₂ /K _{1,Lit}	PA [deg]	θ [″]	K ₂ /K ₁	Remark
Hα 18	K, R1	1999, June 1/2	339.55 ^m	≈0.1004 ^d	≈0.505	–	–	–	not detected
			80.4	1.08	0.7	82.3	1.083	0.737	
Hα 19	K, R1	1999, June 1/2	260.7	1.53	0.5	262.9	1.491	0.462	
SR 2	A2, G2	1990, July 8	156	0.236	0.80	122.4	0.222	0.874	ROX 1
ROXs 2	B2, C	2002, May 24	347.1	0.42	0.57	345.5	0.424	0.598	
IRS 2	B2, C	2002, May 24	77.6	0.42	0.13	78.6	0.426	0.132	
ROXs 5	A1	1993 / 1994	≈130	≈0.13	≈0.5	327.3	0.176	0.408	obs. in <i>H</i> -band
Hα 26	K, R1	1999, June 1/2	23.8	1.15	0.9	25.8	1.135	0.846	
Hα 28	R1	1991 / 1992	358	5.1	0.06	357.8	5.209	0.047	obs. with Gunn z
VSSG 27	C	1995 / 1996	68	1.22	0.24	66.8	1.222	0.244	
Hα 35	K, R1	1999, June 1/2	130.3	2.29	0.3	132.2	2.277	0.272	
GSS37	C, K, R1	1999, June 1/2	67.0	1.44	0.3	69.5	1.438	0.299	
ROXs 16	A1, C	1995 / 1996	–	–	–	24.2	0.098 ^d	0.357	VSS 27, triple?
			106	0.57	0.06	105.4	0.577	0.186	
WL 18	B1	1988, June 2/3	293	3.55	0.2	292.4	3.617	0.162	rel. pos. from B2
VSSG 3	C	1995 / 1996	47	0.25	0.38	53.8	0.243	0.801	
Elias 30	M1, S2	1992, June 13	175	6.700	0.030	175.6	6.388	0.063	SR 21
WL 20	R2	1990 / 1998	270.1	3.17	0.70 ⁿ	269.9	3.198	0.877	phot. 1990, pos. 1998
			232.2	3.66	0.07 ⁿ	232.3	3.619	0.071	in the mid-IR
VSSG 25	C	1995 / 1996	356	0.46	0.5	173.3	0.468	0.887	WL 13
IRS 44	C, H, S2, T	1997	81	0.27	–	246.6	0.256	0.204	
VSSG 17	C	1995 / 1996	269	0.25	0.2	260.2	0.242	0.644	
SR 9	B2, G1, G2	1990, July 9	350	0.59	0.09	353.3	0.638	0.057	
VSSG 14	S2	1992, June 13	89 ^p	0.101 ^{d,p}	0.5	83.6	0.130	0.296	
ROXs 31	A1, C, S1, S2	1986 / 1991	262	0.480	0.8	251.3	0.396	0.655	
ROXs 42B	S2	1992, June/July	89 ^p	0.056 ^{d,p}	0.8	157.9	0.083 ^d	0.350	new triple system
			–	–	–	268.0	1.137	0.002	
ROXs 42C	B2, G2, M2	1990, July 8	135	0.157	0.25	151.0	0.277	0.220	
ROXs 43 A/B	A1, M1, R1, S2	1992, July 11	7	4.800	0.44	11.9	4.523	0.445	
Hα 71	K, S2	1999, June 1/2	35.0	3.56	0.17	35.0	3.560	0.151	calibrator
L1689 - IRS 5	H	2001, July 11	240.3	2.92	0.61	241.2	3.006	0.277	new triple system
			–	–	–	84.4	0.140	0.946	
DoAr 51	B2	2002, May 24	80.8	0.79	0.2	79.3	0.784	0.228	ROXs 47A
A1: Ageorges et al. (1997)		A2: Aitken & Doolittle (1932)			B1: Barsony et al. (1989)		B2: Barsony et al. (2003)		
C: Costa et al. (2000)		G1: Geoffray & Monin (2001)			G2: Ghez et al. (1993)		H: Haisch et al. (2002)		
K: Koresko (2002)		M1: 2MASS			M2: Mathieu et al. (1989)		R1: Reipurth & Zinnecker (1993)		
R2: Ressler & Barsony (2001)		R3: Richichi et al. (1994)			S1: Simon et al. (1987)		S2: Simon et al. (1995)		
T: Terebey et al. (2001)									
^d Below our diffraction limit		^m Mod 180 deg			ⁿ Non-photometric conditions		^p Projected values		

the 2MASS survey. The companion resides at a position angle of $168.4^\circ \pm 1.8^\circ$ with a separation of $4.977'' \pm 0.144''$ and a flux ratio of 0.017 ± 0.003 . In our speckle data obtained at the 17th June 2000 we can exclude at a distance of $0.15''$ ($0.50''$) from the main component companions brighter than 0.04 (0.03) times the flux of the primary. All the remaining ten sources are single stars.

Table B.1. Additional sources.

Object	J2000.0		$K_{2\text{MASS}}$ [mag]	Date	0.15''	0.50''	Remarks
	α	δ					
GY 45	16 26 29.98	−24 38 42.8	8.367	2000, June 17	0.04	0.03	binary, background giant (Luhman & Rieke 1999)
GY 65	16 26 32.91	−24 36 26.4	8.996	2000, June 17	0.08	0.05	background giant (Luhman & Rieke 1999)
BKLT J162637-241602	16 26 37.13	−24 15 59.9	10.757	2001, July 3	0.15	0.08	too faint
VSS 28	16 26 52.80	−23 43 12.7	6.702	2001, July 4	0.07	0.04	binary, background determination, not a member
VSSG 6	16 26 53.86	−24 22 28.0	9.827	2000, June 21	0.09	0.04	background giant (Luhman & Rieke 1999)
GY 232	16 27 13.26	−24 41 33.7	9.592	2000, June 21	0.10	0.07	background giant (Luhman & Rieke 1999)
VSSG 13	16 27 46.69	−24 23 22.1	7.270	2000, June 21	0.03	0.02	field star (Elias 1978)
GY 411	16 27 57.89	−24 37 49.0	9.560	2000, June 21	0.07	0.03	background giant (Luhman & Rieke 1999)
VSSG 16	16 28 03.73	−24 26 32.0	6.504	2000, June 17	0.04	0.03	field star (Elias 1978)
VSSG 15	16 28 09.23	−24 23 20.7	7.140	2000, June 21	0.07	0.03	field star (Elias 1978)
ROXs 47B	16 32 23.28	−24 40 18.5	8.581	2000, June 22	0.04	0.04	foreground (Bouvier & Appenzeller 1992)
DoAr 58	16 34 26.70	−24 13 43.7	7.707	2001, July 4	0.15	0.04	HBC 269, deleted in Herbig & Bell (1988)