

have been found practically constant for absolute magnitudes fainter than $+3^m$, disregarding, of course, the absolutely faint white stars mentioned above.

It would be of interest to see, whether the colour indices of the absolutely faint yellow stars show the same behaviour as their effective wavelengths.

List of stars nearer then 5 parsecs.

Star	α (1900)	δ (1900)	vis. mag. Harvard	parallax	abs. mag.	spectrum	proper motion
	$^h \quad ^m$		m	"	m		" / y
α Canis majoris <i>A</i>	6 40.7	$- 16^\circ 34'$	$- 1.58$	.376	$- 3.70$	<i>A0</i>	1.32
α Aquilae	19 45.9	$+ 8 36$	0.89	.200	$- 2.60$	<i>A5</i>	.65
α Canis minoris <i>A</i>	7 34.1	$+ 5 29$	0.48	.312	$- 2.05$	<i>F3</i>	1.24
☉			$- 26.9$		$- 0.33$	<i>G0</i>	
α Centauri <i>A</i>	14 32.8	$- 60 25$	0.33	.759	$- 0.27$	<i>G6</i>	3.66
α Centauri <i>B</i>	14 32.8	$- 60 25$	1.70	.759	1.10	<i>K4</i>	3.66
τ Ceti	1 39.4	$- 16 28$	3.65	.318	1.16	<i>G7</i>	1.92
ϵ Eridani	3 28.2	$- 9 48$	3.81	.311	1.27	<i>K1</i>	.97
ϵ Indi	21 55.7	$- 57 12$	4.74	.284	2.01	<i>K5</i>	4.67
61 Cygni <i>A</i>	21 2.4	$+ 38 15$	5.57	.306	3.00	<i>K7</i>	5.21
Lac 8760	21 11.4	$- 39 15$	6.65	.251	3.65	<i>Ma</i>	3.53
61 Cygni <i>B</i>	21 2.4	$+ 38 15$	6.28	.306	3.71	<i>K8</i>	5.21
Lac 9352	22 59.4	$- 36 26$	7.44	.292	4.77	<i>Ma</i>	7.02
Gou 32416	23 59.5	$- 37 51$	8.5	.203	5.04	<i>K5</i>	6.11
Groombr 34 <i>A</i>	0 12.5	$+ 43 27$	7.98	.281	5.22	<i>Ma</i>	2.85
Lal 21185	10 57.9	$+ 36 38$	7.60	.403	5.63	<i>Ma</i>	4.77
ZC 5 ^h 243	5 7.7	$- 44 59$	8.3	.319	5.82	<i>K2</i>	8.70
Oe Arg 17415-6	17 37.0	$+ 68 26$	9.2	.247	6.16	<i>Mb</i>	1.31
α Canis majoris <i>B</i>	6 40.7	$- 16 34$	8.44	.376	6.32	(<i>A</i>)	1.32
Σ 2398 <i>A</i>	18 41.8	$+ 59 29$	9.33	.287	6.62	<i>Mb</i>	2.28
Krüger 60 <i>A</i>	22 24.5	$+ 57 12$	9.64	.262	6.73	<i>Mb</i>	.94
Σ 2398 <i>B</i>	18 41.8	$+ 59 29$	10.01	.287	7.30	[<i>M</i>]	2.28
Groombr 34 <i>B</i>	0 12.5	$+ 43 27$	11.05	.281	8.29	[<i>M</i>]	2.85
(Barnard)	17 52.9	$+ 4 28$	9.67	.533	8.30	<i>Mb</i>	10.30
Krüger 60 <i>B</i>	22 24.5	$+ 57 12$	11.34	.262	8.43	[<i>M</i>]	.94
(van Maanen)	0 43.9	$+ 4 55$	12.34	.246	9.29	<i>F0</i>	3.01
(Innes)	11 12.0	$- 57 2$	(12)	.339	(9.65)	—	2.69
α Centauri <i>C</i>	14 22.9	$- 62 2$	(10.5)	(.759)	(9.90)	[<i>M</i>]	(3.66)
α Canis minoris <i>B</i>	7 34.1	$+ 5 29$	(12.5)	.312	(9.97)	—	1.24

Remark on the period of VV Orionis, by *Ejnar Hertzsprung*.

In Comptes Rendus **173**, 969, 1921 H. GROUILLER and J. ELLSWORTH have published new elements of VV Orionis based on 293 visual estimates made by ELLSWORTH (*Bull. de l'Obs. de Lyon* N°. 8 and 9, 1921 *). A correction of $+ .0084$ to the period $1^d 48^m 54^s$

*) For J. D. 2421676 read 1876 and for J. D. 2421938 read 1928. Further the J. D. of the two observations made at or after midnight have to be increased by 0.5.

used by me (*Potsdam Publ. Nr. 73*, 1919) is derived.

However, the new period $1^d 49^m 38^s$ is irreconcilable with my observations, as is easily seen. At J. D. 2420053 and 2420175 I have observed the star during decreasing light before the primary minimum. By comparison between the plots of these two series of observations the time interval between two epochs of equal brightness is found to be $121^d 80^m 2^s \pm 0^m 04^s$