HERTZSPRUNG for his help and advice. At the same time I wish to express my thanks to the Italian "Consiglio Nazionale delle Ricerche", for a grant, by which I have been enabled to study different methods of observation at foreign observatories and in that way to complete my astronomical education.

TABLE 4.

n	phase	brightness	n	phase	brightness	n	phase	brightness	n	phase	brightness
	* I			P	m		* III			P	m
			10	·209	13·35				10	·177	13·13
	P	m	10	·252	13·33		P	m	10	·221	13·22
10	·006	12·91	10	·298	13·36	50	·068	13·00	10	·248	13·24
4	·046	13·02	12	·344	13·38	50	·224	12·99	10	·282	13·15
2	·053	13·30	9	·379	13·39	50	·399	13·00	10	·315	13·30
1	·066	13·81	10	·419	13·44	40	·551	13·00	10	·340	13·29
2	·073	13·97	10	·443	13·45	10	·635	12·99	10	·373	13·33
4	·080	13·97	9	·471	13·50	6	·660	13·06	10	·392	13·32
4	·088	13·99	11	·502	13·51	4	·682	13·14	10	·417	13·29
2	·095	13·96	13	·539	13·50	4	·694	13·18	10	·457	13·47
2	·119	13·22	9	·575	13·47	7	·723	13·35	10	·498	13·57
11	·133	13·01	10	·616	13·42	6	·750	13·45	10	·528	13·59
10	·157	12·96	10	·646	13·40	7	·765	13·39	10	·568	13·68
50	·279	12·93	10	·673	13·36	8	·793	13·34	10	·600	13·74
50	·437	12·92	10	·701	13·38	6	·820	13·17	10	·656	13·78
50	·606	12·92	10	·725	13·37	5	·843	13·05	10	·692	13·84
50	·772	12·91	10	·752	13·33	10	·872	13·03	10	·719	13·85
50	·928	12·92	10	·781	13·36	40	·947	13·00	10	·752	13·85
			10	·815	13·37				10	·802	13·81
	* II		10	·849	13·38		* IV		10	·836	13·80
			10	·888	13·40				10	·873	13·70
10	·018	13·53	10	·922	13·45	10	·026	12·91	10	·901	13·60
10	·071	13·47	10	·954	13·52	10	·045	12·88	10	·931	13·45
10	·112	13·41	10	·984	13·53	10	·090	12·94	10	·965	13·14
10	·147	13·37				10	·127	13·03	10	·989	13·03
10	·178	13·38				10	·161	13·06			

Results of observations of stars in the region of the Hyades, by *G. van Herk.*

1. *The programme* consisted of 213 stars chosen within a circle of about 15° radius round θ Tauri. The list has been compiled by Mr. MARTIN from the A. G. Zones: Alb, Be A, Lpz I and Lpz II, taking all the stars with a magnitude from 7·9—8·6 of which no proper motions were known.

It is our intention to derive proper motions of these stars, the results of which will be published later.

2. *The observations* have been secured by Mr. D. GAYKEMA and the writer at the telescope, the microscopes being read by Messrs. L. GAYKEMA, B. MEKKING, J. PRINS and in a few cases by Mr. A. VAN HOOF of Leuven. The observations were made entirely differentially, the fundamental stars were chosen from *Boss' P. G. C.* In right ascension the stars were followed during several revolutions with the hand driven micrometer. Three revolutions giving 19 contacts were read. As the stars followed each other within very short intervals of time, the pointings in declination were often made somewhat off the centre of the field.

In the case of the bright stars, two wire gratings have been used in front of the object glass, increasing the magnitude by about $2^{m}\cdot5$ or $4^{m}\cdot9$. On 24 nights the instrument was used in the position in which the clamp is on the West side, whereas on 6 nights the instrument was used in the reversed position.

3. *The instrumental constants* were determined in the usual way. The constant of collimation was derived by reversing the instrument during transits of Polaris as well as by pointings on the meridian marks. No appreciable difference was found between these values; the mean of all determinations has been used. The azimuth of the instrument and the run of the microscopes were determined at regular intervals.

It was necessary to determine the inclination of both horizontal threads. This has been done by pointings on Polaris at different positions in the field, noting the time and reading two microscopes on two divisions each.

By computing the influence of the curvature of the parallel of Polaris, the inclination of the thread used can be determined.

Another method used, consisted in settings on equatorial stars at the extreme ends of the field and reading one microscope on one divison after each pointing, there being no influence of the curvature of the parallel in this case. Each determination will be affected by a rather large error, arising from the errors of two pointings and two readings. But the number of observations can easily be increased in such a way as to get a satisfactory result. The use of equatorial stars has some advantages over that of Polaris; it is less affected by changes in the atmosphere and the reductions are easier.

The results for the inclinations of the threads are given as corrections to the readings of the circle.

In the position clamp West and for stars South of the zenith, the pointings on the upper thread require a correction to the readings of $+ 0''\cdot21$ for every revolution the star is observed past the meridian.

In the same sense, the lower thread requires a correction of $+ 0''\cdot07$.

The results from Polaris and from 25 equatorial stars were exactly the same. From the latter an internal m.e. of $\pm 0''\cdot02$ has been derived.

4. *The reductions and the mean errors.*

a. Right ascensions.

The fundamental stars yielded the errors of the clock. For each observing night a rate of the clock has been computed and applied to the results of the individual stars. The remaining residuals of the fundamental stars have been collected for every star. The mean of the residuals of one star is a correction to the right ascension of that star. After the application of these corrections the errors of the clock and its rate have been recomputed for every night. As no great differences with the original values were found, the latter were adopted for the reduction of the programmestars. From the stars with four or more determinations I derived a mean error of $\pm 0^s\cdot031 \sec \delta$ for these corrections. The average mean error of one correction to the clock before the application of the corrections was: $\pm 0^s\cdot044 \sec \delta$ and after: $\pm 0''\cdot030 \sec \delta$. I have tried to find a correlation between the corrections to the clock and the magnitudes or the declinations of the stars from which it was derived. For this purpose the residuals were first plotted against the magnitudes, taking into account the effect of the gratings used. No appreciable magnitude equation was found for any of the observers.

In order to investigate a possible correlation with the declinations I have plotted the residuals of the $\sec \delta$ from their mean of every night as abscissae, against the residuals of the errors of the clock as ordinates. The results of all nights were combined into one figure. Averaging the residuals for every interval of 0·01 of the abscissae I found a straight line practically coinciding with the line of abscissae. No corrections of these kinds have therefore been applied to the programme stars.

From the residuals of the programme stars of one night a night correction has been computed, which has been applied.

From the residuals of the observations of one star, a mean error has been computed with the formula: $\varepsilon^2 = \frac{[v\,v]}{m - n}$, m being the number of residuals, n the number of stars; I find: $\varepsilon = \pm 0^s\cdot037 \sec \delta$.

b. Declinations.

From the readings of the microscopes, corrected for run, for the errors of division, for refraction (taken from ALBRECHT's tables) and, if necessary, for the inclination of the horizontal thread, combined with the apparent declinations of the fundamental stars, the equatorpoints as well as its rate of change was derived for each night. As usual the rates were mostly positive.

From the residuals a correction to the adopted declinations has been computed in exactly the same way as described for the right ascensions.

The mean error, found from 4 or more determinations of this correction, was $\pm 0''\cdot56$.

After applying these corrections, the mean equatorpoint and its rate have been recomputed, with the same effect as with the right ascensions. The average mean error of one equatorpoint before the application of the correction was $\pm 0''\cdot71$, and after: $\pm 0''\cdot53$. Finally I have tried to find a correlation with the δ, the procedure followed, being essentially the same as with the right ascensions. No correlation has been found.

From the programme stars corrections for every night have been computed, which have been applied.

The mean error of one observation was found to be $\pm 0''\cdot61$.

I want to thank Dr. HINS for his advice in preparing this paper.

5. *The catalogue* contains the B. D. number, the right ascension for 1935·0 and the declination for 1935·0.

The column headed ep. gives the differences of the epoch of observations with 1930. The 7th column gives the number of observations. In the case of two numbers, the first belongs to the right ascensions.

The magnitudes m are taken from the A. G.

Nr.	B.D.	α 1935.0	ep. 1930 +	δ 1935.0	ep. 1930 +	n	m
	°	h m s		° ′ ″			
1	13 561	3 26 34.46	3.9	14 1 49.7	3.9	34	8.5
2	19 550	28 42.18	4.0	19 57 21.7	4.0	3	8.3
3	10 454	28 54.10	4.0	10 31 1.7	4.0	3	8.6
4	8 531	30 51.54	4.0	8 26 43.3	3.9	3	8.6
5	8 537	32 41.08	4.0	8 40 18.2	4.0	3	8.4
6	11 497	33 7.83	4.0	12 15 37.3	4.0	43	8.6
7	6 562	34 16.11	4.0	7 3 33.2	4.0	3	8.4
8	13 577	35 27.29	4.0	13 28 17.4	4.0	43	8.6
9	7 526	35 30.12	4.0	7 23 24.3	4.0	3	8.5
10	15 514	35 46.38	4.0	15 38 30.4	3.9	32	8.6
11	8 544	35 54.94	4.1	8 51 16.0	4.1	23	8.6
12	8 546	36 0.86	4.1	9 4 55.8	4.1	42	8.6
13	7 529	36 45.54	4.0	7 52 18.3	4.0	3	8.6
14	6 565	36 51.62	4.0	6 20 12.4	4.0	3	8.6
15	6 568	37 15.70	4.0	6 30 33.8	3.9	3	8.5
16	6 574	39 58.54	4.0	7 7 15.2	3.9	2	8.5
17	14 604	40 21.70	4.0	14 24 57.7	4.0	3	8.3
18	8 565	42 52.74	4.0	8 28 43.2	3.9	3	8.6
19	9 486	43 24.32	4.0	9 19 46.9	4.0	3	8.4
20	9 492	44 36.51	4.0	9 58 22.6	3.9	43	8.5
21	15 537	46 32.48	4.0	15 48 13.0	3.9	3	8.6
22	8 575	46 28.88	4.0	9 4 43.2	4.0	3	8.5
23	6 590	46 40.48	4.0	6 25 29.0	4.0	43	8.4
24	8 581	47 15.45	4.1	8 15 8.4	4.1	43	8.4
25	9 502	47 47.94	4.0	9 51 36.9	3.9	3	8.6
26	15 544	50 7.66	3.9	15 53 49.7	3.9	34	8.6
27	9 512	52 4.91	3.9	9 44 23.8	3.9	34	8.5
28	6 605	53 7.36	4.0	6 31 23.7	3.9	34	8.1
29	6 607	54 9.41	4.0	6 41 14.0	4.0	3	8.4
30	7 575	54 45.76	4.0	7 25 10.1	3.9	3	7.9
31	7 577	54 58.84	4.0	8 10 58.7	4.5	35	8.6
32	11 551	55 51.87	4.0	11 35 26.1	4.0	3	8.6
33	13 625	56 37.61	4.0	14 7 23.0	4.4	35	8.3
34	8 611	57 5.26	4.0	8 44 13.7	4.0	3	8.6
35	10 520	57 16.45	4.1	10 36 52.8	4.1	43	8.6
36	8 615	57 37.11	3.9	8 16 41.4	4.0	23	8.6
37	6 626	4 0 26.83	3.9	6 23 3.0	4.4	35	8.6
38	8 631	2 5.32	3.9	9 1 37.8	4.2	35	8.4
39	14 656	3 53.99	4.0	14 48 49.5	4.2	34	8.6
40	16 559	4 12.35	4.0	16 21 28.9	4.0	3	8.2
41	5 593	4 48.79	4.0	6 2 20.2	4.0	3	8.6
42	6 636	5 18.37	4.0	6 36 36.5	4.2	34	8.6
43	10 542	6 16.06	4.0	10 37 51.8	4.0	3	8.5
44	18 593	6 26.42	4.0	18 18 18.4	4.0	3	8.5
45	6 642	7 34.72	3.9	7 7 14.4	3.9	32	8.6
46	10 551	10 59.50	4.0	10 32 16.8	3.9	3	8.1
47	8 654	11 31.83	4.0	8 20 31.2	4.0	3	8.6
48	7 620	11 33.35	4.0	7 15 7.7	4.0	3	8.4
49	8 656	11 47.56	4.0	8 16 1.6	4.0	3	8.5
50	7 624	12 33.68	3.9	7 24 0.9	4.2	34	8.5
51	10 555	14 3.11	3.9	10 52 55.0	4.2	34	8.6
52	5 622	15 7.01	4.0	5 33 35.8	4.2	34	8.6
53	8 667	15 47.48	3.9	8 22 32.6	3.9	3	8.6
54	10 560	17 7.73	3.9	11 4 53.0	3.9	34	8.6
55	6 667	18 4.90	4.0	6 49 15.9	4.0	3	8.3
56	6 669	18 55.37	4.0	6 22 39.6	4.2	34	8.6
57	2 695	19 3.54	4.0	2 54 38.0	4.0	3	8.6
58	4 683	19 36.74	4.0	4 27 48.6	4.0	3	8.6
59	6 676	20 17.31	3.9	6 13 44.1	3.9	3	8.5
60	5 648	22 1.89	4.0	5 25 22.3	4.2	35	8.6
61	9 575	22 8.40	4.0	9 34 40.3	4.0	3	8.6
62	6 687	24 7.11	4.0	6 28 3.7	4.0	34	8.5
63	6 688	24 51.50	4.0	6 16 40.6	4.0	3	8.6
64	17 738	26 51.44	4.0	17 15 28.4	3.9	3	8.5
65	8 702	26 50.61	4.0	8 47 27.7	4.0	3	8.6
66	7 656	27 38.02	4.0	7 11 34.9	4.0	3	8.5
67	7 657	27 41.02	4.1	8 0 14.3	4.1	3	8.6
68	7 658	27 46.05	4.1	7 18 49.4	4.1	3	8.3
69	13 691	27 49.37	4.0	13 45 52.4	4.0	3	8.4
70	7 667	30 51.01	4.0	8 3 8.2	4.0	3	8.5
71	9 602	30 55.75	4.0	9 52 59.0	4.0	43	8.3
72	18 652	4 31 6.16	4.1	18 16 45.5	4.1	43	8.5
73	7 668	31 19.47	4.0	7 18 19.0	4.0	3	8.6
74	18 658	32 40.65	4.0	18 32 3.0	4.0	34	8.4
75	5 690	34 26.01	4.0	5 38 27.2	4.0	3	8.6
76	7 678	35 4.88	4.0	8 21 48.3	3.9	3	8.4
77	6 728	35 9.36	4.0	6 44 47.9	4.0	43	8.4
78	11 636	35 19.10	4.1	11 52 9.1	4.1	34	8.4
79	9 620	35 26.49	4.1	9 30 40.7	4.0	43	8.5
80	7 693	38 11.41	4.0	7 42 10.0	4.0	34	8.6
81	5 710	38 55.77	4.0	5 51 34.7	4.0	3	8.5
82	7 696	39 23.41	4.0	7 22 13.2	4.0	3	8.5
83	7 697	39 24.75	4.1	7 26 52.1	4.1	3	8.5
84	5 721	41 35.79	4.0	5 45 11.6	4.0	34	8.4
85	7 711	41 54.62	3.9	7 28 15.2	4.0	23	8.6
86	7 714	42 57.86	4.0	7 13 39.3	4.0	3	8.5
87	8 760	43 31.74	4.1	8 56 41.0	4.0	3	8.6
88	8 766	44 40.59	4.0	8 50 45.8	4.0	3	8.6
89	5 734	45 7.61	4.0	5 29 32.6	3.9	3	8.5
90	7 730	46 28.33	4.0	7 34 4.7	4.0	34	8.5
91	8 774	46 33.09	4.1	8 24 33.6	4.0	52	8.6
92	7 733	47 19.12	4.0	7 53 7.4	3.9	3	8.2
93	7 734	47 19.56	4.0	7 34 58.2	4.0	3	8.6
94	7 737	47 31.59	4.1	7 35 16.3	4.1	32	8.6
95	7 741	47 50.17	4.1	7 19 42.7	4.1	3	8.5
96	5 751	48 16.71	4.0	6 3 57.6	4.0	3	8.4
97	19 800	48 45.10	4.1	19 59 48.8	4.1	3	8.6
98	9 671	48 47.60	4.0	9 13 47.2	4.0	3	8.4
99	8 789	49 10.10	4.0	8 15 27.2	3.9	3	8.6
100	6 776	50 32.64	4.0	6 21 21.9	4.0	3	8.6
101	12 672	50 50.44	4.0	12 20 48.6	4.2	4	8.6
102	11 671	51 2.46	4.0	11 49 52.5	4.0	3	8.1
103	5 762	51 2.45	4.1	5 49 52.1	4.1	3	8.5
104	11 680	52 26.76	4.1	11 8 25.6	4.1	3	8.6
105	7 756	52 23.41	4.1	7 18 6.1	4.1	3	8.6
106	8 803	52 35.93	4.0	9 2 5.5	4.0	3	8.6
107	5 771	52 45.93	4.0	5 55 3.1	4.2	34	8.6
108	3 714	53 0.62	4.1	3 51 7.1	4.1	23	8.6
109	8 807	53 2.32	4.1	8 41 41.7	4.1	24	8.6
110	8 809	53 4.68	4.0	8 52 26.7	4.0	3	8.6
111	8 811	53 11.09	4.1	8 43 14.7	4.1	3	8.1
112	9 691	53 30.62	4.0	9 39 41.8	4.3	34	8.6
113	8 817	54 34.25	4.0	8 16 38.8	3.9	3	8.5
114	12 688	55 19.17	4.0	12 19 2.4	4.2	35	8.5
115	7 770	56 47.95	4.0	7 27 20.0	4.2	35	8.5
116	6 802	57 25.94	4.0	6 29 23.0	4.3	34	8.6
117	7 774	57 41.51	4.1	7 46 29.5	4.0	23	8.5
118	6 806	57 43.35	4.1	6 9 2.0	4.1	43	8.5
119	6 808	58 13.17	4.0	7 3 20.8	4.0	3	8.2
120	6 809	58 14.45	4.1	6 39 28.9	4.1	3	8.4
121	14 815	58 26.74	4.0	14 39 25.0	3.9	3	8.4
122	8 839	58 32.62	4.1	8 50 5.6	4.1	3	8.6
123	5 791	59 15.28	4.0	5 44 18.2	4.0	3	8.6
124	6 815	59 28.64	4.0	6 35 24.9	4.0	3	8.5
125	9 718	59 41.72	4.1	9 56 18.5	4.4	34	8.6
126	6 818	59 42.43	4.1	7 7 10.6	4.1	3	8.5
127	8 843	5 0 28.53	4.0	8 55 37.7	3.9	3	8.6
128	8 847	0 37.02	4.1	9 3 6.8	4.1	3	8.6
129	7 783	0 38.47	4.0	7 16 20.1	4.3	34	8.6
130	7 785	0 49.48	4.0	8 6 54.9	4.0	3	8.4
131	6 827	1 37.08	4.1	6 7 28.4	4.1	3	8.6
132	9 725	1 40.81	4.1	9 8 31.7	4.1	4	8.6
133	5 805	2 31.83	4.0	5 33 57.5	4.0	45	8.6
134	10 711	2 44.63	4.0	10 12 2.4	4.0	3	8.6
135	2 860	2 59.22	4.1	2 35 12.3	4.3	34	8.0
136	7 802	3 14.05	4.1	8 0 37.8	4.1	3	8.6
137	7 806	3 30.62	4.1	7 37 33.2	4.1	2	8.6
138	14 836	4 24.70	4.0	14 26 55.0	3.9	3	8.3
139	8 867	4 22.82	4.0	8 19 8.3	4.0	3	8.6
140	9 738	4 29.10	4.0	9 17 28.0	4.0	3	8.6
141	6 845	4 30.41	4.1	6 22 16.3	4.1	3	8.5
142	8 870	4 45.15	4.1	8 43 1.0	4.1	2	8.6

Nr.	B.D.	α 1935·0	ep. 1930 +	δ 1935·0	ep. 1930 +	n	m
	°	h m s		° ′ ″			
143	13 812	5 4 50·39	4·1	13 53 41·2	4·2	23	8·6
144	7 814	4 58·48	4·1	7 39 56·5	4·1	32	8·6
145	10 721	5 24·24	4·1	10 19 59·0	4·1	3	8·5
146	11 727	5 35·64	4·1	11 24 47·0	4·1	3	8·5
147	8 873	5 34·08	4·0	9 4 21·9	4·0	3	8·6
148	5 821	6 14·14	4·0	5 42 4·2	3·9	3	8·5
149	9 749	6 24·36	4·0	9 9 4·8	4·3	4	8·6
150	7 827	7 21·46	4·0	7 30 54·7	4·0	3	8·8
151	6 860	7 47·94	4·1	6 55 35·2	4·1	3	8·6
152	10 728	8 18·53	4·0	10 21 51·6	3·9	3	8·6
153	5 833	8 21·22	4·0	5 55 7·1	4·3	34	8·6
154	11 742	8 43·48	4·1	11 17 6·9	4·1	3	8·6
155	8 886	8 45·20	4·1	9 2 47·6	4·0	3	8·6
156	10 731	8 59·31	4·0	10 19 51·1	4·0	43	8·4
157	8 889	9 28·21	4·1	8 23 28·5	4·1	3	8·6
158	5 847	10 3·34	4·0	5 41 35·3	4·0	34	8·6
159	9 768	10 19·81	4·0	9 37 40·3	4·0	3	8·4
160	11 748	10 42·26	4·1	11 24 49·0	4·1	2	8·6
161	7 846	11 28·81	4·0	7 29 41·1	4·0	3	8·6
162	6 873	11 29·39	4·1	6 11 28·4	4·3	34	8·6
163	8 900	11 38·17	4·0	8 21 35·8	4·0	3	8·6
164	19 880	11 59·62	3·9	19 52 54·2	3·9	23	7·8
165	5 859	12 3·00	4·1	5 41 14·8	4·1	23	8·6
166	8 903	12 3·08	4·1	8 22 8·2	4·1	31	8·6
167	8 904	12 6·39	4·1	8 45 3·9	4·1	2	8·3
168	19 881	12 6·90	4·1	20 3 14·0	4·1	2	8·0
169	10 744	12 13·60	4·1	11 2 36·4	4·1	23	8·6
170	12 756	12 37·02	4·1	12 7 52·9	4·1	3	8·5
171	19 885	12 57·90	4·0	19 13 51·9	4·0	3	8·6
172	19 884	12 58·32	4·1	20 2 51·7	4·4	34	7·7
173	12 760	13 26·67	4·0	12 29 43·1	4·0	3	8·1
174	13 848	13 34·80	4·1	13 29 54·2	4·1	3	7·7
175	6 885	13 38·69	4·0	6 13 53·9	4·2	34	8·6
176	13 850	14 5·52	4·1	13 34 14·2	4·1	3	8·4
177	6 887	14 8·53	4·1	7 4 53·0	4·1	3	8·6
178	7 855	14 10·26	4·1	7 17 12·1	4·1	3	7·9
179	9 792	14 24·08	4·0	10 2 32·5	4·0	3	8·6
180	9 796	15 10·65	4·0	9 57 23·5	4·0	3	8·5
181	8 920	15 33·31	4·0	8 45 8·8	3·9	3	8·6
182	9 806	17 18·86	4·0	9 39 35·3	4·0	3	7·9
183	11 774	17 19·16	4·0	11 37 55·2	4·3	34	8·3
184	12 779	17 27·76	4·0	12 22 0·8	3·9	3	8·4
185	19 906	18 23·94	4·0	19 46 18·9	4·0	3	8·4
186	7 875	18 36·39	4·1	7 56 34·6	4·0	3	8·5
187	16 762	18 48·91	4·1	16 44 31·7	4·1	3	8·4
188	9 811	19 6·07	4·0	9 10 3·3	4·3	34	8·6
189	16 766	19 56·87	4·0	16 26 14·0	3·9	3	8·2
190	12 788	20 4·64	4·0	12 42 57·4	4·0	3	8·4
191	16 767	20 24·17	4·1	16 30 12·1	4·3	34	8·0
192	15 807	20 50·51	4·1	15 54 59·3	4·1	3	8·5
193	9 821	20 59·70	4·1	9 6 43·8	4·1	43	8·6
194	8 950	21 8·26	4·1	8 26 38·3	4·3	25	8·6
195	11 795	21 11·09	4·1	11 38 55·8	4·1	23	8·3
196	10 773	21 23·27	4·0	11 4 11·8	4·0	3	8·2
197	9 823	21 26·10	4·0	9 55 23·8	4·0	3	8·3
198	9 830	22 8·43	4·0	9 10 19·4	4·2	34	8·5
199	15 814	22 18·91	4·1	15 36 55·6	4·1	3	7·5
200	9 831	22 23·47	4·1	9 24 12·3	4·1	3	8·6
201	8 959	23 0·00	4·0	8 11 58·3	4·0	3	8·6
202	13 893	23 7·81	4·0	13 31 37·4	4·0	3	7·9
203	18 853	23 33·34	4·1	18 39 48·3	4·1	3	8·6
204	11 807	23 33·60	4·1	11 28 16·2	4·1	34	8·4
205	13 896	23 40·52	4·1	13 58 28·6	4·1	12	8·5
206	14 914	23 58·63	4·0	14 4 54·3	3·9	3	8·6
207	13 907	25 19·93	4·0	13 5 50·7	4·0	3	8·5
208	13 908	25 47·10	4·0	13 22 35·7	3·9	3	8·4
209	18 867	27 17·73	4·0	18 30 44·2	4·0	3	8·4
210	12 810	27 43·84	4·0	12 30 38·2	3·9	3	8·6
211	14 933	27 48·20	4·0	14 35 46·9	4·0	3	8·2
212	14 948	30 18·95	3·9	14 26 47·0	3·9	34	8·6
213	15 866	33 5·34	3·9	15 34 42·9	4·2	35	8·0

A comparison of the above positions with the A. G. Catalogue revealed the following possible members of the group of the Hyades

probable:

nr.	μ_α	μ_δ
	s	″
46	+ 0·0087	— 0·011
69	+ 67	— 22
170	+ 47	— 16

doubtful:

nr.	μ_α	μ_δ
101	+ 73	— 53
158	+ 36	— 38

Determination of the radial velocity of these stars would give a valuable contribution to their classification.