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## COMMUNICATION FROM THE OBSERVATORY AT LEIDEN.

### Provisional ephemerides of 25 variable stars in the constellations Norma and Ara, by *W. E. Kruytbosch*.

The number of plates with the centre at  $16^{\text{h}}4^{\text{m}}, -54^{\circ}45'$  (1875) covering  $10^{\circ} \times 10^{\circ}$  and taken by Dr. H. VAN GENT with the Franklin-Adams instrument at Johannesburg amounts to 304 at present.

Selected pairs of plates of this region were compared in the blinkmicroscope. Of the 25 variable stars forming the subject of the present note 7 were found in this way by Dr. H. VAN GENT at Johannesburg, 10 by P. TH. OOSTERHOFF here in Leiden and 8 by the writer also in Leiden.

The total number of observations used here is 6285; the average number for each star is 251. The extreme dates of the plates used are J. D. 2425381 and 6476. The normal exposure-time of a plate is 30 min.

The 25 stars are listed in Table 1. Four of these occur in the *C. P. D.*, viz: *a, c, o* and *p* with the numbers  $-51^{\circ}8479$  ( $10^{\text{m}}\cdot 0$ ),  $-54^{\circ}6718$  ( $9^{\text{m}}\cdot 8$ ),  $-53^{\circ}7517$  ( $10^{\text{m}}\cdot 2$ ) and  $-53^{\circ}7685$  ( $9^{\text{m}}\cdot 6$ ) respectively.

For 24 periods a mean error has been given in Table 1. Owing to the small number of epochs these mean errors are in several cases rather uncertain.

The phases have in all cases been computed according to the formula:

$$\text{phase} = \text{recipr. period} \times (\text{J. D. hel. M. astr. T. Grw.} \\ - 2420000)$$

The reciprocal periods actually used for this computation are given in Table 1.

For eclipsing variables the phase of minimum, for variables of  $\delta$  Cep or RR Lyr type the phase of maximum has been given. The phases of maximum derived from a sine curve are marked with an asterisk in Table 1 and explained in the remarks below.

The normal epochs given in J. D. have been chosen near the mean date of the observations.

An attempt was made to estimate the magnitudes of the comparison stars by comparison with plates of the Crux region containing the Selected Area No. 193. The magnitudes thus found are given in Table 3. They are naturally rather uncertain.

The dates of the two plates, by the comparison of which the variability of the star was found in the blink microscope, are given in Table 2 together with the interval in days and the difference in steps between the two images of the variable as derived from the mean lightcurve. The last column of Table 2 contains the abbreviation of the name of the observer by whom the different variables were found: vG = Dr. H. VAN GENT, O = P. TH. OOSTERHOFF, K = the writer.

The surroundings of each variable with indication of the comparison stars are shown on the diagrams on page 240. The size of each square is indicated in the right hand bottom corner. The brightness assigned to each comparison star is given in Table 3.

For 15 variables the epochs of minimum and for 9 the epochs of maximum used for computing the period according to least squares are indicated in Table 4 together with the counting of epochs and the residuals *O-C*. For one variable (*m*) the number of epochs was too small to get reliable results by a least square solution.

For each variable the observations were arranged according to phase and divided into groups as indicated in Table 5, where mean values of phase and brightness are given for each group. The mean values are shown on the diagrams on the pages 239 and 240.

On all these diagrams vertical lines have been drawn for each tenth of the period except in the case of variable *d* as explained below. The phases have on the diagrams in case of eclipsing variables been counted from minimum and the brightness in steps from maximum light. For the eclipsing variable *m* a diagram of the complete period is given besides a diagram of the individual observations at or near the minimum.

The individual objects give rise to the following remarks:

*d*. This Algol star shows two different minima marked with vertical lines on the diagram at  $\text{ph}\cdot 342$  and  $\text{ph}\cdot 919$  respectively. The minima are practically

of the same range, but show a somewhat different shape, the former being wider than the latter.

*e.* The mean light curve has an indication of a secondary minimum at  $^{\text{ph}}76$ .

*g.* A secondary minimum is indicated at  $^{\text{ph}}315$ .

*h.* A first trial to determine the period of this star giving insufficient results, the variable was estimated independently by Mr. OOSTERHOFF. The two series of observations are in good agreement. The period was deduced from the minima. A sine curve was determined by least squares from the mean values of phase and brightness for the two series of observations. The results are given in Table 6. The mean value of the phase of the maximum is  $^{\text{P}}180 \pm ^{\text{P}}015$  (m. e.) corresponding to the mean epoch of maximum J. D.  $2425705^{\text{d}}96 \pm ^{\text{d}}03$  (m. e.)

TABLE 6.

P	O O—C		K O—C	
	s	s	s	s
19	'0428	2'39 — '18	1'53 — '10	
19	'1375	1'74 — '38	1'47 + '01	
19	'2038	2'31 + '28	1'45 — '02	
19	'2415	2'42 + '33	1'53 0	
19	'2862	2'19 — '05	1'75 + '12	
19	'3274	2'37 — '08	1'56 — '20	
19	'4036	3'16 + '19	2'12 + '07	
19	'5366	3'65 — '34	2'69 + '15	
19	'6386	4'29 — '20	2'74 + '02	
19	'6955	4'71 + '14	2'51 — '21	
19	'7478	4'15 — '35	2'35 — '30	
19	'7924	5'06 + '72	2'85 + '31	
19	'8437	4'56 + '50	2'52 + '15	
19	'8793	3'46 — '36	2'13 — '11	
19	'9525	3'04 — '20	2'04 + '10	

$$S_O = + 3.305 - 1.196 \sin 2\pi P - .426 \cos 2\pi P, \text{ max. at } ^{\text{P}}196 \\ \pm .107 \quad \pm .107 \quad \pm .017 \text{ (m. e.)}$$

$$S_K = + 2.093 - .550 \sin 2\pi P - .324 \cos 2\pi P, \text{ max. at } ^{\text{P}}165 \\ \pm .064 \quad \pm .064 \quad \pm .016 \text{ (m. e.)}$$

*m.* This Algol star with a period of  $21^{\text{d}}249$  has a broad, constant minimum about  $^{\text{P}}05$  wide. The descending branch of the lightcurve is steep, while no observations on the ascending branch are available. The individual observations of the minimum are given in Table 7 and on the diagram on page 240, where the two open circles represent the mean values of 5 observations at maximum. There is some indication of the brightness not being constant outside the primary minimum. More observations, especially on the ascending branch, are much wanted in order to get a complete lightcurve of this interesting star.

TABLE 7.

J. D. hel.	Phase	Brightness
2420000		s
5797'261	'831	4'5
97'283	'832	7'1
97'442	'839	7'2
97'464	'840	8'0
5415'244	'852	11'1
15'295	'855	11'4
15'318	'856	11'6
15'450	'862	11'7
15'496	'864	12'4
15'543	'866	11'2
6010'578	'870	13'0
10'600	'871	11'5
5798'258	'878	13'0
13'506	'889	13'2
13'528	'890	11'5
5437'448	'897	11'5
16'222	'898	11'3
16'244	'899	11'4
16'266	'900	13'2
16'288	'901	11'8

*w.* A secondary minimum is found at  $^{\text{ph}}826$ .

*z.* The lightcurve of this  $\delta$  Cep variable has approximately the shape of a sinusoid, the maximum of which has been deduced by least squares. The residuals are given in Table 5. The formula is:

$$S = + 3.089 - .324 \sin 2\pi P + 1.452 \cos 2\pi P, \text{ max. at } ^{\text{P}}4652 \\ \pm .089 \quad \pm .089 \quad \pm ^{\text{P}}0094 \text{ (m. e.)}$$

The following stars have meanwhile been published from Harvard as being variable (*H. B.* 884):

*b* = YY Nor, Algol type, Harvard period  $1^{\text{d}}25333$

A possible relation between the periods of Harvard (*H*) and Leiden (*L*) is  $1/2H + 1/L = 1$ .

*l* = BY Nor, RR Lyrae type, period unknown,

*m* = CE Nor, Algol type, period unknown.

*n* = CP Nor, Algol type, Harvard period  $3^{\text{d}}90$ .

*u* = VZ Arae, RR Lyrae type, Harvard period  $^{\text{d}}38443$ .

The reciprocals of the Harvard and Leiden periods differ by one unit.

I want to express my thanks to Professor HERTZSPRUNG for the help and advice he gave me in preparing this paper.

TABLE 1.

	$\alpha$ (1875)		number of plates	period	m. e.	reciprocal period	phase of epoch	epoch 2420000+	fraction of period occupied by minimum	range			max. min.				
	h	m s								°	'	s	s	s	m	m	
<i>a</i>	15	34	58.7	-51	6.5	230	2.20889	± 0.0004	.452717	.775	5824.333	.15	7.0	19.0	12.0	10.6	12.8
<i>b</i>	15	38	3	-57	1.4	246	1.694903	± 0.00047	.590004	.982	5649.084	.25	2.1	7.6	5.5	12.9	16.2
<i>c</i>	15	42	13.3	-54	59.0	290	.942331	± 0.0005	1.061198	.480	5726.057	.56	-1.1	2.0	3.1	10.4	11.0
<i>d</i>	15	45	2.6	-54	9.2	284	5.70386	± 0.00785	.17532	.342 .919	5772.170	.15 .14	2.2	5.0	2.8	12.7	13.3
<i>e</i>	15	46	53.2	-53	43.6	287	5.9569	± 0.0042	.167873	.266	5738.060	.31	2.4	6.9	4.5	13.4	15.6
<i>f</i>	15	50	56.7	-57	.9	272	2.42711	± 0.0013	.412012	.139	5611.824	.15	.4	2.1	1.7	13.8	14.6
<i>g</i>	15	53	24.2	-51	1.2	258	.830755	± 0.00015	1.203724	.797	5637.340	.21	1.4	3.6	2.2	13.6	14.8
<i>h</i>	15	57	48.9	-58	36.7	285	1.9600	± 0.004	.5102	.180*	5705.965	.15	1.5	2.9	1.4	13.0	13.5
<i>i</i>	15	58	42.9	-57	53.1	280	1.454044	± 0.00057	.687737	.291	5706.180	.22	.5	2.1	1.6	13.8	15.9
<i>k</i>	15	59	24.8	-53	2.3	155	2.07116	± 0.0016	.48282	.122	5577.901	.17	2.2	5.3	3.1	13.9	14.8
<i>l</i>	15	59	49.2	-59	22.5	257	.660700	± 0.00043	1.513546	.880	5720.680	.15	1.5	3.1	1.6	13.7	14.7
<i>m</i>	16	0	37.8	-59	42.1	282	21.248566		.047062	.512	5501.050	.10	1.5	12.0	10.5	11.0	13.9
<i>n</i>	16	3	42.6	-57	58.8	282	2.46628	± 0.0009	.40547	.866	5753.490	.29	.9	4.1	3.2	13.2	15.1
<i>o</i>	16	5	13.1	-53	1.0	293	3.4535	± 0.006	.286	.135	5756.680		2.0	5.1	3.1	11.5	12.0
<i>p</i>	16	8	52.2	-53	41.1	177	2.97183	± 0.0016	.3365	.953	5560.039	.11	2.7	8.0	5.3	10.9	11.8
<i>q</i>	16	13	54.8	-52	49.0	198	.624999	± 0.00020	1.600003	.021	5708.127		1.7	3.2	1.5	14.4	15.3
<i>r</i>	16	14	3.0	-53	12.5	252	.134391	± 0.00001	7.440974	.466	5694.210		1.3	3.7	2.4	14.0	15.0
<i>s</i>	16	16	47.5	-52	56.1	257	3.093067	± 0.00059	.323304	.675	5919.119	.16	1.4	4.1	2.7	14.2	15.2
<i>t</i>	16	21	47.9	-55	16.3	291	35.734	± 0.0001	.120	.028	5827.860		1.1	3.5	2.4	12.5	13.8
<i>u</i>	16	26	55.0	-58	8.1	177	.62563	± 0.0001	1.59838	.120	5678.320		.5	3.8	3.3	13.2	14.8
<i>v</i>	16	27	4.0	-51	9.4	250	1.79942	± 0.0004	.55573	.347	5771.410	.23	.3	3.1	2.8	12.7	13.9
<i>w</i>	16	30	8.1	-56	17.2	275	8.5207	± 0.0033	.117361	.345	5822.590	.35	2.5	8.1	5.6	13.0	14.9
<i>x</i>	16	31	30.1	-57	4.9	244	5.9594	± 0.0024	.1678	.384	5765.000		.8	3.4	2.6	12.1	14.3
<i>y</i>	16	33	2.9	-56	12.9	218	16.4370	± 0.0074	.0608	.899	5623.330	.23	1.9	5.7	3.8	13.1	14.8
<i>z</i>	16	35	45.4	-55	32.8	245	16.016	± 0.022	.0624	.465*	5752.680		1.7	4.5	2.8	12.4	15.4

TABLE 2.

	type	variable found		$\Delta d$	$\Delta s$	found by
		bright	faint			
<i>a</i>	Algol	d 2425739.488	d 2425720.541	19	12.6	v G
<i>b</i>	Algol	5720.520	5713.506	7	5.1	K
<i>c</i>	Algol	6118.479	6129.421	11	1.6	O
<i>d</i>	Algol	5713.506	5720.520	7	2.8	K
<i>e</i>	Algol	6129.399	6125.461	4	4.5	K
<i>f</i>	Algol	5417.439	5420.376	3	1.1	K
<i>g</i>	Algol	5720.541	5739.488	19	1.7	v G
<i>h</i>	$\delta$ Cep	6129.421	6118.479	11	1.3	O
<i>i</i>	Algol	5390.439	5393.493	3	2.0	K
<i>k</i>	Algol	5720.541	5739.488	19	4.0 <sup>1)</sup>	v G
<i>l</i>	RR Lyr	6118.479	6129.421	11	1.8	v G
<i>m</i>	Algol	5720.520	5713.506	7	10.4	K
<i>n</i>	Algol	6129.421	6118.479	11	3.5	O
<i>o</i>	$\delta$ Cep	5739.488	5720.541	19	3.1	v G
<i>p</i>	Algol	5739.488	5720.541	19	5.3	v G
<i>q</i>	$\delta$ Cep	6129.421	6118.479	11	2.0	O
<i>r</i>	WU Ma	6129.421	6118.479	11	2.5	O
<i>s</i>	Algol	6118.479	6129.421	11	3.3	O
<i>t</i>	$\delta$ Cep	5720.520	5713.506	7	2.8	K
<i>u</i>	RR Lyr	6129.421	6118.479	11	3.1	O
<i>v</i>	Algol	6118.479	6129.421	11	5.4	O
<i>w</i>	Algol	6118.479	6129.421	11	6.1	O
<i>x</i>	$\delta$ Cep	5390.439	5393.493	3	1.3	K
<i>y</i>	Algol	5720.541	5739.488	19	3.5	v G
<i>z</i>	$\delta$ Cep	6129.421	6118.479	11	2.5	O

1) The variable being invisible on plate 5102 of J. D. 2425739.488, the brightness of this minimum is assumed to equal that of the mean light-curve, viz: 5<sup>s</sup>.3.









