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G 292.0 + 1.8: evidence for a pseudo- Crab stage in SNR evolution

R. Braun 1, *, W. M. Goss 2, J. L. Caswell 3, and R. S. Roger 4

1 Sterrewacht Leiden, Postbus 9513, 2300 RA Leiden, The Netherlands
2 Kapteyn Astronomical Institute, Postbus 800, 9700 AV Groningen, The Netherlands
3 Division of Radiophysics, CSIRO, PO Box 76, Epping, NSW 2121, Australia
4 Dominion Radio Astrophysical Observatory, Penticton, BC, V2A 6K3, Canada

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Summary. New observations of the young type II supernova remnant (SNR) G292.0 + 1.8 have been made at 843 MHz with 0’’8 resolution using the Molonglo Observatory Synthesis Telescope (MOST) and at 12, 25, 60, and 100 µm with IRAS. All of the available structural and kinematic data suggests that the southwest portion of the SNR is encountering the periphery of a cloud (where n = 0.4 cm⁻³) while the northeast portion is expanding into low density intercloud gas (n < 0.003 cm⁻³). A relatively small mass (< 1.0 M☉) of high velocity ejecta seems responsible for the blast wave. A significant mass of slower ejecta is now encountering the reverse shock in the southwest, giving rise to a center-filled morphology. Such a pseudo- Crab phase may be a common occurrence in type II SNR of intermediate age.

Key words: supernova remnants – G292.0 + 1.8 – dust

1. Introduction

Early radio surveys (Milne, 1969; Shaver and Goss, 1970) led to the discovery of the southern galactic supernova remnant G292.0 + 1.8. Subsequent radio observations (Lockhart et al., 1977) revealed a peculiar, apparently center-filled morphology which led to the suggestion that the SNR may be Crab-like (although the radio spectral index, α = 0.37, is rather steep for objects of this class). Optical observations (Goss et al., 1979; Murdin and Clark, 1979; Van den Bergh, 1979; Braun et al., 1983) have revealed a system of high velocity oxygen rich knots reminiscent of those found in Cas A (Chevalier and Kirshner, 1978), suggesting a relatively young SNR and a massive progenitor. Observations of strong emission lines in the thermal X-ray spectrum (Clark et al., 1980) also suggest a massive progenitor, while ruling out the Crab-like classification. The X-ray image (Touhy et al., 1982) indicates an over-all morphology which is similar to that seen in the radio; bright central knots and an extended plateau.

While a large body of observational data exists for the SNR, the origin of the various emission processes has not yet been satisfactorily accounted for. In this paper we present new radio and infrared observations which we analyze in the context of all available observational constraints and in the framework of numerical simulations of pre-adiabatic SNR evolution. The observations and data reduction techniques are described in Sect. 2. In Sect. 3 we begin by considering the distance, environment and geometry of G292.0 + 1.8 and then go on to discuss the brightness distribution in the various wavelength bands and the physical parameters we derive. Our conclusions are summarized in Sect. 4.

2. Observations and data reduction

2.1. Radio data

The radio observations were carried out using the Molonglo Observatory Synthesis Telescope (MOST, Durbin et al., 1984) of the University of Sydney on 27 August, 1984. The flux density calibration was carried out using the point source 1934–63 with an assumed flux density of 13.5 Jy at 843 MHz. After an observation of 12h with a field of view of 46’’ × 54’’ (α × δ) the rms noise is 1.3 mJy and the synthesized beam 43’’ × 50’’ (α × δ). The measured dynamic range is somewhat better than 100:1. The map was processed by “CLEAN” ing the field using an observation of 1934–63 extending over 12 hours as the point source response. The resulting map is shown in Fig. 1.

2.2. Infrared data

A one degree square field centered on G292.0 + 1.8 was extracted from the IRAS Skyflux maps (HCON3, IRAS, 1985) in the 12, 25, 60, and 100 µm bands for further processing. The gridded point source response is similar in the four bands, with FWHM ~ 6’. Straightforward analysis of the data is hampered by a number of sources of confusion. Besides star images which affect localized areas at 12 and 25 µm, the ubiquitous zodiacal light emission introduces a varying background across the field and worse still, cool dust emission with structure on all observable scale sizes is present in most directions. The problem of background confusion was addressed by using the four maps to carry out a spectral decomposition by least squares into three component maps, on the basis of their distinct spectral character (as discussed by Braun, 1986). The three components which were so extracted were; 1) cool dust, peaking longward of 100 µm, 2) zodiacal light, peaking between 12 and 25 µm and 3) the shock-heated dust emission of the SNR, peaking between 25 and 60 µm. Initial

Send offprint requests to: R. Braun

* Now at National Radio Astronomy Observatory, PO Box O, Socorro, NM 87801, USA

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Fig. 1. Radio brightness distribution. The 843 MHz MOST map (HPBW = 43' × 50') with contours at 20, 40, 60, 100, 150, 200, 300, 500, 700, and 900 mJy/beam. Note the central ridge and extended plateau. The nominal expansion center is indicated by the cross.

Fig. 2. G292.0+1.8 and its environment. The grey scale represents cool dust emission, (with transitions occurring at 49, 59, 61, 63, 65, 68, and 74 MJy/sr at 100 μm) while the contours represent a warmer component (with linear contour levels at 3 to 10 MJy/sr in steps of 1 MJy/sr at 60 μm) as derived from the spectral decomposition (see Sect. 2.2) of IRAS data at 12, 25, 60, and 100 μm. A cool arm of dust and gas extends up from the galactic plane to the SNR at center. No attempt was made to model the emission of the H II region to the north. The nominal expansion center is indicated by the cross.

Fig. 3. Distribution of shock-heated dust. The background corrected (see Sect. 2.2) 60 μm IRAS image (HPBW ~ 6'). Contours are in units of nominal surface brightness. The linear contour levels and grey scale transitions are at 2.25 to 7.5 MJy/sr in steps of 0.75 MJy/sr. Note the overall coincidence with the radio plateau but displacement of peak to the south-west. The nominal expansion center is indicated by the cross.

Table 1. Infrared spectral parameters

<table>
<thead>
<tr>
<th>Component</th>
<th>λpeak [μm]</th>
<th>λJ_λmax [W m⁻² sr⁻¹]</th>
<th>Angular radius [arcmin]</th>
</tr>
</thead>
<tbody>
<tr>
<td>(12 μm)</td>
<td>&lt; 0.42</td>
<td>&lt; 5.47 × 10⁻⁸</td>
<td>36 ± 1</td>
</tr>
<tr>
<td>(25 μm)</td>
<td>1.48 × 10⁻⁶</td>
<td>2.16 × 10⁻⁶</td>
<td>5.5</td>
</tr>
<tr>
<td>(60 μm)</td>
<td>1.20 × 10⁻⁶</td>
<td></td>
<td></td>
</tr>
<tr>
<td>(100 μm)</td>
<td>2.59 × 10⁻⁷</td>
<td></td>
<td></td>
</tr>
</tbody>
</table>

estimates of the component spectra were iteratively adjusted so as to obtain featureless residuals and strictly positive component maps with a minimum degree of correlation with each other. Map areas containing bright stars or H II regions were not represented by the three component spectra and could therefore not be fit simultaneously. Such (isolated) map areas will not be considered in the subsequent analysis. The cool and shock-heated components so derived are overlaid in Fig. 2. The four input maps were...
background corrected by subtracting scaled versions of the cool dust component and a quadratic surface fit to the zodiacal light component from each. Accurate integrated flux densities could then be derived against a nearly featureless background in all four IRAS bands. The 60 µm background corrected map is shown in Fig. 3.

The nominal integrated fluxes were color-corrected by comparison with inband integrations of a model spectrum. The model was generated from an MRN (Mathis et al., 1977) size distribution of silicate and graphite grains at a fixed temperature using the absorption efficiencies derived by Draine and Lee (1984) and tabulated by Draine (1985). The same model was used to derive the wavelength where the spectrum peaks, \( \lambda_{\text{peak}} \), and the peak surface brightness, \( I_{\text{peak}} \) (see Braun, 1986 for a more complete discussion). The nominal and color-corrected fluxes are listed in Table 1 together with the parameters of the spectral peak.

3. Discussion

3.1. Distance, environment, and geometry

The available constraints on the distance to G292.0+1.8 are the H I absorption properties (Caswell et al., 1975) and the optical absorption inferred from [O ii] and [O iii] emission line strengths (Goss et al., 1979; Murdin and Clark, 1979). The neutral hydrogen absorption seen against the radio continuum emission consists of the local contribution at \( V_{\text{lsr}} = 0 \) km s\(^{-1}\) and a narrow peak centered at \( V_{\text{lsr}} = -30 \) km s\(^{-1}\). The fractional absorption in the line, together with the line width and brightness temperature seen in emission imply an H I column density \( N_{\text{HI}} \approx 8 \times 10^{20} \) cm\(^{-2}\) for this feature. However, the Schmidt (1965) rotation curve predicts a most negative velocity of \( V_{\text{lsr}} = -20 \) km s\(^{-1}\) at a distance, \( D = 3.7 \) kpc along this line of sight. The departures from simple rotation, in accordance with the above model, are therefore \( \geq 10 \) km s\(^{-1}\). As the minimal estimate of departure is almost certainly over-optimistic, the range of acceptable lower limits to the distance are those which lie within about 15 km s\(^{-1}\) of the narrow absorption feature, i.e. \( D_{\text{min}} \approx 3.7 \pm 2 \) kpc. The narrowness and non-multiplicity of the profile suggest that the lower part of this range in distance is perhaps most appropriate.

The derivation of optical absorption on the basis of the forbidden oxygen line ratios in collisionally ionized gas is difficult. Current modelling of the intrinsic line ratios (e.g. Raymond, 1979; Shull and McKee, 1979) has illustrated the dependence of line intensities on a wide variety of parameters, while comparison of SNR spectra with these models has illustrated remaining inadequacies, especially in regard to the [O ii] and [O iii] lines (e.g. Dennenfeld, 1982). Simple considerations have however been employed to derive a very approximate visual absorption, \( A_V = 2 \pm 2 \) mag (Murdin and Clark, 1979). In conjunction with a mean interstellar absorption in this direction of 0.5 mag/kpc (Goss et al., 1979 and references therein) a distance, \( D = 4 \pm 4 \) kpc is inferred.

The environment of the SNR is well illustrated by the distribution of cool dust (Fig. 2) derived from the spectral decomposition of the IRAS images in the 12, 25, 60, and 100 µm bands. The galactic plane emission is seen in the south and a probably unrelated H II region complex in the north. An arm of emission extends up from the galactic plane complex to the position of the SNR. The shock-heated dust of the SNR (also in Fig. 2) is seen immediately to the northeast of and partially overlapping the brightest emission from this arm. Since previous studies have shown that moderate to high ISM densities (\( n > 0.2 \) cm\(^{-3}\)) are required for observable emission from shock-heated dust (Braun, 1986) this morphological coincidence is quite suggestive of a physical association. The excess surface brightness of the galactic arm at 100 µm, \( I = 10 \) MJy sr\(^{-1}\) with respect to the diffuse background, corresponds to an H I column density, \( N_{\text{HI}} \approx 10^{20} \) cm\(^{-2}\) (and \( A_V = 0.5 \) mag) assuming an approximately local ISRF (interstellar radiation field, cf. Draine and Anderson, 1985). Since this is the only distinct feature visible at the position of the SNR, and the estimated column density is in good agreement with that seen in H I absorption, it seems likely that this is the same cloud and that it lies immediately in front of and below the remnant. Since 0.5 mag of extinction is then occurring in the immediate vicinity of the SNR, the estimated visual absorption (above) implies the somewhat lower distance, \( D = 3 \pm 4 \) kpc, while the H I value is no longer simply a lower limit.

Considering both of the above constraints suggests the distance, \( D = 3.6 \pm 1.5 \) kpc. At this distance, the 12’ dimensions of the galactic arm correspond to 13 ± 5 pc. Assuming a similar depth along the line-of-sight implies a cloud density, \( n_{\text{H}} = 20 \pm 8 \) cm\(^{-3}\). The cloud ends abruptly at the position of the SNR, suggesting that densities must be significantly lower to the northeast.

3.2. Brightness distribution

The peculiar brightness distribution of G292.0+1.8 at all observed wavelengths has not allowed a simple morphological classification of the SNR type.

The Einstein HRI image (Tuohy et al., 1982) consists of a system of bright knots and arcs superposed on a faint elliptical plateau with no clear indication for limb-brightening at the periphery. The analysis of X-ray data for other young SNRs has indicated that the bright clumps which are often observed are most likely associated with knots of ejecta encountering the interface between diffuse ejecta and the shocked ISM (e.g. Seward et al., 1983), while the outer shock itself is either not seen at all or only faintly under favorable circumstances. The Einstein SSS spectrum (Clark et al., 1980) of G292.0+1.8 supports the view that the bulk of its X-ray emission is due to an enriched, thermal plasma, i.e. shocked ejecta from the inner layers of a massive progenitor.

The optical morphology is well illustrated by the [O ii]\( \lambda 3727 \) image of Murdin and Malin (private communication) shown in Fig. 4. The brightest emission is concentrated in “streamers” extending to the south and east of the bright X-ray knots. Just as in some of the fast moving knots (FMK) in Cas A, the optical spectrum of this material consists almost exclusively of the forbidden lines of O and Ne (Goss et al., 1979; Murdin and Clark, 1979). The spectrum and high radial velocities of this material (Murdin and Clark, 1979; Braun et al., 1983) together imply that it represents knots of high density ejecta undergoing recombination after being shocked on encountering a diffuse high density zone, presumably the inner surface of the diffuse ejecta/shocked ISM layer. The abrupt “turn-on” of the “streamers” to the south and east of the X-ray knots suggests that the optical knots are propagating in these directions and that the arc defined by their northwest edge delineates the inner surface of the high density zone. Braun et al. (1983) have shown that a dominant component of systematic expansion is apparent in the radial velocity data if an expansion center lying to the northwest of this interface is assumed at \( z_{\text{exp}} = 11^{22}21^18 \pm 2^\circ, \delta z_{\text{exp}} = -58^\circ57^225 \pm 15^\circ \). This position is marked in Figs. 1–3.

The radio morphology is well illustrated in the 843 MHz map shown in Figs. 1 and 4. A bright central ridge is elongated along the ridge of [O ii] and X-ray emission, while a faint plateau extends
over roughly 10', coincident with that seen in the X-ray image. A
total flux density at 843 MHz of $S_{843} = 15.5 \pm 0.6$ Jy was found;
in good agreement with the spectrum presented by Lockhart et al.
(1977) ($x = -0.37, S \propto \nu^x$) which predicts $S_{843} = 15.9$ Jy. This
is divided between the bright ridge of (deconvolved) angular size
2.0 $\times$ 1.6 (centered at $\alpha = 11^h 22^m 22^s 9^f 3^g, \delta = -58^\circ 59' 31" \pm 2"$
with p.a. = 120') with a flux density $S_{843} = 6.2 \pm 0.5$ Jy, and the
plateau (approx. 10 arcmin diameter) with flux density $S_{843} = 9$ Jy.
A search for spectral index variations between 843 and 1415 MHz
was carried out with a smoothed version of the map in Fig. 3 and
the Fleurs 21 cm map (Lockhart et al., 1977). Over this restricted
frequency range there is no evidence for significant variations in
spectral index within $\pm 0.2$. The spectrum of both the central and
plateau components is therefore almost certainly nonthermal.
In addition, Shaver (private communication) has searched for the
H109a recombination line using the Parkes 64 m telescope at 6 cm
with a 4 arcmin beam. No line was detected with an upper limit of
0.03 K ($T_B$). If a hypothetical thermal source had an electron
temperature and velocity width typical of galactic H II regions, then
a line to continuum ratio of $\sim 0.05$ would be expected. The upper
limit is, however, an order of magnitude less than this value. Thus,
since the plateau contributes about 50% of the peak intensity as
observed with a 4' beam, it appears that less than 20--30% of the
emission could arise from a conventional H II region.

Similar to the case of X-ray emission, the radio emission from
other young SNRs is dominated by bright clumps associated with
the interaction between ejecta and the inner surface of the diffuse
shell. The region between this interface and the outer shock may be
filled with diffuse emission, while the outer shock itself can only
rarely be distinguished (cf. Braun, 1986).

The distribution of shock-heated dust is illustrated by the
background corrected 60 $\mu$m image shown in Fig. 3. A diffuse
plateau is seen which is approximately coextensive with the
X-ray/radio plateau. The peak of the distribution is, however,
located at $\alpha = 11^h 22^m 18^s \pm 2^s, \delta = -58^\circ 57' 25" \pm 15"$ as the
nominal expansion center.

3.3. Physical parameters

The relative radii of the inner ejecta surface and outer shocks in the
southern portion of G292.0+1.8 indicate that a relatively large
mass ratio $X = M_{ISM}/M_0$ of swept-up to ejected diffuse mass has
been acquired locally. The radial velocity data for the optical knots
is well fit by a spatial-velocity ellipsoid of radius $\theta = 2.6$ and
velocity $V = 2200$ km s$^{-1}$ (Braun et al., 1983) with respect to the
nominal expansion center above. Since the knots are now recombinating, their space velocities must have been $\sim 200$ km s$^{-1}$
greater before being shocked on encountering this surface.
Combining an original space velocity $v_{PMR} = 2400$ km s$^{-1}$ with the
radius $r = 2.7 \pm 1.1$ pc (at the distance $D = 3.6 \pm 1.5$ kpc) yields an
age $\tau = 1100 \pm 460$ yr. The angle averaged radial cut through the
Fig. 5. Angle averaged radial profile through the 843 MHz map of Fig. 1 with respect to the nominal expansion center at $x_0 = 11^\circ 22^\prime 18^\prime$, $\delta_0 = -58^\circ 57^\prime 25^\prime$ (marked in Figs. 1–3) and averaged over position angles, $\phi = 100$ to 225. Note the thick inner shell consistent with radius $\theta \approx 150^\prime$ and the edge of the outer shell at $\theta = 330^\prime$.

radio brightness distribution shown in Fig. 5 indicates a thick shell structure with an implied radius consistent with that found for the optical knots. The faint plateau, however, is seen to extend to a radius of $\theta \approx 5^\prime$, or $r_s = 5.8 \pm 2.4$ pc.

Formulæ describing the pre-adiabatic evolution of supernova remnants are given by Braun (1986) based on previously published one-dimensional numerical simulations. Two useful expressions (his equations B.7 and B.8) are those relating shock radius $r_s$ and velocity $v_s$ to the original shock velocity $v_0$ and current mass ratio $X$:

$$r_s = v_0 (0.75 - 0.127 \log^2 X) X^{-0.199} t,$$  

$$v_s = 0.66 v_0 (0.75 - 0.127 \log^2 X) X^{-0.338} t$$  

Comparison of the current ratio of inner ejecta to shock radii, $r_i/r_s = 0.47$ with the values in Table B.1 of Braun (1986) implies a mass ratio, $X \approx 13.5$. Equation (1) can then be used to derive the original shock velocity, $v_0 = 14,400 \text{ km s}^{-1}$ which corresponds to a current shock velocity, $v_s = 2100 \text{ km s}^{-1}$ via Eq. (2). (It should be noted that the above quantities $X$, $v_0$, and $v_s$ are independent of the distance.) While the original expansion velocity is somewhat higher than that found by Braun (1986) for Cas A ($10,900 \pm 1000 \text{ km s}^{-1}$), it is still within the range of velocities observed for the outer layers of SN II ejecta, 10,000 to 15,000 km s$^{-1}$ (e.g. Searle, 1974).

The infrared spectral data for the shock-heated dust emission (listed in Table 1) corresponds to a spectral peak at $\lambda_{\text{peak}} = 36.0 \pm 1.0 \text{ microns}$ with a maximum surface brightness $\lambda I_{\text{max}} = 2.16 \pm 0.20 \times 10^{-6} \text{ W m}^{-2} \text{ sr}^{-1}$. Draine (1981) and Braun (1986) have discussed how these observable parameters are related to the physical parameters of interest, namely the local shock velocity, $v_s$, pre-shock density, $n_0$, and volume filling factor, $f_v$. The appropriate expressions are (Eqs. (4a), (4c), and (5) from Braun 1986),

$$\lambda I_{\text{max}} \approx 3.34 \times 10^{-6} f_v \left( \frac{r}{400 \text{ yr}} \right)^2 \left( \frac{n_0}{\text{cm}^{-3}} \right)^2 \left( \frac{v_s}{10^5 \text{ km s}^{-1}} \right)^3 \text{ W m}^{-2} \text{ sr}^{-1}$$  

for $n_0 \leq 5 \times 10^2 \text{ cm}^{-3} \text{ yr}$ and $v_s \geq 1800 \text{ km s}^{-1}$

$$\lambda I_{\text{max}} \approx 3.48 \times 10^{-5} f_v \left( n_0 \left( \frac{v_s}{10^5 \text{ km s}^{-1}} \right)^3 \right) \text{ W m}^{-2} \text{ sr}^{-1}$$  

for $n_0 > 5 \times 10^4 \text{ cm}^{-3} \text{ yr}$ and

$$\lambda_{\text{peak}} \approx 37.6 \left( \frac{n_0}{\text{cm}^{-3}} \right)^{-0.18} \left( \frac{v_s}{10^5 \text{ km s}^{-1}} \right)^{-0.3} \mu \text{m}.$$  

Since the shock velocity derived above is based on the southern periphery where infrared emission is brightest (cf. Fig. 3) it should be appropriate for the gas with which the dust is mixed. Equation (4) can then be used to derive the pre-shock gas density, $n_S = 0.37 \text{ cm}^{-3}$. The absence of a discernible reverse shock interface in the northeast suggests that the local mass ratio must be, $X_{\text{NE}} \approx 1$. The ratio of current shock radii can then be derived from equation 1 yielding $r_{\text{SW}}/r_{\text{NE}} \approx 2.1$. The volume encountered by the two hemispheres will scale as this ratio cubed, so that the density must be a factor $X_{\text{SW}} (r_{\text{NE}})^3 / X_{\text{NE}} (r_{\text{SW}})^3 > 130$ less, i.e. $n_{\text{SW}} < 0.003 \text{ cm}^{-3}$. Such a density is consistent with current estimates of the intercloud density in a three phase ISM (e.g. McKee and Ostriker, 1977).

The volume filling factor of the dusty gas can be derived from Eq. (3a), yielding $f_v = 0.2 \pm 0.1$, a value consistent with substantial filling to the southwest. The pre-shock density found above implies a “uniform” swept-up mass $M_{\text{ISM}} = 3 \pi r^2 \mu m_{\text{H}} n_s = 10^{27.1} M_\odot$ (with a mean molecular weight of $\mu = 1.3$) so that the mass of ejecta dynamically coupled to the ISM is $M_\odot = 0.7^{+0.2}_{-0.2} M_\odot$ (using the mass ratio, $X = 13.5$ from above). The current kinetic energy of the swept-up ISM assuming approximately adiabatic internal density and velocity profiles is given by (cf. Braun, 1986),

$$E_{K} = 0.0450 \left( \frac{r}{2 \text{ pc}} \right)^3 \left( \frac{n_0}{\text{cm}^{-3}} \right)^2 \left( \frac{v_s}{10^5 \text{ km s}^{-1}} \right)^2 10^{50} \text{ erg}$$

yielding $E_{K} = 1.8^{+0.3}_{-0.2} 10^{50} \text{ erg}$ corresponding to an original energy $E_0 = E_{K}/0.28 = 6.4^{+0.3+0.1}_{-0.1} 10^{50}$ erg. The cubic dependence on the uncertain distance is responsible for the large uncertainties in the quantities above. However, it is clear that a relatively small mass (in relation to the SN II progenitor mass) $M_\odot < 1 M_\odot$ was involved in powering the blast wave. This is similar to the finding of Braun (1986) for the historical SNR, Tycho, Kepler, and Cas A, all of which indicate (dynamically coupled) ejecta masses $M_\odot \approx 0.35 M_\odot$. A similar mass estimate would be inferred for G292.0+1.8 if a distance $D = 2.9$ kpc were adopted. Such a low distance seems quite plausible. On the other hand, if the far distance $D > 5$ kpc is adopted, an ejecta mass $M_\odot = 1.8 M_\odot$ and initial energy $E_0 = 1.710^{54}$ erg are implied. The quantities derived above are summarized in Table 2.

4. Conclusions

The available morphological and kinematic evidence suggests that G292.0+1.8 is near the edge of an interstellar cloud. The southwestern portion of the SNR appears to be encountering the cloud periphery ($n_0 = 0.37 \text{ cm}^{-3}$) and has evolved to the point that the reverse shock, associated with the bright X-ray/radio rim, is at approximately half the outer shock radius. This is in contrast to younger galactic SNR where the reverse shock is still at $\geq 70\%$ of the outer shock radius (cf. Braun, 1986). The northeastern portion is expanding into an intercloud region where the low densities ($n < 0.003 \text{ cm}^{-3}$) can neither support immersed dust nor generate...
Table 2. Derived physical parameters

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
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</thead>
<tbody>
<tr>
<td>Distance [kpc]</td>
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</tr>
<tr>
<td>Age [yr]</td>
<td>1100 ± 460</td>
</tr>
<tr>
<td>Initial expansion velocity [km s(^{-1})]</td>
<td>14,400</td>
</tr>
<tr>
<td>Diffuse ejecta</td>
<td>0.7(^{+1.3}_{-0.55})</td>
</tr>
<tr>
<td>Mass ([M_\odot])</td>
<td>6.4(^{+11}_{-5})</td>
</tr>
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<td>(E^d_{50}) of diffuse</td>
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<tr>
<td>Shock velocity</td>
<td>2100</td>
</tr>
<tr>
<td>ISM density ([\text{cm}^{-3}])</td>
<td>0.37</td>
</tr>
<tr>
<td>ISM density ([\text{cm}^{-3}])</td>
<td>&lt;0.003</td>
</tr>
<tr>
<td>ISM density ([\text{cm}^{-3}])</td>
<td>&gt;6500</td>
</tr>
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</table>

significant emission through deceleration (except perhaps in hard X-rays). A relatively small mass of diffuse ejecta \(M_\odot \lesssim 1 M_\odot\) has determined the early dynamical evolution of the SNR although the current emission processes (excepting the infrared) are dominated by the slow, clumpy ejecta now encountering the reverse shock. This process has given rise to a center-filled morphology (actually a thick one-sided shell in this instance), which may be a common occurrence for the remnants of type II SN during intermediate evolutionary times (where the mass ratio \(X = 10 \rightarrow 30\)). Such objects are clearly not “Crab-like” as their appearance might suggest (although the slowly expanding clumpy ejecta are likely to form the basis of such objects as well) and can most easily be distinguished from their pulsar driven counterparts by their thermal X-ray spectrum.

The evaporation and/or ablation of the remaining ejecta should contribute to the internal energy of the SNR, except when it is sufficiently clumpy to penetrate the outer shock and escape the system. This suggests a gradual increase in the internal energy available for driving the shock in pre-adiabatic SNR even in the absence of evaporation of engulfed interstellar clouds.

There are indications that the G292.0 + 1.8 the SN energy was abnormally high. A refined distance estimate based on a proper motion study of the fast-moving optical knots would help to resolve this question as well as more fully constrain the geometry.

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