that the number of variables with periods between one and two days is very small. In PRAGER's catalogue for 1936 only 13 variables are found within this interval. Apart from the periods however there are other means to distinguish the RR Lyrae- or cluster-type variables from the δ Cephei-type stars, at least statistically, namely the galactic concentration and the space-velocities.

The galactic concentration of the δ Cephei-type variables is very high, their peculiar velocities on the other hand are only small, while the RR Lyrae-type variables behave in the opposite way. As has been remarked above the galactic latitude of XX Vrg is $+50^\circ$. Four measures of the radial velocity, kindly communicated to me by Dr. A. H. Joy, indicate that this variable has a considerable velocity. Though

it is not impossible, as is proved by W Vrg, that a δ Cephei-type star should show the same features, it seems very probable that XX Vrg should be classified as an RR Lyrae-type variable.

During this investigation a new variable was discovered in the same field. It is marked with an arrow in Figure 1. Its position is:

$$14^h 8^m 9^s, \quad -5^\circ 11' 0'' (1855)$$

or relative to the bright star, B. D. $-5^\circ 38' 43'' (4^m 5)$, $4' 0''$ preceding and $7' 4''$ north. The star has not been measured. On the last plate (J. D. 2427625) its brightness was estimated to be $11^m 1$, on the preceding plate (J. D. 2427596) $13^m 0$. On the remaining plates the star is barely visible or invisible. It probably is a long period variable with a range larger than $2^m 5$.

The short-period variable VV Puppis, by *P. Th. Oosterhoff*.

A discussion of some observations of this interesting variable, obtained with the sixty inch reflector of the Mount Wilson Observatory, has been published in *P. A. S. P.* 47, 322. These observations are given in detail in this note together with two more series of observations, made with the same instrument. The exposure-time was five minutes on Eastman 40 plates. The following comparison stars were used:

	$\Delta \alpha \cos \delta$	$\Delta \delta$	m_{pg}
	(relative to the variable)		
1	+ '80	- 2'48	13'90
2	+ '53	- '99	14'35
3	+ 1'15	- 1'74	14'50
4	+ '08	- 1'71	15'20
5	- 2'26	- '29	15'75
6	- 1'03	+ 1'21	16'20
7	- '50	+ 1'10	16'65

The magnitudes were derived from measures in the Schiltmicrophotometer reduced to the system of MAYALL¹⁾. The magnitude of the faintest com-

¹⁾ *P. A. S. P.* 43, 305.

TABLE I.

J. D. H.A.M.T. Gr. m_{pg} — 2420000	J. D. H.A.M.T. Gr. m_{pg} — 2420000	J. D. H.A.M.T. Gr. m_{pg} — 2420000
d 7889'6383 15'20	d 8159'8947 16'26	d 8160'8920 14'90
'6421 15'07	'8985 15'99	'8959 14'72
'6459 15'26	'9024 15'27	'8997 14'74
'6497 15'31	'9062 15'01	'9035 15'46
'6535 15'56	'9100 14'92	'9073 16'34
'6574 16'08	'9138 14'44	'9111 16'44
'6612 16'24	'9176 14'42	'9150 16'40
'6650 16'12	'9226 14'39	'9205 16'43
'6744 16'26	'9263 15'26	'9243 16'47
'6782 16'28	'9301 16'43	'9282 16'44
'6820 16'26	'9340 16'35	'9320 16'43
'6858 16'25	'9378 16'42	'9358 16'48
'6897 16'24	'9416 16'36	'9396 16'43
'6935 16'36	'9454 16'38	'9434 16'51
'6973 16'20	'9503 16'25	'9490 15'32
'7011 15'84	'9541 16'20	'9528 15'08
'7119 15'27	'9579 16'30	'9566 14'80
'7157 15'36	'9617 16'20	'9604 14'48
'7195 15'36	'9656 16'43	
'7233 15'84	'9694 15'86	
'7275 16'17	'9732 15'10	
'7313 16'09		
'7355 16'28		

FIGURE 1.