New estimates of SV Crucis (Errata: 10 146)

Kort, J.J.M.A. de

TABLE 2 *(continued)*.

B.D.	α 1950·0	δ 1950·0	dependences	B.D. or C.D.	α 1950·0	δ 1950·0	dependences
	h m s	° ′ ″			h m s	° ′ ″	
98 + 8·1826	7 35 1·50	8 16 39·4	·3515	103 + 9·1831	7 56 32·71	9 1 24·6	·1953 ·1889
+ 8·1842	7 37 10·69	8 44 1·8	·4328	+ 9·1835	7 57 36·18	9 6 41·6	·5084 ·5141
+ 8·1845	7 37 27·46	8 16 30·8	·2158	+ 8·1945	7 58 31·56	8 30 59·9	·2963 ·2970
99 + 9·1743	7 38 53·06	9 8 12·7	·1935 ·1924	104 − 9·5709	21 16 28·28	9 22 53·7	·5514 ·5623
LpzII 4066	7 40 58·48	8 23 44·9	·2884 ·2859	− 9·5716	21 18 11·84	9 15 16·9	·1897 ·1843
+ 8·1860	7 42 11·13	8 41 4·1	·5182 ·5218	− 9·5724	21 19 17·73	9 32 25·5	·2590 ·2534
100 + 9·1754	7 41 9·82	9 1 31·5	·2985 ·2968	105 −26·15887	21 52 32·53	25 49 21·8	·2644 ·2665
+ 8·1871	7 44 32·37	8 29 4·8	·3648 ·3634	−25·15655	21 52 39·41	25 15 10·2	·4119 ·4109
+ 9·1764	7 44 33·42	8 49 28·1	·3368 ·3398	−25·15690	21 56 16·71	25 35 36·6	·3236 ·3227
101 + 8·1866	7 43 41·87	8 44 46·5	·3930 ·3905	106 −27·15622	21 44 36·26	27 5 10·7	·6146 ·6151
+ 9·1768	7 45 24·75	8 57 35·5	·3485 ·3499	−26·15817	21 45 36·48	26 27 58·4	·0098 ·0106
+ 8·1884	7 47 8·50	8 40 11·5	·2584 ·2596	−26·15858	21 49 43·98	26 15 29·7	·3756 ·3743
102 + 8·1881	7 46 29·98	8 45 43·7	·4236 ·4159	107 −18·5815	20 54 32·92	18 2 14·2	·0728 ·0732
+ 9·1774	7 46 49·56	9 4 50·9	·2841 ·2857	−18·5816	20 54 42·85	17 55 35·8	·5032 ·5027
+ 8·1890	7 47 36·69	8 42 16·4	·2923 ·2985	−18·5836	20 59 3·98	18 21 11·2	·4240 ·4241

New estimates of SV Crucis, by *J. de Kort*.

The variability of this star of the δ Cephei type was discovered and described by OOSTERHOFF [1]). Estimates have now been made on 458 plates, including the plates discussed by OOSTERHOFF. The following comparison stars have been used:
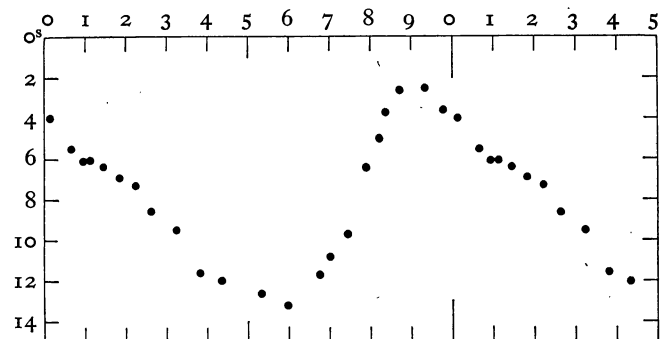
a	b	c	d
^{s}·0	5^s·3	8^s·4	·12^s·3
11^m·3	12^m·1	12^m·6	13^m·2

Of these a, b and c are the comparison stars adopted by OOSTERHOFF, the approximate coordinates of d as read from the diagram in his Figure 3 are 11^h 51^m 20^s·8, − 61°57′·1 (1875). A comparison of these stars with S.A. 193 gave the zero point of the photographic magnitudes, the scale has been obtained from plates taken with a coarse grating in front of the objective.

TABLE 1. Epochs rising branch.

J.D. − 2420000	E	O−C	J.D. − 2420000	E	O−C
3800·95	− 336	− ·70	6028·88	− 18	− ·13
3941·81	− 316	+ ·07	6028·92	− 18	− ·09
3997·68	− 308	− ·09	6113·26	− 6	+ ·20
4285·23	− 267	+ ·28	6113·13	− 6	+ ·07
4285·11	− 267	+ ·16	6119·90	− 5	− ·17
4292·11	− 266	+ ·16	6120·40	− 5	+ ·33
4292·13	− 266	+ ·18	8004·26	+ 264	+ ·04
4292·21	− 266	+ ·26	8004·46	+ 264	+ ·24
4291·82	− 266	− ·13	8003·87	+ 264	− ·35
4291·76	− 266	− ·19	8228·43	+ 296	+ ·07
4585·88	− 224	− ·25	8227·82	+ 296	− ·54
4586·01	− 224	− ·12	8249·10	+ 299	− ·27
5328·77	− 118	+ ·19	8249·41	+ 299	+ ·04
5356·81	− 114	+ ·21	8270·32	+ 302	− ·06
5356·36	− 114	− ·24	8641·97	+ 355	+ ·36
5391·51	− 109	− ·11	8641·93	+ 355	+ ·32
5391·58	− 109	− ·04	8662·63	+ 358	+ ·01
5391·83	− 109	+ ·21	8662·63	+ 358	+ ·01
5616·36	− 77	+ ·60	8662·33	+ 358	− ·29

1) B.A.N. No. 148, 183 (1928), star σ.

The new elements derived for a point on the rising branch, where the brightness 8^s is reached, are

J.D. hel. M.A.T.Grw. 2426155^d·09   + 7^d·00428 E
        ± ·04 (m.e.)   ± ·00018 (m.e.).

This point precedes the equally bright point on the descending branch by ^{P}·476.

The present observations show the slight hump on the descending branch, which is characteristic of periods near 7^d [1]).

For the phases of the mean light curve the following expression has been employed:

$$^{d-1}·14277 \text{ (J.D. hel.M.A.T.Grw.} - 2420000).$$

TABLE 2. Mean light curve.

n	$\bar{P}$	$\bar{s}$	n	$\bar{P}$	$\bar{s}$	n	$\bar{P}$	$\bar{s}$
23	·013	3·97	23	·324	9·50	10	·788	6·42
20	·066	5·49	30	·383	11·59	10	·819	4·97
20	·095	6·10	32	·435	12·00	20	·836	3·73
20	·113	6·06	30	·532	12·63	20	·871	2·60
23	·146	6·39	31	·598	13·21	19	·932	2·52
22	·184	6·92	20	·676	11·68	20	·979	3·58
20	·225	7·29	13	·701	10·81			
20	·262	8·57	12	·744	9·67			

1) B.A.N. No. 96, 115 (1926).

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