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Further remark on the motion of the clusters χ and h Persei
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the new magnitude equation of the direct colour estimates this question deserves a renewed examination.

I have looked up the stars in *B. A. N.* 35, for which the L, P value of c_2/T was below 1.60. For these stars

mean values of c_2/T were calculated, using in addition to O_c and L, P the determinations W, R, K, G, Lb and D, R (*Leiden Ann.* 14, Part 1). Except the white companions to ι Cancri, β Cygni and 31 Cygni, viz:

TABLE 2.

H. R.		Sp.	m_H	$\frac{c_2}{T}$	weight	b	$\frac{v_0}{\sin \lambda}$	$\sin^2 \lambda$
39	γ Pg	B 2	2.87	1.23	177	- 47	+ .02	.99
153	ζ Cs	B 3	3.72	1.32	116	- 10	+ .14	.88
1641	η Ag	B 3	3.28	1.35	116	+ 1	+ .38	.91
1790	γ Or	B 2	1.70	1.20	63	- 15	+ .06	.38
1910	ζ Ta	B 3p	3.00	1.32	116	- 5	+ .14	.62
4787	κ Dr	B 5p	3.88	1.32	94	+ 47	+ .41	.73
5191	η Ur	B 3	1.91	1.37	134	+ 64	+ .33	.59
6092	τ He	B 5	3.91	1.29	85	+ 44	+ .35	.19
6396	ζ Dr	B 5	3.22	1.35	116	+ 34	+ .17	.35
6588	ι He	B 3	3.79	1.31	85	+ 31	+ .02	.08
6787	102 He	B 3	4.32	1.28	85	+ 17	+ .73	.03
7131	δ_1 Ly	B 3	5.51	1.37	85	+ 14	- .60	.04
7298	η Ly	B 3	4.46	1.21	85	+ 12	- .03	.09
7565	12 Vp	B 3	4.91	1.38	85	- 3	+ .61	.19
7613	22 Cy	B 3	4.87	1.40	85	+ 4	- .18	.17
8238	β Cf	B 1	3.32	1.30	191	+ 13	+ .08	.55
8622	10 La	Oe 5	4.64	1.38	148	- 18	+ .02	.69
8781	α Pg	A 0	2.57	1.34	141	- 41	+ .25	.88

H. R. 3474, 7418 and $BD + 46^\circ 2883$, which are still of somewhat doubtful colour, though probably belonging to the whitest stars, I found in this way 17 stars, which are given in Table 2 with the addition of γ Orionis, for which no L, P value has been determined. From Table 2 it appears doubtful whether for any of the stars considered c_2/T is below 1.3 or the effective temperature above 11000°.

These bright white stars do not, like the bright stars extraordinarily yellow for their spectral class, show any sensible concentration towards the milky way, as 7 out of the 17 stars brighter than 5^m have a galactic latitude of more than 30°.

Although the number of stars brighter than 5^m in Table 2 is only 17, the lack of fainter very white stars is noticeable.

Between the limits	1 ^m	2 ^m	3 ^m	4 ^m	5 ^m	
Table 2 contains	2	2	8	5	stars,	
while the numbers	$\frac{1}{2}$	2	8	32		

would correspond to a constant star density in space. It must however be remarked, that a small magnitude equation in c_2/T which may still be present, would considerably affect these figures, as the number of stars increases rapidly with c_2/T .

The median annual parallactic motion reduced to zero magnitude is ".18 corresponding to a parallax of ".043 or an absolute magnitude of $5 \log .043 = -6^m.9$. A star of this absolute magnitude, radiating like a black body with $c_2/T = 1.32$ (mean of the 18 stars), would have a volume 340 times that of our sun. The total volumes of the eclipsing systems of similar spectrum 68 u Herculis and V Puppis, found to be respectively 214 and 815 times that of our sun, are in reasonable agreement with this figure.

Further investigations may considerably alter the two Tables given in the present note. The stars tabulated are merely to be taken as objects deserving special attention for the reasons noted.

Further remark on the motion of the clusters χ and h Persei, by Ejnar Hertzsprung.

Dr. VAN MAANEN calls my attention to the fact, that his paper entitled „The proper motions of stars in and near the double cluster in Perseus” (*Mount Wilson Contrib.* 205) contains also new radial velocities determined by ADAMS and his collaborators of other stars than those given in *A. J.* 648. Thus my appeal to measure the radial velocity of $BD + 57^\circ 568$ was superfluous as this had already been done. The

result, - 39 km/s, answers my question put forward in *B. A. N.* 25 in the negative. Taken together all the 28 radial velocities known of stars in these clusters do not show any trace of perspective effect.

Out of the 17 stars given in Table 2 of *B. A. N.* 25 the radial velocities have been determined for 6 which all agree with the velocity of the clusters.